

STRONG NEBULAR LINE RATIOS IN THE SPECTRA OF $Z \sim 2-3$ STAR-FORMING GALAXIES: FIRST RESULTS FROM KBSS-MOSFIRE¹

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ABSTRACT

We present initial results from a large near-IR spectroscopic survey covering the 15 fields of the Keck Baryonic Structure Survey (KBSS) using the recently-commissioned near-IR multi-object spectrometer MOSFIRE on the W.M. Keck Observatory Keck 1 telescope. Within the current KBSS-MOSFIRE sample, we focus on a representative sub-sample of 179 galaxies with redshifts $2.0 \leq z \leq 2.6$, star-formation rates (SFRs) $5 \lesssim \text{SFR} \lesssim 150 M_{\odot} \text{ yr}^{-1}$, and stellar masses $8.8 \lesssim \log(M_{*}/M_{\odot}) \lesssim 11.5$. Most of the sample has high-quality MOSFIRE spectra in both H and K-band atmospheric windows, allowing sensitive measurements of [O III] $\lambda\lambda 4960, 5008$, H β , [N II] $\lambda 6585$, and H α emission lines. We show unambiguously that the locus of $z \simeq 2.3$ galaxies in the “BPT” nebular diagnostic diagram exhibits an almost entirely disjoint, yet similarly tight, relationship between the line ratios [NII]/H α and [OIII]/H β as compared to local galaxies. Using photoionization models, we argue that the offset of the $z \sim 2.3$ BPT locus relative to that at $z \sim 0$ is caused by both higher ionization parameter *and* harder stellar ionizing radiation field than applies to most local galaxies. Because of the higher overall excitation at $z \sim 2.3$, the position of a galaxy along the star-forming BPT locus is surprisingly *insensitive* to gas-phase oxygen abundance. The locus of [OIII]/H β vs. [NII]/H α is most easily reproduced by models in which the effective ionizing radiation field is characterized by $T_{\text{eff}} = 50000-60000$ K, the gas-phase oxygen abundances lie in the range $0.2 < Z/Z_{\odot} \leq 1.0$, and the ratio of gas-phase N to O is close to the solar value. Such high sustained T_{eff} is not easily produced by standard population synthesis models, but is an expected consequence if massive binaries and/or rapid stellar rotation are important for the evolution of main sequence O-stars; the same phenomena may also account for enhanced production of primary N (and thus high N/O). We assess the applicability of commonly-used strong line indices for estimating the gas-phase metallicity of high redshift galaxies, as well as their likely systematic biases. The empirical correlation between M_{*} and strong-line-inferred oxygen abundance (the “MZR”) at $z \sim 2.3$ is as tight as for local galaxy samples (with intrinsic metallicity scatter of only 0.09 dex at a given M_{*}), but is offset to lower oxygen abundance by 0.34-0.38 dex, apparently independent of M_{*} , compared to a $z \simeq 0$ MZR using the same strong-line metallicity indicator. Over the mass range well-covered by the current sample ($9 \lesssim \log(M_{*}/M_{\odot}) \lesssim 11$), the MZR slope is shallow, with $12+\log(\text{O}/\text{H}) \propto 0.12 \log(M_{*}/M_{\odot})$; a similarly shallow slope is observed for the analogous $z \simeq 0$ MZR when evaluated over the same range in M_{*} . We investigated whether including SFR as an additional parameter in the MZR lowers the scatter, as found for local star-forming galaxy samples (the “fundamental metallicity relation”), finding that the dependence of inferred gas-phase metallicity on SFR at a given M_{*} is much weaker at high redshift than at $z \sim 0$, indicating that $z \sim 2.3$ galaxies do not adhere to the same “fundamental metallicity relation” as star-forming galaxies at low redshift. Taken together, our results suggest that the small observed scatter between M_{*} and inferred metallicity (the MZR) may be a by-product of a more fundamental relationship between a galaxy’s stellar mass and the ionization/excitation of its H II regions.

Subject headings: cosmology: observations — galaxies: evolution — galaxies: high-redshift

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1. INTRODUCTION

In principle, deep near-IR spectroscopy of high- z galaxies offers the possibility of applying the wealth of locally-calibrated and tested rest-frame optical nebular emission line diagnostics to directly probe H II region physics in galaxies as they were forming. In practice, however, this potentially-powerful method—building on well-established techniques developed over the course of several decades for nearby galaxies—has been relatively slow to develop. In spite of substantial observational effort (e.g., Pettini et al. 2001; Shapley et al. 2004; Erb et al. 2006c; Förster Schreiber et al. 2009; Maiolino et al. 2008; Mannucci et al. 2010), samples of high redshift galaxies for which a reasonably complete set of strong lines has been measured remain very small. Moreover, except for gravitationally-lensed examples (e.g., Teplitz et al. 2000; Hainline et al. 2009; Richard et al. 2011; Rigby et al. 2011; Wuyts et al. 2012; Christensen et al. 2012), the low S/N of the near-IR spectra has limited both the dynamic range and the significance of observed line ratios for individual objects. The advent of efficient multi-object near-IR spectrographs on 8-10m class telescopes has long promised to revolutionize nebular spectroscopy of high-redshift galaxies by vastly enlarging the sample sizes and making very deep spectroscopy observationally practical.

The suite of nebular emission lines available in the rest-frame optical (i.e., $0.3 \lesssim \lambda \lesssim 1 \mu\text{m}$) includes probes of density ([OII] $\lambda\lambda$ 3727,3729 and [SII] $\lambda\lambda$ 6718,6732), electron temperature ([O III] $\lambda\lambda$ 4960,5008/[O III] λ 4364) and ionization state (e.g., [O III] $\lambda\lambda$ 4960,5008/[O II] $\lambda\lambda$ 3727,3729) as well as the so-called “strong-line” metallicity indicators, e.g. those based on ([OIII]+[OII])/H β (“R23”; Pagel et al. 1979; Kewley & Dopita 2002), [NII] λ 6585/H α (“N2”) and ([OIII] λ 5008/H β)/([NII] λ 6585/H α) (“O3N2”; Pettini & Pagel (2004) [PP04]). In addition, the Baldwin et al. (1981) (see also Veilleux & Osterbrock 1987) diagnostic line ratios (“BPT”): [N II]/H α and [O III]/H β are commonly used to establish the dominant excitation mechanism of nebular emission in galaxies, providing a relatively “clean” separation of galaxies whose spectra are dominated by AGN-ionized gas from those ionized primarily by the UV radiation field of young stars (e.g., Kewley et al. 2001; Kauffmann et al. 2003; Brinchmann et al. 2008). Using large samples, primarily drawn from the SDSS spectroscopic database, it has been shown that star-forming galaxies occupy a relatively tight locus in the BPT plane. As the earliest samples of high-redshift galaxies with the relevant measurements became available, however, there were already indications that distant star-forming galaxies occupy a region of the BPT plane distinct from that of the vast majority of star-forming galaxies in the local universe (Shapley et al. 2005a; Erb et al. 2006a; Liu et al. 2008; Brinchmann et al. 2008). If the initial observations were to hold up when confronted with much larger samples, it would suggest that using nebular line ratios to measure metallicity and other physical properties of the high- z H II regions may be more complex than might have been hoped.

It is well-known that various nebular diagnostics using strong optical emission lines in galaxy spectra can differ substantially—the most obvious example is systematic differences of up to ~ 0.77 dex in oxygen abundance for ostensibly the same set of low-redshift galaxies (see e.g. Kewley & Ellison 2008; Maiolino et al. 2008.) The very different abundance scales depend to a large extent on whether the calibration has been done using theoretical models (which tend to infer

higher O abundances) or empirically, using sensitive observations of weak electron temperature sensitive emission lines—the so-called “direct”, or “ T_e ” method. The direct method is generally considered to provide more reliable results when available, but has the practical disadvantage that it requires the detection of very weak emission lines, already challenging for nearby galaxies, becoming rapidly more difficult with increasing redshift as the lines become apparently fainter and are redshifted into spectral regions plagued by much higher terrestrial background. It has also been argued that T_e -based metallicities may be biased low due to temperature gradients and/or by the details of the electron energy distribution (e.g., Stasińska 2005; Dopita et al. 2013).

To place the situation for the determination of nebular oxygen abundances in context, at the highest stellar masses, the asymptotic (i.e., maximum) gas-phase metallicity of star-forming galaxies ranges from below solar¹¹ to nearly 3 times solar (see Kewley & Ellison 2008). Given the problematic differences in metallicity scale among the many locally-calibrated and/or theoretically derived “strong-line” indicators, attempts have been made to implement new calibrations for which all of the various strong-line methods yield consistent metallicities when applied to large samples of local galaxies (Kewley & Ellison 2008; Maiolino et al. 2008). The results have been successful, in the sense that it is possible to force the calibrations to give the same results (to within $\simeq 0.03$ dex) for the same sample of galaxies (Kewley & Ellison 2008), thus providing confidence that one can at least measure *relative* oxygen abundances at $z \simeq 0$. However, even putting aside our ignorance of the “correct” H II region abundance scales at $z \simeq 0$, it is a separate issue as to whether the “renormalization” of the strong-line techniques can (or should) be applied to samples of high redshift galaxies—clearly this is a desirable possibility, but it has not yet been demonstrated. The root of the problem, which is the main topic of this paper, is that measuring line ratios and then applying regression formulae established at $z \simeq 0$ will work only if the physics of high- z H II regions resembles that of local star-forming galaxies. If there are substantive physical differences, blind application of local calibrations will introduce systematics in inferred metallicity; the origins of any systematics are likely to be fundamental to understanding what drives star formation in rapidly-evolving galaxies at high redshifts.

At most redshifts $z > 1$, only a subset of optical emission lines used by the so-called “strong-line” techniques in the local universe are accessible to ground-based spectroscopy due to significant gaps in the near-IR atmospheric transmission as well as the increasingly prohibitive thermal background at observed wavelengths of $\gtrsim 2.3\text{--}2.4 \mu\text{m}$. Potentially most-problematic is comparison of metallicities inferred from one set of strong-line indicators for a sample in a particular redshift range, with those based on a different set of lines at a second redshift. In such a case, it would be impossible to distinguish between evolution of gas-phase metallicities and changes (for other physical reasons) in the dependence of the measured line intensity ratios on metallicity.

A better statistical lever-arm, initially independent of the low-redshift calibrations, can be constructed using observations of high redshift galaxies selected in special redshift intervals for which a relatively complete set of the rest-optical nebular lines falls fortuitously within the near-IR atmospheric

¹¹ We adopt the solar metallicity scale of Asplund et al. (2009) throughout this paper, for which solar abundance corresponds to $12+\log(\text{O}/\text{H})=8.69$.

windows for ground-based spectroscopy. Perhaps the best such interval is $2.0 \lesssim z \lesssim 2.6$ (e.g., Erb et al. 2006c), where $H\alpha$, [NII], and [SII] fall in the K band, [OIII] and $H\beta$ fall in H, and [OII] and [NeIII] in J. In large part for this reason, the Keck Baryonic Structure Survey (KBSS; see Rudie et al. 2012; Trainor & Steidel 2012; Rakic et al. 2012) has focused on galaxies in this redshift range over the past several years (e.g., Adelberger et al. 2004; Steidel et al. 2004; Erb et al. 2006a,c,b; Shapley et al. 2005b; Reddy et al. 2008). Figure 1 shows the current KBSS spectroscopic redshift distribution and schematically illustrates the high priority redshift ranges targeted by the current work.

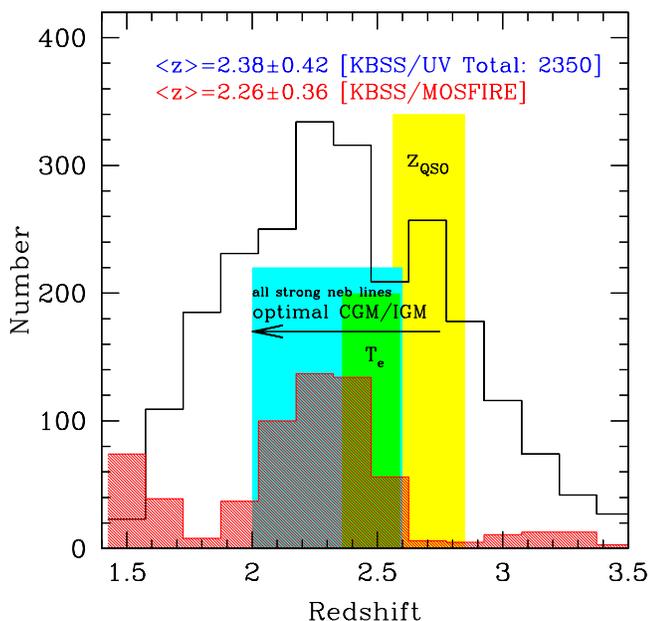


Figure 1. Redshift histogram in the KBSS survey regions as of 2013 September. The unshaded histogram shows the redshift distribution of 2350 spectroscopically confirmed galaxies based on rest-frame UV spectroscopy with Keck/LRIS-B, while the (red) shaded histogram shows the distribution of nebular redshifts obtained with MOSFIRE. The shaded regions schematically illustrate the redshift range emphasized for KBSS-MOSFIRE: the cyan shaded region corresponds to the redshift range over which the targeted suite of strong nebular emission lines fall within the ground-based near-IR atmospheric windows ($2 \lesssim z \lesssim 2.6$), the green shading (labeled “ T_e ”) corresponds to the subset of the cyan region over which the electron temperature sensitive [OIII] λ 4364 line is accessible in H band ($2.36 \leq z \leq 2.57$; see section 2). The yellow shading shows the redshift range of the very bright background QSOs in the KBSS fields; while these are not directly relevant to the topic of this paper, their lines of sight provide extremely sensitive measurements of H I and metals in the circum-galactic (CGM) and intergalactic medium (IGM) surrounding KBSS survey galaxies (see Rudie et al. 2012; Rakic et al. 2012; Trainor & Steidel 2012; Turner et al. 2014). The redshift range over which the background QSOs provide the most complete IGM/CGM measurements is $2 \lesssim z \lesssim z_{\text{QSO}}$, indicated with the arrow.

KBSS provides a wealth of multi-wavelength ancillary data as well as a large sample of spectroscopically-identified galaxies (primarily using Keck/LRIS-B) with $1.5 \lesssim z \lesssim 3.5$. In this paper, we present initial results based on new multiplexed near-IR (rest-frame optical) spectroscopy obtained in the KBSS survey regions, focusing on $2.0 \leq z \leq 2.6$, for the reasons outlined above and summarized in Figure 1.

The paper is organized as follows: section 2 describes the new observations and the properties of the initial KBSS-MOSFIRE sample; section 3 compares the locus of relative

emission line intensities of $z \sim 2.3$ galaxies with samples of galaxies in the local universe, showing very distinct differences between the two. Section 4 attempts to explain the principal cause of the change in the diagnostics, with the aid of photoionization models. Section 5 briefly examines the extent to which the observed strong emission line ratios (the “BPT” diagram) can be used at high redshift to discriminate between hot young stars and AGN as ionizing sources; Section 6 identifies likely local analogs of the high-redshift sources and compares them to the most extreme galaxies in the high redshift sample, as a means of forecasting what more sensitive observations might yield. Section 7 revisits the relationship between stellar mass and inferred metallicity (the “Mass-Metallicity” relation, or “MZR”) at $z \sim 2.3$, and briefly addresses the extent to which the new KBSS-MOSFIRE sample supports the concept of a “fundamental metallicity relation” similar to that observed in the local universe. Finally, section 8 summarizes the main results, and discusses their implications for metal enrichment and star formation in galaxies near the peak of the galaxy formation epoch.

2. OBSERVATIONS AND DATA

All near-IR spectroscopic observations described in this paper were obtained using the Multi-Object Spectrometer for InfraRed Exploration (MOSFIRE; McLean et al. 2010, 2012), the recently commissioned near-IR imaging spectrometer on the Keck 1 10m telescope at the W.M. Keck Observatory on Mauna Kea. Some of the data were obtained during MOSFIRE commissioning science verification in 2012 May and June, with the remainder obtained during early science observations in 2012 September and October and 2013 March, May, and June.

2.1. Target Selection and Survey Strategy

Over the course of MOSFIRE commissioning and early science observations, we developed an observing strategy that takes advantage of the unique capabilities of the instrument in order to achieve multiple scientific goals. The combination of the compact KBSS field geometry (typically $7'.5$ by $5'.5$) with the flexibility of MOSFIRE’s electronically re-configurable cryogenic focal plane mask (the “Configurable Slit Unit”, or CSU) lends itself to a “tiered” approach to the near-IR survey, combining routine and difficult observations on the same masks. Because most of the KBSS fields are only slightly larger than the $6'.1 \times 6'.1$ MOSFIRE field of view, there is significant spatial overlap of every mask within a given field. By repeatedly observing masks with similar footprint but distinct sets of objects, we ensure that all high priority targets are observed and that the geometrical constraints imposed by slit-masks do not limit the sampling of targets on small angular scales. At the same time, if very deep spectra are required to detect weak emission lines (e.g., auroral [OIII] λ 4364, or [NII] λ 6585 in galaxies with very low metallicity), the same target is repeated on multiple masks, thereby accumulating much longer total integration times (more than 10 hours in some cases).

The objects in the parent catalog for each KBSS field were assigned numerical priorities based on multiple criteria: the highest priorities were given to galaxies known from previous spectroscopic observations to lie in narrow redshift range $2.36 \leq z \leq 2.57$ – the range over which the set of strong emission lines (including [S II] λ 6718, 6732) as well as [OIII] λ 4364 are all accessible within the near-IR atmospheric windows (Figure 1). These would be the initial candidates

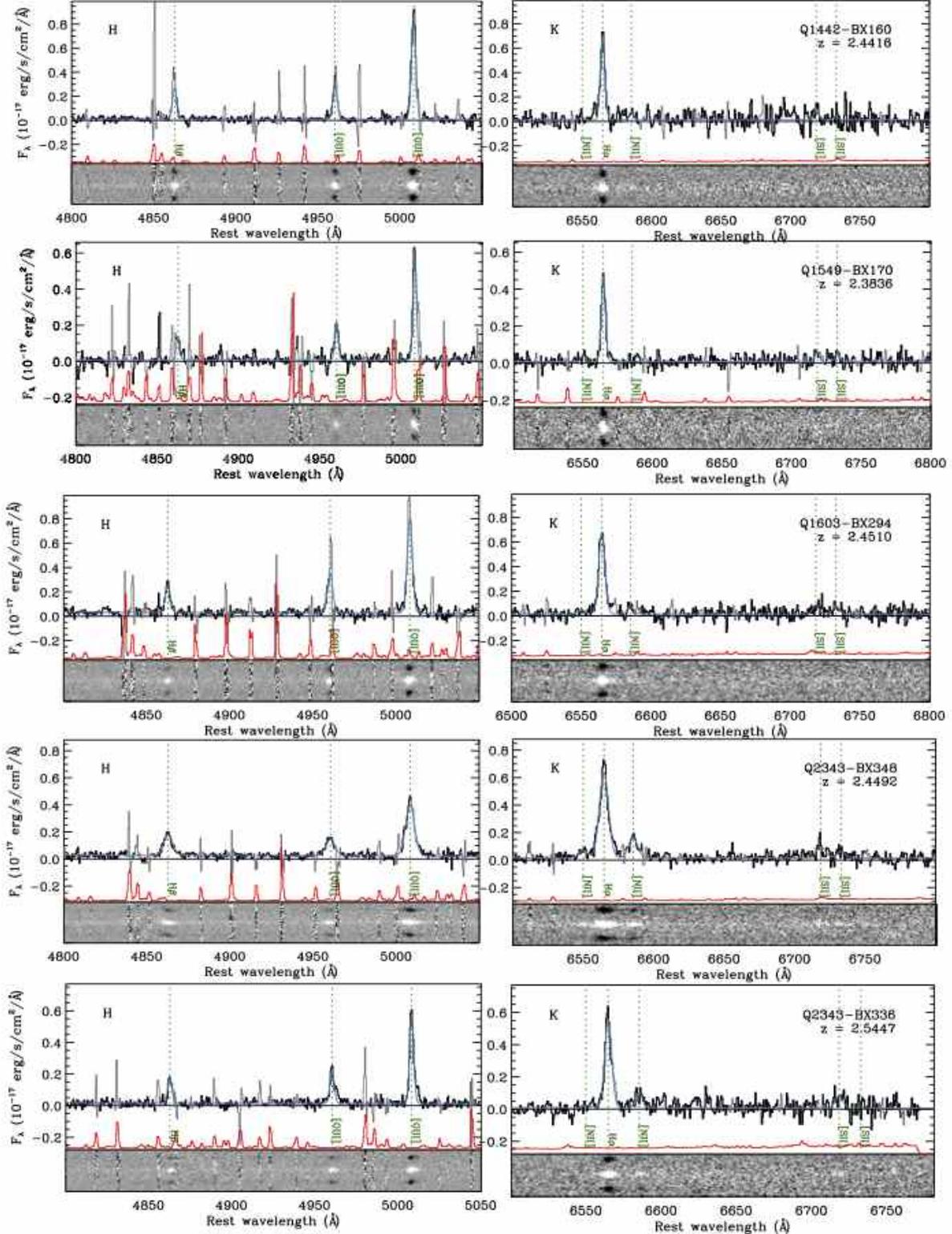


Figure 2. Portions of MOSFIRE H-band (left) and K-band (right) spectra for 10 of the KBSS galaxies listed in Tables 1 and 2. The flux-calibrated spectra are presented *unsmoothed*, with their original pixel sampling, with the wavelength scale shifted to each galaxy’s rest frame. The best-fit line profiles are superposed (blue), while the 1σ error spectrum (red) is offset, but on the same flux scale, as its corresponding galaxy spectrum. The stacked two-dimensional spectra from which the 1-d spectra were extracted are shown in grayscale, over the same range of rest-wavelength. Each reduced 2-D spectrogram exhibits a positive (central) image and 2 flanking negative images due to the differencing of spatially dithered exposures (see section 2.2) that is part of the background subtraction procedure.

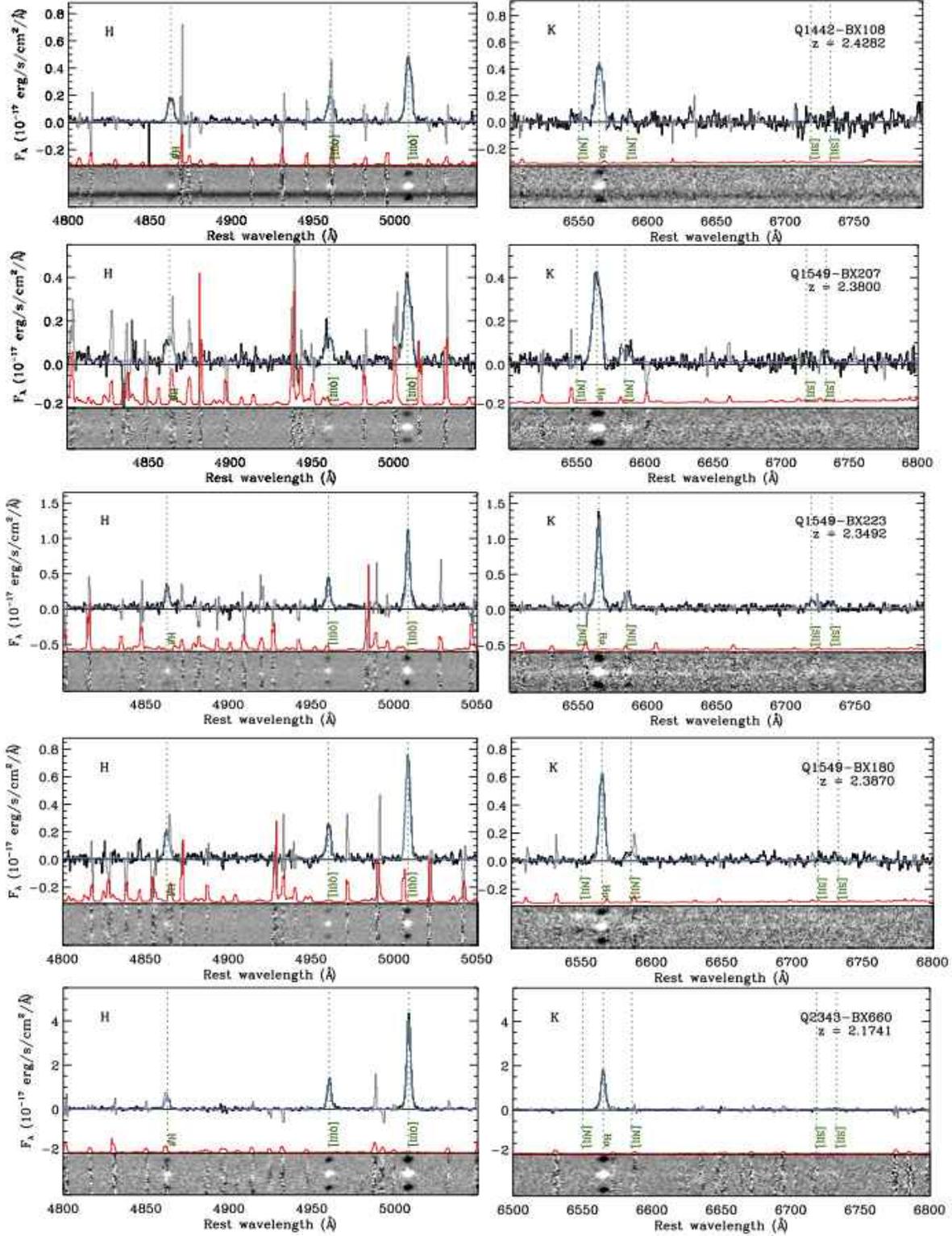


Figure 2. (Continued)

to appear on multiple masks, since (for example) the strength of the T_e -sensitive [OIII] λ 4364 line is expected to be $\gtrsim 50$ times weaker than that of [OIII] λ 5008. The design of a series of masks in a given field proceeded by keeping the highest priority targets on each, and assigning the rest of the available slit “real estate” to different targets according to their relative numerical priorities. A typical mask included 10-15 such fixed targets, out of a total $\simeq 30-35$ slits.

For the other 20-25 targets on each mask, initial priorities were assigned based on the following criteria, from highest to lowest: (1) those with existing high quality UV spectra and known redshifts $2 \lesssim z \lesssim 2.6$, weighted according to their angular separation from the central QSO sightline (2) those flagged as probable high-stellar-mass targets in the redshift range $1.5 \leq z \leq 2.5$, selected using joint optical/near-IR photometric criteria (3) $\mathcal{R} \leq 25.5$ UV color selected galaxies expected to have redshifts within the optimal $2 \lesssim z \lesssim 2.6$ range (in practice, these are the “BX” and “MD” objects defined by Adelberger et al. (2004); Steidel et al. (2004) and Steidel et al. (2003)) (4) rest-UV color-selected galaxies judged likely to have redshifts $z < 2$ or $z > 2.6$ from their rest-UV colors, but not yet confirmed spectroscopically, and (5) UV color-selected galaxies with $\mathcal{R} > 25.5$ (when the depth of the optical photometry allows). Note that category (2) includes galaxies that satisfy the UV color selection criteria and have red optical/IR colors $(\mathcal{R} - K_s)_{AB} > 2$, as well as those with UV colors redder than those of BX/MD galaxies and $(\mathcal{R} - K_s)_{AB} > 2$. Empirically, we have found that the latter criteria, indicated with the prefix “RK”, identify more heavily-reddened galaxies, with relatively large M_* and $1.4 \lesssim z \lesssim 2.5$, that would otherwise not be included in our spectroscopic samples. The “RK” sample and its statistical properties are discussed in more detail by Strom et al (in prep.) The main purpose of including category (2) targets is to improve the sample statistics for galaxies with $\log(M_*/M_\odot) > 10.5$.

Since MOSFIRE mask configurations can be updated easily (and electronically) during an observing run, and the MOSFIRE-DRP (developed by us) produces pipeline-processed 2-D “stacks” in nearly-real-time, the overall efficiency and scientific return of the survey is optimized through quantitative assessment of the data immediately after an observing sequence has completed (generally 20x180s for K band, 30x120s for J or H-band). The results are then used to modify target lists for subsequent masks, performing a running “triage”, in which we remove objects with $z < 2$ or $z > 3$ after they have been successfully identified (replacing them with a new set of targets according to the aforementioned priorities), and evaluate the need for additional integration time for targets in the optimal redshift range.

A significant fraction of targets (both within and outside of the optimal redshift range $2.0 \lesssim z \lesssim 2.6$) requires only 1 hour total integration in a given band to produce spectra of sufficient quality to yield precise nebular redshifts, line widths, and strong-line ratios. However, many targets prove more difficult; our unique iterative procedure is used to ensure that the highest priority or most difficult targets receive the longest total integration times (up to $\gtrsim 10$ hours), that galaxies having useful diagnostics in multiple atmospheric bands are observed in multiple bands, and that minimal time is spent observing objects that do not further the scientific goals.

MOSFIRE observations have been acquired in all 15 KBSS survey regions, though at the time the current sample was finalized a few of the fields had been observed only in H-

band¹².

2.2. MOSFIRE Observational Details

MOSFIRE obtains spectra of up to 46 objects simultaneously within a 6'.1 x 6'.1 field of view at the f/15 Cassegrain focus of the Keck 1 10m telescope. For H and K band spectroscopic observations, a custom made gold-coated reflection grating with 110.5 lines mm^{-1} is used in orders 4 and 3, respectively, providing wavelength coverage of 1.465–1.799 μm (H) and 1.953–2.398 μm (K) for slits in the center of the field of view. Although the MOSFIRE CSU can configure up to 46 slits anywhere across the 6'.1 field, in practice masks observed for our program included 28–34 targets distributed within the central 6'.1 by 3'.0 field of view of the instrument, in order to ensure a large swath of common wavelength coverage for each slit. All masks were designed with slit widths of 0''.7, with lengths in the range 7''.0–23''.0¹³. With 0''.7 slits, MOSFIRE achieves spectral resolution of $R = 3690$ (3620) in the K (H) atmospheric bands, sampled with 2.172 (1.629) \AA pix^{-1} in the dispersion direction, and 0''.18 pix^{-1} spatially. The anamorphic magnification of the spectrometer layout is such that a spectral resolution element is sampled with ~ 2.7 pixels at the MOSFIRE detector.

MOSFIRE observations were acquired using a 2-position nod sequence separated by 3''.0 along slits; individual integrations were 180s and 120s for K and H band, respectively, usually obtained in sequences of \sim one hour total integration time. MOSFIRE’s Hawaii-2RG detector was read out using Fowler sampling with 16 read pairs, resulting in effective read noise of $\simeq 5.3$ electrons (rms). The decision to use 2-3 minute individual integration times between nods was based on significant experimentation with temporal sampling, readout modes, and dither strategies optimized for faint-object spectroscopy; we have recommended the same strategy to other MOSFIRE users via the MOSFIRE web documentation¹⁴.

By design, the integration times used for individual MOSFIRE exposures are sufficient to yield background-limited performance in spectral regions free of strong OH night sky lines, but are short enough to mitigate the effects of the strong and highly-variable OH emission lines on accurate background subtraction. The dark current of the MOSFIRE detector is negligible (< 0.008 electrons s^{-1} pixel^{-1}) relative to the inter-OH background ($\simeq 0.2-0.3$ electrons s^{-1} pixel^{-1} .) Using the ABAB dither sequence of short individual exposures and combining frames taken in positions A and B separately, quite good background subtraction is obtained by simple subtraction (i.e., A-B or B-A) because they have been obtained quasi-simultaneously; residuals are generally seen only in the OH emission lines, which vary significantly on timescales shorter than the 120s (or 180s) nod time. The differencing has the advantage of removing many systematics that would otherwise cause problems for background subtraction, and requires no fitting or resampling of the data. In spite of its Cassegrain location, MOSFIRE is very stable thanks to

¹² Experience has shown that H-band observations, in addition to having the best sensitivity per unit integration time, are most likely to yield spectroscopic redshifts for galaxies without previous spectroscopic identifications, since H α falls in the band for $1.2 \lesssim z \lesssim 1.74$ and [O III] λ 5008 for $1.85 \lesssim z \lesssim 2.59$.

