ARECIBO H\textsc{I} ABSORPTION MEASUREMENTS OF PULSARS AND THE ELECTRON DENSITY AT INTERMEDIATE LONGITUDES IN THE FIRST GALACTIC QUADRANT

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ABSTRACT

We have used the Arecibo telescope to measure the H\textsc{I} absorption spectra of eight pulsars. We show how kinematic distance measurements depend on the values of the Galactic constants $R_0$ and $\Theta_0$, and we select our preferred current values from the literature. We then derive kinematic distances for the low-latitude pulsars in our sample and electron densities along their lines of sight. We combine these measurements with all others in the inner Galactic plane visible from Arecibo to study the electron density in this region. The electron density in the interarm range $48^\circ < l < 70^\circ$ is $0.017^{+0.016}_{-0.003} (68\% \ c.l.) \ cm^{-3}$. This is $0.75^{+0.22}_{-0.19} (68\% \ c.l.)$ of the value calculated by the Galactic electron density model of Cordes & Lazio. The model agrees more closely with electron density measurements toward Arecibo pulsars lying closer to the Galactic center, at $30^\circ < l < 48^\circ$. Our analysis leads to the best current estimate of the distance of the relativistic binary pulsar B1913+16: $d = 9.0 \pm 3 \ kpc$. We use the high-latitude pulsars to search for small-scale structure in the interstellar hydrogen observed in absorption over multiple epochs. PSR B0301+19 exhibited significant changes in its absorption spectrum over 22 yr, indicating H\textsc{I} structure on a $\sim 500 \ AU$ scale.


1. INTRODUCTION

Neutral hydrogen (H\textsc{I}) absorption measurements of pulsar signals at $\lambda = 21 \ cm$ are important probes of various properties of the interstellar medium such as the small-scale structure of cold H\textsc{I} (Dickey et al. 1981; Stanimirović et al. 2003b) and calibrators of the pulsar distance scale and electron density models at large Galactic distances (Weisberg et al. 1979, 1987, 1995; Frail & Weisberg 1990). These measurements are complementary to interferometrically determined parallaxes, which can be utilized on nearer sources (Brisken et al. 2002; Chatterjee et al. 2004). We report new Arecibo\textsuperscript{6} H\textsc{I} absorption observations of eight pulsars, and we use these measurements to study the electron density of the interstellar medium and the small-scale structure of neutral hydrogen clouds.

Sensitivity limitations indicate that few, if any, additional pulsar H\textsc{I} absorption measurements of electron density in the Galactic plane at the intermediate first-quadrant Galactic longitudes accessible to the Arecibo telescope will be made in the next decade or so. Hence, it is an appropriate time to combine the new results with all previous pulsar H\textsc{I} measurements at these longitudes to globally assess the density in the Galactic plane in this region. We use this information to estimate the distance to the relativistic binary pulsar B1913+16. In addition, H\textsc{I} absorption measurements on some of our high-latitude pulsars were originally measured over 20 years ago, during which interval the pulsars have moved many AU through the interstellar medium. Hence, comparison of the old and new absorption spectra yields information on the small-scale structure of the absorbing neutral hydrogen.

The plan of the paper is as follows: The pulsar H\textsc{I} absorption observational technique is sketched in the next section (§ 2). We present H\textsc{I} absorption spectra and kinematic distance results for low-latitude pulsars in § 3. In § 4 we use all measured pulsar distances in the inner Galactic plane visible from Arecibo Observatory to analyze the electron density in that region. In § 4.1 we review the latest work on the galactocentric distance of the Sun and the orbital velocity of the local standard of rest (LSR) in order to select an optimum model to use in refining pulsar kinematic distances and Galactic electron densities. Then in § 4.2 we apply the results of §§ 3 and 4.1 to determine the distance of the relativistic binary pulsar PSR B1913+16. We provide absorption spectra of high-latitude pulsars in § 5, along with analyses of the time variability of absorption in those cases in which earlier epoch data are available. Finally, we discuss our conclusions in § 6.

