

## PROBING SHOCK BREAKOUT WITH SERENDIPITOUS *GALEX* DETECTIONS OF TWO SNLS TYPE II-P SUPERNOVAE

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### ABSTRACT

We report the serendipitous detection by *GALEX* of fast ( $<1$  day) rising ( $\geq 1$  mag) UV emission from two Type II plateau (II-P) supernovae (SNe) at  $z = 0.185$  and  $0.324$  discovered by the Supernova Legacy Survey. Optical photometry and VLT spectroscopy 2 weeks after the *GALEX* detections link the onset of UV emission to the time of shock breakout. Using radiation hydrodynamics and non-LTE radiative transfer simulations, and starting from a standard red supergiant (RSG; Type II-P SN progenitor) star evolved self-consistently from the main sequence to iron core collapse, we model the shock breakout phase and the 55 hr that follow. The small scale height of our RSG atmosphere model suggests that the breakout signature is a thermal soft X-ray burst ( $\lambda_{\text{peak}} \sim 90 \text{ \AA}$ ) with a duration of  $\approx 2000$  s. Longer durations are possible but require either an extended and tenuous nonstandard envelope or an unusually dense RSG wind with  $\dot{M} \sim 10^{-3} M_{\odot} \text{ yr}^{-1}$ . The *GALEX* observations miss the peak of the luminous ( $M_{\text{FUV}} \approx -20$ ) UV burst but unambiguously capture the rise of the emission and a subsequent 2 day long plateau. The postbreakout, UV-bright plateau is a prediction of our model in which the shift of the peak of the spectral energy distribution (SED) from  $\sim 100$  to  $\sim 1000 \text{ \AA}$  and the ejecta expansion both counteract the decrease in bolometric luminosity from  $\sim 10^{11}$  to  $\sim 10^9 L_{\odot}$  over that period. Based on the observed detection efficiency of our study, we make predictions for the breakout detection rate of the *GALEX* Time Domain Survey.

*Subject headings:* radiative transfer — stars: atmospheres — supernovae: general — ultraviolet: ISM

### 1. INTRODUCTION

In recent years, with the advent of sensitive  $\gamma$ -ray and X-ray detectors, high-energy radiation from stellar objects has been observed routinely. This includes in particular the long-soft  $\gamma$ -ray bursts (GRBs), associated at present with a subset of hydrogen- and helium-deficient (Type Ic) SNe (Woosley & Bloom 2006). Their duration of  $\geq 2$  s suggests a compact progenitor with a  $\sim 1$  s light-crossing time, a Wolf-Rayet (WR) star. Observations have revealed that there is in fact a continuum from  $\gamma$ -ray to X-ray signatures in association with such SNe. Hence, some events are related to GRBs in association with relativistic ejecta, while others are associated with X-ray flashes and the SN shock breaking out into the progenitor's dense wind, as in the Type Ic SN 2006aj and the Type Ib SN 2008D (Campana et al. 2006; Waxman et al. 2007; Soderberg et al. 2008; Li 2007, 2008). In such SNe, the subsequent thermal emission from the cooling stellar envelope has also been observed (e.g., SN Ic 1998bw: Galama et al. 1998; SN Ib/c 1999ex: Stritzinger et al. 2002; SNe 2006aj and SN 2008D).

In contrast to Type Ib/c SNe, hydrogen-rich (Type II) SNe originate from main-sequence objects in the mass range  $8\text{--}25 M_{\odot}$ , with iron cores at the time of collapse buried under an envelope of  $\sim 6\text{--}20 M_{\odot}$ . The envelope light-crossing time is  $\sim 0.65 R_3$  hr ( $R_3$  is the progenitor radius in  $10^3 R_{\odot}$ ) and varies from about 5 minutes for a blue supergiant (BSG) progenitor to 1 hr in the largest RSGs (Levesque et al. 2005). In Type II SNe, these two properties prevent a central engine from powering relativistic ejecta and from producing their associated nonthermal  $\gamma$ -ray/X-ray radiation. Hence, a burst of high-energy radiation from a Type II SN is most likely associated with shock emergence at the surface of the progenitor and its immediate aftermath.

