

# Strong size evolution of the most massive galaxies since $z \sim 2$

Ignacio Trujillo,<sup>1</sup>\*† C. J. Conselice,<sup>1</sup> Kevin Bundy,<sup>2</sup> M. C. Cooper,<sup>3</sup> P. Eisenhardt<sup>4</sup>  
and Richard S. Ellis<sup>5</sup>

<sup>1</sup>*School of Physics and Astronomy, University of Nottingham, University Park, Nottingham NG7 2RD*

<sup>2</sup>*Department of Astronomy and Astrophysics, University of Toronto, 50 St. George Street, Rm 101, Toronto ONM5S 3H4, Canada*

<sup>3</sup>*Department of Astronomy, University of California at Berkeley, Berkeley, CA 94720, USA*

<sup>4</sup>*Jet Propulsion Laboratory, California Institute of Technology, Pasadena, CA 91109, USA*

<sup>5</sup>*Caltech MC 105-24, 1201 East California Boulevard, Pasadena, CA 91125, USA*

Accepted 2007 August 25. Received 2007 July 24; in original form 2007 May 10

## ABSTRACT

Using the combined capabilities of the large near-infrared Palomar/DEEP-2 survey, and the superb resolution of the Advanced Camera for Surveys *HST* camera, we explore the size evolution of 831 very massive galaxies ( $M_{\star} \geq 10^{11} h_{70}^{-2} M_{\odot}$ ) since  $z \sim 2$ . We split our sample according to their light concentration using the Sérsic index  $n$ . At a given stellar mass, both low ( $n < 2.5$ ) and high ( $n > 2.5$ ) concentrated objects were much smaller in the past than their local massive counterparts. This evolution is particularly strong for the highly concentrated (spheroid like) objects. At  $z \sim 1.5$ , massive spheroid-like objects were a factor of 4 ( $\pm 0.4$ ) smaller (i.e. almost two orders of magnitudes denser) than those we see today. These small sized, high-mass galaxies do not exist in the nearby Universe, suggesting that this population merged with other galaxies over several billion years to form the largest galaxies we see today.

**Key words:** galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: formation – galaxies: fundamental parameters – galaxies: high-redshift – galaxies: structure.

## 1 INTRODUCTION

In the nearby Universe, the population of galaxies with stellar masses greater than  $10^{11} M_{\odot}$  is dominated by large early-type galaxies (Baldry et al. 2004) with correspondingly large sizes (Shen et al. 2003). These nearby systems contain old and metal-rich stellar populations that formed quickly in the early Universe (Heavens et al. 2004; Feulner et al. 2005; Thomas et al. 2005). However, it has remained unknown whether the stars in these galaxies were all assembled in the same system or formed in lower mass galaxies that later merged.

Two major formation models have been proposed in order to explain the properties of these galaxies: the so-called monolithic collapse model (Eggen, Lynden-Bell & Sandage 1962; Larson 1975; Arimoto & Yoshii 1987; Bressan, Chiosi & Fagotto 1994) and the hierarchical merging scenario (Toomre 1977; White & Frenk 1991). These two mechanisms have both observational and theoretical successes and drawbacks. For instance, evidence in favour of a fast and dissipative monolithic collapse is the fact that the bulk of stars in massive ellipticals are old (Mannucci et al. 2001) and have high  $[\alpha/\text{Fe}]$  ratios (i.e. short star formation time-scales; Worthey, Faber & Gonzalez 1992). In addition, the structural and dynamical prop-

erties of observed spheroid galaxies are well reproduced by cold dissipationless collapse (van Albada 1982; McGlynn 1984; May & van Albada 1984; Aguilar & Merritt 1990; Londrillo, Messina & Stiavelli 1991; Udry 1993; Hozumi, Burkert & Fujiwara 2000; Trenti, Bertin & van Albada 2005), a process that is expected to dominate the last stages of a highly dissipative collapse. On the other hand, supporting a hierarchical merger scenario, observations find a decline in the number of massive galaxies seen at higher redshifts. This decline is moderate since  $z \sim 1$  and much stronger at even higher redshifts (Daddi, Cimatti & Renzini 2000; Pozzetti et al. 2003; Bell et al. 2004; Drory et al. 2004; Daddi et al. 2005; Drory et al. 2005; Faber et al. 2005; Saracco et al. 2005; Pozzetti et al. 2007). Moreover, new renditions of semi-analytical models where merging is the cornerstone of galaxy formation (e.g. De Lucia et al. 2006) are much better able to match the stellar population properties of elliptical galaxies. Finally, it has been found theoretically that the scaling laws followed by elliptical galaxies are robust against merging (Ciotti, Lanzoni & Volonteri 2007).

Exploring the assembly of massive galaxies with cosmic time through number density analysis or merger rate estimations is, however, difficult to conduct since, first, very massive galaxies are scarce and, secondly, their clustering properties make them strongly affected by field-to-field variance associated with limited observed volumes. These two difficulties make claims of the assembly of massive galaxies in the early Universe based on number densities and merger rate estimations highly uncertain (e.g. Cimatti, Daddi & Renzini 2006; Renzini 2007).

\*E-mail: trujillo@iac.es

†Present address: Instituto de Astrofísica de Canarias, Vía Láctea s/n, 38200 La Laguna, Tenerife, Spain

A much more straightforward approach to the assembly problem of massive galaxies is to explore the size evolution of these systems at a given stellar mass. In a ‘monolithic-like’ scenario where a galaxy is fully assembled after the formation of its stars, the stellar mass–size relation should remain unchanged as cosmic time evolves. In the hierarchical merging scenario, however, the stellar mass–size relation will evolve as a result of the increase in size after each galaxy merger. For instance, state-of-the-art hierarchical semi-analytical models predict a very strong (a factor of 1.5–3) evolution in the size of very massive galaxies ( $M_* \geq 10^{11} h_{70}^{-2} M_\odot$ ) nine billion years ago (Khochfar & Silk 2006b). This predicted evolution in size is a strong function of galaxy mass, with the more massive objects expected to have the largest increase in size.

Due to the lack of a large sample of very massive galaxies at high redshift, analysis of the evolution of the stellar mass–size relation of galaxies has mainly explored objects in the  $10^{10} < M_* < 10^{11} M_\odot$  range. Observations within this mass interval do not find a significant evolution of the stellar mass–size relation since  $z \sim 1$  (Barden et al. 2005; McIntosh et al. 2005). Explorations of these objects at even earlier look-back times ( $1 < z < 3$ ) have been also attempted, and show a moderate decline in size for galaxies at a given stellar mass (Trujillo et al. 2004, 2006a). However, the size of very massive galaxies,  $M_* \geq 10^{11} M_\odot$ , is still largely unexplored, with only around a dozen objects studied in detail so far (Daddi et al. 2005; Trujillo et al. 2006b; Longhetti et al. 2007; McGrath, Stockton & Canalizo 2007a). Definitive conclusions regarding the evolution of sizes for massive galaxies thus remain largely unknown. The goal of this paper is to shed some light on this issue by exploring the sizes of a large sample of very massive galaxies since  $z < 2$ . This will help to clarify whether the evolutionary scenario for these objects is hierarchical or monolithic like.

The paper is structured as follows. In Section 2, we give a brief summary of the Palomar/DEEP-2 data, and in Section 3, we describe the determination of the stellar masses. Size measurement technique and robustness estimations for our data are provided in Section 4. In Section 5, we study the selection effects and in Section 6, we present the observed stellar mass–size relations. We compare our results with other samples in Section 7, and finally, we discuss our results in Section 8. In what follows, we adopt a cosmology of  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$  and  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ .

## 2 DESCRIPTION OF THE DATA

We use the Palomar Observatory Wide-field Infrared (POWIR)/DEEP-2 survey (Davis et al. 2003; Bundy et al. 2006; Conselice et al. 2007a,b) to define a sample of 831 galaxies with masses larger than  $10^{11} h_{70}^{-2} M_\odot$  located over  $\sim 710 \text{ arcmin}^2$  in the Extended Groth Strip (EGS). This field (63 *Hubble Space Telescope* tiles) was imaged with the Advanced Camera for Surveys (ACS) in the *V* band (F606W, 2660 s) and *I* band (F814W, 2100 s). Each tile was observed in four exposures that were combined to produce a pixel scale of 0.03 arcsec, with a point spread function (PSF) of 0.12 arcsec full width half-maximum (FWHM). In addition to the *HST* data, optical imaging from the Canada–France–Hawaii Telescope 3.6-m telescope in the *B*, *R* and *I* bands taken with the CFH12K camera was used. Integration times for these observations were 1 h in *B* and *R*, and 2 h in *I*. Limiting magnitudes reached are  $B = 24.5$  ( $AB, 8\sigma$ ),  $R = 24.2$  ( $AB, 8\sigma$ ) and  $I = 23.5$  ( $AB, 8\sigma$ ). The details of the data reduction for this data are provided in Coil et al. (2004).