2. OBSERVATIONS

All observations were made with the 305 m Arecibo telescope from 1998 to 2000. The radio frequency signals near 1420 MHz were mixed to baseband, sampled, and recorded with the Caltech Baseband Recorder (CBR; Jenet et al. 2001), a 10 MHz bandwidth, fast-sampling receiver backend. Every 100 ns, the CBR sampled complex voltages with four-level digitization from the two orthogonally polarized feed channels, and recorded the samples to tape for later processing. (For additional details of the observing techniques and equipment, see Stanimirović et al. 2003a, 2003b.) The data were then corrected for quantization (Jenet & Anderson 1998), Fourier transformed, and folded modulo the apparent pulsar period with the supercomputers of Caltech’s Center for Advanced Computing Research, resulting in data cubes consisting of intensity as a function of pulsar rotational phase (128 phase bins) and radio frequency (2048 frequency channels,
each having 1 km s\(^{-1}\) width). Subsequent processing at Carleton College collapsed the data cubes into two spectra: the pulsar-on spectrum is a sum of the spectra gathered during the pulsar pulse weighted by \(I_{psr}(\phi)\), while the sum of those gathered in the interval between pulses is called the pulsar-off spectrum. Here \(I_{psr}(\phi)\) is the broadband pulsar intensity in a given pulse phase bin \(\phi\).

Two final spectra are formed and displayed for each pulsar. The pulsar absorption spectrum, which represents the spectrum of the pulsar alone less any absorption caused by intervening neutral hydrogen, is the normalized difference of the pulsar-on and pulsar-off spectra. In order to maximize sensitivity, multiple integrals are summed with a weight depending on the square of the pulsar signal strength at the velocity where one finds both no significant absorption, and an emission feature of \(i\) \(A\)-line, \(i\) \(A\)-line and \(i\) \(H\)-line, are both given.

The farthest observed absorption feature is at \(v = 7.4\) km s\(^{-1}\). However, it was shown by Weisberg et al. (1979) that emission spectrum pair. We also add and subtract velocities of 7 km s\(^{-1}\) to our nominal velocities, in order to derive estimates of the uncertainties in distance limits due to streaming and random gas motions in the Galaxy. These procedures are identical to those used in the critical evaluation of all such measurements by Frail & Weisberg (1990) and by all subsequent pulsar \(H\) \(I\) absorption experimenters. Hence, our results are directly intercomparable with these earlier ones. Our derived distance limits are discussed below for each pulsar, and are summarized in Table 1. See § 4.1 for a discussion of modifications to our derived distances for different values of the Galactic constants.

3.1. PSR J1909+0254 = B1907+02

The farthest observed absorption feature is at \(v = 51.9\) km s\(^{-1}\), well before the tangent point. Hence, the pulsar lies beyond a
lower distance limit of 3.8 ± 0.5 kpc. Unfortunately, the H i emission has \( T_b \leq 10 \) K at all velocities where one might test for an upper distance limit via lack of absorption, because the \( b = 2.7^\circ \) line of sight rapidly leaves the hydrogen layer. Since significant absorption could not in any case be guaranteed at the velocity of such weak emission (Weisberg et al. 1979), no upper distance limit can be set.

