Investigating the peculiar emission from the new VHE gamma-ray source H1722+119


(Affiliations can be found after the references)
ABSTRACT

The MAGIC (Major Atmospheric Gamma-ray Imaging Cherenkov) telescopes observed the BL Lac object H1722+119 (redshift unknown) for six consecutive nights between 2013 May 17 and 22, for a total of 12.5h. The observations were triggered by high activity in the optical band measured by the KVA (Kungliga Vetenskapsakademien) telescope. The source was for the first time detected in the very high energy (VHE, $E > 100\text{ GeV}$) $\gamma$-ray band with a statistical significance of $5.9\sigma$. The integral flux above $150\text{ GeV}$ is estimated to be $(2.0 \pm 0.5)\%$ per cent of the Crab Nebula flux. We used contemporaneous high energy (HE, $100\text{ MeV} < E < 100\text{ GeV}$) $\gamma$-ray observations from Fermi-LAT (Large Area Telescope) to estimate the redshift of the source. Within the framework of the current extragalactic background light models, we estimate the redshift to be $z = 0.34 \pm 0.15$. Additionally, we used contemporaneous X-ray to radio data collected by the instruments on board the Swift satellite, the KVA, and the OVRO (Owens Valley Radio Observatory) telescope to study multifrequency characteristics of the source. We found no significant temporal variability of the flux in the HE and VHE bands. The flux in the optical and radio wavebands, on the other hand, did vary with different patterns. The spectral energy distribution (SED) of H1722+119 shows surprising behaviour in the $\sim 3 \times 10^{14} - 10^{18}\text{ Hz}$ frequency range. It can be modelled using an inhomogeneous helical jet synchrotron self-Compton model.

Key words: BL Lacertae objects: individual: H1722+119 – gamma-rays: galaxies – galaxies: active – galaxies: distances and redshifts

1 INTRODUCTION

Active galactic nuclei (AGN) are the most luminous persistent sources of electromagnetic radiation in the Universe. They are compact regions in the centres of galaxies formed around supermassive black holes (SMBHs) that actively accrete matter. Approximately 10 per cent of AGN eject matter through relativistic jets (Dunlop et al. 2003). AGN whose jets are oriented close to the line of sight to the Earth are referred to as blazars. Jets are sources of electromagnetic radiation of all wavelengths from radio to $\gamma$-rays. They can extend to Mpc distances from the nucleus and can be brighter than the rest of the galaxy. The spectral energy distribution (SED) of blazars is characterised by two broad peaks: a “low-energy” peak in the optical to X-ray band and a “high-energy” peak in the X-ray to $\gamma$-ray band. Their characteristics of their emission: flat spectrum radio quasars (FSRQs) and BL Lacertae objects (BL Lacs). FSRQs are known to have prominent broad and narrow optical emission lines, in addition to strong optical and X-ray continuum emission. A quite common feature is the so-called blue bump in the optical-UV band associated with the emission from the accretion disc. BL Lacs, on the other hand exhibit very weak optical emission lines if any at all. The low-energy peak in SED of FSRQs is believed to be a combination of synchrotron radiation of electrons in the jet and thermal emission from the broad line region (BLR), dust torus and accretion disc, while usually in the case of BL Lacs, all the emission is attributed to synchrotron emission. According to leptonic scenarios, the second peak is a result of inverse Compton (IC) scattering of lower-energy photons (so-called seed photons) on relativistic electrons within the jet, while hadronic scenarios assume protons in the jet are accelerated even to energies $\gtrsim 10^{19}\text{ eV}$ and significantly contribute to the emission either through proton-synchrotron emission in relatively strong magnetic fields, or photo-pion production (Böttcher et al. 2013). Ghisellini et al. (2010) showed that the broad-band SEDs of $\gamma$-ray blazars can be, on average, well described by a simple one-zone leptonic model including synchrotron and IC emission components, with the addition of possible external contributions from e.g. the accretion disc or host galaxy emission.

H1722+119 was first observed in the 1970s as a part of the Uhuru X-ray sky survey and the HEAO 1 Large Area Sky Survey (LASS). The resulting X-ray source catalogues (Forman et al. 1978; source name: 4U1722+11l and Wood et al. (1984); source name: 1H1720+117) only identify it as an X-ray source. Almost twenty years later, Griffiths et al. (1989) and Brissenden et al. (1990) independently, classified it as a BL Lac object. Furthermore, H1722+119 was classified as an intermediate-energy-peaked BL Lac with the low-energy peak at $\nu_p = 6.3 \times 10^{15}\text{ Hz}$ (Nieppola et al. 2006). Brissenden et al. (1990) measured a very high level of linear polarization in the optical band, reaching $17.6 \pm 1.0$ per cent in the B band. H1722+119 has been observed in the radio band by the OVRO (Owens Valley Radio Observatory) since 2007 (Richards et al. 2011)\footnote{www.astro.caltech.edu/ovroblazars}. H1722+119 is included in the third Fermi-LAT (Large Area Telescope) catalogue (3FGL; Acero et al. 2015) as a counterpart of the $\gamma$-ray source 3FGL J1725.0+1152, with photon index $\Gamma = 1.89 \pm 0.05$ and $0.1–100\text{ GeV}$ flux of...
The MAGIC telescopes are located at the Observatorio del Roque de los Muchachos in the Canary Island of La Palma, Spain (28°45′ north, 18°54′ west), at 2200 m above sea level. Two 17 m diameter Imaging Atmospheric Cherenkov Telescopes are optimised for observations of γ-rays of energies above 50 GeV. A detailed overview of the MAGIC experiment and the telescope performance is given in Aleksić et al. (2016a,b).