The shock-breakout phenomenon has been studied theoretically (Falk 1978; Klein & Chevalier 1978), with specific attention to the BSG progenitor of SN 1987A (Ensmann & Burrows 1992; Blinnikov et al. 2000) or Type Ib/c SNe (Blinnikov et al. 2002). When the radiation-dominated SN shock wave travels outward through the envelope and reaches regions of moderate optical depth (i.e.,  $\tau$  of a few tens), a radiative precursor initiates the brightening of the object leading to a brief soft X-ray burst. This precursor starts when the radiative diffusion time ( $t_{\text{diff}} \sim 3\tau H_p/c$ ;  $H_p$  is the atmospheric scale height) becomes shorter than the shock-travel time to the surface, located at  $\tau \sim 1$ . The envelope structure and extent are two key ingredients controlling the properties of the radiative precursor. In the absence of a wind, the expected durations are  $\sim 1$  s in a WR star (Blinnikov et al. 2002),  $\sim 100$  s in a BSG (Ensmann & Burrows), and  $\sim 1000$  s in a RSG (Klein & Chevalier 1978). A very optically thick wind, with a scale height comparable to the stellar radius, may, however, dramatically increase the diffusion time and consequently the precursor duration.

In this Letter, we report the serendipitous detection with *GALEX* of rising UV emission in two Type II-P SNe discovered by the Supernova Legacy Survey (SNLS), SNLS-04D2dc (also reported by Schawinski et al. 2008) and SNLS-06D1jd, and associate both events with the breakout phase. This contrasts

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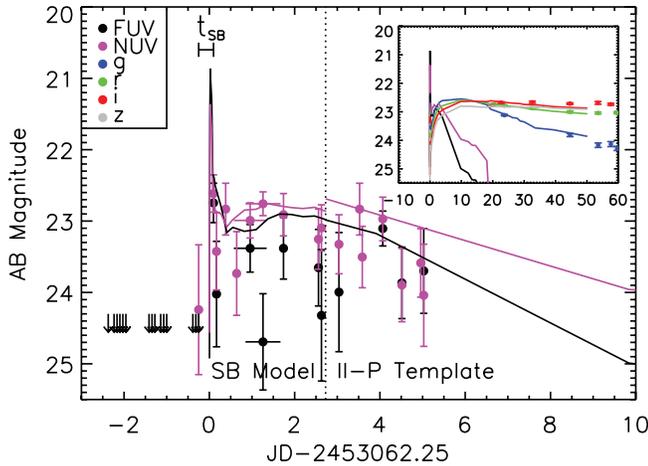


FIG. 1.—*GALEX* and SNLS (*inset*) light curves (colored symbols) of SNLS-04D2dc ( $z = 0.185$ ) shown relative to the estimated breakout time (horizontal bar, labeled  $t_{\text{SB}}$ ) and overplotted with synthetic magnitudes (colored lines; II-P Template: scaled synthetic light curves of SN 2006bp [D08]; SB Model: shock-breakout model). Horizontal error bars show the time coverage of photometry measured from co-added observations. Solid arrows show 95% upper limits. The *GALEX* FUV detector was temporarily not operational during the NUV observations around  $t \sim 3.5$  days.

with previous early UV observations of Type II SNe, i.e., the Type II-pec SN 1987A (Hamuy et al. 1988), the Type IIb SN 1993J (Schmidt et al. 1993; Richmond et al. 1994), and the Type II-P SN 2006bp (Immler et al. 2007; Dessart et al. 2008, hereafter D08), in which they were all *fading* in the UV, since the SN ejecta had already begun their long-term cooling phase through expansion and radiation. In § 2, we describe the *GALEX* and SNLS observations. In § 3, we present our numerical simulations of the breakout and the early postbreakout evolution in a standard RSG progenitor and compare these results to observations. In § 4, we conclude and make predictions for the detection rate by the *GALEX* Time Domain Survey (TDS). Throughout this Letter we adopt  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_M = 0.3$ , and  $\Omega_\Lambda = 0.7$ .