3.2. PSR J1922+2110 = B1920+21 (\( l = 55.3^\circ; b = 2.9^\circ \)), Fig. 2

Weisberg et al. (1987) observed this pulsar. Their H i spectrum was contaminated by interstellar scintillation and was rather noisy, prompting us to reobserve it with greater sensitivity. The highest velocity absorption that Weisberg et al. (1987) could confidently detect was at \( v = 26 \) km s\(^{-1}\), leading to their limit of 1.9 ± 0.7 kpc \( \leq d \). Our new observations clearly exhibit absorption at \( v = 41 \) km s\(^{-1}\), near the tangent point. Hence, we revise the lower distance limit significantly upward: 4.8 ± 1.8 kpc \( \leq d \). The Weisberg et al. (1987) spectrum also showed an absorption feature at this velocity, but it was not sufficiently above the noise to serve as a reliable kinematic distance indicator. The dip in the current absorption spectrum at \( v = -48 \) km s\(^{-1}\) is probably noise and will not be used to set a distance limit. (It appears much more prominently in one of the two circular polarization channels than in the other, and is not visible in the Weisberg et al. [1987] absorption spectrum. Furthermore, the nearby source G55.6+2.3 \( = \) B1923+210 does not exhibit reliable far-side absorption until \( v \approx -60 \) km s\(^{-1}\) [Dickey et al. 1983; Colgan et al. 1988].) Conversely, the lack of absorption in the \( T_b = 41 \) K emission feature at \( v = -65 \) km s\(^{-1}\) sets the upper distance limit of \( d \leq 16.2 \pm 1.0 \) pc.

3.3. PSR J1926+1648 = B1924+16 (\( l = 51.9^\circ; b = 0.1^\circ \)), Fig. 3

While the Fich et al. (1989) model predicts a tangent point velocity of 48 km s\(^{-1}\) in this direction, we detect significant emission and absorption well beyond this velocity, with the last major feature centered near \( v = 61 \) km s\(^{-1}\). These features are due to the Sagittarius arm (Burton 1970). Garwood & Dickey (1989) and Colgan et al. (1988) also observe H i absorption out to these velocities in the nearby sources G50.625–0.031 and G51.4–0.0, respectively. We choose the tangent point as our lower distance limit, so 5.2 ± 1.8 kpc \( \leq d \). The lack of absorption in the \( T_b \sim 42 \) K emission feature at \( v \approx -47 \) km s\(^{-1}\) yields the upper distance limit: \( d \leq 14.9 \pm 0.8 \) kpc.

4. ANALYSIS OF THE ELECTRON DENSITY IN THE INNER GALACTIC PLANE VISIBLE FROM ARECIBO

The currently reported H i absorption distance measurements are probably among the last to be determined at Arecibo in the foreseeable future, as all known pulsars having sufficient flux density to achieve adequate S/N in a reasonable time have now been measured as well as is possible with this procedure. Unfortunately, Galactic H i emission is now the dominant noise source, so that future receiver improvements will not significantly decrease the overall noise. Therefore, new low-latitude pulsars discovered in the future in this longitude range will probably be too faint for the H i absorption technique, even at Arecibo. Consequently, now is a good time to review all such measurements and their implications for the electron density in the inner Galactic plane visible from Arecibo.

The dispersion measure DM of a pulsar, derived from multiple-frequency timing measurements and reported in Manchester et al. (2005),

\[
\langle n_e \rangle \equiv \frac{\int n_e \, ds}{d} = \frac{DM}{d}.
\]

Table 2 lists the electron density limits derived in this fashion from our measurements and from all earlier ones in similar directions; i.e., toward the inner Galactic plane accessible to the Arecibo telescope. Note that the kinematic distances listed in Table 2 were all derived with the standard procedures discussed above in § 3, namely, the uniform criteria established by Frail & Weisberg (1990) and the IAU standard Galactic constants (Kerr & Lynden-Bell 1986). Procedures to modify these values for other choices of Galactic constants are given below in § 4.1.

In Figure 4a we display measured limits on the mean electron density as a function of Galactic longitude \( l \) in the inner Galactic plane visible from Arecibo. Only those pulsars from Table 2 with \( |b| < 3^\circ \) and a lower distance limit \( d_l \geq 1 \) kpc are included in

\[\text{The catalog is maintained and updated at http://www.atnf.csiro.au/research/pulsar/psrcat.}\]
Figure 4a in order to concentrate on kiloparsec-scale averages in the Galactic plane. Almost all of these measurements were derived from Arecibo pulsar absorption spectra. While there is significant scatter, it is apparent that the measured densities tend to decline as the line of sight rotates away from the inner Galaxy, with a notable drop in the region between the Sagittarius arm (the first arm interior to our location) and the local spiral arm (see Fig. 5). Ables & Manchester (1976) were the first to have adequate \( \langle n_e \rangle \) measurements to infer that interarm densities are lower than those in spiral arms. As illustrated in Figure 4a, the numerous measurements made since that epoch serve to confirm and refine their suggestion, at least in this particular interarm region.