The results reported here are based on the observations performed during six nights between 2013 May 17 and 22, triggered by the high optical state detected by the KVA (Kungliga VetenskapsAkademietelescope) (see Section 2.4). Between the first (2005–2009) and the second (2013) observation campaign the MAGIC system underwent a series of major upgrades (see Cortina et al. 2005; Albert et al. 2008a; Goebel et al. 2008; Aleksić et al. 2012a; Sitarek et al. 2013; Aleksić et al. 2016a,b). The current instrument is roughly twice as sensitive in the 100–200 GeV energy range compared to the single MAGIC I telescope in operation in 2009 (Aleksić et al. 2016b).

MAGIC usually observes sources in the so-called wobble mode (Fomin et al. 1994; Aleksić et al. 2016b). H1722+119 was observed in four false-source positions for a total of 12.5 h during 2013. After quality selection based on the stability of the event rates, the effective time amounted to 12.0 h. Observations were performed at zenith angles between 16° and 37°.

The data were analysed within the MARS (MAGIC Analysis and Reconstruction Software) analysis framework (Lombardi et al. 2011; Zanin et al. 2013). The VHE γ-ray signal is estimated using the distribution of the squared angular distance (θ2) between the reconstructed and nominal (ON-source) source positions in the camera coordinates for each event. Background is estimated in the same manner with respect to the OFF-source position. Usually three OFF-source positions are chosen at the same distance from the camera centre as the ON-source position and rotated by 90° each. Noff is the number of events originating within the source region (θ2 < 0.0125 deg2), and Noffs the normalised number of all events from the same region around OFF-source positions. An excess of 337.5 events above 60 GeV with respect to the background was detected, yielding a signal significance of 5.9σ using Eq. 17 of Li & Ma (1983). The θ2 distribution is shown in Fig. 1. The light curve of the integral VHE γ-ray flux above 150 GeV is shown in Fig. 2. There is no evidence of flux variability on a night-by-night basis. A fit with a constant flux of (6.3±1.6)×10−12 ph cm−2 s−1 resulted in χ2/NDf = 3.5/5, where NDf stands for number of degrees of freedom. This is equivalent to (2.0 ± 0.5) per cent of the Crab Nebula VHE γ-ray flux. The measured flux is consistent with the UL set by the MAGIC observations of H1722+119 from the previous campaign (Aleksić et al. 2011). The differential energy spectrum was reconstructed using the Forward Unfolding algorithm presented in Albert et al. (2007a). It can be described by a simple power law function dN/dE = f0(E/E0)−Γ with normalisation f0 = (4.3 ± 0.9stat ± 0.9syst)×10−11 ph cm−2 s−1 TeV−1 and a photon index Γ = 3.3±0.3stat ± 0.2syst, at normalisation energy E0 = 200 GeV.
The fit resulted in \( \chi^2 / \text{NDF} = 3.6/8 \). The systematic uncertainty of the photon index was estimated to be ±0.2, using Eq. 3 of Aleksić et al. (2016b). We estimate the error on the \( f_0 \) to be 20 per cent, which does not include the energy scale uncertainty estimated to be 17 per cent.

### 2.2 Fermi-LAT

The Fermi-LAT is a pair-conversion telescope operating from 20 MeV to > 300 GeV. Details about the Fermi-LAT are given in Atwood et al. (2009). The LAT data reported in this paper were collected from 2013 January 1 (MJD 56293) to December 31 (MJD 56657). During this time, the Fermi observatory operated almost entirely in survey mode. The analysis was performed with the ScienceTools software package version v9r32p5. Only Pass 7 reprocessed events belonging to the ‘Source’ class were used. The time intervals when the rocking angle of the LAT was greater than 52° were rejected. In addition, a cut on the zenith angle (< 100°) was applied to reduce contamination from the Earth limb \( \gamma \) rays, which are produced by cosmic rays interacting with the upper atmosphere. The spectral analysis was performed with the instrument response function P7REP_SOURCE_V15 using an unbinned maximum-likelihood method implemented in the tool gtlike in the 0.1–100 GeV energy range. Isotropic (iso_source_v05.txt) and Galactic diffuse emission (gll_iem_v05_revl.fit) components were used to model the background\(^2\). The normalizations of both components were allowed to vary freely during the spectral fitting.

