

Searching for gas emission lines in *Spitzer* Infrared Spectrograph (IRS) spectra of young stars in Taurus[★]

C. Baldovin-Saavedra^{1,2}, M. Audard^{1,2}, M. Güdel³, L. M. Rebull⁴, D. L. Padgett⁴, S. L. Skinner⁵,
A. Carmona^{1,2}, A. M. Glauser^{6,7}, and S. B. Fajardo-Acosta⁸

¹ ISDC Data Centre for Astrophysics, Université de Genève, 16 chemin d'Ecogia, 1290 Versoix, Switzerland
e-mail: carla.baldovin@unige.ch

² Observatoire Astronomique de l'Université de Genève, 51 chemin de Maillettes, 1290 Sauverny, Switzerland

³ University of Vienna, Department of Astronomy, Türkenschanzstrasse 17, 1180 Vienna, Austria

⁴ Spitzer Science Center, California Institute of Technology, 220-6 1200 East California Boulevard, CA 91125 Pasadena, USA

⁵ Center for Astrophysics and Space Astronomy, University of Colorado, CO 80309-0389 Boulder, USA

⁶ ETH Zürich, 27 Wolfgang-Pauli-Str., 8093 Zürich, Switzerland

⁷ UK Astronomy Technology Centre, Royal Observatory, Blackford Hill, EH3 9HJ Edinburgh, UK

⁸ IPAC, California Institute of Technology, 770 South Wilson Avenue, CA 91125 Pasadena, USA

Received 19 August 2010 / Accepted 14 January 2011

ABSTRACT

Context. Our knowledge of circumstellar disks has traditionally been based on studies of dust. However, gas dominates the disk mass and its study is key to our understanding of accretion, outflows, and ultimately planet formation. The *Spitzer* Space Telescope provides access to gas emission lines in the mid-infrared, providing crucial new diagnostics of the physical conditions in accretion disks and outflows.

Aims. We seek to identify gas emission lines in mid-infrared spectra of 64 pre-main-sequence stars in Taurus. Using line luminosities and other known star-disk-outflow parameters, we aim to identify correlations that will help to constrain gas heating, excitation mechanisms, and the line formation.

Methods. We have based our study on *Spitzer* observations using the Infrared Spectrograph (IRS), mainly with the high-resolution modules. Line luminosities (or 3σ upper limits) have been obtained by fitting Gaussian profiles to the lines. We have further searched for correlations between the line luminosities and different parameters related to the star-disk system.

Results. We have detected H₂ (17.03, 28.22 μ m) emission in 6 objects, [Ne II] (12.81 μ m) emission in 18 objects, and [Fe II] (17.93, 25.99 μ m) emission in 7 objects. [Ne II] detections are found primarily in Class II objects. The luminosity of the [Ne II] line (L_{NeII}) is in general higher for objects known to drive jets than for those without known jets, but the two groups are not statistically distinguishable. L_{NeII} is correlated with X-ray luminosity, but for Class II objects only. L_{NeII} is also correlated with disk mass and accretion rate when the sample is divided into high and low accretors. Furthermore, we find correlations of L_{NeII} with mid-IR continuum luminosity and with luminosity of the [O I] (6300 Å) line, the latter being an outflow tracer. $L_{[\text{FeII}]}$ correlates with \dot{M}_{acc} . No correlations were found between L_{H_2} and several tested parameters.

Conclusions. Our study reveals a general trend toward accretion-related phenomena as the origin of the gas emission lines. Shocks in jets and outflowing material are more likely to play a significant role than shocks in infalling material. The role of X-ray irradiation is less prominent but still present for [Ne II], in particular for Class II sources, while the lack of correlation between [Fe II] and [Ne II] points toward different emitting mechanisms.

Key words. ISM: jets and outflows – infrared: stars – protoplanetary disks – stars: formation – stars: pre-main sequence – stars: protostars

1. Introduction

Protoplanetary disks are formed by a mixture of gas, dust and ices. The bulk of their mass is in gaseous form (mainly H₂) but the opacity is dominated by dust at infrared (IR) and optical wavelengths. Ro-vibrational transitions of H₂ in the IR are difficult to stimulate due to the combination of low temperatures found in protoplanetary disks and the lack of a permanent electric dipole moment of the H₂ molecule. For this reason, our current knowledge of protoplanetary disks is largely based on the study of dust (e.g., Beckwith et al. 1990; Hartmann 2008; Mann & Williams 2009).

One of the key questions in star formation is the disk lifetime, since it provides a strong constraint on the planet formation

timescale. Many studies have been dedicated to this problem (e.g., Carpenter et al. 2009; Meyer 2009). The signatures of protoplanetary disks, namely thermal emission from dust grains in the millimeter and the infrared, and molecular line emission from CO in the sub-millimeter, typically disappear after a couple of Myr (e.g., Duvert et al. 2000; Andrews & Williams 2005). However, it is not clear whether the material has been completely cleared out or if present observations are not sensitive enough to allow the detection of its remnant. Recent evidence indicates that the gas and dust lifetimes are not the same (e.g., Pascucci et al. 2009; Fedele et al. 2010). Detections of quiescent emission of molecular H₂ in weak-lined T Tauri stars (WTTS) that do not show infrared excess, i.e., those that are not expected to have a disk, support the hypothesis that gas lives longer than dust in disks (e.g., Bary et al. 2002, 2008a). To address the question of

[★] Appendices and Tables 1 and 2 are only available in electronic form at <http://www.aanda.org>

disk dissipation timescales, studies of the gas component of protoplanetary disks independent of dust are particularly needed.

To study gas in protoplanetary disks, molecular spectroscopy, in particular emission lines, can be used. As the temperature and density of the gas vary as function of the radius in the disk and the mid-plane distance, spectral lines at different wavelengths probe different disk regions, and gas at different physical conditions. In the sub-millimeter and millimeter regimes, observations probe cold gas ($T \sim 20\text{--}50$ K) at large radii (>20 AU) through emission lines such as CO species and HCO^+ (e.g., Mannings & Sargent 1997; Piétu et al. 2007).

In the near-IR, the H_2 emission arises from ro-vibrational transitions, probing gas at $T \sim 1000$ K (e.g., Weintraub et al. 2000; Carmona et al. 2007; Bary et al. 2008b). After H_2 , CO is the most abundant species found in disks. The different excitation conditions of this molecule probe a wide range of disk radii; the first overtone band at $2.3 \mu\text{m}$ traces hot (~ 4000 K) and dense material ($n > 10^{10} \text{ cm}^{-3}$) of the innermost part of protoplanetary disks. Observations show that the emission has the characteristic shape suggestive of a rotating disk (Carr 1989). The CO ro-vibrational band at $4.7 \mu\text{m}$ is estimated to arise within 1–10 AU from the star (e.g., Najita et al. 2003; Blake & Boogert 2004; Rettig et al. 2004; Salyk et al. 2007; Carmona et al. 2008b) with characteristic temperatures between 100 to a few 1000 K.

The mid-IR traces the warm gas at radii of a few AU up to several tens of AU (the giant planet formation region). At these wavelengths, H_2 emission is observed, arising from pure rotational transitions at 12.28, 17.03, and 28.22 μm (e.g., van den Ancker et al. 1999; Lahuis et al. 2007; Martin-Zaïdi et al. 2007; Bitner et al. 2008). Furthermore, many detections of forbidden atomic lines have been reported, such as [Ne II] (12.81 μm), [Fe I] (24.0 μm), [Fe II] (17.93 and 25.99 μm), and [S III] (18.71 μm), opening up a new window into the diagnosis of gas in protoplanetary disks.

When studying the gas emission from protoplanetary disks, it is important to discriminate among emission produced by (i) gas in jets or outflows; (ii) accreting gas falling onto the star; and (iii) gas in the disk itself. High spectral and spatial resolution observations enable the study of the line profiles and possible velocity shifts with respect to the stellar rest velocity. However, the heating mechanism of gas and the ionization of atomic species is not well understood. Glassgold et al. (2007) predicted that stellar X-rays heat the upper layers of protoplanetary disks within 25 AU from the star, exciting atomic species. Gorti & Hollenbach (2008) considered both stellar X-rays and ultraviolet (UV) irradiation, proposing several emission lines as possible diagnostics of the gas in disks irradiated by high energy photons: forbidden lines of [Ne II] (12.81 μm), [Ne III] (15.55 μm), [Fe II] (25.99 μm), [O I] (63 μm), and pure rotational lines of H_2 . Moreover, Alexander (2008) successfully reproduced the observed [Ne II] line luminosities by modeling the line profile using a photoevaporative wind model, which considered the UV flux from the central star as the only ionizing source. In addition, far-UV (FUV), extreme-UV (EUV), and X-ray photons may be produced by accretion shocks as the material falls into the central star (Hollenbach & Gorti 2009). Forbidden emission lines in the optical and near-IR have been used for a long time to probe the shocked gas from T Tauri stars (e.g., Mundt et al. 1990; Hirth et al. 1997; Hartigan & Hillenbrand 2009; Agra-Amboage et al. 2009). Hollenbach et al. (1989) studied the infrared emission from interstellar shocks and proposed that for low pre-shock densities ($n \leq 10^{4-5} \text{ cm}^{-3}$), forbidden emission lines of [Fe II], [Ne II], [Si II], and [O I] are particularly strong.

Forbidden emission lines are used to trace the electron densities of the environment where they originate; typical densities of $10^3\text{--}10^7 \text{ cm}^{-3}$ are probed. At these densities, the origin is in low-density winds or shocks rather than in the infall of material through accretion funnels. On the other hand, permitted emission lines exhibiting large widths and inverse P-Cygni features are associated with infalling material (Muzerolle et al. 2001).

Up to now, observations have not been conclusive in giving a broadly accepted explanation of the heating and ionizing mechanism exciting the gas component in protoplanetary disks. We aim to contribute to the understanding of this problem by studying a large sample of pre-main sequence stars of low and intermediate mass. We have used observations done by the *Spitzer* Space Telescope (Werner et al. 2004) Infrared Spectrograph (IRS, Houck et al. 2004) in order to study the emission from gas lines. Although higher spectral resolution observations are needed in order to obtain kinematical information for the emitting gas, the high sensitivity of *Spitzer* does enable us to access sources not detectable from the ground. Furthermore, the high number of sources observed allows the extraction of statistically significant information from a large sample. We have made use of archival spectra of 61 pre-main sequence stars plus 3 objects from our own *Spitzer* program. This study is complementary to previous investigations; our work includes objects of different evolutionary stages, while the somewhat related previous study by Güdel et al. (2010) concentrated on a pure Class II sample. Additionally, in our analysis we have made a distinction based on the YSO class, in contrast to previous studies (e.g., Lahuis et al. 2007; Flaccomio et al. 2009).

Sections 2 and 3 are dedicated to the description of the sample and of the data analysis, respectively. Results are presented in Sect. 4 and discussed in Sect. 5 together with the presentation of correlation tests. A brief summary of the results and conclusions are given in Sect. 6. In Appendix A we have included additional information on individual objects whenever one of the studied gas lines was detected.

2. Sample

Our study is based on objects from the Taurus Molecular Cloud (TMC). This star-forming region is known to form mainly low-mass stars, allowing the study of the star formation process without the influence of high-mass young stars, which are known to have strong UV radiation fields. Its nearby distance (137 pc, Torres et al. 2007) enables us to access fainter sources. All the objects have similar ages, lie at nearly the same distance, and are presumably formed out of the same parent cloud.

A further motivation to focus our study on the TMC region is the large amount of multi-wavelength observations available. The TMC was indeed been observed extensively in X-rays with *XMM-Newton*: the XEST survey covered ~ 5 square degrees of the TMC, Güdel et al. 2007. In the optical, the Canada-France-Hawaii Telescope (CFHT) covered 28 square degrees of the TMC (Monin et al. 2010, in prep.), and the Sloan Digital Sky Survey (SDSS) covered 48 square degrees (Padmanabhan et al. 2008). In the IR, *Spitzer* mapped ~ 43 square degrees (*Spitzer* Taurus Legacy Project; Padgett et al. 2007, 2011, in prep.). Multiwavelength observations are crucial in order to test the influence of the stellar X-rays in the heating and excitation of the gas species.

Our study is based on publicly available archived data obtained by IRS with its high resolution modules (short-high – SH – and long-high – LH), with a resolving power of $R \sim 600$ ($\sim 500 \text{ km s}^{-1}$) for wavelength coverage between 9 and

37 μm . Most of the targets studied belong to GTO (Guaranteed Time Observations) programs; several publications exist that focussed on the study of the silicate and ice features (e.g., Watson et al. 2004; Furlan et al. 2006; Furlan et al. 2008). In particular, we used data from the GTO program number 2, from the Legacy programs 171, 179, and from the GO (General Observer) programs 3470, 3716, and 30300. In addition, we included 3 low-mass pre-main sequence stars identified by the *Spitzer* TMC survey for which we obtained IRS high and low-resolution spectra (program 50807, PI. Baldovin-Saavedra). For one object (SST 041412.2+280837) we only have the spectrum taken with the low-resolution modules (long-low – LL – and short-low – SL).

Our sample consists of 64 low and intermediate mass pre-main sequence stars. Their infrared classification is taken from the literature and is based on the slope of the spectral energy distribution (SED) in the infrared. The sample includes 10 Class I objects, 40 Class II, 4 Class III, and 6 intermediate mass Herbig AeBe stars. We found 4 objects classified in the literature as both Class I and Class II. Table 1 summarizes the properties of the stars in our sample, including the bibliographic references for the adopted values.

3. Data analysis

3.1. Initial steps

The data were retrieved from the *Spitzer* archive; we started the analysis from the post-basic calibrated data, i.e., spectra extracted automatically by the pipeline (version S15.3.0 for SH and S17.2 for LH¹). The spectra of the three targets from our own program were treated using the pipeline version S18.7.0².

In a first step we used IRSCLEAN³ to create masks for the rogue pixels from the basic calibrated data images (rogue pixels have abnormally high dark current and/or photon responsivity). Once the images were free from obvious rogue pixels, we re-extracted the spectra using SPICE (*Spitzer* IRS Custom Extraction) regular extraction. Furthermore, in the cases where we detected emission lines or where the detections were not obvious, we also used optimal extraction (details on the regular and optimal extraction can be found in the SPICE User's Guide Version, 16.2.1+).

Some of the spectra also presented residual fringes; these interference patterns are typical for infrared detectors as their surface coating acts as a Fabry-Pérot etalon, leading to a sinusoidal modification of the detector response of spectrally dispersed light. Although they were corrected during the pipeline processing, residual fringes remained in a few cases that we removed using the tool IRSFRINGE³.

3.2. Background

The mid-IR background comes predominantly from zodiacal light at the IRS wavelengths, while diffuse emission from the Taurus cloud contributes only little. Instrumental background (rogue pixels) is also significant for the high-resolution modules. Therefore, sky observations are needed in order to correct

for the sky and instrumental background. In this study, however, most of our source sample had no dedicated background observations. Good rogue pixel removal proved to be crucial for our emission line study.

Some of the GTO observations were performed using a series of 6 pointings per object with a shift corresponding to approximately half of the PSF. In these cases, for the on-source observations we used the 2 pointings with higher flux and averaged them; these pointings are equivalent to the standard staring mode observations, i.e., 2 slit positions with the target located at 30 and 60% of the slit length. For each object, we checked using *Spitzer* Infrared Array Camera (IRAC) 8 μm images that the pointings correspond to the SIMBAD position of the target, and that the flux levels of the 2 used pointings were equivalent.