¹³ MOSFIRE slit lengths are quantized, with lengths $(8.0 \times N - 1.0)$ arcsec, where N is the integer number of masking bars comprising the slit.

¹⁴ http://www2.keck.hawaii.edu/inst/mosfire/exposure_recipes.html

active flexure compensation that maintains the spectral format fixed with respect to the detector at the level of better than 0.05 pixels (rms) over the course of a typical 1-2 hour exposure sequence.

The MOSFIRE data reduction pipeline (DRP; described in more detail below) performs the background subtraction in two stages, of which the first is the simple pairwise subtraction of the interleaved, disregistered stacks just described. This is followed by fitting a 2-D b-spline model to the background residuals only, using a method similar to that described by Kelson (2003). We have found that the combination yields background residual errors consistent with counting statistics even in the vicinity of strong OH emission lines.

Because KBSS-MOSFIRE is an ongoing survey and is being conducted in a “tiered” fashion (as described above), the total integration times for observations presented here span a wide range: 3578-29100s (H-band) and 3578-43700s (K-band). The median (average) total integration times for galaxies appearing in Tables 1 and 2 were 6084s (9920s) and 9840s (12240s) in H and K bands, respectively; objects listed in Table 3 have median (average) K-band integration times of 5368s (7832s).

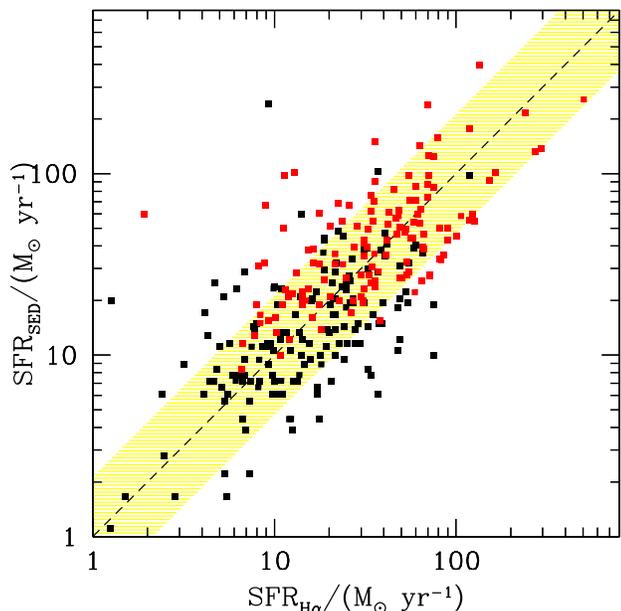


Figure 3. Comparison of star formation rates estimated from SED fitting (SFR_{SED}) with those based on the $\text{H}\alpha$ luminosity ($\text{SFR}_{\text{H}\alpha}$). The dashed line indicates equality between $\text{SFR}_{\text{H}\alpha}$ and SFR_{SED} . Both estimates assume a Chabrier (2003) IMF, with Calzetti et al. (2000) attenuation for the stellar continuum and nebular extinction as given in equations 4 and 5. Galaxies with $E(B-V)_{\text{cont}} > 0.2$ are shown with red points. The median SFR_{SED} and $\text{SFR}_{\text{H}\alpha}$ agree to within $\sim 5\%$, with a median absolute deviation of $\simeq 0.20$ dex (indicated with the shaded region).

MOSFIRE data were reduced using the publicly-available data reduction pipeline (DRP) developed by the instrument team¹⁵. The MOSFIRE DRP produces flat-fielded, wavelength calibrated, rectified, and stacked 2-D spectrograms for each slit on a given mask. The 2-D wavelength solutions in both K and H bands were obtained from the night sky lines on each slit, with typical residuals of 0.08 Å (K) and 0.06 Å (H), or $\simeq 1.1 \text{ km s}^{-1}$ (rms). All spectra were reduced

to vacuum wavelengths, corrected for the heliocentric velocity at the mid-point of each exposure sequence prior to being combined using inverse-variance weighting to form the final 2-D spectra. One-dimensional (1-D) spectra, together with their associated 1σ error vectors, were extracted from the final background-subtracted, rectified spectrograms using “MOSPEC”, an IDL-based 1-D spectral analysis tool developed specifically for analysis of MOSFIRE spectra of faint galaxies (see Strom et al., in prep., for a full description.) Figure 2 shows relevant portions of reduced 1-D and 2-D spectra for 10 of the galaxies in the KBSS-MOSFIRE sample discussed below.

In brief, MOSPEC extracts the 1-D spectrum and its associated 1σ error spectrum, applies a flux calibration determined from wide-slit observations of AOV stars, performs a low-order polynomial continuum fit followed by a simultaneous fit to a user-specified set of emission lines; outputs include redshift, line flux, line width, and the associated uncertainties. The fitted line profiles were constrained to have the same redshift and velocity width within a given observed band (e.g., H-band or K-band); Gaussian profiles were found to provide good fits to the data except in cases with very high S/N ($\gtrsim 50$), where small departures of line shapes from Gaussian may yield statistically significant residuals relative to the models. In such cases, the line intensities and significance were also estimated from direct integration of the line profiles and error vectors within $\pm 3.0\sigma$ of the line center derived from the Gaussian fit. In most such cases, the best-fit Gaussian line profile and direct integration yield line intensities that agree to better than a few percent. For well-detected objects, the statistical uncertainty on measured redshifts was $\sigma_z \simeq 1-10 \text{ km s}^{-1}$ (depending on line width and S/N), and the agreement in redshift between the independently-fitted H and K-band spectra was typically $\Delta z \lesssim 0.0001$, i.e., $\Delta v \lesssim 10 \text{ km s}^{-1}$ (rms).

Balmer emission lines produced by recombination in the galaxy H II regions are expected to be superposed on the corresponding broad stellar absorption lines of the same transitions, which would lead to a systematic underestimate of the Balmer emission line fluxes unless corrections are applied. Unfortunately, the continuum starlight of typical galaxies in the KBSS-MOSFIRE sample is not well-detected, thus any correction relies on population synthesis models of the stellar spectra. For the range of stellar population parameters inferred from the SED fits (section 2.3), the estimated Balmer absorption line strength ranges from $W_{\lambda, \text{abs}} \simeq 3-5 \text{ \AA}$, with the exact value depending on star formation history. Generally, the larger values of $W_{\lambda, \text{abs}}$ are expected for galaxies with the largest $W_{\lambda, \text{em}}$. Since the Balmer absorption lines are of roughly equal strength, but the corresponding $\text{H}\alpha$ and $\text{H}\beta$ emission lines are expected to have an intensity ratio of $\simeq 3$, in practice the correction is always negligible for $\text{H}\alpha$, and is $\lesssim 10\%$ for $\text{H}\beta$ based on experiments using the best-fit SED model as the local continuum but otherwise measuring the lines using a procedure identical to that described above. Given that any correction would be model-dependent, we have chosen here not to apply one to the measured $\log([\text{OIII}]/\text{H}\beta)$, but we remind the reader that the tabulated values could be over-estimated by up to $\Delta[\log([\text{OIII}]/\text{H}\beta)] \lesssim 0.04$ dex; an uncorrected systematic error at that level would not change any of the results described below. In particular, the correction is much smaller for galaxies with the largest $[\text{OIII}]/\text{H}\beta$, since these tend to have the largest emission line equivalent widths.

¹⁵ See <http://code.google.com/p/mosfire/>

The sensitivity of the spectra for detection of relatively narrow nebular emission lines is of course strongly wavelength-dependent even when sky subtraction systematics have been eliminated; the detection sensitivity for a given spectral feature can also be time-dependent, as the intensity of OH emission lines can vary by up to an order of magnitude over the course of an observing night, and a line’s velocity with respect to OH emission lines changes with variations in the heliocentric velocity at the time of the observations. MOSFIRE’s relatively high spectral resolution, dark inter-line background ($0.2\text{--}0.3\text{ e}^- \text{ pix}^{-1}$), high system throughput ($\simeq 40\%$ on the grating blaze in H and K bands), and fast optics (so that background-limited observations are achieved in short integrations times) have all been optimized by design for spectroscopy of faint objects. Thus, we find that, in spectral regions free of strong OH emission and under typical observing conditions, the limiting (5σ , 1 hour) emission line flux (assuming the median line width of $\text{FWHM} \simeq 240\text{ km s}^{-1}$ and a typical spatial extraction aperture of 7 pixels [$\simeq 1''.3$]) is $3.5 \times 10^{-18}\text{ ergs s}^{-1}\text{ cm}^{-2}$ ($4.5\text{--}14 \times 10^{-18}\text{ ergs s}^{-1}\text{ cm}^{-2}$) in H-band (K-band¹⁶). The corresponding limiting fluxes for the median total integration times discussed above (~ 1.7 and ~ 2.7 hours in H, K, respectively) are $\simeq 2.6 \times 10^{-18}$ (H-band) and $\simeq 2.7\text{--}8.5 \times 10^{-18}\text{ ergs s}^{-1}\text{ cm}^{-2}$ (K-band). These sensitivities are within $\simeq 10\%$ of those predicted by the MOSFIRE exposure time calculator XTcalc¹⁷ which we developed during commissioning and subsequently made publicly available.

In this paper, we focus on the subset of KBSS-MOSFIRE galaxies with nebular redshifts $1.95 \lesssim z \lesssim 2.65$ and sufficiently-deep K- and H-band spectra to allow significant detections or useful limits for a minimum set of emission lines¹⁸: $\text{H}\alpha$ $\lambda 6564.61$, $[\text{NII}]$ $\lambda 6585.27$ (in the K band), and $\text{H}\beta$ $\lambda 4862.72$, $[\text{OIII}]$ $\lambda \lambda 4960.30, 5008.24$ (in the H-band). A measurement was considered a “detection” when the statistical significance of both $[\text{OIII}]\lambda 5008$ and $\text{H}\alpha$ was $> 5\sigma$, and that of both $\text{H}\beta$ and $[\text{NII}]\lambda 6585 > 2\sigma$. Undetected $\text{H}\beta$ and/or $[\text{NII}]$ lines were flagged as limits and assigned a flux upper limit of $+2\sigma$. The 82 galaxies for which all features satisfy the criteria for detection are listed in Table 1.

In practice, the most difficult of the BPT lines to detect is the $[\text{NII}]\lambda 6585$ feature, whose intensity is typically only 10% that of $\text{H}\alpha$, and can be substantially weaker in the most metal-poor galaxies (e.g., Erb et al. 2006a; see Figure 5). Because of this, the KBSS-MOSFIRE sample contains a significant number of galaxies for which $[\text{OIII}]$, $\text{H}\alpha$, and $\text{H}\beta$ are detected according to the above criteria, but only upper limits have been measured for $[\text{NII}]\lambda 6585$. These 43 galaxies are listed in Table 2. The current KBSS-MOSFIRE sample with both H and K band measurements includes only one case with a measurement of $[\text{N II}]/\text{H}\alpha$ and $[\text{O III}]$ but only a lower limit on $[\text{O III}]/\text{H}\beta$: Q1603-BX191, which is identified as an AGN (see discussion in section 5.)

For some purposes in what follows below, we have made use of additional KBSS-MOSFIRE galaxies with $1.95 \lesssim z \lesssim 2.65$ and measurements of $\text{H}\alpha$ and $[\text{N II}]$ from MOSFIRE K-band observations, for which comparable H-band observations have not yet been obtained. These 54 galaxies are listed separately in Table 3.

¹⁶ In K-band, the sky continuum level rises monotonically for $\lambda \gtrsim 2.2\mu\text{m}$; the brighter limiting flux is appropriate to the red edge of the band, near $2.4\mu\text{m}$.

¹⁷ <http://www2.keck.hawaii.edu/inst/mosfire/etc.html>.

¹⁸ All wavelengths are in vacuum.

2.3. Stellar Masses and Star Formation Rates

We assigned stellar masses (M_*) to the KBSS-MOSFIRE galaxies using population synthesis SED fits based on photometry in the optical (U_nGR), near-IR (K_s , J, and, for a subset, WFC3-IR F160W), and *Spitzer*/IRAC ($3.6\mu\text{m}$ and/or $4.5\mu\text{m}$, for all but one of the KBSS fields) bands. Prior to performing the SED fits, the near-IR photometry was corrected for the contribution of $\text{H}\alpha$ and $[\text{OIII}]$ emission lines to the broadband fluxes using the spectroscopically measured values. The population synthesis method used is described in detail by (e.g.) Shapley et al. (2005c); Erb et al. (2006c); Reddy et al. (2012b); for the current sample we adopted the best-fit stellar masses using the Bruzual & Charlot (2003) models assuming constant SFR and extinction according to Calzetti et al. (2000). As discussed in detail by Shapley et al. (2005c) and Erb et al. (2006c), typical uncertainties in $\log(M_*/M_\odot)$ are estimated to be $\pm 0.15\text{--}0.25$ dex. Inferred stellar masses and SFRs throughout this paper assume a Chabrier (2003) IMF for ease of comparison with the majority of other galaxy samples considered. For a Salpeter (1955) IMF, both values would be larger by a factor of 1.8.

SFRs were derived from the observed $\text{H}\alpha$ line fluxes after correcting for slit losses and nebular extinction. For simplicity we applied a constant multiplicative factor of 2.0 to correct for flux outside of the MOSFIRE slit; this value was estimated from the reduced 2-D MOSFIRE spectra for each galaxy, assuming that the $\text{H}\alpha$ emission has a symmetric 2-D Gaussian spatial profile. The profile width was estimated from the spatial distribution of $\text{H}\alpha$ emission in the slit direction, from which we calculated the fraction of the total flux falling within the solid angle bounded by the slit width ($0''.7$) and the spatial extraction aperture (median 7.8 pixels, or $1''.4$). This method has the advantage that it accounts for the actual image quality at the slit, averaged over the full duration of an observation, but has the disadvantage that the true 2-D spatial profile of $\text{H}\alpha$ emission is generally unknown and may not be symmetric as assumed. In addition, for cases having significant spectroscopic continuum detections, we estimated the fraction of K-band light lying outside the slit by comparing the detected flux with the expected total based on broadband K_s photometry. By assuming that the $\text{H}\alpha$ flux has the same spatial distribution as the continuum starlight, these estimates provide an alternative estimate of the spectroscopic slit loss correction. We found the two methods to be in reasonable agreement given the uncertainties, with implied correction factors of $\simeq 1.5\text{--}2.5$ and median values close to $\simeq 2$.

We believe that the slit loss correction is the dominant source of uncertainty for individual fluxes, since the instrumental sensitivity corrections (i.e., the absolute flux calibrations) were obtained through a $4''$ slit where slit losses for the calibration stars would be negligible; however, there is an unavoidable $\simeq 3\text{--}5\%$ systematic uncertainty associated with using A0V star spectral models normalized using broad-band photometry (Cohen et al. 2003). Spectroscopic observations of relatively bright stars included on MOSFIRE masks, with the same slit width used for the galaxies, have measured fluxes typically a factor of $\simeq 1.4$ smaller than measured from broad-band photometry of the same stars; this is consistent with slit loss corrections estimated using the same procedure outlined above.¹⁹

¹⁹ As a further test of our flux calibration and slit loss estimates, we compared the observed $\text{H}\alpha + [\text{NII}]\lambda 6585$ fluxes for 18 of the KBSS-MOSFIRE targets in the Q1700 field with those measured for the same galaxies from

Extinction corrections were applied to the $H\alpha$ fluxes using the value of $E(B-V)_{\text{cont}}$ from the SED fits, which assumed the Calzetti et al. (2000) starburst attenuation relation (see e.g. Erb et al. 2006b; Reddy et al. 2010); the present KBSS-MOSFIRE sample has $0 \leq E(B-V)_{\text{cont}} \leq 0.8$ with a median $E(B-V)_{\text{cont}} \simeq 0.2$. It is conventional to assume that nebular emission lines are affected differently by dust compared to the UV stellar continuum, and therefore subject to a different attenuation relation. Calzetti et al. (2000) found a relationship between the reddening of the stars and that of the ionized gas in nearby starburst galaxies,

$$E(B-V)_{\text{neb}} = 2.27 E(B-V)_{\text{cont}} \quad (1)$$

where the color excess for the stellar continuum $E(B-V)_{\text{cont}}$ can be interpreted with the Calzetti et al. (2000) attenuation relation, but $E(B-V)_{\text{neb}}$ is derived from a line-of-sight attenuation relation (e.g., the diffuse Galactic ISM extinction curve of Cardelli et al. 1989)²⁰ In the original calibration of equation 1, a standard Galactic ISM reddening curve was used with measurements of H recombination line ratios to derive $E(B-V)_{\text{neb}}$; for the Cardelli et al. (1989) Galactic extinction curve, $A(0.656 \mu\text{m})_{\text{GAL}}/E(B-V) = 2.52$ [the average ‘‘SMC bar’’ extinction curve of Gordon et al. (2003) has $A(0.656 \mu\text{m})_{\text{SMC}}/E(B-V) = 2.00$] whereas the Calzetti et al. (2000) continuum reddening curve has $A(0.656 \mu\text{m})/E(B-V) = 3.33$. Equation 1 implies that, to use a measurement of the color excess $E(B-V)_{\text{cont}}$ to estimate the attenuation of the $H\alpha$ emission line in magnitudes,

$$A(H\alpha) = 2.52 \times 2.27 E(B-V)_{\text{cont}} = 5.72 \times E(B-V)_{\text{cont}} \quad (2)$$

assuming Galactic extinction, or

$$A(H\alpha) = 2.00 \times 2.27 E(B-V)_{\text{cont}} = 4.54 \times E(B-V)_{\text{cont}} \quad (3)$$

for SMC extinction (Gordon et al. 2003) applied to the nebular emission.

However, the relationship between $E(B-V)_{\text{neb}}$ and $E(B-V)_{\text{cont}}$, and the appropriate extinction curve to be used with either, remains uncertain for high redshift star-forming galaxies. It has been shown that the assumption that $E(B-V)_{\text{neb}} = E(B-V)_{\text{cont}}$ together with the Calzetti et al. (2000) continuum attenuation relation (i.e., that $A(H\alpha) = 3.33 E(B-V)_{\text{cont}}$) yields SFRs consistent with those measured from stacks of X-ray, mid-IR, and far-IR observations of similarly-selected $z \sim 2$ galaxies (Reddy & Steidel 2004; Erb et al. 2006b; Reddy et al. 2010, 2012a). Other analyses, however, suggest higher nebular extinction (see, e.g., Förster Schreiber et al. 2009; Price et al. 2013), particularly for more metal-rich and/or higher mass galaxies, even after

deep, continuum-subtracted narrow-band $H\alpha$ observations, discussed previously by Reddy et al. (2010); Erb et al. (2006b). We found that the median $f_{\text{mos}}/f_{\text{NB}} = 0.50 \pm 0.25$ (error is the inter-quartile range) where f_{mos} is the line flux measured from the MOSFIRE spectra prior to applying slit loss corrections.

²⁰ The relation in equation 1 is often misunderstood to mean that the attenuation of $H\alpha$ emission in magnitudes is higher by a factor of 2.27 than for a continuum photon at the same wavelength; however, it is important to note that equation 1 *assumes* that the two values of the color excess are applied in combination with *different reddening curves*. Although often done, it is incorrect (or at least inconsistent with the original derivation and intended use of equation 1) to use the continuum reddening curve to estimate the attenuation of emission lines. Under the common assumptions that $E(B-V)_{\text{neb}} = C E(B-V)_{\text{cont}}$ (where C is a constant) and that $E(B-V)_{\text{neb}}$ can be multiplied with the selective extinction coefficient at 6564\AA in the Calzetti et al. (2000) attenuation relation, one obtains the same $H\alpha$ attenuation as given by proper interpretation of equation 1 with $1.36 \lesssim C \lesssim 1.72$ (see equations 2 and 3.).

accounting for the extinction curve/color excess interpretation issues mentioned above.

For definiteness, we have assumed the following:

$$A(H\alpha) = 4.54 E(B-V)_{\text{cont}} ; \quad E(B-V)_{\text{cont}} \leq 0.20 \quad (4)$$

$$A(H\alpha) = 5.72 E(B-V)_{\text{cont}} ; \quad E(B-V)_{\text{cont}} > 0.20 \quad (5)$$

equivalent to using an SMC-like extinction curve for galaxies with continuum reddening equal to or below the median value, and Galactic diffuse ISM extinction curve²¹ for those above. Using these corrections, we find that the median $\log(\text{SFR}_{H\alpha})$ and median $\log(\text{SFR}_{\text{SED}})$ for the KBSS-MOSFIRE sample agree to within 0.02 dex; for individual galaxies, the two SFR estimates have a median absolute deviation $\simeq 0.20$ dex (see Figure 3.). Of course, we do not know that agreement between SFR_{SED} and $\text{SFR}_{H\alpha}$ means that either is ‘‘correct’’, but at the very least we can say that they are surprisingly consistent both as an ensemble and on an object-by-object basis. Observations of Balmer line ratios (e.g., $H\alpha/H\beta$) can be used to measure $E(B-V)_{\text{neb}}$ directly, but at present measurements of suitably high S/N are in hand for only a subset of the KBSS-MOSFIRE galaxies, and greater attention to relative calibration between K-band and H-band spectra is required before the line ratios can be confidently used for extinction measurements (and in any case, interpretation of Balmer decrements itself requires the use of an extinction curve).

After the corrections for slit losses and extinction were applied, $H\alpha$ fluxes were converted to luminosities assuming a Λ -CDM cosmology with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and $h = 0.7$ and the Kennicutt (1998) conversion between $H\alpha$ luminosity and SFR (with adjustment to the Chabrier IMF). In what follows, we use the values of $\text{SFR}_{H\alpha}$ listed in Tables 1-3, since they are less strongly covariant with inferred M_* (see, e.g., Reddy et al. 2012b) compared to SFR_{SED} ; however, none of the results presented in this paper would alter significantly if SFR_{SED} were used instead.

Similarly, it is important to emphasize that, with the exception of the inferred SFR (and thus also sSFR), most of the results presented in this paper do not rely on the absolute calibration of emission line fluxes or their extinction corrections; rather, they depend primarily on the intensity ratios of lines observed simultaneously (i.e., in the same atmospheric band) and are sufficiently close to one another in wavelength that differential slit losses or extinction should be negligible.

2.4. Current Sample Statistics

As summarized in Figure 4, the KBSS-MOSFIRE sample includes galaxies with $8.6 \lesssim \log(M_*/M_\odot) \lesssim 11.5$ and $5 \lesssim \text{SFR}_{H\alpha} \lesssim 500 M_\odot \text{ yr}^{-1}$. Specific star formation rates ($\text{sSFR} \equiv \text{SFR}_{H\alpha}/M_*$) range over more than two orders of magnitude, with a median value of 2.5 Gyr^{-1} , in good agreement with median values estimated when SFR is measured using mid and far-IR luminosities in addition to the UV (Reddy et al. 2012b,a). Note that Figure 4 compares histograms of M_* , $\text{SFR}_{H\alpha}$, and sSFR for the sample of 173 galaxies appearing in Tables 1-3 (excluding those flagged as AGN; see section 5) with a ‘‘parent’’ KBSS-MOSFIRE sample of 232 galaxies with the same redshift distribution and $\geq 5\sigma$ $H\alpha$ detections, without regard to whether or not any other lines are significantly detected. Thus, galaxies in the $H\alpha$ sample but not appearing in Tables 1-3 are single-line detections, and of

²¹ The LMC average extinction curve is nearly identical to that of the Galaxy over the relevant wavelength range.

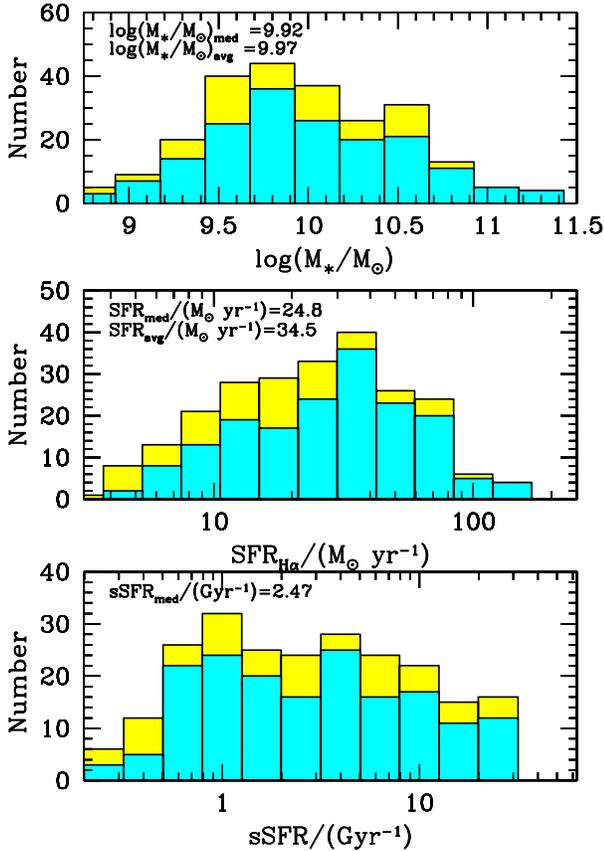


Figure 4. Histograms of stellar mass, star formation rate, and specific star formation rate for the KBSS-MOSFIRE sample with $1.95 \leq z \leq 2.60$. The yellow (light) histogram includes a total of 232 galaxies with $H\alpha$ and stellar mass measurements, independent of whether or not $[NII]\lambda 6585$, $[OIII]$, or $H\beta$ are detected. The cyan (darker) histogram includes 173 of 179 galaxies appearing in Tables 1-3 (excluded are 5 objects identified as AGN [see section 5] and one lacking an M_* measurement.) The mean and median values are given on each panel for the parent sample; the subset with $N2$ and/or $[OIII]/H\beta$ measurements is statistically indistinguishable.

lower overall S/N, usually because their spectra are based on relatively short total integration times.

Of the 174 (non-AGN) galaxies included in Tables 1-3, 124 (70.9%) had prior redshift identifications from optical (rest-frame far-UV) spectra obtained with Keck I/LRIS-B, 22 (12.6%) had been observed previously with LRIS-B without yielding a redshift identification, and 28 (16.1%) had never before been observed spectroscopically. Among all of the KBSS-MOSFIRE observations so far, when the redshift was known from optical spectroscopy to be in the targeted range $2 \leq z \leq 2.6$, more than 90% yielded successful detections of rest-frame optical nebular lines; when the redshift was not known *a priori*, a similar fraction yielded spectroscopic redshifts from the H-band or K-band (or both).

As mentioned above, the selection criteria used for KBSS-MOSFIRE are broader than those of purely rest-UV-color selected samples over the same range of redshifts discussed by (e.g.) Steidel et al. (2004); Erb et al. (2006a,c); Reddy et al. (2010, 2012b); specifically, targets were included whose observed rest-UV and UV/optical colors indicate more heavily reddened galaxies compared to those selected by the “BX” and “MD” criteria. We have also found that the spectroscopic success rate for optically-faint ($R \gtrsim 25$) galaxies within the UV-color selected samples is higher using near-IR

spectroscopy, where most galaxies have strong nebular emission lines, than for optical spectroscopy, where most galaxies have no strong emission lines, and thus identification depends on much weaker absorption lines observed against a faint stellar continuum. Compared to the optical spectroscopic sample of $2 \leq z \leq 2.6$ UV color-selected galaxies in the same 15 KBSS fields (1077 galaxies at the time of this writing), the KBSS-MOSFIRE sample includes a slightly larger fraction of galaxies with $\log(M_*/M_\odot) > 10.5$ (20.1% vs. 17.4%); at lower stellar masses, the KBSS-MOSFIRE sample has a nearly identical fraction of galaxies with $\log(M_*/M_\odot) < 9.5$ (18.3% vs. 18.5%). The median SFR_{SED} is $23.9 M_\odot \text{ yr}^{-1}$ in the KBSS-MOSFIRE sample, and $19.5 M_\odot \text{ yr}^{-1}$ in the full rest-UV spectroscopic sample. In summary, the sensitivity of MOSFIRE for near-IR spectroscopy has produced a spectroscopic sample that is essentially unbiased with respect to the parent photometric sample, at least in terms of SFR and M_* ; this was not the case for the earlier NIRSPEC sample at similar redshifts (Erb et al. 2006a,c).

Realistically, any spectroscopic sample at high redshift, whether based on near-IR or optical spectra, suffers from incompleteness with respect to SFR, which will in turn affect the sample’s distribution of M_* . At the low mass end, for example, even with zero extinction our photometric selection criteria limit the galaxies to $G \lesssim 26$, which corresponds to $SFR \gtrsim 1.3 M_\odot \text{ yr}^{-1}$ using the standard conversion of rest-frame 1500 \AA luminosity to SFR (e.g., Madau et al. 1998) at $z \sim 2.3$; our detection limit for $H\alpha$ corresponds to approximately the same SFR for zero extinction. The same practical limits would apply to even the deepest “ M_* -selected” samples. At the high stellar mass end, larger extinction (nebular and/or UV) may more than compensate for larger overall SFRs, so that the resulting selection function with respect to M_* or SFR becomes potentially complex. The KBSS-MOSFIRE sample is undoubtedly missing high M_* galaxies with very low SFR, which constitute a substantial fraction ($\approx 40\%$) of galaxies with $\log(M_*/M_\odot) > 11.0$ and $z \sim 2.3$ Kriek et al. (2008).

3. THE “BPT” DIAGRAM AT $\langle Z \rangle = 2.3$

3.1. The Locus of Star-forming Galaxies in the BPT Plane

Perhaps the most remarkable aspect of the BPT diagram for local star-forming galaxies is the narrow locus along which most star-forming galaxies are found, sometimes referred to as the “HII region abundance sequence” (Dopita et al. 2000) because the left-hand branch can be interpreted as a sequence in overall ionized-gas-phase metallicity. The tightness of the sequence is controlled by the range within a galaxy sample of some combination of the hardness and intensity of the ionizing stellar radiation field and the properties of the ambient ISM being ionized. At $z \approx 0$, more than 90% of star-forming galaxies fall within ± 0.1 dex of the ridge-line of the sequence (Kewley et al. 2013a); for the SDSS data set used in Figure 5, the scatter in $\log([OIII]/H\beta)$ at fixed $\log([NII]/H\alpha)$ is ≈ 0.11 dex after accounting for measurement errors.

