

Direct detection of exoplanet host star companion γ Cep B and revised masses for both stars and the sub-stellar object

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ABSTRACT

Context. The star γ Cep is known as a single-lined spectroscopic triple system at a distance of 13.8 pc, composed of a K1 III–IV primary star with $V = 3.2$ mag, a stellar-mass companion in a 66–67 year orbit (Torres 2007, ApJ, 654, 1095), and a substellar companion with $M_p \sin i = 1.7 M_{\text{Jup}}$ that is most likely a planet (Hatzes et al. 2003, ApJ, 599, 1383).

Aims. We aim to obtain a first direct detection of the stellar companion, to determine its current orbital position (for comparison with the spectroscopic and astrometric data), its infrared magnitude and, hence, mass.

Methods. We use the Adaptive Optics camera CIAO at the Japanese 8 m telescope Subaru on Mauna Kea, Hawaii, with the semi-transparent coronagraph to block most of the light from the bright primary γ Cep A, and to detect at the same time the faint companion B. In addition, we also used the IR camera Ω Cass at the Calar Alto 3.5 m telescope, Spain, to image γ Cep A and B by adding up many very short integrations (without AO).

Results. γ Cep B is clearly detected on our CIAO and Ω Cass images. We use a photometric standard star to determine the magnitude of B after PSF subtraction in the Subaru image, and the magnitude difference between A and B in the Calar Alto images, and find an average value of $K = 7.3 \pm 0.2$ mag. The separations and position angles between A and B are measured on 15 July 2006 and 11 and 12 Sept. 2006, B is slightly south of west of A.

Conclusions. By combining the radial velocity, astrometric, and imaging data, we have refined the binary orbit and determined the dynamical masses of the two stars in the γ Cep system, namely $1.40 \pm 0.12 M_{\odot}$ for the primary and $0.409 \pm 0.018 M_{\odot}$ for the secondary (consistent with being a M4 dwarf). We also determine the minimum mass of the sub-stellar companion to be $M_p \sin i = 1.60 \pm 0.13 M_{\text{Jup}}$.

Key words. instrumentation: adaptive optics – binaries: spectroscopic – binaries: visual – planetary systems – star: individual: γ Cep

1. Introduction: γ Cep

The bright star γ Cep (also HD 222404 or HIP 116727) located near the north celestial pole ($\alpha = 23^{\text{h}}39^{\text{m}}20.8^{\text{s}}$ and $\delta = +77^{\circ}37'56.2''$ for J2000.0) is known as a single-lined spectroscopic triple: The primary star has spectral type K1 III–IV, is visible to the naked eye ($V = 3.2$ mag) and is located at a distance of 13.8 pc. A low-mass stellar companion with an orbital period of several decades was discovered spectroscopically by Campbell et al. (1988), but until now it has never been imaged directly (e.g. Hatzes et al. 2003; Mugrauer et al. 2006). Campbell et al. (1988) and Walker et al. (1992) discussed the evidence for additional radial velocity variations indicative of a very low-mass, possibly planetary companion. Hatzes et al. (2003) confirmed that there is indeed a third object in the system with a minimum mass of $M_p \sin i = 1.7 M_{\text{Jup}}$, i.e., probably a planet, with a semi-major axis of 2 AU and a 900-day eccentric orbit ($e = 0.12$).

Several investigators have attempted to constrain the orbit as well as the mass of γ Cep B (the stellar companion) on the basis of available spectroscopic data. The estimated period of the binary from these studies has ranged from about 30 to 66 yr (Walker et al. 1992; Griffin et al. 2002). Most recently, Torres (2007) put together not only the spectroscopic data (dating back to 1902), but also *Hipparcos* and ground-based

astrometric data (dating back to 1898) and obtained new results for the primary star (γ Cep Aa), the secondary (B), and the sub-stellar companion (Ab): B orbits A in an eccentric orbit ($e = 0.4085 \pm 0.0065$) with 66.8 ± 1.4 year period and a semi-major axis of 19.02 ± 0.64 AU; the primary has a mass of $1.18 \pm 0.11 M_{\odot}$, temperature of 4800 ± 100 K, and an age of 6.6 Gyr; the unseen stellar companion should have a mass of $0.362 \pm 0.022 M_{\odot}$, i.e. spectral type about M4; the minimum mass of the lowest-mass component in the system is then $M_p \cdot \sin i = 1.43 \pm 0.13 M_{\text{Jup}}$; the astrometric data yield an upper mass limit of $17 M_{\text{Jup}}$ (at the 3σ confidence level), i.e. certainly sub-stellar, probably planetary; all values above are from Torres (2007). The predictions for γ Cep B can be tested by direct imaging.