One central issue lies in understanding the origin of the emission. It could be that the lines detected are not really emitted by the star-disk system, but rather come from the background or a nearby source. Our data are in the form of an echellogram, which is a two-dimensional representation of a spectrum containing both spatial and spectral information. As a first step, we searched each spectrum and searched for evident line emission; when found, we examined in the echellogram two spectral regions with no emission lines adjacent to the position of the line studied. These two regions were used as a measure of local background plus stellar continuum; they were averaged and then subtracted from the line region. In this way, any possible extended emission can be identified; moreover, by removing the continuum, the detection becomes more evident. This method did not result in the discarding of any detections. However, it is not sensitive to potential extended emission from the direct vicinity of the star.

3.3. Line fitting

In order to obtain line luminosities, we used a function consisting of a Gaussian plus a linear component to account for the continuum within a range of typically $\pm 0.6 \mu\text{m}$ of the line. All the parameters of the function (height, center, and width) were left free to vary. The integrated flux was calculated as the area under the Gaussian fit once the continuum was subtracted, and the line luminosity was calculated using a distance to the TMC of 137 pc from Torres et al. (2007)⁴. A detection was claimed when the peak of the line was higher than 3 times the level of the noise in the spectra. The noise level is the standard deviation of the continuum flux (σ) in the wavelength range of interest.

Most of the detected lines were spectrally unresolved; the full width at half-max (*FWHM*) obtained from the fit was similar to the instrumental *FWHM* ($\sim 500 \text{ km s}^{-1}$). In the cases where the lines were not been detected, we obtained 3σ upper limits for their luminosities by multiplying the 3σ noise value of the continuum by the *FWHM*.

4. Results

Our study is focused on the search for pure rotational emission lines from H_2 (12.28, 17.03, 28.22 μm), and forbidden emission lines from [Ne II] (12.81 μm), [Ne III] (15.56 μm), [Fe II] (17.94, 25.99 μm), and [S III] (18.71, 33.48 μm). We detected [Ne II] in 18 objects in our sample, [Fe II] in 7 objects, and H_2 in 6 objects. We did not detect [Ne III] or [S III]. The

⁴ The distance to the TMC has often been considered to be 140 pc. In comparison to the 137 pc used in this study, this can result in a difference of the order of 4% in the line luminosities.

¹ We checked for a few sources that the latest version of the pipeline (S18.7.0) did not change the obtained line luminosities.

² Our observations were reduced from the basic calibrated data, correcting for the co-adding bug in S18.7.0 and for time and position-dependent dark currents on the LH array (DarkSettle 1.1).

³ Available from <http://ssc.spitzer.caltech.edu>

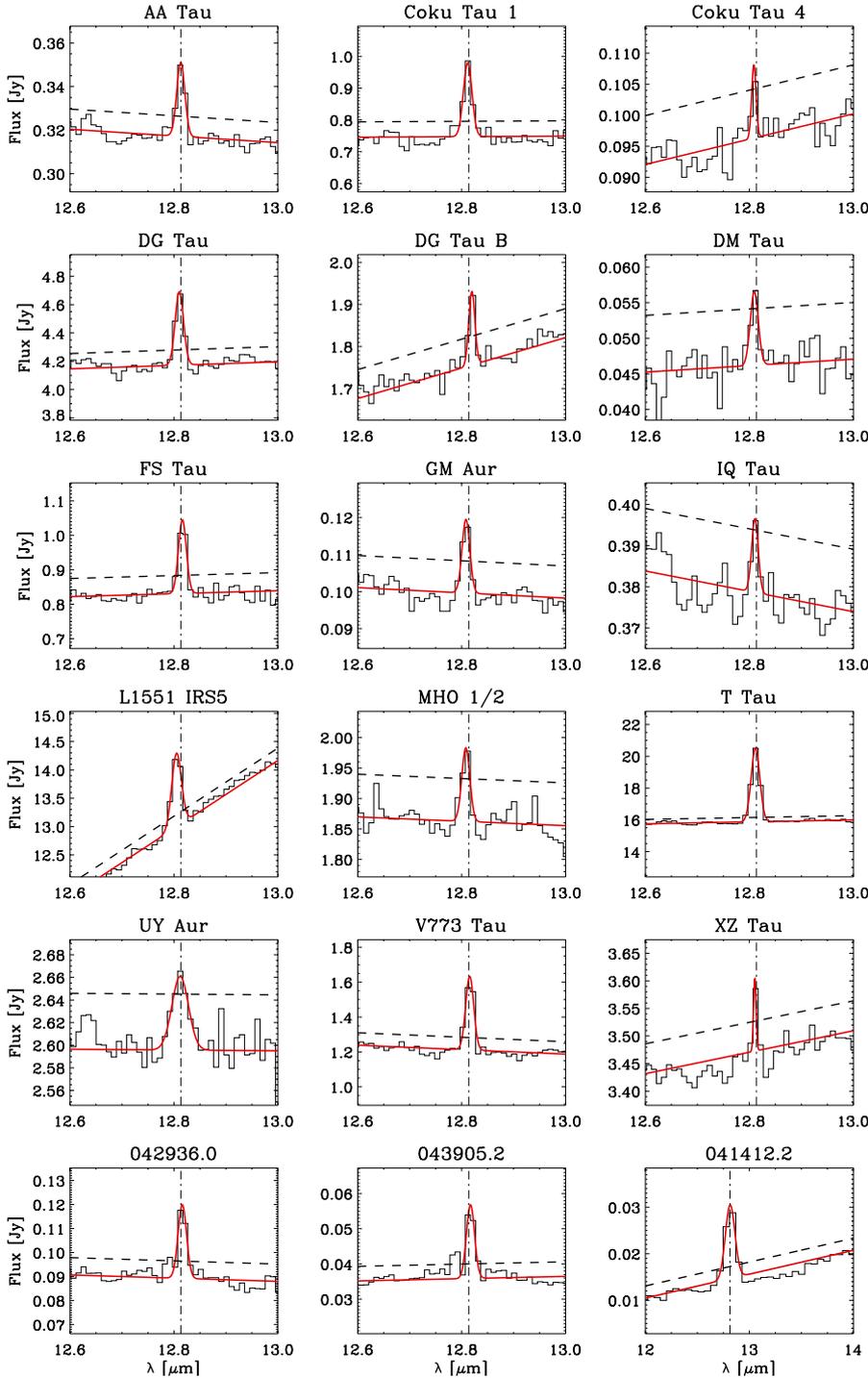


Fig. 1. Detections of [Ne II] at $12.81 \mu\text{m}$. The spectra are plotted as solid histogram, the lines were fitted using a gaussian profile plus a polynomial component for the continuum, the fit is plotted using a solid line. The 3σ detection threshold is plotted as a dashed line. The name of each star is labeled in the top of each plot. We recall that the spectrum of SST 041412.2+280837 was obtained using the low resolution modules (SL and LL).

line luminosities and upper limits are presented in Table 2, with the luminosities expressed in units of $10^{28} \text{ erg s}^{-1}$; bold fonts indicate detected lines, and normal fonts are 3σ upper limits. In Appendix A we provide information on the single objects with detected emission lines.

In Fig. 1 we present the spectra of the targets with detected [Ne II] line emission, whereas we present the spectra surrounding the position of the [Ne II] line for the sources with non-detections in Appendix B. Figures 2 and 3 show the line spectra of the sources with detected emission lines from H_2 and [Fe II]. In all figures, the red line shows the Gaussian fit to the line, the dashed line shows 3σ detection threshold, and the vertical

dashed-dotted line shows the expected position of the [Ne II] line. The title of each panel is the name of the target.

5. Discussion

5.1. [Ne II] emission

The search for [Ne II] emission has gained importance in the last few years thanks to an increasing number of reported detections in the spectra of young stars observed mainly with *Spitzer* (e.g., Pascucci et al. 2007; Lahuis et al. 2007; Espaillat et al. 2007; Flaccomio et al. 2009; Güdel et al. 2010). The sample

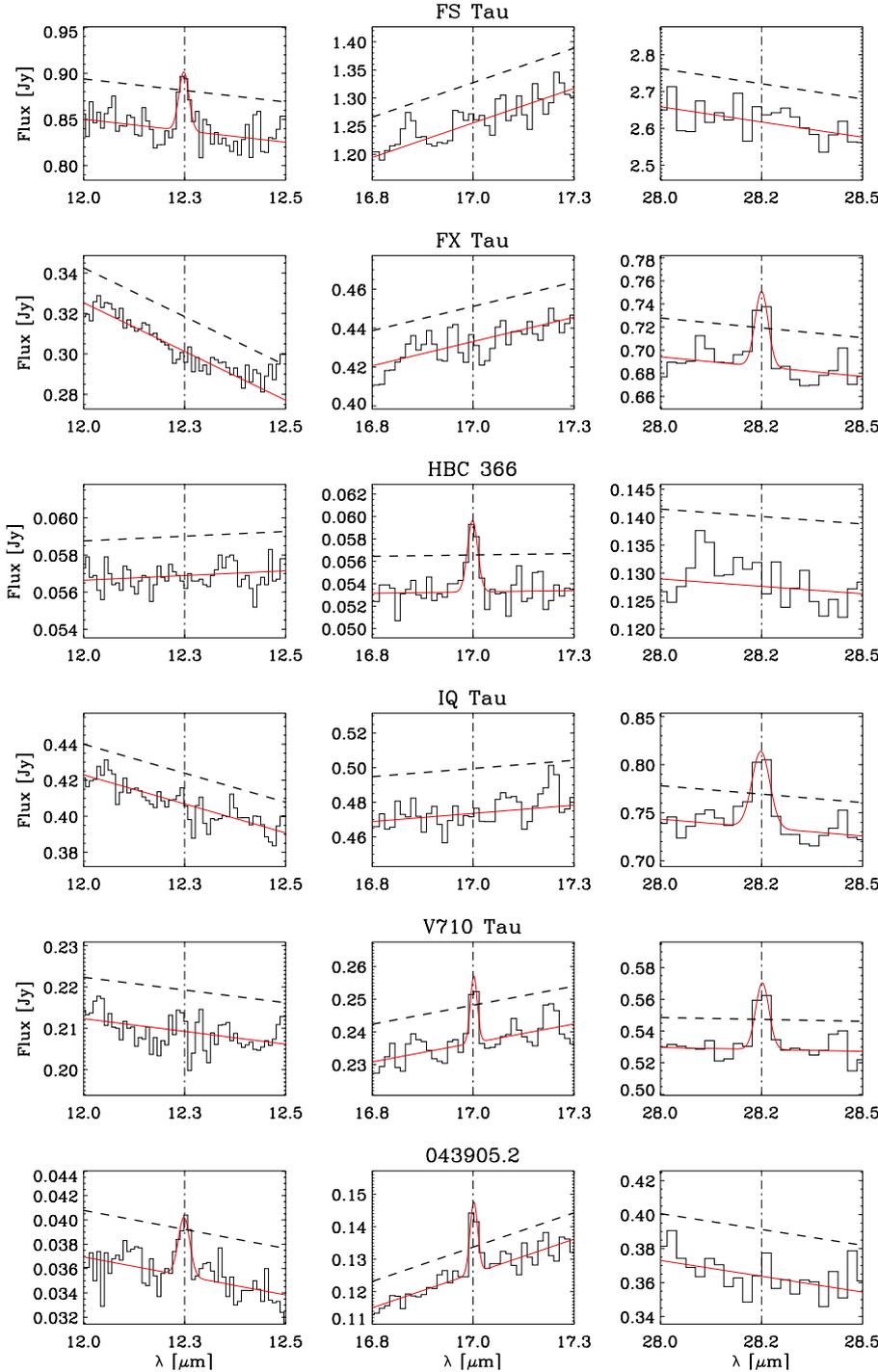


Fig. 2. The plotted spectra show the regions around the H_2 lines at 12.28, 17.03 and 28.22 μm for the sources where we have detected at least one of the lines.

presented in [Pascucci et al. \(2007\)](#) included six targets showing evidence of optically thick disks. The [Ne II] line was detected in four of them. Moreover, [Pascucci et al. \(2007\)](#) reported a correlation between the luminosity of the [Ne II] line and the stellar X-ray luminosity. [Lahuis et al. \(2007\)](#) detected [Ne II] in 15 out of 76 low-mass pre-main sequence stars from different star forming regions, and also reported one detection of [Ne III]. Those authors concluded that the [Ne II] detections and the [Ne II]/[Ne III] line flux ratio are consistent with excitation through X-ray heating of the circumstellar disk from the central star, although excitation through EUV might contribute as well. In addition, [Flaccomio et al. \(2009\)](#) obtained *Spitzer* IRS spectra

for 28 young stellar objects in the Ophiuchus region, obtaining 10 detections of [Ne II] and one detection of [Ne III]. They did not find a correlation between the luminosity of the [Ne II] line and the stellar X-ray luminosity. Instead, the luminosity of the [Ne II] line seemed to be correlated with disk mass accretion rate, consistent with [Espaillat et al. \(2007\)](#). [Güdel et al. \(2010\)](#) based their study on 92 sources observed with *Spitzer* IRS, obtaining detections of [Ne II] in 58 of them. A weak correlation was found between the luminosity of the [Ne II] line and the stellar X-ray luminosity, whereas a good correlation was also found between $L_{[\text{Ne II}]}$ and jet/outflow parameters, such as the luminosity of the [O I] line.

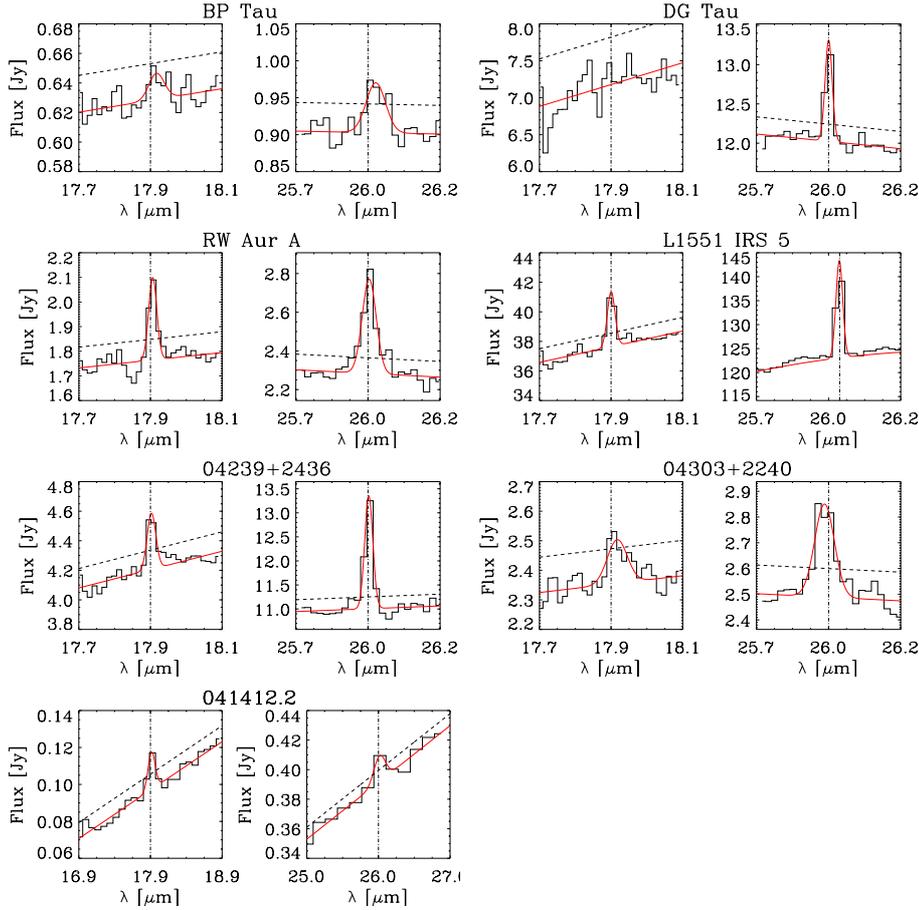


Fig. 3. Sources with [Fe II] detections at 17.94 and 25.99 μm . We recall that the spectrum of SST 041412.2+280837 was obtained using the low resolution modules (SL and LL).

