

SOFIA/FORCAST Galactic Center Legacy Survey: Overview

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ABSTRACT

The Galactic Center contains some of the most extreme conditions for star formation in our Galaxy as well as many other phenomena that are unique to this region. Given our relative proximity to the Galactic Center, we are able to study details of physical processes to a level that is simply not yet possible for more distant galaxies, yielding an otherwise inaccessible view of the nuclear region of a galaxy. We recently carried out a targeted imaging survey of mid-infrared bright portions of the Galactic Center at 25 and 37 μm using the FORCAST instrument on *SOFIA*. This survey was one of the inaugural Legacy Programs from *SOFIA* cycle 7, observing a total area of 403 arcmin² (2180 pc²), including the Sgr A, B, and C complexes. Here we present an overview of the survey strategy, observations, and data reduction as an accompaniment to the initial public release of the survey data. We discuss interesting regions and features within the data including extended features near the circumnuclear disk, structures in the Arched Filaments and Sickle H II regions, and signs of embedded star formation in Sgr B2 and Sgr C. We also feature a handful of less well studied mid-infrared sources located between Sgr A and Sgr C that could be sites of relatively isolated star formation activity. Last, we discuss plans for subsequent publications and future data releases from the survey.

Keywords: Galaxy: center, ISM: H II regions

1. INTRODUCTION

The environment in the Galactic Center (GC) is unlike any other part of our Galaxy. The region contains high molecular gas densities (Guesten & Henkel 1983), high temperatures (Morris et al. 1983; Guesten et al. 1985), and large turbulent motions (Bally et al. 1987), all inside a deep gravitational potential well (Morris & Serabyn 1996). The conditions in the GC mirror those found in the nuclei of luminous infrared galaxies and high-redshift systems near the peak of cosmic star formation history (Kruijssen & Longmore 2013), but its relative proximity ($d = 8.0 \pm 0.5$ kpc; Reid 1993) enables us to study the physical processes there at a level of detail that is simply not possible in more distant systems (see the recent review by Mills 2017).

Observations of the GC often challenge theoretical models of star formation, which often break down in this complex region. For example, while the GC comprises less than 0.01% of the total volume of the Galactic disk, its star formation rate (SFR) ($\sim 0.1 M_{\odot} \text{ yr}^{-1}$; Immer et al. 2012; Barnes et al. 2017) is a considerable fraction of the total SFR of the Galaxy ($\sim 1.2 M_{\odot} \text{ yr}^{-1}$; Lee et al. 2012). However, the global GC SFR is more than an order of magnitude smaller than one would expect based on scaling relations with its dense molecular gas content (Lada et al. 2012; Longmore et al. 2013). The inefficiency of the GC in converting dense gas to stars presents a significant quandary for the region with much broader implications. For example, SFR measurements are universally used as a fundamental diagnostic tool for understanding the underlying physics in galaxies and for understanding galaxy evolution.

Even with the aforementioned star formation deficiency, the stellar inventory of the GC is relatively rich with numerous types of massive evolved stars, such as Luminous Blue Variables and Wolf-Rayet stars (e.g., Figer 2009). The GC is home to three known massive stellar clusters (Lu 2018) - the Arches cluster (Cotera et al. 1996), the Quintuplet cluster (Okuda et al. 1990; Nagata et al. 1990), and the central cluster (Krabbe et al. 1991, 1995) in addition to a large number of massive field stars which are spread throughout the region (e.g., Munro et al. 2006; Mauerhan et al. 2010; Dong et al. 2012).

The origin of these massive field stars is somewhat of a mystery. While several sources may be former cluster members which have been dynamically ejected or removed due to tidal evaporation, the entire population of these sources cannot be accounted for with these mechanisms (Habibi et al. 2014). Instead, some fraction of the field stars likely originates from a more isolated mode of star formation, and there is evidence of ongoing massive star and stellar cluster formation within parts of the GC (e.g., Barnes et al. 2019, and references therein). There have been efforts to identifying possible clusters associated with GC field stars, but these have proven challenging to directly observe (e.g., Steinke et al. 2016; Dong et al. 2017).

Large extinction toward the GC ($A_V \sim 30$; Fritz et al. 2011) makes it impossible to observe massive stars and protostars at optical and ultraviolet wavelengths. Instead, studies of the region have relied on observations at other wavelengths to examine the distribution and birth environment of stars. Numerous studies of star formation in the GC have been conducted with mid-infrared observations between $\sim 3.6 - 24 \mu\text{m}$ (e.g., Ramírez et al. 2008; Yusef-Zadeh et al. 2010; An et al. 2011; Immer et al. 2012). In particular, warm dust emission at $\sim 24 \mu\text{m}$ is a valuable probe for identifying young stellar objects (YSOs) and estimating star formation rates (e.g., Calzetti et al. 2007). However, the most active regions within the inner ~ 200 pc of our Galaxy are strongly saturated in the *Spitzer*/MIPS $24 \mu\text{m}$ data (Yusef-Zadeh et al. 2009). Earlier observations with the *Midcourse Space Experiment* (MSX; Egan et al. 2003) at $21.3 \mu\text{m}$ (Band E) provide unsaturated images, but are relatively low spatial resolution ($\sim 20''$ or ~ 0.8 pc) when compared with *Spitzer*/MIPS at $24 \mu\text{m}$ ($\sim 6''$ or ~ 0.2 pc). The MSX data suffers from significant confusion in these complex regions, which presents a considerable hurdle in our understanding of these very active portions within the GC. Furthermore, the lack of high-quality mid-infrared data in these regions represents an essential, missing piece of the rich multi-wavelength picture of the GC that has emerged over the last decade.

In order to create improved mid-infrared maps of the brightest portions of the inner ~ 200 pc of our Galaxy, we set out to con-

duct a targeted survey of these regions within the GC using *SOFIA*/FORCAST. This survey was selected as one of the inaugural Legacy Programs in *SOFIA* cycle 7 (Program ID: 07-0189; PI: Hankins). Observations were obtained at 25 and 37 μm using the Faint Object infraRed CAmera for the SOFIA Telescope (FORCAST), enabling us to create a high-resolution (FWHM \sim 2.3" or \sim 0.07 pc at 25 μm and FWHM \sim 3.4" or \sim 0.1 pc at 37 μm) mosaic of portions within the GC including the Sgr A complex and other prominent star forming regions such as Sgr B and Sgr C.

In this paper we present a description of the observations along with the survey strategy and an initial look at the data set with various regions of interest highlighted. In section 2, we provide details on the observations and information regarding the creation of the mosaics presented in this work. In section 3, we highlight and discuss regions and sources that stand out as particularly interesting. These overviews are not intended to be a full and complete analysis of the objects discussed but rather to emphasize areas where this data set is providing an enhanced view of regions within the GC. We are planning follow-up papers for several of these features, as discussed in the text below. Finally, we provide a summary and future outlook in section 4.

2. OBSERVATIONS AND DATA REDUCTION

Observations of our mid-infrared GC survey were carried out with the 2.5 m telescope aboard the Stratospheric Observatory for Infrared Astronomy (*SOFIA*; Young et al. 2012) using the FORCAST instrument (Herter et al. 2012). FORCAST is a 256×256 pixel dual-channel, wide-field mid-infrared camera with a field of view (FOV) of $3.4' \times 3.2'$ and a plate scale of 0.768" per pixel. The two channels consist of a short-wavelength camera (SWC) operating at 5 – 25 μm and a long-wavelength camera (LWC) operating at 28 – 40 μm , and the instrument is capable of observing with both cameras simultaneously using a dichroic beam splitter.

2.1. *SOFIA*/FORCAST Survey Plan

The *SOFIA*/FORCAST Galactic Center survey was designed to map out infrared bright regions within the inner \sim 200 pc of the Galaxy.

We chose to conduct the survey using the FORCAST 25 μm filter because it is close in wavelength to the MIPS 24 μm filter and only a slight color-correction is needed to compare photometry between the two data sets where possible. The FORCAST and MIPS data sets are complementary since the higher bright source limit of FORCAST allows for observations of objects and regions where the MIPS data are saturated, while the higher sensitivity of the MIPS probes fainter sources than FORCAST can detect in reasonable integration times. Our nominal sensitivity per field was selected to achieve a 5- σ point source depth of 250 mJy at 25 μm , able to detect class I YSOs (age < 1 Myr) down to a mass of $\sim 6 M_{\odot}$ based on the mean 24 μm flux of YSO models from Robitaille et al. (2006). This depth is also comfortably below the MIPS hard saturation limit for the existing 24 μm GC map (~ 400 mJy) allowing for comparison of common sources between the MIPS and FORCAST surveys.

In addition to the 25 μm filter, we used FORCAST's internal dichroic beam-splitter to enable simultaneous observation with the long wavelength camera which results in only modest loss in the short wavelength channel.¹ To give the greatest wavelength coverage, we selected the longest wavelength filter available on FORCAST (37.1 μm) to be paired with the 25 μm filter observations. The nominal spatial resolution of the 25 and 37 μm filters in this observing mode are 2.3" and 3.4", respectively.²

Field pointings for the survey were planned using the *MSX* band E data (Figure 1). Additionally, we compared the survey footprint with the aforementioned 24 μm *Spitzer*/MIPS data in order to optimize coverage of the hard saturated sources. In total our survey plan covered more than 99% of the hard saturated area in the innermost 200 pc of the Galaxy (Figure 1).