## 2. OBSERVATIONS

We conducted a search for early UV emission in Type II SNe by matching SNLS SN candidates with UV sources in overlapping *GALEX* Deep Imaging Survey (DIS) (Martin et al. 2005) fields. We selected SN candidates that had  $\sim 1.5$  ks *GALEX* DIS observations within 30 days before the optical SN candidate discovery, yielding a sample of 293 SN candidates with 1 to 77 *GALEX* observations within that time period. Two variable UV sources were detected with a  $3''$  matching radius to the position of SN candidates SNLS-04D2dc (R.A. =  $10^{\text{h}}00^{\text{m}}16.681^{\text{s}}$ , decl. =  $+02^{\circ}12'18.52''$  [J2000.0]) and SNLS-06D1jd (R.A. =  $02^{\text{h}}27^{\text{m}}36.189^{\text{s}}$ , decl. =  $-04^{\circ}31'56.62''$  [J2000.0]). The *GALEX* photometry was measured using aperture magnitudes with  $r_{\text{ap}} = 6''$  and a centroid fixed to that of the highest signal-to-noise (S/N) detection, with an aperture correction to account for the fraction of energy enclosed in the aperture of  $m_{\text{tot}} - m_{\text{ap}} = -0.15$  and  $-0.23$  mag in the FUV ( $\lambda_{\text{eff}} = 1539 \text{ \AA}$ ) and NUV ( $\lambda_{\text{eff}} = 2316 \text{ \AA}$ ), respectively (Morrissey et al. 2007).

SNe SNLS-04D2dc and SNLS-06D1jd were discovered on 2004 March 16 and 2006 December 12 by the SNLS real-time difference imaging pipeline (Astier et al. 2006; Sullivan et al. 2006). Follow-up VLT-FORS1 spectroscopy was obtained on 2004 March 19 (JD 2,453,083.5) and 2006 December 22 (JD 2,454,091.5), corresponding to 21.2 and 23.4 days after the

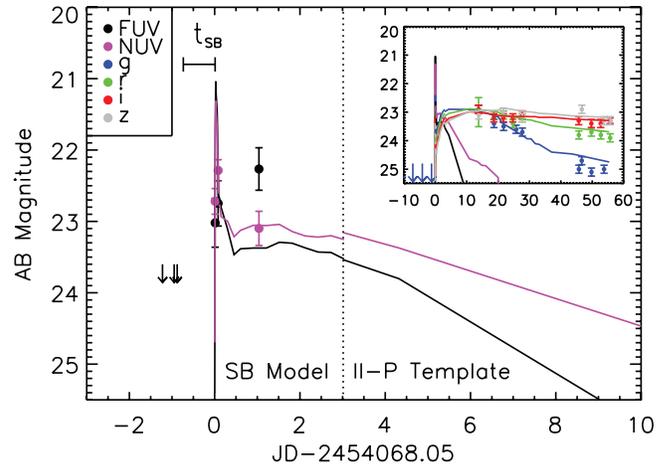


FIG. 2.—*GALEX* and SNLS (*inset*) light curves (colored symbols) of SNLS-06D1jd ( $z = 0.324$ ), shown relative to the estimated breakout time (horizontal bar, labeled  $t_{\text{SB}}$ ) and overplotted with synthetic magnitudes (colored lines; II-P Template: scaled synthetic light curves of SN 2006bp [D08]; SB Model: shock-breakout model). Solid arrows show 95% upper limits.