We have detected the [Ne II] line in 18 objects, representing 28% of our sample. This detection rate is similar to the results of Flaccomio et al. (2009). That study was based only in the Ophiuchus region, obtaining a detection rate of 36%. Lahuis et al. (2007) and Güdel et al. (2010) used a mixed sample with stars from different star-forming regions, obtaining ~ 20 and $\sim 60\%$ detections respectively. The sample from Güdel et al. (2010) did not contain any Class I or Herbig AeBe stars. Furthermore, their data reduction differed from ours, as they used the method introduced in Lahuis et al. (2010) in which the compact and extended component of the line emission are separated during the spectral extraction procedure.

Among the objects in our sample with [Ne II] detections, the distribution in terms of evolutionary state is the following: 11% of the [Ne II] detections correspond to Class I objects, 67% of the detections correspond to pure Class II, and 22% of the detections correspond to objects classified as Class I and II in the literature.

The highest luminosity of the [Ne II] line in our study came from T Tau, which is also one of the most luminous young star in the X-rays ($L_X = 8 \times 10^{30} \text{ erg s}^{-1}$). This system is known to drive an outflow responsible for the [Ne II] emission (Van Boekel et al. 2009). L1551 IRS 5 has also a [Ne II] luminosity comparable to that of T Tau, of order $10^{30} \text{ erg s}^{-1}$. Previous studies in the optical and near-IR have shown that L1551 IRS 5 drives a strong molecular outflow (see Appendix A), so there is a suggestion in this case that the [Ne II] could be produced in an outflow/jet environment. We note that its X-ray luminosity is not well constrained, given that the source is very absorbed (Bally et al. 2003); the value available is associated with the jet. For this reason, in the

correlation tests presented in Sect. 5.1.2, we do not consider the X-ray luminosity of L1551 IRS 5.

Other sources with [Ne II] detections known to drive outflows are AA Tau, DG Tau, DG Tau B, UY Aur, and CoKu Tau 1. In general, these sources show high [Ne II] luminosities ($> 10^{29} \text{ erg s}^{-1}$) except for AA Tau ($L_{[\text{Ne II}]} \sim 3 \times 10^{28} \text{ erg s}^{-1}$). We might conclude that for high [Ne II] luminosities, the heating of gas and excitation of neon is caused by the outflow activity. This trend is actually observed in the study of Güdel et al. (2010); sources known to drive jets or outflows show an excess of [Ne II] emission when compared with sources that do not present jets.

To check the suggested difference between sources driving jets and sources without jets, we have divided our sample into sources showing jet emission (typically forbidden line emission in the optical and near-IR) and sources without known jets. Using the Kaplan-Meier (K-M) estimator available within the ASURV⁵ package (Feigelson & Nelson 1985; Isobe et al. 1986), we have extracted the cumulative distribution of [Ne II] luminosities for these two groups. The K-M estimator is the basis of survival analysis. It is a generalized maximum-likelihood estimator for the population from which the sample is drawn. Although it was first used in right censoring problems (i.e., lower limits), it has been adapted for application to left censoring problems (i.e., upper limits). The basic concept underlying the K-M estimator is the survivor function, which is defined as the probability of a value being larger than or smaller than the upper or lower limit, depending on the type of censoring. The survivor

⁵ Available from <http://astrostatistics.psu.edu/statcodes/asurv>

Table 3. Results from survival analysis correlation tests for the luminosity of the [Ne II] line.

All infrared classes						
Variable	N° Points	N° Upper limits		Probability and correlation parameter		
		Total	[NeII]	Cox	Kendall	Spearman
M_*	53	38	38	0.92 (0.010)	0.62 (0.49)	0.55 (0.084)
L_*	57	42	42	0.56 (0.34)	0.82 (0.23)	0.93 (0.012)
L_X	47	35	33	...	0.09 (1.7)	0.18 (0.18)
L_{IR}	64	46	46	0.04 (4.4)	0.01 (2.4)	0.01 (0.31)
$EW(\text{H}\alpha)$	56	42	42	0.04 (4.3)	0.03 (2.1)	0.01 (0.3)
\dot{M}_{acc}	47	36	32	...	0.62 (0.50)	0.37 (0.13)
$\dot{M}_{\text{acc-high}}$	27	21	20	...	0.13 (1.5)	0.04 (0.39)
$\dot{M}_{\text{acc-low}}$	19	15	12	...	0.20 (-1.3)	0.55 (-0.14)
$L_{[\text{OI}]}$	27	21	16	...	0.02 (2.2)	0.02 (0.46)
$L_X L_{[\text{OI}]}$	22	17	12	...	0.06 (1.9)	0.03 (0.46)
$L_X \dot{M}_{\text{wind}}$	36	28	23	...	0.67 (0.42)	1.00 (0.00)
$L_{[\text{FeII}]}$	64	61	46	...	0.09 (1.7)	0.10 (0.21)
Class II						
M_*	40	26	26	0.10 (2.6)	0.46 (-0.74)	0.27 (-0.17)
L_*	40	26	26	0.94 (0.005)	0.85 (0.19)	0.88 (0.02)
L_X	35	24	22	...	0.04 (2.0)	0.05 (0.34)
L_{IR}	44	28	28	0.03 (4.7)	0.004 (2.8)	0.01 (0.39)
$EW(\text{H}\alpha)$	41	28	28	0.26 (1.3)	0.29 (1.1)	0.23 (0.19)
\dot{M}_{acc}	36	26	24	...	0.90 (0.12)	0.99 (-0.002)
$\dot{M}_{\text{acc-high}}$	19	14	13	...	0.05 (1.9)	0.04 (0.49)
$\dot{M}_{\text{acc-low}}$	17	12	10	...	0.02 (-2.2)	0.05 (-0.49)
$L_{[\text{OI}]}$	24	18	18	...	0.02 (2.3)	0.03 (0.45)
$L_X L_{[\text{OI}]}$	19	14	13	...	0.10 (1.9)	0.09 (0.44)
$L_X \dot{M}_{\text{wind}}$	30	22	19	...	0.61 (-0.51)	0.63 (-0.90)
$L_{[\text{FeII}]}$	44	42	28	...	0.62 (0.50)	0.41 (0.12)

Notes. The first column shows the variable for which the correlation is tested, the second column shows the total number of points considered, the third column shows the total number of upper limits used (dependent and independent variable), the fourth column shows the number of upper limits for the luminosity of the [Ne II] line (dependent variable), and the probabilities (P) of a correlation achieved by chance are presented in Cols. 5–7, where $P \leq 5\%$ are shown in bold fonts. The value of the correlation parameters are shown in parentheses. We have defined $\dot{M}_{\text{acc-high}}$ and $\dot{M}_{\text{acc-low}}$ as follows: $\dot{M}_{\text{acc-high}} \geq 10^{-8} M_{\odot} \text{ yr}^{-1}$ and $\dot{M}_{\text{acc-low}} < 10^{-8} M_{\odot} \text{ yr}^{-1}$. The Cox method cannot be used when upper limits are present in both of the tested variables. See text for more details.

function is a non increasing step function which only jumps at uncensored values. When the sample is large, it approaches the true survivor function of the population. For details we refer to Feigelson & Nelson (1985).

The sub-sample of sources without jets has 38 objects; among them, 9 have [Ne II] detections. The mean of the distribution is $L_{[\text{Ne II}]} = (0.8^{+0.5}_{-0.3}) \times 10^{28} \text{ erg s}^{-1}$. The sub-sample of sources with jets has 26 objects; among them we have 9 detections of the [Ne II] line. The cumulative distribution of [Ne II] luminosities is extended towards higher values, with a mean of $L_{[\text{Ne II}]} = (1.6^{+0.9}_{-0.6}) \times 10^{28} \text{ erg s}^{-1}$, almost twice the mean value of the sub-sample of sources with no jets. However, both distributions overlap. Moreover, the median of the two sub-samples are very close, with $L_{[\text{Ne II}]} = 4 \times 10^{28} \text{ erg s}^{-1}$ and $L_{[\text{Ne II}]} = 5 \times 10^{28} \text{ erg s}^{-1}$ for jet and non-jet sources, respectively. Thus, we cannot distinguish between the two populations, in contrast with Güdel et al. (2010). This difference might be due to the sample selection criteria. In fact, Güdel et al. (2010) chose a sample composed only of Class II objects. Furthermore, their sub-sample of jet-driving sources, unlike ours, did not contain objects with indirect evidence of jets (i.e., high-velocity emission line components).

The number of detections of [Ne II] in our sample (18 objects) has prompted us to search for possible correlations or trends between the luminosity of the line and parameters related to the star-disk system, and outflow/jet indicators. The parameters studied are: stellar mass (M_*), stellar luminosity (L_{star}), X-ray luminosity (L_X), $\text{H}\alpha$ equivalent width, mass accretion rate

(\dot{M}_{acc}), wind mass loss rate (\dot{M}_{wind}), luminosity of the [O I] line in the optical ($L_{[\text{OI}]}$), and luminosity of the [Fe II] line at $26 \mu\text{m}$ ($L_{[\text{FeII}]}$). If correlations exist, they would give new hints on the neon excitation mechanism.

Due to the high number of upper limits for the luminosity of the [Ne II] line in our sample, we have tested the correlations using survival analysis implemented in the ASURV package. We have used 3 statistical tests: Cox's proportional hazard model, generalized Kendall's τ rank correlation, and Spearman's ρ . The Kendall and Spearman methods allow the use of upper limits in one or both of the variables tested. Therefore, we have included upper limits for all the parameters tested. In contrast, the Cox hazard model allows only upper limits in the dependent variable. The Spearman method is not reliable for data sets smaller than 30 data points. Nevertheless, we report for completeness the probabilities obtained with this method even when the number of data points is smaller than 30. Calculations were done using the logarithm of values. Table 3 summarizes our results, in which we have included the values for the correlation parameters between parentheses for completeness. In the case of the Cox hazard method, the parameter reported is the χ^2 of the regression fit. For the Kendall method, the value reported is the ratio between the correlation parameter τ and its standard deviation σ_{τ} , a function of the correlation. For the Spearman's ρ method, the parameter reported is the estimated correlation.

We have tested the correlations using the whole sample (i.e., all infrared classes). In addition, we have created a sub-sample with the Class II sources, given that the majority of the

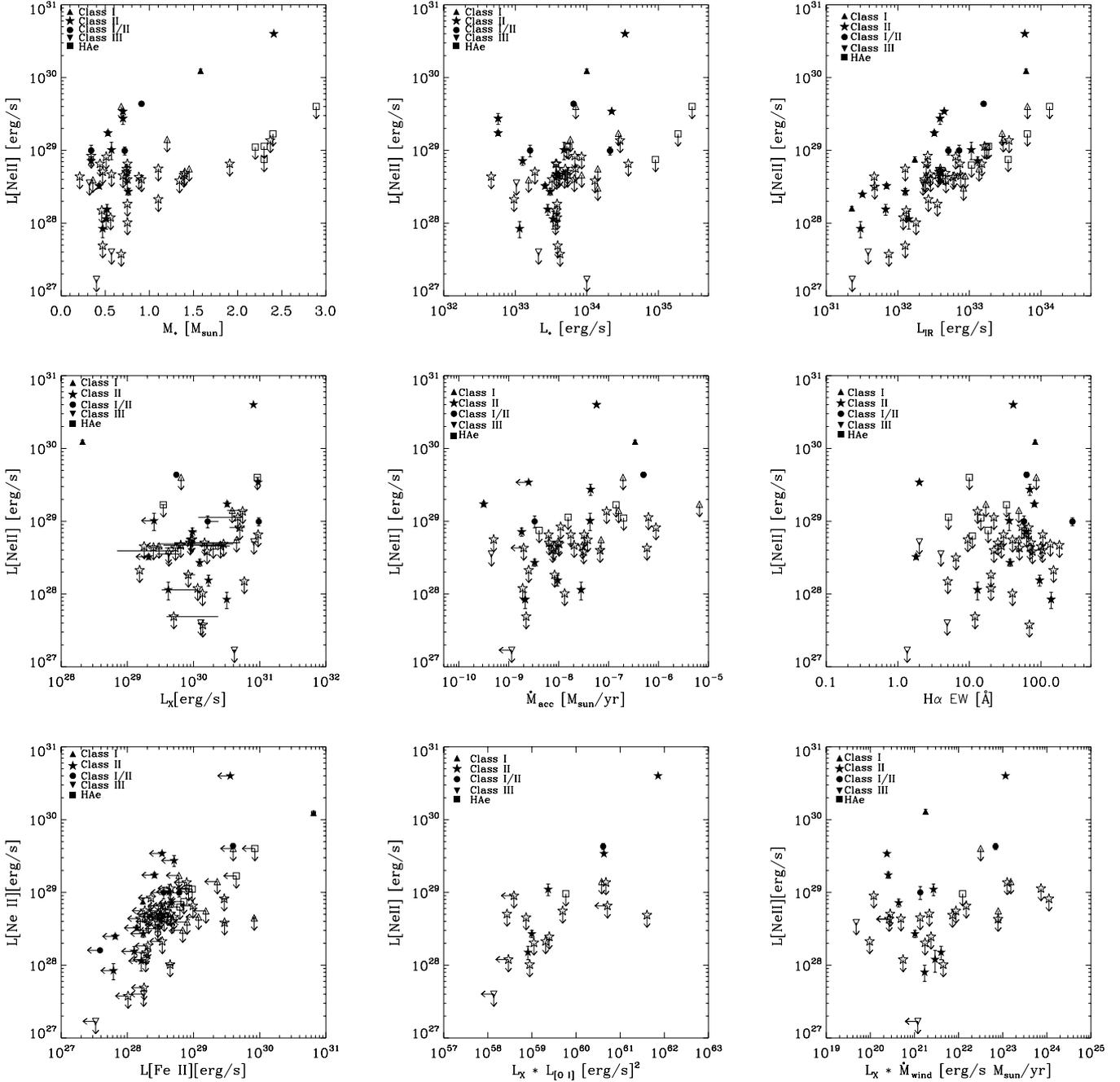


Fig. 4. [Ne II] luminosity versus the different parameters for which correlations have been tested. The different symbols indicate Class I, II, III or Herbig stars as indicated in the legend. Filled symbols indicate [Ne II] detections while empty symbols with arrows indicate upper limits.

[Ne II] detections are from these objects. In Figs. 4 and 5 we present $L_{[\text{Ne II}]}$ as a function of the different parameters studied for the complete sample and for the sub-sample of Class II objects separately. Although we find significant correlations (at a 5% threshold) based on several methods included in the ASURV package, there is a certain chance that the conclusion of correlation is incorrect. To this end, we have used a bootstrap method to determine the chance of finding a correlation by chance. Such a method implies the creation of a sampling distribution from the original data. In practice, we created 5×10^4 random samples in each testing case by keeping the independent variable unchanged and permuting each time the dependent variable ($L_{[\text{Ne II}]}$ in this case) in order to break the links between each pair of points, keeping the number of detections and upper limits unchanged.

