ASTEROIDS IN GALEX: NEAR-ULTRAVIOLET PHOTOMETRY OF THE MAJOR TAXONOMIC GROUPS

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ABSTRACT

We present ultraviolet (UV) photometry (near-UV (NUV) band, 180–280 nm) of 405 asteroids observed serendipitously by GALEX from 2003 to 2012. All asteroids in this sample were detected by GALEX at least twice. Unambiguous visible-color-based taxonomic labels (C type versus S type) exist for 315 of these asteroids; of these, thermal-infrared-based diameters are available for 245. We derive NUV − V color using two independent models to predict the visual magnitude V at each NUV-detection epoch. Both V models produce NUV − V distributions in which the S types are redder than C types with more than 8σ confidence. This confirms that the S types’ redder spectral slopes in the visible remain redder than the C types’ into the NUV, this redness being consistent with absorption by silica-containing rocks. The GALEX asteroid data confirm earlier results from the International Ultraviolet Explorer, which two decades ago produced the only other sizeable set of UV asteroid photometry. The GALEX-derived NUV − V data also agree with previously published Hubble Space Telescope (HST) UV observations of asteroids 21 Lutetia and 1 Ceres. Both the HST and GALEX data indicate that NUV band is less useful than u band for distinguishing subgroups within the greater population of visible-color-defined C types (notably, M types and G types).

Key words: minor planets, asteroids: general – surveys

Supporting material: machine-readable table

1. INTRODUCTION

As in visible wavelengths, ultraviolet (UV) flux from asteroids is reflected sunlight. However, the steep drop in the solar spectrum shortward of ~300 nm (Figure 1) makes asteroids orders of magnitude fainter in the UV than in the visible. For this reason—as well as the strong UV absorption by atmospheric ozone—UV observations of asteroids typically employ the Hubble Space Telescope (HST) or specialized instruments on a space-mission payload physically closer to the asteroid. These constraints have generally prohibited large-sample demographic studies of asteroids in the UV.

Predating HST, the International Ultraviolet Explorer (IUE) targeted 45 asteroids from 1978 to 1992, producing what remains to date the largest published sample of near-UV (NUV) asteroid spectra (Roettger & Buratti 1994), specifically in the range of 230–325 nm. The IUE data show evidence of clustering, principally with respect to UV geometric albedo. This clustering becomes further evident when coarsely defined visible spectral type is included as a categorical parameter for each object (C, S and M types being the classes considered in the original work). Comparing the IUE-derived UV geometric albedos for each class with the visible light geometric albedos demonstrated that the S types, which are redder-colored in the visible (specifically, 400–800 nm), remain redder than C types into the NUV.

The IUE data, combined with previously measured visible spectra, suggested that asteroid reflectances over the entire NUV to visible wavelength range (200–800 nm) are generally consistent with those of silica-bearing rocks (e.g., Wagner et al. 1987). To first order, this trend is characterized by an increase in a rock’s reflectance at longer wavelengths, generally attributable to the decreased fraction of volume-scattered light, i.e., light which penetrates into the mineral grains. The intensity of volume-scattered light varies as exp(−k d/λ), where d is the grain size, k is the imaginary part of the index of refraction, and λ is the wavelength. Once refracted into the grains, volume-scattered light is subject to absorption by various (λ-dependent) interactions with the mineral’s crystalline structure. At sufficiently short wavelengths (i.e., well into the UV region), most incident light penetrates the grains and is absorbed, while the small amount of measured reflected light is predominantly scattered directly from the grain surface. Precise characterization of this transition between surface-dominated and surface plus volume-scattered reflectance—as well as the identification of any additional mineral-specific spectral features—is therefore useful for tying astronomical observations of asteroids to laboratory-measured analogs, including lunar and meteoritic samples.

In this work we aim to verify the IUE’s findings with a newer and larger sample of UV asteroid data from the GALEX, a NASA Small Explorer-class space telescope mission which from 2003 to 2012 conducted a UV imaging survey in a far-UV band (FUV, 130–190 nm) and a NUV band (180–280 nm). Approximately 2/3 of the sky was covered, with avoidance of bright stars and low galactic latitudes. Martin et al. (2005) discuss the extragalactic science program, while Morissey et al. (2005, 2007) discuss the on-orbit performance, survey calibration and data products. GALEX has a 50 cm² effective area, 1.25 diameter circular field of view, and FWHM resolution of 4.5 in the NUV. Programs within the GALEX mission included an all-sky survey (AIS, with ~100 s exposures) and a medium-depth survey (MIS, with ~1500 s exposures), and also a spectroscopic (grism) survey. Figure 1 shows the photometric response functions of the two GALEX bandpasses multiplied by the solar spectrum, with comparison to the ugriz visible bandpasses. Detection of asteroids in the FUV is extremely unlikely (nonetheless, as described below we searched for both NUV and FUV asteroid detections).
methods for predicting the visual magnitude. The first method simply adopts the widely used MPC \(^4\) predicted magnitudes; the second method applies color-dependent phase-function and bond-albedo estimates adapted from the Waszczak et al. (2015) study of lightcurves from the Palomar Transient Factory survey\(^5\) (PTF; Law et al. 2009; Rau et al. 2009).

2. GALEX ASTEROID OBSERVATIONS

Extracting detections of known asteroids from a survey involves a three-dimensional (R.A., decl., time) cross-matching of ephemerides against the survey’s time-stamped image boundaries (e.g., Ofek 2012). We modified software originally used to search for asteroids in PTF (Waszczak et al. 2013, 2015) to instead search for asteroids in GALEX.