2.2. *SOFIA*/FORCAST Cycle 7 Observations

A total of 34 fields were observed during *SOFIA* cycle 7 as part the FORCAST Galactic Center Legacy Program (Program ID: 07-0189;

¹ Throughput of the dichroic for the SWC from 11–25 μm is 85%, while the throughput of the LWC from 25–40 μm is only 40%

² <https://www.sofia.usra.edu/science/proposing-and-observing/observers-handbook-cycle-7/5-forcast>

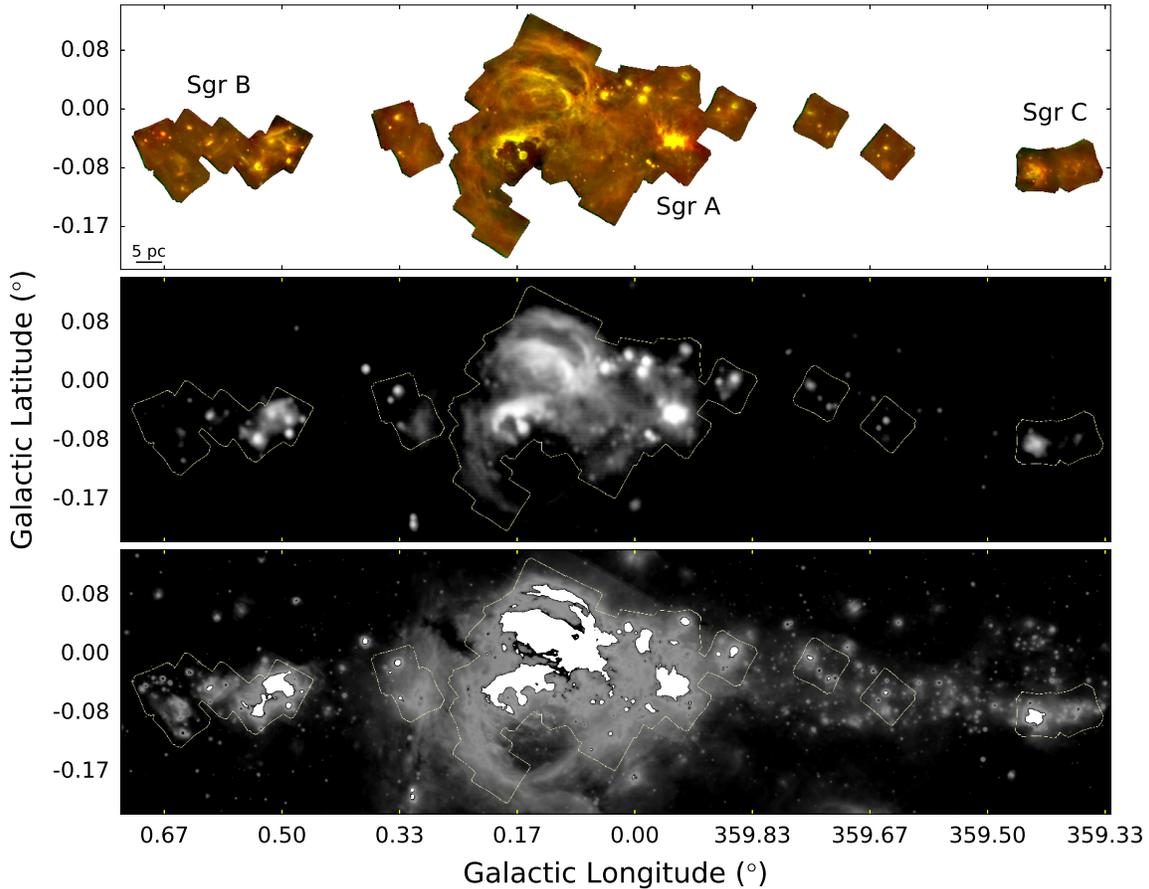


Figure 1. *Upper:* The *SOFIA*/FORCAST GC survey mosaic created using the 25.2 (green) & 37.1 (red) maps. Well known regions including Sgr A, B, and C are labeled. *Middle:* The *MSX* band E data that were used for planning the FORCAST observations. The stretch of this figure shows the approximate depth of the FORCAST observations for comparison. *Bottom:* The 24 μm *Spitzer*/MIPS mosaic of the GC. Hard saturated regions within the data are shown in white. The stretch of this figure demonstrates the relative depth of the MIPS map and the numerous unsaturated sources within the region that have useful data.

PI: Hankins). Here, we provide a brief description of the observations, with further details provided in Table 1. All 34 fields were observed during the annual *SOFIA* southern hemisphere deployment to Christchurch, NZ. Observations for the program were spread over 8 total flights which occurred between 1 July 2019 and 11 July 2019.

Chopping and nodding was used to remove the sky and telescope thermal backgrounds for all observations. This technique requires off fields which are devoid of emission to properly subtract the background emission. The GC is a particularly complex emitting environment, and there were instances where we did not have perfectly blank sky to nod and chop

onto within telescope limits. In these cases, we selected the ‘best’ available off field which minimized the number of sources (most frequently an individual source within the expected detection limits). We employed the C2NC2 observing mode to allow use of the full FORCAST FOV, even though this mode comes with considerable overheads due to the necessary three off-source positions. Dithering was used to remove bad pixels and mitigate response variations. We used different dither patterns for several of the fields in order to minimize observing inefficiencies related to the C2NC2 observing mode, which does not adversely impact the data quality between fields.

On average we achieved 462s of integration for each field in the 25 μm filter and an average of 423s for the 37 μm filter. A listing of integration times for each field individually can be found in Table 1. The estimated total photometric errors for the 25 μm and 37 μm data are $\sim 10\%$. Two fields (7 & 11) have integration times which are substantially below the average, and were cut short due to scheduling issues and/or adjustments that were needed in flight. However, the depth of these images is still suitable for some science goals (corresponding to a 250 mJy point source detection of $\sim 3.7\text{-}\sigma$ rather than $5\text{-}\sigma$). In addition to these low-SNR fields, there were three additional fields that were planned as part of the survey but not observed in cycle 7 due to time and observability constraints.

During the flight series there was an issue with the secondary chopper mirror on the telescope which caused sources to be elongated in the direction of the chop throw. Consequently, in Table 1 we have included information regarding the extent to which each field was affected by the secondary mirror issue. Fields where the elongation is known to be an issue are labeled ‘affected.’ Fields that may also have been impacted, but only slightly are labeled ‘nominal’, while those taken after the issue was corrected are labeled ‘post-fix.’ In the affected data, PSF ellipticities range between $\sim 0.1\text{-}0.4$; however, several of the fields contain few, and in some cases no suitable PSF reference stars which makes measuring the data quality effects challenging. Using a deconvolution method with a suitable PSF model should be able to correct the PSF variation to an extent, although this is a complex process given level of PSF variation in the ‘affected’ fields. We are planning to produce a set of corrected maps with a more uniform PSF to be made public with a future data release.

2.3. *SOFIA/FORCAST Observations from Prior Cycles*

In addition to the observations carried out in *SOFIA* cycle 7, suitable data from earlier cycles were considered for inclusion in our GC survey. Several guaranteed time observations (GTO) with FORCAST were focused on targets within the GC, including the Circumnuclear Ring (Lau et al. 2013), the Sickle H II region (Lau et al. 2014a; Hankins et al. 2016),

and the Sgr A East H II regions (Lau et al. 2014b).

For the survey mosaic presented in this work, we only used observations taken in *SOFIA* cycle 3 or later which includes the Arched Filaments H II region (Hankins et al. 2017) and the H H II regions (Hankins et al. 2019). Observation details for these data can be found in the above referenced works and a description of their incorporation in our survey mosaic can be found in the following subsection.

2.4. *Data Processing*

Observations were processed using the pipeline steps described in Herter et al. (2013). The Level 3 processed data products were downloaded from the *SOFIA* Data Cycle System.³ Images from each individual pointing were mosaicked using the *SOFIA* Data Pipeline software REDUX (Clarke et al. 2015) in order to construct the preliminary FORCAST Level 4 imaging mosaics that are presented in this work.

Creating the mosaic for this data set resulted in several challenges which were addressed as follows. First, observations using FORCAST on bright sources cause a negative signal offset throughout the detector that lowers the background flux levels. To correct for this, background levels were subtracted on a field by field basis to bring the overall background values to approximately zero. Next, as mentioned previously, chopping and nodding requires off fields which are free of detectable mid-IR point sources or extended emission in order to properly subtract sky and telescope backgrounds. In cases where off fields contained sources, this results in negative point sources or regions. To correct for these negative artifacts, we modeled point sources using a 2D Gaussian and if they achieved sufficient SNR ($\gtrsim 5$) the source model was subtracted from the data, leaving behind the fitted background at the location of the artifact. A summary of the locations of removed sources can be found in Table 2.

Next, the issue that required the most effort to correct involved astrometry. Telescope pointing with FORCAST is only accurate to within a few pixels (\sim arc-seconds), therefore astrometry was absolutely calibrated using the

³ <https://dcs.arc.nasa.gov>

Table 1. *SOFIA*/FORCAST Observation Details

Field ID	Field Center (l,b)	Sky Angle (°)*	25 μ m t _{int} (s)	37 μ m t _{int} (s)	Chop Angle (°) [†]	Comment [‡]	AOR ID
1	(359.376, -0.080)	148.3	402	369	88	Affected	07_0189_1
2	(359.429, -0.087)	130.1	385	353	52	Affected	07_0189_2
3	(359.641, -0.062)	83.5	404	315	60	Post-fix	07_0189_3
5	(359.737, -0.018)	99.7	461	360	308	Post-fix	07_0189_5
6	(0.310, -0.058)	153.1	458	415	110	Post-fix	07_0189_6
7	(0.337, -0.024)	138.9	217	192	230	Post-fix	07_0189_7
8	(0.584, -0.053)	174.3	433	404	200	Affected	07_0189_8
9	(0.631, -0.041)	178.2	614	626	220	Nominal	07_0189_9
10	(0.641, -0.092)	165.1	528	480	77	Affected	07_0189_10
11	(0.674, -0.051)	146.2	210	176	130	Affected	07_0189_11
12	(0.534, -0.072)	163.3	414	387	80	Affected	07_0189_35
13	(0.496, -0.051)	184.7	522	532	60	Nominal	07_0189_36
A	(359.867, -0.007)	100.0	478	400	240	Affected	07_0189_12
B	(359.942, 0.027)	130.2	486	454	270	Affected	07_0189_13
C	(359.931, -0.019)	85.6	440	368	265	Affected	07_0189_14
D	(359.934, -0.067)	88.3	458	415	5	Affected	07_0189_15
E	(359.970, -0.019)	143.6	393	319	245	Affected	07_0189_16
F	(359.975, -0.064)	165.5	472	433	352	Affected	07_0189_17
G	(0.015, -0.022)	161.0	492	399	250	Affected	07_0189_18
H	(0.014, -0.078)	188.5	527	590	28	Affected	07_0189_19
I	(0.040, -0.124)	183.4	505	458	45	Affected	07_0189_20
K	(0.051, -0.008)	172.3	505	397	260	Affected	07_0189_22
L	(0.056, -0.053)	183.4	524	451	98	Affected	07_0189_23
M	(0.065, -0.089)	164.8	493	447	118	Post-fix	07_0189_24
O	(0.101, -0.023)	181.0	472	451	78	Affected	07_0189_26
P	(0.102, -0.071)	136.4	525	511	95	Nominal	07_0189_27
Q	(0.145, 0.006)	188.0	518	500	190	Affected	07_0189_28
R	(0.200, 0.023)	174.9	427	387	225	Affected	07_0189_29
S	(0.182, -0.018)	176.6	469	430	180	Affected	07_0189_30
T	(0.223, -0.043)	159.2	492	479	205	Affected	07_0189_31
U	(0.223, -0.089)	148.4	468	437	82	Affected	07_0189_32
V	(0.216, -0.134)	175.4	505	458	75	Post-fix	07_0189_33
W	(0.190, -0.169)	180.9	482	500	80	Post-fix	07_0189_34
X	(0.142, -0.045)	181.6	484	451	60	Post-fix	07_0189_37
Y	(0.175, -0.067)	170.6	504	471	35	Post-fix	07_0189_38

*Sky angle refers to the orientation of the field measured in degrees east of north.

