WASP-5b: a dense, very-hot Jupiter transiting a 12th-mag Southern-hemisphere star


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ABSTRACT

We report the discovery of WASP-5b, a Jupiter-mass planet orbiting a 12th-mag G-type star in the Southern hemisphere. The 1.6-d orbital period places WASP-5b in the class of very-hot Jupiters and leads to a predicted equilibrium temperature of 1750 K. WASP-5b is the densest of any known Jovian-mass planet, being a factor seven denser than TrES-4, which is subject to similar stellar insolation, and a factor three denser than WASP-4b, which has a similar orbital period. We present transit photometry and radial-velocity measurements of WASP-5 (= USNO-B1 0487-0799749), from which we derive the mass, radius and density of the planet: $M_P = 1.58 \pm 0.13 M_J$, $R_P = 1.090 \pm 0.054 R_J$ and $\rho_P = 1.22 \pm 0.19 \rho_J$. The orbital period is $P = 1.6284296 \pm 0.0000048$ d and the mid-transit epoch is $T_C (HJD) = 2454375.62466 \pm 0.000026$.

Key words: planetary systems – stars: individual (WASP-5)

1 INTRODUCTION

Several wide-angle, ground-based surveys for transiting extrasolar planets have now been successful, including HAT (Bakos et al. 2002), TrES (O’Donovan et al. 2004), WASP (Pollacco et al. 2004) and XO (McCullough et al. 2005). In the Southern hemisphere relatively faint ($V = 16-17$) transiting systems have been found by OGLE (Udalski et al. 1992; Pont et al. 2007) and by Weldrake et al. (2007). The WASP-South survey is the first to find much brighter transiting systems in the South (Wilson et al. 2007). We report here the second planet found by the WASP-South and CORALIE collaboration, a dense, very-hot Jupiter transiting a $V = 12.3$ star with an orbital period of 1.6 d.

Transit searches are most sensitive to large, Jupiter-sized planets in very short orbits. WASP-5b, like the recently announced WASP-3b and WASP-4b, has an ultra-short period below 2 d (Pollacco et al. 2007; Wilson et al. 2007). We discuss the characteristics of these ultra-close, highly irradiated planets.
2 OBSERVATIONS

The WASP-South observatory is sited at the Sutherland station of the South African Astronomical Observatory. It consists of an array of 8 cameras, each with a Canon 200-mm f/1.8 lens backed by an c2v 2k×2k CCD, covering 7.8°×7.8°. WASP-South started operating in May 2006 with a strategy of tiling up to eight pointings with a cadence of 5–10 mins, taking two 30-s exposures at each pointing. The WASP hardware, observing strategy, data-reduction pipeline and archive are described in detail by Pollacco et al. (2006). The CORALIE spectra, when co-added, give a signal-to-noise ratio of ∼40, which is suitable for a preliminary photospheric analysis of WASP-5. The analysis

![Figure 1. Discovery light curves for WASP-5 obtained with WASP-South camera 7 (top) and WASP-South camera 8 (bottom, offset by 0.3 mag for clarity).](image1)

![Figure 2. The R-band light curve obtained with the Euler telescope (top) and the SDSS i'-band light curve obtained with Faulkes Telescope South (bottom, offset by 0.025 mag for clarity). In each case the solid line is the optimal MCMC solution (see 4).](image2)