In this paper, we present the first direct imaging detections of γ Cep B (Sect. 2), compare the parameters observed with the predictions (Sect. 3), and newly determine the masses of γ Cep Aa, Ab, and B (Sect. 4).

2. Imaging observations

We observed γ Cep first on 15 July 2006 with the Coronagraphic Imager for Adaptive Optics (CIAO) on the Japanese 8 m telescope Subaru on Mauna Kea, Hawaii, in the K -band. The night was photometric with ~ 0.7 arcsec natural seeing. We took

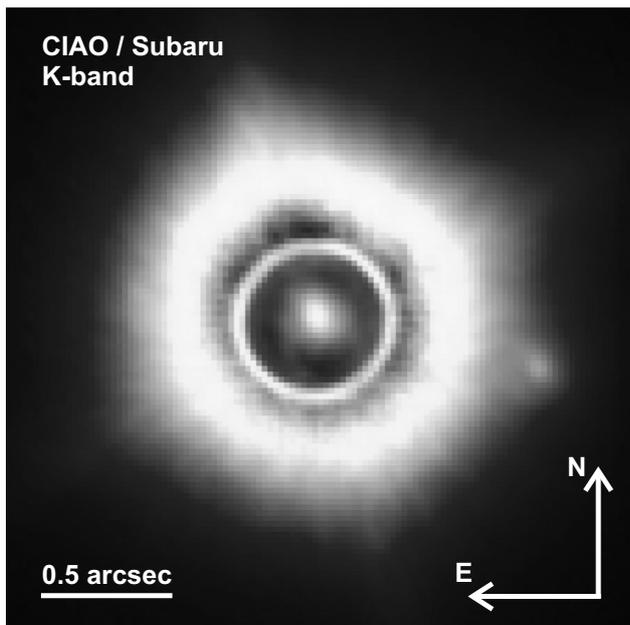


Fig. 1. Subaru CIAO K -band image of γ Cep A and B, where A is below the semi-transparent coronagraph. The total integration time was 30 s, obtained on 15 July 2006.

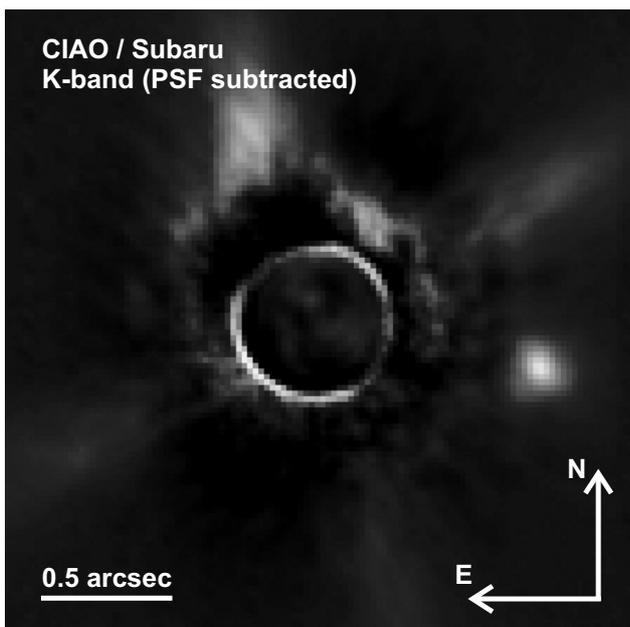


Fig. 2. Subaru CIAO K -band image of γ Cep A and B, now after subtraction of the PSF of γ Cep A. The stellar companion γ Cep B is now clearly detected 0.87 arcsec west of A (slightly south of west).

ten images (always 30 times 0.1 s each, i.e. 30 s total integration time) through a semi-transparent coronagraphic mask (0.6 arcsec diameter, a chrome spot on a sapphire substrate) in order to block most of the light of γ Cep A, but still to be able to detect B.

The data were reduced in the normal way: Dark subtraction, flat-field division, and adding up the individual images. The resulting image is shown in Fig. 1.

The semi-transparent coronagraph has a transmission of between about 1% and 6%, thus still allowing the primary star to shine through. We subtracted the point spread function (PSF) of

remaining light of γ Cep A, see Fig. 2 for the result. γ Cep B is clearly detected in Figs. 1 and 2.

