Herschel observations of gas and dust in comet C/2006 W3 (Christensen) at 5 AU from the Sun

M. de Val-Borro1,2, D. Bockelée-Morvan3, E. Jehin4, P. Hartogh1, C. Opitom4, S. Szutowicz4, N. Biver3, J. Crovisier3, D. C. Lis6, L. Rezac1, Th. de Graauw1, D. Hutsemékers4, C. Jarchow1, M. Kidger8, M. Küppers9, L. M. Lara10, J. Manfroid4, M. Rengel1, B. M. Swinyard11,12, D. Teyssier8, B. Vandenbussche13, and C. Waelkens13

1 Max-Planck-Institut für Sonnensystemforschung, Justus-von-Liebig-Weg 3, 37077 Göttingen, Germany
e-mail: hartogh@mps.mpg.de
2 Department of Astrophysical Sciences, Princeton University, Princeton, NJ 08544, USA
e-mail: valborro@princeton.edu
3 LESIA, Observatoire de Paris, CNRS, UPMC, Université Paris-Diderot, 5 place Jules Janssen, 92195 Meudon, France
4 Institut d’Astrophysique et de Géophysique, Université de Liège, 4000 Liège, Belgium
5 Space Research Centre, Polish Academy of Sciences, Bartycka 18A, 00-716 Warsaw, Poland
6 California Institute of Technology, Pasadena, CA 91125, USA
7 Joint ALMA Observatory, Alonso de Córdova 3107, 763 0355 Vitacura, Santiago, Chile
8 Herschel Science Centre, ESAC, European Space Agency, 28691 Villanueva de la Cañada, Madrid, Spain
9 Rosetta Science Operations Centre, ESAC, European Space Agency, 28691 Villanueva de la Cañada, Madrid, Spain
10 Instituto de Astrofísica de Andalucía (CSIC), Glorieta de la Astronomía s/n, 18008 Granada, Spain
11 RAL Space, Rutherford Appleton Laboratory, Chilton, Didcot OX11 0QX, UK
12 Department of Physics and Astronomy, University College London, Gower St., London WC1E 6BT, UK
13 Instituut voor Sterrenkunde, Katholieke Universiteit Leuven, 3001 Leuven, Belgium

Received 15 January 2014 / Accepted 26 February 2014

ABSTRACT

Context. Cometary activity at large heliocentric distances is thought to be driven by outgassing of molecular species more volatile than water that are present in the nucleus. The long-period comet C/2006 W3 (Christensen) was an exceptional target for a detailed study of its distant gaseous and dust activity.

Aims. We aimed to measure the H2O and dust production rates in C/2006 W3 (Christensen) with the Herschel Space Observatory at a heliocentric distance of ~5 AU and compared these data with previous post-perihelion Herschel and ground-based observations at ~3.3 AU from the Sun.

Methods. We have searched for emission in the H2O and NH3 ground-state rotational transitions, J=K (1→0) at 557 GHz and J=K (1→0) at 572 GHz, simultaneously, toward comet C/2006 W3 (Christensen) with the Heterodyne Instrument for the Far Infrared (HIFI) onboard Herschel on UT 1.5 September 2010. Photometric observations of the dust coma in the 70 µm and 160 µm channels were acquired with the Photodetector Array Camera and Spectrometer (PACS) instrument on UT 26.5 August 2010.

Results. A tentative 4σ H2O line emission feature was found in the spectra obtained with the HIFI wide-band and high-resolution spectrometers, from which we derive a production rate of (2.0 ± 0.5) × 1027 molec s−1. A 3σ upper limit for the ammonia production rate of <1.5 × 1027 molec s−1 is obtained taking into account the contribution from all hyperfine components. The dust thermal emission was detected in the 70 µm and 160 µm filters, with a more extended emission in the blue channel. We fit the radial dependence of the surface brightness with radially symmetric profiles for the blue and red bands. The dust production rates, obtained for a dust size distribution index that explains the fluxes at the photocenters of the 70 µm and 160 µm PACS images, lie in the range 70 kg s−1 to 110 kg s−1. Scaling the CO production rate measured post-perihelion at 3.20 AU and 3.32 AU, these values correspond to a dust-to-gas production rate ratio in the range 0.3–0.4.

Conclusions. The blueshift of the water line detected by HIFI suggests preferential emission from the subsolar point. However, it is also possible that water sublimation occurs in small ice-bearing grains that are emitted from an active region on the nucleus surface at a speed of ~0.2 km s−1. The dust production rates derived in August 2010 are roughly one order of magnitude lower than in September 2009, suggesting that the dust-to-gas production rate ratio remained approximately constant during the period when the activity became increasingly dominated by CO outgassing.

Key words. comets: individual: C/2006 W3 (Christensen) – submillimeter: planetary systems – techniques: photometric – techniques: spectroscopic

1. Introduction

Comets are small solar system bodies with a wide range of orbital periods. They spend most of their time in the cold outer regions of the solar system, beyond the distance at which water condenses into ice grains – the so-called snow line. Therefore, they contain pristine material that reflects the chemical composition in the early solar nebula. Comets are composed of a loose aggregation of volatile icy materials and refractory particles. Observations of gas production in comets provide a unique opportunity to constrain the composition of cometary ices and
to characterize a classification scheme based on their relative chemical abundances (A’Hearn et al. 1995; Biver et al. 2002b; Bockelée-Morvan et al. 2004; Crovisier et al. 2009). Studying the processes responsible for ice sublimation in comets is also important in the understanding of the activity and thermal properties of cometary nuclei. In addition, the simultaneous imaging of the thermal emission from dust particles released from the nucleus in two wavelengths bands can constrain important properties of the dust coma, such as the dust size distribution index and dust production rate (e.g., Jewitt & Luu 1990; Bockelée-Morvan et al. 2010b). Therefore, observations of cometary dust and gas can provide clues about the physical conditions in the solar nebula, and help to establish a link between the materials in the parent interstellar cloud and cometary nuclei (see e.g., Ehrenfreund et al. 2004).

Cometary activity at small heliocentric distances is driven mainly by sublimation of water molecules from the nucleus with CO$_2$/H$_2$O and CO/H$_2$O mixing ratios that show a great diversity and range from ~0.01–0.1 in most cometary atmospheres (e.g., Bockelée-Morvan 2011; de Val-Borro et al. 2013). Although H$_2$O is generally the primary component of the nucleus, the sublimation of water is inefficient at $r_h$ larger than 3 AU to 4 AU and the activity is dominated by molecules such as CO$_2$ and CO that are more volatile than water (see e.g., Bockelée-Morvan et al. 2010b; McKay et al. 2012). Therefore, the evolution of the production rate of these molecular species with heliocentric distance provides important clues for understanding the structure and composition of cometary nuclei (e.g., Biver et al. 2002a).