We first retrieved the metadata of all GALEX images, available from the Space Telescope Science Institute via command-line queries with the CasJobs tool (Li & Thakar 2008). We then indexed all image centers with respect to (R.A., decl.) into uniformly spaced sky cells of \(\theta_{\text{rad}}^2\) radius. For all \(\sim 380,000\) numbered asteroids, we queried JPL’s online service HORIZONS (Giorgini et al. 1996) to generate a 1-day-spaced ephemeris spanning 2003–2012. Using an object-specific search radius equal to \(3^\circ\) (cell radius) plus 0.75 (FOV radius) plus the object’s maximum 1-day motion (\(< 10\) arcmin for most main-belt objects), we matched the ephemeris points against the sky cells. For each matched cell, we filtered out all images in that cell not within the epoch range of the matched ephemeris points, then for each surviving image we re-queried HORIZONS for the precise location at each observed epoch. We next performed a 1.25-radial match of these precise positions against the relevant GALEX image centers.

We found \(\sim 850,000\) predicted detections of numbered asteroids (with no limit on apparent magnitude) in GALEX using this method. For each predicted detection, using CasJobs we queried the GALEX single-visit source list (as opposed to the co-added source list). Multiple matches near the same point occurring more than 6 hr apart were excluded, as were all matches further than 2° from the predicted location. Additionally, to ensure the inclusion of greater than (approximately) 5\(\sigma\) detections, we discarded all matches with \(\text{NUV} > 21\) mag in the shorter exposures (AIS program), and discarded all matches with \(\text{NUV} > 22.7\) mag in the longer exposures (MIS program), following the limiting magnitudes quoted by Morissey et al. (2007).

Following the above procedure and criteria, we extracted a total of 1342 positive NUV detections of 405 unique asteroids which were detected by GALEX at least twice (and no FUV detections, as expected). These detections are listed in Table 1; several histograms detailing these detections appear in Figure 2.

All of the GALEX-observed asteroids are in the main-belt; the sample includes no near-Earth or outer-solar system objects.

3. MODELING VISIBLE MAGNITUDES

In this section we consider two distinct methods of estimating the visible magnitudes corresponding to all GALEX NUV detections; this in turn provides the distribution of the asteroids’ NUV – V color. The general model for an asteroid’s

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\(4\) IAU Minor Planet Center, \(\text{http://minorplanetcenter.net}\).

\(5\) \(\text{http://ptf.caltech.edu}\).
Table 1

Observations of Asteroids Detected in GALEX NUV Images

<table>
<thead>
<tr>
<th>Asteroid Number</th>
<th>Observation Date (UT)</th>
<th>R.A. (deg)</th>
<th>decl. (deg)</th>
<th>Position Residual (&quot;)</th>
<th>NUV Mag</th>
<th>NUV Mag Uncertainty</th>
<th>MPC-predicted Visible Mag (V_{MPC})</th>
<th>Exposure Time (s)</th>
<th>Unique Database ID (objID Key in CasJobs)</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>2011 Oct 12.65797</td>
<td>355.22180</td>
<td>−18.47266</td>
<td>0.5</td>
<td>14.38</td>
<td>0.01</td>
<td>8.0</td>
<td>91</td>
<td>6380556162844065792</td>
</tr>
<tr>
<td>1</td>
<td>2011 Oct 21.42095</td>
<td>353.95617</td>
<td>−18.39556</td>
<td>0.2</td>
<td>14.67</td>
<td>0.01</td>
<td>8.1</td>
<td>80</td>
<td>6380556163951362048</td>
</tr>
<tr>
<td>3</td>
<td>2005 Dec 26.72536</td>
<td>75.12144</td>
<td>−1.31299</td>
<td>1.4</td>
<td>14.43</td>
<td>0.01</td>
<td>7.8</td>
<td>80</td>
<td>6381858059773280256</td>
</tr>
<tr>
<td>3</td>
<td>2011 Apr 17.04995</td>
<td>166.80140</td>
<td>8.52138</td>
<td>1.4</td>
<td>16.30</td>
<td>0.01</td>
<td>9.8</td>
<td>1513</td>
<td>38553297770719936512</td>
</tr>
<tr>
<td>6</td>
<td>2006 Aug 29.56208</td>
<td>309.25038</td>
<td>−19.31464</td>
<td>0.9</td>
<td>14.78</td>
<td>0.01</td>
<td>8.4</td>
<td>112</td>
<td>6379782093773209600</td>
</tr>
<tr>
<td>6</td>
<td>2005 May 7.46079</td>
<td>204.16985</td>
<td>11.58510</td>
<td>0.4</td>
<td>16.67</td>
<td>0.02</td>
<td>10.2</td>
<td>112</td>
<td>6378656257217134592</td>
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<tr>
<td>8</td>
<td>2004 Dec 21.41462</td>
<td>122.52225</td>
<td>19.17684</td>
<td>0.7</td>
<td>15.63</td>
<td>0.02</td>
<td>9.0</td>
<td>92</td>
<td>6377776615736213504</td>
</tr>
<tr>
<td>8</td>
<td>2004 Dec 21.48313</td>
<td>122.51009</td>
<td>19.18345</td>
<td>0.1</td>
<td>15.60</td>
<td>0.02</td>
<td>9.0</td>
<td>87</td>
<td>637777661576976936</td>
</tr>
</tbody>
</table>

Note. Includes 1342 detections of 405 asteroids detected at least twice.
(This table is available in its entirety in machine-readable form.)
apparent visual magnitude \( V \) (log flux) is

\[
V = H + \delta + 5 \log_{10}(r\Delta) - 2.5 \log_{10}(|\phi(\alpha)|),
\]

where \( H \) is the absolute magnitude (a constant), \( \delta \) is a periodic variability term due to rotation (e.g., if the object is spinning and has some asymmetry in shape or albedo), \( r \) and \( \Delta \) are the heliocentric and geocentric distances (in AU), and \( \phi = \phi(\alpha) \) is the phase function, which varies with the solar phase angle \( \alpha \) (the Sun–asteroid–Earth angle). When \( \alpha = 0 \) (i.e., at opposition), \( \phi = 1 \) by definition, while in general \( 0 < \phi < 1 \) for \( \alpha > 0 \) (with \( \phi \) decreasing as \( \alpha \) increases).