We have applied the correlation tests using the ASURV package for each sample. Then we have constructed the distribution of the correlation parameters (i.e., τ for Kendall, ρ for Spearman, and χ^2 for Cox), and we have looked into the percentage of cases with correlation parameters larger than the ones obtained for the original sample. In all correlations tested, the probabilities were generally less than 4%, thus confirming the correlations presented in Table 3.

5.1.1. Correlations with stellar properties

We have tested correlations with the stellar mass and luminosity. In a previous study, Flaccomio et al. (2009) found that the stellar mass does not influence the [Ne II] luminosity. In our sample,

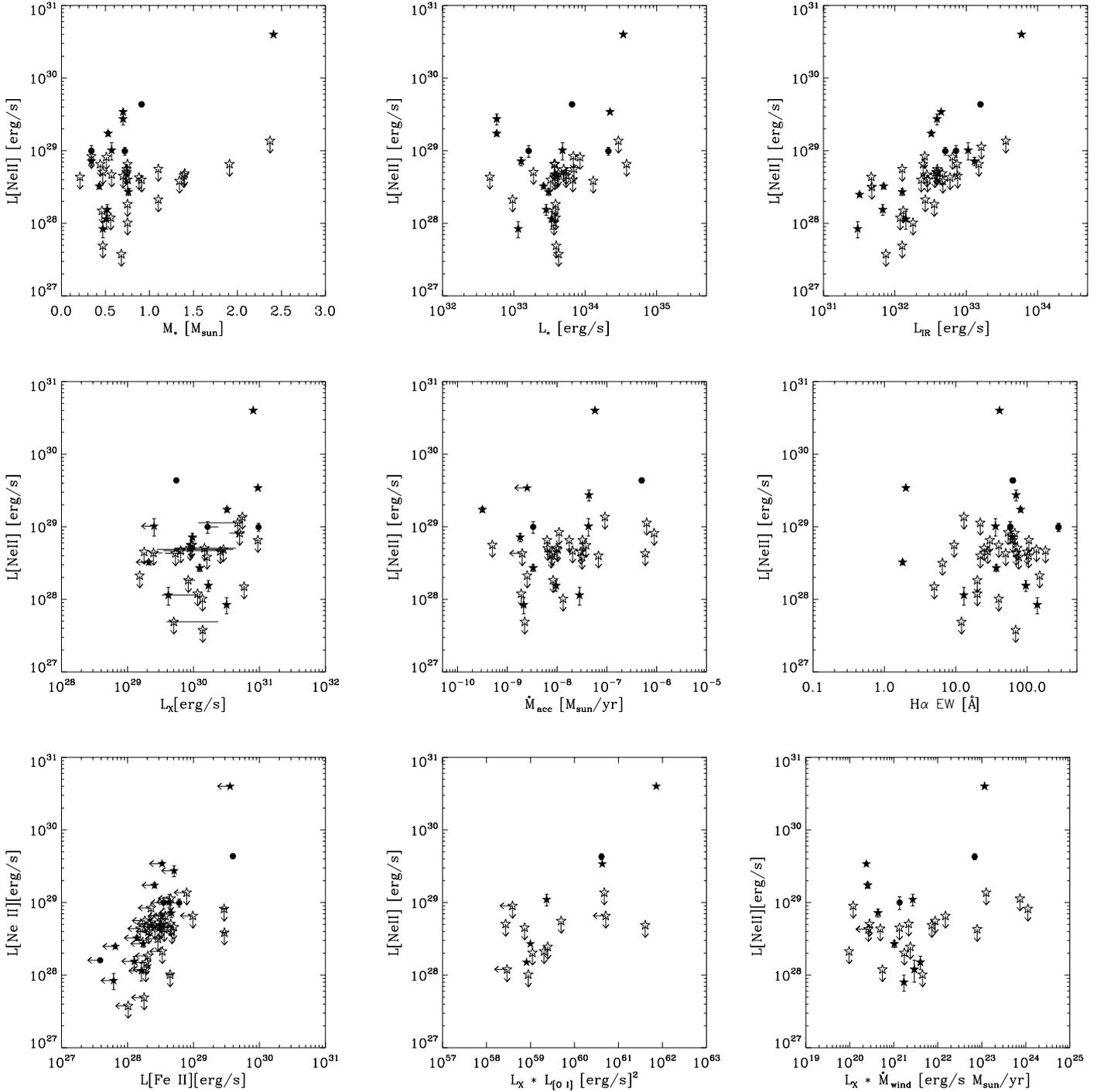


Fig. 5. Same as Fig. 4 for Class II (stars) and sources classified as I and II in the literature (circles).

the detected [Ne II] lines are concentrated in sources at the low end of the mass range (with stellar masses typically $< 1 M_{\odot}$); only two objects with $M_* \geq 1 M_{\odot}$ show [Ne II] emission. In [Flaccomio et al. \(2009\)](#), the detected [Ne II] lines are spread across a larger range of stellar masses. Despite the different mass ranges probed by the two samples, the result is the same: $L_{[\text{Ne II}]}$ is not dependent on the stellar mass. Since the stellar mass and luminosity are related through the relation $L_* \propto M_*^2$ ([Hartmann 2009](#)), we expect to obtain a similar lack of correlation between $L_{[\text{Ne II}]}$ and the stellar luminosity, which is indeed confirmed.

5.1.2. Correlations with high-energy irradiation

The stellar X-ray luminosity is one of the key parameters driving the [Ne II] emission proposed in several theoretical studies, e.g., [Glassgold et al. \(2007\)](#). The first observational study suggesting such a correlation was [Pascucci et al. \(2007\)](#), although this result was based on a sample of just four objects, spanning a small range in L_X . [Espaillat et al. \(2007\)](#) increased the sample to seven objects but did not find any evident correlation between $L_{[\text{Ne II}]}$ and L_X . The same result was obtained recently by

Flaccomio et al. (2009) based on a larger sample and spanning a larger range in L_X . Nevertheless, Güdel et al. (2010) have found a statistically significant dependence on the stellar L_X for their sample confined to Class II objects and spanning a broad range of L_X .

Our sample spans a range of X-ray luminosities between 10^{29} and $\sim 10^{31}$ erg s $^{-1}$, similar to the sample presented by Güdel et al. (2010). The correlation between $L_{[\text{Ne II}]}$ and L_X is non-existent for the complete sample (all infrared classes), i.e., consistent with Flaccomio et al. (2009). But we observe a correlation with L_X for the sub-sample of Class II objects, consistent with Güdel et al. (2010). These results are not contradictory: in fact, the difference between the previous studies lies in the sample selection. Flaccomio et al. (2009) used a mixed sample, including Class I to Class III objects, similar to our sample (although we have included also Herbig AeBe stars), whereas the sample used in Güdel et al. (2010) contains only Class II objects. Therefore, we do expect to find a lack of correlation like in Flaccomio et al. (2009) when using our whole sample, and the same result as in Güdel et al. (2010) when restricting the sample to Class II objects.

5.1.3. Correlations with accretion parameters

The mass accretion rate (\dot{M}_{acc}) is an indirect measurement of the evolutionary state of the sources. We expect \dot{M}_{acc} to decrease as the star evolves toward the main sequence. Espaillat et al. (2007) and Flaccomio et al. (2009) found a positive correlation between $L_{[\text{Ne II}]}$ and \dot{M}_{acc} . Hollenbach & Gorti (2009) studied the irradiation of the disk by high-energy photons (EUV and soft X-rays). They showed that in order for the high-energy photons to penetrate the strong protostellar winds and illuminate the disk beyond 1 AU, low wind mass-loss rates were needed ($\dot{M}_{\text{wind}} < 10^{-9} M_{\odot} \text{ yr}^{-1}$). This critical value was translated in Hollenbach & Gorti (2009) into a critical disk mass accretion rate of $\dot{M}_{\text{acc}} < 10^{-8} M_{\odot} \text{ yr}^{-1}$. Sources with low mass accretion rates would be more suitable to study the effects of high-energy photons. Following this idea, we have split our sample between high and low accretors, setting the limit at $\dot{M}_{\text{acc}} = 10^{-8} M_{\odot} \text{ yr}^{-1}$, as proposed in Hollenbach & Gorti (2009). We have, therefore, defined two sub-samples: $\dot{M}_{\text{acc-high}} \geq 10^{-8} M_{\odot} \text{ yr}^{-1}$, and $\dot{M}_{\text{acc-low}} < 10^{-8} M_{\odot} \text{ yr}^{-1}$. For the sub-sample of Class II objects, we find a positive correlation between $L_{[\text{Ne II}]}$ and \dot{M}_{acc} for the group $\dot{M}_{\text{acc-high}}$ and a negative correlation with $L_{[\text{Ne II}]}$ for the group $\dot{M}_{\text{acc-low}}$. This result is not found when considering the sub-sample of Class II objects with no separation in the mass accretion rate range. We observe a difference between low and high accretors; the anti-correlation found for low-accretors is statistically more significant. But we are cautious about this result: this correlation has been discussed in Güdel et al. (2010) in the sense that jet driving sources are younger and, in general, have higher mass accretion rates than sources without jets. They argued that the segregation between both samples might produce an apparent correlation. Indeed, in our study, the sub-sample of low-accretors ($\dot{M}_{\text{acc-low}}$) is dominated by non-jet sources; 18 out of the 21 sources in this sub-sample do not have evidence of jets. On the other hand, in the sub-sample of high accretors ($\dot{M}_{\text{acc-high}}$), which accounts for 27 sources, 16 present observational evidence of jets/outflows.

The $\text{H}\alpha$ line (seen usually in emission in young stars) is another tracer of accretion. Its equivalent width is commonly used to differentiate between accreting and non-accreting stars. Typically, $EW(\text{H}\alpha) > 10 \text{ \AA}$ for accreting stars, and

$EW(\text{H}\alpha) < 10 \text{ \AA}$ for non-accreting stars (Martin 1997). The result for the test between $L_{[\text{Ne II}]}$ and $EW(\text{H}\alpha)$ is not clear in our sample. Although we found a correlation for the whole sample (i.e., all infrared classes), this result is not repeated for the sub-sample of Class II objects.

Finally, we have tested a correlation of $L_{[\text{Ne II}]}$ with the luminosity in the mid-IR, in order to check if bright emission lines are related to a bright continuum emission. The mid-IR luminosity has been obtained for each object by integrating the continuum flux in the SH range (9 to 19 μm). Since in most cases we do not have background observations, we have over-plotted the SH and SL spectra and checked that the continuum emission is dominated by the source and not by the background. Given the small aperture and wavelength coverage of the SH module, the background contribution is expected to be small (e.g., Furlan et al. 2006). We have obtained a correlation between $L_{[\text{Ne II}]}$ and L_{IR} . This result is confirmed by the bootstrap technique. The origin of this correlation is not clear, although not likely to be the product of bias in the observations. The rms of the continuum increases with L_{IR} , but the signal-to-noise ratio for the detected [Ne II] lines is distributed uniformly over the L_{IR} spanned by our sample, with no preference for either high or low values of the signal-to-noise ratio. This correlation was also found by Lahuis et al. (2010) in a sample of Class I sources.

5.1.4. Correlations with jet/outflow parameters

Forbidden atomic lines of [O I] (6300 \AA) and [Fe II] (1.64 μm) are used to trace the outflow activity in young stars (e.g., Cabrit et al. 1990; Agra-Amboage et al. 2009). Güdel et al. (2010) investigated this correlation to test the relation between [Ne II] emission and outflow activity. Shang et al. (2010) modeled the flux and shape of the [Ne II] and [O I] lines and found a clear correlation between both lines, indicating that neon would also trace jets.

We have tested the correlation between the luminosity of the [Ne II] line and the luminosity of the [O I] line at 6300 \AA compiled from the literature, obtaining a positive result. We have also tested a correlation with the product of L_X and $L_{[\text{O I}]}$, as proposed in Güdel et al. (2010), but a correlation was not found. We have further used the luminosities of the [Fe II] line at 25.99 μm obtained in our sample to study the dependence of [Ne II] on this tracer, and no correlation was found. We emphasize, however, that the tests were done with a large number of upper limits due to the lack of detections in the tested parameters.

Three objects with detected [Ne II] emission also show [Fe II]: L1551 IRS 5 (Class I), DG Tau (Class I/II), and SST 041412.2+280837. Since [Fe II] is a known tracer of shocks, the origin of the [Ne II] emission in these particular cases might be explained by the shock scenario. However, this is not clear for the rest of the objects with [Ne II] emission for which there is no detected emission from [Fe II].

5.1.5. Correlations with $L_X \times \dot{M}_{\text{wind}}$

According to the model presented in Shang et al. (2010), [Ne II] is produced in jets from young stars. Based on the X-wind model theory, they obtained line profiles for [Ne II], [Ne III], and [O I]. They tested the dependency of the line luminosities on critical parameters of this model such as L_X and the wind mass loss rate (\dot{M}_{wind}). The best correlation was found for the product of both parameters $L_X \times \dot{M}_{\text{wind}}$. The relation between wind mass loss rate and accretion rate is $\dot{M}_{\text{wind}} = f \times \dot{M}_{\text{acc}}$, with $f \approx 0.25$

Table 4. Correlations tested for the luminosity of the [Fe II] line.

All infrared classes						
Var.	N° Points	N° Upper L.		Probability		
		Total	[FeII]	Cox	Kendall	Spearman
M_*	53	48	48	0.55 (0.45)	0.17 (1.3)	0.001 (0.44)
L_*	57	51	51	0.77 (0.08)	0.14 (1.5)	0.001 (0.46)
L_X	47	43	40	...	0.66 (0.44)	0.25 (0.17)
L_{IR}	64	57	57	0.81 (0.05)	0.62 (0.49)	0.49 (0.09)
\dot{M}_{acc}	47	42	39	...	0.01 (2.6)	0.0004 (0.52)
$L_{[OI]}$	27	24	24	...	0.16 (1.4)	0.09 (0.33)

Notes. Probabilities (P) of a correlation achieved by chance are presented in Cols. 5–7, where $P \leq 5\%$ are shown in bold fonts. The value of the correlation parameters are shown between parentheses.

(Shang et al. 2004). However, in our sample, no such correlation is found.

5.2. [Fe II] emission

[Fe II] is a tracer of shock activity typically observed in the near-IR at 1.53 and 1.64 μm (Davis et al. 2003; Hartigan & Hillenbrand 2009). Its excitation mechanism can also be stimulated by the irradiation of protoplanetary disks by high-energy photons (Gorti & Hollenbach 2008). We have detected [Fe II] emission in 7 sources. In five cases, we have detected both iron lines (at 17.93 and 25.99 μm). Typically, whenever both lines have been detected, the luminosity of the [Fe II] line at 25.99 μm is roughly a factor of 3 higher than the luminosity of the line at 17.93 μm . This result is consistent with Hollenbach & McKee (1989), where in a shock environment the [Fe II] line at 25.99 μm is a more effective coolant than the line at 17.93 μm . Five out of the seven sources with detected [Fe II] lines (L1551 IRS5, DG Tau, RW Aur, IRAS 04239+2436, and IRAS 04303+2240) have previous indications of jet activity based on near-IR and/or optical observations (see Appendix A). We have tested the correlation between the luminosity of the [Fe II] line (25.99 μm) obtained in this study with different parameters: stellar mass, luminosity, X-ray luminosity, infrared luminosity, mass accretion rate, and luminosity of the [O I] line (see Table 4). In some cases (mass and luminosity) the results from the different correlation tests for $L_{[Fe II]}$ give largely different results. We remind that mostly upper limits are available for the [Fe II] line; indeed, we have only 7 detections for 53 to 64 points, depending on the parameter tested. Furthermore, most correlations were found with the Spearman method, but the Kendall and Cox methods are in general more robust and no correlations were found with these methods. In summary, except for mass accretion rate, no correlation is found for the [Fe II] line luminosity.