Figure 5 shows definitively what had already been suggested by a handful of earlier measurements for galaxies at $z \sim 1-2.5$ (Shapley et al. 2005a; Erb et al. 2006a; Liu et al. 2008): the nebular spectra of high redshift star-forming galaxies occupy an almost entirely distinct region of the BPT diagram as compared to local galaxies. It has been shown (e.g., Kewley et al. 2001; Kauffmann et al. 2003) that, for local galaxies, the locus of points along the star-forming branch of

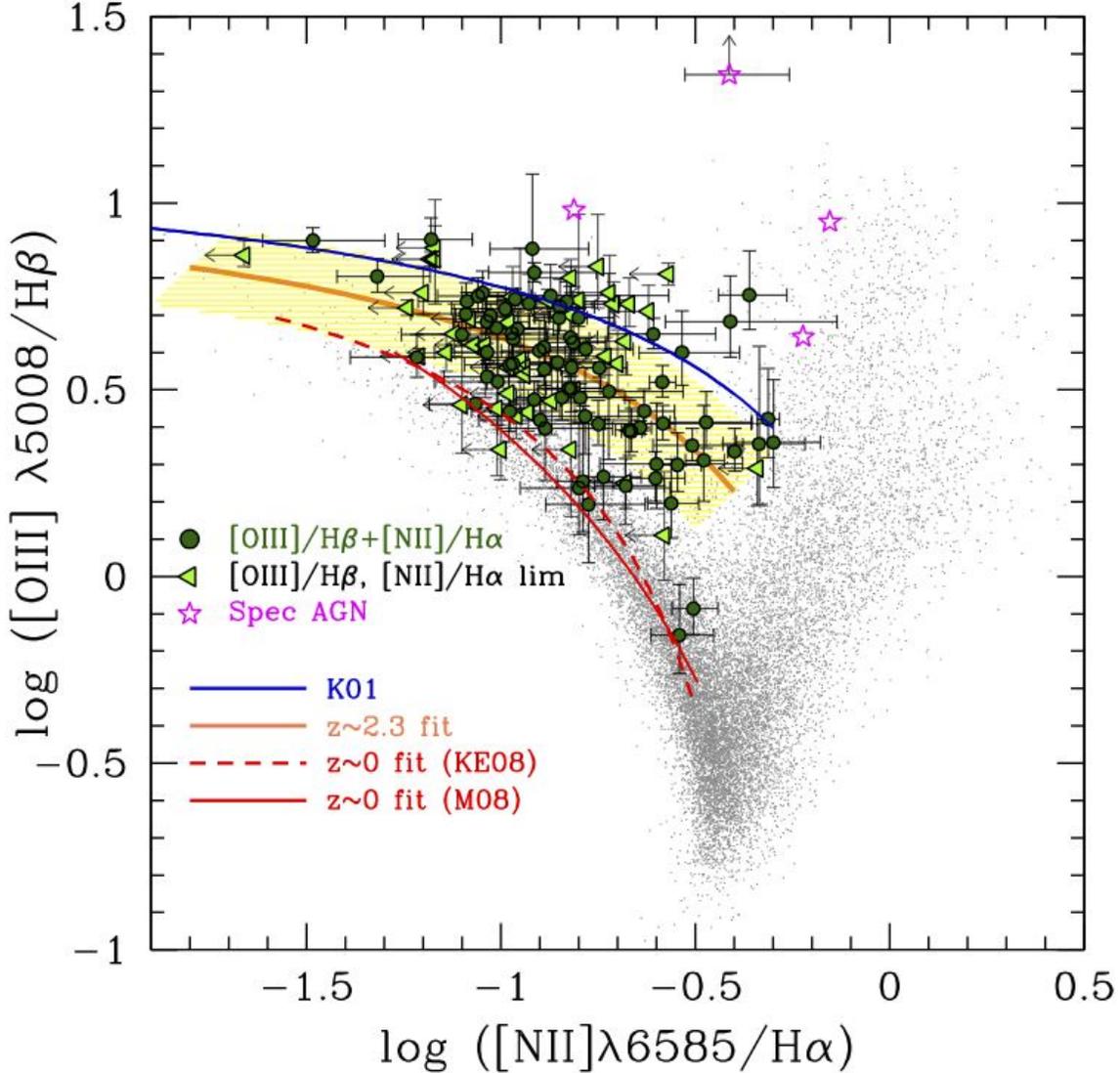


Figure 5. “BPT” diagram for 126 objects with $\langle z \rangle = 2.34 \pm 0.16$ in the KBSS-MOSFIRE survey (large points with error bars) in comparison with the SDSS ($z \simeq 0$) sample (e.g., Tremonti et al. 2004; locus of gray points). The 82 galaxies with measurements in both $[\text{NII}]/\text{H}\alpha$ and $[\text{OIII}]/\text{H}\beta$ are indicated with dark green points, while an additional 43 galaxies with $[\text{OIII}]/\text{H}\beta$ detections and upper limits (2σ) for $[\text{NII}]/\text{H}\alpha$ are light triangles with left-pointing arrows. The red curves trace the “metallicity sequence” of SDSS star-forming galaxies, showing the expected location of galaxies in the BPT plane for oxygen abundances of 0.2–1.0 solar – the solid curve is based on the calibration of Maiolino et al. (2008), while the dashed curve represents the same metallicity sequence implied by the strong-line calibration of Kewley & Ellison (2008). Both curves have been adjusted to the N2 metallicity scale of PP04 for consistency. The blue solid curve is the “maximum starburst” model of Kewley et al. (2001). The orange curve is the best-fit BPT sequence for the KBSS-MOSFIRE sample, with the yellow shaded region tracing the inferred intrinsic dispersion of ± 0.1 dex. Four objects among the 126 have been identified as AGN based on their rest-UV or rest-optical spectra (see discussion in section 5); these are indicated with magenta “stars”.

the BPT diagram can be fit well by a function

$$\log([\text{OIII}]/\text{H}\beta) = \frac{0.61}{\log([\text{NII}]/\text{H}\alpha) + 0.08} + 1.10 \quad (6)$$

(e.g., Kewley et al. 2013a). Fitting the same functional form to the KBSS-MOSFIRE sample in Table 1 yields

$$\log([\text{OIII}]/\text{H}\beta) = \frac{0.66}{\log([\text{NII}]/\text{H}\alpha) - 0.31} + 1.14. \quad (7)$$

Formally, $\chi^2/\nu = 8.88$ for the best-fit model with respect to the data, or a weighted error of $\simeq 0.13$ dex. For comparison to the BPT locus of local star-forming galaxies, it is of in-

terest to estimate the intrinsic scatter (in the absence of measurement errors) of the locus about the best-fit model. To accomplish this, we assumed that the error bars on each point $\sigma_{m,i}$ are the true measurement errors but that the total variance for each point $\sigma_{\text{tot},i}^2 = \sigma_{m,i}^2 + \sigma_{\text{sc}}^2$, where σ_{sc} represents the intrinsic scatter, and is assumed to be a constant (i.e., independent of the measurement errors). The value adopted for σ_{sc} is that which yields $\chi^2/\nu \approx 1$; we find that $\sigma_{\text{sc}} \approx 0.10$ dex – remarkably similar to the scatter observed in the SDSS galaxy sample relative to the best-fit locus (which generally has negligible measurement errors by comparison). Figure 5 (light shaded region) shows that the vast majority of data points

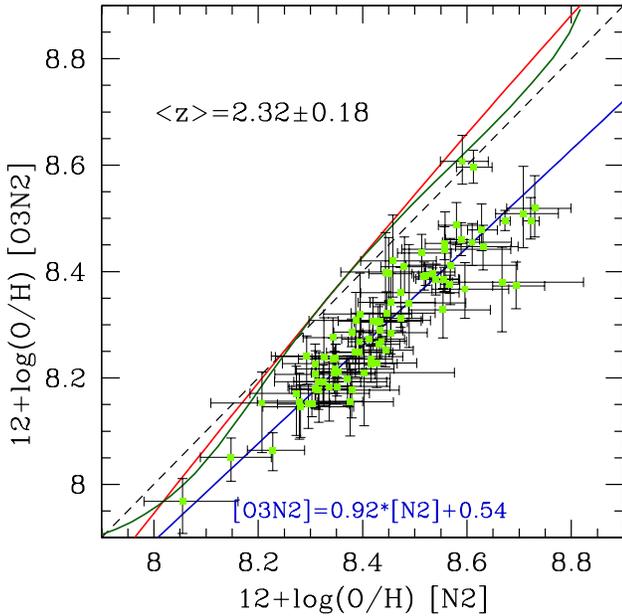


Figure 6. Comparison of the inferred oxygen abundance of the $\langle z \rangle = 2.3$ sample based on the PP04 calibrations of the N2 and O3N2 strong-line indices (points with error bars, where the error bars account for measurement errors only and not for uncertainties in the calibrations). The dashed line indicates the expected relation if the two methods were to give the same value of $12+\log(\text{O}/\text{H})$. The red and green curves represent the best-fit regression formulae for recent re-calibrations of strong line indicators based on the low redshift sample (Kewley & Ellison 2008 and Maiolino et al. 2008, respectively)—the observed scatter in the low-redshift training sets is $\sim 0.03 - 0.04$ dex. The systematically lower values of O3N2-based oxygen abundances as compared to those of N2 are evident, consistent with an offset of $\Delta(12+\log(\text{O}/\text{H})) = 0.15 \pm 0.01$ dex. The blue line is the best linear fit to the relation between the two inferred values at $z = 2.3$. The scatter about this relation, after accounting for measurement errors, is $\simeq 0.04$ dex.

(as well as the points with upper limits on $\log([\text{NII}]/\text{H}\alpha)$) are consistent with a swath in which both $\log([\text{NII}]/\text{H}\alpha)$ and $\log([\text{OIII}]/\text{H}\beta)$ vary by ± 0.1 dex with respect to the best-fit model in equation 7.

Formally, it is difficult to distinguish whether the shift in the locus is primarily due to changes in $[\text{OIII}]/\text{H}\beta$, $[\text{NII}]/\text{H}\alpha$, or both. The shift has implications, independent of its physical origin, for the use of strong-line nebular diagnostics beyond the local universe. As shown with red curves in Fig. 5, the calibrations (or re-calibrations) of the strong line indices imply a one-dimensional curve in the BPT plane, since galaxies of a given value of $12+\log(\text{O}/\text{H})$ map uniquely to values of $[\text{NII}]/\text{H}\alpha$ and $[\text{OIII}]/\text{H}\beta$, with metallicity increasing toward the “right” and “down” along the sequence. The red curves superposed on the $z \simeq 0$ locus in the BPT plane trace the expected metallicity sequence predicted by recently recalibrated strong-line indicators that make use of the same line ratios as appear in the BPT diagram, for galaxies with oxygen abundances from 0.2-1.0 times solar; the solid curve is the best fit regression formula advocated by Maiolino et al. (2008) while the dashed curve is the same locus predicted by the conversion formulae of Kewley & Ellison (2008)²². Not surprisingly, both curves follow the ridge line in BPT space traced by the SDSS sample rather accurately—reflecting the fact that essentially the same set of galaxies was used to establish the best-fit joint calibrations for the relevant strong-

line indices.

The important point is that according to the local calibrations, overall changes in $[\text{O}/\text{H}]$ would simply move objects along these curves; it then follows that any galaxy whose BPT parameters do *not* fall along the calibration sequence cannot yield consistent values of $12+\log(\text{O}/\text{H})$. Stated simply, Figure 5 shows that there is a problem applying a calibration based on local galaxies to a high-redshift sample, even for those that have been re-normalized to consistent metallicity scales for $z \simeq 0$ (e.g., Maiolino et al. 2008; Kewley & Ellison 2008). In practice this means that the “measured” $12+\log(\text{O}/\text{H})$ from strong line ratios will depend systematically on which emission lines are measured. For example, many measurements at $z < 2.6$ rely on the N2 metallicity calibration, since applying it requires observations in only one atmospheric band and the ratio is insensitive to nebular extinction; at $z \gtrsim 3$, on the other hand, estimates are more likely to be based on R23 and other permutations of $[\text{OII}]$, $[\text{OIII}]$, and $\text{H}\beta$, since $[\text{NII}]$ and $\text{H}\alpha$ cannot be observed from the ground. In the latter case, additional issues come into play, e.g., nebular extinction, accurate relative flux calibrations, and the well-known non-monotonic behavior of the line indices.

Figure 6 illustrates the problem in the context of the $z \sim 2.3$ sample: using locally-established metallicity calibrations leads to systematically different metallicities even for the closely-related N2 and O3N2 methods (both calibrations from PP04), which were calibrated primarily using the “direct” or “ T_e ” method and the same set of local H II regions. Interestingly, the *scatter* in the locus of inferred metallicities for the $z \sim 2.3$ sample remains small ($\lesssim 0.04$ dex after accounting for the contribution of measurement errors to the observed scatter), suggesting that a re-calibration at high redshift of the strong-line indicators may produce an equally good, albeit different, mapping of metallicity to strong line intensity ratios²³. The linear regression in Figure 6 serves as an initial estimate of how the conversion might work at $z \sim 2.3$; it will be used in section 7 below.

The question remains whether either of the estimates of $12+\log(\text{O}/\text{H})$ is reliable when applied to galaxies at $z \simeq 2.3$; the answer depends strongly on what factor is primarily responsible for the shift in the BPT sequence at $z \sim 2.3$, and whether it is reasonable to interpret the locus as an abundance sequence as at low redshift. We address this question in section 4 below.

4. PHYSICAL INTERPRETATION OF THE $Z \sim 2.3$ BPT DIAGRAM

The physical cause of the offset of high-redshift galaxies in the BPT plane has recently been explored by a number of authors through examination of the relatively small number of nearby galaxies occupying similar positions in the BPT diagram (e.g., Liu et al. 2008; Brinchmann et al. 2008), using theoretical models (e.g., Kewley et al. 2001; Erb et al. 2010; Kewley et al. 2013a), or a combination of the two (e.g., Shiraazi et al. 2013; Kewley et al. 2013b). Kewley et al. (2013a) in particular have explored in some detail how altering various physical parameters (metallicity, hardness of ionizing radiation field, electron density, prevalence of shocks or AGN versus stellar photoionization) in high redshift HII regions would affect galaxies’ position in the BPT diagram. A common conclusion of most of the recent work is that the main driver of the offset is a higher effective ionization parameter, or the dimen-

²² Both curves were corrected to reflect oxygen abundances consistent with the N2 abundance scale with the PP04 calibration, for consistency.

²³ In section 4.4, we revisit the calibrations of the PP04 N2 and O3N2 metallicity relations and their implications for the high-redshift sample.

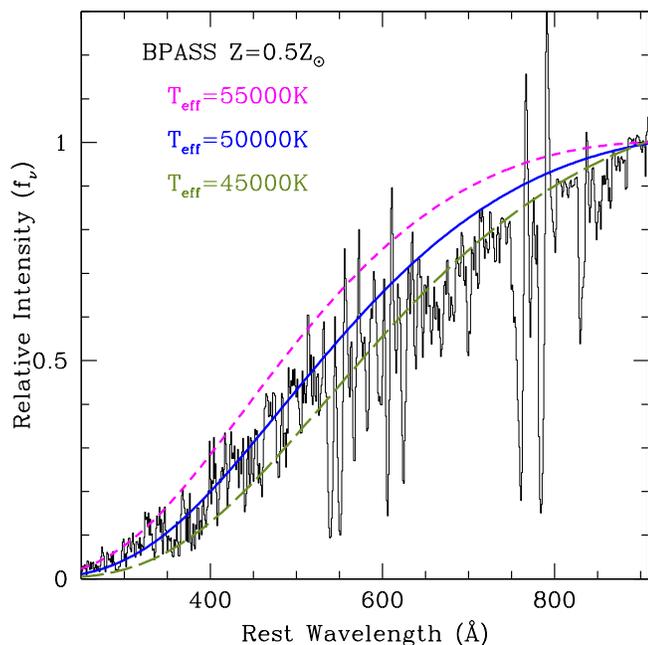


Figure 7. Comparison of blackbody spectra with $T_{\text{eff}} = 45000\text{--}55000\text{ K}$ and a “Binary Population and Spectral Synthesis” (BPASS) population synthesis model with continuous star formation, $Z = 0.5 Z_{\odot}$, an age of 10^8 years, and including the effects of binaries (Eldridge & Stanway 2009). The spectra have been normalized to match at rest wavelength of 912 \AA (1 Ryd). The $T_{\text{eff}} = 50000\text{ K}$ blackbody (blue solid curve) is a good match to the theoretical spectrum, whose metallicity $12+\log(\text{O}/\text{H})=8.4$ is typical of those inferred for the KBSS-MOSFIRE sample.

sionless ratio of the number density of H-ionizing photons to that of H atoms in the H II gas,

$$\Gamma \equiv \frac{\Phi}{n_H} \approx \frac{\Phi}{n_e} \quad (8)$$

where n_H is the number density of hydrogen atoms and Φ is the equivalent density of photons capable of ionizing hydrogen impinging on the face of the gas layer. Γ is analogous to the commonly-used parameter U (or $q = U/c$, where c is the speed of light), except that the latter are generally used in the context of a spherical geometry as in the case of an idealized H II region Stromgren sphere surrounding a point-like ionizing source such as a single O-star. Γ is intended to be more general, and to avoid the connotation of a particular geometrical configuration. The effective ionization parameter Γ obviously depends on both the shape and intensity of the radiation field and the physical density in the ionized gas; the former will in turn depend on the physical density of star formation and the ionizing-luminosity-weighted effective temperature mix of the stars producing the ionizing photons. It (Γ) will also depend on the relative three-dimensional distribution of massive stars, ionized gas, and neutral ISM within a galaxy; in some locations, a packet of gas may be ionized by multiple sources impinging from different distances and directions, each of which has been subject to different modulations of intensity and shape by intervening material.

4.1. Photoionization Models

To gain some intuition, we ran a large grid of model H II regions using CLOUDY²⁴ in which gas-phase metallicity, ionization parameter, physical density, and the effective temperature (T_{eff}) of the stellar ionizing sources were allowed to vary. We initially assumed solar abundance ratios for all elements, but allowed the overall metallicity to range from 0.2 to 1.0 times solar. For simplicity, we began by modeling the UV radiation field shape with a blackbody of temperature $T_{\text{eff}} = 45,000\text{ K}$, motivated by the shape of the rest-frame far-UV spectra of $z \sim 3$ LBGs in a very deep spectroscopic survey (Steidel et al 2014, in prep.). We then extracted the predicted intensity ratios of nebular emission lines, as well as the corresponding values of the N2 and O3N2 estimates of $12 + \log(\text{O}/\text{H})$, for comparison with the model metallicity. We sought the range of model parameters that could reproduce the high redshift BPT data, including the observed trend in the metallicity indicators (Figure 6). We found that higher effective temperatures for the ionizing radiation field were needed to reproduce the $[\text{OIII}]/\text{H}\beta$ ratios of the bulk of the $z \sim 2.3$ galaxies, so the grids were expanded to include blackbody energy distributions with $40,000\text{ K} \leq T_{\text{eff}} \leq 60,000\text{ K}$.

Note that we have deliberately chosen not to use theoretical stellar models in the CLOUDY runs because of the large uncertainties in the ionizing spectra of O stars and the very high density and complex morphology of star formation within the high-redshift galaxies. Instead, we emphasize that we are interested in constraining the effective shape of the ionizing radiation field and the average ratio of ionizing photons to ISM density within the ionized regions *required to reproduce the observations*. In spite of the relative simplicity of our models, we argue that assuming blackbody ionizing spectra is reasonable. For example, Figure 7 shows that blackbody ionizing spectra represent a reasonable approximation to the shape of the 1-4 Ryd stellar continuum of modern stellar population synthesis models (Eldridge & Stanway 2009.) Similarly, we show below that the low-redshift BPT sequence can be adequately reproduced assuming a blackbody ionizing spectrum with $T_{\text{eff}} \simeq 42000\text{ K}$.

For the moment we treat T_{eff} of the ionizing sources independently of the metallicity of the ionized gas producing the observed emission lines. The rationale for doing this, explained in more detail below, is once again that, given the uncertainty in modeling massive star populations as a function of stellar metallicity, we choose to fix the input spectrum in order to better understand the sensitivity of the strong-line indicators to *gas phase* metallicity. Similarly, it seems prudent not to assume that other physical conditions in high-redshift H II regions are similar to local ones until it has been shown to be the case; for this reason, our models include no assumptions about dust, depletion of elements onto dust grains, or non-solar abundance ratios of any elements relative to O.

Figure 8 shows model BPT diagrams where the solid curves and shading are as in Figure 5 but the KBSS-MOSFIRE data points have been suppressed for clarity. Two versions of the model are plotted, representing the approximate range of electron density n_e inferred from the density-sensitive line ratios $[\text{SII}]\lambda 6718/[\text{SII}]\lambda 6732$ and/or $[\text{OII}]\lambda 3727/[\text{OII}]\lambda 3729$ for the sub-sample of KBSS-MOSFIRE galaxies with appropriate measurements (to be described in detail in future work). We focus on the results for $n_e = 1000\text{ cm}^{-3}$ for the purpose of discussion; the main effect of the lower-density model grid is

²⁴ Calculations were performed with version 13.02 of CLOUDY, last described by Ferland et al. (2013).

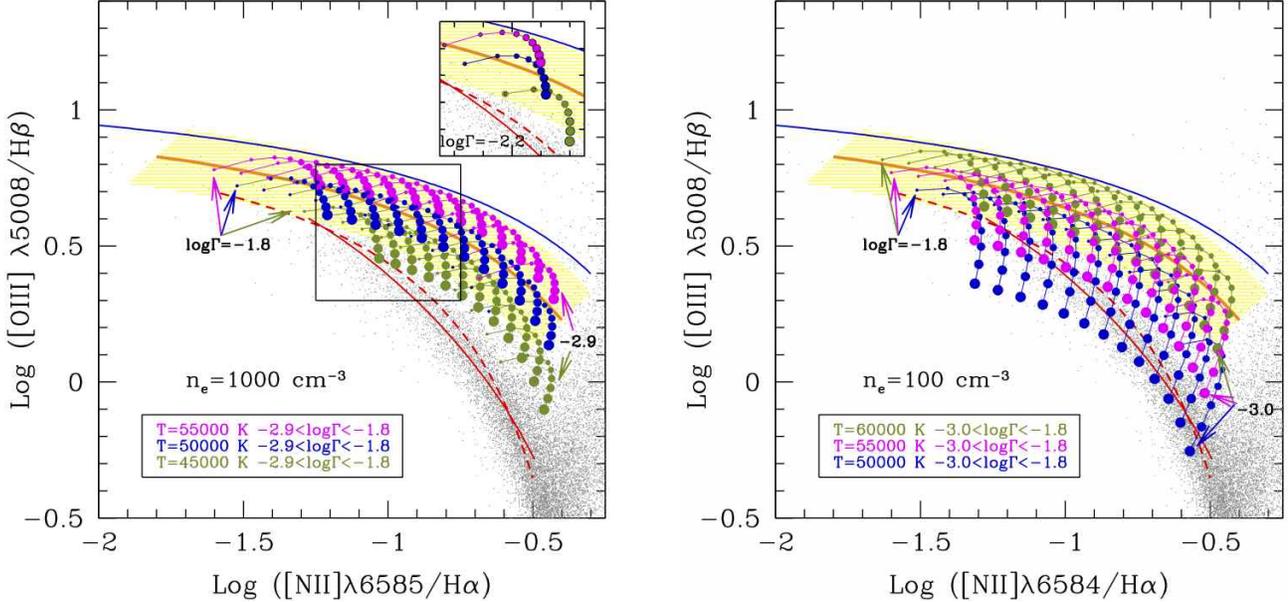


Figure 8. (Left) Predicted locus in the BPT diagram for CLOUDY models described in the text; the colors are coded according to the assumed shape of the ionizing radiation field, for ionization parameters in the indicated range and assuming $n_e \sim 1000 \text{ cm}^{-3}$. The shaded region and the solid and dashed curves are the same as in Figure 5. The curves between points, with the same color coding as the points themselves, connect model runs with the same value of $\log \Gamma$, at intervals of $\Delta \log \Gamma = 0.1$ dex in between each. For each value of $\log \Gamma$, the connected points range in metallicity $Z/Z_\odot = 0.2 - 1.0$ in steps of $\Delta(Z/Z_\odot) = 0.1$, where the point size scales with Z/Z_\odot . The inset panel re-plots the region within the black box, but for a single value of the ionization parameter. This view shows the modest dependence on gas-phase metallicity at fixed Γ and T_{eff} . (Right) As in the left panel, but for models with $n_e = 100 \text{ cm}^{-3}$

to require slightly higher values of T_{eff} to reproduce a given value of $\log([\text{OIII}]/\text{H}\beta)$.

The left-hand panel of Figure 8 shows that the locus of models with $T_{\text{eff}} = 50000 \text{ K}$ and $-2.9 \leq \log \Gamma \leq -1.8$, with metals $Z/Z_\odot = 0.2 - 1.0$, follows very closely the global fit to the KBSS-MOSFIRE BPT data presented above (equation 7); if the $T_{\text{eff}} = 45000$ and $T_{\text{eff}} = 55000$ grids are included, the correspondence with the full distribution of the $z \sim 2.3$ galaxies is remarkably good. Similarly, Figure 9 shows that the same range of model parameters predicts a relationship between the N2 and O3N2 indices in excellent agreement with the observations (cf. Figure 6.)

Unfortunately, in the context of the models that work well to reproduce the observations, neither N2 nor O3N2 is particularly sensitive to the oxygen abundance in the *ionized gas*, which of course is known *a priori* for the models. In fact, the position of the model locus on the BPT diagram is nearly independent of *gas-phase* oxygen abundance over the modeled range (0.2-1.0 times solar); the position along the BPT sequence is sensitive primarily to ionization parameter Γ , while the maximum value of $\log([\text{OIII}]/\text{H}\beta)$ reached in the BPT diagram is closely related to T_{eff} of the assumed ionizing radiation field. Figure 8a shows that the “metallicity sequence”, such as it is, is a very subtle effect, in which a factor of 5 change in gas-phase metallicity moves the locus primarily vertically, but only by $\lesssim \pm 0.05$ dex (and the trend with metallicity is not monotonic). Other strong-line methods would be equally problematic; for example, we find that the same range of model parameters predicts that $\log([\text{OIII}] + [\text{OII}]/\text{H}\beta)$, the ratio upon which the “R23” method depends, is also essentially independent of input gas-phase metallicity (Figure 10). The implication is that, if the models are reasonable, essentially all galaxies in the KBSS-

MOSFIRE sample are consistent with having anywhere from 0.2-1.0 times solar nebular oxygen abundance, and that the strong-line ratios are probably not measuring $12 + \log(\text{O}/\text{H})$ of the ionized gas at high redshifts. The implications for strong-line metallicity measurements are discussed in section 4.2 below.

4.2. Implications for Strong-Line Metallicity Calibrations

The utility of strong lines for estimating metallicity at high redshift can still be salvaged as long as T_{eff} and/or Γ are monotonically correlated with *stellar* metallicity, likely to closely trace the gas-phase metallicity for such young stars. Such correlations are expected at some level, though they are arguably more model-dependent. It is known that early O-type Magellanic Cloud stars of a given spectral classification have T_{eff} higher by several thousand K compared to their Galactic counterparts (e.g., Massey et al. 2005). For example, a systematic shift from $T_{\text{eff}} \simeq 42,000 \text{ K}$ to $T_{\text{eff}} = 50,000 \text{ K}$ for the stars dominating the ionizing radiation field could produce a vertical shift in the BPT diagram of $\Delta(\log([\text{OIII}]/\text{H}\beta)) \simeq 0.3$ dex (see, e.g., Figure 8). In general, stellar metallicity is expected to affect the shape of the ionizing radiation field (Shields & Tinsley 1976) over the critical range 1–4 Rydberg relevant for ionizing H, He, O, N, and S and producing the observed nebular lines. Harder ionizing UV spectra are expected at lower overall metallicity due to reduced metallic line blanketing in the stellar photospheres; metallicity also strongly affects the prevalence, composition, and structure of massive star stellar winds (e.g., Kudritzki & Puls 2000), which in turn have implications for the degree of stellar rotation. These effects and their possible implications are discussed further in section 4.3 below.

If one admits the possibility that the strong line ratios at

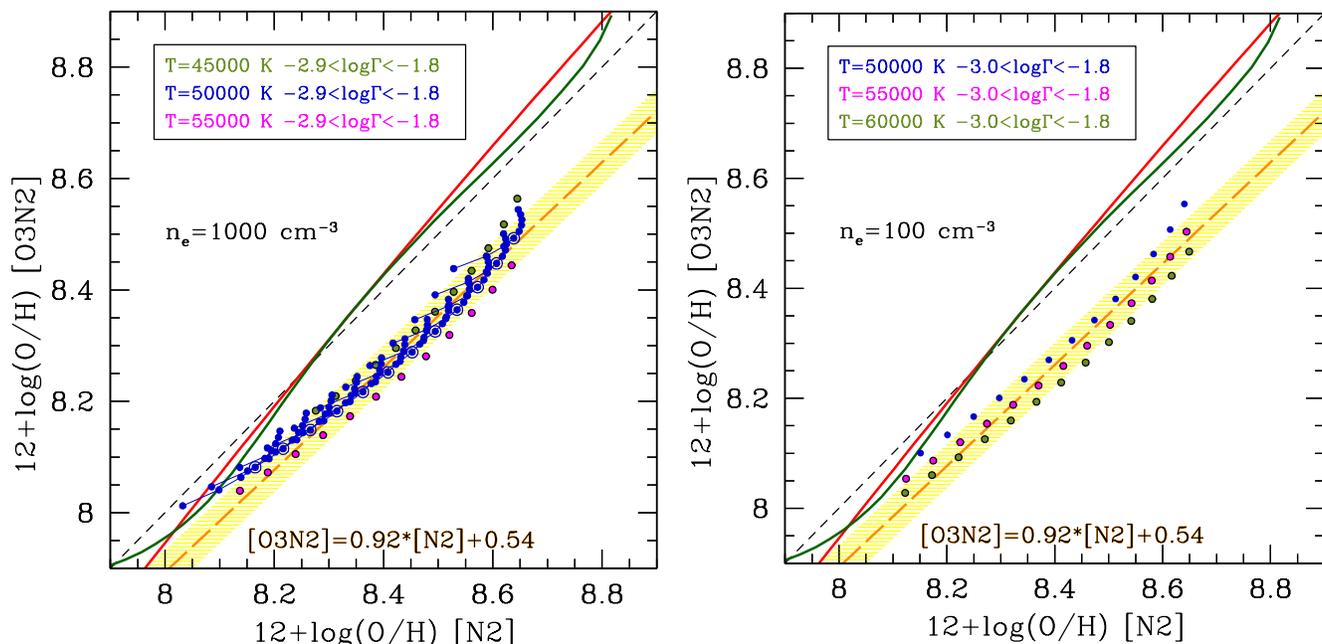


Figure 9. (Left) The predicted run of the N2 and O3N2 metallicity estimates for the same set of model parameters as in Figure 8, which should be compared with Figure 6. The yellow shaded region represents the $\simeq 0.04$ dex scatter inferred from the $z \sim 2.3$ observations relative to the best-fit linear relation (orange dashed line). Note that the run of values for these metallicity estimators is actually almost entirely due to a variation in ionization parameter, rather than gas-phase metallicity, and the scatter about the linear relation is dominated by the differences in T_{eff} considered. As in Figure 8a, models with $T_{\text{eff}} = 50,000$ K are the best overall fit to the observations. (Right) Same as the left-hand panel, for models with $n_e = 100 \text{ cm}^{-3}$. Here only the points for $Z/Z_{\odot} = 0.5$ are plotted, for clarity, to better illustrate the Γ and T_{eff} dependence (at fixed gas-phase metallicity).

high redshifts are driven primarily by factors only indirectly related to gas-phase metallicity, why do they appear to work at low redshift? For the empirically-calibrated O3N2 and N2 relations, for example, the oxygen abundances are rather directly related to galaxy positions along the BPT sequence. Lines of constant O3N2 index are very close to being perpendicular to the low-redshift BPT star-forming locus. However, as we show in Figure 12 and discuss below, a large part of this behavior may be attributable to variations in N/O, and not O/H.

In this context, it is of interest to ask whether the low redshift BPT sequence can be reproduced using simple photoionization models similar to those applied above to the high redshift data. As shown in Figure 12, the characteristic shape of the low-redshift BPT locus can be reproduced by assuming ionizing sources with $T_{\text{eff}} \simeq 42000$ K, ionization parameter extending to slightly lower values than those required for the $z \sim 2.3$ locus, and metallicities $Z/Z_{\odot} \simeq 0.5 - 0.7$, substantially lower than usually ascribed to the low-excitation branch of the BPT diagram. The typical metallicities required to match the nearly-vertical portion of the BPT locus depend to a large degree on assumptions built into models—most commonly, the dependence of (N/O) on (O/H). There is ample evidence in local H II regions for strong systematic variation of N/O with O/H when both have been determined by the direct T_e method; for example, data compiled by Pilyugin et al. (2012) show that for $12 + \log(\text{O}/\text{H}) \lesssim 8.2$, $\log(\text{N}/\text{O}) \simeq -1.45$ ($[\text{N}/\text{O}] = -0.6$), but that for higher oxygen abundance, $\log(\text{N}/\text{O}) \simeq -1.45 + 1.8[12 + \log(\text{O}/\text{H}) - 8.2]$ (see Figure 11, where we show a third-order polynomial fit to the Pilyugin et al. 2012 data set). Qualitatively similar results have been obtained by many other studies (e.g., Vila Costas & Edmunds 1993; van Zee et al. 1998; Henry & Worthey 1999; Andrews & Martini 2013). The generally-

accepted interpretation is that, for low oxygen abundances, H II regions are chemically very young, so that only primary N is present (hence the plateau at $\log(\text{N}/\text{O}) \simeq -1.5$), while at higher O/H, N/H includes contributions from both primary and secondary N enrichment, with the result that N/O increases more rapidly than O/H. Pilyugin et al. (2012) show that the solar ratio of N/O ($\log(\text{N}/\text{O}) \simeq -0.86$) is reached when $12 + \log(\text{O}/\text{H}) \simeq 8.5$, with N/O becoming super-solar [$\log(\text{N}/\text{O}) \simeq -0.6$] for solar (O/H) ($12 + \log(\text{O}/\text{H}) = 8.69$). Other results suggest more gradual changes in (N/O) with (O/H) (e.g., Pérez-Montero & Contini 2009) or a significantly higher (O/H) at the transition from primary to primary+secondary N (e.g., Andrews & Martini 2013)—suggesting that the precise behavior depends on the nature of the calibration sample and the details of the methods used to measure the abundances. Figure 11 illustrates the substantial range of N/O versus O/H from the recent literature (by no means exhaustive).