All asteroids for which we have extracted GALEX observations have known orbits, meaning \( r, \Delta, \) and \( \alpha \) are accurately and precisely known at all observed epochs. Our two methods for estimating \( V \) differ in their assumptions regarding (and observational data used to constrain) \( H \) and \( \phi \). In both cases we do not attempt to model the rotational term \( \delta \), but rather incorporate \( \delta \) into the uncertainty of \( V \) using lightcurve amplitude estimates from the literature. In particular, 388 of the 405 GALEX-observed asteroids have an amplitude lower-limit estimate available in the Lightcurve Database (Warner et al. 2009; Harris et al. 2012).

In the following sections we refer to two different albedo quantities. The visible-band geometric albedo \( p_V \) relates to the visible-band bond albedo \( A_{\text{bond}} \) and the phase function \( \phi \) (of Equation (1)) according to

\[
p_V = \frac{A_{\text{bond}}}{2} \left( \int_0^\pi \phi(\alpha) \sin(\alpha) \, d\alpha \right)^{-1} = \frac{A_{\text{bond}}}{q},
\]

The above equation also defines the phase integral \( q \). The bond albedo \( A_{\text{bond}} \) is defined as the total visible light energy reflected or scattered by the asteroid (in all directions) divided by the total visible light energy incident upon the asteroid (from the Sun). Assuming the asteroid has a circular cross-section of diameter \( D \), this can be expressed as

\[
A_{\text{bond}} = \frac{\int_0^\pi f(\alpha) \sin(\alpha) \, d\alpha}{\left(\frac{f_{\text{Sun}}}{4\pi A^2} \times \pi(D/2)^2\right)},
\]

where \( f(\alpha) = 10^{-V(\alpha)/2.5} \) is the asteroid’s flux as a function of phase angle, with \( V(\alpha) = H - 2.5 \log_{10}|\phi(\alpha)| \) being Equation (1) evaluated at \( \delta = 0 \) and \( r = \Delta = 1 \) AU (similarly, \( f_{\text{Sun}} = 10^{-V_{\text{Sun}}(2.5)} \)).

### 3.1. \((H, G)\) from MPC Data

The first method for estimating \( V \) adopts the Minor Planet Center’s computed absolute magnitudes \((H_{\text{MPC}})\), which are regularly updated by the MPC’s automated processes and utilize the Lumme–Bowell \( G \)-parameter model for \( \phi \) (Bowell et al. 1989, pp. 524–556). This same \((H, G)\) model then predicts the apparent magnitude \( V_{\text{MPC}} \) as a function of solar phase angle.

The \( H_{\text{MPC}} \) values are fit to photometry provided by a variety of surveys/individuals, many of whom may use slightly different absolute calibration standards or filters with slightly different specifications. A small fraction of asteroids have fitted \( G \) values; Harris & Young (1988) present mean \( G \) values for several major taxonomic classes, with \( G = 0.15 \) being an average between the C types \((G \approx 0.08)\) and the S types \((G \approx 0.23)\). For the majority of asteroids the MPC uses an assumed \( G = 0.15 \) with this model. Waszczak et al. (2015) compares the \( H_{\text{MPC}} \) values with \( H \) magnitudes derived from a model that includes rotation and the more modern \((H, G_{12})\) phase function of Muinonen et al. (2010). Among bright asteroids the relative difference is typically between 0.3% and 3%, corresponding to (on average) an \( \sim 0.07 \) mag discrepancy.

Though \( H_{\text{MPC}} \) values are available for all 405 GALEX-observed asteroids, we only consider the subset of 315 asteroids having visible-band color indices of either less than 0.25 (“C types”) or greater than 0.75 (“S types”). Of these, 41 asteroids have \( G_{\text{MPC}} \approx 0.15 \).
3.2. $(D, A_{\text{bond}}, G_{12})$ from PTF, Infrared, and Color Data

Our second means of estimating visual magnitudes applies only to asteroids having both a color index and a diameter estimate constrained from thermal fluxes in an infrared survey. In this approach we use the $G_{12}$-parameter model for $\phi$ (Muinonen et al. 2010), and we replace $H$ with its equivalent expression$^{7}$ in terms of the diameter $D$, bond albedo $A_{\text{bond}}$, and phase integral $q$:

$$H = -5 \log_{10} \left( \frac{D^2 A_{\text{bond}}/q}{1329 \text{km}} \right),$$

where the phase integral $q$ is a linear function of $G_{12}$:

$$q(G_{12}) = \begin{cases} 0.2707 - 0.236G_{12} & \text{if } G_{12} < 0.2; \\ 0.2344 - 0.054G_{12} & \text{otherwise}. \end{cases}$$

We again define “C types” as all asteroids with color indices less than 0.25 and “S types” as all with color indices greater than 0.75. For $S$ types we then consider diameters derived from any of four infrared surveys (see Footnote 6), while for C types we specifically require that the asteroid have been observed in the WISE 4-band cryogenic survey (Wright et al. 2010; Masiero et al. 2014 and references therein). Both the WISE W1-band geometric albedo $p_{W1}$ and the PTF-derived bond albedo$^{8}$ $A_{\text{bond}}$ show evidence of bimodality among objects with color indices less than 0.25 (Figure 3 panel (D)). Thus, we divide the C types into low-bond-albedo ($C_{\text{low}}$) and high-bond-albedo ($C_{\text{high}}$) subgroups based on their $p_{W1}$ as reported by Masiero et al. (2014). In Section 6 we show that the $C_{\text{high}}$ types most closely correspond to what other authors have called M types.