5.3. H_2 emission

Three pure rotational lines of H_2 are covered by the *Spitzer* IRS: 12.28 μm (S2), 17.03 μm (S1), and 28.22 μm (S0). Detections of H_2 in the near-IR (S(1) transition at 2.12 μm) were reported in a handful of objects (e.g., Itoh et al. 2003; Bary et al. 2002, 2008b; Carmona et al. 2008a; Panić et al. 2009). In addition to H_2 disk emission, H_2 emission may also originate in shock environments; in particular the near-IR lines and the mid-IR line at 12.28 μm are associated with shock densities of $n \sim 10^{4-5} \text{ cm}^{-3}$ (Hollenbach et al. 1989; Hollenbach & McKee 1989).

We have detected H_2 in six sources: one of the newly identified Taurus members SST 043905.2+233745, which is a

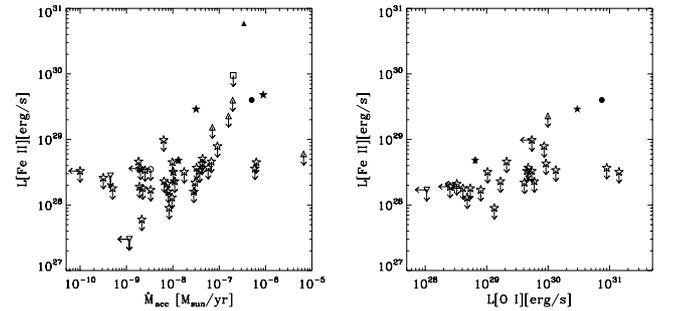


Fig. 6. [Fe II] luminosity versus mass accretion rate and luminosity of the [O I] line. The symbol convention is the same used in Fig. 4.

Class I source, the Class II sources FS Tau A, FX Tau, IQ Tau, V710 Tau, and the Class III HBC 366. FX Tau and V710 Tau are both binary systems composed of a classical and a weak-lined T Tauri star. These are the first detections of H_2 reported toward these sources.

We have tested correlations between the total luminosity of the H_2 line with the following parameters: L_* , M_* , L_{IR} , L_X , \dot{M}_{acc} , and $L_{[OI]}$. The total luminosity of the H_2 line was calculated by adding the luminosities of the three lines accessible in the IRS wavelength range. Our analysis does not reveal any correlation with the above parameters. Although correlations with H_2 were found for both the mid-IR and [O I] luminosities, this was only the case for one test method out of three. We prefer, thus, to conclude that no correlation was found.

6. Summary and conclusions

We have searched for gas in a large sample of pre-main sequence stars (64 objects) spanning a broad range of evolutionary stages. Our sample is dominated by Class II objects but it also includes Class I, III, and Herbig AeBe stars. We have detected 36 gas emission lines; in particular, we have obtained 18 detections of the [Ne II] line. Nine out of the 18 [Ne II] detections are reported here for the first time. For 4 targets there are previously reported detections of [Fe II] based on near-IR observations, which we complement with the mid-IR information. The six sources in which we detect H_2 emission do not have previously reported detections.

The [Ne II] detections are concentrated towards Class II sources; we have not detected [Ne II] lines towards Class III or Herbig AeBe stars. Two Class I sources show [Ne II] line emission. These sources have higher accretion rates, and are more likely to show jet/outflow phenomena, but they

Table 5. Correlations tested for the luminosity of the H₂ line.

Var.	All infrared classes					
	N° Points	N° Upper L.			Probability	
		Total	H ₂	Cox	Kendall	Spearman
M_*	53	48	48	0.97 (0.001)	0.79 (−0.26)	0.57 (−0.078)
L_*	57	52	52	0.17 (1.9)	0.61 (−0.51)	0.41 (−0.11)
L_X	47	42	39	...	0.86 (−0.17)	0.54 (−0.095)
L_{IR}	64	58	58	0.03 (4.5)	0.09 (1.7)	0.09 (0.01)
M_{acc}	47	44	40	...	0.22 (−1.2)	0.43 (−0.15)
$L_{[O I]}$	27	26	26	...	0.99 (0.00)	0.04 (0.40)

Notes. Probabilities (P) of a correlation achieved by chance are presented in columns 5, 6, and 7, where $P \leq 5\%$ are shown in bold fonts. The value of the correlation parameters are shown between parentheses.

also are more embedded in the circumstellar material. If neon line emission is produced by jets, we would expect to have a larger detection rate for Class I objects. We have attempted to disentangle sources with and without jets, but we cannot statistically separate the two populations.

We have tested correlations between the [Ne II] line with parameters probing different stellar-disk properties; high-energy photon irradiation, accretion, and jets/outflows. We have found that the luminosity of the [Ne II] line is dependent on the continuum mid-IR luminosity. The dependence with L_X obtained in Güdel et al. (2010) is confirmed for the sub-sample of Class II sources. A good correlation is also found for the $L_{[O I]}$, echoing the finding by Güdel et al. (2010) and clearly supporting a trend with jet parameters.

In the jet/outflow scenario proposed by Hollenbach & McKee (1989), the [Ne II] line is proposed as an effective coolant. For high density shocks ($n \sim 10^{5-6} \text{ cm}^{-3}$), the [Fe II] line (25.99 μm) is more effective, with expected luminosities up to 2 orders of magnitude higher. Only in lower density environments ($n \sim 10^{3-4} \text{ cm}^{-3}$) and high velocities ($>80 \text{ km s}^{-1}$) the luminosity of the [Ne II] surpasses the luminosity of [Fe II]. Among the sources with [Ne II] detections, three also show [Fe II] emission: DG Tau, L1551 IRS 5 (both known to drive jets), and SST 041412+289837 (for which little information is available in the literature). In these cases, a plausible explanation for the [Ne II] emission is outflow activity. The question arises for sources with [Ne II] detections that do not show [Fe II] emission: if [Ne II] is excited in shocks produced by outflows, what is the explanation for the absence of [Fe II] emission? This might be due to different reasons; perhaps lower signal-to-noise was obtained in the [Fe II] wavelength range, or the conditions in the emitting region are such that the neon line is brighter, or simply for those sources, neon is not excited by shocks, but its line emission originates, e.g., from the disk atmosphere.

The [Fe II] emission line has been detected in a smaller number of sources, in comparison to the [Ne II] line. The statistical analysis has shown that the luminosity of [Fe II] is dependent on the mass accretion rate. This indicates a link with outflow phenomena, since the outflow rate is proportional to the mass accretion rate. On the other hand, no correlation between the [Fe II] luminosity and the luminosity of the outflow-tracing [O I] (6300 Å) line was found.

Furthermore, *Spitzer* IRS can access three H₂ emission lines. We have obtained few detections, and, despite the large number of upper limits, we have tested different correlations. Our results do not show any dependency on the parameters tested.

Our results show a trend towards accretion-related phenomena as the origin for the forbidden lines. Several mechanisms could be playing a role in the heating and excitation mechanism

of gas in protoplanetary disks. A sample with several Class I, and Class III objects would be more suitable to study possible trends with evolutionary state. But no matter what mechanism produces the neon emission, it is likely originating very close to the star. High spatial and spectral resolution observations are crucial in order to determine its origin.

Acknowledgements. This work is based on observations made with the *Spitzer* Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. C.B.-S., M.A., and A.C. acknowledge support from the Swiss National Science Foundation (grants PP002-110504 and PP00P2-130188). We thank Fred Lahuis for useful discussion on the data reduction, Maria Suveges for insightful discussion on bootstrap techniques, and an anonymous referee for suggestions and comments that improved the manuscript.

References

- Agra-Amboage, V., Dougados, C., Cabrit, S., Garcia, P. J. V., & Ferruit, P. 2009, *A&A*, 493, 1029
- Alexander, R. D. 2008, *MNRAS*, 391, L64
- Andrews, S. M., & Williams, J. P. 2005, *ApJ*, 631, 1134
- Bacciotti, F., Ray, T. P., Mundt, R., Eisloffel, J., & Solf, J. 2002, *ApJ*, 576, 222
- Bally, J., Feigelson, E., & Reipurth, B. 2003, *ApJ*, 584, 843
- Bary, J. S., Weintraub, D. A., & Kastner, J. H. 2002, *ApJ*, 576, L73
- Bary, J. S., Matt, S. P., Skrutskie, M. F., et al. 2008a, *ApJ*, 687, 376
- Bary, J. S., Weintraub, D. A., Shukla, S. J., Leisenring, J. M., & Kastner, J. H. 2008b, *ApJ*, 678, 1088
- Beckwith, S. V. W., Sargent, A. I., Chini, R. S., & Guesten, R. 1990, *AJ*, 99, 924
- Bieging, J. H., & Cohen, M. 1985, *ApJ*, 289, L5
- Bitner, M. A., Richter, M. J., Lacy, J. H., Herczeg, G. J., & Greathouse, T. K. 2008, *ApJ*, 688, 1326
- Blake, G. A., & Boogert, A. C. A. 2004, *ApJ*, 606, L73
- Bouvier, J., Chelli, A., Allain, S., et al. 1999, *A&A*, 349, 619
- Briceño, C., Hartmann, L., Stauffer, J., & Martín, E. 1998, *AJ*, 115, 2074
- Cabrit, S., Edwards, S., Strom, S. E., & Strom, K. M. 1990, *ApJ*, 354, 687
- Calvet, N., D'Alessio, P., Hartmann, L., et al. 2002, *ApJ*, 568, 1008
- Calvet, N., D'Alessio, P., Watson, D. M., et al. 2005, *ApJ*, 630, L185
- Carmona, A., van den Ancker, M. E., Henning, T., et al. 2007, *A&A*, 476, 853
- Carmona, A., van den Ancker, M. E., Henning, T., et al. 2008a, *A&A*, 478, 795
- Carmona, A., van den Ancker, M. E., Henning, T., et al. 2008b, *A&A*, 477, 839
- Carpenter, J. M., Bouwman, J., Mamajek, E. E., et al. 2009, *ApJS*, 181, 197
- Carr, J. S. 1989, *ApJ*, 345, 522
- Carr, J. S., & Najita, J. R. 2008, *Science*, 319, 1504
- Cohen, M., & Kuhl, L. V. 1979, *ApJS*, 41, 743
- Corcoran, M., & Ray, T. P. 1997, *A&A*, 321, 189
- Cox, A. W., Hilton, G. M., Williger, G. M., Grady, C. A., & Woodgate, B. 2005, *BAAS*, 37, 1287
- Davis, C. J., Ray, T. P., Desroches, L., & Aspin, C. 2001, *MNRAS*, 326, 524
- Davis, C. J., Whelan, E., Ray, T. P., & Chrysostomou, A. 2003, *A&A*, 397, 693
- Duchêne, G., Ghez, A. M., McCabe, C., & Weinberger, A. J. 2003, *ApJ*, 592, 288
- Duvert, G., Guilloteau, S., Ménard, F., Simon, M., & Dutrey, A. 2000, *A&A*, 355, 165
- Dyck, H. M., Simon, T., & Zuckerman, B. 1982, *ApJ*, 255, L103
- Eisloffel, J., & Mundt, R. 1998, *AJ*, 115, 1554
- Espaillet, C., Calvet, N., D'Alessio, P., et al. 2007, *ApJ*, 664, L111