Figure 4b exhibits the measured limits on \( \langle n_e \rangle \) normalized by the mean electron densities predicted by the NE2001 model. The measured-to-model electron density ratio hovers near 1 (with some inevitable scatter) for \( l \leq 48^\circ \), indicating that the model fits the data adequately at those longitudes. However, at the higher longitudes \( 48^\circ < l < 70^\circ \) discussed above, there is a tendency for the typical measured-to-model density ratio to lie below 1, suggesting that the NE2001 model densities are somewhat high in this interarm region.

In order to further explore the electron density in this interarm region, we performed Monte Carlo simulations to assess the best value and uncertainty on \( \langle n_e \rangle \) and on the measured-to-model \( \langle n_e \rangle \) ratio for the pulsars in this region. Of the 10 pulsars with measured distance limits and limits on \( \langle n_e \rangle \) in this range (see Table 2), we discarded the data from PSR J1935+1616, since it lacks a measured upper distance and lower density limit. For each of the remaining nine pulsars, all possible distances between the measured upper and lower limits were rendered equally probable in our simulations by choosing distances randomly from a uniform distribution between the measured distance limits. Note that the assumed uniform probability is the simplest reasonable hypothesis for measurements such as these which have only a lower limit.

Therefore, it is useful to discuss improved measurements of the ratio of measured electron density to the NE2001 model value, vs. longitude.

\[ \frac{d}{l} \]

were used; see text for procedures to rescale with other values. Pulsars in Table 2 are not shown. Top: Measured electron density limits (error bars) and NE2001 model electron densities (circles) vs. longitude. Bottom: Ratio of measured electron density to the NE2001 model value, vs. longitude.

4.1. The Effects of Improved Galactic Constants \( R_0 \) and \( \Theta_0 \)

Much progress has been made on Galactic structure and kinematics in the \( \sim 20 \) years since the IAU constants were defined. Therefore, it is useful to discuss improved measurements of \( R_0 \) and \( \Theta_0 \), and to investigate the influence of these better values on the kinematic distances and electron densities discussed in §§ 3 and 4.

From remarkable proper motion and radial velocity measurements of the orbit of star S2 about Sgr A’ through much of its 15 yr period, Eisenhauer et al. (2005) find that the distance to the Galactic center \( R_0 = 7.62 \pm 0.32 \) kpc. Reid & Brunthaler (2004) have made equally stunning interferometric measurements of the proper motion of Sgr A’, which yield an angular velocity of the LSR about the Galactic center, \( \Omega_0 = \Theta_0 / R_0 = 29.45 \pm 0.15 \) km s\(^{-1}\) kpc\(^{-1}\). Both of these measurements are notable in that they rest on far fewer assumptions than earlier determinations, and hence should be freer of systematic errors. Consequently, we adopt them in what follows. An even newer estimate of \( R_0 \) from infrared measurements of red clump giants in the Galactic bulge (Nishiyama et al. 2006) yields a value of \( 7.52 \pm 0.10 \) (statistical) \( \pm 0.35 \) (systematic) kpc, which is consistent with our above adopted value.

Consider material traversing a circular orbit in the plane at galactocentric radius \( R \). Given a flat rotation curve with circular velocity \( \Theta_0 \), its radial velocity with respect to the LSR is \( v_{\text{rad}} = \Theta_0 \sin l \).