Waszczak et al. (2015) computed $A_{\text{bond}}$ and $G_{12}$ values for ~1600 asteroid lightcurves in the PTF survey. Using that work’s data (Figure 3) we compute median $A_{\text{bond}}$ and $G_{12}$ values (and associated scatter) for the $S$, $C_{\text{low}}$ and $C_{\text{high}}$ taxonomic groups. Table 2 summarizes the definitions and assumed $A_{\text{bond}}$ and $G_{12}$ values of these groups. There are 245 WISE-observed asteroids (out of the 405 in Table 1) which have color and diameter data available, allowing them to be modeled by this method. To each WISE-observed asteroid we assign the appropriate $A_{\text{bond}}$ and $G_{12}$ value based on its class membership, then use its diameter to compute a model absolute magnitude ($H_{\text{PTF}}$) using Equation (4). Together with the assumed $G_{12}$ value, this $H_{\text{PTF}}$ then predicts the apparent magnitude $V_{\text{PTF}}$ at each WISE-observed solar phase angle.

3.3. Rotational Uncertainty in V

Both the $V_{\text{MPC}}$ and $V_{\text{PTF}}$ model magnitudes discussed here lack an estimate of the rotational term ($\delta$ in Equation (1)). We account for this by incorporating a term for rotational modulation into the reported uncertainty of $V$. Of the 315 asteroids with $V_{\text{MPC}}$ values, 302 have an amplitude lower limit listed in the Lightcurve Database (Warner et al. 2009; Harris et al. 2012), while for the 245 asteroids with $V_{\text{PTF}}$ predictions there are 239 with reported amplitudes. As shown for instance by Waszczak et al. (2015), asteroids in the relevant size range typically have amplitudes less than ~0.4 mag. For the few objects in our sample lacking an amplitude limit, we assume a value of 0.2 mag.

Assuming an asteroid’s rotational phase $\phi$ to be random at the time of a GALEX detection (i.e., with a probability distribution of the form $P(\phi) \propto \text{constant}$), then the probability distribution of a basic sinusoidal $\delta$ (i.e., one of the form $\delta = \delta_0 \sin \phi$) can be shown to have the form

$$P(\delta) \propto \frac{1}{\sqrt{\delta_0^2 - \delta^2}},$$

where $\delta_0$ is the amplitude. We use Equation (6) as a probability density function to generate, for each modeled $V$, a set of $10^5$

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**Table 2**

<table>
<thead>
<tr>
<th>Name</th>
<th>Color Index</th>
<th>$G_{12}$</th>
<th>$A_{\text{bond}}$ Median</th>
<th>$A_{\text{bond}}$ Scatter</th>
<th>$G_{12}$ Median</th>
<th>$G_{12}$ Scatter</th>
</tr>
</thead>
<tbody>
<tr>
<td>S</td>
<td>&gt;0.75</td>
<td>N.A.</td>
<td>0.056</td>
<td>0.016</td>
<td>0.36</td>
<td>0.16</td>
</tr>
<tr>
<td>$C_{\text{high}}$</td>
<td>&lt;0.25</td>
<td>&gt;0.125</td>
<td>0.038</td>
<td>0.022</td>
<td>0.42</td>
<td>0.20</td>
</tr>
<tr>
<td>$C_{\text{low}}$</td>
<td>&lt;0.25</td>
<td>&lt;0.125</td>
<td>0.010</td>
<td>0.003</td>
<td>0.84</td>
<td>0.16</td>
</tr>
</tbody>
</table>

$^a$ Scatter is defined here as $0.5 \times (84\text{th percentile} - 16\text{th percentile}).$
simulated $\delta$ values. These simulated $\delta$ are added to an equal number of model $V$ magnitudes computed by random (Gaussian distribution) sampling of the component terms: in the case of $V_{\text{MPC}}$ we just assume a fixed $H_{\text{MPC}}$ uncertainty of 0.1 mag, whereas for the $V_{\text{PTF}}$ values we randomly sample all three of $A_{\text{bond}}$, $G_{12}$, and $D$, using the scatter values in Table 2 for the first two and the literature-reported diameter uncertainty for $D$. The 16th to 84th percentile spread in the distribution of combined $\delta + V$ values then becomes the quoted uncertainty for $V$.

4. NUV − $V$ COLOR DISTRIBUTION

Having computed the model $V$ magnitudes, we obtain the NUV − $V$ color for each GALEX asteroid detection and the corresponding uncertainty. The latter contains an additional rotational uncertainty component (now associated with the NUV observation), again determined by repeated sampling of Equation (4) as described above. Since all the asteroids we consider have more than one GALEX NUV detection, we compute the inverse-variance-weighted average NUV − $V$ color for each asteroid (plotted in Figures 4 and 6); the uncertainty in this average is the inverse quadrature sum of the individual uncertainties.

In Figures 4 and 6 (and the accompanying analysis) we have omitted all asteroids with NUV − $V$ uncertainties of greater than 0.5 mag. As a result, the sample size of asteroids with NUV − $V_{\text{MPC}}$ estimates is 297 (out of the 315 quoted in Section 3.1), while the sample with NUV − $V_{\text{PTF}}$ estimates is 223 (out of the 245 quoted in Section 3.2). Figure 5 graphically summarizes the sample selection criteria in a flowchart. In Figures 4, 6 and 7, the errorbars on the color indices were computed by a bootstrapping process described in the appendix of Waszczak et al. (2015).