We analysed a region of interest of 10° radius centred at the location of H1722+119. We evaluated the significance of the \( \gamma \)-ray signal from the source by means of the maximum-likelihood test statistic \( T S = 2 (\log L_1 - \log L_0) \), where \( L \) is the likelihood of the data given the model with \( (L_1) \) or without \( (L_0) \) a point source at the position of H1722+119 (e.g. Mattox et al. 1996). The source model used in gtlike includes all sources from the 3FGL catalogue that fall within 15° of the source. A first maximum-likelihood analysis was performed to remove from the model the sources having \( T S < 10 \) and/or the predicted number of counts based on the fitted model \( N_{\text{pred}} < 3 \). A second maximum-likelihood analysis was performed on the updated source model. In the fitting procedure, the normalization factors and the photon indices of the sources lying within 10° of H1722+119 were left as free parameters. For the sources located between 10° and 15° from our target, we kept the normalization and the photon index fixed to the values from the 3FGL catalogue.

All uncertainties in measured HE \( \gamma \)-ray flux reported throughout this paper are statistical only. The systematic uncertainty in the flux is dominated by the systematic uncertainty in the effective area (Ackermann et al. 2012), which amounts to 5–10 per cent in the 0.1–100 GeV energy range, and therefore smaller than the typical statistical uncertainties for this analysis\(^3\).

Integrating over the period 2013 January 1 – December 31 in the 0.1–100 GeV energy range, using a power law model as in the 3FGL catalogue, the fit yielded a TS = 335, with an average flux of \((4.1 \pm 0.7) \times 10^{-9} \) ph cm\(^{-2}\) s\(^{-1}\), and a photon index of \( \Gamma = 1.99 \pm 0.08 \) at the decorrelation energy \( E_0 = 3063 \) MeV. During the same period, the fit with a log-parabola model, \( dN/dE \propto (E/E_0)^{-\alpha} \log(E/E_0) \), in the 0.1–100 GeV energy range results in TS = 355, with a spectral slope \( \alpha = 1.99 \pm 0.08 \), and a curvature parameter \( \beta = 0.01 \pm 0.01 \), indicating no significant curvature of the \( \gamma \)-ray spectrum. The LAT light curve and the photon index evolution with 2-month time bins is shown in Fig. 3. The 2-month bin width is a trade-off between the significance of the source (TS > 25 in each time bin) and the shortest time-scale to be probed. For each time bin, the spectral parameters for H1722+119 and for all the sources within 10° from the target were left free to vary. No significant increase of the \( \gamma \)-ray activity was observed by LAT\(^6\).

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\(^2\) fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html

\(^3\) fermi.gsfc.nasa.gov/ssc/data/access/LAT\_sweats.html

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**Figure 1.** \( \theta^2 \) distribution of signal (black points) and background (grey area) events. \( N_{\text{on}} \) is the number of all events with \( \theta^2 < 0.0125 \text{deg}^2 \) (vertical dashed line) with respect to the source position in the camera, and \( N_{\text{off}} \) the normalised number of all events from the same region around OFF-source position.

**Figure 2.** MAGIC nightly light curve for energies above 150 GeV. Horizontal error bars represent the duration of each observation. Vertical arrows represent ULS for points whose relative error of the excess is > 0.5. The horizontal dashed line is a constant flux fit with parameters stated in the text.
in 2013 May and June – the time bin contemporaneous to the MAGIC observations. Both the flux and the photon index are well fitted by a constant with the following parameters: flux: \( p_0 = (3.43 \pm 0.56) \times 10^{-8} \text{ph cm}^{-2} \text{s}^{-1} \) and \( \chi^2/\text{NDF} = 3.87/5 \), photon index: \( \Gamma = 1.96 \pm 0.07 \) and \( \chi^2/\text{NDF} = 1.68/5 \). Leaving the photon index free to vary during 2013 May, we obtained \( TS = 37 \), with an average flux of \( (8.5 \pm 3.3) \times 10^{-8} \text{ph cm}^{-2} \text{s}^{-1} \), and a photon index of \( \Gamma = 1.4 \pm 0.3 \). The hint of hardening of the LAT spectrum may be an indication of the shift of the IC peak to higher energies during the MAGIC detection. By considering only the period May 13–26, including the MAGIC observation period, the maximum-likelihood analysis results in a \( TS = 15 \), with a flux of \( (0.9 \pm 0.4) \times 10^{-8} \text{ph cm}^{-2} \text{s}^{-1} \) (assuming \( \Gamma = 1.4 \)). By means of the \texttt{gtscalprob} tool, we estimated that the highest energy photon emitted from H1722+119 (with probability > 90 per cent of being associated with the source) was observed on 2013 April 28 (MJD 54610), with an energy of 65.7 GeV. Another photon with an energy of 51.6 GeV was observed on 2013 December 31 (MJD 56657).