Because the throughput of the coronagraph is not well known and because the star A may still be in the non-linear regime, we cannot derive the magnitude of B with respect to A from the image itself. A photometric standard star is needed. We therefore observed the standard star FS 150 (9.94 ± 0.01 mag, UKIRT web page) in the same night (8×3 s) to obtain the magnitude of γ Cep B by aperture photometry (after PSF subtraction). The result for B is $K = 7.32 \pm 0.11$ mag. Hence, the magnitude difference between γ Cep A and B is 6.28 ± 0.24 mag (A has $K_s = 1.04 \pm 0.21$ mag, 2MASS Catalog, Skrutskie et al. 2006).

By Gaussian centering, we can determine the position of B relative to A, i.e. to measure the separation and position angle. For a companion (or candidate) to a very bright primary, it is very useful to observe with a *semi-transparent* coronagraph, so that the primary star is still visible together with the companion (for separation measurement). For the astrometric analysis, we use the most recent pixel scale for CIAO, namely 21.3 ± 0.1 mas/pixel (Mayama et al. 2006) and the orientation of the detector being 0.81 ± 0.25 deg from north as measured by us from images of three sources in the V1686 Cyg region taken in the same night.

The resulting separation and position angle are listed in Table 1.

After our detection with Subaru we observed γ Cep with the Ω Cass IR imager on the Calar Alto 3.5 m telescope in Spain. Due to the fact that our instrument of choice, the AO system ALFA, was not functioning at the time, we observed γ Cep in a backup programme by taking several thousand short (0.1 s) integration images (speckle), which we then combined by shift+add, after background correction and flatfielding. We detected γ Cep B on both 11 Sept. (7391 images) and on 12 Sept. 2006 (18140) using the K' (K -prime) filter together with Br γ . In addition to measuring the separation between B and A from the resulting images in each night, we determined also the magnitude difference relative to the primary, which was neither saturated nor in the non-linear regime of the detector. See Fig. 3 for the final images of the two nights and Table 1 for the results of our measurements. We used the Hipparcos multiples HIP 109474, HIP 91115, HIP 92961, HIP 104536, observed in the same two nights, for astrometric calibration of the pixel scale (38.681 ± 0.059 mas/pixel) and the detector position angle ($21.362 \pm 0.037^\circ$).

Our measurements may indicate a small change in the position angle of 0.63 ± 0.36 deg between July and September 2006, and a more significant increase in separation of 19 ± 6 mas, consistent with the predictions for orbital motion, see Fig. 4.

For the secondary star γ Cep B, imaged for the first time here, we obtain a magnitude of $K = 7.3 \pm 0.2$ mag, the average of the Subaru and the two Calar Alto data points.

3. Comparison with predictions

Our new measurements of the relative position of γ Cep B with respect to the primary star allow us to test the predictions from the orbital solution for the binary produced recently by Torres (2007). That study combined all available measurements of the radial velocity of the primary from a variety of sources (including the high-precision studies that uncovered the sub-stellar companion, along with classical measurements from the literature and from new observations) and used also the intermediate astrometry from the *Hipparcos* mission (“abscissa residuals”) as

Table 1. Astrometry and photometry for γ Cep A and B.

Obs date HJD	Telescope instrument	PA [°]	O–C [°]	separation [arcsec]	O–C [mas]	ΔK [mag]
2453932.046	Subaru 8.3 m CIAO	256.91 ± 0.27	-0.20	0.870 ± 0.005	-0.005	6.28 ± 0.24
2453989.520	CA 3.5 m Ω Cass	256.16 ± 0.35	+0.09	0.891 ± 0.006	+0.005	6.28 ± 0.13
2453990.514	CA 3.5 m Ω Cass	256.39 ± 0.34	+0.33	0.887 ± 0.005	+0.001	6.12 ± 0.11

Remark: Observed minus computed (O–C) residuals are explained in Sect. 3.