The Herschel Space Observatory (Pilbratt et al. 2010) has proven to be the most sensitive facility for directly observing water emission in distant comets and studying the chemical composition of cometary material (see e.g., Hartogh et al. 2011; Bockelée-Morvan et al. 2012; Lis et al. 2013). Water emission has been detected in the distant comet 29P/Schwassmann-Wachmann 1 (Bockelée-Morvan et al. 2010a). Moreover, recent attempts have been made to directly detect sublimating water using Herschel in comets C/2006 W3 (Christensen) (Bockelée-Morvan et al. 2010b) and C/2012 S1 (ISON) (O’Rourke et al. 2013a) at large heliocentric distances, and also in active objects orbiting within the main asteroid belt, that is, main-belt comets (MBCs; see e.g., de Val-Borro et al. 2012b; O’Rourke et al. 2013b).

Comet C/2006 W3 (Christensen) is a long-period comet that was discovered in November 2006 at a distance of 8.6 AU from the Sun. It passed perihelion on 6 July 2009 at a heliocentric distance of 3.13 AU. As noted previously, water ice sublimation is expected to be ineffective at large heliocentric distances and, owing to the low outgassing rate that requires very sensitive observations, it has only been directly detected previously in comet C/2006 W3 (Christensen) at infrared wavelengths in the pre-perihelion observations at 3.66 AU and 3.13 AU by the Akari satellite on UT 21.1 December 2008 and UT 16.8 June 2009, respectively (Otsuibo et al. 2012).

At the time of the previous Herschel post-perihelion observations of comet C/2006 W3 (Christensen) at 3.3 AU from the Sun (Bockelée-Morvan et al. 2010b, hereafter Paper I) the Heterodyne Instrument for the Far Infrared (HIFI; de Graauw et al. 2010) was not available due to a single event upset in the memory of the Local Oscillator Control Unit microcontroller caused by a cosmic-particle impact. Had HIFI observations been obtained at that time, water emission in the coma would have been readily detected. Nonetheless, an upper limit to the water production rate of <1.4 × 10$^{28}$ molecules s$^{-1}$ was derived from spectroscopic observations with the Photodetector Array Camera and Spectrometer (PACS; Poglitsch et al. 2010) on UT 8.8 November 2009, and an upper limit of <4 × 10$^{27}$ molecules s$^{-1}$ was obtained from observations with the Spectral and Photometric Imaging REceiver (SPIRE; Griffin et al. 2010) on UT 6.6 November 2009. Additionally, the dust size distribution and dust production rate were calculated from the spectral energy distribution measured by PACS (see Paper I). Complementary production rate measurements of several species (namely CO, CH$_3$OH, HCN, H$_2$S, and OH) were obtained from the ground in pre-perihelion observations with the Nançay radio telescope and post-perihelion observations with the Institut de Radioastronomie Millimétrique (IRAM) 30-m telescope (Paper I).

Here we present the analysis of more Herschel outbound observations of comet C/2006 W3 (Christensen) at 5 AU from the Sun. The paper is structured as follows: in Sect. 2, we summarize the observations performed with HIFI and PACS, and the data reduction method. Section 3 presents the results of the data analysis. In Sect. 3.1, we compare the gas production rate obtained by HIFI with measurements by other facilities at various heliocentric distances. The derived dust production rates using a dust thermal emission model are compared with the values obtained in the previous PACS observations in Sect. 3.2. Finally, we discuss the main results of this work in Sect. 4 and present the conclusions in Sect. 5.

2. Observations

The Herschel Space Observatory is a 3.5-m telescope built and operated by the European Space Agency (ESA) covering the far-infrared and submillimetre wavelength ranges (Pilbratt et al. 2010). Comet C/2006 W3 (Christensen) (hereafter referred to as C/2006 W3) was observed with PACS and HIFI, two of the focal-plane instruments onboard Herschel, on UT 26.5 August and 1.5 September 2010, respectively, when the object was at $r_h$ ~ 5.0 AU and a distance of $\Delta$ = 4.5–4.7 AU from the spacecraft. The observations were performed within the framework of the Herschel guaranteed-time key program “Water and related chemistry in the solar system” (Hartogh et al. 2009). These data complement previous outbound observations of comet C/2006 W3 obtained four months post-perihelion with PACS during the science demonstration phase on UT 1.8 and 8.8 November 2009 and with SPIRE on UT 6.6 November 2009, that have been reported previously in Paper I.

Table 1 summarizes the observing circumstances and distances to comet C/2006 W3 during the observations obtained with the HIFI and PACS instruments. We used the ephemeris provided by JPL’s HORIZONS online solar system data service$^1$ to obtain the position and solar position angle of the comet with respect to the satellite (Giorgini et al. 1996).

The pointing information of Herschel observations acquired in the period between operational days (ODs) 320 and 761 suffered from a shift in the reconstructed astrometry of up to 8′′ due to an introduced change in the temperature of the star-tracker. This effect was modified by a new model that was uploaded to the satellite on OD 762. For the PACS and HIFI observation of C/2006 W3 obtained on ODs 469 and 475, respectively, new pointing products have been generated by reprocessing the data with the new astrometry generated with the Herschel interactive processing environment (HIPE) v10.3.0 that corrects the pointing offset problem.

$^1$ http://ssd.jpl.nasa.gov/?horizons
2.1. HIFI observations

Double-sideband (DSB) heterodyne systems such as HIFI are sensitive to two frequency ranges located on either side of the local oscillator frequency. Using HIFI, we aimed to simultaneously detect the ground-state rotational transitions \(J_{K-K} (1_{0}^1-0_{1}^0)\) of ortho-\(\text{H}_2\) and \(J_{K} (1_{0}^1-0_{1}^0)\) of ortho-\(\text{NH}_3\) at 557 GHz and 572 GHz, in the lower and upper sidebands (LSB and USB) of band 1b, with a total integration time of 48 min. The observation was performed in the single-point frequency-switching observing mode without a reference position on the sky that maximizes the on-target integration time, although it introduces strong standing waves in the baseline. These standing waves have to be removed and are a source of additional systematic uncertainties. A frequency throw of 94.5 MHz was applied in this observation. We simultaneously used the wide-band spectrometer (WBS) and the high-resolution spectrometer (HRS) with frequency resolutions of 1.1 MHz and 120 kHz, after resampling to a uniform frequency grid. Nonetheless, note that both the WBS and HRS spectra are oversampled and the spectrometers have a spacing of the frequency grid of \(~500\) kHz and 60 kHz, respectively. In the HIFI observation the \(\text{H}_2\) and \(\text{NH}_3\) emission line frequencies were targeted in “subband” 4 of the WBS, whereas the HRS has a single subband in the high-resolution observing mode.