Both the $V_{\text{MPC}}$ and $V_{\text{PTF}}$ model magnitudes produce a bimodal NUV − $V$ color distribution, with the $S$ types having the redder NUV − $V$ color (panels (A) and (B) of both Figures 4 and 6). Median and scatter of NUV − $V$ for the various classes appear in Table 3. To formally ascertain the inequality of the two distributions, we use the two-sided Kolmogorov–Smirnov (KS) test (Massey 1951), which compares two empirical distributions. In particular this test computes a statistic quantifying the extent to which the cumulative distribution function differs in the two
distributions being compared. For the VMPC model we find the C-type NUV – V color distribution differs from that of the S-type distribution at an 11.6σ significance level (Figure 4 panel (B)). For the VPTF model (Figure 6 panel (B)) we find the C types (C_{low} and C_{high} combined) differ from the S types at an 8.1σ level, while the C_{low} and C_{high} types only differ at a 1.9σ level (this difference is thus not statistically significant).

An important characteristic of our sample is that the C types outnumber the S types by a ratio of 3:1 in the VMPC sample and a ratio of 2:1 in the VPTF sample (cf. panel (C) of Figures 4 and 6). This ratio is a combination of (1) the inherent difference in the population sizes of the two types (above a given diameter cut-off), a detection bias due to S types dominating the inner main-belt and thus typically having brighter apparent
magnitudes for a given size and albedo, and (3) the difference in the S and C types’ NUV albedo (discussed in Section 5).

In panels (C)–(F) of Figure 4 the sample size decreases from \( N = 297 \) down to \( N = 223 \) asteroids as we consider only those objects in the \( V_{\text{MPC}} \) sample that also have available diameters (this is equivalently the \( V_{\text{PTF}} \) sample considered in Figure 6). We compute the MPC-based visible bond albedo \( A_{\text{MPC}} \) using Equation (2) together with the asteroid’s \( H_{\text{MPC}} \) and \( G_{\text{MPC}} \) values. In particular, there are 38 asteroids (out of the 223 with diameters) with a measured \( G_{\text{MPC}} = 0.15 \); for the remainder we assume \( G_{\text{MPC}} = 0.15 \) for consistency with the manner in which the \( V_{\text{MPC}} \) are computed. Analogous to Equation (5), the phase integral for the \( G \)-model (required for computation of \( A_{\text{MPC}} \) via Equation (4)) is

\[
q(G) = 0.290 + 0.684G,
\]

as given by Bowell et al. (1989, pp. 524–556). With the \( G_{\text{MPC}} = 0.15 \) assumption for the majority of the asteroids in our sample, the \( A_{\text{MPC}} \) values are not expected to be as accurate as the \( A_{\text{bond}} \) values computed for instance by Waszczak et al. (2015), wherein distinct \( q \) values were fitted to each object on the basis of a lightcurve. Nonetheless, it is instructive to compute \( A_{\text{MPC}} \), e.g., to check for consistency with the class-median \( A_{\text{bond}} \) values, and to exploit as a second taxonomic metric in addition to visible color.

Figure 4 panel (E) shows that NUV − \( V \) correlates with \( A_{\text{MPC}} \) (\( \rho_{\text{Spearman}} = 0.698, >10\sigma \) significance), similar to how NUV − \( V \) correlates with the color index in panel (A) (\( \rho_{\text{Spearman}} = 0.491, >10\sigma \) significance). Unlike the color index however, the separation between the \( C_{\text{low}} \) and \( C_{\text{high}} \) subgroups is qualitatively evident in this plot. Figure 4 panel (F) combines all three parameters; note the axes are the same as Figure 3 panel (B), with \( A_{\text{MPC}} \) replacing \( A_{\text{bond}} \) and the data consisting of \( GALEX \)-observed asteroids rather than PTF-observed asteroids.

Figure 6 panel (F) confirms (independently of Figure 3 panel (D)) the validity of using \textit{WISE} \emph{W}1-band geometric albedo as a proxy for visible bond albedo to separate \( C_{\text{low}} \) from \( C_{\text{high}} \)—the two classes robustly differ in their \( A_{\text{MPC}} \) distributions (9.5\( \sigma \) KS-test significance). However, the class-median \( A_{\text{MPC}} \) values of the \( C_{\text{low}} \), \( C_{\text{high}} \), and \( S \) types are 100%, 67%, and 63% greater than their class-median PTF-based \( A_{\text{bond}} \) values in Table 2. This reflects the differing values of \( H \) and \( q \) produced by the \( G \) and \( G_{12} \) models, as well as the fact that we apply class-specific \( G_{12} \) values, whereas \( G_{\text{MPC}} = 0.15 \) is assumed for the majority of asteroids, regardless of their class.

Consideration of both the \( V_{\text{MPC}} \) and \( V_{\text{PTF}} \) model magnitudes provides two independent means of computing NUV − \( V \); this helps rule out the effect of potential systematic errors unique to either one of the \( V \) models, as well as possible biases in the distinct observational data sets upon which each \( V \) is based. In

\[\text{Table 3} \]

<table>
<thead>
<tr>
<th>Class</th>
<th>( \text{NUV} - V_{\text{MPC}} )</th>
<th>( \text{NUV} - V_{\text{PTF}} )</th>
</tr>
</thead>
<tbody>
<tr>
<td>S</td>
<td>( 6.52 \pm 0.25 )</td>
<td>( 6.71 \pm 0.21 )</td>
</tr>
<tr>
<td>C</td>
<td>( 5.90 \pm 0.19 )</td>
<td>( 6.03 \pm 0.22 )</td>
</tr>
<tr>
<td>( C_{\text{high}} )</td>
<td>( ... ... )</td>
<td>( 6.14 \pm 0.33 )</td>
</tr>
<tr>
<td>( C_{\text{low}} )</td>
<td>( ... ... )</td>
<td>( 6.02 \pm 0.19 )</td>
</tr>
</tbody>
</table>

\( ^\text{a} \) Scatter is defined here as \( 0.5 \times (84^{\text{th}} \text{ percentile} - 16^{\text{th}} \text{ percentile}) \).
Assume photometry from two filters (1 and 2) produce the color measurement $m_1 - m_2$. This color relates to the solar flux distribution $S(\lambda)$, the albedos in each band ($A_1$ and $A_2$) and the filter responses $F_1(\lambda)$ and $F_2(\lambda)$ according to

$$10^{(m_1 - m_2)/2.5} = \frac{\int F_1(\lambda)\lambda^{-2}d\lambda \int S(\lambda)F_2(\lambda)A_2 d\lambda}{\int S(\lambda)F_1(\lambda)A_1 d\lambda \int F_2(\lambda)\lambda^{-2}d\lambda}$$

(8)

which we adapted from a similar equation in Pickles (1998).