In the EGS region, the Palomar near-infrared (near-IR)  $K_s$ -band imaging has a typical depth of greater than  $K_{AB} = 22.5 \text{ mag}$  ( $5\sigma$ ) and a *J*-band depth of  $J_{AB} = 23.4 \text{ mag}$  ( $5\sigma$ ). About 70 per cent

of the galaxies brighter than  $R_{AB} = 24.1 \text{ mag}$  in this field were targeted by the DEEP-2 Galaxy Redshift survey, using the Keck 10-m telescope, for a total spectroscopic redshift completeness of  $\sim 56$  per cent. In our case, 410 objects have a spectroscopic redshift ( $\sim 50$  per cent of the full data set), with a spectroscopic completeness of 61 per cent for galaxies with  $z < 1$ , and a completeness of 30 per cent for galaxies with  $z > 1$ . We supplemented our spectroscopic redshift catalogue with photometric redshifts with an accuracy for all galaxies in the total *K*-band survey of  $\delta z/(1+z) \approx 0.07$  in the range  $0.2 < z < 1.4$  and  $\delta z/(1+z) < 0.22$  for galaxies at  $1.4 < z < 2$  (Conselice et al. 2007a). However, it is worth noting that, for the massive galaxies we use in this paper, the accuracy is  $\delta z/(1+z) = 0.025$  at  $z < 1.4$  and likely similarly lower for  $z > 1.4$  galaxies [i.e.  $\delta z/(1+z) \approx 0.08$ ].

## 3 DETERMINATION OF STELLAR MASSES

The determination of stellar masses for each of our galaxies follows a standard multicolour (BRIJK) stellar population fitting technique, producing uncertainties of  $\approx 0.2$  dex. The largest systematic source of error comes from the assumed initial mass function (IMF), in this paper, we have used the Chabrier IMF for all stellar mass measurements. The details of these estimations are presented in Bundy et al. (2006) and Conselice et al. (2007b). However, because the measurement of stellar masses is an integral part of this paper, we give a brief description of how our masses are measured, and what systematics might be present.

The basic mass determination method we use consists of fitting a grid of model spectral energy distributions constructed from Bruzual & Charlot (2003, hereafter BC03) stellar population synthesis models, with different star formation histories. We use an exponentially declining model to characterize the star formation, with various ages, metallicities and dust contents included. These models are parametrized by an age and an e-folding time. These parametrizations are fairly simple, and it remains possible that stellar mass from older stars is missed under brighter younger populations. While the majority of our systems are passively evolving older stellar populations, it is possible that up to a factor of 2 in stellar mass is missed in any star bursting blue systems. However, stellar masses measured through our technique are roughly the expected factor of 5–10 smaller than dynamical masses at  $z \sim 1$  using a sample of disc galaxies (Conselice et al. 2005), demonstrating their inherent reliability.

We calculate the likely stellar mass, age and absolute magnitudes for each galaxy at all star formation histories, and determine stellar masses based on this distribution. Distributions with larger ranges of stellar masses have larger resulting uncertainties. It turns out that while parameters such as the age, e-folding time, metallicity, etc. are not likely accurately fit through these calculations due to various degeneracies, the stellar mass is robust. Typical errors for our stellar masses are 0.2 dex from the width of the probability distributions. There are also uncertainties from the choice of the IMF. Our stellar masses utilize the Chabrier IMF, which can be converted to Salpeter IMF stellar masses by adding 0.25 dex. There are additional random uncertainties due to photometric errors. The resulting stellar masses thus have a total random error of 0.2–0.3 dex, roughly a factor of 2.

There is furthermore the issue of whether or not our stellar masses are overestimated based on using the BC03 models. It has recently been argued by Maraston (2005) and Bruzual (2007) that the exclusion of an updated treatment of thermal pulsating asymptotic giant branch (TP-AGB) stars in the BC03 models results in calculated stellar masses too high by a factor of a few. While we consider an uncertainty of a factor of 2 in our stellar masses, we must consider

whether our sample is in the regime where the effects of TP–AGB stars will influence our mass measurements. This has been investigated recently in Maraston (2005) and Bruzual (2007) who have both concluded that galaxies stellar masses computed using newer TP–AGB star prescriptions are up to roughly 50–60 per cent lower than without. This is particularly true for masses determined in the rest-frame infrared.

This problem has also been recently investigated independently by Kannappan & Gawiser (2007) who come to similar conclusions, but do not advocate one model over another. Furthermore, the effect of TP–AGB stars is most pronounced in the rest-frame IR, and for young stellar population ages. Our survey is  $K$ -selected, and the observed  $K$  band is used as the flux in which the masses are computed. The rest-frame wavelength probed with the observed  $K$  band ranges from 0.7 to 1.5  $\mu\text{m}$  where the effects of TP–AGB stars are minimized. The ages of our galaxies are also older than the ages where TP–AGB stars have their most effect (Maraston 2005; Bruzual 2007). It is also worth noting that the effects of TP–AGB stars are more important when normalizing stellar masses further into the red. In this paper, we find a strongest evolution at higher redshifts where we are probing the rest-frame optical at observed  $K$ , and where the effects of the TP–AGB stars are minimized (e.g. Bruzual 2007). We investigate the effects of TP–AGB stars in our estimations by determining how our stellar masses change between using the BC03 models and the updated Charlot & Bruzual (in preparation) models using the new TP–AGB methods. We find for 1330 massive galaxies in the Palomar sample a difference of  $\sim 0.07$  in  $\log M_*$ , which is neglectable. A similar conclusion has been recently achieved by McGrath et al. (2007b) using a sample of galaxies at  $z \sim 1.5$  with stellar masses similar to those explored here. On comparing the new rendition of models from Charlot & Bruzual (in preparation) versus BC03, they found that the inferred galaxy masses are slightly ( $\sim 10$  per cent) smaller using the 2007 models.

Although we do have *Spitzer* data for our sources, we do not use this imaging for two reasons. The first is that by normalizing the stellar masses with *Spitzer* IRAC magnitudes, we are in a regime where the TP–AGN stars are more pronounced, and thus would affect the stellar mass measurements to a degree even greater than using the observed  $K$  band (Bruzual 2007; Kannappan & Gawiser 2007). Secondly, the large PSF of the IRAC images makes it difficult to obtain accurate photometry for many of our sources due to contamination from other galaxies with overlapping PSFs. While it is possible to correct for this, the large resulting random photometry uncertainties make the stellar masses less certain.

Another possible source of uncertainty is the photometric redshifts we use for our sample. While at  $z < 1.4$  about half of our sample has spectroscopic redshifts, at  $z > 1.4$  all of our systems have photometric redshifts. We can, however, determine the accuracy for those systems at  $z < 1.4$ . The agreement is very good for our massive systems with  $\delta z / (1 + z) = 0.025$ . This results in another  $< 20$  per cent uncertainty in the stellar mass measurements. For  $z > 1.4$ , if we assume  $\delta z / (1 + z) \approx 0.08$ , the uncertainty in the stellar masses would be  $\sim 32$  per cent. Overall, however, these uncertainties cannot account for the trends seen later in this paper.

#### 4 SIZE ESTIMATION

The structural parameters used in this paper were measured using the ACS  $I$ -band filter. Sizes (as parametrized by the half-light or effective radius along the semimajor axis  $a_e$ ) were estimated using the GALFIT code (Peng et al. 2002). Sizes were circularized,  $r_e = a_e \sqrt{1 - \epsilon}$ , with  $\epsilon$  being the ellipticity of the object.

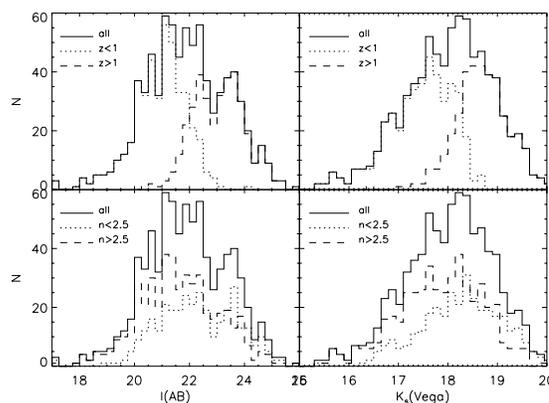
GALFIT convolves Sérsic (1968)  $r^{1/n}$  galaxy models with the PSF of the images, and determines the best fit by comparing the convolved model with the galaxy surface brightness distribution using a Levenberg–Marquardt algorithm to minimise the  $\chi^2$  of the fit. The Sérsic model is a flexible parametric description of the surface brightness distribution of the galaxies and contains the exponential ( $n = 1$ ) and de Vaucouleurs ( $n = 4$ ) models as particular cases. In addition, this model is used in the structural analysis of the SDSS galaxy sample (our local comparison sample; Blanton et al. 2003; Shen et al. 2003).

The Sérsic index  $n$  measures the shape of surface brightness profiles. In the nearby Universe, galaxies with  $n < 2.5$  are mostly disc-like objects, whereas galaxies with  $n > 2.5$  are mainly spheroids (Ravindranath et al. 2002). We use this Sérsic index criterion to split our sample at higher redshifts and facilitate a comparison with the local galaxy population. During the fit, neighbouring galaxies were excluded using a mask, but in the case of closely neighbouring objects with overlapping isophotes, the objects were fitted simultaneously. The results of our fitting are shown in the Appendix (Table A1).