[†]Chop angle refers to the direction of the chop throw, which is also the direction of PSF elongation in the fields affected by the secondary mirror issue described in text.

[‡]Data quality comments refer to the chopper mirror issue described in Section 2.2. Data labeled ‘Affected’ have noticeable PSF elongation in the direction of the Chop Angle. Data labeled ‘Nominal’ or ‘Post-Fix’ are unaffected by the mirror issue.

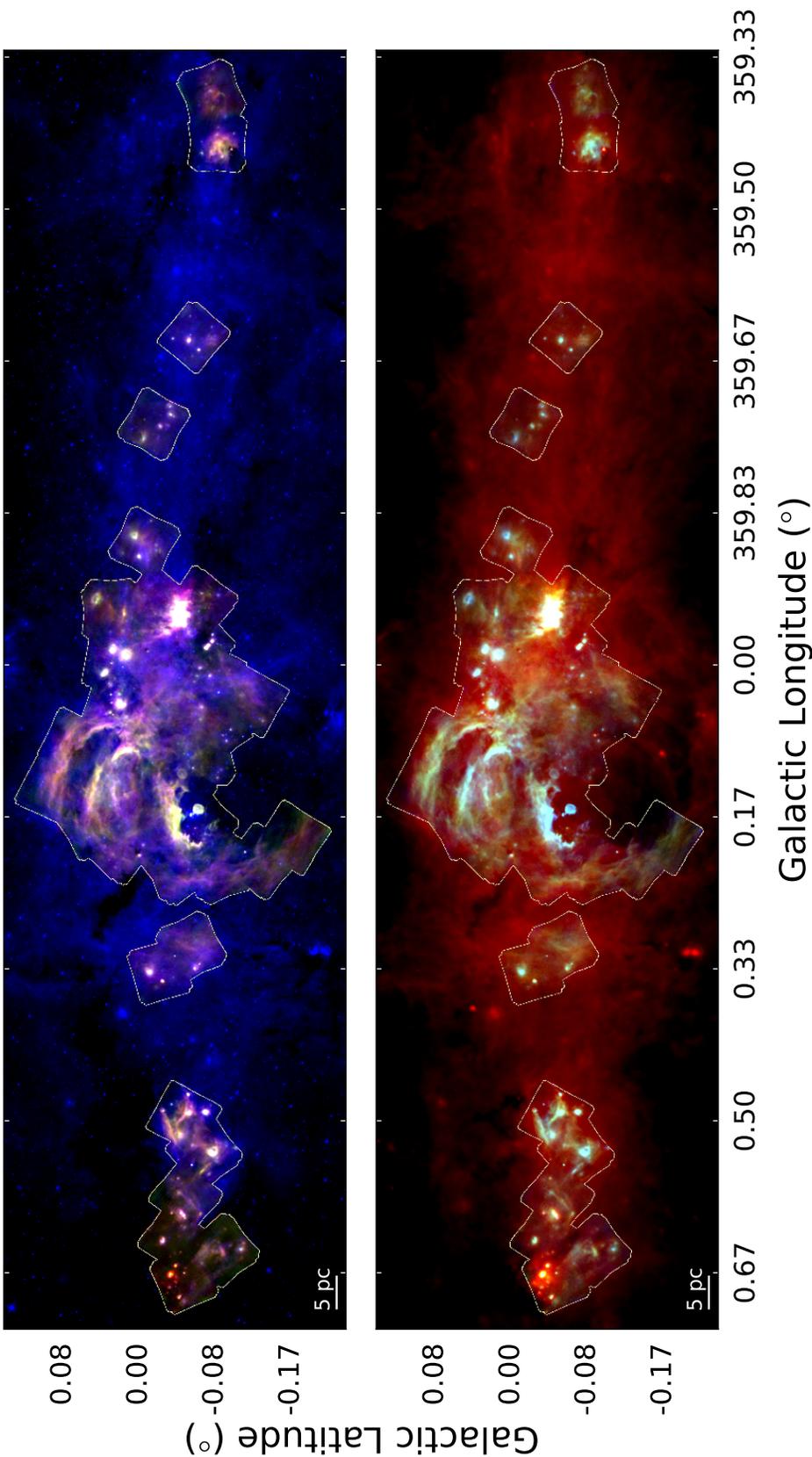


Figure 2. Comparison of the *SOFIA*/FORCAST mosaics with other infrared survey maps. *Top:* A false-color map of the survey region using *Spitzer*/IRAC 8 μm (blue), *SOFIA*/FORCAST 25 μm (green), and 37 μm (red). *Bottom:* A second false-color map of the survey region using *SOFIA*/FORCAST 25 μm (blue), 37 μm (green), and *Herschel*/PACS 70 μm (red). Both figures show an outline of the *SOFIA*/FORCAST survey footprint in white.

Table 2. Known Data Artifacts Removed in Post-Processing

Location (l,b)	Filter(s)	Type*	Comment [†]
(0.103, -0.086)	Both	point	
(0.098, -0.079)	Both	point	
(0.090, -0.070)	Both	point	
(359.933, 0.024)	25 μm	point	
(359.894, -0.008)	25 μm	point	dithered
(359.896, -0.012)	25 μm	point	dithered
(359.892, -0.015)	25 μm	point	dithered
(0.053, -0.086)	25 μm	point	
(0.098, 0.055)	25 μm	point	
(0.103, 0.074)	25 μm	point	
(0.126, 0.111)	Both	point	
(0.115, 0.081)	25 μm	point	
(0.482, -0.077)	Both	point	nearby source

*Type specifies if the source is point-like or extended.

[†]Comments on removed sources. The ‘dithered’ source is the same artifact appearing in multiple locations. The ‘nearby source’ label denotes a nearby positive source where the photometry may be impacted.

available *Spitzer* and *MSX* data by matching up the centroids of point sources in common between those maps and the *SOFIA* data. Precise alignment was difficult in the case of aligning to the lower resolution *MSX* 21 μm data or the saturated regions of the *Spitzer* MIPS 24 μm data. Although, a handful of bright 8 μm point sources were also detected at 25 μm and aided in the alignment of several of fields.

Slight changes in the focal plane distortion across the array and limited calibration data, also contributed to some elongation and smearing of sources on scales similar to the elongation caused by the secondary chopping mirror issue as discussed above. This distortion is variable across the array with the worst effects at the edges resulting in some misalignment of up to a few pixels (\sim arc-seconds). Based on these issues we estimate the astrometry of the final *SOFIA* mosaic is at worst 3 pixels or $\sim 2''$. However, comparing common sources between the FORCAST and *Spitzer* suggests the FORCAST mosaic astrometry is typically bet-

ter than 1 pixel.⁴ More calibration data is expected to be taken in the next *SOFIA* cycle, in which case the overall astrometric precision of the maps may be improved in a future data release.

Finally, changes to the data reduction pipeline between *SOFIA* cycle 7 and earlier cycles presented several issues for combining the various data sets into the mosaics presented in this work. These issues were largely overcome by reprocessing the earlier cycle 3 and 4 data using REDUX with modifications to improve backwards compatibility. While we present initial mosaics produced as part of this program, there is still work needed to optimally combine the older and newer data sets. For example, the new processing for the 37 μm observations has resulted in a 3-pixel wide gap in a small section between two fields in the Arched Filaments near (0.116, 0.073), and there is a similar 1-pixel wide gap in a section between another set of fields in the 25 μm map located near (0.148, 0.065). Neither of these gaps were present in earlier versions of the data at the time they were processed in cycles 3 and 4 and may be a problem with differences in distortion correction between the cycles. For the present data release, we have used neighboring pixels to interpolate over these ‘missing’ pixels in the small gap areas. As part of a future data release, we are planning updates to the REDUX package that will help with these and other issues in order to improve future versions of the survey mosaics.

2.5. Additional Data From The Literature

In our initial study of the *SOFIA*/FORCAST mosaics we compared the maps with a number of other prominent GC surveys from the literature including the *HST* Paschen- α Survey of the Galactic Center (Wang et al. 2010), the GLIMPSE *Spitzer*/IRAC survey (Churchwell et al. 2009), and the *Herschel* Hi-GAL survey (Molinari et al. 2010). These data sets provide a number of useful comparisons that we will discuss in subsequent sections. We also referenced maps and data from a number of other works in the literature which were focused on individual regions within the mosaics,

⁴ The median offset for a modest-sized sample (N=25) of randomly selected sources between the data sets is ~ 0.4 pixels or $\sim 0.3''$.

and these are discussed in the relevant sections in the text.

3. DISCUSSION

3.1. Comparison with other IR surveys

The GC is probably one of the most surveyed portions of sky and a number of high quality data sets exist at nearly all available wavelengths. In this section, we focus on comparisons between the FORCAST mosaics and earlier survey maps produced by *Spitzer*/IRAC at 8 μm and *Herschel*/PACS at 70 μm . Three-color combinations of these data sets can be found in Figure 2.