*GALEX* detection, respectively. The SNe were classified as Type II-P SNe based on the identification of broad H I emission lines (most noticeably H $\beta$  and H $\gamma$  in SNLS-06D1jd, but also H $\alpha$  in SNLS-04D2dc) and a  $\sim 40$  day plateau in optical brightness. The host galaxy redshifts are measured from nebular emission lines and the spectra and photometry are fitted with templates of a late-type galaxy at  $z = 0.185$  with  $m_{\text{r}} = 19.6$  for SNLS-04D2dc and a starburst galaxy at  $z = 0.324$  with  $m_{\text{r}} = 22.1$  for SNLS-06D1jd. Both VLT spectra match well the observations of SN 1999em at  $\sim 17$  days after breakout (Dessart & Hillier 2006), suggesting that the observed UV brightening is contemporaneous to the breakout phase.

The host galaxies are detected in the FUV and NUV with  $m_{\text{FUV}} = 23.42 \pm 0.07$  ( $23.80 \pm 0.08$ ) and  $m_{\text{NUV}} = 22.94 \pm 0.06$  ( $23.51 \pm 0.07$ ) for SNLS-04D2dc (SNLS-06D1jd) measured from deep co-added images constructed from observations taken before the SN detection with  $t_{\text{exp}} = 37.6$  ks (64.0 ks). We show the light curves of SNe SNLS-04D2dc in Figure 1 and SNLS-06D1jd in Figure 2, including *GALEX* FUV and NUV photometry (after subtraction of the host galaxy contribution) and SNLS difference imaging photometry in the  $g'$  ( $\lambda_{\text{eff}} = 4872 \text{ \AA}$ ),  $r'$  ( $\lambda_{\text{eff}} = 6282 \text{ \AA}$ ),  $i'$  ( $\lambda_{\text{eff}} = 7776 \text{ \AA}$ ), and  $z'$  ( $\lambda_{\text{eff}} = 8870 \text{ \AA}$ ) bands. Note the  $\geq 12$  day gap between the UV and optical observations of both SNe. In order to relate the UV and optical photometry, we create synthetic *GALEX*- and SNLS-filter light curves using the multipoch non-LTE synthetic spectra computed for the Type II-P SN 2006bp ( $d = 17.5 \text{ Mpc}$ ; D08) and adjust to the distance, redshift, and extinction of SNLS-04D2dc and SNLS-06D1jd. For each, we adopt the onset of UV emission as the breakout time, yielding JD  $2,453,062.25^{+0.10}_{-0.25}$  (JD  $2,454,068.05^{+0.01}_{-0.74}$ ) for SNLS-04D2dc (SNLS-06D1jd). For SN 2006bp, we adopt JD 2,453,834.2 (Quimby et al. 2007). The optical colors during the plateau of both SNLS SNe are well fitted with a reddening of  $E(B - V) \sim 0.1$ . However, an offset of  $+0.3$  ( $-0.4$ ) for SNLS-04D2dc (SNLS-06D1jd) was added to obtain the best match to the observations, suggestive of a difference in intrinsic brightness of these SNe compared to SN 2006bp or an uncertainty in the extinction.

## 3. MODELS AND RESULTS

We model the UV and optical light curves from 0.05 to 55.6 hr after shock breakout in three consecutive steps. First, using