Clearly, assumptions concerning the behavior of N/O versus O/H directly affect the predicted locations of models in the BPT plane for a given Γ and T_{eff} —for example, lowering N/O by (e.g.) 0.2 dex relative to solar at a given O/H essentially shifts the entire sequence by 0.2 dex in N2, i.e., toward the left in the BPT diagram²⁵. For the same reason, the dependence of N/O on O/H within samples used for local (empirical) calibration of strong line abundance indicators is “built-in” to any method that makes use of the N2 line ratio—

²⁵ In addition to the behavior of (N/O) vs. (O/H), another common assumption in photoionization models is that gas-phase N and O are depleted by amounts similar to those observed in Galactic H II regions (e.g., Esteban et al. 2004), typically $\simeq 0.07 - 0.09$ dex for each. In our models, we have made no attempt to account for O or N outside of the gas phase; assuming the Orion nebula depletions would raise the inferred total abundance of oxygen by $\sim 20\%$ relative to those shown in Figure 12.

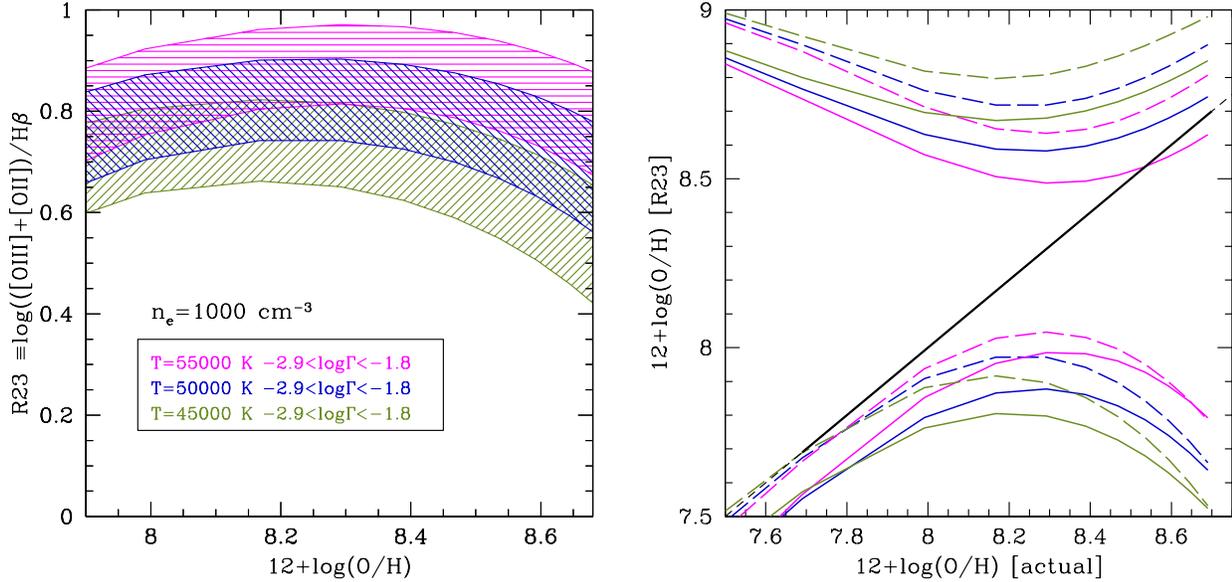


Figure 10. (Left) Predictions of the photoionization models for the “R23” parameter as a function of gas-phase metallicity, where the shaded regions in each color show the effects of varying Γ over the range indicated. (Right) R23-based metallicities versus actual gas-phase metallicity predicted by the model grids. Two curves are plotted for each color-coded set of models: the solid curve corresponds to $\log\Gamma = -1.8$ and the long-dashed curve corresponds to $\log\Gamma = -2.9$. The two distinct sets of curves for a given T_{eff} reflect the well-known double-valued behavior of R23-based metallicities. The lower set of curves uses the lower branch calibration of McGaugh (1991), while the upper set of curves uses the upper-branch calibration of Kobulnicky & Kewley (2004)—both calibrations use O32 to correct for ionization parameter variations in addition to the R23 parameter. The black solid line indicates equality between the measured and actual values of $12+\log(\text{O}/\text{H})$. Note that neither calibration performs well in the $Z \sim 0.2 - 1.0 Z_{\odot}$ range (i.e., $12+\log(\text{O}/\text{H}) = 8.0-8.7$.)

even when no explicit reference is made to N/O. Similarly, the ratio $\log([\text{NII}]\lambda 6585/[\text{OII}]\lambda 3729)$, often used as an indicator of oxygen abundance, is only weakly dependent on O/H when N/O is held fixed, and the mapping of strong line ratios to O/H depends almost entirely on what has been assumed for (N/O) as a function of O/H. The right-hand panel of Figure 12 shows the effects in the BPT plane of the assumption of modest dependence of N/O on O/H,

$$\log(\text{N}/\text{O}) = -1.1 + 0.7(X - 8.0) \quad (9)$$

where $X = 12 + \log(\text{O}/\text{H})$. This relation, which predicts $\log(\text{N}/\text{O}) = -0.62$ for solar O/H, is consistent with the local sample of giant extragalactic H II regions and H II galaxies compiled by Pérez-Montero & Contini (2009), as well as with our inferences for the $z \sim 2.3$ KBSS-MOSFIRE sample as discussed below.

Many commonly-used models have encoded the assumptions into the model grids, and in some cases they assume a very rapid increase of N/O over the most relevant range of O/H (e.g., Charlot & Longhetti 2001; Dopita et al. 2013.) The issue of how N/O affects strong line methods of measuring O/H is discussed in detail by Pérez-Montero & Contini (2009); see section 4.4. In spite of the well-established (though not necessarily numerically agreed-upon) trends of N/O as a function of O/H in nearby H II regions and emission line galaxies, similar measurements are not yet available for galaxies at high redshift.

An interesting, and potentially important, issue emerging from the KBSS-MOSFIRE data and the photoionization models that reproduce them is the implication for (N/O) in the H II regions of the high redshift galaxies: models that simultaneously produce the observed values of $[\text{NII}]/\text{H}\alpha$ and $[\text{OIII}]/\text{H}\beta$ ratios do not obviously require non-solar (N/O). One possibility would be that the models resemble the data by acci-

dent, and that, if the correct N/O dependence on O/H were in the models, a different set of parameters would be required to match the data. However, there also exists evidence that (N/O) may behave differently in the high redshift galaxies compared to local H II regions with the same range in $12+\log(\text{O}/\text{H})$. For example, from stacks of SDSS spectra in bins of SFR, Andrews & Martini (2013) found qualitatively similar (N/O) behavior compared to those of Pilyugin et al. (2012) for galaxies with $\log(\text{SFR}/M_{\odot}\text{yr}^{-1}) \leq 1$, with a “plateau” in N/O at low O/H, and a rapid increase in N/O above $12+\log(\text{O}/\text{H}) \simeq 8.5$ (note that the transition occurs at a value of $12+\log(\text{O}/\text{H})$ 0.2 dex higher than Pilyugin et al. (2012) in spite of the fact that both samples were based on direct metallicity measurements—see Figure 11). However, for the sub-samples of galaxies with SFR similar to those of the KBSS-MOSFIRE sample ($\log(\text{SFR}/M_{\odot}\text{yr}^{-1}) \gtrsim 1$), N/O does *not* appear to vary with O/H and, moreover, *is consistent with solar* (see their Figure 14).

We have also considered the ratio $\text{N2S2} \equiv \log([\text{NII}]\lambda 6585/([\text{SII}]\lambda 6718 + \lambda 6732))$, proposed by Pérez-Montero & Contini (2009) as a sensitive measure of (N/O) in H II regions. N2S2 has the advantage of being insensitive to extinction and (in the case of KBSS-MOSFIRE) involves lines measured simultaneously in the K-band spectra. We find that $\text{N2S2} \simeq -0.1 \pm 0.1$, i.e., nearly constant, in both individual spectra for which both [N II] and [S II] are detected, and in spectral stacks formed from subsets of the KBSS-MOSFIRE $z \sim 2.3$ sample. The photoionization models that reproduce the observed N2 and [OIII]/H β ratios predict N2S2 in the observed range if $\log(\text{N}/\text{O}) \simeq -1.0 \pm 0.1$, independent of gas-phase oxygen abundance for $Z/Z_{\odot} = 0.1 - 1$.

Thus, at present we do not see evidence for the low values of $\log(\text{N}/\text{O})$ ($\simeq -1.5$) that might be expected for very young

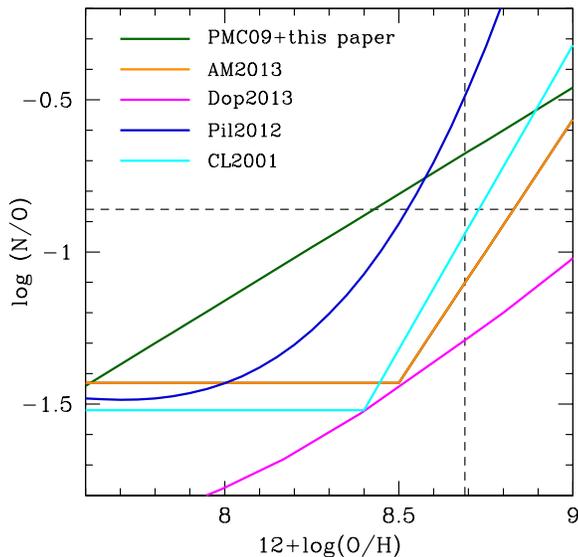


Figure 11. Plot showing examples of N/O versus O/H trends and/or assumptions from the recent literature (PMC09=Pérez-Montero & Contini 2009; AM2013 = Andrews & Martini 2013; Dop2013 = Dopita et al. 2013; Pil2012=Pilyugin et al. 2012; CL2001=Charlot & Longhetti 2001.) The solar values of $12+\log(\text{O}/\text{H})$ and $\log(\text{N}/\text{O})$ are indicated with vertical and horizontal dashed lines, respectively. At low oxygen abundance, several of the results show the “plateau” near $\log(\text{N}/\text{O}) \simeq -1.5$, usually attributed to primary N enrichment. The PMC09 result is based on a collection of both extragalactic H II regions and emission line galaxies having T_e -based measurements of N and O abundances; Pil2012 is based on a literature sample of extragalactic H II regions with T_e measurements, AM2013 is based on stacks of SDSS galaxies in bins of M_* , also based on the direct method. CL2001 is a parametrization of data presented by Henry & Worthey (1999), while Dop2013 is a new fit of data presented by van Zee et al. (1998) (note that Dopita et al. (2013) assumes a much lower value of the solar N/H than our assumption of $\log(\text{N}/\text{H}) = 7.83$).

systems in which only primary N has enriched the ISM, such as damped Lyman α systems (see e.g., Pettini et al. 2008), nor of a strong dependence of (N/O) on (O/H) as expected as secondary N production progresses. Rather, most of the galaxies in our sample appear to have $\log(\text{N}/\text{O})$ within ~ 0.2 dex of $\log(\text{N}/\text{O})_{\odot}$. Of course, we cannot make strong statements about the galaxies with only upper limits on $\log([\text{NII}]/\text{H}\alpha)$.

Figure 13 illustrates another potentially important set of issues, when viewed together with Figure 12: the higher overall excitation of the high redshift BPT sequence has the effect of “compressing” the predicted metallicity dependence of the models, and it is possible to match both the low redshift and high redshift BPT sequences with the same model (with the same modest dependence of N/O on O/H) where the only change was to increase T_{eff} from 42000 K to 55000 K. Figures 12 and 13 show why the lower portion of the BPT sequence, so prominent in the low-redshift galaxy samples where it is extremely sensitive to N/O versus O/H, may be much less apparent in high redshift sample. At both low and high redshifts, it remains the case that the BPT sequence is primarily a sequence in Γ , with the vertical position (i.e., in $\log([\text{OIII}]/\text{H}\beta)$) of the leftward “bend” being primarily sensitive to the radiation field shape. Comparison of the right-hand panels of Figures 12 and 13 shows that a galaxy’s position in the BPT plane may be significantly less dependent on metallicity (N/O or O/H) as compared to local galaxies. It also suggests that the highest metallicity objects at high redshift

might be expected in the region between the two “branches” of the BPT diagram, where they might ordinarily be classified as AGN (see section 5).

4.3. Is $T_{\text{eff}} = 50000$ K reasonable?

While more sophisticated modeling is beyond the scope of this paper, we note that the upper envelope of the KBSS-MOSFIRE $z = 2 - 2.6$ sample (Figure 5) is reasonably well-bounded by the “maximum starburst” model of Kewley et al. (2001), whose main distinguishing characteristic is a much “harder” stellar ionizing radiation field between 1 and 4 Ryd compared to standard stellar models. Indeed, the “maximum starburst” curve looks almost identical to the $T_{\text{eff}} = 55,000$ K blackbody models in the left panel of Figure 8 over the suggested range of ionization parameter. The main point is that high ionization parameter *and* hard (i.e., high T_{eff}) ionizing spectra are both required to easily match the observations.

Stellar models capable of producing the inferred harder radiation field have been proposed— particularly those including binaries and/or rotation (e.g., Eldridge & Stanway 2009; Brott et al. 2011; Levesque et al. 2012; see also Figure 7). In fact, the shape of the ionizing spectra of individual massive stars, and the expected net radiation field from massive star populations, remain very uncertain, both theoretically and observationally. Indeed, there are other areas where tension exists between observations and stellar evolution models— for example, the very blue observed colors of a fraction of star-forming galaxies at $z \gtrsim 2.7$ which indicate little or no “break” shortward of the rest-frame Lyman limit (e.g., Iwata et al. 2009; Nestor et al. 2011; Mostardi et al. 2013), or the possible shortfall of H-ionizing photons from known galaxy populations at redshifts relevant for reionization. This issue is discussed further in section 8 below.

Increased importance of binarity and rotation (which are expected to be strongly coupled, since mass loss in binary systems naturally produces more rapidly-rotating stars— see, e.g. Eldridge & Stanway 2012) among massive stars could reasonably explain a number of qualitative properties observed in the high-redshift H II regions. Aside from producing massive stars that evolve toward *hotter* T_{eff} while on the main sequence, rapid rotation also results in larger UV luminosities and longer main sequence lifetimes at a given mass and metallicity (Brott et al. 2011). The effects become much stronger in stars with metallicities comparable to that of the LMC (assumed in the models to have $12+\log(\text{O}/\text{H})=8.35$, similar to the mean inferred metallicity of the $z \simeq 2.3$ sample). Models suggest that all rapid rotators with $M \gtrsim 30M_{\odot}$ spend a substantial fraction of their lifetimes with $T_{\text{eff}} \sim 50000 - 60000$ K, much hotter than their slowly-rotating counterparts (see Figure 7 of Brott et al. 2011). Rapidly-rotating massive stars also produce much more N during their main-sequence evolution, possibly affecting the gas-phase N/O in the surrounding nebula (see section 4.2 for further discussion). *The implication is that it may not be necessary to invoke unusual numbers of Wolf-Rayet stars, extremely young stellar population ages, or extremely “top-heavy” initial mass functions (IMFs) to understand the high T_{eff} that appear to be common in high redshift galaxies.* It thus seems entirely plausible that the BPT sequence observed for $z \sim 2.3$ galaxies could be driven by changes in stellar evolution that are favored in high redshift star-forming galaxies compared to most galaxies in the local samples.

Of course, there are other potentially important physical processes that could alter the positions of galaxies in the BPT

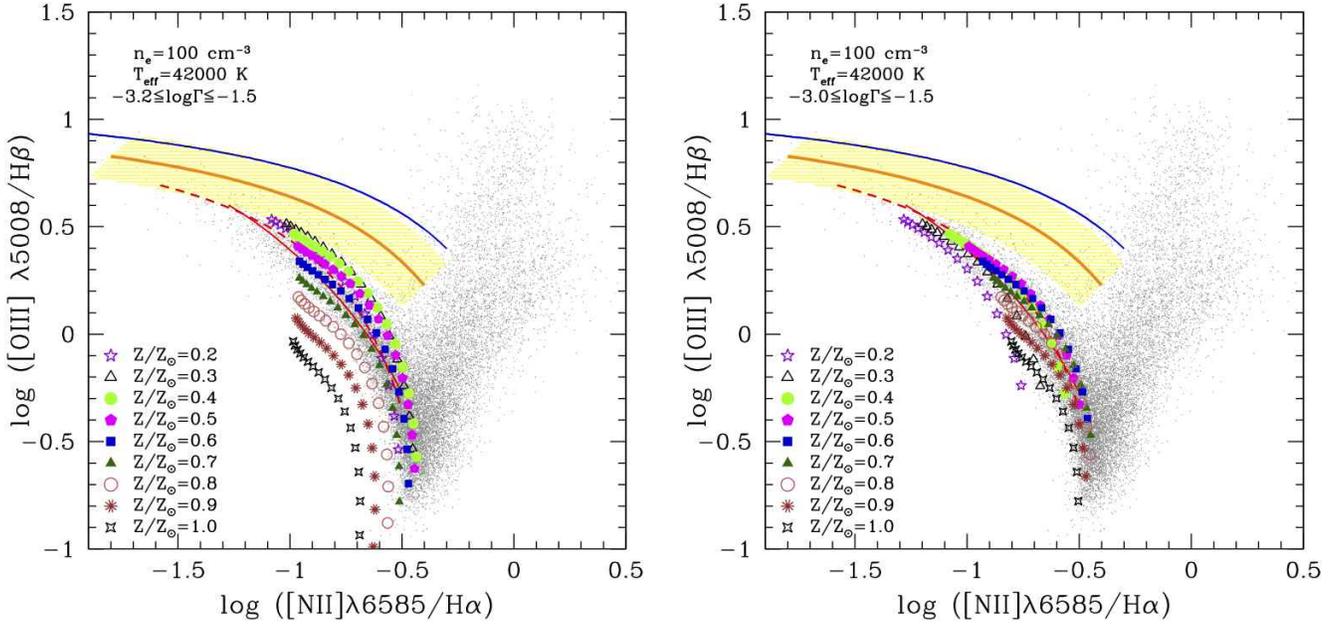


Figure 12. BPT diagram showing example photoionization model grids capable of reproducing the salient features of the star-forming galaxy sequence at low redshift. The point colors and symbols depend on assumed O/H relative to the solar values, with the ionization parameter Γ spanning the same range at each metallicity, in steps of $\Delta \log \Gamma = 0.1$. The assumed radiation field shape is that of a $T_{\text{eff}} = 42000$ K blackbody. (*Left:*) Models for which the solar ratio of N/O (i.e., independent of O/H) has been assumed (*Right:*) Models which have N/O varying with O/H as in equation 9. Note that the assumptions about N/O have two effects: one is to shift the model loci onto the low-redshift BPT locus. The other is to introduce a dependence on O/H of the position along the BPT sequence, at least for $Z/Z_{\odot} \lesssim 0.5$. Note that in both model sets, no assumptions have been made about depletion of N or O onto grains (both assume zero depletion).

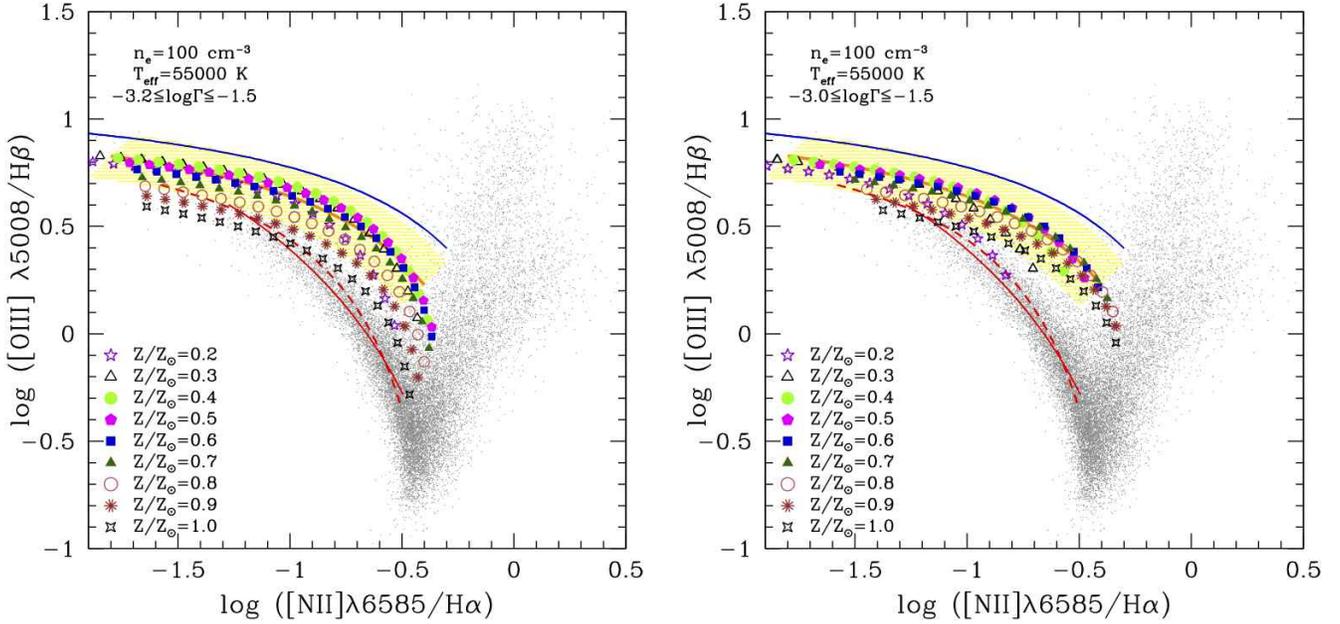


Figure 13. Predictions for the locations on the BPT diagram as in Figure 12, where the only change is that T_{eff} of the ionizing sources has been increased from 42000 K to 55000 K. As in Figure 12, the left-hand panel is for solar N/O, while the right-hand panel assumes that N/O is dependent on O/H according to equation 9.

plane. For example, it has been shown recently that shifts in the BPT diagram can result from H II regions in which radiation pressure dominates over gas pressure, so that the standard assumption of constant electron density breaks down and the structure of the H II zone is fundamentally altered (Verdolini et al. 2013; Yeh et al. 2013.) However, the bulk of $z \sim 2.3$ BPT locus is most easily explained by photoionization, as described above. Shocks certainly may also play a role, though few if any of the KBSS-MOSFIRE sample occupy the regions of the BPT plane where shock-ionization-dominated galaxies are found in the local universe— in particular, the “LINER” region, in the upper right portion of the BPT diagram, is at present completely unoccupied. For the sake of simplicity, since the sample in hand does not appear to require shocks to explain the observations in the context of the BPT diagram, we will not consider them further.

4.4. N2 and O3N2 Calibrations, Revisited

Thus far we have used the calibrations presented by PP04 for mapping the N2 and O3N2 indices to oxygen abundances determined from direct T_e measurements. We have seen above that both of these calibrations are sensitive (through the N2 index) to the behavior of N/O as a function of O/H, so that differences in this behavior between the local calibration set and the high-redshift galaxies could produce systematic differences in inferred $12+\log(\text{O}/\text{H})$. Systematically higher N/O at a given O/H in the high redshift sample could potentially account for the N2-inferred oxygen abundances being systematically higher than the corresponding O3N2 values (see Figure 6).

According to the linear versions of the PP04 calibrations,

$$12+\log(\text{O}/\text{H})_{\text{N2}} = 8.90 + 0.57 \times \text{N2} \quad (10)$$

and

$$12+\log(\text{O}/\text{H})_{\text{O3N2}} = 8.73 - 0.32 \times \text{O3N2} \quad (11)$$

so the dependence of the inferred metallicity on N2 is shallower in the case of O3N2. In addition, both PP04 fits intentionally cover a wide range line indices— considerably wider than the range observed in the current KBSS-MOSFIRE sample— in order to calibrate the index over the widest possible metallicity range. It may be useful in the case of the $z \sim 2.3$ sample to restrict the calibration data set to the same range of N2 and O3N2 index observed, since it allows estimation of the calibration uncertainties most relevant to the high redshift sample (see section 7.1 below).

We have repeated the fits to the N2 and O3N2 metallicity calibrations of PP04, using *the same data set and measurement errors as PP04*, with the following exceptions: first, we limited the regression to the range of N2 and O3N2 line indices observed in the KBSS-MOSFIRE $z \sim 2.3$ sample, and second, we have included only the data points for which the oxygen abundance was measured using the direct T_e method, to reduce the effect of systematics in the overall metallicity scales. The results of the least-squares fits are as follows (see Figure 14:)

For N2,

$$12+\log(\text{O}/\text{H})_{\text{N2}} = 8.62 + 0.36 \times \text{N2} \quad (12)$$

$$(\sigma = 0.13 \text{ dex}; \sigma_{\text{sc}} = 0.10 \text{ dex})$$

where the new fit includes only the PP04 data for which $-1.7 \leq \text{N2} \leq -0.3$ (92 measurements).

For O3N2,

$$12+\log(\text{O}/\text{H})_{\text{O3N2}} = 8.66 - 0.28 \times \text{O3N2} \quad (13)$$

$$(\sigma = 0.12 \text{ dex}, \sigma_{\text{sc}} = 0.09 \text{ dex})$$

where the fit was restricted to the range $-0.4 \leq \text{O3N2} \leq 2.1$, again including only direct T_e -based oxygen abundances (65 measurements). In both relations, σ is the weighted error between the data points and the best fit, and σ_{sc} is an estimate of the *intrinsic* scatter calculated in a manner analogous to that used above for the $z \sim 2.3$ BPT sequence fits. The values of σ should be compared to those obtained by PP04, $\sigma = 0.18$ dex and $\sigma = 0.14$ dex for N2 and O3N2, respectively. Both calibrations become tighter when considered over the smaller range of line index, with $\sigma_{\text{sc}} \simeq 0.10$ dex and $\sigma_{\text{sc}} \simeq 0.09$ dex for N2 and O3N2, respectively. These values should be viewed as the minimum uncertainties in the calibration between the N2 and O3N2 line indices and $12+\log(\text{O}/\text{H})$ at $z \simeq 0$.

Using the revised regression formulae in equations 12 and 13 (shown in Figure 14) lowers the systematic offset between the two indicators when applied to the $z \sim 2.3$ sample, primarily because the coefficient in front of the N2 index in equation 12 is substantially reduced relative to that in equation 10, while the new calibration of [O3N2] (equation 13) is nearly identical to the original PP04 solution (equation 11).

Although we have used the data set assembled by PP04, a more recent calibration of O3N2-based oxygen abundances by Pérez-Montero & Contini (2009) finds an almost identical linear fit to that of PP04, $12+\log(\text{O}/\text{H}) = 8.74 - 0.31 \times \text{O3N2}$, using a larger sample of T_e -based measurements. In addition, these authors examined how the strong-line calibration of O/H would be affected systematically by N/O; they find that the overall scatter is reduced substantially if a term dependent on N/O is included,

$$12+\log(\text{O}/\text{H}) = 8.33 - 0.31 \times \text{O3N2} - 0.35 \log(\text{N}/\text{O}) \quad (14)$$

Equation 14 becomes identical to that of PP04 if $\log(\text{N}/\text{O}) = -1.17$ ($[\text{N}/\text{O}] = -0.3$) and matches the normalization of equation 13 at the median O3N2 of the calibration data set (corresponding to $12+\log(\text{O}/\text{H}) = 8.32$) if $\log(\text{N}/\text{O}) = -1.06$ ($[\text{N}/\text{O}] = -0.2$), close to the values inferred for the high redshift sample.

Thus, there is some cause for optimism that, at least in the case of O3N2, the inferred oxygen abundances are not likely to be strongly biased by differences in N/O between the calibration data set compared to that of the high-redshift sample.

4.5. Direct Metallicity Calibration at $z \sim 2.3$

At present, there is only a handful of direct metallicity measurements at $z > 1.5$ (Villar-Martín et al. 2004; Yuan & Kewley 2009; Erb et al. 2010; Rigby et al. 2011; Christensen et al. 2012; James et al. 2014; Bayliss et al. 2013), some of which are limits only and/or quite uncertain. In any case, as an ensemble they remain insufficient to discern any systematic trends. At minimum, a cross-check on strong-line abundance estimates at high redshifts will require a substantial sample of galaxies covering a range of implied Γ for which each galaxy has both measurements of the doublet ratio of $[\text{OII}]\lambda\lambda 3727, 3729$ and/or $[\text{SII}]\lambda\lambda 6718, 6732$ (for estimates of n_e) and the ratio $[\text{OIII}]\lambda 4364/[\text{OIII}]\lambda 5008$ in addition to the strong lines. Figure 15 shows that, in the context of the models, measurement of the weak $\lambda 4364$ feature would (as expected) provide a relatively model-independent measure of *gas phase* oxygen

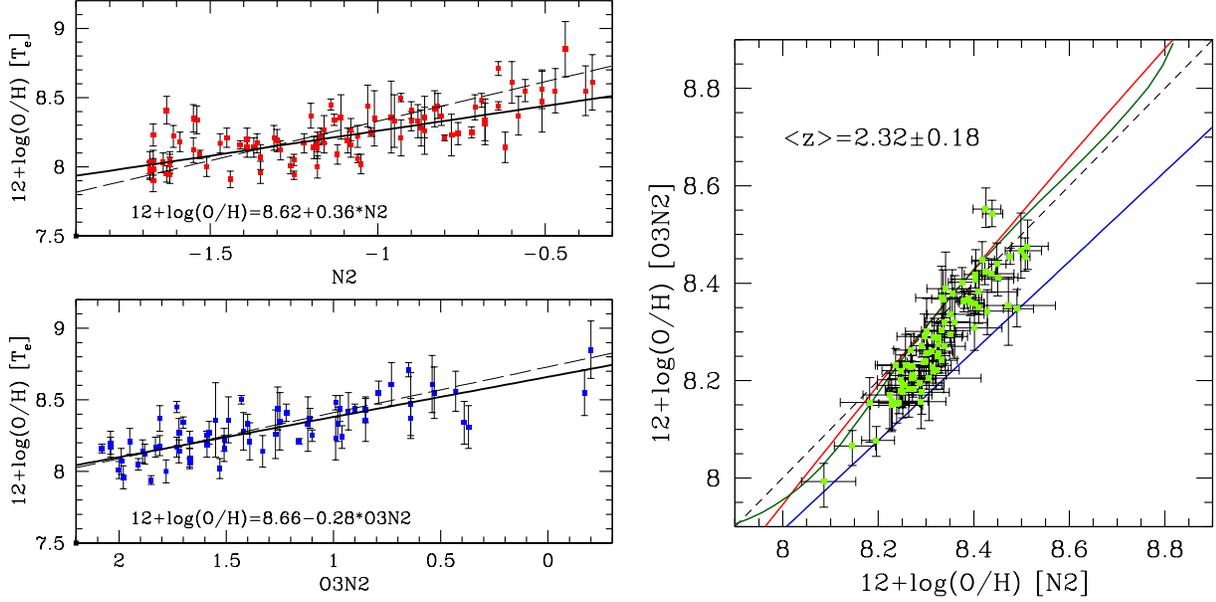


Figure 14. (Left) Linear regressions between T_e -based metallicity and the N2 line index ($N2 \equiv \log([NII]\lambda 6585/H\alpha)$) (top) and O3N2 $\equiv \log([OIII]\lambda 5008/H\beta) - \log([NII]\lambda 6585/H\alpha)$ (bottom) for a subset of the data used by PP04, as described in the text. The modified fits, given as an equation in each panel, are shown by the heavy solid lines. The best linear fits given by PP04 (and used in this paper) are shown with lighter, dashed lines. (Right) Same as Figure 6, but assuming the modified calibrations for O3N2 and N2 metallicity measurements given in equations 12 and 13. The blue solid line is best fit linear relationship from Figure 6, using the original PP04 calibrations.