The data reduction was carried out with the standard HIFI pipeline v10.3.0 using the HIPE software package to obtain calibrated level-2 products (Ott 2010). We corrected the temperature scale using the HIFI forward efficiency, i.e., the fraction of radiation received from the forward hemisphere of the beam as compared to the total radiation received by the antenna, of 0.96. Additionally a main beam efficiency of 0.75 in band 1b has been taken into account to correct for the fraction of power collected in the main Gaussian beam of the telescope with respect to the total power (see e.g., Roelfsema et al., 2012).

2.2. PACS observations

In photometer mode, the PACS instrument simultaneously images two of its three filters centered on 70 \(\mu\)m, 100 \(\mu\)m and 160 \(\mu\)m (referred to as red, green, and blue bands) that cover the 60–85 \(\mu\)m, 85–125 \(\mu\)m, and 125–210 \(\mu\)m ranges. Two bolometer arrays provide a field of view of 1.75 \(\times\) 3.5\(^\prime\) in each of the bands.

The PACS maps presented here were taken in the red and blue bands with two orthogonal scanning directions with respect to the detector array, using the medium slewing speed of the spacecraft along parallel lines of 20\(^\prime\)/s (as compared with the slow-scan mode of 10\(^\prime\)/s for the observation obtained during the Herschel science verification phase presented in Paper I).

The integration time was 20.6 min for each scanning direction. In the earlier PACS observations from November 2009, we used three scan legs with a 99 length and a \(\sim 2.5\) leg separation, while the more recent observations have five scan legs with a 3\(^\prime\) length and 20\(^\prime\) separation. The pixel sizes are 6\(^\prime\)\(\times\)6\(^\prime\) and 3\(^\prime\)\(\times\)3\(^\prime\) for the red and blue channels, but for maps smaller than 5\(^\prime\) in size, as in the case of the C/2006 W3 observations from August 2010, the PACS photometry pipeline resamples the images to pixel sizes of 1\(^\prime\) and 2\(^\prime\) in the blue and red bands. The PACS data from November 2009 and August 2010 were reduced with the HIPE pipeline version v10.3.0.

3. Data analysis

3.1. HIFI data analysis

The data reduction for the HIFI data follows the workflow described for example in Biver et al. (2012a,b) and de Val-Borro et al. (2012b). After processing the raw data to calibrated level-2 products with the HIFI pipeline, the frequency-switched data are folded by averaging the original spectrum with an inverted copy that is shifted by the local oscillator throw of 94.5 MHz (e.g., Liszt 1997). A strong baseline ripple is present, introduced by multiple standing waves in the frequency-switching observing mode. To obtain a reliable estimate of the noise in the emission-free part of the spectrum, the baseline has to be removed, which was accomplished by fitting a linear combination of sine waves using the Lomb-Scargle periodogram technique (Lomb 1976; Townsend 2010). The spectrum can also be processed using an empirical mode-decomposition (EMD) method (see Rezac et al., 2014; and Rezac et al., in prep., for a quantitative comparison of Lomb-Scargle and EMD techniques with other procedures). We used a masking window of \((-1, 1)\) km s\(^{-1}\) around the transition frequency in each of the orthogonal polarizations spectra to fit the standing waves in the baseline.

The procedure used to remove the residual standing waves of frequency-switched observations introduces an additional systematic error source of about 50% of the root mean square (rms) scatter of the baseline-subtracted spectrum (see e.g., Bockelée-Morvan et al., 2012; O’Rourke et al., 2013b). Moreover, an additional uncertainty, similar to that arising from the baseline removal, is related to the flux calibration error associated with the beam efficiency error, calculated from Mars mapping observations, and sideband gain ratio between the upper and lower sidebands (see Roelfsema et al., 2012). We have adopted a conservative value of 15% flux error in band 1b resulting from a combination of these uncertainties (see e.g., Higgins et al., 2014). All these systematic errors were considered in the derived line intensity and production rate uncertainties we present below.

---

Table 1. HIFI and PACS observing circumstances of comet C/2006 W3 (Christensen)

<table>
<thead>
<tr>
<th>Date(^a)</th>
<th>OD(^b)</th>
<th>Inst.</th>
<th>ObsID</th>
<th>Exp. (^{\text{c}}) (\text{min})</th>
<th>Angle(^c) (^{\circ})</th>
<th>Scan size (\text{'} \times \text{'})</th>
<th>Speed(^d) (^{/s})</th>
<th>(\eta^{e}) (\text{AU})</th>
<th>(\Delta^{f}) (\text{AU})</th>
<th>(\phi^{g}) (^{\circ})</th>
</tr>
</thead>
<tbody>
<tr>
<td>2009-11-01.836</td>
<td>171</td>
<td>PACS</td>
<td>1342 186 621(^b)</td>
<td>9.4</td>
<td>135</td>
<td>9''9 * 74'</td>
<td>10</td>
<td>3.33</td>
<td>3.56</td>
<td>16:31</td>
</tr>
<tr>
<td>2009-11-01.843</td>
<td>171</td>
<td>PACS</td>
<td>1342 186 622(^b)</td>
<td>9.4</td>
<td>45</td>
<td>9''9 * 74'</td>
<td>10</td>
<td>3.33</td>
<td>3.56</td>
<td>16:31</td>
</tr>
<tr>
<td>2010-08-26.532</td>
<td>469</td>
<td>PACS</td>
<td>1342 203 478</td>
<td>20.6</td>
<td>45</td>
<td>3''0 * 17'</td>
<td>20</td>
<td>4.96</td>
<td>4.54</td>
<td>11:22</td>
</tr>
<tr>
<td>2010-08-26.547</td>
<td>469</td>
<td>PACS</td>
<td>1342 203 479</td>
<td>20.6</td>
<td>135</td>
<td>3''0 * 17'</td>
<td>20</td>
<td>4.96</td>
<td>4.54</td>
<td>11:22</td>
</tr>
<tr>
<td>2010-09-01.529</td>
<td>469</td>
<td>HIFI</td>
<td>1342 204 014</td>
<td>48.0</td>
<td></td>
<td></td>
<td></td>
<td>5.00</td>
<td>4.68</td>
<td>11:45</td>
</tr>
</tbody>
</table>

Notes. \(^{a}\) UT mid-date of the observation with fractional days. \(^{b}\) Herschel operational day. \(^{c}\) Orientation angle of the scan map with respect to the detector array. \(^{d}\) Slowing speed of the satellite along the scan line legs. \(^{e}\) Heliocentric distance. \(^{f}\) Distance to Herschel. \(^{g}\) Solar phase angle (Sun–C/2006 W3–Earth). \(^{h}\) Data from Paper I.
Fig. 1. Folded spectra for the horizontal polarization mixer (blue line), vertical polarization mixer (red line), and averaged spectrum (black line) of the ortho-water $J_{K,K'} (1_{01} - 1_{00})$ line at 556.936 GHz acquired with the WBS on UT 1.5 September 2010. The fitted baselines are obtained by masking a $(-1, 1)$ km s$^{-1}$ window around the rest frequency of the water line and using a linear combination of sine functions. The vertical axis is the calibrated main-beam brightness temperature, the upper horizontal axis is the lower sideband frequency, and the lower horizontal axis shows the Doppler velocity with respect to the nucleus rest frame.