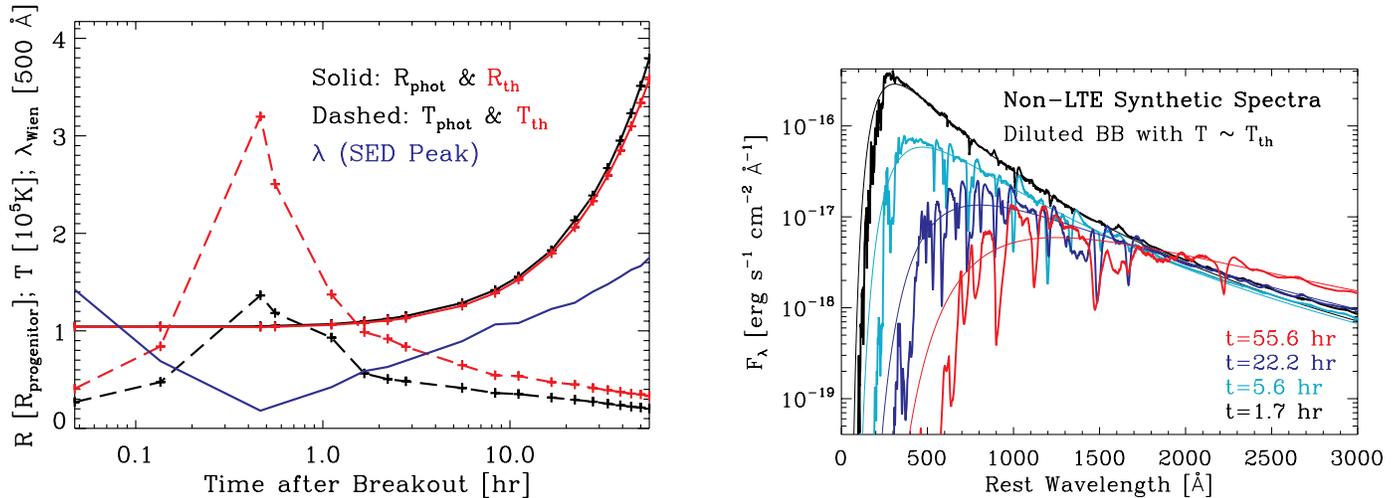


Fig. 3.—*Left*: KEPLER/CMFGEN predictions for the postbreakout evolution of the wavelength of the SED peak (blue line; in units of 500 Å), together with the gas temperature ( $T_{\text{phot}}$ : dashed line) and radius ( $R_{\text{phot}}$ : solid line) at the photosphere (black line) and at the thermalization depth (red line). *Right*: Multi-epoch postbreakout synthetic non-LTE (thick line) and diluted blackbody (thin line;  $T_{\text{BB}} \sim T_{\text{th}}$ ) spectra, scaled to the luminosity distance of SNLS-06D1jd.

the one-temperature Lagrangian radiation hydrodynamics code KEPLER (Weaver et al. 1978), the evolution at solar metallicity of a  $15 M_{\odot}$  main-sequence star is computed at high resolution ( $\Delta M \sim 10^{-6} M_{\odot}$  at the surface) until iron core collapse. At that time, the star is a RSG with a radius  $R_* = 865 R_{\odot}$  and a total mass of  $12.58 M_{\odot}$  (Woosley & Heger 2007). The star above a mass cut of  $1.45 M_{\odot}$  is then artificially exploded with a piston to yield an asymptotic SN ejecta kinetic energy of  $1.2 \times 10^{51}$  ergs, typical of a core-collapse SN.<sup>12</sup> We finally postprocess a sequence of 20 KEPLER models with the non-LTE nonrelativistic steady-state code CMFGEN (Hillier & Miller 1998; Dessart & Hillier 2005a). The solution for level populations and the radiation field at  $\sim 10^5$  frequencies is, thus, iterated to convergence on the stationary hydrodynamical structure, but allowing the temperature to relax in regions exterior to a Rosseland mean optical depth of  $\sim 20$ . This ensures that our assumption of steady-state is consistent, i.e., the diffusion time at the RSG surface being much less than the SN age. This is the first time the phase contemporaneous to shock breakout is modeled with a non-LTE multiangle multifrequency radiative transfer code, which allows for both line and continuum sources of opacities and treats multiple species and ionization states, including H I, He I–II, C II–V, N II–V, O II–VI, Ne IV–VIII, and Fe IV–XI and corresponding to a total of  $\sim 1000$  levels.