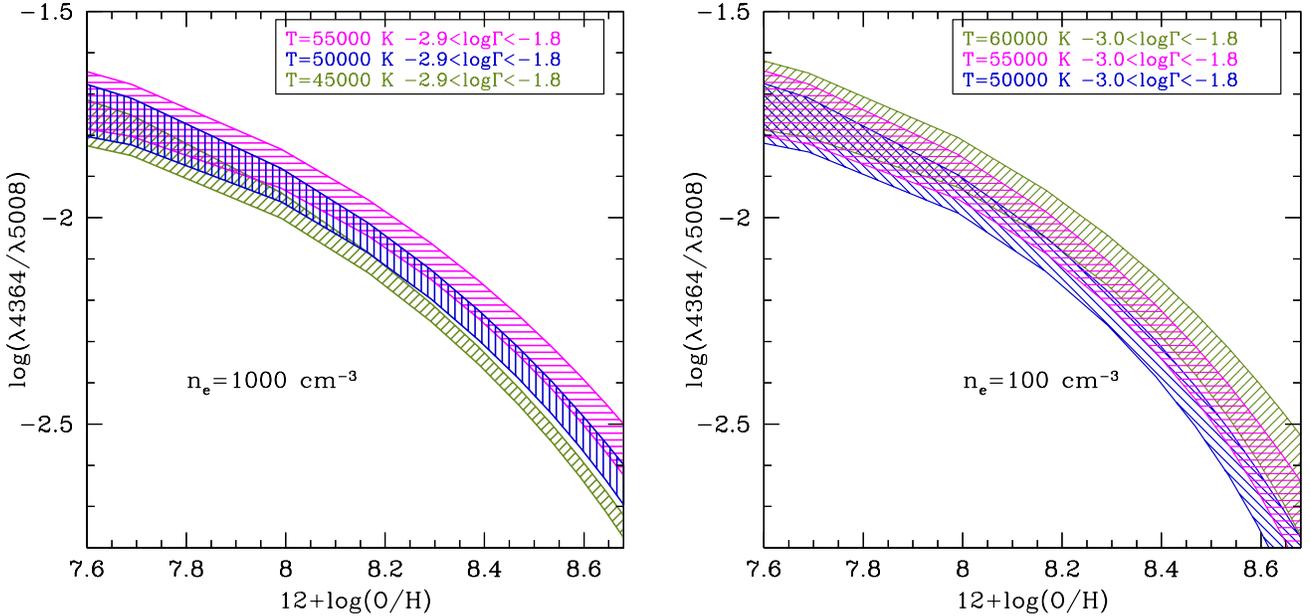


Figure 15. The expected intensity of the T_e -sensitive auroral line $[OIII]\lambda 4364$ relative to $[OIII]\lambda 5008$ for the range of photoionization models as in Figures 8 and 9. The dependence on the radiation field intensity and shape (as well as on n_e) is modest over this range, so that measurements should yield accurate values of $12 + \log(O/H)$ suitable for calibrating the strong-line ratios.

abundance in the high redshift galaxies²⁶. Figure 15 also indicates that uncertainties in the radiation field shape (i.e., T_{eff}) may limit the precision of measuring oxygen abundances to $\sim \pm 0.1$ dex; qualitatively, this may be understood as due to the dependence of equilibrium T_e on the mean energy per ionizing photon, at fixed metallicity.

As for the general detectability of the [OIII] λ 4364 feature, its predicted strength relative to H β ranges from 0.03 to 0.1 for $0.05 \lesssim Z/Z_{\odot} \lesssim 0.5$; the median observed H β flux in the current KBSS-MOSFIRE sample is $f(\text{H}\beta) \simeq 7.5 \times 10^{-18}$ ergs s⁻¹ cm⁻², with median S/N $\simeq 8.6$. Thus, typical KBSS-MOSFIRE spectra are within a factor of a few of the expected flux level $\sim 1 \times 10^{-18}$ ergs s⁻¹ cm⁻², and the best individual spectra, with $(\text{S/N})_{\text{H}\beta} > 40$, should allow for detections. The results of our analysis of the KBSS-MOSFIRE data on this topic will be presented in future work (see also section 6.) Importantly, all of the strong lines (of [O II], [O III], H α , H β , [N II], and [S II]) plus [O III] λ 4364 can be observed from the ground only over the more restrictive $2.36 \lesssim z \lesssim 2.57$ (see Figure 1; by design, 58% of the present KBSS-MOSFIRE sample (Tables 1-3) falls in this range. Thus, we expect that analysis of the highest-quality spectra, together with stacks formed from those of typical quality, should allow for the necessary calibration tests using direct T_e method measurements of gas-phase oxygen abundances— at least for $z \simeq 2.36 - 2.57$.

4.6. Physical Interpretation: Summary

Sections 4.1-4.4 above have attempted to highlight the *caveats* associated with interpreting the ratios of strong emission lines produced in the H II regions of high redshift galaxies in the context of what is known from much more extensively studied “local” star-forming galaxies.

The “offset” in the position of the locus of star-forming galaxies in the high redshift sample compared to the BPT sequence of local star-forming galaxies appears to have contributions, in rough order of importance, from 1) harder stellar ionizing radiation field, needed to explain the preponderance of large observed [OIII]/H β in the high redshift sample 2) higher ionization parameters than inferred for most low-redshift star-forming galaxies 3) a shallower dependence of (N/O) on (O/H) than is typically inferred for galaxies in the local universe, with (N/O) close to the solar value over the full range of inferred (O/H) (see also Masters et al. 2014, which independently reached a similar conclusion). The implications of the BPT shift for measurements of gas-phase abundances from strong emission lines remain uncertain, but the generally higher level of excitation, and the less-pronounced behavior of (N/O) vs. (O/H), has the combined effect of *reducing* the degree to which the strong line ratios are sensitive to gas-phase (O/H). However, the inferred higher T_{eff} , enhanced (N/O), and higher Γ may all be direct consequences of pronounced differences in the evolution of massive main sequence stars in sub-solar metallicity environments at high redshifts. If so, the differences will have very broad implications— perhaps more important than measurement of gas-phase metallicities.

At present, we suggest that metallicities inferred from strong line ratios should be used with caution until they have

²⁶ Alternatively, the UV OIII λ 1661,1666 intercombination feature, which is predicted to be somewhat stronger than λ 4364 over most of the range in physical conditions covered by the models, can be used instead, although its use introduces a much stronger dependence on accurate nebular extinction estimates—see section 6 for examples.

been calibrated directly (i.e., at high redshift) using T_e -based measurements, which has become feasible with the advent of multiplexed near-IR spectroscopy. Based on the (currently limited) observational constraints together with inferences from photoionization models, we suggest that the most reliable of the commonly-used strong line indices is O3N2, whose calibration onto the “ T_e ” abundance scale appears stable with respect to changes in the (low- z) samples used for calibration, and is only moderately sensitive to the behavior of N/O with O/H, unlike N2. The commonly-used R23 method is unfortunately of very limited use over the actual range of $12 + \log(\text{O}/\text{H})$ that appears to be most relevant at $z \sim 2.3$, $8.0 \lesssim 12 + \log(\text{O}/\text{H}) \lesssim 8.7$.

5. AGN VERSUS STELLAR IONIZATION

The BPT diagram has been used most often in the literature as a means of separating galaxies whose nebular spectra are produced predominantly by HII regions from those for which a significant contribution of the observed line emission is likely to have been excited by AGN. The basic principle in distinguishing the star-forming galaxy sequence from that of the so-called “mixing sequence” (see, e.g., Kewley et al. 2001; Kauffmann et al. 2003; Kewley et al. 2013a) is that AGN generally have much harder far-UV spectra than stellar populations, resulting in a tendency to produce higher [OIII]/H β relative to [NII]/H α . In addition, regions near the centers of galaxies harboring AGN tend to be relatively metal-rich, which together with emission from slow shocks that often accompany such activity, pushes [NII]/H α toward high values. The expectations for how high-redshift AGN might behave in the BPT plane have been explored in some detail by Kewley et al. (2013a), who pointed out that AGN in low metallicity hosts (which appear to be extremely rare at $z \simeq 0$) could conceivably be found with high [OIII]/H β but low [NII]/H α . If so, they could fall near to the star-forming sequence in the BPT diagram, possibly leading to ambiguities in classification of objects falling above the $z \simeq 0$ star-forming sequence. This is clearly an important issue to address here, since we have shown that nearly all star-forming galaxies at $z \sim 2.3$ are found in that region.

Fortunately, a deep survey at high redshifts such as KBSS provides some advantages for AGN identification over wide-field samples at $z \simeq 0$ such as SDSS. One is that the identification of active galactic nuclei in distant galaxies has been revolutionized in recent years thanks to pointed, very deep X-ray surveys with CHANDRA and mid-IR photometry with SPITZER/IRAC, which image, respectively, the X-rays produced in AGN accretion disks and emission from AGN-heated dust. In addition, most of the galaxies in our KBSS-MOSFIRE sample have already been observed in the rest-frame far-UV using LRIS; the far-UV provides access to emission lines of much higher ionization species than are easily observed in the rest-frame optical (e.g., C IV λ 1548, 1550, N V λ 1238,1242.) The presence of nebular emission in such species clearly indicates AGN excitation, since the relevant ionization potentials are too high to have been produced by hot stars (see, e.g., Steidel et al. 2002; Hainline et al. 2011.) By combining what is known from UV spectroscopy with multi-wavelength observations sensitive to the presence of AGN, one would normally not need to rely on the BPT diagram as the primary means of discrimination. Nevertheless, it is useful to examine the handful of objects identified as likely AGN to see where they lie in BPT space.

Three objects in the current sample were identified as AGN

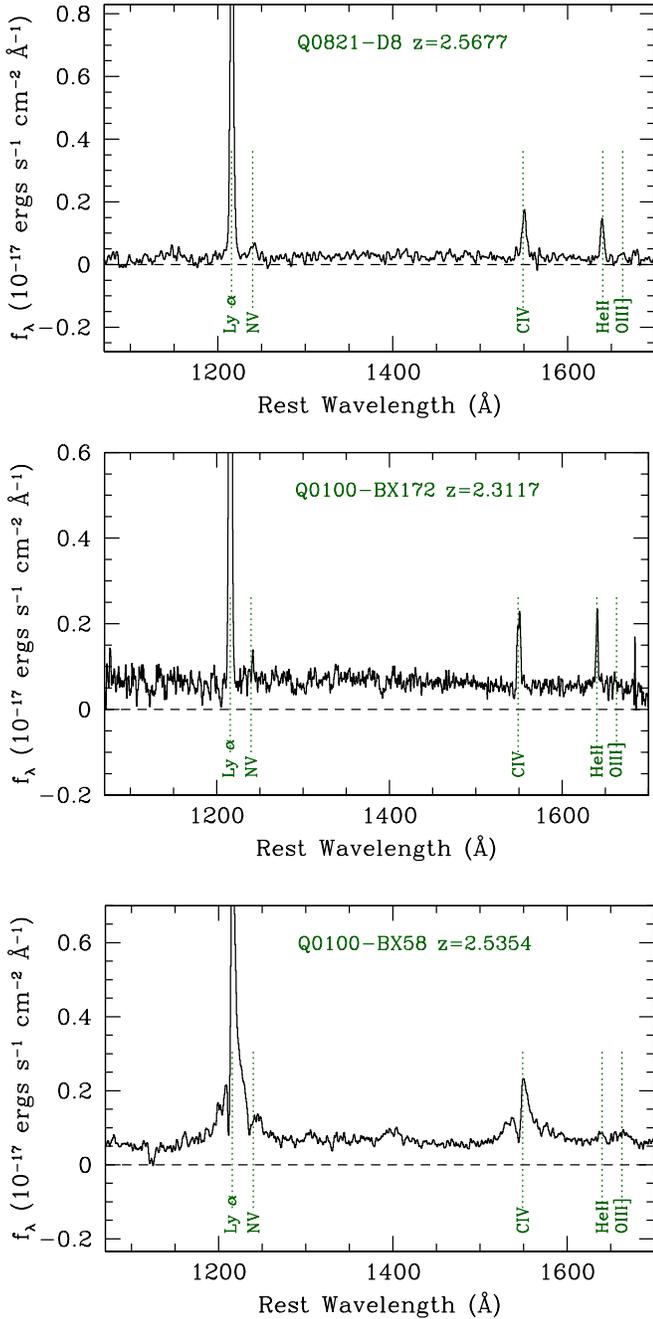


Figure 16. Rest-UV spectra for 3 objects identified as spectroscopic AGN prior to being observed with MOSFIRE; all were obtained using Keck/LRIS-B (Steidel et al. 2004). Their rest-frame optical spectra as observed with MOSFIRE are shown in Figure 17. Emission lines of N V, C IV, and He II clearly indicate the presence of AGN; BX58 (bottom panel) is clearly a broad-lined AGN, albeit quite faint ($\mathcal{R} = 23.4$).

based on existing rest-frame UV spectra (Figure 16.): Q0821-D8, Q0100-BX172, and Q0105-BX58. The first two are narrow-lined AGN flagged by strong emission lines of N V, C IV, and He II $\lambda 1640$; Q0105-BX58 is a faint broad-lined AGN. These objects are indicated in Figure 5, where all 3 lie well away from the bulk of the rest of the sample in the BPT plane (though in rather different locations with respect to one another). In addition to the unusual BPT line ratios, the rest-optical emission lines (even for the narrow-lined objects) are substantially broader than typical among the star-forming

galaxy sample; see Figure 17.

Two additional objects, both among those lacking previous rest-UV spectra, have been classified as AGN: one is Q1603-BX191, the only object in the current KBSS-MOSFIRE sample with a significant measurement of $[\text{NII}]/\text{H}\alpha$ but only a lower limit on $[\text{OIII}]/\text{H}\beta$ (i.e., $\text{H}\beta$ is not detected significantly)—see Figure 5. It is also distinguished by having very broad emission lines, with $\sigma = 223 \text{ km s}^{-1}$ (after correction for the instrumental resolution), placing it in the 98th percentile among the ~ 300 objects with $2 \lesssim z \lesssim 2.6$ observed at $\text{H}\alpha$. Q0821-RK5 is similarly flagged as an AGN because of its observed $\log([\text{NII}]/\text{H}\alpha) = +0.02$ (it has not yet been observed in the H-band, and so does not appear in Figure 5). Such large $[\text{NII}]/\text{H}\alpha$ ratios are reached only by galaxies in the “AGN” portion of the BPT diagram for local galaxies, and thus are very likely to harbor AGN; this is supported also by its very red color $(\mathcal{R} - K_s)_{\text{AB}} = 3.59$, its very broad and diffuse emission lines ($\sigma = 246 \text{ km s}^{-1}$, or 99th percentile of all measured $\text{H}\alpha$ line widths) and its huge inferred stellar mass ($\log(M_*/M_\odot) = 11.79$, by far the largest in the sample) likely due to contamination by hot dust emission in the observed near-IR (see Hainline et al. 2012). In this particular case the SED is ambiguous due to the lack information for observed wavelengths $> 2 \mu\text{m}$, since Q0821 is the only KBSS field lacking deep IRAC coverage.

Among the KBSS galaxies with no evidence for AGN (e.g., those falling within or consistent with the shaded region in Figure 5), there appears to be a rather sharp upper limit of $\log([\text{OIII}]/\text{H}\beta) \lesssim 0.9$. This threshold is exceeded only by two of the objects flagged as AGN. As discussed below in section 6, this upper envelope in $[\text{OIII}]/\text{H}\beta$ implies a maximum T_{eff} for sources dominating the ionizing radiation field, $T_{\text{eff,max}} \sim 55000 - 60000 \text{ K}$ based on our modeling. AGN can exhibit a wide range of energy distributions over the important 1-4 Ryd range, but their shape is approximately a power law rather than an exponential (as in the case of stars) over this interval. Figure 18 helps to illustrate the difference between the assumed blackbody spectra and power-law AGN spectra. Clearly lines from species with ionization potential above $\simeq 50 \text{ eV}$ (or recombination lines from them) would be more unambiguous signatures of AGN excitation compared to $[\text{OII}]$, $[\text{OIII}]$, and $[\text{NII}]$ transitions. On the other hand, Figure 18 shows that these ions (particularly $[\text{OIII}]$) are extremely sensitive to T_{eff} —e.g., increasing T_{eff} from 45000 K to 55000 K approximately triples the number density of photons capable of ionizing $\text{O II} \rightarrow \text{O III}$ relative to those that can ionize $\text{H I} \rightarrow \text{H II}$. Although the possible range of AGN SEDs is large, and while AGN would tend to produce larger ratios of $[\text{OIII}]/\text{H}\beta$, they would nevertheless be unlikely to produce a *consistent* upper envelope at a particular value ($[\text{OIII}]/\text{H}\beta \simeq 0.9$) as observed in the $z \sim 2.3$ sample.

Thus, we conclude that AGN excitation plays a significant role in only a small fraction of the KBSS-MOSFIRE sample. The 5 identified AGN have been excluded from most analyses in this paper, since their strong line ratios are unlikely to be related to stellar processes.

6. LOCAL ANALOGS OF KBSS-MOSFIRE GALAXIES

It is potentially instructive to examine what is known about a relatively local population of galaxies that in many respects resembles both typical and extreme members of the KBSS-MOSFIRE sample at $z \simeq 2.3$: the so-called “green pea” (GP) galaxies (e.g., Cardamone et al. 2009) are relatively rare $z \simeq 0.2$ objects selected by their distinctive colors caused by

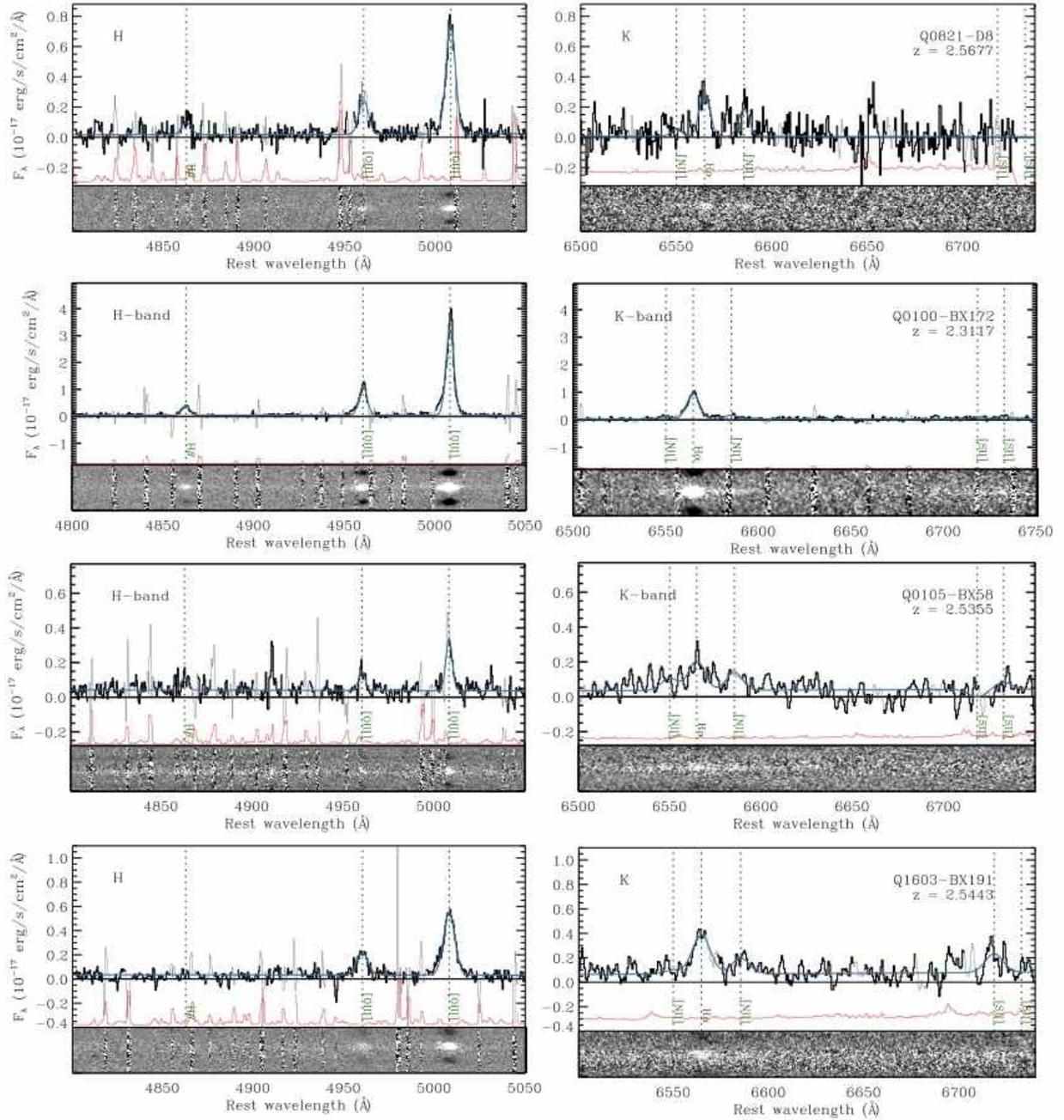


Figure 17. (Top:) MOSFIRE spectrum of galaxy Q0821-D8, one of 3 objects in the current sample to have been identified as AGN based on their rest-frame UV spectra. Note the much larger ratio of $[\text{NII}]\lambda 6585/\text{H}\alpha$ compared to any of the objects shown in Figure 2, as well as its corresponding location in the BPT plane (Figure 5). (2nd row:) Q0100-BX172, (see Figure 16) is also clearly an AGN from its UV spectrum; in the rest-optical MOSFIRE spectra, its line ratios place it close to the locus of galaxies but still significantly higher in $[\text{O III}]/\text{H}\beta$ (Figure 5). (3rd row:) Q0105-BX58, identified unambiguously as a broad-lined AGN from its rest-UV spectrum shown in Figure 16. (Bottom:) H and K band spectra of Q1603-BX191, which had not been observed in the rest-UV, but is identified as an AGN on the basis of its large $[\text{OIII}]/\text{H}\beta$ lower limit together with relatively high $[\text{NII}]/\text{H}\alpha$ (see Figure 5)

unusually large $[\text{O III}]$ equivalent widths. As shown in Figure 19, these rapidly star-forming, compact galaxies occupy much of the same region of the BPT plane as the high redshift objects. The sample of 9 GPs in Figure 19 is comprised of 6 “extreme” GPs studied by Jaskot & Oey (2013) and 3 “normal” GPs with very deep follow-up spectroscopy (Amorín et al. 2012). These 9 galaxies have a complete set of strong nebular lines as well as metallicity measurements based on the direct T_e method; they serve as a possible “preview” of the efficacy of the strong line metallicity measurements for similar

galaxies at higher redshifts.

First, we note that our simple photoionization models [more sophisticated, but less general, models were presented by Jaskot & Oey (2013) and Amorín et al. (2012) in interpreting their data] presented in section 4 above are able to reproduce both the position of the GPs in the BPT plane (left panel of Figure 19) and the behavior of the N2 and O3N2 indices measured from the GPs’ strong line ratios (right panel of Figure 19). The corresponding direct method metallicities (indicated in Figure 19b beside each point) show good

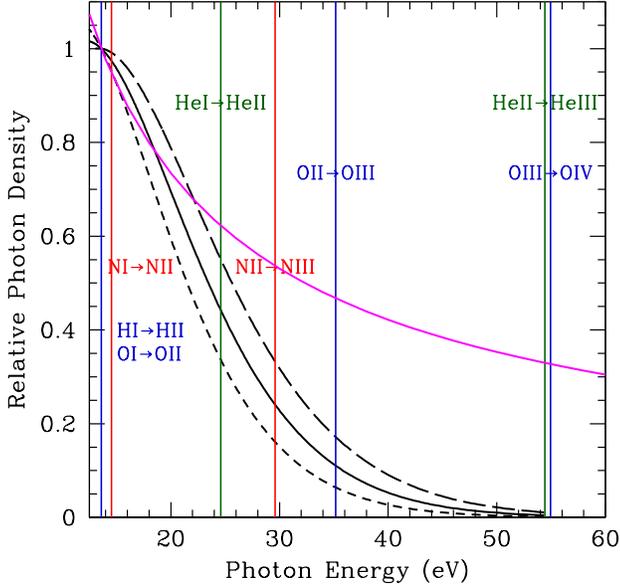


Figure 18. Illustration of the relative photon density of ionizing photons under various assumptions about the ionizing radiation field. The vertical lines indicate the ionization potential for ions most relevant to the observations. The black curves are for blackbody spectra with $T_{\text{eff}} = 45000$, 50000 , and 55000 K (short-dashed, solid, and long-dashed lines, respectively). The magenta curve is for a power law spectrum of the form $f_\nu \propto \nu^{-\alpha}$, with $\alpha = 0.8$. agreement when the O3N2 and N2-based numbers are near $12+\log(\text{O}/\text{H}) \simeq 8.0$ (true for all 6 of the “extreme green peas”), but the 3 “normal” GP galaxies suggest a possible issue²⁷: their direct metallicities are comparable to or even lower than those of the “extreme” examples, but the strong-line indices imply higher metallicities— and the discrepancy may become marginally worse along the sequence followed by both the $z \sim 2.3$ galaxies and by the ionization parameter sequence in the photoionization models. Part of the disagreement relative to the local strong-line calibrations, as pointed out by Amorín et al. (2010), is due to the fact that the GP galaxies appear to have higher (N/O) than typical local galaxies of the same oxygen abundance. The 3 examples in Figure 19 all have $\log(\text{N}/\text{O}) \simeq -1.0$, close to the solar ratio. We recall from section 4 above that roughly solar (N/O) was also inferred for most of the $z \simeq 2.3$ galaxies on the basis of the photoionization models. The “extreme” GPs, on the other hand, appear to have (N/O) consistent with that of local metal-poor dwarf galaxies, with $\log(\text{N}/\text{O}) \sim -1.5$, normally interpreted as systems in which only primary N enrichment has occurred (e.g., van Zee et al. 1998).

6.1. “Extreme” Galaxies at $z \simeq 2.3$

In fact, the extreme GPs have properties very similar to our most extreme $z \sim 2.3$ galaxies— most of which have only upper limits on $[\text{N II}/\text{H}\alpha]$ (light green triangles in Figure 19a) as well as the highest values of $[\text{O III}]/\text{H}\beta$ in the sample, with $\log([\text{O III}]/\text{H}\beta) \simeq 0.9$. Thus, it appears that the extreme

²⁷ We note that the GP whose direct metallicity is most discrepant with the strong line estimators (GP232539; Amorín et al. 2012) has $z = 0.277$, which places the weak $[\text{O III}]\lambda 4364$ feature at an observed wavelength of ~ 5572 Å, so that it is possible that its measured intensity has been affected by residuals from the strong 5577 Å night sky emission line; there is a positive residual at the position of the (weaker) NaD night sky line in the spectrum.

galaxies at both low and high redshift share a common “upper envelope” in the BPT diagram, previously discussed in section 5.

Three galaxies in the current KBSS-MOSFIRE sample (Q2343-BX418, Q2343-BX660, and Q0207-BX74) are found near the EGPs in the BPT diagram, and have particularly good MOSFIRE (J, H, and K bands) as well as LRIS-B (rest-frame UV) spectra; they are indicated in Figure 19. Q2343-BX418 ($z = 2.3053$) was studied in detail by Erb et al. (2010) as a prototypical example of a UV-bright galaxy with little or no reddening, strong Lyman α emission, and unusually strong rest-UV nebular lines of $[\text{O III}]\lambda\lambda 1661, 1666$ and $[\text{C III}]\lambda\lambda 1906, 1909$. Q2343-BX660 ($z = 2.1741$) and Q0207-BX74 ($z = 2.1889$) have similar rest-UV spectra to that of BX418, as shown in Figure 20, as well as similar rest-optical strong line ratios. All 3 galaxies have $\log(M_*/M_\odot) \sim 9.0$ (very similar to the GP sample introduced above), $\text{SFR} \simeq 30 - 50 M_\odot \text{ yr}^{-1}$, and among the lowest inferred oxygen abundances and highest sSFRs in the current KBSS-MOSFIRE sample. Both Q2343-BX418 and Q2343-BX660 are known to be compact, both in $\text{H}\alpha$ emission (from Keck/OSIRIS laser guide star AO IFU observations; Law et al. 2009) and in the rest-optical continuum (from HST/WFC3 F160W observations; Law et al. 2012). Q0207-BX74 appears to be similarly compact, though as yet no AO or HST observations are available.

Erb et al. (2010) used measurements of rest-UV $[\text{O III}]\lambda 4364$ intercombination lines (at the redshift of BX418, $[\text{O III}]\lambda 4364$ does not fall within one of the ground-based atmospheric windows) together with rest-optical nebular emission based on Keck/NIRSPEC spectra to measure T_e and thus “direct” metallicities, finding $12+\log(\text{O}/\text{H}) = 7.8 \pm 0.1$, where some of the uncertainty stems from a non-detection of $[\text{O II}]\lambda 3727, 3729$ (so that the contribution of O^+ to O/H could not be determined). We have re-observed Q2343-BX418 with MOSFIRE in J, H, and K bands, covering, in addition to the BPT line ratios (Table 1 and Figure 5), the $[\text{O III}]\lambda\lambda 3726, 3729$ doublet, detected with $\text{S/N} \sim 30$. The observations in the 3 near-IR bands were obtained on the same night and carefully cross-calibrated to remove any differential slit losses using observations of a calibration star placed on one of the slits for all 3 observations. We used these observations to calculate the electron density n_e from the $[\text{O II}]\lambda 3727, 3729$ doublet ratio, the Balmer decrement ($\text{H}\alpha/\text{H}\beta$), the ratio

$$\text{O32} \equiv [\text{O III}](\lambda 4960 + \lambda 5008) / [\text{O II}](\lambda 3727 + \lambda 3729) \quad (15)$$

and, using the new $[\text{O III}]\lambda 5008$ measurement together with the rest-UV measurement of the $[\text{O III}]\lambda\lambda 1661, 1666$ intercombination feature presented by Erb et al. (2010), the electron temperature $T_e[\text{O III}]$, from which $(\text{O}^{++}/\text{H}^+)$ was derived. The ratio (O^+/H^+) was inferred assuming $T_e([\text{O II}]) \simeq T_e([\text{O III}])$ as indicated by photoionization models, yielding a direct measure of $12+\log(\text{O}/\text{H})$ assuming that $(\text{O}/\text{H}) = (\text{O}^{++}/\text{H}^+) + (\text{O}^+/\text{H}^+)$. The results are summarized in Table 4. We have assumed zero nebular extinction as in Erb et al. (2010), supported by both the Balmer decrement and the SED fitting results, and find $12+\log(\text{O}/\text{H}) = 8.08 \pm 0.05$, $\simeq 0.3$ dex higher than that obtained by Erb et al. (2010). The difference is attributable to a lower derived T_e driven by a larger $[\text{O III}]\lambda 5008$ flux from the new H-band spectrum, as well as the detection of the $[\text{O II}]\lambda\lambda 3727, 3729$ doublet, which allowed the contribution to (O/H) from O^+ to be included. We also note that the measured $\text{O32} \simeq 10$ is in excellent agreement with the photoionization models that repro-

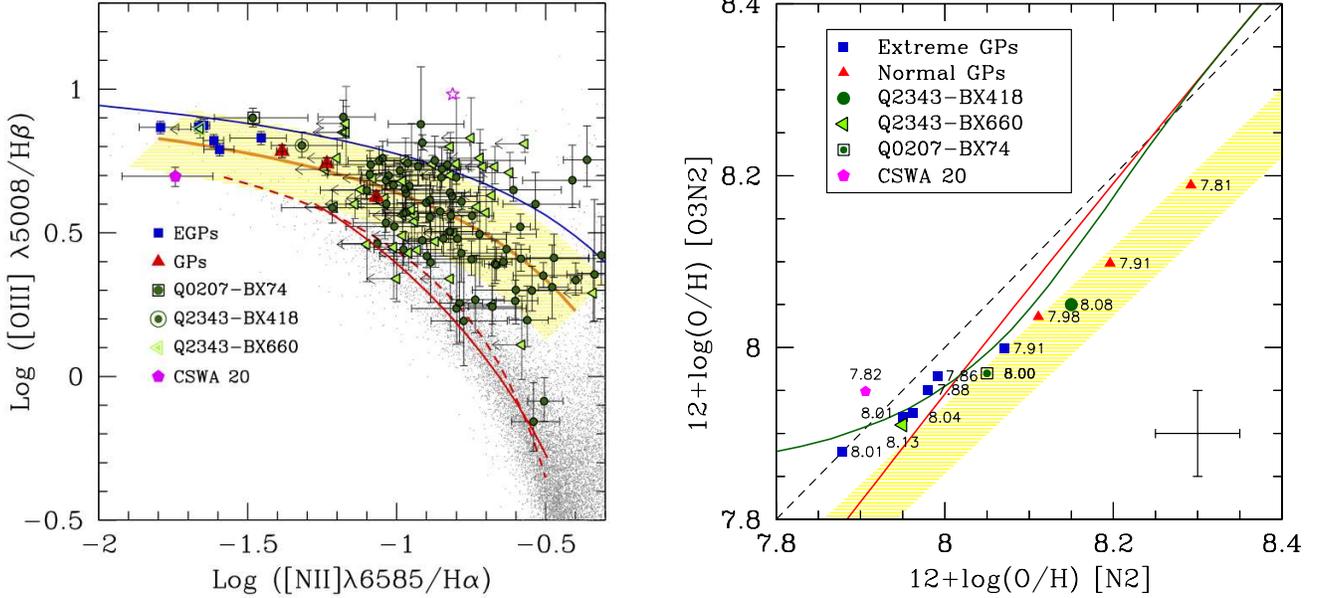


Figure 19. (Left) Plot analogous to Figure 8, but adding 9 $z \simeq 0.2$ “green pea” galaxies (EGP: dark blue squares; GP: red triangles). (Right) Analogous to Figures 6 and 9, showing the metallicities that would be inferred using the N2 and O3N2 indices for the GPs (blue squares), EGPs (red triangles), the 3 $z \sim 2$ galaxies with direct metallicity measurements (this work; green symbols), and CSWA 20, a $z = 1.4$ lensed galaxy (James et al. 2014; magenta pentagon.) For each point, the value of $12+\text{log}(\text{O}/\text{H})$ measured from the direct method is indicated (see text for discussion, and Figure 21. The errorbar in the righthand panel show the typical uncertainties on the T_e -based metallicity determinations; the formal uncertainties for the N2- and O3N2-based determinations are similar or smaller.)

duce BX418’s position on the BPT diagram (see section 4).