### 2.3 Swift

The \textit{Swift} satellite (Gehrels et al. 2004) performed three observations of H1722+119 on 2008 May 31, 2013 January 15 and May 20. The observations were carried out with all three on-board instruments: the X-ray Telescope (XRT; Burrows et al. 2005, 0.2–10.0 keV), the Ultraviolet/Optical Telescope (UVOT; Roming et al. 2005, 170–600 nm) and the Burst Alert Telescope (BAT; Barthelmy et al. 2005, 15–150 keV). The hard X-ray flux of this source is below the sensitivity of the BAT instrument for the short exposure of these observations (see Table 1), therefore the data from this instrument are not included in this work. The source is not present in the \textit{Swift}/BAT 70-month hard X-ray catalogue (Baumgartner et al. 2013).

The XRT data were processed with standard procedures (\texttt{xrtpipeline v0.12.8}), filtering, and screening criteria using the HEASoft package (v6.14). The data were collected in photon counting (PC) mode on 2008 May 31 and 2013 January 15, and windowed timing (WT) mode on 2013 May 20. The source count rate in PC mode was low (< 0.5 counts s\(^{-1}\)); thus pile-up correction was not required. Source events were extracted from a circular region with a radius of 20 pixels (1 pixel = 2.36 arcsec), while background events were extracted from a circular region with a radius of 50 pixels in PC mode and 20 pixels in WT mode away from the source region and from other bright sources. Ancillary response files were generated with \texttt{xrtmkarf}, and account for different extraction regions, vignetting and point-spread function corrections. We used the spectral redistribution matrices in the Calibration data base maintained by HEASARC. We fitted the spectrum with an absorbed power law using the photoelectric absorption model \texttt{tbabs} (Wilms et al. 2000), with a neutral hydrogen column density fixed to its Galactic value (\( N_{\text{H}} = 8.88 \times 10^{20} \text{cm}^{-2} \); Kalberla et al. 2005). We noted that the X-ray flux of the source observed in 2013, i.e. \( (1.2) \times 10^{-11} \text{erg cm}^{-2} \text{s}^{-1} \) (Fig. 3), is a factor between 3 and 5 higher than the flux in 2008 May 31, \( (3.6 \pm 0.7) \times 10^{-12} \text{erg cm}^{-2} \text{s}^{-1} \), with no significant spectral change (Table 1). For the observation in 2013 May 20, with the higher statistics, we fit the spectrum also with a log-parabola model, obtaining a spectral slope \( \alpha = 2.2 \pm 0.1 \), a curvature parameter of \( \beta = 0.54^{+0.31}_{-0.14} \), and a \( \chi^2_{\text{red}}/\text{NDF} = 1.124/63 \). The F-test shows an improvement of the fit with respect to the simple power law \( \chi^2_{\text{red}}/\text{NDF} = 1.245/64 \) with a probability of 99.9 per cent. UVOT data in the \textit{v}, \textit{b}, \textit{u}, \textit{w1}, \textit{m2}, and \textit{w2} filters were reduced with the HEASoft package v6.16 executing the aperture-photometry task \texttt{uvotsource}. We extracted the source counts from an aperture of 5\('\) centered on the source, and the background counts from a circle with 10 arcsec radius in a nearby source-free region. Magnitudes were converted into dereddened flux densities by adopting the extinction value E(B–V) = 0.1497 from Schlafly & Finkbeiner (2011), the mean Galactic extinction curve in Fitzpatrick (1999) and the magnitude-flux calibrations by Bessell & Brett (1988). The UVOT flux densities collected in 2013 are reported in Fig. 3. The flux densities observed in 2008 May 31 are a factor of 2 lower with respect to the 2013 observations.

Checking the source SED (Fig. 5), we noticed that a monotonic connection between optical-UV and X-ray spectrum was not possible, which motivated us to analyse possible sources of errors affecting the data. An aperture correction procedure was executed for the 2013 May 20 UVOT images in all filters, estimating from field stars the correction to magnitudes extracted with full-width at half maximum (FWHM) circular apertures to obtain the counts within the standard apertures (5\('\) radii). A first photometry of the object with apertures of 2.5\('\) (the FWHM for all filters) was executed. Then 6 to 8 non-saturated field stars were selected and photometry was performed for each of them using UVW1, UVW2 and standard apertures, and non-contaminated background annular regions of radii 26\('\) to 33\('\). The weighted mean among all magnitude differences obtained with the two different apertures was calculated and finally subtracted from the object magnitude at FWHM apertures. Results are however compatible with the standard procedure, with a low flux in the \textit{w1} filter. Another check was required because the source colour, \( b - v \sim 0.7 \), is out of the range to which the average count rate to flux ratios (CFR) estimated by Breeveld et al. (2011) are applicable. Therefore, we explored UVOT calibration issues following the procedure described in Raiteri et al. (2010). We fit the source spectrum with a power law and convolved it with the filter effective areas and appropriate physical quantities to derive source-dependent effective wavelengths (EW) and CFR. However, the results obtained are very similar to those given by Breeveld et al., the largest variations being an increase of the EW by 2 and 3 per cent in the \textit{w1} and \textit{w2} band, respectively, and of the CFR by 4 per cent in \textit{w1}. More significant differences were found when comparing the convolved Galactic extinction values with those obtained by simply evaluating them at the filter EW. We obtained a 7 per cent decrease of the extinction in the \textit{w2} and \textit{m2} bands, and an 8 per cent increase in the \textit{w1} band. The optical part of the SED remained essentially unchanged after the re-calibration procedure. The results after the re-calibration procedure are shown in Fig. 5.
2.4 KVA