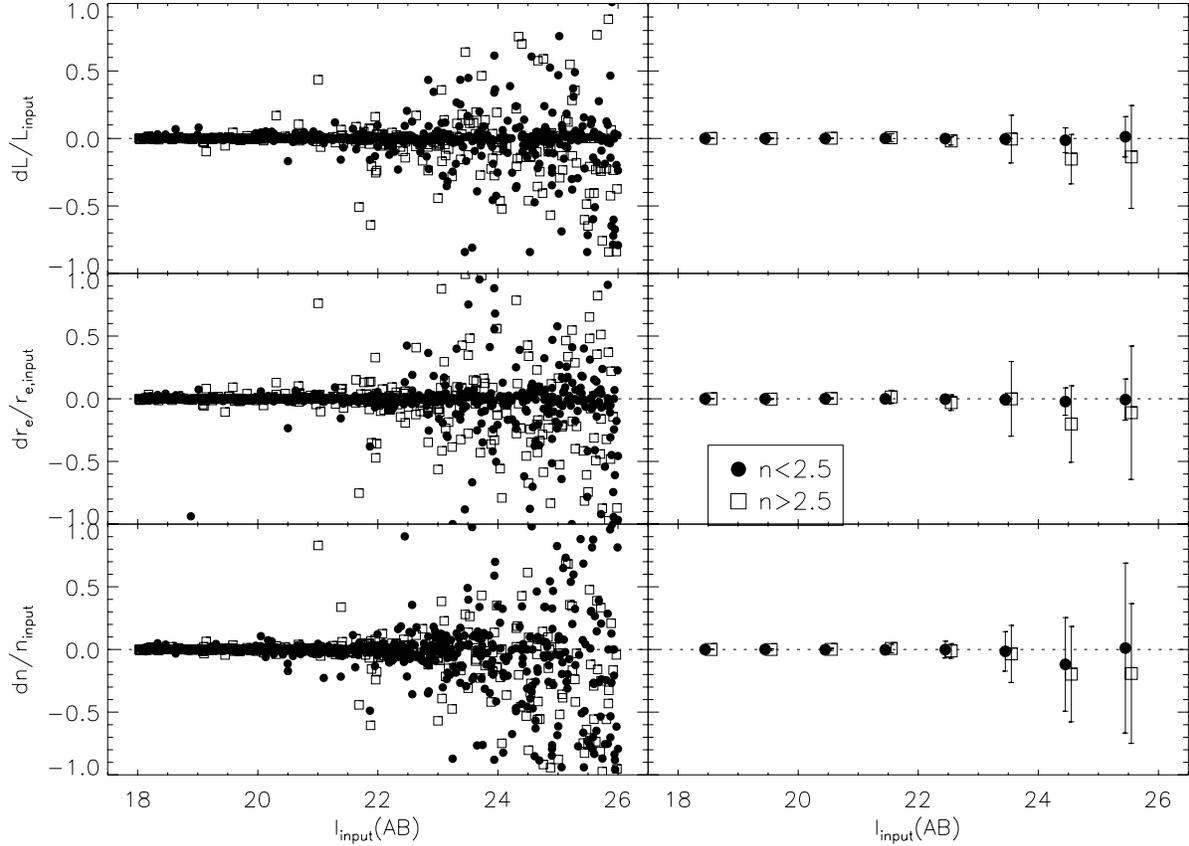
#### 4.1 Testing the structural parameters estimates: simulations

The results presented in this paper rely on our ability to measure accurate structural parameters. To gauge the accuracy of our parameter determination realistic simulations were conducted. We have created 1000 artificial galaxies uniformly generated at random in the following ranges, matching the observed distribution of our galaxies:  $18 \leq I_{AB} \leq 26$  (see Fig. 1),  $0.03 \leq r_e \leq 2.55$  arcsec (which at  $z \sim 1$  equals  $0.25$ – $20.5 h_{70}^{-1}$  kpc),  $0.5 \leq n \leq 8$  and  $0 \leq \epsilon \leq 0.8$ . To simulate the real conditions of our observations, we add a background sky image taken (randomly at each time) from a piece of the ACS  $I$ -band image. Finally, the galaxy models were convolved with the observed PSF. The same procedure was used to retrieve the structural parameters in both the simulated and actual images.

The results of these simulations are shown in Figs 2 and 3. 50 per cent of our galaxies are brighter than  $I(AB) = 21.9$  mag, and for these galaxies, we find  $dr_e/r_e < 5$  per cent and  $dn/n < 7$  per cent. For 95 per cent of our galaxies which are at  $I(AB) < 24.2$  mag, the uncertainties are  $dr_e/r_e < 30$  per cent and  $dn/n < 38$  per cent. As expected, at fainter apparent magnitudes the structural parameters are recovered with larger uncertainties. Only for magnitudes fainter than  $I(AB) = 24$  mag is there a small bias ( $\sim 20$  per cent) of



**Figure 1.** Histograms showing the magnitude distributions in  $I(AB)$  and  $K_s(\text{Vega})$  bands of the galaxies in our sample. The sample is split by redshift and Sérsic index.



**Figure 2.** Relative errors derived from the difference between the input and recovered structural parameters [(output–input)/input] according to our simulations on the *HST* F814W (*I* band) imaging. Filled symbols are used to indicate less concentrated objects ( $n_{\text{input}} < 2.5$ ), whereas open symbols imply highly concentrated objects ( $n_{\text{input}} > 2.5$ ). The right-hand panels show the mean systematic difference and  $1\sigma$  error bars.

the index  $n$  towards smaller indices for galaxies with  $n_{\text{input}} > 2.5$ . Recovery of the structural parameters with larger  $n$  is more affected than those with lower values. As shown in Fig. 3, we do not find any bias of our sizes or shape parameter as a function of the size of the objects. Finally, we have also explored whether the variation of the PSF along the image can affect the recovery of the size of the galaxies. Using different stars in the images as a PSF, we find that the estimation of the sizes is robust to changes in the selected PSF to analyse the data: the scatter between the sizes is  $\lesssim 10$  per cent ( $1\sigma$ ).

#### 4.2 Potential sources of systematic errors: $K$ -correction effects and AGN contamination

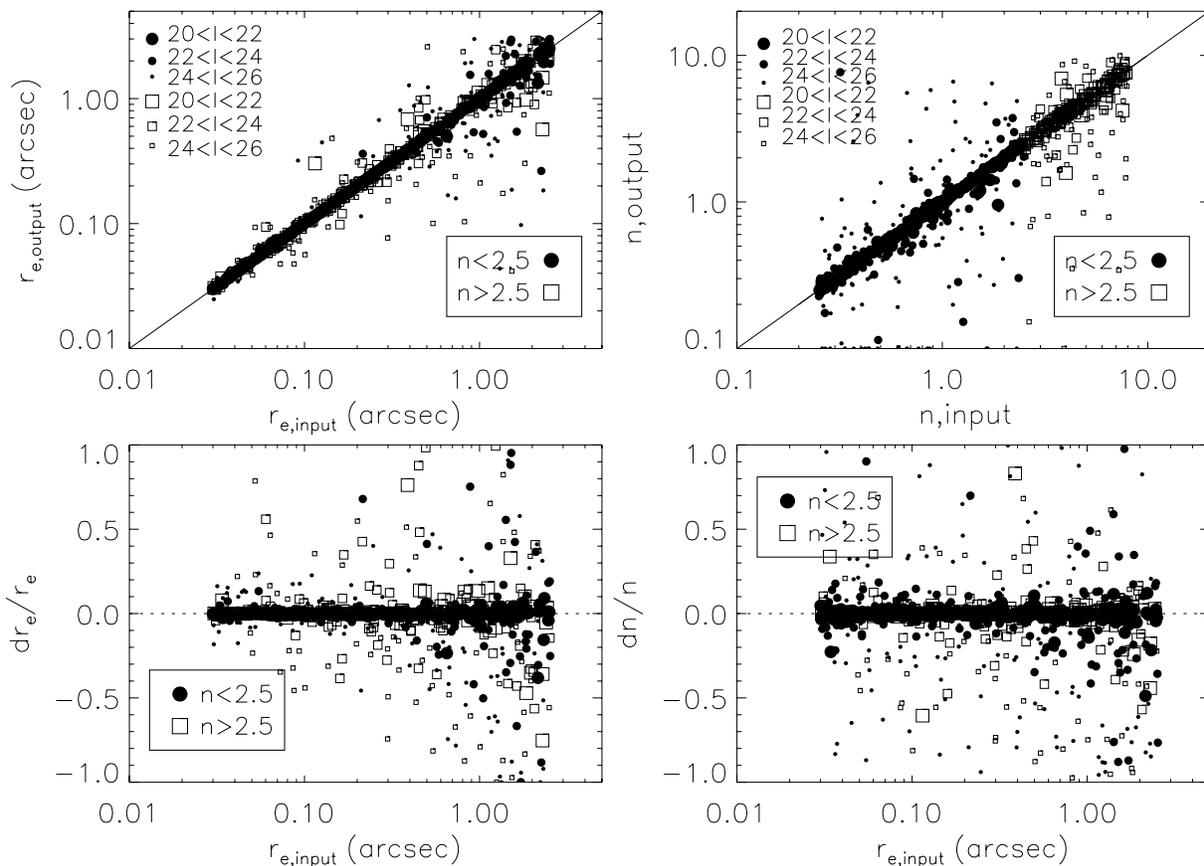
The sizes of the galaxies presented in this paper were measured using the ACS observed *I* band, which implies that for galaxies at  $z > 1.3$  sizes are retrieved in the rest-frame ultraviolet. To check whether this  $K$ -correction effect can affect our size estimates, we have compared our sizes measured in the *I* band with sizes obtained in the Near-Infrared Camera and Multi-Object Spectrometer (NICMOS) *H* band (*F160W*) for a subset of 27 galaxies (with  $0.8 < z < 1.8$  and median  $z \sim 1.2$ ) that were observed at both wavelengths. These NICMOS data consist of 63 pointings of camera 3 ( $52 \times 52$  arcsec,  $0.203$  arcsec pixel $^{-1}$ ) in the EGS field. Each pointing is the combination of four subpixel dithered exposures, with a total exposure time of 2600 s. The final mosaic was assembled using a drizzle task and has a pixel scale of 0.10 arcsec.