The one-temperature treatment adopted in KEPLER forces the gas and the radiation to the same temperature and, thus, prevents our modeling of the radiative precursor. However, the small atmospheric scale height  $H_p \sim 0.01 R_*$  in our RSG progenitor surface suggests that the precursor will last  $\sim 1000$  s, light-travel time effects smearing the observed signal over  $\sim R_*/c \sim 2000$  s. Longer durations are possible but would require a larger atmospheric scale height, perhaps reaching  $H_p \sim R_*$ . This may be produced by an extended low-density envelope, as adopted by Schawinski et al. (2008), although this appears incompatible with the RSG envelope structure computed with KEPLER, the 3D hydrodynamics simulations of RSG envelopes (Freytag et al. 2002), and the quantitative spectroscopic analyses of RSG stars

(Josselin & Plez 2007). Alternatively, this may be produced by an RSG in a superwind phase, i.e., an RSG with an unusually high mass-loss rate  $\dot{M} \sim 10^{-3} M_{\odot} \text{yr}^{-1}$  compared to the more standard values of a few  $\times 10^{-5} M_{\odot} \text{yr}^{-1}$  (Salasnich et al. 1999).

In Figure 3a, we show in black the evolution of the temperature  $T_{\text{phot}}$  and radius  $R_{\text{phot}}$ , defined by  $\int_{R_{\text{phot}}}^{\infty} \kappa_t dR = 2/3$ , where  $\kappa_t = \kappa_a + \sigma$  and  $\kappa_a$  ( $\sigma$ ) is the absorptive (scattering) opacity.  $T_{\text{phot}}$  reaches up to  $\sim 1.35 \times 10^5$  K only  $\sim 0.5$  hr after breakout, associated with a brightening phase at constant radius (light-travel time effects and the radiative precursor would yield a  $\sim 2000$  s signal instead). Note that at this time, the “RSG” has the surface temperature of a white dwarf and radiates at  $10^{11}$  times the rate of the Sun! The temperature, decreasing almost as quickly as it rose and reaching down to 40,000 K after  $\sim 5$  hr, dominates the bolometric variation since mass shells have hardly moved during this early phase. At 55 hr after breakout,  $T_{\text{phot}} \sim 20,000$  K, and  $R_{\text{phot}}$  has increased by only a factor of 3.8 to  $\sim 2.3 \times 10^{14}$  cm. KEPLER predicts a steep density distribution at the photosphere, scaling as  $1/R^{200}$  at breakout to  $1/R^{50}$  55 hr later, a property also inferred from line profile morphology in SN 2006bp by D08. With electron scattering dominating the opacity, the emergent SN flux suffers dilution (Dessart & Hillier 2005b), reflecting better the temperature at the thermalization radius  $R_{\text{th}}$  (defined by  $\int_{R_{\text{th}}}^{\infty} \sqrt{\kappa_a \kappa_t} dR = 2/3$ ) rather than at the photosphere (Fig. 3a). These temperatures differ by up to a factor of 2–3 at all times, and the SED is much harder than the blackbody distribution at  $T_{\text{phot}}$ , with a peak (blue curve in Fig. 3a) at  $\sim 90$  Å (900 Å) at 0.5 hr (55.6 hr) after breakout. The non-LTE synthetic spectra shown in Figure 3b illustrate these properties and reveal for the first time the line features predicted under these breakout conditions (these characteristics will be discussed in a future paper).

In Figures 1 and 2, we show the model *GALEX* and SNLS light curves created from our synthetic spectra for SNLS-04D2dc and SNLS-06D1jd. The CMFGEN models ( $1.67 \text{ hr} \leq t \leq 55.6 \text{ hr}$ ) are offset by  $-1.5$  ( $-1.4$ ) mag for SNLS-04D2dc (SNLS-06D1jd) to match the SN 2006bp template light curves ( $t > 2$  days) that were scaled to fit the optical observations, suggesting that our models underestimate the observed luminosity by a factor of  $\sim 4$ . We find no significant variability in the photon data between the first and second halves of the first two  $\sim 20$  minute