The KVA telescope is located at the Observatorio del Roque de los Muchachos, La Palma (Canary Islands, Spain), and is operated by the Tuorla Observatory, Finland. The telescope consists of a 0.6-m f/15 Cassegrain devoted to polarimetry, and a 0.35-m f/11 Schmidt-Cassegrain auxiliary telescope for multicolour photometry. The telescope has been successfully operated remotely since the fall of 2003. The KVA is used for optical support observations for MAGIC by making R-band photometric observations, typically one measurement per night per source.

H1722+119 has been regularly monitored by the KVA since 2005. At the beginning of May 2013, after an extended optical high state, the source reached an R-band magnitude of 20.56, but we saw no regularity in this variation. The flux varied significantly, the ratio between the highest and lowest flux being 1.6, but we did not observe any correlation with the Fermi-LAT data or any other CDMS sources.

The data were reduced by the Tuorla Observatory Team as described in Nilsson et al. (2016, in prep.). The data were corrected for Galactic extinction using the total absorption A_V = 0.374 mag from Schlafly & Finkbeiner (2011).

3.3 OVRO

The 40-m radio telescope at the OVRO observes at the 15 GHz band. In late 2007, it started regular monitoring of a sample of blazars in order to support the goals of the Fermi-LAT telescope (Richards et al. 2011). This monitoring program includes about 1800 known or potential γ-ray-loud blazars, including all candidate γ-ray blazar survey (CGRaBS; Healey et al. 2008) sources above declination −20°. The sources in this program are observed in total intensity twice per week with a 4 mJy (minimum) and 3 per cent (typical) uncertainty on the flux density. Observations are performed with a dual-beam (each 2.5 arcmin full-width half-maximum) Dicke-switched system using cold sky in the off-source beam as the reference. Additionally, the source is switched between beams to reduce atmospheric variations. The absolute flux density scale is calibrated using observations of 3C 286, adopting the flux density (3.44 Jy) from Baars et al. (1977). This result is in about a 5 per cent absolute scale uncertainty, which is not reflected in the plotted errors in Fig. 3.

The OVRO nightly light curve for 2013 is shown in Fig. 3. Although there is some indication of a short time variability, we fitted the whole sample with a linear function (F [Jy] = p_0 + p_1 × (t - 56300)[MJD]) to point out the general trend of increasing flux. The fit parameters are p_0 = (5.2 ± 0.2) × 10^{−2} Jy, p_1 = (1.07 ± 0.10) × 10^{−4} Jy/day and \chi^2/NDF = 39.68/36. We also considered the possibility of a constant flux, but that assumption was discarded with \chi^2/NDF = 146.9/37 (p = 5 × 10^{−15}).

3 RESULTS

3.1 Redshift and the intrinsic VHE γ-ray spectrum

VHE γ-rays can be absorbed by the extragalactic background light (EBL), through photon-photon interactions, resulting in e^+ e^− pair-production. The flux attenuation is directly dependent on the redshift of the source and energy of γ-rays. A redshift-estimation method described in Prandini et al. (2010) uses this fact. It relies on the assumption that both HE and VHE γ-rays are created by the same physical processes and in the same region, and that the intrinsic spectrum in the VHE range cannot be harder than the spectrum in the HE range. The method uses only the spectral slope of the HE range data. The VHE γ-ray spectrum was de-absorbed using the EBL model from Franceschini et al. (2008). Because spectral points obtained by the MAGIC telescopes at energies below 100 GeV are usually affected by larger systematic uncertainties, they are not used for the fit of the VHE spectrum. The reconstructed redshift is obtained by applying a simple empirical formula to the value estimated from de-absorption. Applying this method to the MAGIC and contemporaneous (2013 May; MJD 56413 – 56442) Fermi-LAT data we obtained the reconstructed redshift of H1722+119 to be z = 0.34 ± 0.15_{stat} + 0.05_{marg}, the error marked as meth being a result of the method as described in Prandini et al. (2010). The UL on the redshift was set to 1.06 for the 95 per cent confidence level. Applying the log likelihood ratio test to set the UL, as described in e.g. Mazin & Goebel (2007), we obtained a value of 0.95. When de-absorbed for redshift values greater than 0.95, the spectrum shape becomes parabolic in a log(dN/dE) vs log E representation, with apparent minimum at \approx 200 GeV. Our reconstructed redshift value is in agreement with the latest Landoni et al. (2014) and Farina et al. (2013) results. We used our result combined with the LL from Farina et al. (z > 0.4) to de-absorb the VHE γ-ray flux. For z = 0.4, and using the EBL model from Franceschini et al. (2008), we found the parame-