- Fedele, D., van den Ancker, M. E., Henning, T., Jayawardhana, R., & Oliveira, J. M. 2010, *A&A*, 510, A72
- Feigelson, E. D., & Nelson, P. I. 1985, *ApJ*, 293, 192
- Flaccomio, E., Stelzer, B., Sciortino, S., et al. 2009, *A&A*, 505, 695
- Forrest, W. J., Sargent, B., Furlan, E., et al. 2004, *ApJS*, 154, 443
- Furlan, E., Hartmann, L., Calvet, N., et al. 2006, *ApJ*, 165, 568
- Furlan, E., McClure, M., Calvet, N., et al. 2008, *ApJS*, 176, 184
- García Lopez, R., Natta, A., Testi, L., & Habart, E. 2006, *A&A*, 459, 837
- Glassgold, A. E., Najita, R., J., & Igea, J. 2007, *ApJ*, 656, 515
- Gorti, U., & Hollenbach, D. 2008, *ApJ*, 683, 287
- Güdel, M., Briggs, K. R., Arzner, K., et al. 2007, *A&A*, 468, 353
- Güdel, M., Lahuis, F., Briggs, K. R., et al. 2010, *A&A*, 519, A113
- Gullbring, E., Hartmann, L., Briceño, C., & Calvet, N. 1998, *ApJ*, 492, 323
- Hartigan, P., & Hillenbrand, L. 2009, *ApJ*, 705, 1388
- Hartigan, P., Edwards, S., & Ghandour, L. 1995, *ApJ*, 452, 736
- Hartigan, P., Edwards, S., & Pierson, R. 2004, *ApJ*, 609, 261
- Hartmann, L. 2008, *Phys. Scr. T*, 130, 014012
- Hartmann, L. 2009, *Accretion Processes in Star Formation*, 2nd ed., ed. A. King, D. Lin, S. Maran, J. Pringle, & M. Ward (Cambridge University Press)
- Herbig, G. H., & Bell, K. R. 1995, *VizieR Online Data Catalog*, 5073
- Hirth, G. A., Mundt, R., & Solf, J. 1997, *A&AS*, 126, 437
- Hollenbach, D., & Gorti, U. 2009, *ApJ*, 703, 1203
- Hollenbach, D., & McKee, C. F. 1989, *ApJ*, 342, 306
- Hollenbach, D. J., Chernoff, D. F., & McKee, C. F. 1989, in *Infrared Spectroscopy in Astronomy*, ed. B. H. Kaldich, ESA SP, 290, 245
- Houck, J. R., Roellig, T. L., van Cleve, J., et al. 2004, *ApJS*, 154, 18
- Ireland, M. J., & Kraus, A. L. 2008, *ApJ*, 678, L59
- Isella, A., Tatulli, E., Natta, A., & Testi, L. 2008, *A&A*, 483, L13
- Isobe, T., Feigelson, E. D., & Nelson, P. I. 1986, *ApJ*, 306, 490
- Itoh, Y., Sugitani, K., Ogura, K., & Tamura, M. 2003, *PASJ*, 55, L77
- Kenyon, S. J., Hartmann, L. W., Strom, K. M., & Strom, S. E. 1990, *AJ*, 99, 869
- Kenyon, S. J., Brown, D. I., Tout, C. A., & Berlind, P. 1998, *AJ*, 115, 2491
- Köhler, R. 2008, *J. Phys. Conf. Ser.*, 131, 012028
- Krist, J. E., Stapelfeldt, K. R., Burrows, C. J., et al. 1998, *ApJ*, 501, 841
- Krist, J. E., Stapelfeldt, K. R., Hester, J. J., et al. 2008, *AJ*, 136, 1980
- Lahuis, F., van Dishoeck, E. F., Blake, G. A., et al. 2007, *ApJ*, 665, 492
- Lahuis, F., van Dishoeck, E. F., Jørgensen, J. K., Blake, G. A., & Evans, N. J. 2010, *A&A*, 519, A3
- Lavalley, C., Cabrit, S., Dougados, C., Ferruit, P., & Bacon, R. 1997, *A&A*, 327, 671
- Lim, J., & Takakuwa, S. 2006, *ApJ*, 653, 425
- Luhman, K. L., Allen, P. R., Espaillat, C., Hartmann, L., & Calvet, N. 2010, *ApJS*, 186, 111
- Maheswar, G., Manoj, P., & Bhatt, H. C. 2002, *A&A*, 387, 1003
- Mann, R. K., & Williams, J. P. 2009, *ApJ*, 694, L36
- Mannings, V., & Sargent, A. I. 1997, *ApJ*, 490, 792
- Manoj, P., Bhatt, H. C., Maheswar, G., & Muneer, S. 2006, *ApJ*, 653, 657
- Martin, E. L. 1997, *A&A*, 321, 492
- Martin, E. L., Rebolo, R., Magazzu, A., & Pavlenko, Y. V. 1994, *A&A*, 282, 503
- Martin-Zaidi, C., Lagage, P., Pantin, E., & Habart, E. 2007, *ApJ*, 666, L117
- McCabe, C., Ghez, A. M., Prato, L., et al. 2006, *ApJ*, 636, 932
- Meyer, M. R. 2009, in *IAU Symp. 258*, ed. E. E. Mamajek, D. R. Soderblom, & R. F. G. Wyse, 111
- Moriarty-Schieven, G. H., Wannier, P. G., Tamura, M., & Keene, J. 1992, *ApJ*, 400, 260
- Mundt, R., & Fried, J. W. 1983, *ApJ*, 274, L83
- Mundt, R., Buehrke, T., Solf, J., Ray, T. P., & Raga, A. C. 1990, *A&A*, 232, 37
- Muzerolle, J., Calvet, N., & Hartmann, L. 2001, *ApJ*, 550, 944
- Najita, J., Carr, J. S., & Mathieu, R. D. 2003, *ApJ*, 589, 931
- Najita, J. R., Doppmann, G. W., Bitner, M. A., et al. 2009, *ApJ*, 697, 957
- Neuhäuser, R., Sterzik, M. F., Schmitt, J. H. M. M., Wichmann, R., & Krautter, J. 1995, *A&A*, 297, 391
- Nguyen, D. C., Jayawardhana, R., van Kerkwijk, M. H., et al. 2009, *ApJ*, 695, 1648
- Padgett, D., McCabe, C., Rebull, L., et al. 2007, in *BAAS*, 38, 780
- Padmanabhan, N., Schlegel, D. J., Finkbeiner, D. P., et al. 2008, *ApJ*, 674, 1217
- Panic, O., Wilner, D., Qi, C., van Dishoeck, E. F., & Merin, B. 2009, in <http://www.eso.org/sci/meetings/disks2009/talks/panic.pdf>
- Pascucci, I., Hollenbach, D., Najita, J., et al. 2007, *ApJ*, 663, 383
- Pascucci, I., Apai, D., Luhman, K., et al. 2009, *ApJ*, 696, 143
- Petrov, P. P., Gahm, G. F., Gameiro, J. F., et al. 2001, *A&A*, 369, 993
- Piétu, V. V., Dutrey, A., & Guilloteau, S. 2007, *A&A*, 467, 163
- Pontoppidan, K. M., Salyk, C., Blake, G. A., & Käuffl, H. U. 2010, *ApJ*, 722, L173
- Prato, L., Lockhart, K. E., Johns-Krull, C. M., & Rayner, J. T. 2009, *AJ*, 137, 3931
- Pyo, T. S., Hayashi, M., Kobayashi, N., Terada, H., & Tokunaga, A. T. 2009, *ApJ*, 694, 654
- Rebull, L. M., Padgett, D. L., McCabe, C., et al. 2010, *ApJS*, 186, 259
- Reipurth, B., Yu, K. C., Heathcote, S., Bally, J., & Rodríguez, L. F. 2000, *AJ*, 120, 1449
- Rettig, T. W., Haywood, J., Simon, T., Brittain, S. D., & Gibb, E. 2004, *ApJ*, 616, L163
- Robitaille, T. P., Whitney, B. A., Indebetouw, R., & Wood, K. 2007, *ApJS*, 169, 328
- Rodríguez, L. F., Anglada, G., & Raga, A. 1995, *ApJ*, 454, L149
- Salyk, C., Blake, G. A., Boogert, A. C. A., & Brown, J. M. 2007, *ApJ*, 655, L105
- Shang, H., Lizano, S., Glassgold, A., & Shu, F. 2004, in *Star Formation in the Interstellar Medium: In Honor of David Hollenbach*, ed. D. Johnstone, F. C. Adams, D. N. C. Lin, D. A. Neufeld, & E. C. Ostriker, *ASP Conf. Ser.*, 323, 299
- Shang, H., Glassgold, A. E., Lin, W., & Liu, C. 2010, *ApJ*, 714, 1733
- Shukla, S. J., Weintraub, D. A., & Kastner, J. H. 2008, *ApJ*, 683, 893
- Torres, R. M., Loinard, L., Mioduszewski, A. J., & Rodríguez, L. F. 2007, *ApJ*, 671, 1813
- Van Boekel, R., Güdel, M., Henning, T., Lahuis, F., & Pantin, E. 2009, *A&A*, 497, 137
- van den Ancker, M. E., Wesselius, P. R., Tielens, A. G. G. M., van Dishoeck, E. F., & Spinoglio, L. 1999, *A&A*, 348, 877
- Watson, D. M., Kemper, F., Calvet, N., et al. 2004, *ApJS*, 154, 391
- Watson, D. M., Leisenring, J. M., Furlan, E., et al. 2009, *ApJS*, 180, 84
- Weintraub, D. A., Kastner, J. H., & Bary, J. S. 2000, *ApJ*, 541, 767
- Welty, A. D. 1995, *AJ*, 110, 776
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, *ApJS*, 154, 1
- White, R. J., & Ghez, A. M. 2001, *ApJ*, 556, 265
- White, R. J., & Hillenbrand, L. A. 2004, *ApJ*, 616, 998
- Woitats, J. 2003, *A&A*, 406, 685

Table 1. Stellar properties used in this study.

Source	Class	Mass (M_{\odot})	L (L_{\odot})	$EW(H\alpha)$ (\AA)	$\log(\dot{M}_{\odot} \text{ yr}^{-1})$	L_X (10^{30} erg/s)	$L_{\text{mid-IR}}$ (10^{32} erg/s)	$\log(L_{\text{[O III]}}/L_{\odot})$	Jet ^a	Ref.
AA Tau	II	0.76	0.80	37	-8.48	1.24	1.27	-4.68	√	1, 22, 25
BP Tau	II	0.75	0.95	40	-7.88	1.36	1.80	-4.77	×	1, 22
CW Tau	II	1.40	1.10	135	-7.99	2.84	4.62	-2.43	√	1, 23, 26
CZ Tau (A, B)	III	0.32	0.27	4	-9.35	2.24	3.11	...	×	1, 2
DF Tau (A, B)	II	0.34	1.77	54	-7.96	...	2.62	-3.81	√	2, 3, 26
DG Tau	I/II	0.91	1.70	63	-6.13	0.55	15.8	-2.71	√	1, 4, 7, 22, 27
DG Tau B	I/II	0.72	5.50	276	...	9.69	5.09	...	√	1, 31
DK Tau	II	0.74	1.30	40	-7.42	0.92	3.93	-3.85	×	1, 22
DL Tau	II	0.75	0.85	105	-7.17	...	2.87	-4.26	√	1, 2, 5, 25
DM Tau	II	0.47	0.30	139	-8.67	3.18	0.30	...	×	1, 5
DN Tau	II	0.56	1.00	20	-8.72	1.15	1.19	<-5.18	×	1, 2, 22
DO Tau	II	0.70	1.05	100	-7.52	0.18	7.56	-3.97	√	3, 6, 7, 22, 25
DR Tau	II	0.75	0.96	106	-7.49	...	7.41	-3.91	√	2, 3, 6, 22, 25
FS Tau A	II	0.53	0.15	81	-9.50	3.22	3.24	...	×	1, 6
FV Tau (A, B)	II	0.88	0.98	50	-6.23	0.53	5.93	...	×	1
FX Tau	II	0.47	1.02	10–15	-8.65	0.50	1.27	-4.85	×	1, 22
FZ Tau	II	0.57	0.98	181	-7.69	0.64	2.81	...	×	1, 2
GG Tau	II	0.44	0.98	30	-7.76	...	2.50	-4.57	×	3
GI Tau	II	0.75	1.00	20	-8.08	0.83	3.57	-4.46	×	1, 23
GK Tau (A, B)	II	0.74	1.40	25	-8.19	1.47	3.80	-4.36	×	1, 22
GV Tau (A, B)	I	0.68	1.82	86	-6.71	0.65	64.7	...	×	1
HL Tau	I	1.20	1.53	49	-6.80	3.84	29.7	-3.59	√	1, 22, 25
HN Tau	II	1.10	0.25	150	-8.60	0.15	2.67	-3.46	√	1, 2, 23, 25
HP Tau	II	1.39	1.40	28	...	2.55	3.63	...	√	1, 27
HQ Tau	III	2	...	8.16	3.75	...	×	1, 8
IQ Tau	II	0.51	0.88	13	-7.55	0.42	1.43	...	×	1
RY Tau	II	2.37	7.60	13	-7.04	5.52	35.8	-3.65	√	1, 22
T Tau	II	2.41	8.90	41	-7.24	8.05	59.5	-2.63	√	1, 22, 27
UX Tau A	II	1.10	1.81	9.5	-9.30	0.95	1.27	<-4.97	×	2, 6, 9, 22
UZ Tau	II	0.47	0.49	73	-8.01	0.89	3.98	...	×	1
VY Tau	III	0.57	0.55	4.9	...	1.29	0.38	<-5.56	×	2, 7, 10, 22
XZ Tau (A, B)	II	0.34	0.33	62	-8.74	0.96	13.1	...	√	1, 2, 25
Coku Tau-1	II	0.70	0.15	70	-7.36	...	3.90	...	√	1, 31
Coku Tau-3	II	0.46	0.98	5	...	5.85	1.31	...	×	1
Coku Tau-4	II	0.43	0.68	1.8	...	<0.21	0.70	...	×	2, 9, 10
AB Aur	HAe	2.40	49.00	33	-6.85	0.35	64.6	-3.36	√	1, 11, 28
GM Aur	II	0.52	0.74	96	-8.02	1.68	0.68	-4.90	×	3, 5, 7, 22
RW Aur A	II	1.34	3.37	76	-7.50	1.60	4.70	-3.11	√	4, 23, 25, 32
SU Aur	II	1.91	9.90	63	-8.20	9.46	15.0	<-3.85	×	1, 22
UY Aur (A, B)	II	0.57	1.25	36	-7.37	<0.25	10.6	-3.62	√	2, 9, 12, 23, 25
Haro 6-28	II	0.21	0.12	48–92	-8.10	0.25	0.47	...	×	1, 12
Haro 6-37	II	0.91	1.74	22	-8.12	...	2.35	<-5.07	×	2, 6, 22
L1551 IRS5	I	1.58	2.60	83	-6.47	0.021	62.0	...	√	1, 13, 14, 30
MHO-1/2	I/II	0.34	0.42	58	<-8.48	1.64	7.20	...	×	1
V710 Tau	II	0.68	1.10	69	...	1.38	0.75	...	×	1, 6
V773 Tau	II	0.58	5.60	2	<-8.6	9.49	4.46	-3.93	×	1, 6, 22
V892 Tau	HAe	2.89	77.00	7–13	...	9.21	132	...	√	1, 29
HBC 366	III	0.40	2.60	1.3	<-8.94	4.14	0.23	...	×	1, 16
HD 31648	HAe	2.30	23.7	18.3	-8.39	...	34.8	...	×	17
HD 35187	HAe	2.30	1.37	5.1	-7.81	...	18.5	...	√	11, 18
HD 36112	HAe	2.20	1.45	14.5	-6.70	...	17.4	...	×	18, 19
IC2087IR	II	22	-6.20	4.67	16.3	...	×	1
IRAS 04016+2610	I	1.45	3.70	41–56	-7.15	4.47	13.6	...	√	1, 21, 27
IRAS 04101+3103	HAe	11	10.7	...	×	15
IRAS 04108+2803 B	I	0.36	0.4	0.42	2.46	...	√	1, 27
IRAS 04239+2436	I	...	1.27	8.76	...	√	13, 27

Table 1. continued.

Source	Class	Mass (M_{\odot})	L (L_{\odot})	$EW(H\alpha)$ (\AA)	$\log(\dot{M}_{\odot} \text{ yr}^{-1})$	L_X (10^{30} erg/s)	$L_{\text{mid-IR}}$ (10^{32} erg/s)	$\log(L_{[\text{O III}]} / L_{\odot})$	Jet ^a	Ref.
IRAS 04278+2253	I	...	7.20	17	-5.18	...	28.6	...	×	4, 20, 21
IRAS 04299+2915	II	6.5	0.47	...	×	15
IRAS 04303+2240	II	0.51	2.20	67	-6.05	5.01	6.43	...	×	1
IRAS 04361+2547	I	...	3.70	<0.30	8.28	...	√	1, 9, 27
IRAS 04365+2535	I	...	2.20	1.99	6.13	...	√	1, 27
SST 041412.2+280837	I	1.73	...	×	33
SST 042936.0+243555	II	0.32	...	×	24, 33
SST 043905.2+233745	I/II	0.23	...	×	24, 33

Notes. ^(a) √ indicates presence of jet, × indicates that no jet has been reported. ^(b) “...” indicates no information available.

References. 1 Güdel et al. (2007); 2 Watson et al. (2009); 3 Gullbring et al. (1998); 4 White & Hillenbrand (2004); 5 Herbig & Bell (1995); 6 White & Ghez (2001); 7 Güdel priv. comm. 2009; 8 Nguyen et al. (2009); 9 Neuhäuser et al. (1995); 10 Cohen & Kuhi (1979); 11 Garcia Lopez et al. (2006); 12 compiled in McCabe et al. (2006); 13 Prato et al. (2009); 14 Kenyon et al. (1998); 15 Kenyon et al. (1990); 16 compiled in Martin et al. (1994); 17 Mannings & Sargent (1997); 18 Manoj et al. (2006); 19 Isella et al. (2008); 20 Bitner et al. (2008); 21 Furlan et al. (2008); 22 Cabrit et al. (1990); 23 Hartigan et al. (1995); 24 Luhman et al. (2010); 25 Hirth et al. (1997); 26 Hartigan et al. (2004); 27 Moriarty-Schieven et al. (1992); 28 Corcoran & Ray (1997); 29 Maheswar et al. (2002); 30 Davis et al. (2003); 31 Eisloffel & Mundt (1998); 32 Güdel et al. (2010); 33 Rebull et al. (2010).

Table 2. Line luminosities derived in this study.