Figure 1 shows the folded and baseline-subtracted WBS spectra of the $J_{K,K'} (1_{01} - 1_{00})$ transition of H$_2$O measured in horizontal and vertical polarizations (hereafter H and V). We averaged the H and V polarization data together to improve the signal-to-noise ratio (S/N) of the emission feature and used the same method to remove the standing waves in the baseline (black line shown in Fig. 1). A tentative detection of the H$_2$O line is present in the folded spectra at the position of the transition rest frequency for both polarizations. However, there are similar features in the spectrum introduced by the standing waves in the baseline ripple. The line feature is present in the folded H and V polarization spectra with a similar intensity and S/N of 3.5, although not exactly at the same frequency; the H spectrum is blueshifted by about $(-240 \pm 90)$ m s$^{-1}$.

Using the same method as in the WBS spectrum to fit and subtract the baseline, an emission feature appears in the HRS data that is consistent with the line area detected by the WBS within $1\sigma$ uncertainties including the systematic component from the baseline subtraction. Figure 2 shows the reduced spectra of the water $J_{K,K'} (1_{01} - 1_{00})$ line measured by the HRS. Note that to use the spectrometer’s native frequency resolution and to increase the S/N of the spectra, we resampled over two adjacent channels with a rectangular window function, which gives an effective spectral resolution of 120 kHz. After removing the baseline by a linear combination of sinusoids, the integrated intensities of the line are (11.4 $\pm$ 3.5) mK km s$^{-1}$ and (16.2 $\pm$ 4.6) mK km s$^{-1}$ in the main-beam brightness temperature scale for WBS and HRS. The agreement in the line intensities within uncertainties is to be expected since it is the same signal coming from the mixer including the standing waves, that is multiplexed to the WBS and HRS. This also confirms that the spectral feature has not been introduced by either of the two backends.

The line centroids in the HRS data with the original frequency resolution are shifted by $(-230 \pm 120)$ m s$^{-1}$ and $(-150 \pm 66)$ m s$^{-1}$ with respect to the rest frequency of the H$_2$O transition in the H and V spectra, respectively, compared to a blueshift of $(-240 \pm 90)$ m s$^{-1}$ and $(-10 \pm 70)$ m s$^{-1}$ for the WBS. The difference between the H and V line centroids in the HRS spectra, $(-80 \pm 140)$ m s$^{-1}$, is smaller than in the estimation in the WBS data, $(-230 \pm 120)$ m s$^{-1}$, while they agree within $\sim 1\sigma$ uncertainties. Although the HRS data has higher frequency resolution than the WBS data, the S/N in the emission line in the HRS spectra is very low to allow a good estimate of the difference between the H and V line velocity offsets.

We simultaneously searched for H$_2$O and NH$_3$ emission in the HIFI observation of C/2006 W3 in the LSB and USB. The averaged spectra around the ammonia transition observed with the two HIFI backends and centered on the barycentric position of the hyperfine component frequencies are shown in Fig. 3. We removed the residual standing waves in the calibrated spectra using the same procedure as described for the H$_2$O emission line. There is no indication of an NH$_3$ line emission in either the WBS or HRS data. The derived 3$\sigma$ upper limits to the NH$_3$ line intensity including systematic uncertainties are $<13$ mK km s$^{-1}$ and $<17$ mK km s$^{-1}$ for the WBS and HRS spectra.

3.1.1. Origin of the difference in the H- and V-polarization spectra

The difference in the H$_2$O line position for the orthogonal polarizations in the HIFI WBS data can be partly explained by the pointing offset between H and V polarization spectra of $\sim$6” in band 1b. A comparison of the pointing offset of the H and V beams with the full width at half maximum (FWHM) of the averaged beams is shown in Fig. 4 ($\sim$38” at these frequencies, corresponding to an aperture of $\sim$130 000 km at a distance...
of 4.68 AU from Herschel). For Herschel observations of solar system objects, the tracking mode is activated to follow their motion on the sky. In the case of comet C/2006 W3 the HORIZONS ephemeris was used to determine the tracking coefficients for the instrument boresight that are used in the pointing commands. For HIFI’s band 1, the orientation of the polarization grids is such that the H polarization is basically parallel to the ecliptic plane and the V polarization lies perpendicular to it, in the case of sources close to the ecliptic plane. The projected trajectory of the comet nucleus on the sky during the observation passes farther away from the center of the averaged H beams (mean angular distance ~4″8) than from the center of the averaged V beams (mean angular distance ~2″3) as shown in Fig. 4.

As noted previously, a pointing correction due to a warm star-tracker during the HIFI observation was taken into account by the HIPE pipeline that resulted in an offset of \( \cos(\delta_e) \times (\delta_e - \delta_o) = 1°7 \) and \( (\alpha_e - \alpha_o) = 2°7 \) on average over the observation, where \( (\alpha_e, \delta_e) \) and \( (\alpha_o, \delta_o) \) are the original and estimated coordinates after correction, respectively. Thus, a total offset of \(~2°2\) towards the northeast as compared to the original pointing, i.e., the comet position, was introduced that explains the fact that the V polarization beam is closer to the comet nucleus as shown in Fig. 4. Since we have not obtained a mapping observation with HIFI, we are not able to recover the absolute pointing offset from the morphology of the brightness distribution in the coma. However, the 1σ pointing accuracy within the absolute pointing error (APE) is known for the epoch of the C/2006 W3 observation, and estimated to be 1°6 after the warm star tracker correction is taken into account. The absolute pointing uncertainty introduces an additional error in the production rate estimates of about 1% that has been taken into account in their absolute standard deviations.

Owing to the small solar phase angle of ~11°, the blueshift of the line in the H polarization spectrum might be caused by preferential outgassing in the direction towards the Sun within the H aperture. The smaller line shift in the V spectrum, consistent with zero offset within uncertainties, suggests a different outgassing geometry of the material comprised within the V beam. This indicates that there may be anisotropies in the coma structure in C/2006 W3 combined with possible variations in the radial and azimuthal components of the gas expansion velocity. Since the line area of the folded V spectrum is larger than the folded H spectrum and the V beam is pointing closer to the nucleus position, the difference in the line intensities in both polarizations may also arise from a higher total column density averaged within the V beam than in the H beam because of the relative pointing offset. Using the spherically symmetric outgassing model described in Sect. 3.1.2, we expect that the line intensities should decrease by ~1% and 3% at 2° and 5° offset, respectively. These values are roughly consistent with the difference in the line intensity of the folded H and V spectra, although the deviation in the negative phase of the unfolded spectra is much larger, probably because of the poor S/N.