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Table 1. Log and fitting results of Swift/XRT observations of H1722+119 using a power law model with N_H fixed to Galactic absorption.

<table>
<thead>
<tr>
<th>Observation Date (MJD)</th>
<th>Net Exposure Time in s</th>
<th>Photon index Γ</th>
<th>Flux 2-10 keV ×10^{-12} erg cm^{-2} s^{-1}</th>
<th>\chi^2_{red} (NDF)</th>
</tr>
</thead>
<tbody>
<tr>
<td>2008-05-31 (54617)</td>
<td>1733</td>
<td>2.13 ± 0.22</td>
<td>3.6 ± 0.7</td>
<td>1.257 (11)</td>
</tr>
<tr>
<td>2013-01-15 (56307)</td>
<td>812</td>
<td>2.28 ± 0.15</td>
<td>17.8 ± 2.3</td>
<td>0.9040 (17)</td>
</tr>
<tr>
<td>2013-05-20 (56432)</td>
<td>1983</td>
<td>2.31 ± 0.08</td>
<td>10.5 ± 0.7</td>
<td>1.245 (64)</td>
</tr>
</tbody>
</table>

4. http://users.utu.fi/kani/1m
5. The highest flux of 7.65 ±0.11 mJy was measured in 2014 June, and the lowest in 2008 April (2.18 ± 0.05 mJy).
Moll et al. The axis of Villata & Raiteri 3.1, and references therein; Villata & Ferrari Raiteri et al. 2015) are included to show the over-

2.3 Mignone et al. AO 0235+16 Rieger at 4 from the same night was used to obtain a spectral point

the source SED during the two different brightness states.

EBL model from Swift sic VHE part of the spectrum. We also show the

The SED for 2013 May is shown in Fig. 3.3 SED modelling

tical spectral points from

GHz. The

b

curved, with a peak in the

Release (. dashed line in the bottom panel of Fig. 3), while the HE γ-ray flux was consistent with being constant, and the optical flux varied with no apparent regularity. Therefore, we cannot claim any connection between emissions in different energy bands.

3.2 Multifrequency light curve

As already mentioned in Section 2.4, MAGIC observations were triggered by an extended optical high state, which was the historical R-band maximum at that time. The VHE γ-ray flux (Fig. 2) is compatible with a constant flux and the previously established UL based on combined data taken over several years. None the less, we cannot reach a firm conclusion on whether the source was flaring in the VHE γ-rays during MAGIC observations or not. Observations performed during six consecutive nights were not sufficient to study a longer term variability. However, we compared the HE γ-ray light curve for the entire year to the optical light curve over period 2013 March 22 to October 05, and the radio light curve over period 2013 January 21 to October 05. The Fermi-LAT data were divided into 2-month time bins with the photon index left free to vary (Fig. 3), while each point in the KVA and OVRO light curves represents a single measurement. Comparison of fluxes for the entire period of collected data revealed that the OVRO data show an obvious trend of increasing flux on a time-scale of one year (dashed line in the bottom panel of Fig. 3), while the HE γ-ray flux was consistent with being constant, and the optical flux varied with no apparent regularity. Therefore, we cannot claim any connection between emissions in different energy bands.

3.3 SED modelling

The SED for 2013 May is shown in Fig. 5. Only data contemporaneous to MAGIC observations have been considered for modelling the SED for 2013. On 2013 May 22 (MJD 56434), one observation was performed by the OVRO at 15 GHz. The Swift data collected on 2013 May 20 (MJD 56432) were considered, and the R-band observation by the KVA from the same night was used to obtain a spectral point at 4.56 × 10^{18} Hz. The Fermi-LAT spectrum was calculated in the period 2013 May 1 – June 30. The MAGIC spectral points were obtained with Schmelling’s method as described in Albert et al. (2007a). Based on the arguments discussed in Section 3.1, we adopted a value for redshift of 0.4, and EBL model from Franceschini et al. (2008) to get the intrinsic VHE part of the spectrum. We also show the Swift data collected on 2008 May 31 (MJD 54617), in order to compare the source SED during the two different brightness states. Archival data from the 2MASS (Two Micron All Sky Survey) All-Sky Catalog of Point Sources (Cutri et al. 2003) and ALLWISE (Wide-field Infrared Survey Explorer) Data Release (Cutri & et al. 2013) are included to show the overall behaviour in the IR-optical band, however these were not considered for modelling. We also show the near-IR and optical spectral points from Landoni et al. (2014), which were also not used for modelling.