<table>
<thead>
<tr>
<th>BJD−2,400,000</th>
<th>RV</th>
<th>σ_RV</th>
<th>BS^a</th>
</tr>
</thead>
<tbody>
<tr>
<td>54359.614570</td>
<td>19.785</td>
<td>0.016</td>
<td>−0.023</td>
</tr>
<tr>
<td>54362.654785</td>
<td>19.942</td>
<td>0.015</td>
<td>0.033</td>
</tr>
<tr>
<td>54364.676219</td>
<td>19.753</td>
<td>0.021</td>
<td>0.026</td>
</tr>
<tr>
<td>54365.682733</td>
<td>20.203</td>
<td>0.013</td>
<td>0.015</td>
</tr>
<tr>
<td>54372.781658</td>
<td>19.750</td>
<td>0.013</td>
<td>−0.047</td>
</tr>
<tr>
<td>54374.817984</td>
<td>20.016</td>
<td>0.040</td>
<td>0.091</td>
</tr>
<tr>
<td>54376.714797</td>
<td>20.248</td>
<td>0.014</td>
<td>−0.024</td>
</tr>
<tr>
<td>54377.762440</td>
<td>19.747</td>
<td>0.013</td>
<td>−0.002</td>
</tr>
<tr>
<td>54379.627523</td>
<td>19.907</td>
<td>0.014</td>
<td>−0.060</td>
</tr>
<tr>
<td>54380.682557</td>
<td>19.830</td>
<td>0.010</td>
<td>−0.024</td>
</tr>
<tr>
<td>54387.644986</td>
<td>19.818</td>
<td>0.015</td>
<td>0.027</td>
</tr>
</tbody>
</table>

^a Bisector spans; \( \sigma_{BS} = 2 \sigma_{RV} \)

We used a line-bisector analysis to look for asymmetries in the spectral line profiles, as could be caused by contamination from an unresolved eclipsing binary. Queloz et al. (2001), Mandushev et al. (2003). Such a binary would produce bisector spans that vary in phase with the photometric period with an amplitude comparable to the radial-velocity amplitude. This is not seen in our data (Fig. 3 bottom panel), supporting the conclusion that the radial-velocity variations are due to a planet.

3 STELLAR PARAMETERS

The CORALIE spectra, when co-added, give a signal-to-noise ratio of ∼40, which is suitable for a preliminary photospheric analysis of WASP-5. The analysis
Table 2. Stellar parameters for WASP-5

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>R.A. (J2000)</td>
<td>23°57′23.74″</td>
</tr>
<tr>
<td>Dec. (J2000)</td>
<td>−41°16′37″5″</td>
</tr>
<tr>
<td>(T_{\text{eff}}) (K)</td>
<td>5700 ± 150</td>
</tr>
<tr>
<td>(\log g) (ergs)</td>
<td>4.3 ± 0.2</td>
</tr>
<tr>
<td>([\text{M}/\text{H}])</td>
<td>0.0 ± 0.2</td>
</tr>
<tr>
<td>(v\sin i) (km s(^{-1}))</td>
<td>3.4 ± 0.7</td>
</tr>
<tr>
<td>(\log A(\text{Li}))</td>
<td>&lt; 1.0</td>
</tr>
<tr>
<td>(M_\star) (M(_\odot))</td>
<td>0.99 ± 0.08</td>
</tr>
<tr>
<td>(R_\star) (R(_\odot))</td>
<td>0.97 ± 0.06</td>
</tr>
<tr>
<td>Spec. Type</td>
<td>G4V ± 2 subtypes</td>
</tr>
<tr>
<td>Age (Gyr)</td>
<td>1.7–4.4</td>
</tr>
<tr>
<td>Distance (pc)</td>
<td>300 ± 50</td>
</tr>
</tbody>
</table>

was performed using the uclsyn spectral synthesis package and atlas9 models without convective overshooting (Castelli, Gratton & Kurucz 1997). The \(H_\alpha\), Na i D and Mg i b lines were used as diagnostics of both \(T_{\text{eff}}\) and \(\log g\). The metallicity was estimated using the photospheric lines in the 6000–6200 Å region. The parameters obtained from this analysis are listed in Table 2.