We found that sizes measured in the NICMOS images, and those measured with the ACS *I*-band data, are well correlated

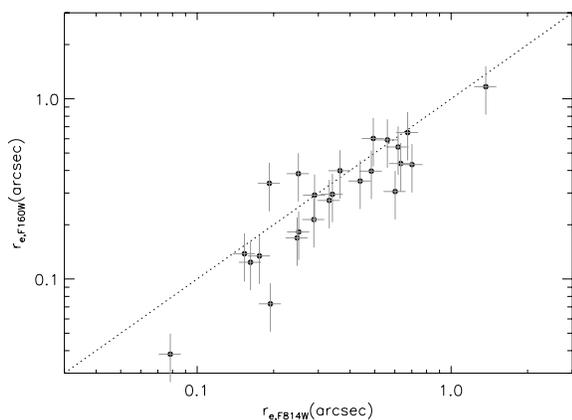
(see Fig. 4). The scatter between both measurements is 32 per cent ( $1\sigma$ ). There seems to be a systematic (although statistically non-significant) bias of  $19 \pm 7$  per cent between both measurements: sizes measured in the redder band (*H* band) are slightly smaller than those measured in the bluer band (*I* band). Consequently, if any  $K$ -correction effect is affecting our results, the size evolution found here using the *I*-band filter would be an upper limit.

Interestingly, this trend of smaller sizes at redder bands  $dr_e(\lambda)/d \log \lambda = -0.6(\pm 0.2)$  is in qualitative agreement to what is found in nearby galaxies  $-0.18 > dr_e(\lambda)/d \log \lambda > -0.25$  (Barden et al. 2005; McIntosh et al. 2005). However, due to the large uncertainty on the bias (and consequently in the size corrections), and because we do not know how this correction could evolve with redshift, we avoid making any  $K$ -correction in our results. In any case, it is important to note that the trend towards smaller sizes observed using the *H* band does not strongly affect the main result of this paper. In fact, the evolution found in this paper would be even stronger using the near-IR band sizes at high redshift.

Another source of concern in the size determination is the presence of an active galactic nuclei (AGN) in the centre of the galaxy which can bias our measurement towards smaller sizes (Daddi et al. 2005). Deep (200 ks per pointing) X-ray observations ( $L_{2-10 \text{ keV}} \geq 10^{42} \text{ erg s}^{-1}$  at  $z \sim 1$ ) from the *Chandra* telescope in this field (Conselice et al. 2007b; Nandra et al. 2007) only detect emission in 35 objects (i.e.  $\sim 4.2$  per cent our sample). We remove these potential AGN from our sample in what follows, but our essential results are unchanged.



**Figure 3.** Galaxy size-measurement bias: the figure shows a comparison between input and recovered structural parameter values in our simulations with the F814W ( $I$  band) ( $AB$ ) imaging. Top left: relation between measured and the input intrinsic half-light radius (before seeing convolution). Top right: relation between measured and input seeing deconvolved Sérsic index  $n$ . Bottom left: relative error between the input and the measured seeing deconvolved effective radius [ $dr_e/r_e = (r_{e,output} - r_{e,input})/r_{e,input}$ ] versus the input effective radius. Bottom right: relative error between the input and the measured seeing deconvolved Sérsic index  $n$  [ $dn/n = (n_{output} - n_{input})/n_{input}$ ] versus the input effective radius. Filled symbols are used to indicate less concentrated objects ( $n_{input} < 2.5$ ), whereas open symbols imply highly concentrated objects ( $n_{input} > 2.5$ ).



**Figure 4.** Comparison between the size estimates using the ACS  $I$ -band filter versus the NICMOS  $H$  band for 27 galaxies in our sample at  $0.8 < z < 1.8$  where NICMOS data are available.

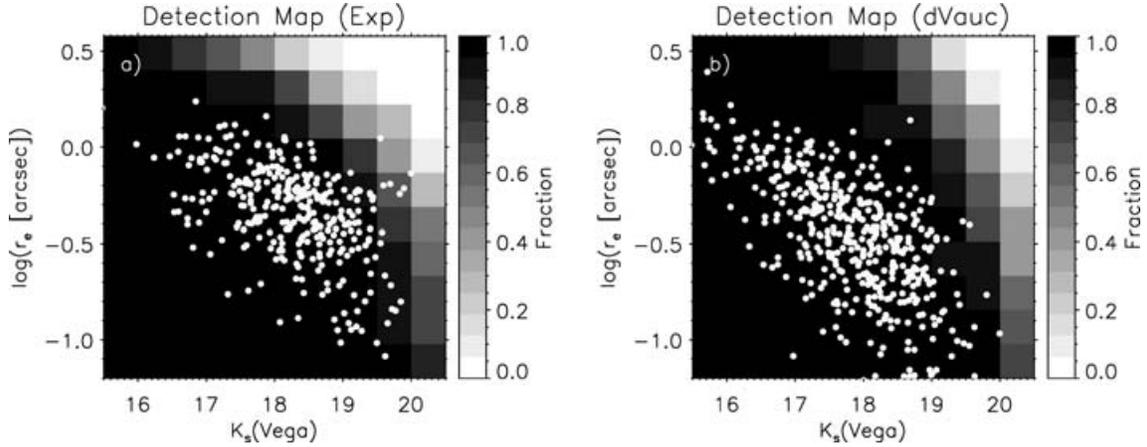
## 5 SELECTION EFFECTS

In practice, any galaxy survey has a surface brightness limit beyond which the sample is incomplete. Characterizing this limit is particularly important for high- $z$  samples, where the effects of the

cosmological surface brightness dimming are severe. At a given total flux limit, the surface brightness limit translates into an upper limit on the size in which a galaxy can be detected. We have explored whether our  $K_s$  band selected sample could be incomplete at large sizes by examining the detectability of our galaxies as a function of their apparent magnitudes and sizes.

To determine the detection map of the Palomar  $K_s$ -band image, we have created two sets of  $10^4$  mock sources each with intrinsic exponential or de Vaucouleurs profiles uniformly distributed as follows:  $K_s$  band total magnitudes between 15.5 and 20.5 (Vega) mag (see Fig. 1), effective radius  $r_e$  between 0.0625 and 3.75 arcsec (this will be equivalently to  $0.5\text{--}30 h_{70}^{-1}$  kpc at  $z \sim 1$ ) and ellipticities between 0 and 0.8. The simulated sources are placed randomly on the real image and extracted as for the real source detection. We construct from these simulations detection maps giving the number of recovered sources over the number of input artificial sources per input magnitude and input  $\log r_e$  bin (see Fig. 5). At it is expected, galaxies with a de Vaucouleurs profile, and consequently more centrally concentrated, are easier to detect at a given magnitude.

Over-plotted on the detection maps are the distribution of our sample galaxies.  $K_s$ -band magnitudes are measured using the  $MAG_{AUTO}$  output from *SEXTRACTOR*. The sizes are those estimated in the  $I$  band. We are making the implicit assumption that sizes are similar in both

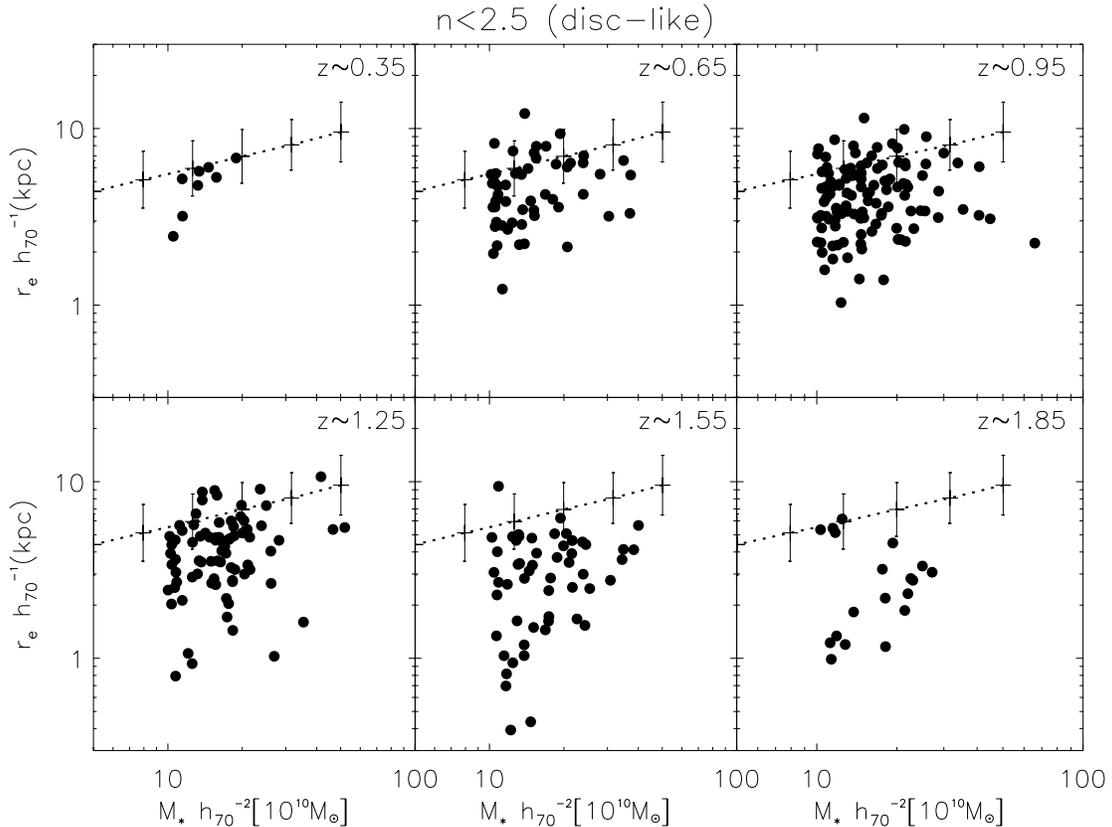


**Figure 5.** (a) Detection map for simulated sources with exponential ( $n = 1$ ) profiles placed at random in our  $K_s$ -band image of the Palomar field. The grey-scale map reflects the ratio between input and recovered objects per input magnitude and  $\log r_e$  bin. Over-plotted on the map is the distribution of the sample of  $K_s$  band selected objects in the Palomar field with  $n < 2.5$  (as measured in the  $I$  band). Sizes of real objects are those derived in the  $I$ -band image. (b) Same as (a) but for simulated sources with de Vaucouleurs ( $n = 4$ ) profiles placed at random in our  $K_s$ -band image of the Palomar field. Over-plotted on the map is the distribution of the sample of  $K_s$  band selected objects in the Palomar field with  $n > 2.5$  (as measured in the  $I$  band).