Considering that the integrated line of the folded and averaged WBS spectrum has an S/N of ~4, it is expected that a line feature is present in both switching phases at the position of the frequency throw in the unfolded spectrum if the detected signal is real. This aspect should also be reproducible in both polarizations with an S/N between 2–3. To test this possibility, we searched for emission features in the unfolded WBS spectra for H and V polarizations in Fig. 5. We found evidence of an emission line feature in the positive and negative phases of the frequency-switched line frequency in the V polarization, but a feature is only present in the positive phase for the H polarization, although there is no physical reason to expect that the source has an intrinsic polarization so that the line is present in only one of the polarizations. Generally, the H polarization has been found to provide better stability and a lower noise level than the V polarization (e.g., Lis et al. 2013). Since the S/N in our observation is limited, it is possible that we do not see a detection in one of the phases for the H polarization although a signal may be present.
3.1.2. H$_2$O and NH$_3$ production rates

To compute the production rates we used a spherically symmetric radiative transfer model that takes into account nonthermal processes and has previously been applied to analyze _Herschel_ data and ground-based millimeter/submillimeter observations of line emission in cometary atmospheres (see e.g., Hartogh et al. 2010, 2011; de Val-Borro et al. 2010, 2012a, 2013). The excitation of water molecules was calculated with a code based on _ratran_ (Hogerheijde & van der Tak 2000; Hogerheijde et al. 2009), which includes collisions between neutrals and electrons and radiation-trapping effects. We considered the effects of pumping of the vibrational excited states by infrared solar radiation. To compute the upper limit to the NH$_3$ production rate, the collisions with neutrals and electrons were modeled following Biver et al. (2012a,b).

We adopted similar parameters as in the models used to derive production rates from C/2006 W3 observations with IRAM (Paper I). Since the electron density in the coma is poorly constrained, an electron density scaling factor of $n_e = 0.2$ with respect to the standard profile derived from observations of comet 1P/Halley was used (see e.g., Biver et al. 2007; Paper I; Hartogh et al. 2010). This value was found to explain the brightness distribution of the $J_{K_{\text{C}}}$ ($1_{01}$–$1_{01}$) water emission line obtained with mapping observations (e.g., Biver et al. 2007; Hartogh et al. 2010, 2011; Bockelée-Morvan et al. 2012). The radial number density profile for each molecule was obtained using the standard spherically symmetric Haser distribution (Haser 1957):

$$n(r) = \frac{Q}{4\pi r^2 v_{\text{exp}}} \exp \left( -\frac{r\beta}{v_{\text{exp}}} \right),$$

where $Q$ is the total production rate of the molecule in molec s$^{-1}$, $v_{\text{exp}}$ is the expansion velocity in the coma, $r$ is the nucleocentric distance, and $\beta$ denotes the photodissociation rate including dissociation and ionization effects by the radiation from the Sun. The expansion velocity is assumed to be 0.4 km s$^{-1}$ in comet C/2006 W3. This value is similar to the IRAM observations scaled to 5 AU and with the expected expansion velocity derived from the half width at half maximum (HWHM) of a Gaussian fit to the emission line observed at the Nançay radio telescope (Tseng et al. 2007).

Generally, the actual outgassing geometry in cometary atmospheres can be different from the Haser density pattern used in our model. Nonetheless, the systematic errors introduced by the assumption of a spherically symmetric distribution are on the same order as the statistical errors. Provided that the model uses the same average parameters, a spherically symmetric outflow provides a similar fit to the production rate as a nonisotropic distribution consisting of a combination of several jets with variable velocity, although the shape of the derived line profile for a nonisotropic outflow may differ substantially.

Table 2 shows the line intensity and rms noise of the brightness temperature and integrated line area for each spectrometer
that were used to derive the H$_2$O production rates and a 3σ upper limit for the NH$_3$ outgassing rate. Theoretically, the HRS and WBS noise levels should be consistent taken into account the square root of the ratio of their respective resolutions, and the HRS efficiency, $\eta = 0.81$. The data have a ratio of the HRS to WBS noise that is slightly larger than the expected theoretical value.

Since CO is the most abundant species in the atmosphere of C/2006 W3 (e.g., Otsufo et al., 2012), the main collisional excitation mechanism for H$_2$O molecules is collisions with CO molecules. The collisional excitation between H$_2$O and CO molecules was modeled using a cross section $\sigma_{CO-H_2O} = 0.5 \times \sigma_{H_2O-H_2O}$, where $\sigma_{H_2O-H_2O} = 5 \times 10^{-14}$ cm$^2$ (see Boeckele-Morvan 1987, and references therein). We also investigated a higher collision rate for H$_2$O of 1.5 $\times$ 10$^{-13}$ cm$^2$, but with a $Q_{CO}/Q_{H_2O}$ ratio of ~6 at 5 AU from the Sun, the retrieved $Q_{H_2O}$ decreases by only 5%.

Using the standard deviation of the brightness temperature in the WBS backend, we derive a 3σ upper limit for the line intensity of the NH$_3$ line of <13 mK km s$^{-1}$ over a (−1, 1) km s$^{-1}$ window, which converts into 14.8 mK km s$^{-1}$ taking into account losses of hyperfine structure of the line to get 100% of the line emission. For this value, an upper limit of the production rate, $Q_{NH_3} < 1.5 \times 10^{27}$ molec s$^{-1}$, is derived, considering collisions with CO and using a cross section $\sigma_{CO-NH_3} = 3 \times 10^{-14}$ cm$^2$ (i.e. six times larger than $\sigma_{H_2O-H_2O}$). From the 3σ upper limit for the line intensity of the HRS NH$_3$ spectrum of <17 mK km s$^{-1}$, we derive an upper limit to the production rate of <1.9 $\times$ 10$^{27}$ molec s$^{-1}$ taking into account losses of hyperfine structure. Although the collisional cross section between CO and NH$_3$ is poorly determined, the dominating factor for deriving an NH$_3$ production rate of <1.5 $\times$ 10$^{27}$ molec s$^{-1}$ is radiative self-absorption.

There is a 60% difference in the production rate calculated in local thermodynamic equilibrium (LTE), while including only collisions between water and NH$_3$, the production rate $Q_{NH_3}$ changes by 10%. Therefore, a non-LTE treatment like the escape probability method or a Monte Carlo radiative transfer approach is required to accurately compute the NH$_3$ population levels (Biver et al., 2012a).