We also used the GALEX NUV flux and magnitudes from the NOMAD, DENIS and 2MASS catalogues to estimate the effective temperature using the Infrared Flux Method (IRFM) (Blackwell & Shallis 1977). This gives \(T_{\text{eff}} = 5610 ± 250\) K, which is in close agreement with the spectroscopic analysis. These results imply a spectral type in the range G2V–G6V. The Li i 6708 Å line is not present in the coadded spectrum, giving an upper limit on the lithium abundance of \(\log A(\text{Li}) \lesssim 1.2\). This implies an age of \(\gtrsim 2\) Gyr for a star of this effective temperature (Sestito & Randich 2005). The value of \(v\sin i\) = 3.4 km s\(^{-1}\) obtained from the CORALIE spectra indicates an age of 1.7–4.4 Gyr (Bouvier 1997). Comparison of the temperature and \(\log g\) with the stellar evolution models of Girardi et al. (2000) gives maximum-likelihood values of \(M_\star = 0.99 ± 0.08\) M\(_\odot\) and \(R_\star = 0.97 ± 0.06\) R\(_\odot\). The distance of WASP-5 (\(d = 300 ± 50\) pc) was calculated using the distance modulus, the NOMAD apparent visual magnitude (\(V = 12.3\)) and the absolute visual magnitude of a G4V star (\(V = 4.9\), Gray 1992). We assumed \(E(B-V) = 0\), which is reasonable at such a distance and in a direction out of the galactic plane.

4 SPECTROSCOPIC ORBITAL SOLUTION & PLANETARY PARAMETERS

The CORALIE radial-velocity measurements were combined with the WASP-South and EulerCAM photometry in a simultaneous Markov-chain Monte-Carlo (MCMC) analysis to find the system parameters. This process is described in detail in Collier Cameron et al. (2007) and Pollacco et al. (2007). The FTS data were not used owing to their higher red noise. The best-fitting solution to the radial-velocity data gave an eccentricity of \(e = 0.03 ± 0.03\). Our adopted MCMC solution fixes the eccentricity at \(e = 0\), which is expected for such a short period planet.

In optimising the MCMC solution we used constraints (Gaussian priors) on the stellar parameters, setting \(M_\star = 0.99 ± 0.08\) M\(_\odot\) and \(\log g_\star = 4.30 ± 0.20\). In order to balance the weights of the photometry and radial velocities in the MCMC analysis, we added a systematic error of 7 m s\(^{-1}\) to the radial velocities (as might arise, for example, from stellar activity) to reconcile \(\chi^2\) with the number of degrees of freedom; the error bars in Fig. 3 do not incorporate this error. The optimal MCMC parameters are shown in Table 3.

5 DISCUSSION

The discovery of WASP-3b (Pollacco et al. 2007), WASP-4b (Wilson et al. 2007) and WASP-5b, all transiting, Jupiter-mass planets with orbital periods of less than 2 d, increases the number of such systems known from four to seven. It is notable that they show a large disparity in density (Fig. 4). WASP-5b is the densest of these systems, whereas WASP-4b is the least dense, with a density a third that of WASP-5b. This illustrates that planets with a similar orbital period and radiation environment must still have different compositions or past histories.

Figure 3. Top: The radial-velocity measurements of WASP-5 obtained with CORALIE. The centre-of-mass velocity has been subtracted. The solid line is the MCMC solution (1), it includes the Rossiter–McLaughlin effect resulting from a \(v\sin i\) of 3.4 km s\(^{-1}\), as obtained from the CORALIE spectra. Middle: The radial-velocity residuals about the MCMC fit. Bottom: Bisector spans for the 11 CORALIE spectra.