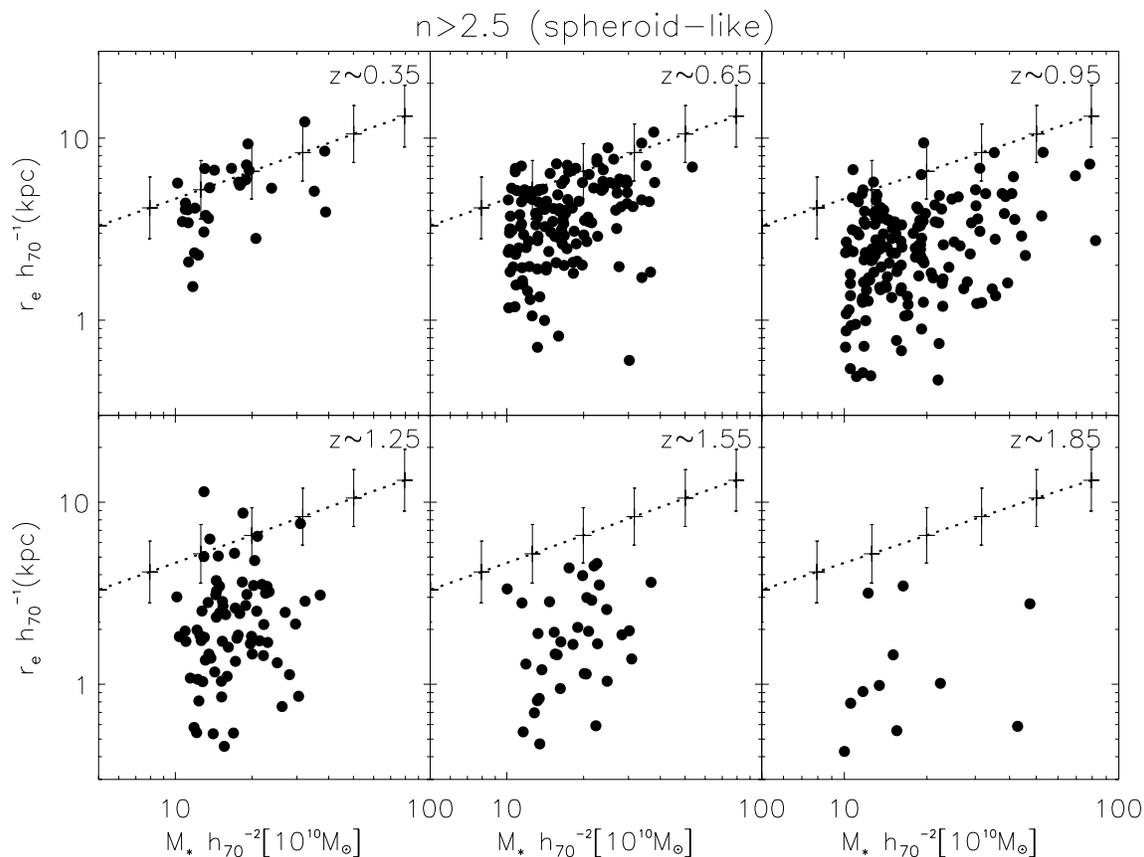
bands (i.e. that  $K$ -correction effects are not relevant, see previous section). At a given magnitude, the observed size distribution declines more rapidly to larger sizes than the detection limit. This indicates that our sample is not affected by incompleteness for the largest galaxies at a given magnitude.

## 6 THE OBSERVED STELLAR MASS–SIZE RELATION

The accuracy of the structural parameters and the completeness of our sample has been demonstrated in the previous section. In this



**Figure 6.** Stellar mass size distribution for our less concentrated (disc like) galaxies. Over-plotted on the observed distribution of points are the mean and dispersion of the distribution of the Sérsic half-light radius of the SDSS late-type ( $n < 2.5$ ) galaxies as a function of the stellar mass. We use the SDSS sample as the local reference ( $z \sim 0.1$ ). SDSS sizes were determined also using a circularized Sérsic model and masses were retrieved using a Kroupa IMF. SDSS sizes were measured using the observed  $r'$  band, closely matching the  $V$ -band rest-frame filter at  $z \sim 0.1$ . For clarity, individual error bars are not shown. The mean size relative error is  $< 11$  per cent. Uncertainties in the stellar masses are  $\sim 0.2$  dex.



**Figure 7.** Stellar mass size distribution of our high-concentrated (spheroid like) galaxies. Over-plotted on the observed distribution of points are the mean and dispersion of the distribution of the Sérsic half-light radius of the SDSS early-type ( $n > 2.5$ ) galaxies as a function of the stellar mass. For clarity, individual error bars are not shown. The mean size relative error is  $< 30$  per cent. Uncertainties in the stellar mass are  $\sim 0.2$  dex.

section, we discuss the stellar mass–size relations for our massive galaxy sample. These relations are presented in Figs 6 ( $n < 2.5$ ; disc-like objects) and 7 ( $n > 2.5$ ; spheroid-like objects). In each of these figures, we have divided our sample (over  $0.2 < z < 2$ ) into six redshift slices with a width of  $dz = 0.3$ .

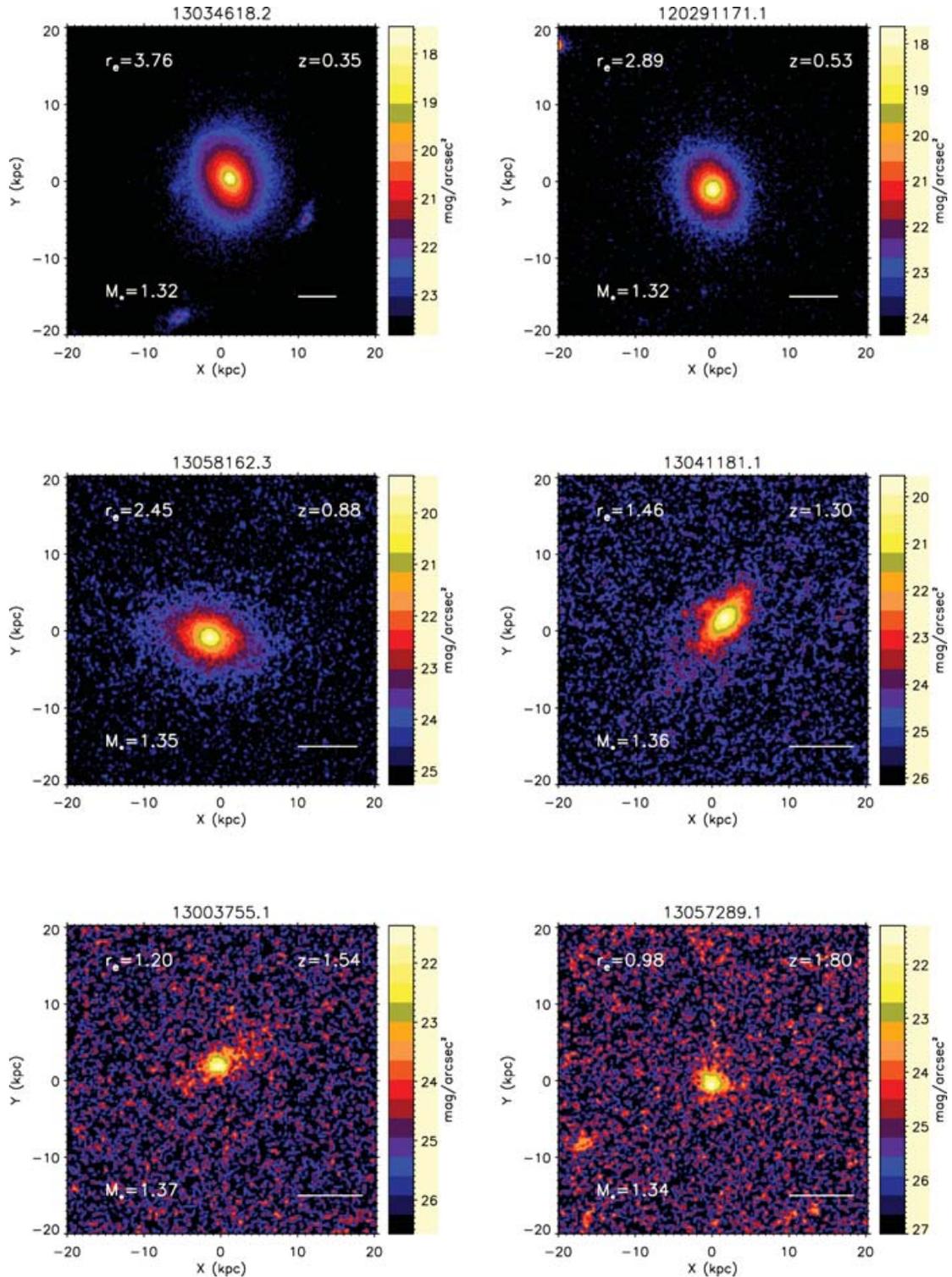
Over-plotted on our observed distributions are the mean and dispersion of the distribution of the Sérsic half-light radii from the SDSS (York et al. 2000) galaxies. We use the SDSS sample as the local reference. As it has been done for our high- $z$  galaxies, the SDSS galaxy’ sizes were determined from a Sérsic model fit (Blanton et al. 2003). The characteristics of the SDSS sample used here are detailed in Shen et al. (2003). The mean of the SDSS galaxies’ redshift distribution used for comparison is 0.1. We use the sizes and the shapes estimated in the observed SDSS  $r$  band as this closely matches the V-band rest-frame filter at  $z \sim 0.1$ . SDSS stellar masses were derived using a Kroupa (2001) IMF, which gives the same masses as the Chabrier IMF used to measure the higher redshift sample stellar masses.