![Figure 4](image)

**Fig. 4.**—Mean electron density in the inner Galactic plane visible from Arecibo Observatory. The IAU (Kerr & Lynden-Bell 1986) values of \( R_0 \) and \( \Theta_0 \) were used; see text for procedures to rescale with other values. Pulsars in Table 2 with \( d_l < 1 \) kpc or \( |b| > 3^\circ \) are not shown. Top: Measured electron density limits (error bars) and NE2001 model electron densities (circles) vs. longitude. Bottom: Ratio of measured electron density to the NE2001 model value, vs. longitude.

\[ \Theta_0(R_0/R - 1) \sin l. \]

With \( R = (R_0^2 + d^2 - 2R_0d \cos l)^{1/2} \), the distance as a function of \( v_{\text{rad}} \) is given by

\[ d = R_0 \left( \cos l \pm \sqrt{1 + \frac{v_{\text{rad}}^2}{\Theta_0^2 \sin^2 l} - \sin^2 l} \right). \]

Hence, our new adopted Galactic parameters, giving a small change in \( \Theta_0 \) but a relatively large change in \( R_0 \), result in a recalibration of the distance in a particular direction, accomplished by rescaling it with an approximately constant factor of \( R_{0,\text{new}} / R_{0,\text{old}} = 7.62 / 8.50 = 0.896 \). Kinematic distance limits and Galactic distance models that use the old IAU Galactic constants should be adjusted by this factor to reflect the improved constants.

Figure 6 displays distance—radial velocity curves toward \( l = 50^\circ \) derived with the old and new flat Galactic rotation curves. Figure 6a illustrates that the old and new model distances differ primarily by a constant multiplicative factor, as indicated in the previous analysis. It follows from the above that when distance is normalized by \( R_0 \), the old and new radial velocity—distance curves are almost identical (Fig. 6b). Such normalized distances are useful in the analysis of the acceleration of pulsars in the Galactic gravitational field, as is discussed in the next section (§ 4.2).

As described above, the new Galactic constants would result primarily in the measured distances being rescaled by the ratio of new to old \( R_0 \). The electron densities, including model densities that have been calibrated via pulsar kinematic distance measurements, would then be rescaled by the inverse of this ratio (see eq. [1]). Hence, Figure 4a would be thus rescaled; while Figure 4b would remain unchanged, since it displays the ratio of two densities, each of which should be rescaled by the same factor.

4.2. The Distance to the Relativistic Binary Pulsar B1913+16

The relativistic binary pulsar B1913+16 lies in the heart of the region of the Galaxy that we are studying, at \((l, b) \sim (50^\circ, 2^\circ)\).
The orbital decay of this system due to gravitational wave emission provides a strong test of relativistic gravitation (Taylor & Weisberg 1989). Currently, the precision of the relativistic test is limited by uncertainties in the Galactic acceleration of the pulsar and the solar system (Damour & Taylor 1991), which are currently dominated by uncertainties in the pulsar distance \( d \), and the solar system (Damour & Taylor 1991), which are currently dominated by uncertainties in the pulsar distance \( d \) and the solar system. This is illustrated by the center-left panel of Figure 4, which shows that the measured electron density appears to have a local minimum near the pulsar’s longitude. The relatively low value of \((n_e)\) at \( l \sim 50°\) is not surprising, since we noted in § 4 that the line of sight in this direction traverses a long path between spiral arms (see Fig. 5). The tightest limits on \((n_e)\) in the whole inner plane visible from Arecibo are for PSR B1915+13 at \( l = 48°\), which happens also to be the pulsar closest to PSR B1913+16. Hence, we adopt the measured limits from B1915+13, including an allowance for up to 7 \( \text{km s}^{-1} \) noncircular velocities (not shown in Table 2 or Fig. 4): \((n_e) \sim (0.0188 \pm 0.006) \times (8.5 \text{ kpc}/R_0)^{-3}\). Note that this value is congruent with and has modestly tighter limits than the electron density derived in this interarm region in § 4.