Using NUV as band 1 and V as band 2 in Equation (6), we use the colors in Table 3 (specifically, the $V_{\text{PTF}}$-based colors) to obtain the albedo ratio $A_{\text{NUV}}/A_V$, with uncertainties coming from the associated scatter in the colors. In Figure 7 we plot these albedo ratios for the C types and S types, incorporating an additional uncertainty component from the transformation from $r$ to $V$ (see Waszczak et al. 2015 for a discussion of this transformation in the context of asteroids). The end-computed values are $(A_{\text{NUV}}/A_r)_{C} = 0.63^{+0.14}_{-0.12}$ and $(A_{\text{NUV}}/A_r)_{S} = 0.33^{+0.07}_{-0.06}$. The relative albedo values in the SDSS bands included for comparison in Figure 7 are taken directly from Table 2 for V band. Assuming (from Table 2) that C types have $p_V = 0.054 \pm 0.019$, then $(p_{\text{NUV}})_{C} = 0.035 \pm 0.019$, and assuming S types have $p_V = 0.264 \pm 0.031$, then $(p_{\text{NUV}})_{S} = 0.088 \pm 0.028$. That is, S types have a higher mean NUV albedo than C types, similar to the visible.

6. COMPARISON TO HST DATA

6.1. Lutetia

Weaver et al. (2010) obtained HST photometry of asteroid 21 Lutetia in UV and visible bands. Lutetia has been classified by various authors as an M-type asteroid; in the context of this work its color index is 0.05 (making it a C type) and its $A_{\text{FPOC}} = 0.06$ suggest it to be a C$_{\text{high}}$ type in particular, though in this work’s system we formally would require a WISE $p_{\text{W1}}$ measurement to classify it as such. In the following section we show that M types (a group in the Tholen taxonomic system) and our C$_{\text{high}}$ types are largely the same population.

The HST Lutetia photometry revealed a steep drop in albedo around $\sim 300$ nm and nearly constant albedo in the $200-300$ nm region at a factor $\sim 0.6$ times the visible ($r$-band-equivalent) albedo. The HST observations of Lutetia thus generally agree with the C-type albedo trend (Figure 9), the main difference being the location of the UV albedo drop-off (the bluest two ECAS bands also demonstrate this difference between M types and C types, e.g., see Figure 2 of Bus et al. (2002, pp. 169–182). The Rosetta spacecraft’s flyby of Lutetia enabled FUV observations with the on-board Alice UV imaging spectrograph (Stern et al. 2011); the longest wavelengths of the FUV data ($\sim 190$ nm) yield an albedo consistent with the constant value measured in the $200-300$ nm range by HST, namely $\sim 0.6$ times that of the visible albedo.
6.2. Ceres

HST photometry of asteroid 1 Ceres has also been obtained in the UV and visible (Parker et al. 2002; Li et al. 2006). With a color index of 0.01, Ceres is also a C type in our classification scheme, though its \( A_{\text{MPC}} = 0.033 \) makes its placement in our \( C_{\text{low}} \) versus \( C_{\text{high}} \) groups ambiguous (see Figure 6 panel (F)). Like Lutetia, Ceres lacks a reported \( p_W \) so that we cannot formally classify it as either \( C_{\text{low}} \) or \( C_{\text{high}} \).

Ceres was observed by GALEX and thus is included in our MPC-data-based analysis; our measured \( \Delta_{\text{NUV}} - V_{\text{MPC}} = 6.45 \pm 0.19 \) for Ceres make it a clear outlier from the C-type NUV \( - V_{\text{MPC}} \) distribution (Figure 4 panel (B)). In the Tholen taxonomic system Ceres is classified as a G type; in the following section we show that other G types exhibit similarly high \( \Delta_{\text{NUV}} - V_{\text{MPC}} \) values but less anomalous NUV \( - V_{\text{PTF}} \).

The Parker et al. (2002) HST data show that around \( \sim 300 \text{ nm} \) Ceres’ albedo drops to as low as \( \sim 0.3 \) times the visible-band albedo—compared to the factor of \( \sim 0.6 \) seen for GALEX C types and the Lutetia data—but that around \( \sim 200 \text{ nm} \) it appears to rise again to a more typical C-type UV albedo. Roettger & Buratti (1994) did not observe this unusually deep absorption feature near 300 nm in their IUE spectrum of Ceres; if real this feature could partially explain the anomalous NUV \( - V_{\text{MPC}} \) we observe for G types in GALEX. Figure 9 shows Ceres data in the three \( HST \) bands observed by Li et al. (2006), none of which sample the 300 nm region containing the putative absorption band, though these three bands do generally match the GALEX C-type data.

7. C-TYPE SUBGROUPS

C types deserve further consideration for several reasons: (1) C types outnumber S types in the GALEX samples by a factor of several, (2) our division of C types into \( C_{\text{low}} \) and \( C_{\text{high}} \) merits interpretation in more conventional taxonomic systems, and (3) both of the \( HST \)-observed asteroids in the previous section are known members of C-type subgroups, the UV properties of which are worth confirming with additional group members.