We performed similar analyses for Q2343-BX660 and Q0207-BX74. As summarized in Table 4, we infer high values of $n_e = 300 - 1600 \text{ cm}^{-3}$, very high ionization level (based on O32), and direct oxygen abundances $12+\text{log}(\text{O}/\text{H}) \simeq 8.00 - 8.15$. Once again the measurement of $T_e[\text{OIII}]$ is based on the rest-UV OIII] doublet strength relative to $[\text{OIII}]\lambda 5008$, and the oxygen abundance includes the contribution from O^+ ; the Keck/LRIS-B spectra used to measure the UV features are shown in Figure 20.

The disadvantage of using the UV $[\text{O III}]\lambda 4364$ to measure T_e is that it is much more sensitive to the nebular extinction correction. Fortunately, two of the 3 galaxies for which the measurements are available (BX418 and BX660) are consistent with zero nebular extinction based on the observed $\text{H}\alpha/\text{H}\beta$ ratio (the “Balmer decrement”), which are each consistent with the “Case B” expectation of $\text{H}\alpha/\text{H}\beta = 2.86$; both are also consistent with zero extinction based on their SED fits.

For Q0207-BX74, based on the observed Balmer decrement, we obtain $E(B-V)_{\text{neb}} = 0.18$, while the stellar continuum (from SED fitting) has $E(B-V)_{\text{cont}} = 0.13$ assuming the Calzetti et al. (2000) attenuation relation. We corrected the relevant line fluxes in Table 4 assuming the former and the Cardelli et al. (1989) extinction curve. Note that effect of the dust correction to the observed ratio $\text{OIII}](\lambda 1661 + \lambda 1666)/[\text{OIII}]\lambda 5008$ was to increase it by a factor of 2.04, increasing the inferred T_e from $\sim 12140 \text{ K}$ to $\sim 14300 \text{ K}$ and lowering the inferred oxygen abundance by ~ 0.23 dex, from 8.23 to 8.00.

Figure 21 summarizes the comparison of the direct metallicity estimates for the same 13 galaxies as in Figure 19. In addition to the N2 and O3N2-based estimates, we have applied the low-metallicity branch of the R23 calibration of

McGaugh (1991) [as expressed by Kobulnicky et al. 1999] to the measurements of O32 and $([\text{OIII}]_{\text{tot}} + [\text{OII}]_{\text{tot}})/\text{H}\beta$, to estimate R23-based metallicities, shown in the rightmost panel of Figure 21. Figure 21 suggests that, at least for this sample, O3N2 provides a slightly better approximation to the direct method metallicities, with a relative offset of $\text{log}(\text{O}/\text{H})_{\text{O3N2}} - \text{log}(\text{O}/\text{H})_{\text{dir}} = 0.00 \pm 0.11$ dex, while $\text{log}(\text{O}/\text{H})_{\text{N2}} - \text{log}(\text{O}/\text{H})_{\text{dir}} = 0.04 \pm 0.14$ dex and $\text{log}(\text{O}/\text{H})_{\text{R23}} - \text{log}(\text{O}/\text{H})_{\text{dir}} = 0.16 \pm 0.08$ dex. We caution that these statistics are based on a very small sample, confined to low metallicities where the various indicators appear to be in reasonable agreement with one another; however, the finding that the O3N2 index provides the least-biased estimate of direct-method metallicities for galaxies offset from the BPT excitation sequence is consistent with results presented by Liu et al. (2008) for $z \simeq 0$ SDSS galaxies; these authors found that N2 systematically over-estimates $12+\text{log}(\text{O}/\text{H})$ compared to the direct method.

On balance, it seems most likely that the O3N2 index yields more reliable values of $12+\text{log}(\text{O}/\text{H})$ than those derived from the N2 index; however, the current KBSS-MOSFIRE sample with N2 measurements is a factor of more than two larger than that with O3N2 measurements. In section 7 below, we use the relative calibrations between the two metallicity estimates (which has a small scatter— see Figure 6) to “convert” the N2-only measurements onto the O3N2 metallicity scale in the context of discussing the relationship between stellar mass, oxygen abundances, and star formation rates at $z \sim 2.3$.

7. THE M_* -METALLICITY RELATION AT $\langle z \rangle = 2.3$

A correlation between stellar mass M_* and nebular oxygen abundance has now been well-established at $z \simeq 0$ using both strong-line metallicity measurements (e.g., Tremonti et al. 2004; Kewley & Ellison 2008) as well as direct T_e -based measures (Andrews & Martini 2013). However, even in the lo-

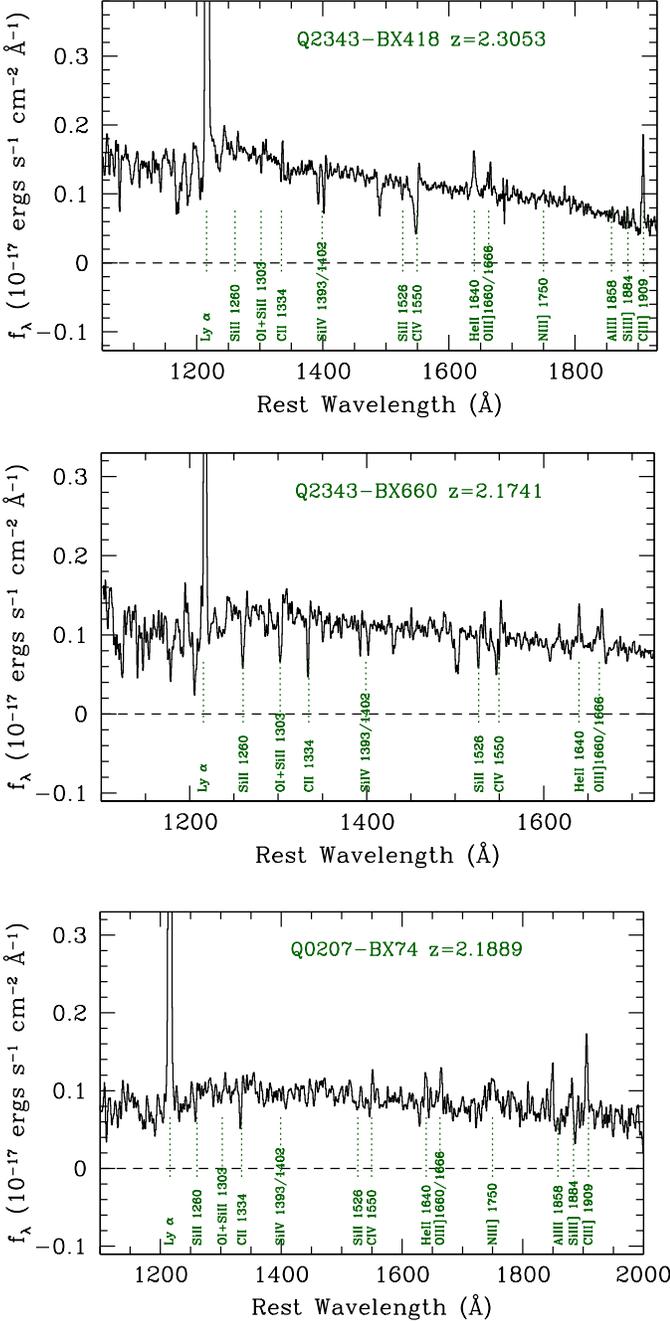


Figure 20. Rest-UV Keck/LRIS-B spectra of (top) Q2343-BX418 (see Erb et al. 2010), (middle) Q2343-BX660 and (bottom) Q0207-BX74. Note the presence of unusually strong lines of OIII] $\lambda\lambda$ 1661, 1666 and the [CIII]+CIII] blend near 1908 Å.

cal universe the quantitative behavior of the mass-metallicity relation (MZR) depends substantially on the method used to measure $12+\log(\text{O}/\text{H})$. Apparently similar behavior, with substantial offsets in the sense that galaxies are inferred to have lower ionized gas metallicities at a given M_* , has been observed for relatively small samples of $z > 2$ galaxies (Erb et al. 2006a; Maiolino et al. 2008; Law et al. 2009; Förster Schreiber et al. 2009; Newman et al. 2013; Kewley et al. 2013b). As we have seen above, we do not yet know whether the strong-line determinations of metallicity at high redshift are directly comparable to those obtained at low redshift,

even using the same diagnostic. We showed in the previous section that there is evidence, albeit limited, (see Figure 21 and Liu et al. 2008) that the O3N2 calibration is least biased with respect to direct-method metallicities for galaxies falling “above” the low-redshift BPT ionization sequence.

Thus, in addition to the N2-based MZR we have assembled a similar MZR based on O3N2. Where O3N2 measurements exist (79 galaxies, from Table 1), we use them directly (light green points in the lefthand panel of Figure 23). Since the KBSS-MOSFIRE galaxies having both O3N2- and N2-based metallicities fall along a well-defined sequence with small intrinsic scatter (Figure 6), we converted galaxies with only N2 metallicities (53 galaxies; Table 3) (or upper limits; 42 galaxies; Table 2) into their corresponding O3N2-based values using the linear regression noted in Figure 6:

$$12+\log(\text{O}/\text{H})_{\text{O3N2}} = 0.92 [12+\log(\text{O}/\text{H})_{\text{N2}}] + 0.54 \quad (16)$$

where the values of $12+\log(\text{O}/\text{H})$ are those based on the O3N2 and N2 calibrations of PP04. Figures 22 and 23 show both versions of the resulting KBSS-MOSFIRE $z \sim 2.3$ MZR, as well as previous results from Erb et al. (2006a)²⁸. The latter measurements were based on spectral stacks in bins of stellar mass. The righthand panels of Figure 22 and 23 are identical to the corresponding lefthand panels, except that the KBSS-MOSFIRE data are binned by M_* for clarity; the binned data points are collected in Table 5 for convenience.

The new KBSS-MOSFIRE sample in Figure 22 includes a total of 174 galaxies, of which 132 have N2 measurements from Tables 1 and 3, and 42 have measurements of $\log([\text{OIII}]/\text{H}\beta)$ along with 2σ upper limits on N2 (Table 2). The long-dashed (green) line in each panel of Figures 22 and 23 are the best linear fits to the full data set, where the 2σ upper limits on $12+\log(\text{O}/\text{H})$ were treated as detections, with the following parameters:

$$12+\log(\text{O}/\text{H})_{\text{N2}} = 8.44 + 0.12 [\log(M_*/M_\odot) - 10]; \quad (17)$$

$$\sigma = 0.13 \text{ dex}$$

and

$$12+\log(\text{O}/\text{H})_{\text{O3N2}} = 8.30 + 0.12 [\log(M_*/M_\odot) - 10]; \quad (18)$$

$$\sigma = 0.11 \text{ dex}$$

where the σ values are the weighted errors between the data points and the fits. Fits including only the detections and not the upper limits are statistically indistinguishable.

We also tried solutions similar to those advocated by Moustakas et al. (2011) (see also Andrews & Martini 2013), of the form

$$12+\log(\text{O}/\text{H}) = A - \log \left[1 + \left(\frac{M_\Gamma}{M_*} \right)^\gamma \right] \quad (19)$$

where A is the asymptotic value of $12+\log(\text{O}/\text{H})$ at high M_* , M_Γ is a characteristic stellar mass, and γ is the low-mass slope of the MZR. The best fits with no priors on the 3 parameters essentially reproduce the linear functions of equations 17 and 18. If we fix $A = 8.80$, the asymptotic metallicity of both low- z MZRs shown in Figure 22, the best fit parameters are

$$\text{N2: } A = 8.80 \quad M_\Gamma = 10.20; \quad \gamma = 0.33 \quad (20)$$

$$\sigma = 0.13 \text{ dex}$$

²⁸ The values of $12+\log(\text{O}/\text{H})_{\text{N2}}$ from Erb et al. (2006a) were converted to $12+\log(\text{O}/\text{H})_{\text{O3N2}}$ using equation 16

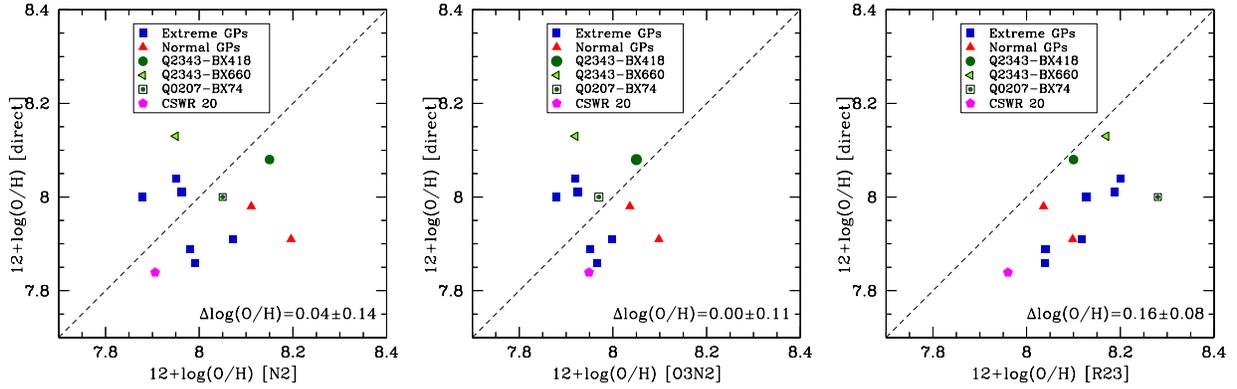


Figure 21. Comparison between direct T_e measurements of $12+\log(\text{O}/\text{H})$ and values suggested by three different strong-line indicators for the same set of galaxies shown in Figure 19b. Left: N2, Center: O3N2, Right: R23, using the lower-branch calibration of McGaugh (1991) (see text, and Table 4, for details). The mean and rms scatter in the difference between individual measurements of $12+\log(\text{O}/\text{H})$ and the direct method measurement is given in the lower right of each panel.

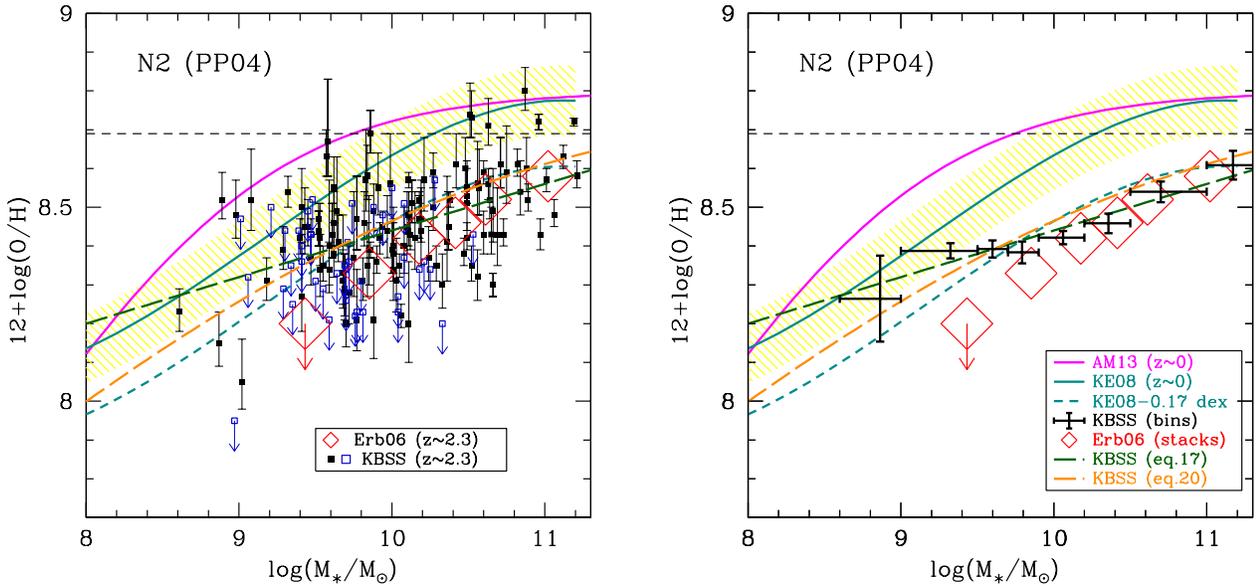


Figure 22. Observed relation between stellar mass (M_*) and oxygen abundance inferred from the PP04 N2 index calibration, for $z \sim 2.3$ KBSS-MOSFIRE galaxies. The sample includes 132 galaxies with $[\text{NII}]/\text{H}\alpha$ measurements (black points) and 42 with 2σ upper limits on $[\text{NII}]/\text{H}\alpha$ (blue open squares, with downward arrows). The open diamonds are from stacked spectra in bins of M_* from Erb et al. (2006a), also with $z \sim 2.3$. The long-dashed green line is the best-fit linear relation between $12+\log(\text{O}/\text{H})_{\text{N2}}$ and $\log(M_*/M_\odot)$ (equation 18), and the dashed orange curve is the best fit of the form given in equation 19, with parameters in equation 20 (see text for discussion). The solid turquoise curve and light shading (representing the approximate scatter in the SDSS sample) is the robust polynomial representation of the best-fit MZR for SDSS galaxies using the PP04 N2 calibration, as parametrized by Kewley & Ellison (2008). The short-dashed turquoise curve is the same MZR, shifted to lower metallicity by 0.17 dex. The solid magenta curve is the best-fit MZR from Andrews & Martini (2013) (see equations 19 and 22), where metallicities were determined using the direct method based on stacked SDSS spectra in bins of stellar mass. (Right) Same as left panel, but with data points binned by stellar mass, with values indicated by error bars. In the y-direction, the error bars reflect uncertainty in the bi-weight mean within each bin, with the x-direction error bars indicating the limits of the M_* bin. The x-location of each point is determined by the median M_* within the bin. Table 5 summarizes the data shown. Note that the fitted curves are the same as in the lefthand panel, based on the fits to the full sample of individual data points as described in the text.

and

$$\text{O3N2} : A = 8.80 \quad M_T = 11.35 ; \quad \gamma = 0.24 \quad (21)$$

$$\sigma = 0.11 \text{ dex}$$

The non-linear fits are shown as dashed orange curves in Figures 22 and 23. For comparison, the best fit MZR parameters at low redshift for direct method metallicities (Andrews & Martini 2013) are

$$\text{Direct} : A = 8.80 ; \quad M_T = 8.90 ; \quad \gamma = 0.64 ; \quad (22)$$

these parameters produce the continuous magenta curves in Figures 22 and 23. Also shown for comparison in Figure 22 and 23 are the $z \simeq 0$ MZR of Maiolino et al. (2008), converted, respectively, to the PP04 N2 and O3N2 calibrations (solid turquoise curves with light (yellow) shading indicating the rms scatter).

Erb et al. (2006a) first showed, based on composite spectra formed from bins of M_* (large open diamonds in Figure 22), that the $z \sim 2.3$ MZR lies substantially below the $z \simeq 0$ relation. The amplitude of the shift in metallicity depends on

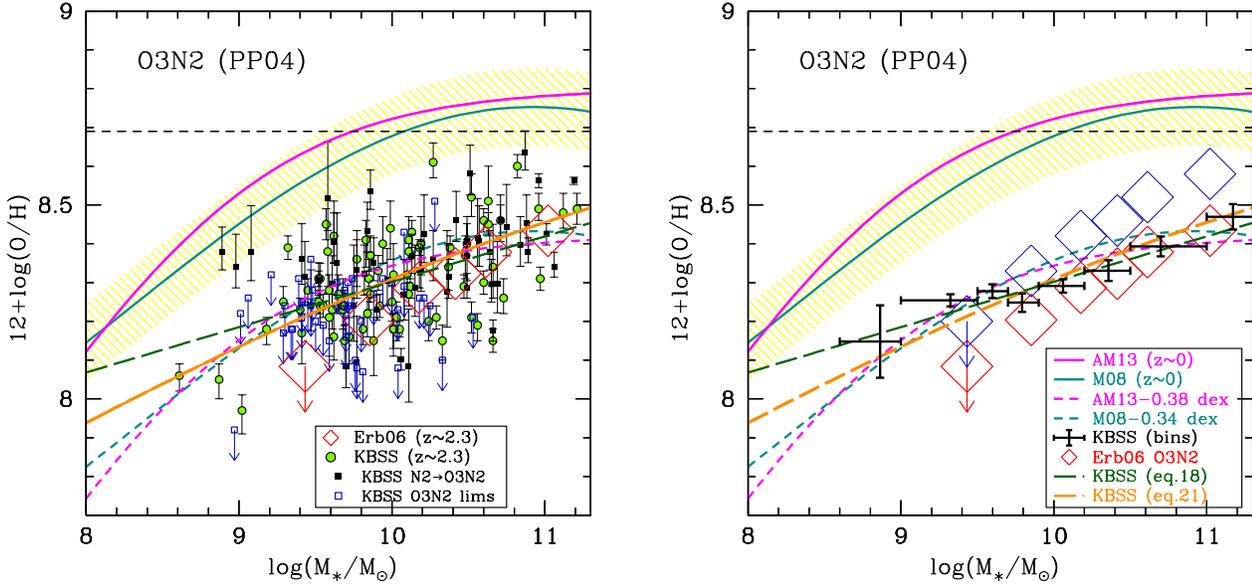


Figure 23. (Left) Same as Figure 22, but for metallicities inferred from the PP04 O3N2 calibration. Galaxies with individual O3N2 measurements are indicated with light green solid dots, and equation 16 was used to convert galaxies with N2 but lacking O3N2 measurements (solid square dots with error bars). Blue squares represent N2-based limits as in Figure 22, converted to O3N2-based limits using equation 16. The scatter in the MZR is slightly smaller when the O3N2 values are used (see also equations 17 through 21). The solid turquoise curve and yellow shaded region represents the $z \sim 0$ MZR of Maiolino et al. (2008), with metallicities converted to the PP04 O3N2 calibration using the equations given in those two papers; the short-dashed turquoise curve is the same MZR shift to lower metallicity by 0.34 dex. Similarly, the short-dashed magenta curve is the Andrews & Martini (2013) MZR, shifted to lower metallicity by 0.38 dex. The red diamonds are from ?Right) As in Figure 22, where individual data points in the lefthand panel have been binned for clarity (see Table 5). The red diamonds are the points from Erb et al. (2006a), converted to O3N2-based metallicities using equation 16; the large blue diamonds are the same N2-based points as in Figure 22, shown for comparison.).

the method used to measure it – Erb et al. (2006a) found that the shift of the $z \sim 2.3$ metallicities (measured using N2) relative to the MZR of Tremonti et al. (2004) was -0.56 dex, but decreased to $\simeq -0.3$ dex when the PP04 N2 calibration was applied to the local sample. Using the PP04 O3N2 calibration (as in Figure 23), the offset between the $z \sim 0$ O3N2-based MZR and the KBSS-MOSFIRE sample is in good agreement (discussed further below), $\Delta \log(O/H) \simeq -0.34$ dex; however, we find that a smaller offset applies to the N2-based MZR, $\Delta 12 + \log(O/H)_{N2} \simeq -0.17$ dex, due primarily to higher inferred metallicity at a given M_* for N2-based metallicities in the new sample.

It appears from Figure 22 that the KBSS-MOSFIRE MZR exhibits a shallower dependence of the N2 index on M_* than the Erb et al. (2006a) sample (both used the PP04 N2 calibration to convert to metallicity). The new MZR results at $z \sim 2.3$ are almost entirely independent of the Erb et al. (2006a) sample, in the sense that all of the nebular line measurements are based on new observations with MOSFIRE, and only 10 of 174 galaxies in the new sample are in common with the Erb et al. (2006a) NIRSPEC sample. Referring to Figure 22, at high M_* the locus of individual galaxies from KBSS-MOSFIRE agrees well with the result from the stacked spectra of Erb et al. (2006a), while for $\log(M_*/M_\odot) \lesssim 10.5$, the Erb et al. (2006a) results are systematically lower in $12 + \log(O/H)_{N2}$ than the corresponding locus of individual measurements.

In order to better assess the significance of the difference, we also constructed composite spectra in similar bins of M_* (to be discussed in detail elsewhere; Strom et al., in preparation) for the KBSS-MOSFIRE sample. We find that a linear fit to the N2-based MZR derived from composite spectra

within each bin of M_* results in a slope and normalization indistinguishable from the least-squares fit to the individually-measured galaxies (equation 17). Thus, the origin of the difference is not related to binning/stacking, nor to the inclusion of spectra with individual N2 upper limits in the stack. Aside from pure sample variance, some part of the discrepancy might be explained by very different spectral resolution and S/N (both are considerably higher for the MOSFIRE spectra), when one accounts for the fact that weak emission lines are harder to distinguish from the continuum level, so that systematic errors in the zero level of the spectra can have a large effect on inferred line strength near the detection limit. In any case, the relation between the strong-line metallicity– using either the N2 or O3N2 indices– and $\log(M_*/M_\odot)$ is quite shallow over the range in M_* spanned by the KBSS-MOSFIRE sample, with best-fit linear slope of $\gamma = 0.12$. In view of the $z \sim 0$ MZRs, this is not surprising– over the same range in stellar mass ($9 \lesssim \log(M_*/M_\odot) \lesssim 11$) the observed slopes are similarly shallow. For example, a linear approximation to the Andrews & Martini (2013) MZR for $9 < \log(M_*/M_\odot) < 11$ has $\gamma = 0.13$, while the Maiolino et al. (2008) O3N2 MZR curve shown in Figure 23 has $\gamma = 0.15$ over the same mass range.

Focusing on the O3N2 MZR for definiteness (and because it appears to be closest to the T_e /direct metallicity scale at $z \sim 2.3$ and at $z \sim 0$ as discussed above; see also Figure 23), we find that the two functional forms for the $z \sim 0$ MZR shown provide reasonable fits to the observed $z \sim 2.3$ data after adjusting the overall metallicity normalization. The dashed magenta and turquoise curves in Figure 23 show the Andrews & Martini (2013) and Maiolino et al. (2008) MZR fits with $\Delta(12 + \log(O/H)) = -0.38$ and -0.34 dex, respectively. With

these adjustments, either curve provides a fit of similar quality (i.e., $\sigma \simeq 0.12$ dex) to those in equations 18 and 21 above. Thus, we find that over the current stellar mass range probed, the MZR at $z \sim 2.3$ may be reasonably well-described by the $z \sim 0$ MZR with a $\simeq 0.36$ dex reduction in oxygen abundance, *independent of stellar mass*. Qualitatively, the same conclusion was reached by Erb et al. (2006a).

7.1. Scatter in the $z \sim 2.3$ MZR

As for the BPT locus discussed in section 3 above, we estimate the intrinsic scatter in the $z \simeq 2.3$ MZR by comparing the RMS scatter σ_{tot} between the fits and the data in Figures 22 and 23 with that expected from measurement errors ($\sigma_{m,i}$) alone. We solve for σ_{sc} (assumed to be constant) such that $\sigma_{\text{tot},i}^2 = \sigma_{m,i}^2 + \sigma_{\text{sc}}^2$, where the subscripts i refer to each individual data point, such that $\chi^2/\nu \approx 1$. For the O3N2-based MZR (Figure 23), we infer intrinsic scatter of $\sigma_{\text{sc}} = 0.09$ dex. Dividing the galaxy sample in half near the median $\log(M_*/M_\odot) = 10.0$ and estimating σ_{sc} separately for each sub-sample, $\sigma_{\text{sc}}[\log(M_*/M_\odot) > 10] = 0.09$ dex and $\sigma_{\text{sc}}[\log(M_*/M_\odot) < 10] = 0.08$ dex, suggesting that there is no significant difference in intrinsic scatter of O3N2-inferred oxygen abundance versus M_* over the range probed thus far by KBSS-MOSFIRE.

We note that the scatter in the $z \sim 2.3$ MZR, $\sigma_{\text{sc}} \simeq 0.09$ dex, is remarkably similar to that measured for galaxies over the same range of stellar mass in the low redshift samples when the same strong-line index is used to derive the metallicities ($\sigma_{\text{sc}} = 0.08 - 0.10$ dex; Kewley & Ellison 2008.)

An obvious point, relevant at both $z \sim 0$ and $z \sim 2.3$, is that the scatter in the inferred metallicity at a given stellar mass is *smaller* than could be reasonably expected even if the “true” oxygen abundance were a perfect monotonic function of M_* . The scatter in the empirical strong-line metallicity calibration, estimated by PP04 to be $\simeq 0.18$ dex for N2 and $\simeq 0.14$ dex for O3N2, both *exceed* the inferred intrinsic MZR scatter of $\simeq 0.09$ dex. We showed in section 4.4 that the calibration errors of the N2 and O3N2 methods could be reduced compared to the numbers given by PP04 by restricting the range of line index included in the linear fit. *However, even the reduced calibration uncertainties would still account for 100% of the observed scatter in the MZR even if M_* were perfectly correlated with oxygen abundance.* Taken at face value, this suggests that the relative intensities of the strong emission lines, which we have argued are modulated primarily by ionization parameter and by the hardness of the UV radiation field, are *more strongly* correlated with M_* than is the oxygen abundance. We will return to the possible implications of this “ M_* -excitation” relation in section 8 below.