Figs 6 and 7 show that at a given stellar mass galaxies are progressively smaller at higher redshift. This evolution is particularly strong for more concentrated ( $n > 2.5$ ) galaxies. We find the remarkable result that there are no spheroid-like objects at  $z > 1.5$  on the local relation.

To illustrate this size decrease at progressively higher  $z$ , Fig. 8 shows the appearance of six of our  $n > 2.5$  (i.e. spheroid like) galaxies with the same stellar mass, but at different look-back times. The objects shown in Fig. 8 have the mean structural properties of

the galaxy population at each of the different redshift slices (see the stellar mass–size relations in the above redshift intervals in Figs 6 and 7). To allow a fair comparison between the sizes of objects at very different redshifts, the limiting surface brightness in each panel is changed according to the cosmological surface bright dimming  $\sim (1+z)^4$ . For this reason, the surface brightness of the object at  $z = 0.35$  is shown down to  $\sim 3.25$  mag brighter than the object at  $z = 1.85$ . This figure visually illustrates that the most massive galaxies are progressively smaller at increasing redshift.

To quantify the size evolution, we measure the ratio between the observed size and the expected size at a given stellar mass by comparing with the SDSS (Shen et al. 2003) distribution at different redshifts. To estimate the expected size from SDSS at a given stellar mass, we interpolate linearly between the SDSS points when necessary. The evolution of the median and the dispersion of the above ratio are shown in Fig. 9 and listed in Table 1. The observed size evolution of  $M_* > 10^{11} M_\odot$  galaxies is stronger than the one found in previous work using less massive galaxies (see the detailed comparison of this issue in the next section). Following recent claims (i.e. Maraston et al. 2006) that stellar masses could be systematically overestimated by a factor of few at high  $z$ , we have repeated our analysis under the assumption that our masses could be overestimated by a factor of 2. In this case, at  $z \sim 1.5$ , our galaxies will still be more compact than present-day galaxies of the same stellar mass by a factor of 1.6 for  $n < 2.5$  and a factor of 3 for  $n > 2.5$ . In other words, even a systematic effect of factor of 2 in the stellar mass determination cannot avoid a significant evolution in the size

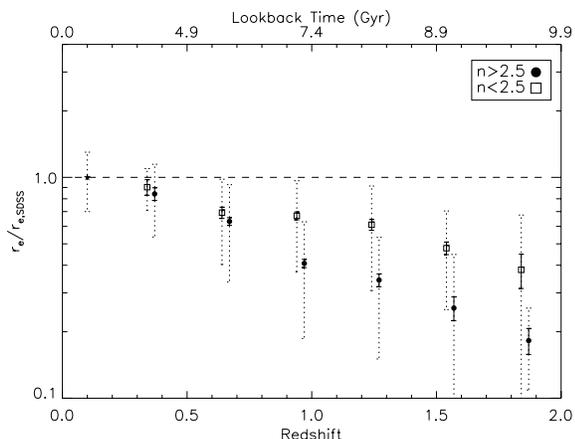


**Figure 8.** Visual structural evolution of equal stellar mass galaxies at different look-back times. The panel shows six concentrated ( $n > 2.5$ , spheroid like) representative galaxies in our sample at different redshift (or look-back time, from top-left to bottom-right: 3.9, 5.3, 7.4, 9.0, 9.7 and 10.3 Gyr back, respectively). Effective radii are given in kpc and stellar masses in  $10^{11} M_\odot$  units. Galaxies are shown with different surface brightness limits to account for the cosmological surface brightness dimming. The solid line indicates 1 arcsec angular size.

of these galaxies. In fact, to prevent a significant evolution of the sizes of galaxies at a given stellar mass, we found that, at  $z \sim 1.5$ , our stellar masses would need to be overestimated by a factor of  $\sim 10$  for galaxies with  $n < 2.5$  and overestimated by a factor of  $\sim 50$  for galaxies with  $n > 2.5$ .

## 7 COMPARISON WITH OTHER SAMPLES

An interesting point to explore is how the size evolution found here for the most massive galaxies compares with the size evolution observed for galaxies with lower masses. This comparison can be done



**Figure 9.** Size evolution of the most massive galaxies with look-back time. The evolution with redshift of the median ratio between the sizes of the galaxies in our sample and the galaxies of the same stellar mass in the SDSS local comparison sample is shown. Solid points indicate the size evolution of spheroid-like ( $n > 2.5$ ) galaxies. Open squares show the evolution for disc-like ( $n < 2.5$ ) galaxies. The small error bars indicate the uncertainty ( $1\sigma$ ) at the median position. The large error bars represent the dispersion in the distributions.

in detail in the redshift range  $0 < z < 1$  using the data from one of the largest sample currently available: the GEMS survey (Rix et al. 2004). The stellar mass–size relation of this survey has been already derived for late-type (Barden et al. 2005) and early-type (McIntosh et al. 2005) galaxies. The GEMS late- and early-type separation criteria are based on the Sérsic index  $n$ . Late types are defined through  $n < 2.5$ , and early types through  $n > 2.5$  and a colour within the ‘red sequence’ (Bell et al. 2004). The comparison of the distribution of their data points with ours is shown in Fig. 10.

It is encouraging to see that these two independent analyses and data sets match well where there is overlap in their stellar mass (i.e. for the small subset of GEMS galaxies with  $M_* > 10^{11} M_\odot$ ). These two data sets allow us to compare the difference in size evolution in the  $0 < z < 1$  redshift interval at two different mass ranges  $10^{10} < M_* < 10^{11} M_\odot$  (GEMS) and  $M_* > 10^{11} M_\odot$  (Palomar). This is shown in Fig. 11. From this comparison, we can see that more massive galaxies evolve in size much faster than lower mass objects (particularly for disc-like galaxies). This mass dependent evolution was hinted in previous works (Trujillo et al. 2006a) but our current large data set shows this more clearly and robustly.

At higher redshift, the amount of data is more limited. At  $1 < z < 2$ , our results are in good agreement with recent findings (based on a few objects) of massive compact galaxies at high  $z$  (Waddington et al. 2002; Daddi et al. 2005; di Serego Alighieri et al. 2005; Trujillo et al. 2006b; Longhetti et al. 2007). For example, using the 10 of

the most massive galaxies ( $1.2 < z < 1.7$ ) in the MUNICS survey, Trujillo et al. (2006b) found that these galaxies were a factor of  $4^{+1.9}_{-1.0}$  smaller than local counterparts. At even higher  $z$  (i.e.  $z \sim 2.5$ ), there has been also recent claims of very compact ( $r_e \lesssim 1$  kpc) massive galaxies (Trujillo et al. 2006a; Zirm et al. 2007; Toft et al. 2007).

## 8 DISCUSSION

As shown in Section 4, the compact nature of our most massive galaxies at high  $z$  cannot be interpreted as a  $K$ -correction or AGN effect. In addition, there is no observational evidence in the local Universe for galaxies as massive and compact as the ones in our sample. These two observational facts raise the following two questions: first, how can these objects be so dense in the past? Secondly, how do these objects evolve in stellar mass and/or size in order to reach the current local relation?