1 http://www.inscience.ch/transits/
Table 3. System parameters for WASP-5

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>$P$ (days)</td>
<td>1.6284296</td>
</tr>
<tr>
<td>$T_C$ (HJD)</td>
<td>2454375.62466</td>
</tr>
<tr>
<td>$T_{14}$ (days)$^a$</td>
<td>0.0979</td>
</tr>
<tr>
<td>$R_p^2/R_\star^2$</td>
<td>0.01192</td>
</tr>
<tr>
<td>$b = a \cos i/R_\star$ (R$_\odot$)</td>
<td>&lt; 0.46</td>
</tr>
<tr>
<td>$i$ (days)</td>
<td>&gt; 85.0</td>
</tr>
<tr>
<td>$e$</td>
<td>0 (adopted)</td>
</tr>
<tr>
<td>$K_1$ (km s$^{-1}$)</td>
<td>0.2778</td>
</tr>
<tr>
<td>$\gamma$ (km s$^{-1}$)</td>
<td>20.0105</td>
</tr>
<tr>
<td>$M_\star$ (M$_\odot$)</td>
<td>0.972</td>
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<tr>
<td>$R_\star$ (R$_\odot$)</td>
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<tr>
<td>$T_{eff}$ (K)</td>
<td>5883</td>
</tr>
<tr>
<td>$\log g_\star$ (cgs)</td>
<td>4.403</td>
</tr>
<tr>
<td>$M_p$ (M$_J$)</td>
<td>1.58</td>
</tr>
<tr>
<td>$R_p$ (R$_J$)</td>
<td>1.090</td>
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<tr>
<td>$\rho_\star$ ($\rho_\odot$)</td>
<td>1.22</td>
</tr>
<tr>
<td>$a$ (AU)</td>
<td>0.02683</td>
</tr>
<tr>
<td>$\log g_\star$ (cgs)</td>
<td>3.484</td>
</tr>
<tr>
<td>$T_J$ (K)</td>
<td>1753</td>
</tr>
<tr>
<td>$v$ sin $i$ (km s$^{-1}$)</td>
<td>8.50</td>
</tr>
</tbody>
</table>

$^a$ $T_{14}$: transit duration, the time between 1st and 4th contact

For example, WASP-5b will receive the same insolation and thus has a similar estimated equilibrium temperature (1753 K) to TrES-4 (1720 K), which has a longer orbit about a more luminous star [Mandushev et al. 2007]. Despite this, WASP-5b is denser by a factor of 7 (1.22 $\rho_\odot$ compared with 0.167 $\rho_\odot$ for TrES-4) [Burrows et al. 2003] postulate that the bloating of the under-dense planets could be due to an enhanced atmospheric opacity, which could be absent in WASP-5b.

Although it is denser than other planets with a similar orbital period, WASP-5b is still within the expected theoretical range [Fortney, Marley & Barnes 2007]; depending on the age of the system, a core mass of $\sim$50–100 M$_\oplus$ is expected, with a younger planet requiring a more massive core. However, if the bloating of some planets were due to an extra energy source that is common to all hot Jupiters, then the core of a dense planet such as WASP-5b would have to be more massive to compensate [Fortney et al. 2004; Guillot et al. 2000]. If the core mass is toward the higher end of the estimate and the metallicity is no more than solar then there could be a discrepancy with the relationship between core mass and stellar metallicity postulated by [Burrows et al. 2003]. High resolution spectra with better S/N are required to reduce the uncertainties in the metallicity and age.

WASP-5b ($\log g_\star = 3.48$) has the highest surface gravity of planets with orbital periods less than 5 d. Thus it could be a good system for studying evaporation as a function of surface gravity. In comparison HD 290458b is heavily bloated ($\log g_\star = 2.97$) and appears to be undergoing evaporation [Vidal-Madjar et al. 2003; Ben-Jaffel 2005; Vidal-Madjar et al. 2008]. Given its high equilibrium temperature, WASP-5b is also expected to have a prominent secondary transit, which should be detectable with Spitzer (e.g. Harrington et al. 2007). Further, with the moderately bright host star having a moderate rotation rate of $v \sin i =$ 3.4 km s$^{-1}$, the Rossiter–McLaughlin effect (40 ± 8 m s$^{-1}$) should be detectable [Gaudi & Winn 2003], allowing us to check the alignment between the planetary orbit and the stellar spin axis.

REFERENCES


Figure 4. Planet density versus orbital period for planets with periods less than 5 d. The dense, hot Neptune, GJ 436b, and under-dense TrES-4 are labelled; the five WASP planets are labelled numerically. The error bars on the densities were calculated from the uncertainties on the masses and radii, which were assumed to be uncorrelated. Data taken from [http://www.inscience.ch/transits/](http://www.inscience.ch/transits/).