The morphology of the 8 μm emission contains numerous similarities, but also important differences when compared to the 25 and 37 μm emission. The origin of the 8 μm emission differs from that at 25 and 37 μm ; 8 μm emission primarily traces very small transiently heated grains, particularly Polycyclic Aromatic Hydrocarbons (PAHs) which have strong line emission in this wavelength region, while longer wavelength emission traces emission from larger, presumably silicate grains. Throughout the region, there is a diffuse 8 μm component corresponding to most of the IR bright regions at 25 and 37 μm , with the 8 μm emission appearing more extended in comparison. There are notable exceptions to this trend, however, which likely points to variations in extinction over the survey area. For example, Sgr B2 is relatively dark at 8 μm compared to many other regions within the survey, and it is well established that Sgr B2 suffers from significant local extinction (e.g., [Scoville et al. 1975](#))

Comparison of the 70 μm emission and the 25 and 37 μm emission highlights a number of notable features. The emission at each of these wavelengths is primarily thermal continuum with the 25 μm emission tracing relatively warm ~ 100 K dust and the 70 μm emission tracing cooler ~ 40 K dust. Many of the H II regions throughout the GC contain significant emission from warm dust which results in bright emission at 25 μm . As can be seen in Figure 2, this highlights the importance of the 25 μm data when studying recent star formation both in giant H II regions like the Arched Filaments and Sickle as well as in smaller H II regions which may be associated with individual massive stars. Furthermore, the spatial resolution of FORCAST enables us

to resolve many ‘bubble-like’ structures down to a physical scale of ~ 0.1 pc. For isolated sources, this is sufficient to resolve all but the most compact H II regions at the distance of the GC.

Another striking feature of the comparison map at 25, 37, and 70 μm is the number of compact red sources spread throughout the region. A prime example of this can be seen in Sgr B2, which contains several of the most luminous 70 μm sources in the GC (see §3.4). In the FORCAST data we clearly see these same features with strong emission at 37 μm , suggesting deeply embedded star formation.

3.2. The Sgr A Complex

One of the most interesting regions within the FORCAST mosaics is the Sgr A complex. As noted earlier, some of the earliest FORCAST observations of the GC were focused on the Circumnuclear Disk (CND; [Lau et al. 2013](#)); but the limited field of view prevented study of any large extended structures that might be associated. Our survey significantly expands the coverage of this region allowing us to examine various warm dust features extending from the position of the CND out to several parsecs. Among the highlights from this region is the mapping out of mid-infrared emission at the position of the NW X-ray lobe, extending from the CND ([Ponti et al. 2015](#)). In addition, we are able to see several other extended structures which have been noted in radio and ionized gas maps of the region, enabling us to compare the warm dust and ionized gas throughout the region (e.g., [Zhao et al. 2016](#)).

We note remarkably detailed structures in the warm dust emission in this region (Figure 3). Outside of the brightest mid-IR features tracing the ‘minispiral’ feature and the inner edge of the CND, also referred to as the Circumnuclear Ring (CNR), we observe fainter extended structures protruding from the CNR which correspond to the well-known “wing” features (often referred to as the NW and SE Wings; [Zhao et al. 2016](#)). Similar streamers are also detected in this region in molecular gas ([Liu et al. 2012](#)), and the emission observed by FORCAST may be related to these molecular clouds. Dust near the NW wing appears to extend for several arcminutes, terminating near the position of the source H1 in the H H II regions. This structure also appears to trace the

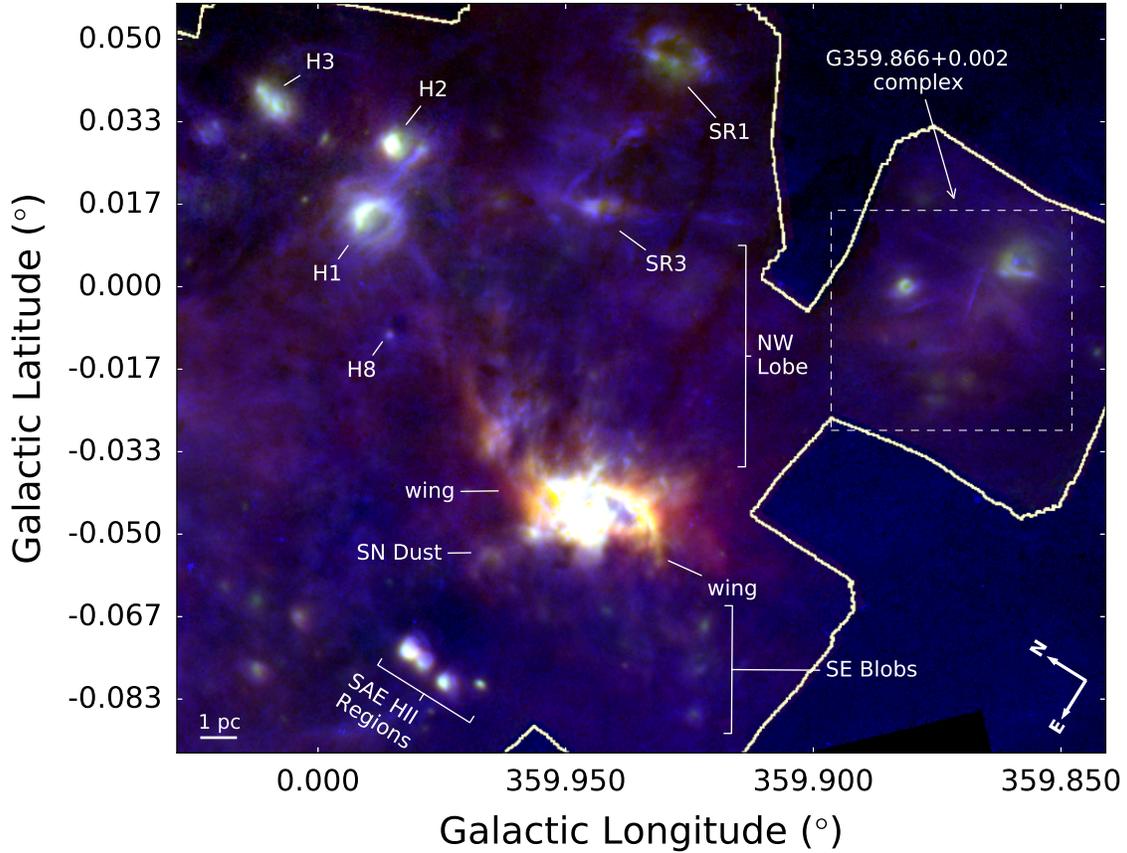


Figure 3. A false color map of the Sgr A complex created with the *HST* Paschen- α emission (blue), *SOFIA*/FORCAST 25 μm (green), and 37 μm (red) data. Several of the features discussed in text are labeled for reference. Sgr A East is abbreviated ‘SAE’ for brevity in the figure label. The dashed box shows the region around the G359.866+0.002 complex which is featured in a subsequent figure and an outline of the FORCAST survey footprint is also overlaid as a solid white outline.

edge of the NW X-ray lobe, although more detailed analysis is needed to understand possible interaction between the x-ray emitting region and the dusty infrared clouds.

At the position of the NW lobe we find relatively faint, extended mid-IR emission that extends northward to a few relatively bright 25 μm objects. These 25 μm sources appear to be associated with the ‘smoke rings’ discussed in Zhao et al. (2016). In the infrared, only one of these sources has a ring-like appearance (smoke ring #1; SR1 in Figure 3), and might be associated with the nearby massive star CXOGC J174516.7-285824 (Mauerhan et al. 2010). The IR emission of the third smoke ring (SR3) appears to ‘fill in’ the ring-like structure observed in radio and ionized gas, and suggests this feature may simply be a portion of a molecular cloud with prominent ionized edges. The second smoke ring structure from Zhao et al. (2016) has little to no detectable emission in

the mid-infrared wavelengths presented in this work.

To the east of the CNR we see the dust emission associated with the Sgr A East supernova remnant which was reported in Lau et al. (2015). To the south of the CNR there are a number of relatively faint sources and structures. The location of several of these features appear to correspond to the location of the ‘SE blobs’ reported by Zhao et al. (2016) and may be related to disrupted cloudlets in the region (see Figure 7 from Liu et al. 2012). The origin of material in this region is of particular interest because it could signify outflow from the CNR (e.g., Wang et al. 2010). We will explore these and other IR features in greater detail in a future paper focusing on Sgr A.

The H II region complex G359.866+0.002 (Figure 3) lies to the equatorial southwest of the CNR with a projected separation of $\sim 2.5'$ (~ 6 pc). While this complex is not typically

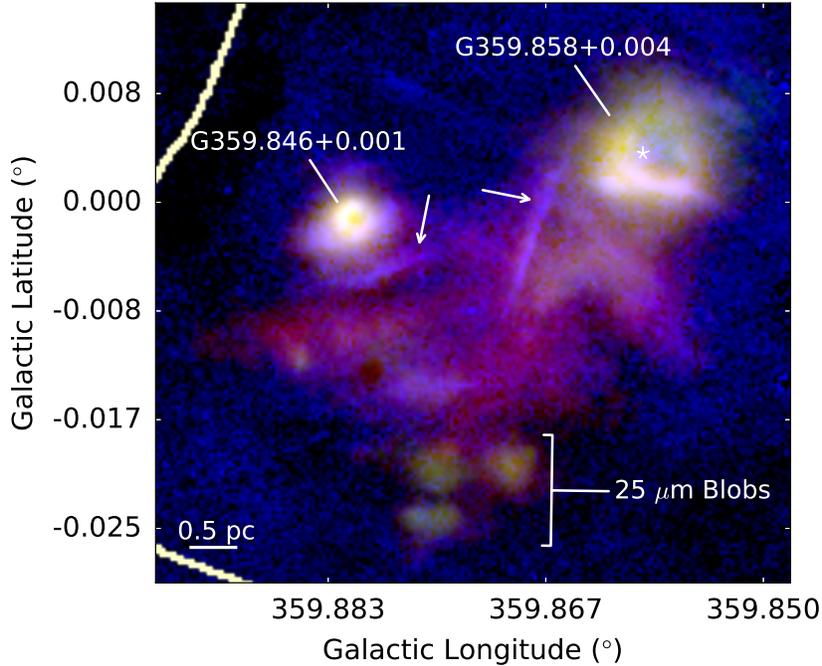


Figure 4. A false color map of the G359.866+0.002 complex created with the *HST* Paschen- α emission (blue), *SOFIA*/FORCAST 25 μm (green), and 37 μm (red) data. White arrows point to a few of the most prominent ionized gas filaments in the region. Several additional dusty ridges can be seen at 37 μm . The two dusty nebulae discussed in text are labeled and the location of the emission line star 2MASS J17451618-2903156 is marked with a white star. Three extended sources to the south are labeled as the 25 μm blobs.

discussed in the context of Sgr A, this region contains a number of interesting features and is the southernmost portion of our coverage of Sgr A. G359.866+0.002 was found as a collection of Paschen- α emitting sources in Wang et al. (2010) where it was featured for its linear ionized gas features. In the FORCAST data, we observe dust emission at 25 and 37 μm that traces many of these same narrow ionized gas features (Figure 4). These structures could point to locally strong magnetic fields or shocks that are sculpting the gas and dust emission. We also note two extended nebulae in the complex that are prominent at 25 μm , indicating dust that is likely being heated by luminous nearby stars. The 25 μm nebulae (G359.858+0.004) may be associated with the nearby emission line star 2MASS J17451618-2903156 from (Mauerhan et al. 2010), while the other 25 μm source (G359.846+0.001) is cataloged as a YSO candidate based on ISO GAL data (Immer et al. 2012), though this source is not present in the YSO catalog published in Yusef-Zadeh et al. (2009) and was not observed as part of the *Spitzer*/IRS sample of Galactic center YSOs presented in An et al. (2011). To the south of these sources there are a collection

of dusty ridges visible at 37 μm some of which have associated ionized gas emission and others which do not. Further to the south there is a collection of interesting extended 25 μm sources, which we refer to as 25 μm ‘blobs’. As part of a future data release for this program, we will produce source catalogs for the 25 and 37 μm data sets and publish additional information on possible YSO candidates determined from infrared colors of the *SOFIA* data.