Figures 10 and 11 detail the distribution of GALEX-observed asteroids belonging to six classes each from the Tholen and Bus/Binzel taxonomic systems (Tholen 1989, pp. 1139–1150; Bus & Binzel 2002), the latter is sometimes referred to as the SMASII system after the survey data with which it was derived. These two classification systems were created on the basis of different visible-band color data; a comparison of their group definitions is given in Table 1 of Bus et al. (2002, pp. 169–182). We consider only the subset of GALEX-observed asteroids having both \( V_{\text{MPC}} \) and \( V_{\text{PTF}} \) model magnitudes and omit subgroups containing less than three objects. In the following subsections we briefly comment on these subgroups.

One key interpretation of these data—supported also by the \( HST \) data in Figure 9—is that NUV-band albedo is not very useful for discriminating C-type subgroups, e.g., M types versus G types, whereas \( u \) band appears to be more diagnostically useful in this regard. The \( u \)-band discrepancy between these subgroups was remarked most notably by Zellner et al. (1998) in the ECAS data, but it was unknown at that time (indeed, up until now) whether the discrepancy in UV albedo became more or less pronounced shortward of \( \sim 300 \text{ nm} \). The NUV data indicate that the discrepancy lessens in the NUV, as M types do in fact exhibit a step-down in albedo (between NUV and \( u \) bands), similar to the step down the G types exhibit within \( u \) band.

7.1. X Complex

The Tholen system’s X-type group includes asteroids with relatively flat visible color, including no substantial absorption in the blue (in contrast to, e.g., the \( u \)-band drop-off seen in G types). The subgroups within the X group include E, M, and P types and are distinguishable only by albedo.

The twelve M types in our sample all have \( A_{\text{MPC}} > 0.03 \) and \( p_W < 0.125 \), the latter formally makes them all \( C_{\text{high}} \) types in this work’s classification system. The M types have NUV \( - V_{\text{MPC}} = 5.89 \pm 0.15 \) and NUV \( - V_{\text{PTF}} = 6.12 \pm 0.21 \), neither of which significantly differ from the C-type averages given in Table 3. This is consistent with the above-noted observation that M-type Lutetia’s NUV \( - V \) is similar to that of the GALEX C types, despite an obvious difference in \( u \)-band (Figure 9).

Assuming all 29 of the \( C_{\text{high}} \) types in the GALEX sample are in fact M types, then the \( C_{\text{high}} \) types’ slightly higher NUV \( - V_{\text{PTF}} = 6.14 \pm 0.33 \) (compared to NUV \( - V_{\text{PTF}} = 6.03 \pm 0.22 \) for the whole C type group) agrees well with the M types’ slightly higher average.

Complementary to the M types, the 11 P types in our sample all have \( A_{\text{MPC}} < 0.03 \) and \( p_W < 0.125 \), the latter formally makes them all \( C_{\text{low}} \) types. The P types have NUV \( - V_{\text{MPC}} = 5.74 \pm 0.17 \) and NUV \( - V_{\text{PTF}} = 5.91 \pm 0.13 \). These values are less than both models’ C-type averages as well as less than the \( C_{\text{low}} \) average, suggesting our \( C_{\text{low}} \) group includes more diverse objects than just P types (e.g., the five F types also all have \( A_{\text{MPC}} \) consistent with \( C_{\text{low}} \)).

There are 14 GALEX-observed asteroids listed simply as X types in the Tholen system (presumably because no visible albedo was available at the time of classification); Figure 10 shows that these are in fact distributed across both the \( C_{\text{low}} \) and \( C_{\text{high}} \) albedo ranges.

In the Bus/Binzel system, the X complex consists of four subgroups: Xc, Xk, X and Xe, these being differentiated by their spectral slope and presence of various absorption features. In the GALEX sample the most numerous of these are the Xc types, which have the least red visible color and seem to include both high and low visible albedo members. Both the Xe and Xk types have higher visible color indices (with larger uncertainties in the color). As with the Tholen X types, we see no systematic trends with respect to the NUV properties of these subgroups.

7.2. G Types

Three GALEX-observed asteroids are categorized as G types. Like G-type Ceres, these have intermediate \( A_{\text{MPC}} \) and an above-average NUV \( - V_{\text{MPC}} = 6.22 \pm 0.11 \). In contrast, however, the G-type NUV \( - V_{\text{PTF}} = 5.79 \pm 0.10 \) lies slightly below the C-type average. The reason for this discrepancy is that all three G types in this sample have \( p_W < 0.125 \) and so are formally classed as \( C_{\text{low}} \) objects, as a result their assumed \( A_{\text{bond}} = 0.01 \) in the computation of \( V_{\text{PTF}} \) may be too low. On the other hand, Oszkiewicz et al. (2011) fit \( G_{12} = 0.88 \pm 0.2 \) to Ceres’ phase function, suggesting that the assumed \( G_{12} = 0.84 \pm 0.14 \) for \( C_{\text{low}} \) types (Table 2) is a more valid assumption for G types than the \( C_{\text{high}} \) value of \( G_{12} = 0.42 \pm 0.20 \). Hence the G types seem not to fit well into either of our \( C_{\text{low}} \) or \( C_{\text{high}} \) groups, and hence are not accurately modeled by our \( V_{\text{PTF}} \).

The three-asteroid G-type sample’s higher than average NUV \( - V_{\text{MPC}} \) agrees with the Cere’s \( HST \) data (Figure 9), which as discussed above could be indicative of an absorption
feature at ∼300 nm unique to G types (Li et al. 2006), the precise shape and location of which remains unresolved in the broadband photometry considered here.

Tholen’s G types are represented in the Bus/Binzel system by the Cg and Cgh groups; however no asteroids in our GALEX sample have either of these SMASSII labels.