Compared with the typically observed $Q_{NH_3}/Q_{H_2O}$ mixing ratio of about 0.5−1% in comets at a heliocentric distance of 1 AU (see e.g., Biver et al., 2012a, and references therein; Biver et al., in prep.), the derived upper limit in comet C/2006 W3 of <1.5 $\times$ 10$^{27}$ molec s$^{-1}$ is on the order of the water production rate, giving a mixing ratio $Q_{NH_3}/Q_{H_2O} < 0.75$. However, because NH$_3$ is much more volatile than water (with a sublimation temperature, $T_{sub}$, of 78 K to be compared with $T_{sub} = 152$ K for H$_2$O) we could expect an enhanced ratio $Q_{NH_3}/Q_{H_2O}$ closer to unity at 5 AU from the Sun. Nonetheless, a mixing ratio of ~1% cannot be excluded from the HIFI data.

### 3.1.3. Comparison with pre- and post-perihelion observations

The gas production rate at large heliocentric distances is controlled by sublimation of highly volatile species like carbon dioxide and carbon monoxide, which have sublimation temperatures of ~70 K. Although in most cometary atmospheres CO$_2$ has been found to dominate CO with a wide variety of mixing ratios (e.g., Otsufo et al., 2010, 2012), the activity of C/2006 W3 was mostly driven by CO emission from the measurements obtained around 3.1−3.7 AU. Scaling the CO production rate obtained from post-perihelion measurements of the CO (2−1) line at IRAM by $r_{sub}^{-1/2}$ as measured in comet C/1995 O1 (Hale-Bopp) (Biver et al. 1997); we estimate a mixing ratio $Q_{CO}/Q_{H_2O}$ of 6.5 at 5 AU. These IRAM observations were obtained almost simultaneously with the PACS and SPIRE measurements from November 2009. The CO production rate measured by IRAM is about a factor of two higher than the value derived by the AKARI infrared observations at $r_h = 3.13$ AU (Otsufo et al., 2012). This might be explained by the difference in the field of view of the observations and uncertainties in the coma distribution and excitation parameters (e.g., Otsufo et al., 2012). Additionally, infrared detections of the hypervolatiles CO, CH$_3$ and C$_2$H$_4$ were obtained in comet C/2006 W3 using the Cryogenic Infrared Echelle Spectrograph (CRIRES) on the Very Large Telescope (VLT) at heliocentric distances of 3.25 AU and 4.03 AU, as initially reported in Biver et al. (2010, 2013).

Figure 6 shows the H$_2$O, CO$_2$ and CO production rates as a function of heliocentric distance as measured by various instruments. The water production rate measured by HIFI and the production rate at the subsolar point on the nucleus shown by the solid curve in Fig. 6 agree but this similitude does not require water emission to come entirely from the nucleus. The solid curve represents the maximum value of water production rate because it is assumed that there is no dust on the surface and the thermal inertia of the surface ice is very low. The absorbed solar energy goes into sublimation, thermal re-radiation, and diffusion of heat into the interior of the nucleus in our model.

However, this simple model does not include release of trapped CO molecules during crystallization of amorphous water
3.2. PACS data analysis

We analyzed the PACS photometric observations in the same way as described in Paper I. The proper motion of the comet was relatively fast at the time of the observations, of about 10/′h, and the total integration time was ∼10 min to 20 min for each scanning direction in the PACS observations. Because of this, the Astronomical Observation Requests (AORs) for these observations identified comet C/2006 W3 as a moving solar system object and the nonsidereal tracking was activated by the mission planning system that lead to a continuous tracking of the motion of the comet. The HIPE pipeline that has been used to reduce the PACS observations with this special tracking mode brings all the individual frames together for each individual scan, correcting for the observed proper motion in this observing mode. For each scanning direction, individual maps were projected onto the same World Coordinate System (WCS) centered on the comet nucleus by shifting the data arrays onto a uniform grid with multivariate interpolation using a fixed Gaussian kernel. We subtracted the background emission and averaged the maps with orthogonal scanning direction, correcting for the proper motion of the target during the observations for each of the two epochs (see e.g., Santos-Sanz et al. 2012) using the ephemeris computed by the HORIZONS system (Giorgini et al. 1996). The position of the nucleus at mid-time of the observations was used to center each of the images.

Figure 7 shows the blue and red maps obtained with PACS in November 2009 – presented in Paper I – and August 2010, centered on the peak position of the images after merging the observation pairs with orthogonal scan directions for each channel. The nucleus position at mid-time of the observations is shown by the yellow circles. PACS photometer observations obtained in the scan mapping mode can be affected by offsets due to the 1σ APE of the telescope of about 1′′ to 2′′ at the time the C/2006 W3 observations were acquired (see Pilbratt et al. 2010; Valtchanov et al. 2013, and references therein). This value can explain the offset between the nucleus position and the estimated position of the peak shown in Fig. 7. On the other hand, the nominal relative pointing error (RPE) measured once the telescope is commanded to the target position using short observing intervals of ∼1 min during the performance verification (PV) phase has been found to lie in the range 0′′2–0′′3 for solar system objects within 68% confidence level. This error can generally be neglected compared with the FWHM beam sizes of 5′′ and 11′′ for the 70 μm and 160 μm PACS scan-mapping observations.

3.2.1. Dust coma distribution and radial profiles

The surface brightness of the 70μm image has a broader distribution than the 160 μm image in both epochs, although the difference is significantly smaller in the August 2010 observations (see lower panel in Fig. 7). The coma distribution clearly broadens in the direction towards the Sun in the first observations, possibly due to outflow of large dust grains. In the more recent PACS maps, the red band image is elongated towards the southwest close to the direction of the velocity vector of the comet, and the blue image is roughly symmetric. Note that the solar phase angle was 11° at the time of the latter observations, so the observing geometry is not the most appropriate to see an extended dust tail towards the antisolar direction. Arc-like asymmetries and spirals are observed in the 4.5 μm Spitzer images of C/2006 W3 when the comet was at 3.1 AU from the Sun. However, these can be explained in part by gas structures (Reach et al. 2013).
Fig. 7. Maps at 70 µm (blue band; left column) and 160 µm (red band; right column) of comet C/2006 W3 (Christensen) observed with PACS on UT 1.5 November 2009 (upper row, see Paper I) and UT 26.5 August 2010 (bottom row) centered on the peak position indicated by the white cross. The pixel sizes are 3.2′′ in the blue map and 6.4′′ in the red map for the November 2009 observations, and 1′′ in the blue map and 2′′ in the red map for the August 2010 data. Contour levels have a step of 0.1 on logarithmic scale in all the images. The position of the nucleus at the mid-time of the observations according to the HORIZONS ephemeris is shown by the yellow circle. The projected directions toward the Sun and velocity vector of the comet are indicated by the solid and dashed arrows.