3.3. The Arched Filaments and Sickle H II Regions

The Arched Filaments and Sickle H II regions are some of the most prominent structures in the GC at both infrared and radio wavelengths (e.g., Yusef-Zadeh & Morris 1987). Earlier work has been presented on *SOFIA*/FORCAST observations of both of these regions (e.g., Lau et al. 2014a, 2016; Hankins et al. 2016, 2017), however, our discussion of the present GC survey would be incomplete without some mention of these important regions. In this section we highlight how the combined archival and new survey data can improve our picture of the physical processes at

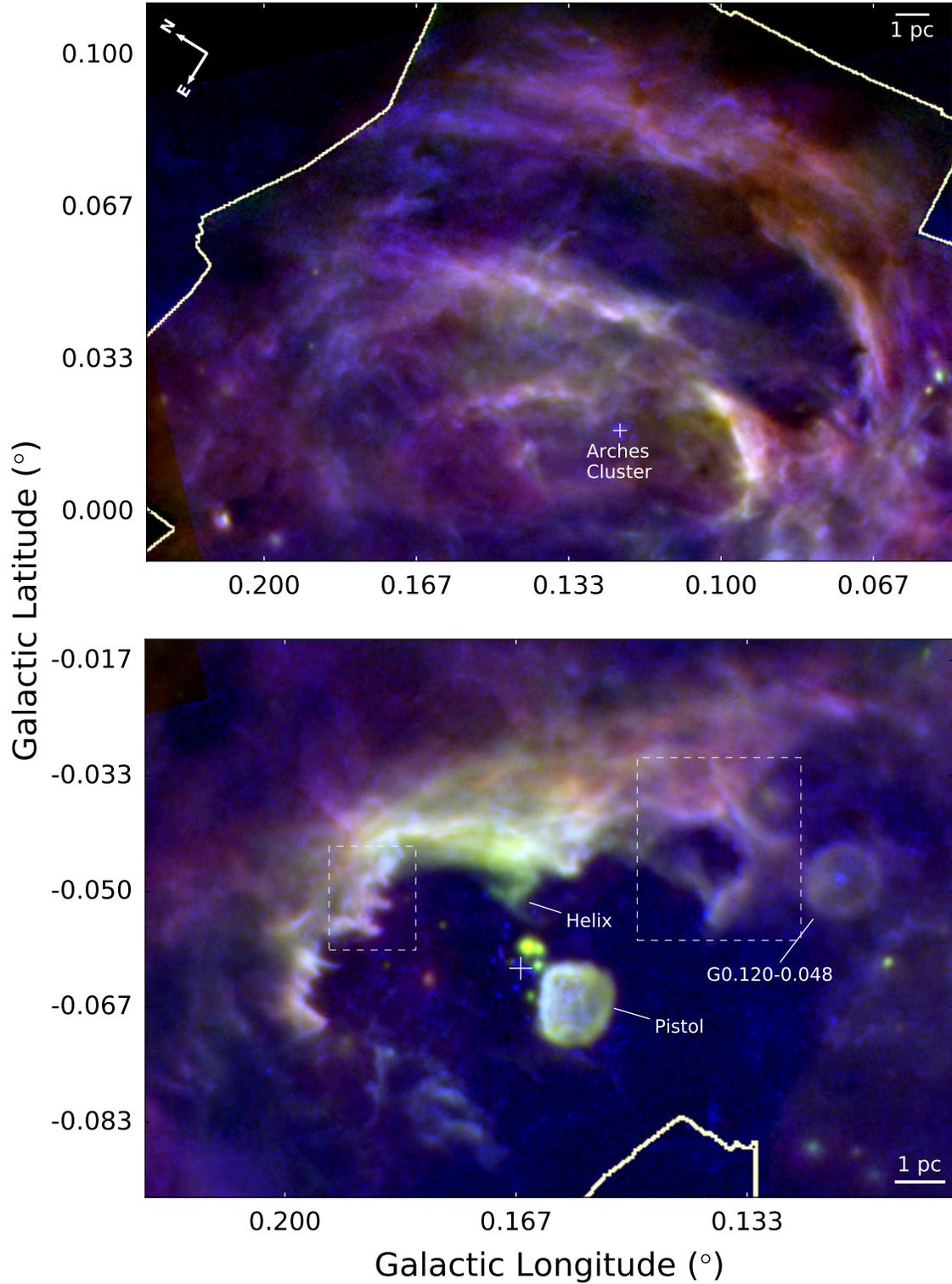


Figure 5. A false-color map of the Arched Filaments (Top) and the Sickle H II Region (Bottom) created with the *HST* Paschen- α emission (blue), *SOFIA*/FORCAST 25 μm (green), and 37 μm (red) data. The approximate locations of the Arches and Quintuplet clusters are marked with a white crosshair. Several infrared sources associated with the Quintuplet cluster can be seen near the marker, while the Arches cluster has a distinct lack of bright mid-infrared counterparts. Additional sources discussed in the text are labeled for reference, while the two dashed boxes indicate regions that are featured in a subsequent Figure. An outline of the FORCAST survey footprint is also overlaid as a solid white outline in both plots.

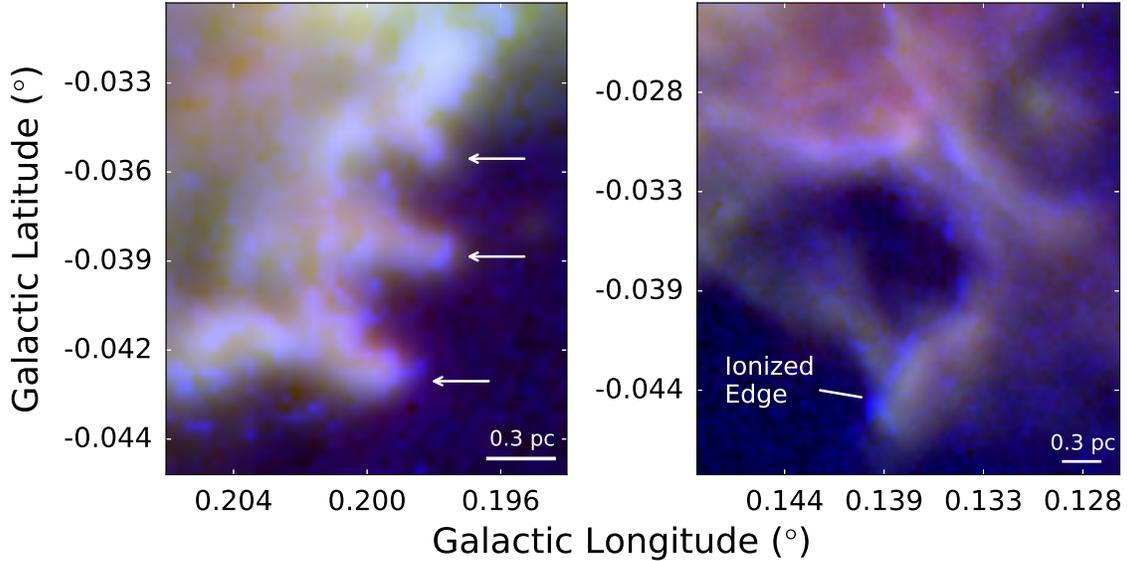


Figure 6. A ‘zoom-in’ on two interesting regions within the Sickle H II region. The left panel shows several of the finger-like structures in the ‘blade’ region. Arrows point out a few of the more prominent of these features which appear to measure $\sim 0.1\text{--}0.3$ pc in size. The right panel shows the ‘ladder’ region which shows prominent ionized gas emission (blue) on the edges of the dusty clouds. Both panels were created using the same 3-color map as figure 5.

work in the Arched filaments and Sickle (Figure 5)

The Arched Filaments complex is a collection of filamentary features that have a distinctive arched morphology due to our viewing geometry of the system (Lang et al. 2001, 2002). These structures are well known for their relation to the non-thermal emission in the GC Radio Arc, which meets with and undergoes a brightness discontinuity at the northern edge of the Arched Filaments (Lang et al. 2001) and appears to travel through the position of the Sickle H II region (Lang et al. 1997). Dust temperatures within the Arched Filaments H II region are consistent with the Arches cluster being the primary heating source (Hankins et al. 2017), as well as the primary source of ionizing radiation (e.g. Lang et al. 2001). Our previous FORCAST study of this region did not reveal any signs of ongoing star formation within the infrared bright filament structures (Figure 5 upper panel; Hankins et al. 2017). Furthermore, we do not detect any infrared point sources near the location of the Arches cluster at 25 or 37 μm which points to a lack of evolved, dust forming stars due to its relatively young age ($\sim 2\text{--}3$ Myr Figer et al. 1999; Stolte et al. 2002).

The Sickle H II region has a number of interesting features as seen both in the morphology of the Sickle proper, and in the numerous in-

frared sources primarily related to the Quintuplet cluster. First, the northern portion of the Sickle cloud (the ‘blade’) shows several ‘finger’-like structures (Figure 5 lower panel and Figure 6) reminiscent of features in M16 (Cotera et al. 2006). It is unclear if these structures are actively star forming like their M16 counterparts, although we may simply lack sufficient sensitivity and spatial resolution to adequately determine this at the GC distance. The morphological differences between the northern and eastern parts of the Sickle suggests they are impacted by the magnetic field organization within the cloud. Notably the perpendicular part of the cloud (the ‘handle’) does not display similar ‘fingers’. Although, there is a relatively faint, wispy, helical structure protruding from the ‘handle’ which Lau et al. (2016) suggest is related to outflow from a massive star that appears near the cloud edge.