<sup>12</sup> However, originally, about half of the energy deposited by the shock wave is in the form of heat and trapped radiation (internal energy), so that at shock breakout, i.e., 1.47 days after the piston trigger, the bulk of the SN ejecta in our model has not yet reached its coasting velocity.

observations during the rise of the UV emission from SNLS-04D2dc and SNLS-06D1jd, with  $\Delta m < 0.5$  mag. Although the lack of significant intraobservation variability and the  $\sim 2$  mag amplitude of variability predicted by our models for the  $\leq 1$  hr long breakout signature implies that the *GALEX* observations do not coincide with the peak of the UV burst, the onset of the fast brightening UV emission is suggestive evidence that we have detected the breakout phase and constrains the time of shock breakout to an accuracy of 0.35 (0.75) days for SNLS-04D2dc (SNLS-06D1jd), i.e., 7.1 hr (13.7 hr) in the SN rest frame. Within the errors, the subsequent 2 day long UV plateau is well reproduced by the model, although we note the poor match to the noisy FUV data in the plateau of SNLS-04D2dc and to the FUV/NUV color of SNLS-06D1jd at  $\sim 1$  day. Future, better sampled, higher S/N UV light curves of breakout events in Type II-P SNe will provide more accurate constraints, foster a better understanding, and, thus, permit more definite conclusions on this ephemeral breakout phase and the RSG envelope structure.

#### 4. DISCUSSION AND CONCLUSION

*GALEX* DIS serendipitously detected fast rising UV emission in two SNLS Type II-P SNe, which we associate with the breakout phase in a RSG star. The observed amplitude of the UV signal is smaller than that of the  $\leq 1$  hr long breakout burst predicted by our models, which suggests that the observations did not capture the peak of the burst (we see no strong variability within the 20 minute observations), although our theoretical peak flux is likely overestimated by the one-temperature treatment in KEPLER. The important conclusion from these models is that the observed UV brightening by  $\sim 1$  mag is predicted to occur only during this early, breakout, phase and thus is a direct probe of the time of shock breakout in these SNe. Moreover, we both observe and predict that the immediate postbreakout in the UV is followed by a plateau phase of about 2 days, supporting a RSG-like progenitor (more compact progenitors, e.g., BSGs, cool much faster after breakout, as seen for SN 1987A) and providing an interesting clock to gauge the time since breakout. In SN 2006bp, the SN was fading in the UV right from the start of data collection, and a time since explosion of 1.7 days was indeed inferred (D08). The postbreakout UV plateau is the result of two competing effects, with the strong decrease in bolometric luminosity compensated by the shift of the SED to longer wavelength as the ejecta expand and cool (Fig. 3a). At maximum light, the peak of the SED in our models is at 90 Å (in the rest frame), and so the UV rise reveals only the evolution of the tail of the SED, not that of the peak.

The volumetric rate of Type II-P SNe is close to that of Type Ia SNe (Dahlen et al. 2004); thus, we use the measured fraction of Type Ia SNe in the total sample of SNLS SN candidates from Sullivan et al. (2006),  $f_{\text{Ia}} = 0.16$ , to estimate the fraction of Type II-P SNe in our sample. If we assume that the shock breakout occurs  $15 \pm 3$  days before the optical discovery of the SN, then the number of candidates in our study that had a probability of at least one *GALEX* observation during the 2 day UV plateau following shock breakout is  $n_{\text{obs}} = 167$ . By contrast, the number of candidates that had a probability of having a *GALEX* observation that coincided with a 1 hr long flash at the time of shock breakout is  $\leq 1$ . We now estimate the number of Type II-P SNe observed during the UV-plateau phase as  $n_{\text{II-P}} \approx f_{\text{Ia}} n_{\text{obs}} \sim 27$  and get a detection efficiency of  $f_{\text{det}} = 2/n_{\text{II-P}} \sim 0.07$ . *GALEX* TDS will operate in parallel with the Pan-STARRS Medium Deep Survey (PS MDS), an optical time domain survey that will monitor  $50 \text{ deg}^2 \text{ night}^{-1}$  with a depth comparable to SNLS. *GALEX* TDS plans to monitor  $\sim 7 \text{ deg}^2$  of the PS MDS with cadences ranging from 98.6 minutes (two consecutive orbits) to several days for  $\sim 8$  months of the year to a depth of  $m_{\text{lim}} \sim 23$  mag. PS MDS is expected to have a SN Ia detection rate of  $r_{\text{Ia}} \approx 23 \text{ month}^{-1}$  per  $7 \text{ deg}^2$  field (J. Tonry 2008, private communication), which for our observed detection efficiency ( $f_{\text{det}} \sim 0.07$ ) and  $n_{\text{months}} = 8$  implies a favorable detection rate of UV emission from the shock breakout phase of Type II-P SNe by *GALEX* TDS of  $r_{\text{det}} \approx r_{\text{Ia}} n_{\text{months}} f_{\text{det}} \approx 10 \text{ yr}^{-1}$ .