Fig. 8. Left panel: surface brightness in comet C/2006 W3 (Christensen) as a function of distance from the position of the center of the peak measured with PACS on UT 1.8 November 2009 (shown by the blue and red dots for the 70 µm and 160 µm bands with a pixel size of 3′′2 and 6′′4). Right panel: surface brightness radial profiles observed on UT 26.5 August 2010 (blue and red dots denote the 70 µm and 160 µm bands with projected pixel size of 1′′ and 2′′). The green crosses and circles are resampled profiles with 1′′ and 2′′ bins for the blue and red band images from November 2009, and 0′′5 and 1′′ bins for the blue and red band images from August 2010. Error bars are 1σ statistical uncertainties.

The radial profiles of surface brightness with respect to the peak position in the blue and red bands are shown in Fig. 8. We fitted the resampled radial profiles by convolving the function $\rho^{-3}$ with a fit to the point spread functions (PSFs) of Vesta consisting of two Gaussian functions. Figure 9 shows the radial profiles of surface brightness in the reference PSFs from
We studied the gas and dust activity in the long-period comet C/2006 W3 (Christensen) using remote-sensing observations obtained with the Herschel Space Observatory at 5.0 AU from the Sun. A tentative detection of the ortho-H$_2$O ground-state transition is observed in the HIFI spectra with $4\sigma$ significance. Even though the H$_2$O line is only marginally detected, the derived H$_2$O production rate and line shape and velocity shift are consistent with the emission feature having a cometary origin. We derive a water production rate of $(2.0 \pm 0.5) \times 10^{27}$ molecules s$^{-1}$ using a spherically symmetric radiative transfer model. Concurrently, we aimed to detect the ground-state rotational transitions of ortho-NH$_3$ in the USB of HIFI’s band 1b, from which a $3\sigma$ upper limit for the ammonia production rate of $<1.5 \times 10^{27}$ molecules s$^{-1}$ is derived, corresponding to a mixing ratio $Q_{NH_3}/Q_{H_2O} < 0.75$.

Water sublimes fully at heliocentric distances <2.5 AU and is responsible for the activity of cometary nuclei at these distances.

### 4. Discussion

We studied the gas and dust activity in the long-period comet C/2006 W3 (Christensen) using remote-sensing observations obtained with the Herschel Space Observatory at 5.0 AU from the Sun. A tentative detection of the ortho-H$_2$O ground-state transition is observed in the HIFI spectra with $4\sigma$ significance. Even though the H$_2$O line is only marginally detected, the derived H$_2$O production rate and line shape and velocity shift are consistent with the emission feature having a cometary origin. We derive a water production rate of $(2.0 \pm 0.5) \times 10^{27}$ molecules s$^{-1}$ using a spherically symmetric radiative transfer model. Concurrently, we aimed to detect the ground-state rotational transitions of ortho-NH$_3$ in the USB of HIFI’s band 1b, from which a $3\sigma$ upper limit for the ammonia production rate of $<1.5 \times 10^{27}$ molecules s$^{-1}$ is derived, corresponding to a mixing ratio $Q_{NH_3}/Q_{H_2O} < 0.75$.
distances form the Sun, while its sublimation becomes inefficient at larger distances where the nucleus surface temperature may be lower than the H₂O sublimation temperature (Crovisier & Encrenaz 2000). Thus, outgassing activity at heliocentric distances >3 AU is mostly driven by molecular species more volatile than water such as carbon monoxide. To date, the largest heliocentric distance at which OH emission – which is regarded as a proxy for H₂O – has been detected is 4.7 AU, in pre-perihelion observations of the exceptionally active comet C/1995 O1 (Hale-Bopp) with the Nançay radio telescope (Colom et al. 1999).

Whether the observed water emission in C/2006 W3 originates directly in the nucleus or it is produced by icy grains in the coma can be partly assessed from the shape of the emission line. The characteristic blueshift of the water line detected by the H polarization WBS spectrum is indicative of preferential emission from the daytime hemisphere in the direction towards the Sun, taking into account the small solar phase angle of ~11° at the time of the observations. The smaller blueshift of the line in the V polarization WBS spectrum suggests that the material in the coma has a nonspherical outgassing geometry with possibly variable gas velocity. However, from the HIFI observations we cannot exclude the possibility that water sublimation occurs in small icy grains outflowing at a speed close to 0.1 km s⁻¹ from the nucleus surface.

Production rates of several species measured with ground-based and space observatories at different heliocentric distances were compared with the H₂O production rate derived from the Herschel observations. The gas production rate in C/2006 W3 is driven by sublimation of hypervolatile carbon monoxide molecules, while CO₂ has been found to dominate CO with a wide variety of mixing ratios in most comets (e.g., Ootsubo et al. 2010, 2012, A'Hearn et al. 2011). Since CO is more volatile than CO₂, CO may be depleted in the surface layers of comets after multiple passages through the inner solar system, and the CO enrichment measured in C/2006 W3 may indicate that of the primordial material of the early solar nebula (Ootsubo et al. 2012). We find that the evolution of the water production rate is consistent with an estimate of the water vapor outgassing from the subsolar point. However, more sophisticated thermal models considering the chemical differentiation of the nucleus and emission from icy grains that were ejected from a localized region

Table 4. Dust model parameters and dust production rates of comet C/2006 W3 (Christensen) derived from PACS 70 μm and 160 μm data for carbon and olivine grains.

<table>
<thead>
<tr>
<th>Date</th>
<th>rₜ (AU)</th>
<th>Qₓ CO (molec s⁻¹)</th>
<th>dₘax (mm)</th>
<th>vₜₜₜₜ (m s⁻¹)</th>
<th>αₛ_min (m s⁻¹)</th>
<th>αₛ_max (m s⁻¹)</th>
<th>α</th>
<th>Qᵧ CO 70 (μm) (kg s⁻¹)</th>
<th>Qᵧ CO 160 (μm) (kg s⁻¹)</th>
<th>Qᵧ olivine 70 (μm) (kg s⁻¹)</th>
<th>Qᵧ olivine 160 (μm) (kg s⁻¹)</th>
</tr>
</thead>
<tbody>
<tr>
<td>2009-11-08</td>
<td>3.35</td>
<td>3 × 10²⁸</td>
<td>0.59c</td>
<td>5.7</td>
<td>174</td>
<td>−3</td>
<td>1520</td>
<td>1092</td>
<td>923</td>
<td>1190</td>
<td></td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>−3</td>
<td>923</td>
<td>884</td>
<td>676</td>
<td>965</td>
<td></td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>−4</td>
<td>650</td>
<td>900</td>
<td>606</td>
<td>1220</td>
<td></td>
</tr>
<tr>
<td>2010-08-26</td>
<td>4.96</td>
<td>4 × 10²⁷</td>
<td>0.07</td>
<td>6.3</td>
<td>106d</td>
<td>−3</td>
<td>89</td>
<td>56</td>
<td>63</td>
<td>76</td>
<td></td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>−4</td>
<td>82</td>
<td>68</td>
<td>70</td>
<td>102</td>
<td></td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>−4</td>
<td>98</td>
<td>112</td>
<td>96</td>
<td>182</td>
<td></td>
</tr>
<tr>
<td>2010-08-26</td>
<td>4.96</td>
<td>8 × 10²⁷</td>
<td>0.14</td>
<td>6.3</td>
<td>131d</td>
<td>−3</td>
<td>172</td>
<td>102</td>
<td>113</td>
<td>109</td>
<td></td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>−4</td>
<td>139</td>
<td>104</td>
<td>109</td>
<td>130</td>
<td></td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>−4</td>
<td>140</td>
<td>153</td>
<td>133</td>
<td>221</td>
<td></td>
</tr>
</tbody>
</table>