The source SED shows surprising behaviour in the frequency range ~ 3 × 10^{14} – 10^{18} Hz. Indeed, the de-absorbed (as described in Section 2.3) optical–UV spectrum appears curved, with a peak in the b band and a steep slope in the UV. H1722+119 is a bona fide BL Lac type of source. None of the observations used in this work, nor information found in the literature revealed anything about the nature of the black hole environment, nor the host galaxy. There is no evidence of emission at any frequency from the accretion disc, dust torus or BLR. In fact, Landoni et al. (2014) detected no intrinsic or intervening spectral lines, and set the ratio of beamed to thermal emission to be ≥ 400.

Therefore we do not expect the presence of thermal emission from an accretion disc and the host galaxy that can justify the shape of the optical–UV spectrum. Moreover, this shape prevents a monotonic connection with the X-ray spectrum, as would be expected if both the optical–UV and X-ray emissions were produced by a synchrotron process in the same jet region. The connection between the UV and X-ray spectrum in both states requires an inflection point that is hard to reproduce with simple, one-zone models. This instead can easily be obtained in the framework of a curved jet. Indeed, curved/helical geomeas and often have been observed in blazar jets (see e.g. Villata & Raiteri 1999, and references therein). A helical-jet morphology can arise as the result of perturbations induced by orbital motion in a binary black hole system or precession of the black hole spin axis (Villata & Raiteri 1999, and references therein; Rieger 2004), 3-D magnetohydrodynamical (MHD) simulations by Nakamura et al. (2001); Moll et al. (2008); Migone et al. (2010) show how kink instabilities lead to a wiggled, in particular helical, jet structure. MHD equilibrium of helical jets was investigated by Villata & Ferrari (1995).

We adopted the helical jet model originally proposed by Villata & Raiteri (1999) to explain the observed SED variations of Mkn 501, and later applied also to other objects like S4 0954+65 (Raiteri et al. 1999), AO 0235+16 (Ostorero et al. 2004), BL Lacertae (Raiteri et al. 2009: 2010) and PG 1553+113 (Raiteri et al. 2015). The axis of the helical-shaped jet is assumed to lie along the z-axis of a 3-D reference frame. The pitch angle is ζ and ψ is the angle defined by the helix axis with the line of sight. The non-dimensional length of the helical path can be expressed in terms of the z coordinate along the helix axis:

$$l(z) = \frac{z}{\cos \zeta}, \quad 0 \leq z \leq 1,$$

which corresponds to an azimuthal angle φ(z) = az, where the angle a is a constant. The jet viewing angle varies along the helical path as

$$\cos \theta(z) = \cos \psi \cos \zeta + \sin \psi \sin \zeta \cos(\phi - az),$$

where φ is the azimuthal difference between the line of sight and the initial direction of the helical path. The jet is inhomogeneous: each slice of the jet can radiate, in the plasma rest reference frame, synchrotron photons from a minimum frequency ν′_{s,min} to a maximum one ν′_{s,max}, which follow the laws:

$$\nu_{s,i}(l) = \nu_{s,i}(0) \left( 1 + \frac{l}{l_i} \right)^{-c_i}, \quad c_i > 0,$$

where l_{i} are length scales, and i = min, max. The model takes into account IC scattering of the synchrotron photons by the same relativistic electrons emitting them (i.e. SSC). Consequently, each portion of the jet emitting synchrotron radiation between ν′_{s,min} and ν′_{s,max} will also produce IC radiation between ν′_{e,min} and ν′_{e,max}, with

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Figure 3. Multifrequency light curve during 2013. MAGIC nightly light curve is the same as in Fig. 2. ULs were omitted for clarity. The HE γ-ray light curve and the evolution of the spectral index, as measured by Fermi-LAT, are shown in 2-month time bins, while one night time-scale is used in light curves in other energy bands. Horizontal dashed lines in both panels represent the fit with a constant. The Swift/XRT filters are represented with the following markers: v – full circle, b – full square, u – upward triangle, w1 – downward triangle, m2 – empty circle, and w2 – empty square. The dashed line in the OVRO panel represents a linear function fit. All the fit parameters are given in the text. The data were collected (from top to bottom) by MAGIC (E > 150 GeV), Fermi-LAT (0.1 GeV < E < 100 GeV), Swift/XRT, Swift/UVOT, KVA (R-band) and OVRO (15 GHz).
\[ \nu'_s(l) = \frac{1}{\gamma} \nu(l)^3 \nu'_s(l). \]