7.3. B Types

Members of the Tholen B and F classes, represented also by the Bus/Binzel B and Cb classes, all are classified as C_low types in the GALEX sample based on their $p_{W1}$. Unlike the G types, the B types are not anomalous in NUV − V_MPC, meaning the B
types likely lack the G types’ strong absorption at 300 nm. The B types also are characterized by slightly higher $A_{MPC} = 0.026$ compared to the C low average $A_{MPC} = 0.020$. Hence, like the G types, the B types show a lower than average NUV − $V_{PTF}$ symptomatic of an underestimated $A_{bond}$ and therefore too dim of a predicted $V_{PTF}$.

8. SUMMARY

We present NUV-band photometry of 405 asteroids observed serendipitously by GALEX from 2003 to 2012. Using a compilation of visible-band color data, we select the subset of these GALEX-observed asteroids belonging to the C-type or S-type classes. We then compute the visual-band magnitude...
### Table 4

Glossary of Acronyms and Symbols Used Repeatedly in This Work

<table>
<thead>
<tr>
<th>Acronym</th>
<th>Full Meaning/Description</th>
</tr>
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<tbody>
<tr>
<td>α</td>
<td>solar phase angle (angle drawn by a light ray as it travels from the Sun to an asteroid to the Earth)</td>
</tr>
<tr>
<td>$A_{\text{band}}$</td>
<td>visible-band bond albedo</td>
</tr>
<tr>
<td>$A_{\text{MPC}}$</td>
<td>visible-band bond albedo computed using $H$ and $G$ values from the MPC together with an infrared-derived diameter</td>
</tr>
<tr>
<td>$A_{\text{NUV}}$</td>
<td>near-ultraviolet-band bond albedo</td>
</tr>
<tr>
<td>$C_{\text{high}}$</td>
<td>types C-type asteroids with a high near-infrared albedo ($p_{\text{W1}} &gt; 0.125$), a group consisting almost exclusively of M types</td>
</tr>
<tr>
<td>$C_{\text{low}}$</td>
<td>types C-type asteroids with a low near-infrared albedo ($p_{\text{W1}} &lt; 0.125$), a group consisting of P types and many X types</td>
</tr>
<tr>
<td>$D$</td>
<td>asteroid diameter</td>
</tr>
<tr>
<td>ECAS</td>
<td>Eight-Color Asteroid Survey</td>
</tr>
<tr>
<td>FUV</td>
<td>far-ultraviolet band ($\lambda \leq 180$ nm)</td>
</tr>
<tr>
<td>$G$</td>
<td>an older photometric phase-function model parameter (Bowell et al. 1989)</td>
</tr>
<tr>
<td>$G_{12}$</td>
<td>a newer photometric phase-function model parameter (Muinonen et al. 2010)</td>
</tr>
<tr>
<td>G types</td>
<td>certain asteroids, including Ceres, that are a subgroup of the C-type asteroid taxonomic class</td>
</tr>
<tr>
<td>GALEX</td>
<td>Galaxy Evolution Explorer satellite</td>
</tr>
<tr>
<td>$H$</td>
<td>visible-band absolute magnitude ($V$ magnitude asteroid would have if observed 1 AU from both the Sun and Earth, at zero phase angle)</td>
</tr>
<tr>
<td>HST</td>
<td>Hubble Space Telescope</td>
</tr>
<tr>
<td>IUE</td>
<td>International Ultraviolet Explorer satellite</td>
</tr>
<tr>
<td>MPC</td>
<td>Minor Planet Center, <a href="http://minorplanetcenter.net">http://minorplanetcenter.net</a></td>
</tr>
<tr>
<td>NUV</td>
<td>near-ultraviolet band (180–200 nm), and/or measured magnitude in this band</td>
</tr>
<tr>
<td>PTF</td>
<td>Palomar Transient Factory survey</td>
</tr>
<tr>
<td>$p_{V}$</td>
<td>visible-band geometric albedo</td>
</tr>
<tr>
<td>$p_{\text{W1}}$</td>
<td>near-infrared (W1-band from WISE) geometric albedo</td>
</tr>
<tr>
<td>$\rho_{\text{Spearman}}$</td>
<td>Spearman’s correlation coefficient</td>
</tr>
<tr>
<td>SDSS</td>
<td>Sloan Digital Sky Survey</td>
</tr>
<tr>
<td>SMASS</td>
<td>Small Main-belt Asteroid Spectroscopic Survey</td>
</tr>
<tr>
<td>$V$</td>
<td>visible-band (~600 nm) astronomical magnitude</td>
</tr>
<tr>
<td>$V_{\text{MPC}}$</td>
<td>predicted $V$ magnitude based on MPC-hosted observational data and the $G$ phase-function model</td>
</tr>
<tr>
<td>$V_{\text{PTF}}$</td>
<td>predicted $V$ magnitude based on color-class-averaged albedos and phase-functions data derived from PTF data and $G_{12}$ phase-function model</td>
</tr>
<tr>
<td>$W$</td>
<td>visible-band (~600 nm) astronomical magnitude</td>
</tr>
<tr>
<td>WISE</td>
<td>Wide-field Infrared Explorer satellite</td>
</tr>
</tbody>
</table>

(Using two different models) corresponding to each GALEX detection in an effort to study the NUV − $V$ color. For both $V$ models, the derived NUV − $V$ color distribution is bimodal, with $S$ types having the redder color, just as they do within the visible band. The average C-type NUV − $V$ agrees with HST observations of the asteroids Lutetia and Ceres, both of which are members of the visible-color-defined C-type group. Slight differences in the measured NUV − $V$ among known taxonomic subgroups of the C types may indicate membership in either the M-type or G-type subgroups, though the 300–400 nm region ($u$-band) is more diagnostic of this division.

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Glossary of Acronyms and Symbols

For the reader’s convenience, Table 4 summarizes the various acronyms and mathematical symbols used in this work.
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