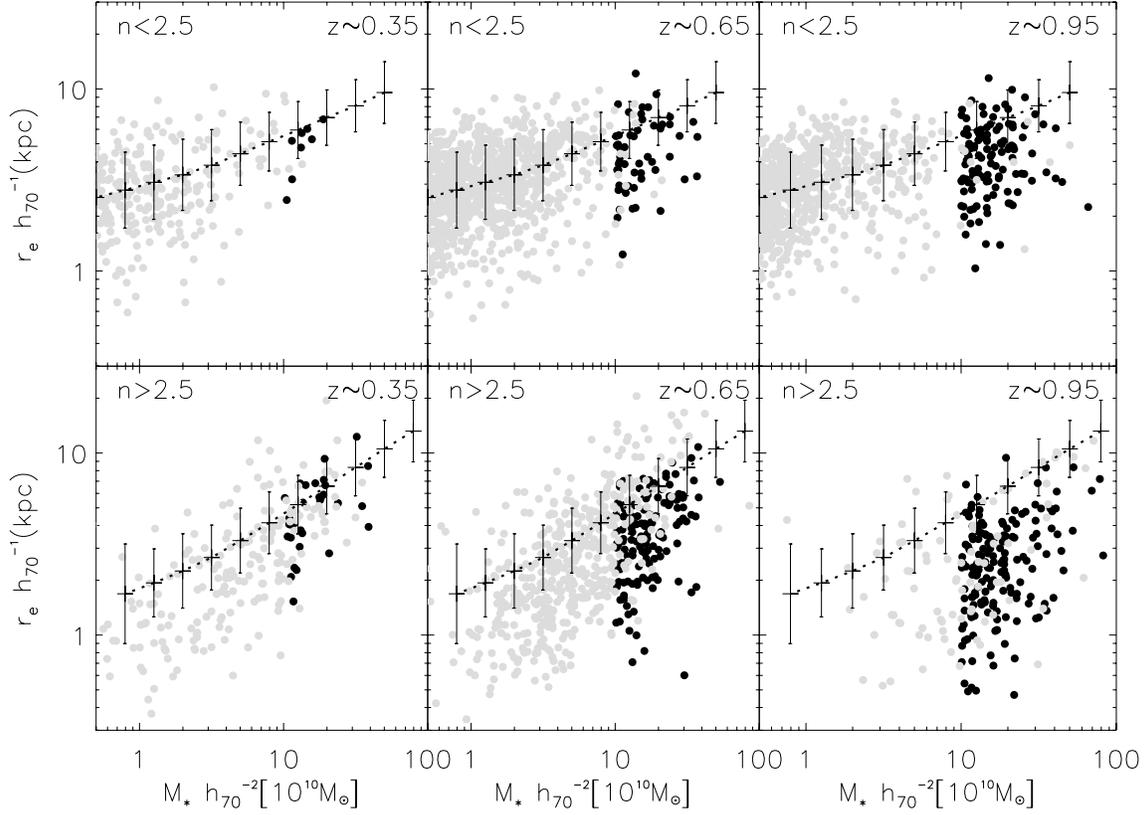
Addressing the first question, recent theoretical results suggest that major galaxy mergers in the early Universe had a much larger component of cold gas available than in the present (Khochfar & Silk 2006a). These wet (dissipative) mergers generate very efficient and massive starbursts creating a very compact massive remnant. Consequently, the very dense nature of our objects at high  $z$  could reflect the much denser condition of the Universe at the time of their formation.

Focusing on the second question, at lower redshifts the available amount of gas is less, and new ‘dry’ (dissipationless) mergers (van Dokkum 2005; Bell et al. 2006) would be the dominant mechanisms of size and stellar mass growing (Ciotti & van Albada 2001; Khochfar & Burkert 2003; Nipoti, Londrillo & Ciotti 2003; Domínguez-Tenreiro et al. 2006; Naab, Khochfar & Burkert 2006; Ciotti et al. 2007). Dry mergers are not efficient at forming new stars, but are efficient in increasing the size of the objects. A particular effective size evolutionary mechanism ( $r_e \sim M_*^{1.3}$ ) has been recently found (Boylan-Kolchin, Ma & Quataert 2006) in mergers of galaxies with radial orbits along large-scale structure filaments. This kind of mechanism would be able to evolve our compact galaxies (a factor of 4 smaller at  $z \sim 1.5$ ) to the local relation with just two major (equal mass) mergers. Since  $z \sim 2$  very few gas rich mergers occur in massive galaxies (e.g. Conselice et al. 2003; Conselice 2006), but a few dry mergers are possible based on pair counts (e.g. Lin et al. 2004). In addition, dry mergers of similar mass objects (and consequently, potentially similar ages and metallicities) will also help to understand the age-uniformity found in the local massive (spheroid) galaxies. On the other hand, in agreement with our results here, new semi-analytical models (Somerville et al. 2007) find that disc-like galaxies (i.e. those presumably with a relative quiet live) evolve only mildly in size since high  $z$ .

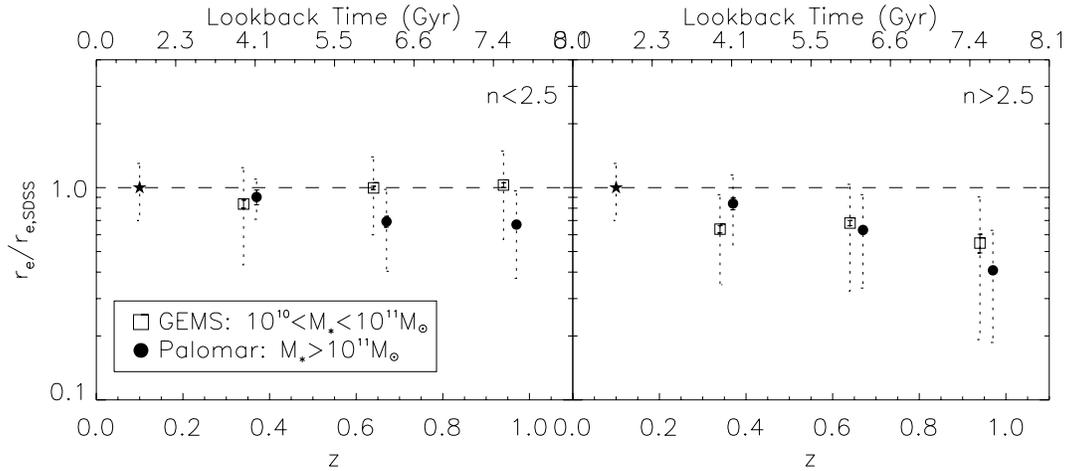
Summarizing our results, we find that the size evolution of the most massive galaxies is not consistent with a scenario whereby

**Table 1.** Size evolution of the most massive ( $M_* > 10^{11} M_\odot$ ) galaxies in the Universe.

Redshift range	$n < 2.5$	$n < 2.5$	$n > 2.5$	$n > 2.5$
	$\langle r_e / r_{e, \text{SDSS}} \rangle (\pm 1\sigma)$	dispersion	$\langle r_e / r_{e, \text{SDSS}} \rangle (\pm 1\sigma)$	dispersion
0.1 (SDSS)	1	0.30	1	0.30
0.2–0.5	0.90 (0.07)	0.19	0.84 (0.06)	0.30
0.5–0.8	0.69 (0.04)	0.29	0.63 (0.03)	0.30
0.8–1.1	0.67 (0.03)	0.29	0.41 (0.02)	0.22
1.1–1.4	0.61 (0.03)	0.30	0.34 (0.02)	0.19
1.4–1.7	0.48 (0.03)	0.23	0.26 (0.03)	0.19
1.7–2.0	0.38 (0.07)	0.30	0.18 (0.02)	0.07



**Figure 10.** Stellar mass size evolution with redshift of disc-like ( $n < 2.5$ ) and spheroid-like ( $n > 2.5$ ) galaxies up to  $z \sim 1$ . The Palomar massive galaxies sample is shown (solid points) in comparison with the lower stellar masses galaxies from the GEMS survey (grey symbols). Over-plotted are the stellar mass–size relation from the SDSS (Shen et al. 2003).



**Figure 11.** Redshift evolution of the ratio between the observed size and the present-day mean size at a given stellar mass. Open squares represent those galaxies within the mass ranges  $10^{10} < M_* < 10^{11} M_{\odot}$  (GEMS), whereas solid points are  $M_* > 10^{11} M_{\odot}$  (Palomar) galaxies. Meaning of bars as in Figure 9.

massive galaxies were fully assembled in the early Universe, and have subsequently evolved passively until today (i.e. a pure ‘monolithic’ scenario). In fact, our findings agree with a scenario where a fraction of the most massive galaxies possibly formed in a short ‘monolithic-like’ collapse and then evolved through major gas-rich or gas-poor merging. In this sense, the two scenarios would be just different phases of galaxy formation and evolution. Our results, consequently, point to a scenario whereby the stellar populations of the

most massive galaxies, we observe today were located in different ‘primordial’ massive galaxy pieces in the early Universe.

#### ACKNOWLEDGMENTS

Authors are grateful to M. Barden and D. McIntosh for providing us with the GEMS data points to allow us a comparison of our results with their less massive galaxies. We acknowledge useful discussions

with O. Almaini and A. Aragón-Salamanca. We also thank the useful suggestions from an anonymous referee. The Palomar and DEEP-2 surveys would not have been completed without the active help of the staff at the Palomar and Keck observatories. The Palomar Survey was supported by NSF grant AST-0307859 and NASA STScI grant HST-AR-09920.01.A. Support for the ACS imaging of the EGS in GO program 10134 was provided by NASA through NASA grant HST-GO-10134.13-A from the Space Telescope Science Institute.