The electron Lorentz factor ranges from \( \gamma_{\text{min}} = 1 \) to \( \gamma_{\text{max}}(l) \), which has a similar dependence as in Eq. 3, with power \( c \), and length scale \( l_s \). As photon energies increase, the classical Thomson scattering cross section gradually shifts into the extreme Klein-Nishina one, which makes Compton scattering less efficient. In our approximation the power law dependence of the observed flux density on the frequency and a cubic dependence on the Doppler beaming factor is averaged with \( K_{\text{KN}} \).

\[ F_\nu(\nu) \approx \frac{\nu^\alpha}{(1 + \frac{l}{l_s})^{-\alpha - 1}} \quad \text{with} \quad \alpha_0 = \frac{\gamma^2}{K_{\text{KN}}} \]

The helical jet model is intended to be a geometrical, dimensionless, model used to describe blazar variability and should not be interpreted as a physical model. However, absolute dimensions can be derived by comparison with VLBA images. In our model, the 15 GHz radio emission comes from the outer 20 per cent of the jet (see Fig. 4). This is assumed to be the jet region where the bulk of the observed 15 GHz radiation comes from. The results of the MOJAVE program\(^6\) show that the size of the emitting core is \( \lesssim 0.1 \text{ pc} \). Therefore the unit length in our model would correspond to a dimension of the order of one tenth of a parsec.

**Summary and Conclusions**

MAGIC detected VHE \( \gamma \)-radiation from BL Lac H1722+119 after observations were triggered by a high flux in R-band measured by the KVA. The MAGIC observations were performed during six consecutive nights and show no flux variability. No significant increase of the activity was observed at HE by Fermi-LAT in 2013 May neither on short timescales nor compared to the average 2013 flux. An indication of spectral hardening at HE in 2013 May might explain a VHE \( \gamma \)-ray flux high enough to be detected by the MAGIC telescopes. Changes in flux were significant in optical and radio data. However, on a time-scale of a year, the radio flux seems to increase, while there was no significant variability.

\[ \begin{array}{c|c}
\zeta & 30^\circ \\
\alpha & 40^\circ \\
\psi & 25^\circ \\
\phi & 25^\circ, 31^\circ \\
\alpha_0 & 0.5 \\
\Gamma_h' & 10 \\
\log \nu'_{\text{min}}(0) & 16.7 \\
\log \nu'_{\text{max}}(0) & 19.3 \\
\epsilon_{\text{min-max}} & 3.3 \\
\log l_{\text{min}} & -2.3 \\
\log l_{\text{max}} & -1.8 \\
\log \gamma_{\text{max}}(0) & 4.8 \\
c_1 & 1.5 \\
\log l_{\beta} & -1 \\
c_{1,1,1} & 3 \\
c_{2,2,2} & 2.3 \\
\log K_{\beta} & -19.25 \\
\log K_{c} & -24.55
\end{array} \]

\( \zeta = \text{observer inclination angle} \)

\( \alpha = \gamma_{\text{min}} \)

\( \psi = \text{v} \)

\( \phi = \gamma_{\text{max}}(l) \)

\( \alpha_0 = \text{power law index of the bulk Lorentz factor} \)

\( \Gamma_h = \text{beaming factor} \)

\( \nu'_{\text{min}}(0) = \text{minimum frequency} \)

\( \nu'_{\text{max}}(0) = \text{maximum frequency} \)

\( \epsilon_{\text{min-max}} = \text{energy index range} \)

\( \log l_{\text{min}} = \text{minimum log of the distance} \)

\( \log l_{\text{max}} = \text{maximum log of the distance} \)

\( \log \gamma_{\text{max}}(0) = \text{maximum log of the Lorentz factor} \)

\( c_1 = \text{constant} \)

\( \log l_{\beta} = \text{log of the distance} \)

\( c_{1,1,1} = \text{constant} \)

\( c_{2,2,2} = \text{constant} \)

\( \log K_{\beta} = \text{log of the distance} \)

\( \log K_{c} = \text{log of the distance} \)

\( \text{Table 2. Main parameters of the helical model to fit the SEDs of H1722+119. The only difference between the high and low state is the angle } \phi, \text{ which changes from } 25^\circ \text{ to } 31^\circ. \)

\( \text{www.physics.purdue.edu/astro/MOJAVE/sourcepages/1722+119.shtml} \)
of the HE flux. Therefore, we cannot claim that the two components originate from the same physical region. While the radio flux was highest towards the end of the year, the highest optical flux was observed in April, followed by a quite sudden drop, which did not occur in radio data.

The SED shows an interesting feature in the optical–UV band. A break between the optical–UV parts of the spectrum is clearly visible. It also prevents a smooth connection between optical, UV and X-ray bands. We proposed the helical jet model to the 2013 data, while the fit to the 2008 data is indicated by the red dash-dot line. Both models represent the intrinsic VHE emission. Only blue and red points were considered for modelling.

Further multifrequency observations will be crucial to investigate in detail this new VHE emitting blazar and its emission mechanisms.

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