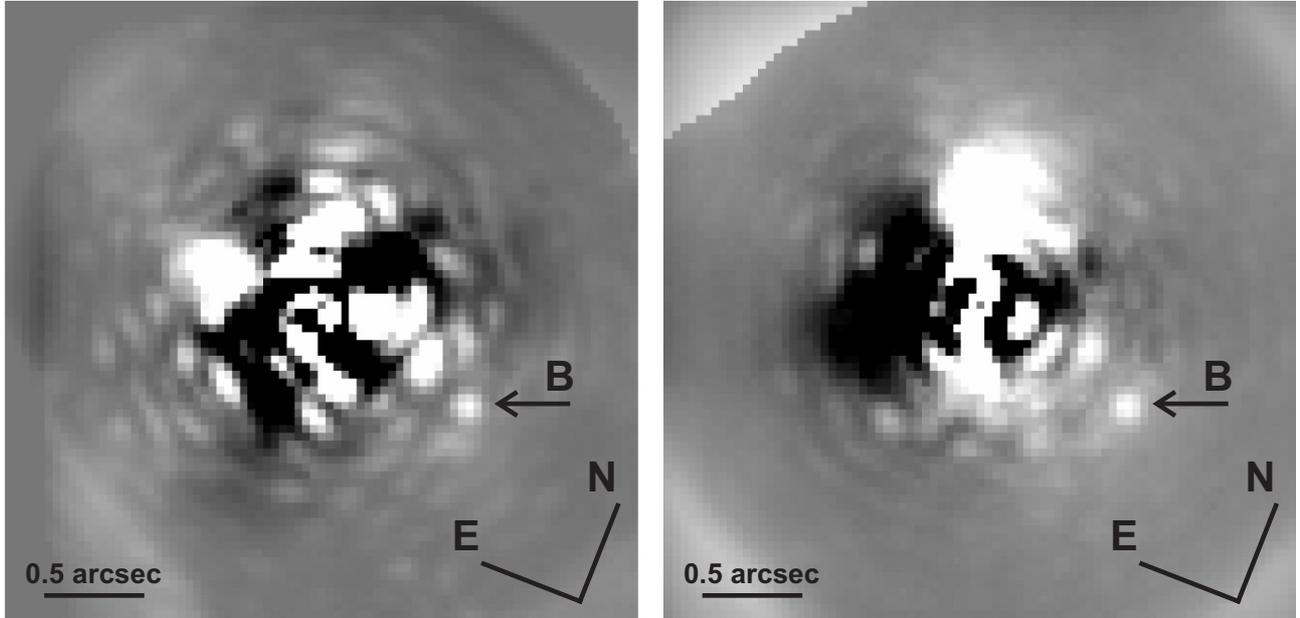


Fig. 3. Ω Cass IR speckle images of γ Cep A and B, obtained on 11 Sept. 2006 (left) and 12 Sept. 2006 (right) with the Calar Alto 3.5 m telescope with ~ 12 to 30 min total integration time (through K' and Br γ filters).

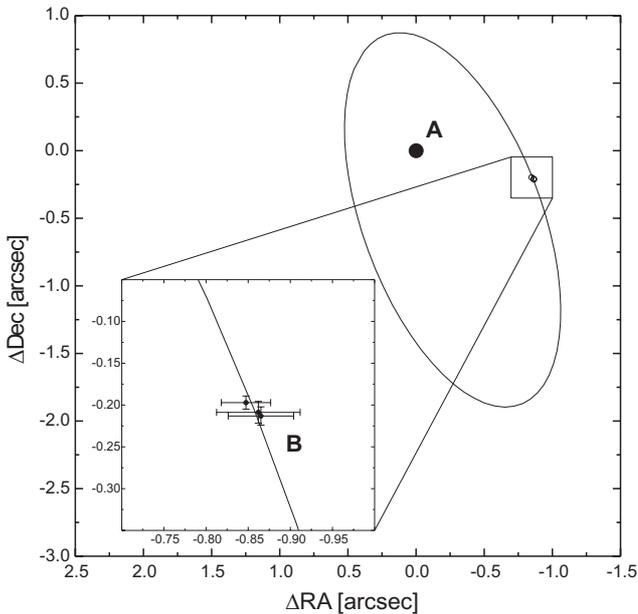


Fig. 4. Orbit and orbital motion of γ Cep B around A as already observed by our imaging within two month. The inlay shows our three imaging data points of B in orbit around A.

to be established for the first time, in a combined solution that included also the perturbing effects of the sub-stellar companion.

A comparison of our measurements with the published orbit (Torres 2007) indicates systematic (observed minus computed) O–C residuals of $\sim 8^\circ$ in position angle and 50–60 mas in separation (positive in both cases). These are many times larger than the measurement errors, although the predicted positions themselves have uncertainties of about 3.1 degrees in position angle and 34 mas in separation for the dates of our observations. These reflect the uncertainty in the orbital elements as well as the estimated error of the adopted primary mass, which was external to the orbital solution of Torres (2007) and was established from stellar evolution considerations. Therefore, the residuals correspond to 1.5–2.5 σ , which, while not overly large, may nevertheless indicate the need for some adjustment in the orbit given that they are systematic in nature. The new O–C values given in Table 1 (this paper) are computed for the new orbit (Table 2) after taking into account our three imaging observations. They are much smaller than in Torres (2007), i.e. a clear improvement as expected, which was the motivation for the imaging observations. The relative brightness of the secondary was predicted by Torres (2007) to be $\Delta K \sim 6.4$ mag, which is not far from the values we measure (see Table 1).