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#### REFERENCES

- Astier, P., et al. 2006, *A&A*, 447, 31  
 Blinnikov, S., et al. 2000, *ApJ*, 532, 1132  
 Blinnikov, S. I., et al. 2002, *Nucl. Astrophys.*, 144  
 Campana, S., et al. 2006, *Nature*, 442, 1008  
 Dahlen, T., et al. 2004, *ApJ*, 613, 189  
 Dessart, L., & Hillier, D. J. 2005a, *A&A*, 437, 667  
 ———. 2005b, *A&A*, 439, 671  
 ———. 2006, *A&A*, 447, 691  
 Dessart, L., et al. 2008, *ApJ*, 675, 644 (D08)  
 Ensmann, L., & Burrows, A. 1992, *ApJ*, 393, 742  
 Falk, S. W. 1978, *ApJ*, 225, L133  
 Freytag, B., Steffen, M., & Dorch, B. 2002, *Astron. Nachr.*, 323, 213  
 Galama, T., et al. 1998, *Nature*, 395, 670  
 Hamuy, M., Suntzeff, N. B., Gonzalez, R., & Martin, G. 1988, *AJ*, 95, 63  
 Hillier, D. J., & Miller, D. L. 1998, *ApJ*, 496, 407  
 Immler, S., et al. 2007, *ApJ*, 664, 435  
 Josselin, E., & Plez, B. 2007, *A&A*, 469, 671  
 Klein, R. I., & Chevalier, R. A. 1978, *ApJ*, 223, L109  
 Levesque, E. M., et al. 2005, *ApJ*, 628, 973  
 Li, L.-X. 2007, *MNRAS*, 375, 240  
 ———. 2008, *MNRAS*, in press (arXiv:0803.0079)  
 Martin, D. C., et al. 2005, *ApJ*, 619, L1  
 Morrissey, P., et al. 2007, *ApJS*, 597, 173, 682  
 Quimby, R. M., et al. 2007, *ApJ*, 666, 1093  
 Richmond, M. W., et al. 1994, *AJ*, 107, 1022  
 Salasnich, B., Bressan, A., & Chiosi, C. 1999, *A&A*, 342, 131  
 Schawinski, K., et al. 2008, *Science*, in press (arXiv:0803.3596)  
 Schmidt, B. P., et al. 1993, *Nature*, 364, 600  
 Soderberg, A. M., et al. 2008, *Nature*, 453, 469  
 Stritzinger, M. et al. 2002, *AJ*, 124, 2100  
 Sullivan, M., et al. 2006, *AJ*, 131, 960  
 Waxman, E., Mészáros, P., & Campana, S. 2007, *ApJ*, 667, 351  
 Weaver, T. A., et al. 1978, *ApJ*, 225, 1021  
 Woosley, S. E., & Bloom, J. S. 2006, *ARA&A*, 44, 507  
 Woosley, S. E., & Heger, A. 2007, *Phys. Rep.*, 442, 269