Gas and dust within the Sickle are most likely ionized and heated by the highly luminous Quintuplet cluster ($L = 3 \times 10^7 L_{\odot}$ Figer et al. 1999). Unlike the Arches cluster, the Quintuplet cluster contains 5 IR bright ‘Quintuplet Proper Members’ from which the cluster name is derived (Okuda et al. 1990; Nagata et al. 1990). These objects are likely Wolf-Rayet stars which are undergoing active dust production (Tuthill et al. 2006; Najarro et al. 2017), and represent only a few of the massive stars

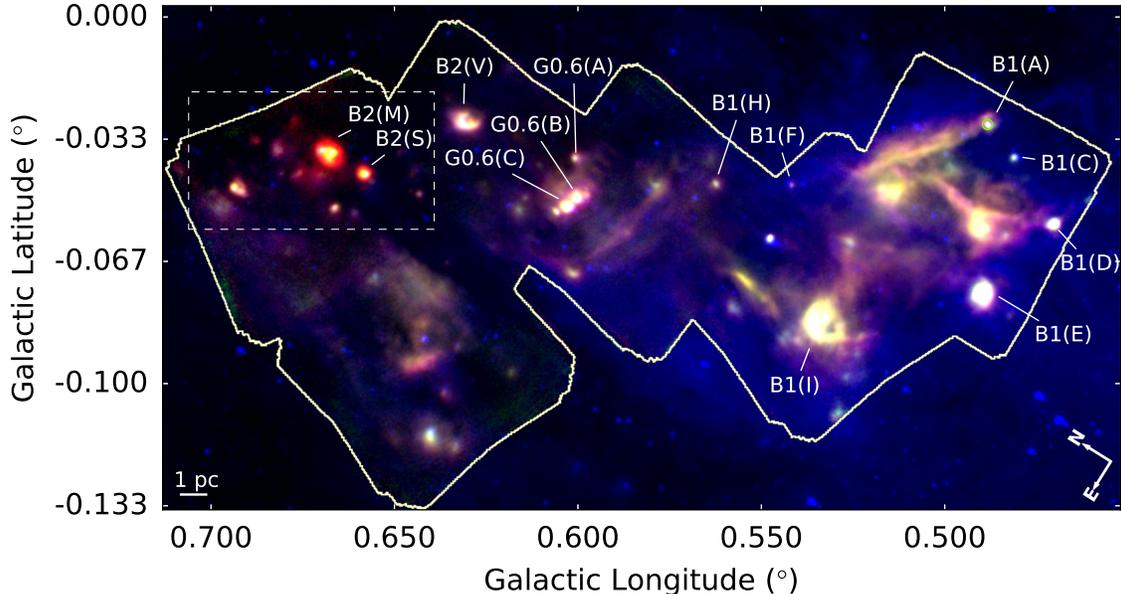


Figure 7. A false-color map of the Sgr B complex using *Spitzer*/IRAC 8 μm (blue), *SOFIA*/FORCAST 25 μm (green), and 37 μm (red) data. We use the IRAC 8 μm data in this figure rather than the *HST* Paschen- α survey because the Sgr B region is outside of the *HST* survey footprint. Sources of interest are labeled throughout the region following the naming convention of [Mehring et al. \(1992, 1993\)](#) with the exception of Sgr B2 Main and South. Sources in the G0.6-0.0 region are abbreviated ‘G0.6’, and the location of the massive YSO, SSTGC 726327, in Sgr B1 is marked with a green circle. The dashed box shows a portion of the map that is featured in a subsequent plot, and the footprint of the *SOFIA*/FORCAST survey is also shown as a solid white outline for reference.

belonging to the cluster (Figure 5 lower panel). Our survey observations also show dust associated with two extended nebulae surrounding the candidate Luminous Blue Variable (LBV) stars, the Pistol star and G0.120-0.048 ([Lau et al. 2014a](#)). While there is a third known LBV candidate in this region, qF362 ([Figer et al. 1999](#)), there is no obvious infrared counterpart associated with this source in the FORCAST maps. The data obtained as part of our survey of the Sickie region will be explored in depth in a subsequent paper focused on the ‘finger’-like structures and other prominent features of the H II region discussed in this section.

3.4. The Sgr B1 & Sgr B2 Complexes

The Sgr B complex is the easternmost region mapped in our survey. Sgr B is located along the galactic plane with a projected separation between $\sim 16'$ – $23'$ (~ 40 – 50 pc) from Sgr A as measured from the near and far side of Sgr B. This complex is most frequently discussed in terms of three distinct parts: Sgr B1, Sgr B2, and G0.6-0.0 which lies between Sgr B1 and Sgr B2 (Figure 7). Sgr B2 is one of the most massive and active star forming re-

gions in our Galaxy with a few hundred sources which are likely a mix of young stellar objects and slightly more developed stars that have produced H II regions ([de Pree et al. 1995, 1996](#); [Ginsburg et al. 2018](#)). In contrast, Sgr B1 is a more evolved star-forming region kinematically linked to Sgr B2, which contains a population of less extinct (embedded) H II regions ([Mehring et al. 1992, 1993](#); [Lang et al. 2010](#)). There is, however, recent evidence to suggest that these H II regions are a result of evolved massive stars passing through the dense medium in Sgr B1, as opposed to young O/B-stars that have formed within the cloud ([Simpson et al. 2018](#)).

The FORCAST observations of Sgr B reveal a number of interesting features throughout the region. Qualitatively, Sgr B1 shows considerable extended emission with large filamentary structures, shells, and bubbles most of which have radio counterparts, while Sgr B2 appears as a collection of compact reddened sources. A few of the brightest mid-infrared sources in this region include the well-known complexes Sgr B2 Main and Sgr B2 South as well as Sgr B2 V (e.g., [Etxaluze et al. 2013](#)). Near Sgr B2 Main and South we note a several relatively faint

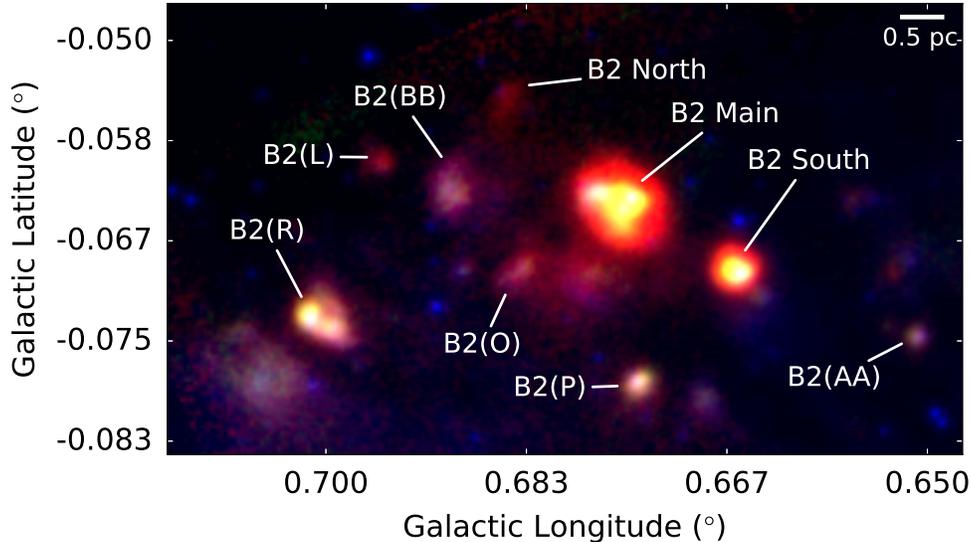


Figure 8. A zoom in on Sgr B2 with various sources of interest labeled following the naming convention from [Mehringer et al. \(1993\)](#) with the exception of Sgr B2 Main, North, and South. This figure was created using the same three-color map as in Figure 7.

IR sources which appear to be counterparts to known compact and ultra-compact H II regions (Figure 8). The presence of these sources provide further evidence for recent star formation activity outside of the brightest complexes in Sgr B2, which is consistent with recent ALMA observations that show evidence of extended star formation outside of the main clusters in Sgr B2 ([Ginsburg et al. 2018](#)). Study of the brightest and faintest mid-IR sources in this region will require use of both the FORCAST data set and the Spitzer/MIPS 24 μm data. As part of a future data release we are planning to create a combined map between these two datasets which will allow for study of regions like Sgr B2 where large dynamic range is needed to effectively dissect the region.

Although the FORCAST data of Sgr B1 does show a bright compact source coincident with one of the massive YSOs (SSTGC 726327) identified in [An et al. \(2011\)](#), the other candidates from that paper within our observations are unremarkable. There is, however, a strong ridge of warm dust emission immediately adjacent to one of the Wolf-Rayet stars found in [Mauerhan et al. \(2010\)](#) that does not have a strong radio counterpart and might be evidence of the impact of an orbiting massive star not formed in situ. [Simpson et al. \(2018\)](#) suggested, based on maps of the [O III] 52 and 88 μm lines in Sgr B1 with FIFI-LS, that the region may not, in fact, be forming stars as we

see in Sgr B2, but rather the observed emission may be the result of passing massive stars, such as is seen in the regions surrounding the Arches and Quintuplet clusters.

3.5. Sgr C & Neighbors

The Sgr C complex is the westernmost region mapped in our survey. Sgr C is located along the Galactic plane with a projected separation of $\sim 15'$ (~ 35 pc) from Sgr A. Our observations of Sgr C focused on the main Sgr C H II region with an additional adjacent pointing to the west containing the position of source ‘C’ from [Liszt & Spiker \(1995\)](#). Earlier measurements of ionized and molecular gas suggest that the H II region has a shell-like morphology which is likely created by massive stars that have blown out a cavity in the surrounding gas ([Lang et al. 2010](#)). The brightest portions of the extended 25 and 37 μm emission we observe in this region are associated with the Sgr C H II region (Figure 9). However, the morphology of the dust appears quite complex compared to a simple shell configuration.