4. Improved masses

well as ground-based astrometry reaching back more than a century. This enabled the scale and orientation of the binary orbit

The assumption of a value for the primary mass in the Torres (2007) solution was necessary because of the lack of direct measurements of the relative position between γ Cep A and B, since

Table 2. Global orbital solution for γ Cep.

Parameter and value	
Adjusted quantities from outer orbit (A+B)	
P_{AB} (yr)	67.5 ± 1.4
γ (km s^{-1})	-42.943 ± 0.046
K_A (km s^{-1})	1.932 ± 0.014
e_{AB}	0.4112 ± 0.0063
ω_A (deg)	161.01 ± 0.40
T_{AB} (yr)	1991.605 ± 0.031
a''_{AB} (arcsec)	1.467 ± 0.046
a''_A (mas)	332.4 ± 7.7
i_{AB} (deg)	119.3 ± 1.0
Ω_{AB} (deg)	18.04 ± 0.98
Adjusted quantities from inner orbit (Aa+Ab)	
P_A (days)	902.9 ± 3.5
K_{Aa} (m/s)	27.0 ± 1.5
e_A	0.115 ± 0.058
ω_{Aa} (deg)	63 ± 27
T_A (HJD-2 400 000)	53147 ± 71
Derived quantities	
μ_α^* (mas yr^{-1})	-63.86 ± 0.98
μ_δ (mas yr^{-1})	$+150.76 \pm 0.43$
π (mas)	72.69 ± 0.41
a_{AB} (AU)	20.18 ± 0.66
M_{Aa} (M_\odot)	1.40 ± 0.12
M_B (M_\odot)	0.409 ± 0.018
$f(M_p)$ ($10^{-9} M_\odot$)	1.81 ± 0.31
$M_p \sin i$ (M_{Jup})	1.60 ± 0.13
a_{Aa-p} (AU)	2.044 ± 0.057

Remark: The symbols have the same meaning as in Torres (2007).

the secondary had never been imaged. With our present results this assumption can now be dropped, and this allows for a direct measurement of the dynamical masses of both stars with no need for external constraints. In order to update the orbital solution and solve for the masses we have added to the data considered by Torres (2007) our 3 measurements of the separation and position angle of γ Cep with their corresponding uncertainties. We have modified the procedures to include the relative semimajor axis of the binary, a''_{AB} , as an adjustable parameter in addition to the other variables considered in the original solution. The results are shown in Table 2, which includes the most relevant elements of the inner and outer orbits and several derived quantities. The notation follows that of Torres (2007). Most of the orbital elements change very little (well within the errors). The most significant change is in the position angle of the ascending node

(Ω_{AB}), which increases by about 5° . This accounts for a large fraction of the residuals in PA noted above. The other significant change is in the scale of the orbit, which increases by about 6%. This in turn accounts for the residuals in the angular separation. The residuals from this new orbit are listed in Table 2, and are now all within 1σ , indicating a satisfactory fit.

The resulting masses of the two stars are $M_{Aa} = 1.40 \pm 0.12 M_\odot$ and $M_B = 0.409 \pm 0.018 M_\odot$. Both the primary and secondary are somewhat more massive than assumed by Torres (2007). Consequently the minimum mass of the sub-stellar companion also increases to $M_p \sin i = 1.60 \pm 0.13 M_{\text{Jup}}$, a value intermediate between those proposed by Torres (2007) and Hatzes et al. (2003), and the upper mass limit from Torres (2007) changes from $17 M_{\text{Jup}}$ to $19 M_{\text{Jup}}$. The brightness of the stellar companion B is fully consistent with a single M4-type dwarf, neither multiple itself nor a white dwarf.

Given the new parameters in Table 2, the critical semi-major axis for long-term stability of planets according to the Holman & Wiegert (1999) criterion is at only 3.85 ± 0.38 AU. The size of a planet-forming circumstellar disk in the close binary is not larger than 3.4 AU according to the Pichardo et al. (2005) criteria. The planet candidate γ Cep Ab orbits around Aa between 1.8 and 2.3 AU, which is inside the stable region, so that it is not necessary to invoke migration after its formation.

Continued measurements of the relative position between γ Cep A and B over the coming years, which we intend to carry out, should further improve the stellar mass determinations and help establish also the true mass of the sub-stellar companion.

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