Armed with our new estimate of \((n_e)\), we use equation (1) to find the distance to PSR B1913+16, which has DM \( \sim 169 \text{ pc cm}^{-3} \):

\[
d_{\text{B1913+16}} = (10.0 \pm 3.2) \left( \frac{R_0}{8.5 \text{ kpc}} \right) \text{ kpc.} \tag{3}
\]

Our result can also be compared with the Cordes & Lazio (2002) NE2001 model distance of 5.90\( (R_0/8.5 \text{ kpc}) \), which is lower for the reasons discussed above. Damour & Taylor (1991) used techniques similar to ours to estimate that \( d_{\text{B1913+16}} \approx (9.2 \pm 1.4)\( (R_0/8.5 \text{ kpc}) \). While the error bars of their and our estimates overlap significantly, we are not able to justify tightening them to the degree that Damour & Taylor (1991) did. Finally, adopting what we judged in § 4.1 to be the current best measurement of \( R_0 \) \( (R_0 = 7.62 \pm 0.32 \text{ kpc};\) Eisenhauer et al. 2005) leads to our best estimate of the distance to PSR B1913+16: 

\[
d_{\text{B1913+16}} = 9.0 \pm 3 \text{ kpc.}
\]

As noted above, the uncertainty in the general relativistic orbital decay rate of this pulsar is currently dominated by the uncertainty in the quantity \( \delta d_{\text{B1913+16}} \). For these purposes, it is useful to give

\[
\delta_{\text{B1913+16}} = \frac{d_{\text{B1913+16}}}{R_0} = 1.18 \pm 0.38. \tag{4}
\]

### 5. High Galactic Latitude Pulsars and Searches for Temporal Variability of Absorption

Relatively nearby pulsars are excellent targets for studies of the tiny-scale atomic structure (Heiles 1997). Comparison of their \( \text{H} \) absorption spectra at multiple epochs as they move across the sky permits us to study structure down to AU scales. We discuss here our measurements of five such pulsars. All but two have been previously observed, so we search for changes between the previous and current epochs. The pulsars’ observing parameters, distances, and transverse velocities are listed in Table 3.

The pulsars discussed in this section are all located at high Galactic latitudes (in all cases, \(|b| > 25°\)). Since their lines of sight leave the Galactic hydrogen layer fairly quickly, they are not amenable to the kinematic distance technique discussed in § 3. Conversely, we cannot search for temporal absorption variations in the low-latitude pulsars of \( \frac{3}{2} \) 9 because two have never before...
been observed, and the third’s absorption spectrum was too noisy at the earlier epoch to yield meaningful differences. (Indeed, we reobserved it in order to more confidently measure its kinematic distance because the noise in the earlier spectrum precluded a secure result.)

5.1. \textit{J0304+1932 = B0301+19} \hfill (l = 161.1°; b = −33.3°),

\textit{Figs. 7 and 8}

The H\textsc{i} absorption spectrum of this pulsar was first observed in 1976–1977 by Dickey et al. (1981). That spectrum and our new one are plotted together in Figure 7. Both epochs’ spectra have a similar \textit{1 km s}^{-1} resolution. Both the \textit{integrated} optical depth (the equivalent width) and the depth of most individual absorption components have changed significantly between the two epochs. Although the early observations were gathered with a spectrometer that digitized the signal to only one bit, its response to rapidly varying pulsar signals was very well characterized by Weisberg (1978). Careful quantization corrections have also been made for the current epoch (Jenet & Anderson 1998). Hence, we believe that both spectra are accurate and that the change is real.

The channel spacing of the old and new spectra are 1.056 and 1.031 \textit{km s}^{-1}, respectively. In order to study the absorption changes more carefully, we slightly coarsened the spacing of the new data via interpolation to match the channel spacing and center frequency of the old data. We also ensured that the old and new channel frequencies matched by verifying that the (essentially noise-free) \textit{emission} spectra were consistent to less than a channel width. Figure 8 displays the old and (resampled) new absorption spectra (\textit{top}) and their difference (\textit{bottom}). The difference spectrum also displays a \pm 2 \sigma noise envelope in order to assess the significance of variations. (The noise envelope grows at the central velocities because the H\textsc{i} emission line itself contributes significantly to the system noise [Johnston et al. 2003; Stanimirović et al. 2003b].) The difference spectrum clearly shows a general trend whereby the absorption line depth is greater at the later epoch in most of the central channels, with the largest single-channel difference being significant at the 2.6 \sigma level.