Name	H ₂ 12.2786	[Ne II] 12.8135	[Ne III] 15.5551	H ₂ 17.0348	[Fe II] 17.9359	[S III] 18.7130	[Fe II] 25.9883	H ₂ 28.2188	[S III] 33.4810
AA Tau	1.2	2.7 ± 0.3	2.2	2.8	1.7	2.2	1.7	1.9	1.9
BP Tau	2.1	1.0	1.3	3.5	2.3	1.8	4.8 ± 0.4	1.9	1.9
CW Tau	11.2	4.9	9.5	8.2	10.9	4.9	3.2	6.2	8.8
CZ Tau (A,B)	4.2	3.8	3.0	3.9	6.5	5.4	2.8	3.2	3.6
DF Tau (A,B)	5.2	8.4	2.4	10.0	3.6	4.6	2.3	3.0	5.8
DG Tau	13.6	43.0 ± 4	9.5	9.8	10.3	10.6	40 ± 2	15.2	9.7
DG Tau B	5.1	10 ± 2	12.3	11.6	8.5	10.1	6.1	3.8	12.9
DK Tau	5.7	5.6	6.7	7.9	6.5	2.7	3.3	3.1	7.4
DL Tau	7.7	4.0	5.5	5.9	6.6	5.6	4.7	2.0	8.9
DM Tau	1.1	0.8 ± 0.2	1.2	0.8	1.2	1.0	0.6	1.5	0.3
DN Tau	1.3	1.2	0.8	0.9	0.8	0.9	1.9	1.4	2.0
DO Tau	4.6	4.5	3.7	4.3	3.2	4.6	2.2	3.9	8.0
DR Tau	12.6	6.5	12.9	18.0	8.8	8.8	3.7	3.3	14.2
FS Tau A	10 ± 2	17 ± 2	4.9	5.2	7.9	7.0	2.6	4.1	4.0
FV Tau (A,B)	6.7	4.3	4.2	6.6	6.3	1.2	3.6	2.1	5.4
FX Tau	1.5	1.1	3.8	2.5	3.3	5.3	1.8	2.3 ± 0.4	1.2
FZ Tau	8.2	4.7	5.2	4.7	8.0	7.0	2.6	6.6	5.8
GG Tau	7.0	6.5	4.1	8.2	5.1	3.9	3.2	3.8	8.9
GI Tau	2.0	2.0	2.9	3.2	2.8	2.1	0.9	2.2	3.1
GK Tau	3.0	5.0	2.5	2.8	3.4	2.1	2.3	3.9	2.9
GV Tau (A,B)	50.5	40.0	41.6	21.4	76.4	30.2	39.7	56.8	51.6
HL Tau	40.0	14.1	14.5	19.8	14.7	13.9	22.8	17.1	35.2
HN Tau	6.3	2.1	4.1	3.3	5.6	4.0	3.4	2.3	6.0
HP Tau	6.8	4.6	4.8	3.9	8.8	6.5	5.0	2.6	6.9
HQ Tau	5.4	5.2	6.4	5.9	10.7	3.1	4.0	3.0	8.4
IQ Tau	1.5	1.2 ± 0.4	1.7	1.6	1.4	2.5	1.6	3.9 ± 0.6	1.7
RY Tau	36.6	13.7	33.7	25.2	17.0	23.3	7.9	7.7	40.9
T Tau(N,S)	49.5	400 ± 9	21.9	54.6	31.2	25.4	36.9	109	119
UX Tau A	6.9	5.6	2.5	2.9	7.5	7.4	1.8	1.8	1.3
UZ Tau	9.1	5.1	4.5	7.3	8.6	9.1	4.5	9.7	8.8
VY Tau	0.3	0.4	0.3	0.2	0.5	0.5	1.7	1.3	1.7
XZ Tau (A,B)	9.5	7.2 ± 0.9	5.4	9.2	15.9	9.9	4.6	9.6	9.6
Coku Tau-1	4.1	27 ± 1	3.9	3.7	5.7	5.2	5.1	7.1	6.0
Coku Tau-3	1.5	1.5	0.5	0.8	1.5	1.3	1.9	1.6	1.3
Coku Tau-4	1.3	3.2 ± 0.2	0.8	0.9	0.9	1.6	1.4	1.1	2.1
AB Aur	29.7	16.8	18.5	25.1	18.2	34.9	44.3	25.7	51.8
GM Aur	2.6	1.5 ± 0.3	1.0	1.0	1.4	2.2	1.3	1.1	4.0
RW Aur	14.6	3.8	7.9	12.9	10 ± 2	6.6	29 ± 1	8.1	5.0
SU Aur	25.2	6.5	6.8	7.8	18.6	17.4	9.8	11.2	10.3
UY Aur (A,B)	14.1	11 ± 2	9.9	8.7	10.0	5.4	4.3	8.3	13
Haro 6-28	4.4	4.4	10.5	3.7	3.7	5.5	1.6	1.4	3.9
Haro 6-37	5.4	4.0	2.2	4.9	5.5	5.7	2.1	3.0	4.0
L1551 IRS5	30.4	129 ± 9	75.8	23.3	200 ± 19	49.8	597 ± 22	20.1	13.1
MHO-1/2	8.4	10 ± 2	4.7	6.9	19.7	12.8	3.5	7.7	11.3
V710 Tau A	0.9	0.4	0.6	1.0 ± 0.3	0.5	1.1	1.0	2.1 ± 0.4	3.5
V773 Tau	3.6	34 ± 2	4.7	8.9	6.8	4.1	3.3	3.6	6.4
V892 Tau	100.4	40.1	141.4	53.7	79.9	56.9	84.3	64.3	166.8
HBC 366	0.4	0.2	0.1	0.5 ± 0.1	0.3	0.2	0.3	0.5	1.0
HD 31648	45.3	7.5	9.9	15.4	35.1	15.5	7.1	5.3	10.4
HD 35187	11.6	11.4	10.8	9.7	16.9	13.5	8.4	8.3	11.1
HD 36112	21.0	11.1	11.2	7.7	15.8	18.3	9.5	5.4	1.8
IC2087IR	20.9	11.3	16.3	11.8	7.8	7.2	4.5	6.4	5.9
IRAS 04016+2610	9.7	5.6	15.9	9.1	8.5	11.4	15.3	8.0	21.1
IRAS 04101+3103	4.8	6.3	6.1	3.8	9.5	7.1	6.2	8.7	4.9
IRAS 04108+2803 B	3.1	3.9	4.6	4.6	6.5	4.5	8.0	6.6	10.0
IRAS 04239+2436	4.6	4.4	6.0	7.3	24 ± 3	8.3	83 ± 3	16.4	38.8

Table 2. continued.

Name	H ₂ 12.2786	[Ne II] 12.8135	[Ne III] 15.5551	H ₂ 17.0348	[Fe II] 17.9359	[S III] 18.7130	[Fe II] 25.9883	H ₂ 28.2188	[S III] 33.4810
IRAS 04278+2253	24.2	17.1	19.1	5.7	21.9	13.4	6.0	9.7	17.8
IRAS 04299+2915	3.6	3.2	3.3	3.6	4.6	2.0	2.9	2.7	6.4
IRAS 04303+2240	14.8	8.2	11.5	13.9	21 ± 2	9.8	48 ± 1	16.9	11.1
IRAS 04361+2547	7.5	3.0	7.8	7.0	8.9	6.6	6.8	34.3	26.4
IRAS 04365+2535	9.4	4.6	8.4	3.7	10.5	6.5	11.9	12.5	19.0
SST 041412.2+280837	4.0	7.5 ± 0.6	7.1	3.2	5.3 ± 0.7	2.8	3.5 ± 0.5	4.7	6.6
SST 042936.0+243555	0.9	2.5 ± 0.2	0.7	0.8	0.5	0.8	0.6	0.8	1.0
SST 043905.2+233745	0.8 ± 0.2	1.6 ± 0.1	0.4	1.2 ± 0.1	0.8	0.09	0.5	1.9	2.6

Notes. Detections plus 1σ errors are in bold characters, 3σ upper limits are presented in normal font. All values are given in units of 10^{28} erg s⁻¹. The wavelength of the emission lines are given in μm .

Appendix A: Results for individual objects

We start with L1551 IRS 5 because it is a good example to illustrate the procedure used to study the line emission. The rest of the objects are ordered as in Table 1.

L1551 IRS 5 is a Class I protostar driving a molecular outflow. High angular resolution observations in the radio have revealed a binary with separation of $0''.3$ (Bieging & Cohen 1985). One of the binary components is now known to be a binary itself with separation of $0''.09$ (Lim & Takakuwa 2006). We do not have any further information on the nature of the components of the system; *Spitzer* observations do not resolve the components (its spatial resolution is $\sim 4''$ at $13\ \mu\text{m}$).

This target is part of a GTO program. The observations were done with 6 different pointings offset in spatial direction. The pointings correspond to three sets of standard nod observations. Therefore we could obtain three average spectra located at three positions; one centered at the star position, and the two others at the offset positions. We can use these data to compare the continuum-subtracted flux at the different positions in order to have a rough estimate for the contribution of the extended emission, whenever the line is detected at multiple positions. In Fig. A.1 we show the spectrum of L1551 IRS 5 around the position of the [Ne II] line. Each plot in the left panel represents the average of the spectra between 2 nods at a given position. In this specific case, we have three different positions identified as “Offset 1”, “Offset 2”, and “Centered”. We have labeled each position in the upper left side of each plot. The maximum flux in the continuum is achieved for the “Centered” position. The line is detected in all three sets of pointings and we point out that it is shifted by $0.1\ \mu\text{m}$ with respect to the expected wavelength of the [Ne II] line. Such a wavelength shift corresponds to a velocity shift of $\sim 230\ \text{km s}^{-1}$. This result is consistent with the velocity of the [Fe II] jet ($285\ \text{km s}^{-1}$) derived by Davis et al. (2003) from near-IR observations. Pyo et al. (2009) have observed the [Fe II] jet in the near-IR and derived a position angle (PA) of 260° for the northern component and 235° for the southern component. *Spitzer* observations have been obtained using a slit PA of -56.7° and -141.5° for the SH and LH modules respectively.

In the right panel of Fig. A.1 we have plotted the continuum-subtracted flux. In this way we could verify that although there is an extended component in the emission, the maximum, i.e., the highest line peak, corresponds to the position centered on the object. But we recall that extended emission close to the star can be present and not be resolved by *Spitzer*.

In the case of [Fe II] at 17.9 and $25.9\ \mu\text{m}$, the lines are detected in all three pointings (Fig. A.2). Both lines are centered at the expected wavelength, in contrast to near-IR observations (Davis et al. 2003) and the [Ne II] line. In both cases, there is an important contribution from extended emission inferred from the high peak of the continuum-subtracted line in the position labeled Offset 2, suggesting diffuse emission as the main contributor to the line luminosity. However, the highest luminosity is still obtained from the pointing centered at the position of the star.

Davis et al. (2003) have reported a detection of an H_2 emission line at $2.12\ \mu\text{m}$, but we do not detect H_2 in the mid-IR spectra obtained with *Spitzer*.

The emission lines from L1551 IRS 5 are the brightest of the sample, with luminosities on the order of $10^{30}\ \text{erg s}^{-1}$. We have obtained VISIR/VLT observations in order to study the kinematics and profile of the [Ne II] line. We have not detected [Ne II] emission at the position of the star, which might suggest a jet origin for the [Ne II] emission, although we have detected

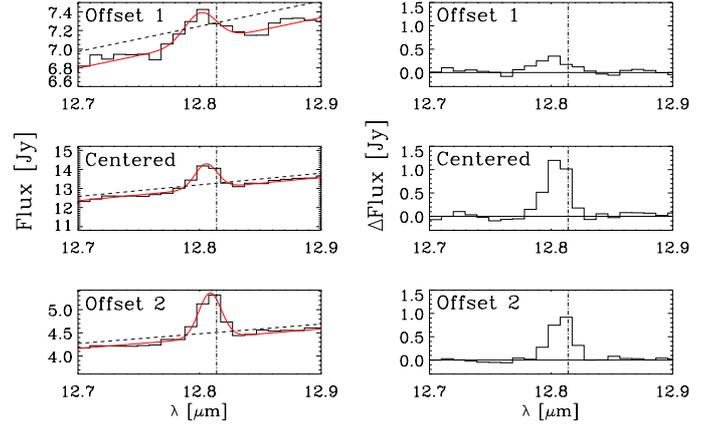


Fig. A.1. Spectrum of L1551 IRS 5 surrounding the position of the [Ne II] line at $12.8\ \mu\text{m}$. The plots on the left panel show the average of 2 pointings at each position. We have labeled the positions as Offset 1, Centered, and Offset 2. The spectrum obtained for the position “Centered” has the highest flux in the continuum so it was used for the line analysis. The expected position of the [Ne II] line is overplotted with a black dash-dotted line, we note that the line is shifted. The right panel shows the difference between the line flux and the continuum flux in order to show the real strength of the line.

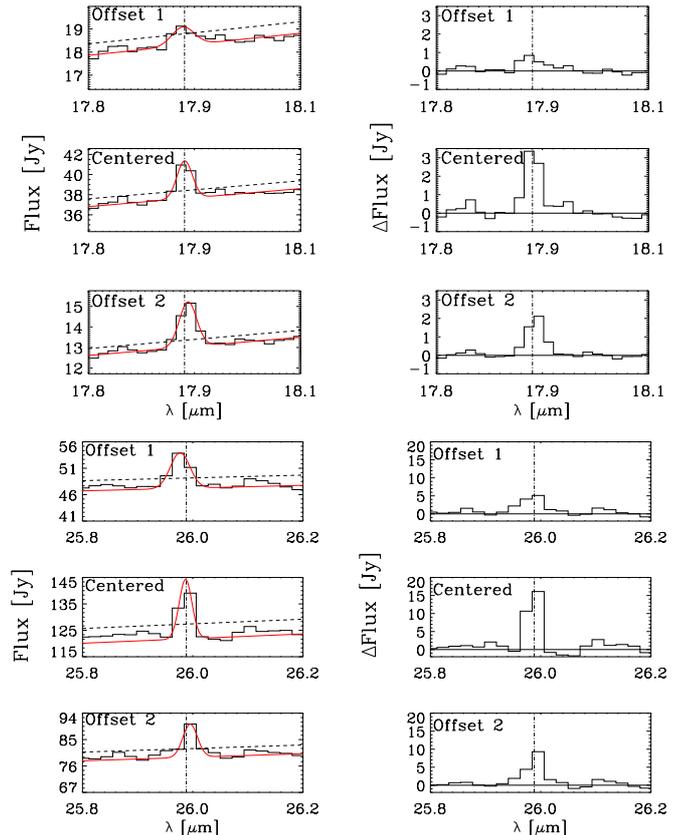


Fig. A.2. Spectrum of L1551 IRS 5 surrounding the position of the [Fe II] line at 17.9 and $25.9\ \mu\text{m}$. The right panel shows the difference between the line flux and the continuum flux in order to show the real strength of the line. The plots on the left panel show the average of 2 pointings at each position. Labels follow the convention used for Fig. A.1.

the system in imaging mode using the [Ne II] filter (Baldovin-Saavedra et al. 2011, in prep.).

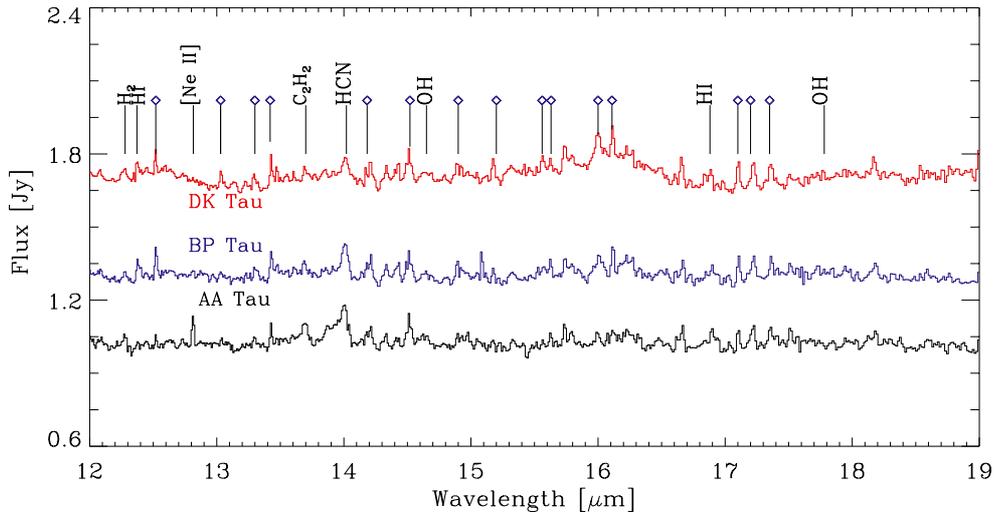


Fig. A.3. Archival SH spectra of DK Tau and BP Tau showing a variety of emission lines from water and organic molecules. We present for comparison the spectrum of AA Tau previously presented in Carr & Najita (2008). The empty diamonds show the position of some water lines (see Pontoppidan et al. 2010). The spectra were background subtracted and shifted to allow comparison.