7.2. The MZR-SFR Relation at $z \simeq 2.3$

At $z \simeq 0$, the scatter in the correlation between M_* and metallicity can be reduced significantly by including an additional parameter that accounts indirectly for the cold gas content of the galaxies, most commonly using the SFR. A parametrization of this dependence of the form

$$12 + \log(\text{O}/\text{H}) \propto \mu_* \equiv \log M_* - \alpha \log(\text{SFR}/M_\odot \text{yr}^{-1}), \quad (23)$$

where α is a constant that minimizes the scatter in metallicity at a given μ_* , was introduced by Mannucci et al. (2010) as a convenient “projection” of what they called the “fundamental metallicity relation” (FMR). According to Mannucci et al. (2010), the FMR is a thin two-dimensional surface in

the space defined by M_* , Z , and SFR, upon which all star-forming galaxies lie, independent of redshift for $z \lesssim 2.5$. To first order, the projection of the FMR parametrized by equation 23 accounts for the clearly observed trend (at $z \simeq 0$) that galaxies with higher SFR have lower gas-phase oxygen abundances at fixed M_* . In the context of the FMR, high redshift galaxies, known to have much higher gas fractions and SFRs than most local star-forming galaxies (e.g., Erb et al. 2006c; Tacconi et al. 2010, 2013), would also be expected to have correspondingly lower (O/H) at a given M_* . Thus, the value of α in equation 23 adjusts the actual M_* to the mass μ_* expected for a galaxy with $\log(\text{SFR}/M_\odot \text{yr}^{-1}) = 0$ and the observed metallicity. For this form of the projected FMR, Mannucci et al. (2010) found that $\alpha = 0.32$ minimized the scatter in their $z \sim 0$ sample. An even stronger dependence on SFR of the M_* - Z relation has been suggested by Andrews & Martini (2013), who found $\alpha = 0.66$ for a local galaxy sample whose oxygen abundances were determined using the “direct” method.

However, the typical $z \sim 2.3$ galaxy in our sample has SFR $\simeq 25 M_\odot \text{yr}^{-1}$ (Figure 4), well beyond the range of SFR well-sampled by the $z \sim 0$ data set used by Mannucci et al. (2010), and near the high SFR extreme of the $z \simeq 0$ sample used by Andrews & Martini (2013). It is an interesting test of the universality of the FMR to see how the high redshift sample compares. Figure 24 shows the MZR-SFR relationship for $\alpha = 0.32$ and $\alpha = 0.66$ (left and right panels, respectively). *It appears that neither value of α brings the $z \sim 2.3$ MZR-SFR relation into agreement with those established at $z \simeq 0$; moreover, neither value of α reduces the scatter of the high-redshift MZR nor aligns the $z \sim 2.3$ data with the $z \sim 0$ FMR expectations.*

Alternatively, one can use the full form of the FMR, which predicts the value of $12 + \log(\text{O}/\text{H})$ given observed values of M_* and SFR. Figure 25 shows the *residuals* between the observed metallicity and that predicted by the Mannucci et al. (2010) FMR (left panel), as well as the updated FMR suggested by Mannucci et al. (2011), which alters the FMR shape for $\mu_* < 9.5$ (right panel). Figure 25 shows that both versions of the FMR, when applied to the $z \sim 2.3$ sample, result in significant M_* -dependent residuals. Qualitatively, the FMR calibrations systematically *under-predict* metallicity at low M_* and systematically *over-predict* it at high M_* ²⁹. Moreover, ignoring the M_* -dependent systematic differences, the scatter in the (M2011) FMR residuals at a given M_* , $\sigma_{\text{res}} = 0.12$ dex, is not significantly smaller than for the $z \sim 2.3$ MZR in Figure 22, where no account has been taken of the SFR.

It is possible that systematic differences in methods or measurements used for the estimation of SFRs between the low-redshift and high-redshift samples could be responsible in part for systematic deviations from the local MZR-SFR relation. However, within the high redshift sample we find that there is no MZR-SFR solution that significantly reduces the scatter in inferred metallicity at a given M_* .

Thus, *assuming that local calibrations of N2 or O3N2 strong-line indices and oxygen abundances are applicable at high redshift*, we conclude that the relationship between SFR, metallicity, and M_* at $z \sim 2.3$ is significantly different from that which applies to galaxies in the same range of stellar mass

²⁹ Very recently, Cullen et al. (2014) have used spectroscopic stacks in bins of stellar mass for a sample of 93 emission-line-selected galaxies at $z \sim 2$ from 3D-HST, and also finding significant residuals with respect to the locally-calibrated FMR.

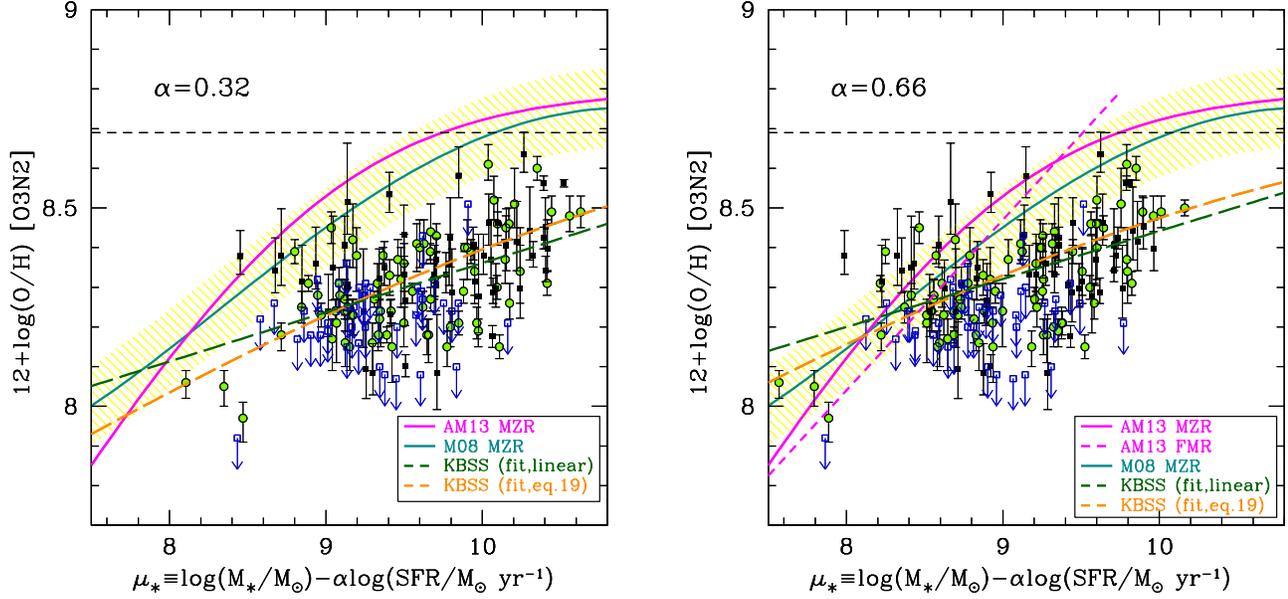


Figure 24. The MZR-SFR relation at $\langle z \rangle \approx 2.3$. Data points are as in Figure 23, with the adjustment from M_* to μ_* . (Left:) The projection of the FMR suggested by Mannucci et al. (2010) to minimize the scatter in metallicity (determined from calibrations given by Maiolino et al. (2008)) at a given $\mu_* = \log(M_*/M_\odot) - \alpha \log(\text{SFR})$, with $\alpha = 0.32$. SFR is expressed in solar units. The same $z \sim 0$ MZR fits as in Figure 23 are also given for comparison. (Right:) As in the left-hand panel, but for $\alpha = 0.66$, found by Andrews & Martini (2013) to minimize the scatter in direct oxygen abundance at a given value of μ . As in Figure 23, the dashed green and orange curves are best fits to the data points of the form given in equations 18 and 19. The dashed magenta line is the linear FMR fit given by Andrews & Martini (2013), which has $12 + \log(\text{O}/\text{H}) \propto 0.43 \log \mu_*$ assuming $\alpha = 0.66$.

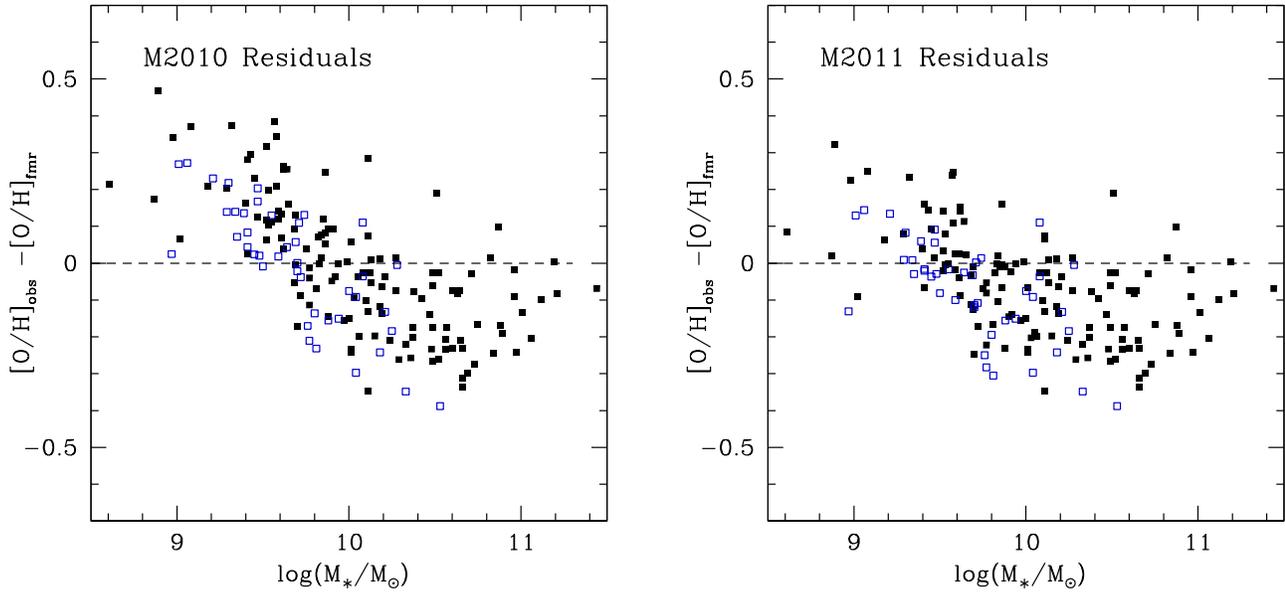


Figure 25. (Left) Residuals of the Mannucci et al. (2010) FMR metallicity predictions with respect to metallicity calculated from the O3N2 line index for the $z \sim 2.3$ KBSS-MOSFIRE sample as in Figure 23. Open blue squares have the same meaning as in Figure 23. (Right) As in the left-hand panel, using the Mannucci et al. (2011) FMR, which adjusts the calibration for low stellar masses. The revised calibration reduces the systematic under-prediction of metallicity at low M_* (see also Belli et al. 2013).

at $z \sim 0$. *It is not obvious within the current $z \sim 2.3$ sample that SFR and metallicity are strongly linked*; as discussed in the previous section, the current MZR data are more consistent with a “vertical” shift in metallicity by ~ -0.36 dex in $12+\log(\text{O}/\text{H})$, rather than the more “horizontal” shift expected by current $z \sim 0$ FMR calibrations.

8. SUMMARY AND DISCUSSION

We have presented near-IR spectroscopy for an initial sample of 179 star-forming galaxies with $2.0 \leq z \leq 2.6$ observed in the 15 fields of the Keck Baryonic Structure Survey. All spectra were obtained using MOSFIRE, the recently-commissioned near-IR multi-object spectrometer on the Keck 1 10m telescope, during the first year of its operation. In addition to the large size of the galaxy sample, the quality of the spectra of individual galaxies is much higher, and the dynamic range within the sample much larger, than has been possible to achieve previously. In this paper, we have explored the quantitative use of the strong nebular emission lines in the rest-frame optical spectra of high-redshift galaxies, re-examining their utility for diagnosing the physical conditions in galaxies during the peak of their most active star-forming phase. The main conclusions are as follows:

1. At $z \sim 2.3$, galaxies occupy an almost entirely distinct, but similarly tight, locus in the BPT diagram compared to the vast majority of star-forming galaxies in the local universe (Figure 5). The shift in the observed locus can be qualitatively explained if essentially all high-redshift H II regions are characterized by both harder ionizing radiation fields and higher ionization parameters than apply for all but the most extreme local galaxies.

2. Since all strong-line metallicity indicators and their calibrations are “tuned” to reproduce the tight sequence in the BPT diagram for local galaxies, the shift of the high redshift locus means that the same calibrations among the strong-line indicators cannot be used at high redshift without introducing systematics in the metallicity scale. Since ground-based observations are confined to redshift intervals within which particular strong nebular lines fall in the near-IR atmospheric windows, galaxy samples at different redshifts will necessarily depend on different subsets of the strong lines. It is entirely possible, perhaps even likely, that calibration issues could mimic global changes in metallicity or other physical conditions in HII regions with redshift. As an example, we show that metallicities inferred from the N2 and O3N2 indices of $z \sim 2.3$ galaxies differ systematically from each other, with an offset that increases with increasing inferred metallicity, but averages $\simeq 0.15$ dex, in the sense that N2-inferred metallicities are higher (Figure 6).

3. Using simple photoionization models (with minimal assumptions about the details of the ionizing sources) we find that the observed locus of $z \sim 2.3$ galaxies in the BPT diagram, as well as the behavior of the N2 and O3N2 indices with respect to one another, can be reproduced remarkably well if the shape of the net ionizing radiation field in high-redshift HII regions resembles a blackbody with effective temperature $T_{\text{eff}} = 50000 \pm 5000$ K and ionization parameter in the range $-2.9 \lesssim \log \Gamma \lesssim -1.8$ (Figures 8 and 9.) In the context of the models, most of the variation along the principle axis of the BPT locus is produced by changes in Γ , while the overall normalization of $[\text{OIII}]/\text{H}\beta$ is modulated primarily by the effective temperature of the ionizing radiation field. For the high inferred level of ionization, a galaxy’s position in BPT space is nearly independent of the ionized gas metallicity over the

range $0.2 \leq Z/Z_{\odot} \leq 1.0$ —so that any observed metallicity dependence of the strong-line ratios is more likely caused by correlations between the radiation field shape and intensity with the metallicity of the stars themselves. In addition, we find the $z = 2.3$ BPT locus is most easily reproduced if the (N/O) ratio in the ionized gas is close to the solar ratio over the full observed range of (O/H). Such high N/O, as well as high T_{eff} , may both be a consequence of the effects of binaries and rapid rotation on massive main sequence stars. Such effects are predicted to be greatly enhanced at sub-solar metallicities that appear to be the rule at high redshift.

4. The KBSS-MOSFIRE sample contains a small number of AGN (Figures 16 and 17), most of which had been previously identified based on emission lines of high ionization species in their rest-frame UV spectra. The positions of AGN on the BPT diagram (Figure 5) appear distinct from the vast majority of objects which show no evidence in their rest-UV, rest-optical, or other multi-wavelength measurements for energetically significant contamination by AGN. The highest-excitation star-forming galaxies in the KBSS-MOSFIRE sample exhibit a maximum $\log([\text{OIII}]\lambda 5008/\text{H}\beta) \simeq 0.9$, consistent with the predictions of the photoionization models with UV ionizing radiation field in the 1-4 Ryd range resembling a blackbody with $T_{\text{eff}} = 55000 - 60000$ K and $\log \Gamma \gtrsim -2.0$ (Figure 8). This upper envelope appears to be the same for the most extreme star-forming galaxies in the local universe, where they are many orders of magnitude rarer.

5. We have drawn attention to the similarities between the most extreme galaxies (in terms of their position on the BPT diagram) in the $z \sim 0$ and $z \sim 2.3$ samples. In particular, the so-called “green pea” galaxies at $z \simeq 0.2$ appear to have strong line ratios placing them directly on the $z \sim 2.3$ BPT locus, while the “extreme green peas” are coincident with the highest excitation galaxies observed in the $z \simeq 2.3$ sample. Comparison of the published samples of green peas having accurate direct (T_e) metallicity measurements with the $z \sim 2.3$ galaxies is also interesting. The strong-line metallicity indices of the GPs follow the same trend as observed among the $z \sim 2.3$ sample and as predicted by our photoionization models (cf. Figure 6, 9, and 19). The corresponding direct measures of oxygen abundances for the GPs and a small subset of the $z \sim 2.3$ sample suggest that among the commonly-applied strong-line calibrations, the least-biased with respect to the direct (T_e) metallicity measurements is O3N2. The differences in N2- and O3N2-based oxygen abundances described above imply that N2 generally over-estimates metallicities at $z \sim 2.3$ (by $\sim 0.09 - 0.18$ dex for $0.2 - 1.0 Z_{\odot}$). The systematic differences can be reduced considerably (but not entirely eliminated) by restricting the low-redshift calibration data sets to the range of line indices observed among the high redshift sample (Figure 14.) We propose a simple empirical relation for converting $12+\log(\text{O}/\text{H})_{\text{N2}}$ to the corresponding O3N2-based value appropriate at $z \sim 2.3$.

6. As shown previously using stacked spectra (Erb et al. 2006a), there is a relationship between M_* and the strong-line indices (N2 or O3N2) in place at $z \sim 2.3$ qualitatively similar to those observed at $z \simeq 0$ (Figures 22 and 23). If one converts the observed line indices into oxygen abundances using the locally-established calibrations (i.e., under the assumption that the line indices can be used to measure metallicity), the best-fit $z \sim 2.3$ MZR is shallower than most previous studies have suggested, $12+\log(\text{O}/\text{H}) \simeq 8.30 + 0.12 [\log(M_*/M_{\odot}) - 10]$. (Figures 22 and 23.) However, we show that the $z \simeq 0$ MZR has slopes

consistent with that of the $z \sim 2.3$ sample when evaluated over the same range of M_* and using the same (O3N2) metallicity calibration. Further, we find that the $z \sim 2.3$ MZR is also well-fit by the $z \sim 0$ MZRs with metallicity normalization shifted lower by $\Delta \log Z = 0.34 - 0.38$ dex, *independent of stellar mass*. As for the locus in the BPT diagram, the intrinsic scatter in the MZR (i.e., scatter of inferred metallicity at a given M_*) is both small and remarkably similar at $z \sim 2.3$ and $z \sim 0$ when a consistent metallicity calibration is applied to both: $\sigma_{\text{sc}} \simeq 0.09$ dex. Over the well-covered range of M_* observed in the current KBSS-MOSFIRE sample ($9 \lesssim \log(M_*/M_\odot) \lesssim 11$), there is no obvious M_* dependence of the MZR scatter.

7. We investigated whether accounting for galaxy SFR in the MZR relation can further reduce the scatter, as it appears to do at $z \sim 0$ —i.e., whether there is a MZR-SFR relation, or “fundamental metallicity relation” (FMR) at $z \sim 2.3$. We find that the KBSS-MOSFIRE $z \sim 2.3$ sample is inconsistent with the recently-suggested parametrizations of the FMR as calibrated locally—i.e., the effect of SFR on the observed metallicity at fixed stellar mass (if present) must be different at $z \sim 2.3$ (Figures 24 and 25). *At present, the residuals between the observed metallicities and those predicted by the FMR suggest that accounting for SFR in the $z \sim 2.3$ MZR has no measurable effect on the metallicity scatter at a given stellar mass.*

8. We have pointed out that the small scatter in the $z \sim 2.3$ MZR ($\sigma \simeq 0.11$ dex including measurement errors; estimated intrinsic scatter $\sigma \simeq 0.09$ dex) is uncomfortably small compared with the uncertainties inherent in the calibrations of the strong-line metallicity methods, even when the latter are recalibrated only over the range of line indices covered by the $z \sim 2.3$ observations. When taken together with the photoionization models showing that the observed line ratios at $z \sim 2.3$ are more strongly affected by ionizing radiation field intensity and shape than by ionized gas metallicity, it suggests that the more fundamental correlation (of which the MZR is a by-product) is between M_* and the metallicity of the massive stars that determine the ionization/excitation state of the gas in their surroundings.

8.1. Implications and Future Work

Among the issues raised, we regard the following as unresolved and particularly interesting to pursue with future work:

8.1.1. Metallicity measurements at high redshift

There is currently very limited evidence that *any* strong-line abundance estimates at high redshift reliably measure gas-phase metallicity, as is universally assumed. However, the prospects for improving the situation are good, since instruments now exist (like MOSFIRE) capable of obtaining sufficiently sensitive spectra of high redshift galaxies to measure weak lines such as [OIII] λ 4364 whose strengths relative to strong lines provide direct information on physical conditions in the ionized gas. We have shown that it should be feasible to obtain such measurements at $z \sim 2.3$ for individual galaxies with metallicities as high as $Z \sim 0.5 Z_\odot$; it may be possible, using spectral stacks, to extend the calibrations to higher metallicity (see, e.g., Andrews & Martini 2013). Because of the remaining uncertainty associated with converting strong-line ratios to oxygen abundance, until direct metallicity cross-checks have been completed, we suggest that galaxies should not be “pre-screened” for deep follow-

up based on their strong-line-implied abundances (see section 6). One should also obtain, wherever possible, measurements of the rest-UV OIII] intercombination lines available from deep ground-based optical spectroscopy (section 6.1), which in some cases may be more sensitive, albeit more dependent on nebular extinction corrections, than measurement of [OIII] λ 4364 in the rest-frame optical.

8.1.2. The dominant ionizing sources in high-redshift H II regions

In order to temporarily avoid uncertainties associated with the details of the predicted ionizing spectra of massive stars in high-redshift galaxies, we have modeled the net radiation field shape using a single-temperature blackbody, which can be thought of as the effective temperature of whatever stars are dominating the radiation field for photon energies between 1 and 4 Ryd. It appears that successful population synthesis models used for future more detailed models of the ionized gas in $z \sim 2.3$ galaxies must be capable of producing, in steady state, a net luminosity-weighted spectrum resembling a $\simeq 50,000 - 60,000$ K blackbody in the far-UV. This may have implications for the high-mass end of the stellar initial mass function (IMF), as well as for the details of the models for the most massive stars. Satisfying the constraint that the stars must produce nebulae with the observed properties may also have implications for the production and transfer of ionizing photons from young galaxies at high redshift.

8.1.3. The slope and normalization of the MZR at $z \sim 2.3$

According to our preferred form of the $z \sim 2.3$ MZR presented in section 7, the average metallicity of the dominant star-forming galaxy population changes by only $\simeq 0.3$ dex over more than 2 orders of magnitude in M_* , $9 \lesssim \log(M_*/M_\odot) \sim 11.3$. This shallow dependence of (strong-line-inferred) metallicity on stellar mass ($12 + \log(\text{O}/\text{H}) \propto 0.12 \log M_*$) is comparable to what is observed *over the same range in M_** at $z \simeq 0$. Because the strong nebular lines appear to be relatively *insensitive* to the ionized gas metallicity, one should be cautious in treating inferred oxygen abundances as direct indications of metallicity in the dominant gas reservoirs of galaxies. Similarly, one should also be cautious interpreting changes in strong-line ratios (e.g., as a function of position within galaxies, or scatter among galaxies of similar stellar mass) as differences in *gas-phase* metallicity—they are perhaps more likely to signal changes in the ionizing sources and their distribution, which may have a different origin.

8.1.4. Fundamental correlations between nebular line ratios and galaxy properties

As discussed in section 7.1, the tightness of the relationship between inferred oxygen abundance and M_* is difficult to understand given the uncertainties in the calibration of the strong line indices onto direct (T_e) based oxygen abundance. The observed relationship is more easily understood if a) there is a relatively narrow range of radiation field effective temperature across all galaxy masses probed in the current sample and b) there is a monotonic relationship between effective ionization parameter Γ and M_* . Understanding *why* ionization level and excitation are so strongly linked to galaxy mass is a key goal for future work.

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Table 1
KBSS-MOSFIRE Galaxies with both [NII]/H α and [OIII]/H β Measurements

Name	z_{neb}	$\log M_*$ (M_{\odot})	$\text{SFR}_{\text{H}\alpha}$ ($M_{\odot} \text{ yr}^{-1}$)	$\log([\text{NII}]/\text{H}\alpha)$	$\log([\text{OIII}]/\text{H}\beta)$	$12 + \log(\text{O}/\text{H})$ (N2) ^b	$12 + \log(\text{O}/\text{H})$ (O3N2) ^c
Q0100-BX172*	2.3118	-0.81 ^{+0.04} _{-0.04}	+0.98 ^{+0.01} _{-0.01}	8.44 ^{+0.02} _{-0.02}	8.16 ^{+0.01} _{-0.01}
Q0100-BX205	2.2912	9.88	24.8	-1.22 ^{+0.29} _{-0.17}	+0.59 ^{+0.06} _{-0.05}	8.21 ^{+0.16} _{-0.10}	8.15 ^{+0.06} _{-0.09}
Q0100-BX210	2.2769	10.10	18.6	-0.80 ^{+0.16} _{-0.12}	+0.48 ^{+0.04} _{-0.04}	8.45 ^{+0.09} _{-0.07}	8.32 ^{+0.04} _{-0.05}
Q0100-BX224	2.1076	9.62	30.9	-0.85 ^{+0.09} _{-0.07}	+0.72 ^{+0.07} _{-0.06}	8.42 ^{+0.05} _{-0.04}	8.23 ^{+0.03} _{-0.03}
Q0100-BX90	2.2850	9.97	37.2	-0.85 ^{+0.07} _{-0.06}	+0.69 ^{+0.05} _{-0.04}	8.42 ^{+0.04} _{-0.04}	8.24 ^{+0.03} _{-0.03}
Q0100-MD19	2.1076	10.26	4.1	-0.54 ^{+0.09} _{-0.07}	+0.16 ^{+0.13} _{-0.10}	8.59 ^{+0.05} _{-0.04}	8.61 ^{+0.05} _{-0.04}
Q0100-RK17	2.1048	11.44	153.3	-0.36 ^{+0.05} _{-0.05}	+0.50 ^{+0.23} _{-0.15}	8.70 ^{+0.03} _{-0.03}	8.45 ^{+0.07} _{-0.05}
Q0105-BX129	2.2880	9.53	25.8	-0.98 ^{+0.16} _{-0.12}	+0.57 ^{+0.03} _{-0.03}	8.34 ^{+0.09} _{-0.07}	8.24 ^{+0.04} _{-0.05}
Q0105-BX147	2.3857	10.00	12.1	-1.10 ^{+0.12} _{-0.16}	+0.65 ^{+0.07} _{-0.03}	8.27 ^{+0.14} _{-0.09}	8.17 ^{+0.05} _{-0.08}
Q0105-BX163	2.2912	10.24	11.0	-0.99 ^{+0.20} _{-0.14}	+0.72 ^{+0.02} _{-0.02}	8.34 ^{+0.12} _{-0.08}	8.18 ^{+0.04} _{-0.07}
Q0105-BX57	2.2588	10.17	52.8	-0.89 ^{+0.04} _{-0.03}	+0.55 ^{+0.05} _{-0.04}	8.39 ^{+0.02} _{-0.02}	8.27 ^{+0.02} _{-0.02}
Q0105-BX58*	2.5351	-0.22 ^{+0.09} _{-0.07}	+0.64 ^{+0.07} _{-0.06}	8.77 ^{+0.05} _{-0.04}	8.45 ^{+0.03} _{-0.03}
Q0105-BX77	2.2930	10.22	12.3	-0.91 ^{+0.16} _{-0.12}	+0.81 ^{+0.03} _{-0.03}	8.38 ^{+0.09} _{-0.07}	8.18 ^{+0.04} _{-0.05}
Q0105-BX79	2.1228	10.52	25.0	-0.30 ^{+0.12} _{-0.09}	+0.36 ^{+0.12} _{-0.05}	8.73 ^{+0.07} _{-0.05}	8.52 ^{+0.06} _{-0.05}
Q0105-BX89	2.2277	9.83	52.0	-0.97 ^{+0.20} _{-0.14}	+0.64 ^{+0.06} _{-0.05}	8.35 ^{+0.11} _{-0.08}	8.22 ^{+0.05} _{-0.07}
Q0105-BX95	2.0305	10.38	31.0	-0.72 ^{+0.08} _{-0.07}	+0.23 ^{+0.15} _{-0.11}	8.49 ^{+0.05} _{-0.04}	8.43 ^{+0.05} _{-0.04}
Q0142-BX122	2.4176	9.53	33.8	-0.84 ^{+0.16} _{-0.12}	+0.48 ^{+0.07} _{-0.06}	8.42 ^{+0.09} _{-0.07}	8.31 ^{+0.04} _{-0.06}
Q0142-BX242	2.2811	9.68	48.9	-1.04 ^{+0.15} _{-0.11}	+0.54 ^{+0.04} _{-0.03}	8.31 ^{+0.08} _{-0.06}	8.23 ^{+0.04} _{-0.05}
Q0142-BX81	2.5026	9.46	46.4	-0.78 ^{+0.12} _{-0.10}	+0.61 ^{+0.03} _{-0.03}	8.45 ^{+0.07} _{-0.05}	8.29 ^{+0.03} _{-0.04}
Q0142-MD20	2.5005	9.56	70.3	-0.47 ^{+0.10} _{-0.09}	+0.41 ^{+0.08} _{-0.07}	8.63 ^{+0.07} _{-0.05}	8.45 ^{+0.04} _{-0.04}
Q0207-BX150	2.1148	10.38	20.8	-0.78 ^{+0.15} _{-0.11}	+0.43 ^{+0.13} _{-0.10}	8.45 ^{+0.09} _{-0.06}	8.34 ^{+0.06} _{-0.06}
Q0207-BX285	2.1504	9.77	6.1	-0.75 ^{+0.20} _{-0.23}	+0.41 ^{+0.06} _{-0.06}	8.47 ^{+0.11} _{-0.06}	8.36 ^{+0.05} _{-0.07}
Q0207-BX65	2.1918	10.18	8.0	-0.80 ^{+0.15} _{-0.15}	+0.24 ^{+0.18} _{-0.12}	8.44 ^{+0.13} _{-0.09}	8.40 ^{+0.07} _{-0.08}
Q0207-BX74	2.1889	9.01	48.9	-1.48 ^{+0.19} _{-0.13}	+0.90 ^{+0.03} _{-0.03}	8.05 ^{+0.11} _{-0.07}	7.97 ^{+0.04} _{-0.06}
Q0821-BX45	2.1803	10.66	64.2	-1.05 ^{+0.06} _{-0.05}	+0.76 ^{+0.06} _{-0.05}	8.30 ^{+0.03} _{-0.03}	8.15 ^{+0.03} _{-0.03}
Q0821-BX77	2.2944	9.61	27.3	-0.83 ^{+0.14} _{-0.10}	+0.74 ^{+0.05} _{-0.05}	8.43 ^{+0.08} _{-0.06}	8.23 ^{+0.04} _{-0.05}
Q0821-D8*	2.5678	-0.15 ^{+0.07} _{-0.06}	+0.95 ^{+0.12} _{-0.10}	8.81 ^{+0.04} _{-0.03}	8.38 ^{+0.04} _{-0.04}
Q0821-RK29	2.4683	10.71	79.2	-0.51 ^{+0.13} _{-0.10}	+0.35 ^{+0.04} _{-0.04}	8.61 ^{+0.03} _{-0.06}	8.46 ^{+0.04} _{-0.04}
Q1442-BX108	2.4281	9.69	70.9	-1.07 ^{+0.17} _{-0.12}	+0.46 ^{+0.01} _{-0.01}	8.29 ^{+0.09} _{-0.07}	8.24 ^{+0.04} _{-0.05}
Q1442-BX133	2.1053	9.67	10.9	-1.01 ^{+0.27} _{-0.16}	+0.52 ^{+0.16} _{-0.12}	8.33 ^{+0.15} _{-0.09}	8.24 ^{+0.07} _{-0.09}
Q1442-BX138	2.4336	9.61	45.7	-0.92 ^{+0.15} _{-0.11}	+0.88 ^{+0.20} _{-0.14}	8.38 ^{+0.08} _{-0.06}	8.16 ^{+0.07} _{-0.06}
Q1442-BX154	2.4297	9.57	22.6	-0.41 ^{+0.27} _{-0.17}	+0.68 ^{+0.12} _{-0.10}	8.67 ^{+0.16} _{-0.10}	8.38 ^{+0.07} _{-0.09}
Q1442-BX160	2.4417	9.54	32.4	-0.97 ^{+0.10} _{-0.08}	+0.57 ^{+0.02} _{-0.02}	8.35 ^{+0.06} _{-0.06}	8.24 ^{+0.03} _{-0.03}
Q1442-BX235	2.4445	10.60	39.2	-0.83 ^{+0.10} _{-0.08}	+0.50 ^{+0.04} _{-0.04}	8.43 ^{+0.06} _{-0.05}	8.31 ^{+0.03} _{-0.03}
Q1442-BX270	2.3578	9.52	14.5	-0.80 ^{+0.09} _{-0.08}	+0.69 ^{+0.03} _{-0.02}	8.44 ^{+0.05} _{-0.04}	8.25 ^{+0.03} _{-0.03}
Q1442-BX69b	2.1478	10.12	17.5	-0.68 ^{+0.10} _{-0.08}	+0.24 ^{+0.07} _{-0.06}	8.51 ^{+0.06} _{-0.05}	8.44 ^{+0.04} _{-0.04}
Q1442-MD13	2.4529	10.55	90.1	-1.01 ^{+0.13} _{-0.10}	+0.67 ^{+0.01} _{-0.01}	8.32 ^{+0.07} _{-0.06}	8.19 ^{+0.03} _{-0.04}
Q1442-MD53	2.2932	10.96	53.9	-0.31 ^{+0.03} _{-0.03}	+0.42 ^{+0.14} _{-0.12}	8.72 ^{+0.02} _{-0.02}	8.50 ^{+0.04} _{-0.03}
Q1442-MD57	2.4438	9.85	23.9	-0.36 ^{+0.10} _{-0.08}	+0.75 ^{+0.02} _{-0.02}	8.69 ^{+0.06} _{-0.05}	8.37 ^{+0.03} _{-0.04}
Q1549-BX180	2.3871	9.31	41.7	-0.63 ^{+0.08} _{-0.07}	+0.44 ^{+0.06} _{-0.05}	8.54 ^{+0.04} _{-0.04}	8.39 ^{+0.03} _{-0.03}
Q1549-BX197	2.4354	9.61	30.7	-0.78 ^{+0.15} _{-0.11}	+0.19 ^{+0.24} _{-0.17}	8.46 ^{+0.08} _{-0.06}	8.42 ^{+0.09} _{-0.07}
Q1549-BX207	2.3802	9.64	58.0	-0.85 ^{+0.12} _{-0.09}	+0.57 ^{+0.12} _{-0.10}	8.41 ^{+0.07} _{-0.05}	8.27 ^{+0.05} _{-0.05}
Q1549-BX223	2.3493	9.51	123.4	-0.75 ^{+0.07} _{-0.06}	+0.56 ^{+0.03} _{-0.03}	8.47 ^{+0.04} _{-0.04}	8.31 ^{+0.02} _{-0.03}
Q1549-BX240	2.0419	10.62	19.5	-0.34 ^{+0.12} _{-0.09}	+0.36 ^{+0.26} _{-0.16}	8.71 ^{+0.07} _{-0.05}	8.51 ^{+0.09} _{-0.06}
Q1549-BX51	2.2897	9.71	19.1	-1.09 ^{+0.21} _{-0.14}	+0.70 ^{+0.08} _{-0.07}	8.28 ^{+0.12} _{-0.08}	8.16 ^{+0.05} _{-0.07}
Q1603-BX127	2.5522	9.59	55.9	-0.87 ^{+0.30} _{-0.18}	+0.75 ^{+0.08} _{-0.07}	8.40 ^{+0.17} _{-0.10}	8.21 ^{+0.06} _{-0.10}
Q1603-BX190	2.0217	10.00	56.3	-0.89 ^{+0.17} _{-0.12}	+0.40 ^{+0.21} _{-0.14}	8.40 ^{+0.07} _{-0.07}	8.32 ^{+0.08} _{-0.07}
Q1603-BX277	2.5499	9.90	48.6	-0.61 ^{+0.16} _{-0.12}	+0.65 ^{+0.04} _{-0.04}	8.55 ^{+0.09} _{-0.07}	8.33 ^{+0.04} _{-0.05}
Q1603-BX294	2.4511	10.03	40.5	-1.04 ^{+0.20} _{-0.14}	+0.60 ^{+0.02} _{-0.02}	8.31 ^{+0.11} _{-0.08}	8.21 ^{+0.04} _{-0.04}
Q1603-BX379	2.1767	10.10	41.3	-0.74 ^{+0.10} _{-0.08}	+0.27 ^{+0.16} _{-0.11}	8.48 ^{+0.06} _{-0.05}	8.41 ^{+0.06} _{-0.05}
Q1603-MD26	2.5512	9.85	63.2	-0.90 ^{+0.25} _{-0.16}	+0.42 ^{+0.04} _{-0.04}	8.39 ^{+0.14} _{-0.09}	8.31 ^{+0.05} _{-0.08}
Q1603-MD42	2.3808	9.99	9.9	-0.60 ^{+0.16} _{-0.13}	+0.30 ^{+0.12} _{-0.12}	8.56 ^{+0.08} _{-0.08}	8.44 ^{+0.07} _{-0.08}
Q1603-MD85	2.4513	10.47	9.9	-0.53 ^{+0.15} _{-0.11}	+0.60 ^{+0.11} _{-0.09}	8.60 ^{+0.08} _{-0.06}	8.37 ^{+0.05} _{-0.06}
Q1623-BX429	2.0159	9.94	10.4	-0.79 ^{+0.08} _{-0.07}	+0.26 ^{+0.20} _{-0.14}	8.45 ^{+0.05} _{-0.04}	8.40 ^{+0.07} _{-0.05}