Notes: (a) Dust velocities for sizes αₛ_min and α₄_max, with α₄_min = 0.1 μm. (b) Dust production rate derived assuming α₄_min = 0.1 μm. (c) In Paper I, the maximum size computed by our dust model was multiplied by 1.5 to better match state-of-the-art hydrodynamic simulations (see Crifo et al. 2005, and references therein). (d) The velocity of 1-μm sized particles is 45 m s⁻¹ and 60 m s⁻¹ for Qₓ CO = 4 × 10²⁷ molec s⁻¹ and 8 × 10²⁷ molec s⁻¹, respectively.

---

**Fig. 9.** Radial PSF profiles from PACS observations of the main-belt asteroid Vesta in the blue (70 μm; left panel) and red (160 μm; right panel) channels. The green crosses are the resampled data using a 1° bin with 1σ statistical uncertainties. Solid lines are fits to the resampled PSF profiles using a sum of two Gaussian functions.
may lead to slightly different profiles and smaller active area (Kossacki & Sztutowicz, in prep.).

The thermal emission from the dust in the coma is detected by PACS photometric observations in the blue and red channels. The asymmetry in the red image indicates an anisotropic emission from the subsolarsector on the nucleus, consistent with the H$_2$O line blueshift. Dust production rates were computed for low (4 × 10$^{27}$ molec s$^{-1}$) and high (8 × 10$^{27}$ molec s$^{-1}$) CO production rates and for the case of amorphous carbon and olivine grains using the dust model described in Paper I. By comparing the flux density at the optocenter in the blue and red bands with the dust thermal emission in the model, we derived the dust production rates. For a dust size distribution index that explains the fluxes at the photocenters of the 70 μm and 160 μm PACS images, the dust production rate is in the range of 70 kg s$^{-1}$ to 110 kg s$^{-1}$ at the time of the PACS observations in August 2010, for the low and high outgassing cases. These values of the production rate correspond to a dust-to-gas production rate ratio of 0.3–0.4.

The dust production rates derived in the August 2010 observations are roughly one order of magnitude lower than those in September 2009 from the PACS observations (see Table 4 and Paper I), while the CO production rate similarly varies in the September 2009 from the PACS observations (see Table 4 and Paper I). This indicates that the dust-to-gas production rate ratio remained approximately constant during the time when the gas activity of the comet became increasingly dominated by CO outgassing. We find that the PACS data of August 2010 is best fitted by a size index between −4 and −3, consistent with the value derived from the August 2009 PACS data (Paper I).

5. Conclusions

Combining observations of rotational emission lines and dust thermal emission provide valuable constraints on the properties of the gaseous and dust activity in comets. The Herschel Space Observatory was a unique facility for obtaining sensitive observations of water emission in comets at large heliocentric distances (e.g., Bockelée-Morvan et al. 2010a; O’Rourke et al. 2013a), and imaging the thermal dust coma in two channels simultaneously at far infrared wavelengths (see e.g., Paper I; Meech et al. 2011).

We presented HIFI and PACS observations of comet C/2006 W3 that will be an additional constraint in understanding the physical processes responsible for the distant activity in comets. These observations will complement previous measurements of C/2006 W3 obtained with other facilities at smaller heliocentric distances (see e.g., Paper I; Ootsu et al. 2012; McKay et al. 2012; Reach et al. 2013; Bonev et al. 2013) that confirm that the activity of this object is dominated by sublimation of molecular species more volatile than H$_2$O. These data will also add another set of data points to the available observations of the distant activity of other comets such as 29P/Schwassmann-Wachmann 1 (Bockelée-Morvan et al. 2010a) and C/1995 O1 (Hale-Bopp) (e.g. Biver et al. 2002a; Rauer et al. 2003).

Acknowledgements. HIFI has been designed and built by a consortium of institutes and university departments from across Europe, Canada, and the United States under the leadership of SRON, Netherlands Institute for Space Research, Groningen, The Netherlands, and with major contributions from Germany, France, and the US. Consortium members are: Canada: CSA, UWaterloo; France: CESR, LAB, LERMA, IRAM; Germany: KOSMA, MPIfR, MPS; Ireland, NUI Maynooth; Italy: ASI, IFSI-INAF, Osservatorio Astrofisico di Arcetri-INAF; Netherlands: SRON, TUD; Poland: CAMK, CBK; Spain: Observatorio Astronómico Nacional (IGN), Centro de Astrobiología (CSIC-INTA). Sweden: Chalmers University of Technology – MCZ, RSS & GARD; Onsala Space Observatory; Swedish National Space Board, Stockholm University – Stockholm Observatory, Switzerland; ETH Zurich, FHNW, USA: Caltech, JPL, NHSC. PACS has been developed by a consortium of institutes led by MPE (Germany) and including UVIE (Austria); KU Leuven, CSL, IMEC (Belgium); CEA, LAM (France); MPIA (Germany); INAF-IFSI/OA/OAP/OAT, LENS, SISSA (Italy); IAC (Spain). This development has been supported by the funding agencies BMVIT (Austria), ESA-PRODEX (Belgium), CEA/CNES (France), DLR (Germany), ASI/INAF (Italy), and CICYT/MCYT (Spain). Support for this work was provided by NASA through an award issued by JPL/Caltech. We thank M. Sánchez-Portal from the Herschel Science Center, and H. Lutz from the PACS Instrument Control Center (ICC) for their help with the reduction and analysis of the Herschel data. The anonymous referee is thanked for providing constructive comments and help in improving the contents of the paper.

References

Biver, N., Bockelée-Morvan, D., Colom, P., et al. 2002a, Earth Moon Planets, 90, 1058
Bockelée-Morvan, D. 2011, in IAU Symp., 280, 261
Bockelée-Morvan, D., Biver, N., Crovisier, J., et al. 2010a, in BAAS, 42, 946
Bonev, N., Mumma, M. J., Villanueva, G. L., et al. 2013, in BAAS, 45, 413.18