The \textit{22 yr} time baseline is unique in being one of the longest extant showing absorption variations, leading to a length scale of \textit{~500 AU}, which is second only to the PSR 1557–50 scale of
6. CONCLUSIONS

We have determined the H i absorption spectra of eight pulsars. The three low-latitude pulsars yield kinematic distances and electron densities in the inner Galactic plane visible from Arecibo. These observations mark the completion of a two-decade effort to accurately measure the H i absorption spectrum of all such pulsars that are strong enough to be accessible to this technique. Therefore, we have combined our new measurements with all others in this direction to study the electron density in this region. We find that the mean electron density in the plane in the interarm range $48^\circ < l < 70^\circ$ is $0.017_{-0.007}^{+0.012}$ (68% c.l.) cm$^{-3}$, which is $0.75_{-0.22}^{+0.46}$ (68% c.l.) of the Cordes & Lazio (2002) model value currently used by most researchers. At the lower longitudes accessible to Arecibo ($30^\circ < l < 48^\circ$), the Cordes & Lazio (2002) model appears to conform generally to the measurements, aside from expected local variations.

As part of the process, we show how to modify kinematic distances and electron densities as a function of the Galactic constants $R_0$ and $\Theta_0$. We review recent efforts to determine the values of these constants and select the best current answers. Applying all of these results to the relativistic binary pulsar B1913+16, we find a dispersion measure distance of $d = (10.0 \pm 3.2)(R_0/8.5$ kpc), or $d = 9.0 \pm 3$ kpc if we adopt our current best choice for $R_0$ (Eisenhauer et al. 2005).

The five high-latitude pulsars are most useful in multiepoch studies of small-scale structure in the interstellar medium. Two were observed for the first time by us in this experiment. Of the other three, two showed no measurable absorption at either of two epochs separated by 22 yr, while one, PSR B0301+19, exhibited a significant change in absorption profile over that time span, indicating H i structure on a $\sim 500$ AU scale.

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~1000 AU (Johnston et al. 2003). Further implications of these results for small-scale structure in the interstellar medium will be analyzed in a separate paper (S. Stanimirović et al. 2008, in preparation).

5.2. J1239+2453 = B1237+25 ($l=252.5^\circ; b=86.5^\circ$), Fig. 9

Dickey et al. (1981) also measured the H i absorption spectrum of this pulsar in 1976–1977. The original and current pulsar spectra show no absorption. These results are not surprising, given the Galactic polar line of sight and the very weak H i emission in this direction. The pulsar moved ~2200 AU between the two observations.

5.3. J1537+1155 = B1534+12 ($l=19.8^\circ; b=48.3^\circ$), Fig. 10

This pulsar is a member of a double neutron star binary system. Its pulse timing parallax gives $d > 0.67$ kpc (Stairs et al. 2002). Under the assumption that general relativity provides the correct description of gravitational wave emission, the excess observed orbital period change not attributable to gravitational wave damping yields $d = 1.02 \pm 0.05$ kpc (Stairs et al. 2002). There is no statistically significant absorption evident in our pulsar spectrum along this high Galactic latitude line of sight. No previous H i absorption measurements have been made on this source.

5.4. J1543+0929 = B1541+09 ($l=17.8^\circ; b=45.8^\circ$), Fig. 11

No statistically significant absorption is present in this high Galactic latitude pulsar’s spectrum. This is the first absorption spectrum from this pulsar.

5.5. J2305+3100 = B2303+30 ($l=97.7^\circ; b=-26.7^\circ$), Fig. 12

The broad, shallow dip in the displayed spectrum disappears in one of our two orthogonal polarizations, leading us to conclude that it is not real. This pulsar was also observed in 1976–1977 by Dickey et al. (1981). In that case as well, no significant absorption was seen. The pulsar traversed ~1600 AU in the intervening time.
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