Table 1
KBSS-MOSFIRE Galaxies with both [NII]/H α and [OIII]/H β Measurements (*cont.*)

Name	z_{neb}	$\log M_*$ (M_\odot)	$\text{SFR}_{\text{H}\alpha}$ ($M_\odot \text{ yr}^{-1}$)	$\log([\text{NII}]/\text{H}\alpha)$	$\log([\text{OIII}]/\text{H}\beta)$	$12 + \log(\text{O}/\text{H})$ (N2) ^b	$12 + \log(\text{O}/\text{H})$ (O3N2) ^c
Q1700-BX490	2.3958	10.04	117.4	-0.96 ^{+0.05} _{-0.05}	+0.74 ^{+0.03} _{-0.03}	8.35 ^{+0.03} _{-0.03}	8.18 ^{+0.02} _{-0.02}
Q1700-BX505	2.3083	10.65	26.6	-0.72 ^{+0.19} _{-0.13}	+0.50 ^{+0.32} _{-0.18}	8.49 ^{+0.11} _{-0.08}	8.34 ^{+0.11} _{-0.08}
Q1700-BX563	2.2910	10.32	66.6	-1.06 ^{+0.14} _{-0.10}	+0.75 ^{+0.05} _{-0.04}	8.30 ^{+0.08} _{-0.06}	8.15 ^{+0.04} _{-0.05}
Q1700-BX710	2.2947	10.51	54.9	-0.97 ^{+0.09} _{-0.08}	+0.65 ^{+0.10} _{-0.08}	8.35 ^{+0.05} _{-0.04}	8.21 ^{+0.04} _{-0.04}
Q1700-BX711	2.2948	8.60	33.8	-1.18 ^{+0.11} _{-0.09}	+0.90 ^{+0.06} _{-0.05}	8.23 ^{+0.06} _{-0.05}	8.06 ^{+0.03} _{-0.04}
Q1700-BX752	2.4002	10.59	37.7	-0.55 ^{+0.06} _{-0.05}	+0.30 ^{+0.09} _{-0.07}	8.59 ^{+0.04} _{-0.03}	8.46 ^{+0.03} _{-0.03}
Q1700-BX763	2.2920	10.11	21.5	-0.81 ^{+0.08} _{-0.07}	+0.63 ^{+0.07} _{-0.06}	8.44 ^{+0.05} _{-0.04}	8.27 ^{+0.03} _{-0.03}
Q1700-BX913	2.2904	10.23	34.0	-0.93 ^{+0.14} _{-0.11}	+0.73 ^{+0.08} _{-0.07}	8.37 ^{+0.08} _{-0.06}	8.20 ^{+0.04} _{-0.05}
Q1700-BX917	2.3066	10.73	54.2	-0.82 ^{+0.07} _{-0.06}	+0.64 ^{+0.14} _{-0.11}	8.43 ^{+0.04} _{-0.03}	8.26 ^{+0.03} _{-0.04}
Q1700-MD69	2.2879	11.21	106.4	-0.56 ^{+0.07} _{-0.06}	+0.20 ^{+0.12} _{-0.09}	8.58 ^{+0.04} _{-0.04}	8.49 ^{+0.04} _{-0.04}
Q2206-BX88	2.1808	10.29	22.7	-0.96 ^{+0.09} _{-0.08}	+0.66 ^{+0.05} _{-0.04}	8.35 ^{+0.05} _{-0.04}	8.21 ^{+0.03} _{-0.03}
Q2206-BX140	2.3515	10.14	13.7	-0.90 ^{+0.13} _{-0.10}	+0.61 ^{+0.08} _{-0.07}	8.39 ^{+0.07} _{-0.06}	8.25 ^{+0.04} _{-0.05}
Q2343-BX182	2.2876	9.80	27.0	-1.04 ^{+0.11} _{-0.09}	+0.68 ^{+0.04} _{-0.03}	8.31 ^{+0.06} _{-0.05}	8.18 ^{+0.03} _{-0.04}
Q2343-BX222	2.2872	10.63	38.2	-0.60 ^{+0.08} _{-0.06}	+0.26 ^{+0.07} _{-0.06}	8.56 ^{+0.04} _{-0.04}	8.45 ^{+0.03} _{-0.03}
Q2343-BX231	2.4988	10.10	505.1	-0.59 ^{+0.03} _{-0.03}	+0.52 ^{+0.04} _{-0.04}	8.57 ^{+0.02} _{-0.02}	8.38 ^{+0.02} _{-0.02}
Q2343-BX336	2.5444	10.11	59.6	-0.82 ^{+0.08} _{-0.07}	+0.56 ^{+0.02} _{-0.02}	8.43 ^{+0.04} _{-0.04}	8.29 ^{+0.02} _{-0.02}
Q2343-BX348	2.4492	10.49	75.0	-0.64 ^{+0.03} _{-0.03}	+0.40 ^{+0.02} _{-0.02}	8.53 ^{+0.02} _{-0.02}	8.40 ^{+0.01} _{-0.01}
Q2343-BX389	2.1711	10.97	81.1	-0.82 ^{+0.07} _{-0.06}	+0.50 ^{+0.07} _{-0.06}	8.43 ^{+0.04} _{-0.04}	8.31 ^{+0.03} _{-0.03}
Q2343-BX418	2.3053	8.88	32.7	-1.32 ^{+0.14} _{-0.10}	+0.80 ^{+0.05} _{-0.04}	8.15 ^{+0.08} _{-0.06}	8.05 ^{+0.04} _{-0.05}
Q2343-BX442	2.1770	11.11	81.0	-0.48 ^{+0.05} _{-0.05}	+0.31 ^{+0.15} _{-0.11}	8.63 ^{+0.03} _{-0.02}	8.48 ^{+0.03} _{-0.04}
Q2343-BX445	2.5445	9.69	47.8	-1.09 ^{+0.18} _{-0.13}	+0.74 ^{+0.02} _{-0.02}	8.28 ^{+0.10} _{-0.07}	8.15 ^{+0.04} _{-0.06}
Q2343-BX460	2.3946	9.18	34.6	-1.03 ^{+0.11} _{-0.09}	+0.70 ^{+0.03} _{-0.03}	8.32 ^{+0.06} _{-0.05}	8.18 ^{+0.03} _{-0.04}
Q2343-BX496	2.3934	9.28	24.1	-0.93 ^{+0.14} _{-0.11}	+0.76 ^{+0.15} _{-0.11}	8.39 ^{+0.06} _{-0.06}	8.25 ^{+0.04} _{-0.04}
Q2343-BX537	2.3394	9.58	34.7	-0.98 ^{+0.13} _{-0.10}	+0.44 ^{+0.06} _{-0.05}	8.34 ^{+0.07} _{-0.06}	8.28 ^{+0.04} _{-0.05}
Q2343-BX587	2.2428	10.18	38.6	-0.67 ^{+0.03} _{-0.03}	+0.39 ^{+0.06} _{-0.05}	8.52 ^{+0.02} _{-0.01}	8.39 ^{+0.02} _{-0.02}
Q2343-BX601	2.3768	10.47	65.4	-0.91 ^{+0.07} _{-0.06}	+0.47 ^{+0.05} _{-0.05}	8.38 ^{+0.04} _{-0.03}	8.29 ^{+0.03} _{-0.03}
Q2343-D29	2.3866	9.83	6.1	-0.58 ^{+0.23} _{-0.15}	+0.41 ^{+0.05} _{-0.04}	8.57 ^{+0.13} _{-0.09}	8.41 ^{+0.03} _{-0.07}
Q2343-D35	2.3989	10.82	35.9	-0.51 ^{+0.06} _{-0.06}	-0.09 ^{+0.08} _{-0.07}	8.61 ^{+0.04} _{-0.03}	8.60 ^{+0.03} _{-0.03}

^a Error bars are 1σ based on measurement uncertainties only.

^b Oxygen abundance assuming the “N2” calibration of PP04.

^c Oxygen abundance assuming the “O3N2” calibration of PP04.

* Object is identified as an AGN on the basis of both its rest-UV and rest-optical spectra.

Table 2
KBSS-MOSFIRE Galaxies with [OIII]/H β Measurements and [NII]/H α Limits^a

Name	z_{neb}	$\log M_*$ (M_{\odot})	$\text{SFR}_{\text{H}\alpha}$ ($M_{\odot} \text{ yr}^{-1}$)	$\log([\text{NII}]/\text{H}\alpha)$	$\log([\text{OIII}]/\text{H}\beta)$	$12 + \log(\text{O}/\text{H})$ (N2) ^b	$12 + \log(\text{O}/\text{H})$ (O3N2) ^c
Q0100-BX167	2.2891	9.39	8.2	< -0.81	$0.50^{+0.12}_{-0.09}$	< 8.44	< 8.31
Q0100-BX277	2.1059	10.18	9.8	< -0.98	$0.49^{+0.20}_{-0.13}$	< 8.34	< 8.26
Q0100-BX57	2.2707	9.41	16.9	< -0.88	$0.74^{+0.03}_{-0.03}$	< 8.40	< 8.21
Q0105-MD12	2.5052	9.55	65.7	< -1.07	$0.62^{+0.01}_{-0.01}$	< 8.29	< 8.19
Q0142-BX138	2.4176	9.48	4.8	< -0.67	$0.73^{+0.07}_{-0.06}$	< 8.52	< 8.28
Q0142-BX214	2.3865	9.69	24.6	< -0.98	$0.68^{+0.03}_{-0.03}$	< 8.34	< 8.20
Q0207-BX211	2.5469	9.71	37.2	< -0.87	$0.47^{+0.12}_{-0.09}$	< 8.41	< 8.30
Q0207-BX87	2.1924	10.04	20.2	< -1.18	$0.85^{+0.02}_{-0.02}$	< 8.23	< 8.08
Q0821-BX207	2.4135	9.59	35.5	< -1.21	$0.59^{+0.02}_{-0.02}$	< 8.21	< 8.15
Q0821-BX221	2.3959	9.76	18.1	< -1.20	$0.76^{+0.07}_{-0.06}$	< 8.22	< 8.10
Q0821-BX80	2.4451	10.24	29.4	< -0.98	$0.56^{+0.13}_{-0.10}$	< 8.34	< 8.24
Q0821-BX92	2.4165	9.21	8.8	< -0.70	$0.57^{+0.09}_{-0.08}$	< 8.50	< 8.32
Q0821-MD5	2.5369	9.88	5.6	< -0.71	$0.73^{+0.06}_{-0.06}$	< 8.50	< 8.27
Q0821-RK27	2.4487	10.07	47.1	< -0.68	$0.25^{+0.15}_{-0.11}$	< 8.51	< 8.43
Q1442-BX116	2.0463	9.74	41.5	< -0.68	$0.63^{+0.11}_{-0.09}$	< 8.51	< 8.31
Q1442-BX172	2.4496	10.07	42.2	< -0.93	$0.44^{+0.06}_{-0.05}$	< 8.37	< 8.29
Q1442-BX199	2.2937	9.28	30.8	< -1.08	$0.69^{+0.03}_{-0.05}$	< 8.29	< 8.17
Q1442-BX290	2.4319	9.68	23.7	< -0.97	$0.43^{+0.07}_{-0.06}$	< 8.35	< 8.28
Q1442-BX295	2.4513	9.34	12.5	< -1.14	$0.60^{+0.05}_{-0.05}$	< 8.25	< 8.18
Q1442-BX69	2.0887	10.53	14.1	< -0.82	$0.80^{+0.05}_{-0.05}$	< 8.43	< 8.21
Q1549-BX170	2.3837	9.76	14.5	< -1.17	$0.85^{+0.16}_{-0.12}$	< 8.23	< 8.08
Q1603-BX173	2.5491	9.01	21.6	< -0.75	$0.83^{+0.14}_{-0.11}$	< 8.47	< 8.22
Q1603-BX255	2.4350	10.21	35.7	< -0.94	$0.54^{+0.11}_{-0.09}$	< 8.36	< 8.26
Q1623-BX376	2.4092	10.00	18.8	< -0.62	$0.71^{+0.07}_{-0.06}$	< 8.55	< 8.30
Q1623-MD102	2.4072	10.28	13.3	< -0.58	$0.11^{+0.17}_{-0.12}$	< 8.57	< 8.51
Q1700-BX585	2.3068	9.06	16.1	< -1.01	$0.45^{+0.30}_{-0.18}$	< 8.32	< 8.26
Q1700-BX625	2.0754	9.80	9.8	< -1.04	$0.62^{+0.13}_{-0.10}$	< 8.31	< 8.20
Q1700-BX708	2.3989	9.69	42.8	< -1.13	$0.65^{+0.08}_{-0.07}$	< 8.26	< 8.16
Q1700-BX713	2.1379	9.63	13.0	< -1.00	$0.35^{+0.09}_{-0.08}$	< 8.33	< 8.30
Q1700-BX717	2.4357	9.46	9.8	< -0.82	$0.34^{+0.31}_{-0.18}$	< 8.43	< 8.36
Q1700-BX772	2.3417	9.71	13.7	< -0.95	$0.58^{+0.30}_{-0.17}$	< 8.36	< 8.24
Q1700-MD77	2.5078	9.47	61.6	< -0.82	$0.70^{+0.13}_{-0.10}$	< 8.43	< 8.24
Q2206-BM64	2.1945	9.33	31.1	< -0.97	$0.75^{+0.08}_{-0.07}$	< 8.35	< 8.18
Q2206-BX128	2.3480	10.03	36.0	< -1.10	$0.46^{+0.18}_{-0.13}$	< 8.27	< 8.23
Q2206-BX145	2.2352	9.44	5.5	< -0.72	$0.76^{+0.10}_{-0.10}$	< 8.49	< 8.26
Q2206-BX153	2.4457	9.29	32.4	< -0.58	$0.74^{+0.08}_{-0.15}$	< 8.57	< 8.31
Q2206-BX199	2.0765	9.80	16.0	< -1.17	$0.88^{+0.06}_{-0.05}$	< 8.23	< 8.07
Q2343-BX473	2.5437	10.32	32.5	< -1.24	$0.72^{+0.02}_{-0.02}$	< 8.20	< 8.10
Q2343-BX660	2.1741	8.96	37.0	< -1.66	$0.86^{+0.04}_{-0.03}$	< 7.95	< 7.92
Q2343-D25	2.1868	9.49	10.8	< -1.04	$0.60^{+0.17}_{-0.12}$	< 8.31	< 8.21
Q2343-D35b	2.4140	< -0.57	$0.81^{+0.03}_{-0.03}$	< 8.58	< 8.29
Q2343-MD37	2.4244	9.94	5.3	< -0.73	$0.59^{+0.05}_{-0.05}$	< 8.49	< 8.31
Q2343-MD86	2.3976	9.40	6.6	< -0.94	$0.56^{+0.04}_{-0.04}$	< 8.37	< 8.25

^a Upper limits on $\log([\text{NII}]/\text{H}\alpha)$ are 2σ ; error bars are otherwise 1σ based on measurement errors only.

^b 2σ upper limit on oxygen abundance assuming the “N2” calibration of PP04.

^c 2σ upper limit on oxygen abundance assuming the “O3N2” calibration of PP04.

Table 3
KBSS-MOSFIRE Galaxies with [NII]/H α and Stellar Mass Measurements

Name	z_{neb}	$\log M_*$ M_{\odot}	$\text{SFR}_{\text{H}\alpha}$ $(M_{\odot}\text{yr}^{-1})$	$\log([\text{NII}]/\text{H}\alpha)$	$12 + \log(\text{O}/\text{H})$ (N2 PP04)
Q0100-BX148	1.9635	9.47	24.0	$-0.93^{+0.10}_{-0.07}$	$8.37^{+0.06}_{-0.04}$
Q0100-BX95	2.2096	10.28	28.4	$-0.70^{+0.08}_{-0.06}$	$8.50^{+0.05}_{-0.04}$
Q0100-RK21	2.0629	10.51	134.3	$-0.29^{+0.14}_{-0.06}$	$8.74^{+0.08}_{-0.06}$
Q0105-BX132	2.2117	10.76	22.8	$-0.55^{+0.12}_{-0.08}$	$8.59^{+0.07}_{-0.05}$
Q0105-BX180	2.0296	8.97	8.0	$-0.74^{+0.15}_{-0.11}$	$8.48^{+0.09}_{-0.06}$
Q0105-BX186	2.1999	10.57	20.0	$-0.61^{+0.10}_{-0.07}$	$8.55^{+0.06}_{-0.04}$
Q0105-BX52	1.9716	10.63	90.0	$-0.66^{+0.06}_{-0.05}$	$8.52^{+0.04}_{-0.03}$
Q0105-BX93	2.0301	10.05	48.6	$-1.19^{+0.06}_{-0.05}$	$8.22^{+0.03}_{-0.03}$
Q0142-BX119	2.2236	10.85	20.7	$-0.63^{+0.08}_{-0.06}$	$8.54^{+0.05}_{-0.04}$
Q0142-BX188	2.0600	9.85	11.5	$-0.57^{+0.08}_{-0.06}$	$8.58^{+0.05}_{-0.04}$
Q0142-BX241	2.2843	9.82	29.4	$-0.77^{+0.08}_{-0.06}$	$8.46^{+0.05}_{-0.04}$
Q0142-BX248	2.4978	9.64	70.2	$-0.72^{+0.14}_{-0.09}$	$8.49^{+0.08}_{-0.05}$
Q0142-BX61	2.0702	9.62	33.9	$-0.61^{+0.08}_{-0.06}$	$8.55^{+0.05}_{-0.04}$
Q0207-BX150	2.1147	10.38	20.8	$-0.78^{+0.12}_{-0.08}$	$8.46^{+0.07}_{-0.05}$
Q0207-BX285	2.1503	9.78	6.1	$-0.74^{+0.11}_{-0.08}$	$8.48^{+0.06}_{-0.05}$
Q0207-BX67	2.1955	9.77	35.5	$-0.99^{+0.10}_{-0.07}$	$8.33^{+0.06}_{-0.04}$
Q0821-BX101	2.4455	10.87	72.1	$-0.18^{+0.11}_{-0.09}$	$8.80^{+0.06}_{-0.05}$
Q0821-BX38	2.1843	10.88	30.1	$-0.53^{+0.14}_{-0.10}$	$8.60^{+0.08}_{-0.06}$
Q0821-BX45	2.1803	10.66	64.2	$-1.05^{+0.06}_{-0.05}$	$8.30^{+0.03}_{-0.03}$
Q0821-BX52	2.1770	10.55	10.3	$-0.60^{+0.19}_{-0.13}$	$8.56^{+0.11}_{-0.08}$
Q0821-BX77	2.2944	9.62	27.3	$-0.83^{+0.13}_{-0.10}$	$8.43^{+0.08}_{-0.06}$
Q0821-MD38	2.0920	10.49	51.6	$-0.62^{+0.08}_{-0.07}$	$8.55^{+0.05}_{-0.04}$
Q0821-RK5 ^a	2.1834	11.79	41.8	$+0.02^{+0.06}_{-0.05}$	$8.91^{+0.03}_{-0.03}$
Q1442-BX154	2.4297	9.57	22.6	$-0.41^{+0.27}_{-0.17}$	$8.67^{+0.16}_{-0.09}$
Q1442-BX159	1.9956	9.97	6.8	$-0.81^{+0.09}_{-0.07}$	$8.44^{+0.05}_{-0.04}$
Q1442-BX270	2.3578	9.53	14.5	$-0.80^{+0.09}_{-0.08}$	$8.44^{+0.05}_{-0.04}$
Q1442-BX69b	2.1478	10.12	17.5	$-0.68^{+0.10}_{-0.08}$	$8.51^{+0.06}_{-0.05}$
Q1442-MD53	2.2932	10.96	53.9	$-0.31^{+0.03}_{-0.03}$	$8.72^{+0.02}_{-0.02}$
Q1442-MD57	2.4438	9.86	23.9	$-0.36^{+0.10}_{-0.08}$	$8.69^{+0.06}_{-0.05}$
Q1549-BX221	2.3410	9.43	63.3	$-0.79^{+0.17}_{-0.12}$	$8.45^{+0.10}_{-0.07}$
Q1549-BX227	2.0574	10.69	24.7	$-0.83^{+0.12}_{-0.08}$	$8.43^{+0.07}_{-0.05}$
Q1549-BX35	2.0620	9.41	28.5	$-0.71^{+0.18}_{-0.13}$	$8.50^{+0.10}_{-0.07}$
Q1549-BX42	2.2196	9.09	12.5	$-0.67^{+0.13}_{-0.08}$	$8.52^{+0.07}_{-0.05}$
Q1549-MD21	2.5343	11.05	100.5	$-0.74^{+0.04}_{-0.03}$	$8.48^{+0.02}_{-0.02}$
Q1603-BX191 ^a	2.5445	10.41	144.0	$-0.41^{+0.15}_{-0.10}$	$8.66^{+0.09}_{-0.06}$
Q1700-BX565	2.4741	10.22	17.1	$-0.57^{+0.10}_{-0.13}$	$8.57^{+0.06}_{-0.08}$
Q1700-BX649	2.2946	10.18	17.9	$-0.75^{+0.06}_{-0.05}$	$8.47^{+0.04}_{-0.03}$
Q1700-BX684	2.2921	10.36	20.6	$-0.66^{+0.12}_{-0.09}$	$8.52^{+0.07}_{-0.05}$
Q1700-BX691	2.1899	10.89	53.4	$-0.67^{+0.04}_{-0.03}$	$8.52^{+0.03}_{-0.02}$
Q1700-BX810	2.2923	10.09	54.7	$-0.87^{+0.10}_{-0.07}$	$8.40^{+0.06}_{-0.04}$
Q1700-BX879	2.3065	9.88	35.9	$-0.71^{+0.07}_{-0.05}$	$8.50^{+0.04}_{-0.03}$
Q1700-BX909	2.2935	10.65	58.4	$-0.83^{+0.07}_{-0.05}$	$8.43^{+0.04}_{-0.03}$
Q1700-BX929	2.3063	10.12	15.9	$-1.23^{+0.16}_{-0.19}$	$8.20^{+0.11}_{-0.06}$
Q1700-BX939	2.2971	9.74	39.5	$-0.94^{+0.17}_{-0.10}$	$8.37^{+0.09}_{-0.06}$
Q1700-BX950	2.2953	10.14	11.6	$-0.85^{+0.12}_{-0.08}$	$8.42^{+0.07}_{-0.05}$
Q1700-BX951	2.3054	10.49	22.3	$-0.69^{+0.11}_{-0.07}$	$8.51^{+0.07}_{-0.04}$
Q1700-BX984	2.2963	10.19	32.5	$-0.76^{+0.08}_{-0.06}$	$8.47^{+0.05}_{-0.04}$
Q2206-BX102	2.2101	11.19	119.4	$-0.32^{+0.03}_{-0.02}$	$8.72^{+0.01}_{-0.01}$
Q2206-BX110	2.0722	10.43	14.2	$-0.51^{+0.14}_{-0.11}$	$8.61^{+0.08}_{-0.06}$
Q2206-BX166	1.9742	9.90	31.2	$-0.95^{+0.07}_{-0.06}$	$8.36^{+0.04}_{-0.03}$
Q2206-BX168	2.1965	9.70	15.8	$-1.22^{+0.14}_{-0.13}$	$8.20^{+0.08}_{-0.08}$
Q2206-BX189	2.0782	11.02	82.9	$-0.49^{+0.13}_{-0.09}$	$8.62^{+0.08}_{-0.05}$
Q2206-BX68	2.0972	10.49	21.4	$-0.85^{+0.07}_{-0.06}$	$8.42^{+0.04}_{-0.03}$
Q2343-BX484	2.1874	10.36	14.3	$-0.87^{+0.17}_{-0.10}$	$8.41^{+0.10}_{-0.06}$

^a Objects identified as AGN on the basis of their MOSFIRE spectra.

Table 4
Properties of “Extreme” KBSS-MOSFIRE Galaxies

Object	[OII] Ratio ^a	n_e ^b	[OIII] Ratio ^c	H α /H β	T _e (K)	O32 ^d	[OIII] _{tot} /H β ^e	12 + log(O/H) ^f
Q0207-BX74 ^g	1.68 ± 0.11	1575 ⁺³⁰⁰ ₋₂₀₀	0.037 ± 0.005	3.46 ± 0.25	14300 ± 400	8.23 ± 0.41	10.06 ± 0.45	8.00 ± 0.05
Q2343-BX418	1.13 ± 0.05	580 ⁺⁸⁰ ₋₇₀	0.023 ± 0.003	2.81 ± 0.20	12830 ± 500	9.66 ± 0.38	8.67 ± 0.42	8.08 ± 0.05
Q2343-BX660	0.93 ± 0.04	300 ⁺⁴⁰ ₋₄₀	0.022 ± 0.004	2.77 ± 0.20	12650 ± 500	10.98 ± 0.50	9.58 ± 0.45	8.13 ± 0.06

^a Intensity ratio [OII] λ 3726/[OII] λ 3729.

^b Electron density in cm⁻³ determined from the intensity ratio of the [OII] doublet.

^c Measured intensity ratio OIII(λ 1661 + λ 1666)/[OIII] λ 5008.

^d Ratio [OIII](λ 4960 + λ 5008)/[OII](λ 3726 + λ 3729).

^e Ratio of [OIII](λ 4960 + λ 5008)/H β .

^f Inferred oxygen abundance from the direct T_e method.

^g Line intensity ratios (other than H α /H β) corrected for nebular extinction assuming E(B - V)_{neb} = 0.18 and the Cardelli et al. (1989) attenuation relation.

Table 5
 $z \sim 2.3$ Binned Mass-Metallicity Relation

Bin Range log(M _* /M _⊙)	Median log(M _* /M _⊙)	N _{gal}	12 + log(O/H) _{N2} ^a (PP04)	12 + log(O/H) _{O3N2} ^b (PP04)
8.6 – 9.0	8.86	5	8.26 ± 0.11	8.15 ± 0.09
9.0 – 9.5	9.32	27	8.39 ± 0.02	8.25 ± 0.02
9.5 – 9.7	9.60	27	8.39 ± 0.02	8.28 ± 0.02
9.7 – 9.9	9.79	27	8.38 ± 0.03	8.25 ± 0.02
9.9 – 10.2	10.06	34	8.42 ± 0.02	8.29 ± 0.02
10.2 – 10.5	10.36	21	8.46 ± 0.03	8.33 ± 0.03
10.5 – 11.0	10.70	29	8.54 ± 0.03	8.39 ± 0.03
11.0 – 11.6	11.17	6	8.61 ± 0.04	8.47 ± 0.03

^a Bi-weight mean and associated uncertainty of oxygen abundance determined using the PP04 N2 calibration.

^b Bi-weight mean and associated uncertainty of oxygen abundance determined using the PP04 O3N2 calibration.