## REFERENCES

- Aguilar L. A., Merritt D., 1990, *ApJ*, 354, 33  
 Arimoto N., Yoshii Y., 1987, *A&A*, 173, 23  
 Baldry I. K., Glazebrook K., Brinkmann J., Ivezić Ž., Lupton R. H., Nichol R. C., Szalay A. S., 2004, *ApJ*, 600, 681  
 Barden M. et al., 2005, *ApJ*, 635, 959  
 Bell E. F. et al., 2004, *ApJ*, 608, 752  
 Bell E. F. et al., 2006, *ApJ*, 640, 241  
 Blanton M. R. et al., 2003, *ApJ*, 592, 819  
 Boylan-Kolchin M., Ma C.-P., Quataert E., 2006, *MNRAS*, 369, 1081  
 Bressan A., Chiosi C., Fagotto F., 1994, *ApJS*, 94, 63  
 Bruzual G., 2007, preprint (astro-ph/0703052)  
 Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000 (BC03)  
 Bundy K. et al., 2006, *ApJ*, 651, 120  
 Cimatti A., Daddi E., Renzini A., 2006, *A&A*, 453, L29  
 Ciotti L., van Albada T. S., 2001, *ApJ*, 552, L13  
 Ciotti L., Lanzoni B., Volonteri M., 2007, *ApJ*, 658, 65  
 Coil A. L., Newman J. A., Kaiser N., Davis M., Ma C.-P., Kocevski D. D., Koo D. C., 2004, *ApJ*, 617, 765  
 Conselice C. J., 2006, *ApJ*, 639, 120  
 Conselice C. J., Bershadsky M. A., Dickinson M., Papovich C., 2003, *AJ*, 126, 1183  
 Conselice C. J., Bundy K., Ellis R., Brinchmann J., Vogt N., Phillips A., 2005, *ApJ*, 628, 160  
 Conselice C. J. et al., 2007a, *ApJ*, 660, L55  
 Conselice C. J. et al., 2007b, *MNRAS*, in press (astro-ph/0708.1040)  
 Daddi E., Cimatti A., Renzini A., 2000, *A&A*, 362, L45  
 Daddi E. et al., 2005, *ApJ*, 626, 680  
 Davis M. et al., 2003, *SPIE*, Vol. 4834, p. 161  
 De Lucia G., Springel V., White S. D. M., Croton D., Kauffmann G., 2006, *MNRAS*, 366, 499  
 di Serego Alighieri S. et al., 2005, *A&A*, 442, 125  
 Domínguez-Tenreiro R., Oñorbe J., Sáiz A., Artal H., Serna A., 2006, *ApJ*, 636, L77  
 Drory N., Bender R., Feulner G., Hopp U., Maraston C., Snigula J., Hill G. J., 2004, *ApJ*, 608, 742  
 Drory N., Salvato M., Gabasch A., Bender R., Hopp U., Feulner G., Pannella M., 2005, *ApJ*, 619, L131  
 Eggen O. J., Lynden-Bell D., Sandage A. R., 1962, *ApJ*, 136, 748  
 Faber S. M. et al., 2007, *ApJ*, 665, 265  
 Feulner G., Gabasch A., Salvato M., Drory N., Hopp U., Bender R., 2005, *ApJ*, 633, L9  
 Heavens A., Panter B., Jimenez R., Dunlop J., 2004, *Nat*, 428, 625  
 Hozumi S., Burkert A., Fujiwara T., 2000, *MNRAS*, 311, 377  
 Kannappan S. J., Gawiser E., 2007, *ApJ*, 657, L5  
 Khochfar S., Burkert A., 2003, *ApJ*, 597, L117  
 Khochfar S., Silk J., 2006a, *MNRAS*, 370, 902  
 Khochfar S., Silk J., 2006b, *ApJ*, 648, L21  
 Kroupa P., 2001, *MNRAS*, 322, 231  
 Larson R., 1975, *MNRAS*, 173, 671  
 Lin L. et al., 2004, *ApJ*, 617, L9  
 Londrillo P., Messina A., Stiavelli M., 1991, *MNRAS*, 250, 54  
 Longhetti M. et al., 2007, *MNRAS*, 374, 614  
 McGlynn T. A., 1984, *ApJ*, 281, 13  
 McGrath E. J., Stockton A., Canalizo G., 2007a, *ApJ*, in press (astro-ph/0707.1046)  
 McGrath E. J., Stockton A., Canalizo G., Masanori I., Maihara T., 2007b, *ApJ*, submitted (astro-ph/0707.1050)  
 McIntosh D. et al., 2005, *ApJ*, 632, 191  
 Mannucci F., Basile F., Poggianti B. M., Cimatti A., Daddi E., Pozzetti L., Vanzi L., 2001, *MNRAS*, 326, 745  
 Maraston C., 2005, *MNRAS*, 362, 799  
 Maraston C., Daddi E., Renzini A., Cimatti A., Dickinson M., Papovich C., Pasquali A., Pirzkal N., 2006, *ApJ*, 652, 85  
 May A., van Albada T. S., 1984, *MNRAS*, 209, 15  
 Naab T., Khochfar S., Burkert A., 2006, *ApJ*, 636, L81  
 Nandra K. et al., 2007, *ApJ*, 660, L11  
 Nipoti C., Londrillo P., Ciotti L., 2003, *MNRAS*, 342, 501  
 Peng C. Y., Ho L. C., Impey C. D., Rix H. W., 2002, *AJ*, 124, 266  
 Pozzetti L. et al., 2003, *A&A*, 402, 387  
 Pozzetti L. et al., 2007, *A&A*, preprint (astro-ph/0704.1600)  
 Ravindranath S. et al., 2002, *ApJ*, 604, L9  
 Renzini A., 2007, preprint (astro-ph/0702148)  
 Rix H.-W. et al., 2004, *ApJ*, 152, 163  
 Saracco P. et al., 2005, *MNRAS*, 357, L40  
 Sérsic J.-L., 1968, *Atlas de Galaxias Australes*. Observatorio Astronomico, Cordoba  
 Shen S., Mo H. J., White S. D. M., Blanton M. R., Kauffmann G., Voges W., Brinkmann J., Csabai I., 2003, *MNRAS*, 343, 978  
 Somerville R. et al., 2007, *ApJ*, in press (astro-ph/0612428)  
 Thomas D., Maraston C., Bender R., de Oliveira C. M., 2005, *ApJ*, 621, 673  
 Toft S. et al., 2007, submitted (astro-ph/07.07.4484)  
 Toomre A., 1977, in Tinsley B. M., Larson R. B., eds, *Evolution of Galaxies and Stellar Populations*. Yale Univ. Obs., New Haven, p. 401  
 Trenti M., Bertin G., van Albada T. S., 2005, *A&A*, 433, 57  
 Trujillo I. et al., 2004, *ApJ*, 604, 521  
 Trujillo I. et al., 2006a, *ApJ*, 650, 18  
 Trujillo I. et al., 2006b, *MNRAS*, 373, L36  
 Udry S., 1993, *A&A*, 268, 35  
 van Albada T. S., 1982, *MNRAS*, 201, 939  
 van Dokkum P. G., 2005, *AJ*, 130, 2647  
 Waddington I. et al., 2002, *MNRAS*, 336, 1342  
 White S. D. M., Frenk C. S., 1991, *ApJ*, 379, 52  
 Worthey G., Faber S. M., Gonzalez J. J., 1992, *ApJ*, 398, 69  
 York D. et al., 2000, *AJ*, 120, 1579  
 Zirm A. W. et al., 2007, *ApJ*, 656, 66

## APPENDIX A: STRUCTURAL PARAMETERS OF THE PALOMAR/DEEP-2 SAMPLE MASSIVE GALAXIES

Table A1 contains a sample of the information about the structural parameters of the sample of galaxies analysed in this paper. Column 1 indicates the galaxy identification, Column 2 lists the apparent  $K$ -band magnitude in the Vega system, Column 3 is the effective radius along the semimajor axis, Column 4 indicates the value of the Sérsic index of the fit, Column 5 is the ellipticity of the source, Column 6 is the stellar mass of the galaxy in units of  $10^{10} h_{70}^{-2} M_{\odot}$ , Column 7 is the measured redshift of the object, Columns 8 and 9 are the RA and Dec. of the source and, finally, Column 10 specifies whether the redshift was determined spectroscopically (1) or photometrically (0). The full version is available online (see the Supplementary Material section).

## SUPPLEMENTARY MATERIAL

The following supplementary material is available for this article:

**Table A1.** Properties of the Palomar/DEEP-2 sample galaxies.

Galaxy ID	$K_s$ (Vega mag)	$a_e$ (arcsec)	$n$	$\epsilon$	$M_*$ ( $10^{10} h_{70}^{-2} M_{\odot}$ )	$z$	RA (J2000)	Dec. (J2000)	Spectroscopic
11051619	16.75	0.90	3.60	0.38	11.94	0.46	214.258 865	52.399 944	0
12004106	17.13	0.80	6.72	0.11	26.72	0.74	214.379 730	52.443 295	1
12004308	16.95	0.89	1.23	0.58	18.03	0.64	214.266 815	52.412 926	0
12004443	17.06	0.48	5.04	0.21	26.96	0.77	214.318 298	52.429 092	1
12004485	17.59	0.49	2.27	0.60	13.20	0.68	214.317 200	52.434 250	1
12004493	17.62	0.26	2.19	0.54	11.28	0.63	214.249 069	52.427 074	1
12004496	16.52	0.96	3.83	0.17	23.86	0.49	214.313 721	52.425 831	0
12004511	17.31	0.52	4.30	0.15	14.87	0.55	214.274 643	52.413 155	0
12007872	17.70	1.21	0.35	0.63	13.47	0.80	214.427 551	52.479 439	0

**Table A1.** Properties of the Palomar/DEEP-2 sample galaxies.

This material is available as part of the online article from: <http://www.blackwell-synergy.com/doi/abs/10.1111/j.1365-2966.2007.12388.x>

(this link will take you to the article abstract).

Please note. Blackwell Publishing are not responsible for the content or functionality of any supplementary materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a  $\text{\TeX/L\TeX}$  file prepared by the author.