We also observe the Sgr C Molecular Cloud as an infrared-dark source at 25 and 37 μm . The molecular cloud has a velocity of -55 km s^{-1} ([Kendrew et al. 2013](#)), indicating that it is likely associated with the H II region, which has a measured velocity of -65 km s^{-1} from recombination line emission ([Lang et al. 2010](#)). Toward the dark cloud we find a number of bright mid-infrared sources (Figure 10). The

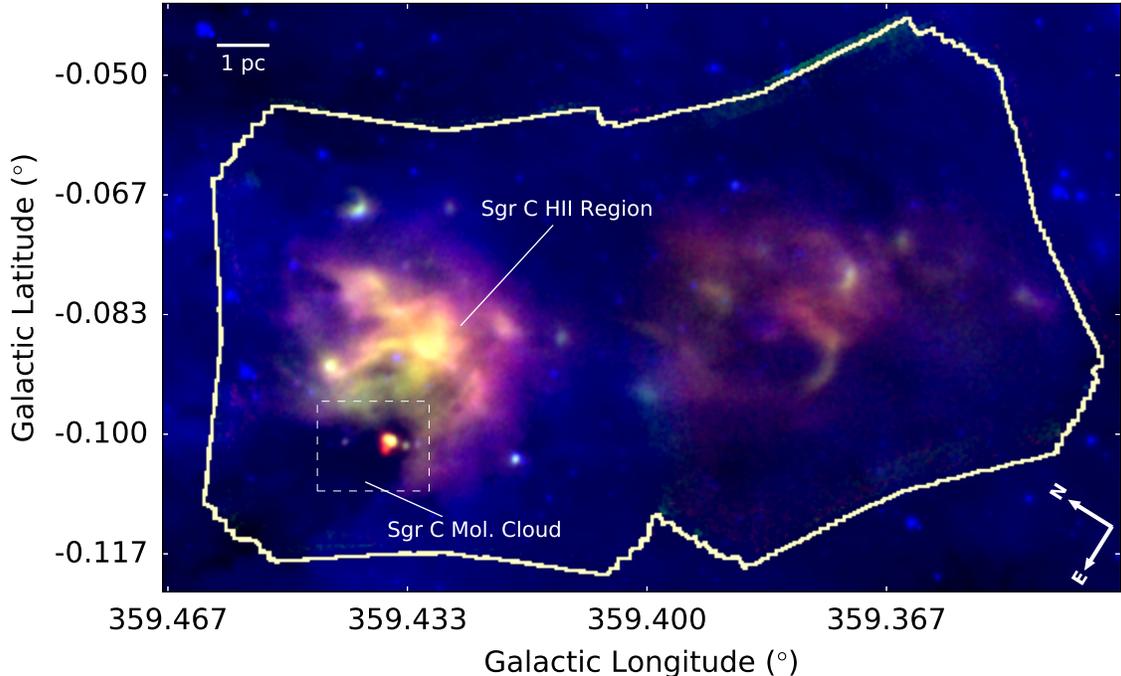


Figure 9. A false-color map of Sgr C using *Spitzer*/IRAC $8\ \mu\text{m}$ (blue), *SOFIA*/FORCAST $25\ \mu\text{m}$ (green), and $37\ \mu\text{m}$ (red) data. We use the IRAC $8\ \mu\text{m}$ data in this figure rather than the *HST* Paschen- α survey because the Sgr C region is outside of the *HST* survey footprint. A few regions of interest discussed in the text are labeled and the *SOFIA*/FORCAST survey footprint is also shown for reference.

most luminous of these sources (Sgr C H3) is reminiscent of the bright, red sources we observe in Sgr B2. This object has been previously cataloged as an ultra-compact H II region by Forster & Caswell (2000), and Kendrew et al. (2013) show that this source has additional substructure, including two dusty protostellar cores and signatures of an outflow which are indicative of ongoing high-mass star formation. There are two additional infrared sources near Sgr C H3, which are also cataloged as H II regions in Lu et al. (2019) and can be found in Figure 10.

The Sgr C complex is also known for its prominent non-thermal filament (NTF) (e.g., Roy 2003). Our survey did not cover the location of the filament because of our focus on the mid-infrared bright regions. However, the NTF should be discussed along with the H II region, because of the likely association of these features. In fact, large H II regions may be key to producing bright non-thermal features at radio wavelengths by providing a vast supply of free electrons that are accelerated to relativistic velocities via locally strong magnetic fields (Serabyn & Morris 1994; Uchida et al. 1996). There are numerous similarities be-

tween the well-known GC Radio Arc and the Sgr C NTF that warrant further study. This is another area where combining the MIPS and FORCAST data is needed to study these features in detail, and will be of interest with the future release of high level data products from this survey.

Between Sgr A and Sgr C there is a smattering of compact mid-infrared sources which have not garnered as much attention as other portions of the GC. These sources appear to be in somewhat isolated environments, which speaks to the apparent dichotomy between regions located at negative and positive Galactic longitudes, the latter of which are thought to be more actively forming stars (Longmore et al. 2013; Barnes et al. 2017). These isolated western sources are of interest because several are cataloged as candidate YSOs in the ISOGAL survey (Immer et al. 2012). Our observations of a handful of these sources show a diverse range of morphologies and colors. For example, G359.753-0.006 shows a prominent bow shock-like morphology (Figure 11). Several other nearby sources (G359.738-0.024, G359.726-0.037, G359.716-0.035, G359.64477-0.056, and G359.655-0.067, labeled respec-

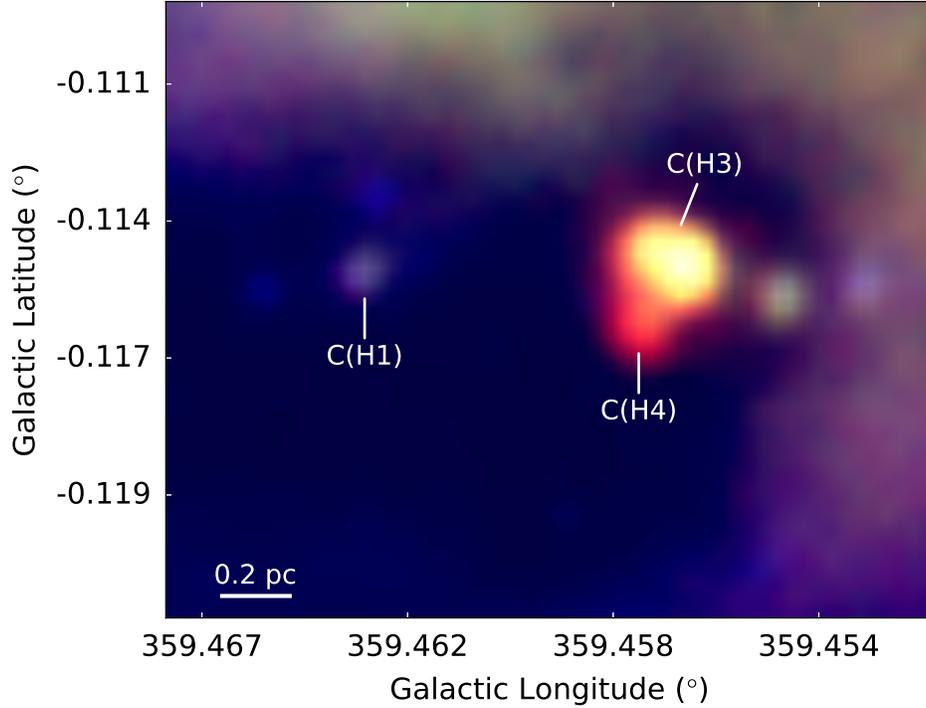


Figure 10. A zoom in on the Sgr C molecular cloud with sources labeled as in Lu et al. (2019). This figure was created using the same three-color map as in Figure 9.

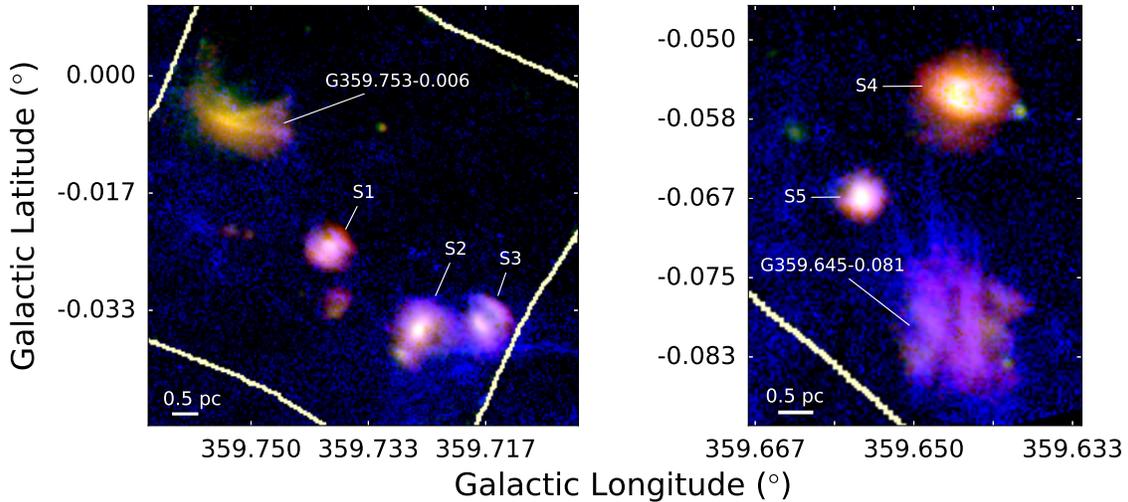


Figure 11. A collection of mid-IR sources located between Sgr A & Sgr C. Both panels are false-color images created using *HST* Paschen- α emission (blue), *SOFIA*/FORCAST 25 μm (green), and 37 μm (red) data. A few regions of interest discussed in the text are labeled and the *SOFIA*/FORCAST survey footprint is also shown for reference. These infrequently studied sources are discussed in section 3.5.

tively as S1, S2, S3, S4, and S5 in Figure 11) are more compact but still resolved in the FORCAST observations. These sources each show ionized gas emission and are possibly compact or ultra-compact H II regions. Interestingly, S2 and S3 may be interacting with a nearby

ionized gas filament which is visible in the Paschen- α data.

One of the most interesting sources in this region is G359.645-0.081, which is a relatively faint extended source. It is brighter at 37 μm compared to 25 μm , indicating that it is somewhat cool and is plausibly the edge of a molec-

ular cloud. However, the amount of extended ionized gas emission in addition to the relatively cool dust emission is somewhat peculiar compared to other extended sources in our survey. Prior observations of this region with *Spitzer*/IRS show a few high ionization species including [Ne V] 24.32 μm and [O IV] 25.9 μm near this position (Simpson 2018), indicative of a hard ionization source. Certainly this is a complex and unusual region worthy of further study.

4. SUMMARY

In this paper we have presented observations and initial results from the *SOFIA*/FORCAST Survey of the GC. Our survey focused on some of the brightest infrared regions in the GC which trace recent star formation. These data provide the highest spatial resolution mapping of the Galactic Center at 25 and 37 μm to date (FWHM \sim 2.3" and FWHM \sim 3.4", respectively), and cover several interesting regions which were badly saturated in the *Spitzer*/MIPS 24 μm data. Ultimately, our primary science objective for the survey is to better characterize star formation in the GC, in particular the well-known star formation rate discrepancy in this region. Examining this topic will require high-level data products, including source catalogs and enhanced maps that will be produced as a part of this program and released to the broader astronomical community in a future data release.

As an accompaniment to the initial public data release for the survey, we have produced this paper to describe the observations and interesting highlights from the data set. We have presented short summaries of several featured regions and sources including extended struc-

tures near the CND, the Arched filaments and Sickie H II regions, and embedded star formation in Sgr B2 and Sgr C. While primarily qualitative, these case studies illustrate the scope and utility of the survey data. We are planning future studies around several of these topics which will be presented in later works or possibly accompany future data releases.

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This research has made use of the VizieR catalogue access tool, CDS, Strasbourg, France. The original description of the VizieR service was published in A&AS 143, 23.

REFERENCES

- An, D., Ramírez, S. V., Sellgren, K., et al. 2011, *ApJ*, 736, 133
- Bally, J., Stark, A. A., Wilson, R. W., & Henkel, C. 1987, *ApJS*, 65, 13
- Barnes, A. T., Longmore, S. N., Battersby, C., et al. 2017, *MNRAS*, 469, 2263
- Barnes, A. T., Longmore, S. N., Avison, A., et al. 2019, *MNRAS*, 486, 283
- Calzetti, D., Kennicutt, R. C., Engelbracht, C. W., et al. 2007, *ApJ*, 666, 870
- Churchwell, E., Babler, B. L., Meade, M. R., et al. 2009, *PASP*, 121, 213
- Clarke, M., Vacca, W. D., & Shuping, R. Y. 2015, *Astronomical Society of the Pacific Conference Series*, Vol. 495, *Redux: A Common Interface for SOFIA Data Reduction Pipelines*, ed. A. R. Taylor & E. Rosolowsky, 355

- Cotera, A., Stolovy, S., Ramirez, S., et al. 2006, in *Journal of Physics Conference Series*, Vol. 54, *Journal of Physics Conference Series*, 183–189
- Cotera, A. S., Erickson, E. F., Colgan, S. W. J., et al. 1996, *ApJ*, 461, 750
- de Pree, C. G., Gaume, R. A., Goss, W. M., & Claussen, M. J. 1995, *ApJ*, 451, 284
- . 1996, *ApJ*, 464, 788
- Dong, H., Wang, Q. D., & Morris, M. R. 2012, *MNRAS*, 425, 884
- Dong, H., Lacy, J. H., Schödel, R., et al. 2017, *MNRAS*, 470, 561
- Egan, M., Price, S., Kraemer, K., et al. 2003, ADA418993
- Etxaluze, M., Goicoechea, J. R., Cernicharo, J., et al. 2013, *A&A*, 556, A137
- Figer, D. F. 2009, *Massive-star formation in the Galactic center*, ed. M. Livio & E. Villaver, 40–59
- Figer, D. F., Kim, S. S., Morris, M., et al. 1999, *ApJ*, 525, 750
- Forster, J. R., & Caswell, J. L. 2000, *ApJ*, 530, 371
- Fritz, T. K., Gillessen, S., Dodds-Eden, K., et al. 2011, *ApJ*, 737, 73
- Ginsburg, A., Bally, J., Barnes, A., et al. 2018, *ApJ*, 853, 171
- Guesten, R., & Henkel, C. 1983, *A&A*, 125, 136
- Guesten, R., Walmsley, C. M., Ungerechts, H., & Churchwell, E. 1985, *A&A*, 142, 381
- Habibi, M., Stolte, A., & Harfst, S. 2014, *A&A*, 566, A6
- Hankins, M. J., Lau, R. M., Mills, E. A. C., Morris, M. R., & Herter, T. L. 2019, *ApJ*, 877, 22
- Hankins, M. J., Lau, R. M., Morris, M. R., & Herter, T. L. 2017, *ApJ*, 837, 79
- Hankins, M. J., Lau, R. M., Morris, M. R., et al. 2016, *ApJ*, 827, 136
- Herter, T. L., Adams, J. D., De Buizer, J. M., et al. 2012, *ApJL*, 749, L18
- Herter, T. L., Vacca, W. D., Adams, J. D., et al. 2013, *PASP*, 125, 1393
- Immer, K., Schuller, F., Omont, A., & Menten, K. M. 2012, *A&A*, 537, A121
- Kendrew, S., Ginsburg, A., Johnston, K., et al. 2013, *ApJL*, 775, L50
- Krabbe, A., Genzel, R., Drapatz, S., & Rotaciuc, V. 1991, *ApJL*, 382, L19
- Krabbe, A., Genzel, R., Eckart, A., et al. 1995, *ApJL*, 447, L95
- Kruijssen, J. M. D., & Longmore, S. N. 2013, *MNRAS*, 435, 2598
- Lada, C. J., Forbrich, J., Lombardi, M., & Alves, J. F. 2012, *ApJ*, 745, 190
- Lang, C. C., Goss, W. M., Cyganowski, C., & Clubb, K. I. 2010, *ApJS*, 191, 275
- Lang, C. C., Goss, W. M., & Morris, M. 2001, *AJ*, 121, 2681
- . 2002, *AJ*, 124, 2677
- Lang, C. C., Goss, W. M., & Wood, O. S. 1997, *ApJ*, 474, 275
- Lau, R. M., Hankins, M. J., Herter, T. L., et al. 2016, *ApJ*, 818, 117
- Lau, R. M., Herter, T. L., Morris, M. R., & Adams, J. D. 2014a, *ApJ*, 785, 120
- . 2014b, *ApJ*, 794, 108
- Lau, R. M., Herter, T. L., Morris, M. R., Becklin, E. E., & Adams, J. D. 2013, *ApJ*, 775, 37
- Lau, R. M., Herter, T. L., Morris, M. R., Li, Z., & Adams, J. D. 2015, *Science*, 348, 413
- Lee, E. J., Murray, N., & Rahman, M. 2012, *ApJ*, 752, 146
- Liszt, H. S., & Spiker, R. W. 1995, *ApJS*, 98, 259
- Liu, H. B., Hsieh, P.-Y., Ho, P. T. P., et al. 2012, *ApJ*, 756, 195
- Longmore, S. N., Bally, J., Testi, L., et al. 2013, *MNRAS*, 429, 987
- Lu, J. R. 2018, *Astrophysics and Space Science Library*, Vol. 424, *Massive Young Clusters Near the Galactic Center*, ed. S. Stahler, 69
- Lu, X., Zhang, Q., Kauffmann, J., et al. 2019, *ApJ*, 872, 171
- Mauerhan, J. C., Cotera, A., Dong, H., et al. 2010, *ApJ*, 725, 188
- Mehring, D. M., Palmer, P., Goss, W. M., & Yusef-Zadeh, F. 1993, *ApJ*, 412, 684
- Mehring, D. M., Yusef-Zadeh, F., Palmer, P., & Goss, W. M. 1992, *ApJ*, 401, 168
- Mills, E. A. C. 2017, *arXiv e-prints*, arXiv:1705.05332
- Molinari, S., Swinyard, B., Bally, J., et al. 2010, *PASP*, 122, 314
- Morris, M., Polish, N., Zuckerman, B., & Kaifu, N. 1983, *AJ*, 88, 1228
- Morris, M., & Serabyn, E. 1996, *ARA&A*, 34, 645
- Muno, M. P., Bower, G. C., Burgasser, A. J., et al. 2006, *ApJ*, 638, 183
- Nagata, T., Woodward, C. E., Shure, M., Pipher, J. L., & Okuda, H. 1990, *ApJ*, 351, 83
- Najarro, F., Geballe, T. R., Figer, D. F., & de la Fuente, D. 2017, *ApJ*, 845, 127

- Okuda, H., Shibai, H., Nakagawa, T., et al. 1990, *ApJ*, 351, 89
- Ponti, G., Morris, M. R., Terrier, R., et al. 2015, *MNRAS*, 453, 172
- Ramírez, S. V., Arendt, R. G., Sellgren, K., et al. 2008, *The Astrophysical Journal Supplement Series*, 175, doi:10.1086/524015
- Reid, M. J. 1993, *ARA&A*, 31, 345
- Robitaille, T. P., Whitney, B. A., Indebetouw, R., Wood, K., & Denzmore, P. 2006, *ApJS*, 167, 256
- Roy, S. 2003, *A&A*, 403, 917
- Scoville, N. Z., Solomon, P. M., & Penzias, A. A. 1975, *ApJ*, 201, 352
- Serabyn, E., & Morris, M. 1994, *ApJL*, 424, L91
- Simpson, J. P. 2018, *ApJ*, 857, 59
- Simpson, J. P., Colgan, S. W. J., Cotera, A. S., Kaufman, M. J., & Stolovy, S. R. 2018, *ApJL*, 867, L13
- Steinke, M., Oskinova, L. M., Hamann, W.-R., et al. 2016, *A&A*, 588, A9
- Stolte, A., Grebel, E. K., Brandner, W., & Figer, D. F. 2002, *A&A*, 394, 459
- Tuthill, P., Monnier, J., Tanner, A., et al. 2006, *Science*, 313, 935
- Uchida, K. I., Morris, M., Serabyn, E., & Guesten, R. 1996, *ApJ*, 462, 768
- Wang, Q. D., Dong, H., Cotera, A., et al. 2010, *MNRAS*, 402, 895
- Young, E. T., Becklin, E. E., Marcum, P. M., et al. 2012, *ApJL*, 749, L17
- Yusef-Zadeh, F., Lacy, J. H., Wardle, M., et al. 2010, *ApJ*, 725, 1429
- Yusef-Zadeh, F., & Morris, M. 1987, *ApJ*, 320, 545
- Yusef-Zadeh, F., Hewitt, J. W., Arendt, R. G., et al. 2009, *ApJ*, 702, 178
- Zhao, J.-H., Morris, M. R., & Goss, W. M. 2016, *ApJ*, 817, 171