AA Tau is a Class II source for which [Ne II] is the only line we detected (of the lines we studied). However, Carr & Najita (2008) have presented the spectrum of AA Tau showing multiple emission lines from organic molecular species (HCN, C₂H₂, and CO₂), water vapor, OH, and [Ne II]. Follow-up observations (at $R \sim 80\,000$, Najita et al. 2009) have confirmed that the [Ne II] emission is centered at the stellar velocity, which means that its origin lies in the disk rather than in a jet or outflow. Nevertheless, due to the high inclination of the system ($\sim 75^\circ$, Bouvier et al. 1999), the absence of velocity shifts in the emission line does not rule out a jet origin for the emission. Cox et al. (2005) have indeed reported a microjet detected using coronagraphic imaging with the Hubble Space Telescope (HST).

BP Tau is a Classical T Tauri star (CTTS) that was observed together with background observations. We detect [Fe II] at $25.99\ \mu\text{m}$, with a marginal detection at $17.93\ \mu\text{m}$. There is no record in the literature of previous [Fe II] detections. We have plotted in Fig. A.3 the SH spectra of AA Tau together with BP Tau and DK Tau. Both spectra show spectral features similar to the ones observed in AA Tau and reported by Carr & Najita (2008). The presence of molecular line emission in the spectra of young stars were recently reported in Pontoppidan et al. (2010). Although the spectra of BP Tau and DK Tau are quite similar to the spectrum of AA Tau, we do not detect [Ne II] emission from these two stars. Güdel et al. (2010) detected [Ne II] emission in BP Tau using the same data set as this study, although the line is weak ($\sim 10^{27}\ \text{erg s}^{-1}$). We observe an excess at the position of the [Ne II] line (see Fig. A.3), but below our 3σ detection threshold. A similar situation occurs for the H₂ line at $12.27\ \mu\text{m}$ (see Fig. A.3). The presence of numerous faint emission lines likely makes our detection threshold slightly too high.

DG Tau is classified as both Class I and II in the literature. This young star drives a collimated Herbig-Haro (HH) jet observed in the near-IR. The *Spitzer* observations were made in the same way as L1551 IRS5, with 3 sets of 2 nod pointings. The [Ne II] line is detected in all of them. The observations were made using a slit PA of -57.1° and -141.9° for the SH and LH modules respectively. The jet PA is -138° (Lavalley et al. 1997), almost coincident with the LH module PA. When comparing the

difference between the line and the continuum flux for each set of pointings, we note that the bulk of the emission comes from the position centered on the coordinates of the star, with the ratio being roughly 2:1 between the centered and the offset positions. A detection of [Ne II] has been previously reported in Güdel et al. (2010), based on a different data set than this study. We detect [Fe II] emission at $25.99\ \mu\text{m}$; this line is detected in all pointings, again with a higher contribution from the position centered on the coordinates of the star. Emission from [Fe II] in the near-IR has been previously reported (e.g., Bacciotti et al. 2002; Davis et al. 2003). In particular, Davis et al. (2003) do not detect any [Fe II] emission within $1''$ from the source, suggesting a jet origin for the [Fe II] emission.

DG Tau B has an edge-on disk and a bipolar jet detected in the optical (Mundt & Fried 1983; Eisloffel & Mundt 1998) and radio (Rodriguez et al. 1995). The jet is oriented perpendicular to the disk. The IRS spectrum has been obtained at only one position. We have detected [Ne II] emission towards this object.

DM Tau is a Class II object with [Ne II] line emission. The IRS observations were done with a background observation, so we are confident that the detected line comes from the source itself. Based on *Spitzer* observations made with the low-resolution modules (SL and LL), Espaillat et al. (2007) have derived a luminosity of the [Ne II] line of $1.3 \times 10^{28}\ \text{erg s}^{-1}$. In contrast, we have obtained a luminosity of $(8.0 \pm 2.0) \times 10^{27}\ \text{erg s}^{-1}$, which is 40% lower than what is presented in Espaillat et al. (2007). This object is particularly interesting because it belongs to the category of transitional disks (Calvet et al. 2002). These objects show a lack of emission at wavelengths shorter than $10\ \mu\text{m}$, which is sometimes interpreted as an evidence of an inner hole and planet growth.

FS Tau A is a close binary with separation of $0''.25$ (Krist et al. 1998). The system is surrounded by a filamentary nebula. *Spitzer* observations have been done in 3 sets of 2 nod pointing; the [Ne II] line is detected in all of them. By analyzing the flux difference between the continuum and the line, we have found that the height of the line in the centered position is comparable to the height of the line at one of the offset positions, indicating the emission is likely extended. In addition we have detected

H₂ at 12.28 μm , while there is no detection of the lines at 17.03 and 28.22 μm . At 28.22 μm , there could be a marginal detection, but only below the detection threshold of 3σ . In the case of the 12.28 μm line, it has been detected in 2 out of the 3 sets of pointings. Comparing the continuum-subtracted line flux at the different positions, we determine that the line flux is higher for a position shifted with respect to the star position, indicating an important contribution from extended emission.

FX Tau is a binary with 0'.89 separation composed of a classical and a weak-lined T Tauri star (White & Ghez 2001). Assuming that the mid-IR emission is dominated by the CTTS, we have included this system in the Class II objects for our analysis. We have detected H₂ emission at 28 μm towards this object, and it is the only line detected. No previous detections have been reported.

IQ Tau is a CTTS showing [Ne II] line emission. There is no information in the literature on possible jet or outflow emission associated with this object. Carmona et al. 2011 (in prep.) detected [O I] emission towards this object employing optical spectroscopy ($R \sim 40\,000$), observing a velocity shift of the line which indicates the presence of an outflow.

T Tau is a triple system visible in the optical and bright in X-rays. T Tau S is an infrared companion located 0'.6 away (Dyck et al. 1982) not detected in X-rays. The southern component is a binary itself with a projected separation of $\sim 0'.13$ (Köhler 2008). We have detected the [Ne II] line towards the T Tau system, consistent with previous studies (e.g., van den Ancker et al. 1999; Güdel et al. 2010). Based on the analysis of the continuum-subtracted line, there is a contribution from extended emission, in agreement with previous results. The first detections of [Ne II], [Fe II], and [O I], among other gas tracers, come from van den Ancker et al. (1999), based on Infrared Space Observatory (ISO) observations. However, we do not detect [Fe II] emission, which would be expected to be present in an outflow environment. Van Boekel et al. (2009) studied this system ($R \sim 30\,000$) and have determined that the emission is extended and associated with an outflow from the North and South components. Only a small fraction of the emission is related to the X-ray bright T Tau N.

XZ Tau is a close binary system with separation of 0'.3. Both stars have collimated, bipolar jets (Krist et al. 2008). Due to the close separation, we cannot resolve the system with *Spitzer*. Observations have been made in 3 sets of 2 nod pointings; we detect [Ne II] only at the pointing with the highest flux in the continuum, centered at the position of the stars. This is the first detection of [Ne II] reported toward this source.

CoKu Tau 1 is a Class II object. It is a binary with separation of 0'.24 (Robitaille et al. 2007) in which we have detected [Ne II] emission. Previous observations in H α show evidence of an outflow (Eisloffel & Mundt 1998), but we do not detect the [Fe II] emission usually observed in outflows.

CoKu Tau 4 is a Class II object known to have a large excess at $\sim 30 \mu\text{m}$ and no excess at $< 8 \mu\text{m}$ (Forrest et al. 2004). This was interpreted as a clear signature of a transitional disk, but recent observations (Ireland & Kraus 2008) have shown that the clearing of the disk at shorter wavelengths is actually due to the presence of a stellar-mass companion. The binary projected separation is calculated to be ~ 7.8 AU. We have detected [Ne II] emission towards this binary.

GM Aur is a Class II source with [Ne II] emission; this object is also a transitional disk (Calvet et al. 2005). We have dedicated background observations, so we can discard the possibility of the emission coming from the background or a nearby source. The first detection of [Ne II] in the *Spitzer* spectrum was reported by

Carr & Najita (2008). Follow-up observations of the [Ne II] line at $R \sim 80\,000$ (Najita et al. 2009) have confirmed that the emission is from a disk rather than an outflow. Espaillat et al. (2007) have derived a 5σ upper limit for the luminosity of the [Ne II] line of $7.4 \times 10^{28} \text{ erg s}^{-1}$. This result is based on observations made with *Spitzer* IRS low-resolution modules ($R \sim 60\text{--}120$). Our calculations based on the high resolution module give a luminosity of $(1.5 \pm 0.3) \times 10^{28} \text{ erg s}^{-1}$.

RW Aur is an optically bright T Tauri star; Petrov et al. (2001) suggested that this is a binary system. The secondary, RW Aur B, is located 1'.2 away from the primary. We have detected both [Fe II] lines (17.93 and 25.99 μm) in the spectrum of this system, which is not spatially resolved by *Spitzer*. The observations were made in 3 sets of 2 nod pointings. The line at 17.93 μm is detected in one of the offset positions, indicating that the emission is likely to be extended. The [Fe II] line at 25.99 μm is detected in all the three positions: centered on the position of the star, and two in positions displaced from the star position. The contribution from extended emission is important, but the line flux is higher at the position of the source. Previous observations have shown emission of [Fe II] in the near-IR from the jet associated with this system (Hartigan & Hillenbrand 2009).

UY Aur (A, B) is a binary system with separation 0'.88 (Hirth et al. 1997). Observations were made in 3 sets of 3 nod pointings. We have detected [Ne II] emission at the 3σ limit. The line is detected only in the pointing centered in the position of the system. We note that in this case the line is broad, with a *FWHM* 1.6 times higher than the IRS instrumental *FWHM*. This system is known to drive a high velocity ($\sim 300 \text{ km s}^{-1}$) [O I] bipolar outflow (Hirth et al. 1997). We do not observe any shift of the position of the line; the broadening could indicate either a strong wind or a jet with a blue and red-shifted component.

MHO 1/2 is a binary with a separation of 3'' (Briceño et al. 1998). Since the *Spitzer* observations do not resolve the system, we cannot know in which star the emission originates. The system has been observed in 3 sets of 2 nod pointings. The [Ne II] line is detected only at one of the three positions.

V773 Tau is a multiple system. V773 Tau AB is a WTTS spectroscopic binary with a separation of 0'.003, i.e., ~ 0.37 AU at the distance of Taurus (Welty 1995). V773 Tau C is a visual companion at 0'.17 from V773 Tau AB, classified as a CTTS. A fourth member of this system, V773 Tau D, has an angular separation of $\sim 0'.22$ with respect to V773 Tau A (Woitas 2003) and it has been classified as an infrared companion due to the shape of its SED (Duchêne et al. 2003). These objects are believed to be T Tauri stars embedded in dusty envelopes and undergoing episodes of intense accretion. But their true nature is not yet understood. The *Spitzer* observations do not resolve the components of the system, but according to the SED presented in Duchêne et al. (2003) we assume that V773 Tau D is responsible for the bulk of the mid-IR emission and therefore for the level of continuum observed. However, we cannot be sure that the [Ne II] emission comes from it. In fact, it could also be produced in shocks from one of the components or in the interstellar environment in a similar way as the case of T Tau. The stellar properties for the system are compiled in Table 1. The H α emission and X-ray luminosity probably come from V773 Tau AB, the WTTS binary, while the mass accretion rate value adopted is the one derived for V773 Tau C, which is likely similar to V773 Tau D. We assume that in the infrared regime, the emission is dominated by the CTTS, so we have included the system in the group of Class II sources. In the correlation tests presented in Sect. 5, we have not used the H α equivalent width, which is derived for the WTTS binary.

HBC 366 is a WTTS. We have detected one of the H₂ lines (17.03 μm) in this spectrum. No other gas lines have been detected.

IRAS 04239+2436 is a binary of 0'.3 separation, driving a highly collimated near-IR [Fe II] jet (Reipurth et al. 2000). The observations were done in three sets of two nod pointings. We detect both iron lines (17.93 and 25.99 μm) in the IRS spectrum. The two lines are detected in all three sets of pointings, but when analyzing the continuum-subtracted line, the highest contribution comes from the position centered on the star. The luminosity of the line at 25.99 μm is more than a factor of 3 higher than the luminosity at 17.93 μm . Previous observations (Davis et al. 2001; Davis et al. 2003) have detected emission of H₂ (2.122 μm) and [Fe II] (1.644 μm). The *Spitzer* spectrum of this source does not show any evidence of H₂ emission.

IRAS 04303+2240 is a CTTS. We have detected both [Fe II] lines (at 17.93 and 25.99 μm), with the line at 17.93 μm detected at the 3σ level. Although there is an excess at the wavelength of the [Ne II] line, it is formally not detected (at the 3σ detection limit). The ratio between the luminosity of both [Fe II] lines is 1:2. Observations were done in 3 sets of 2 nod pointings; the line at 17.93 μm is detected only in the position centered at the star. The line at 25.99 μm is detected in two out of three positions; the contributions of the centered position and the off-set position are similar. Both lines are broadened by a factor of two compared to the intrinsic line width of the instrument. There is no information in the literature about previous detections of [Fe II] towards this source or about jets or outflows that could explain the [Fe II] emission, or winds explaining the broadening of the lines.

V710 Tau is a binary system formed by a CTTS and a WTTS. Observations in the X-rays made with *Chandra* (Shukla et al. 2008) and in the near-IR (White & Ghez 2001) have shown that the separation of the binary is 3''.2. We detect H₂ at 17.03 and 28.22 μm . Considering that the star responsible for the continuum at mid-IR wavelengths is the CTTS, V710 Tau A, we have used the H α equivalent width and X-ray luminosity of that star in the present study, and included the system in the Class II group, but the lines could also be produced by the WTTS.

SST 041412.2+280837 was included in a sample of the low-mass star population of Taurus for which we have obtained *Spitzer*-IRS spectroscopy. It was previously known as IRAS 04111+2800G and is classified as a Class I object (Rebull et al. 2010). This is the only object in our sample for which we have used low-resolution spectra (SL and LL) since we do not have the SH and LH spectra. We have detected [Ne II] and [Fe II] at 17.9 and 25.9 μm towards this object.

SST 042936.0+243555 is a newly discovered member of Taurus (Rebull et al. 2010). It belongs to the sample of low luminosity sources for which we have obtained IRS spectra. Rebull et al. (2010) and Luhman et al. (2010) both classify this object as a Class II. The spectrum shows [Ne II] emission.

SST 043905.2+233745 also belongs to the sample of newly discovered Taurus members (Rebull et al. 2010). It is a Class I object as classified by Luhman et al. (2010), but a Class II object according to Rebull et al. (2010). We have detected emission lines from [Ne II] and H₂ at 17.0 μm .

Appendix B: [Ne II] non-detections

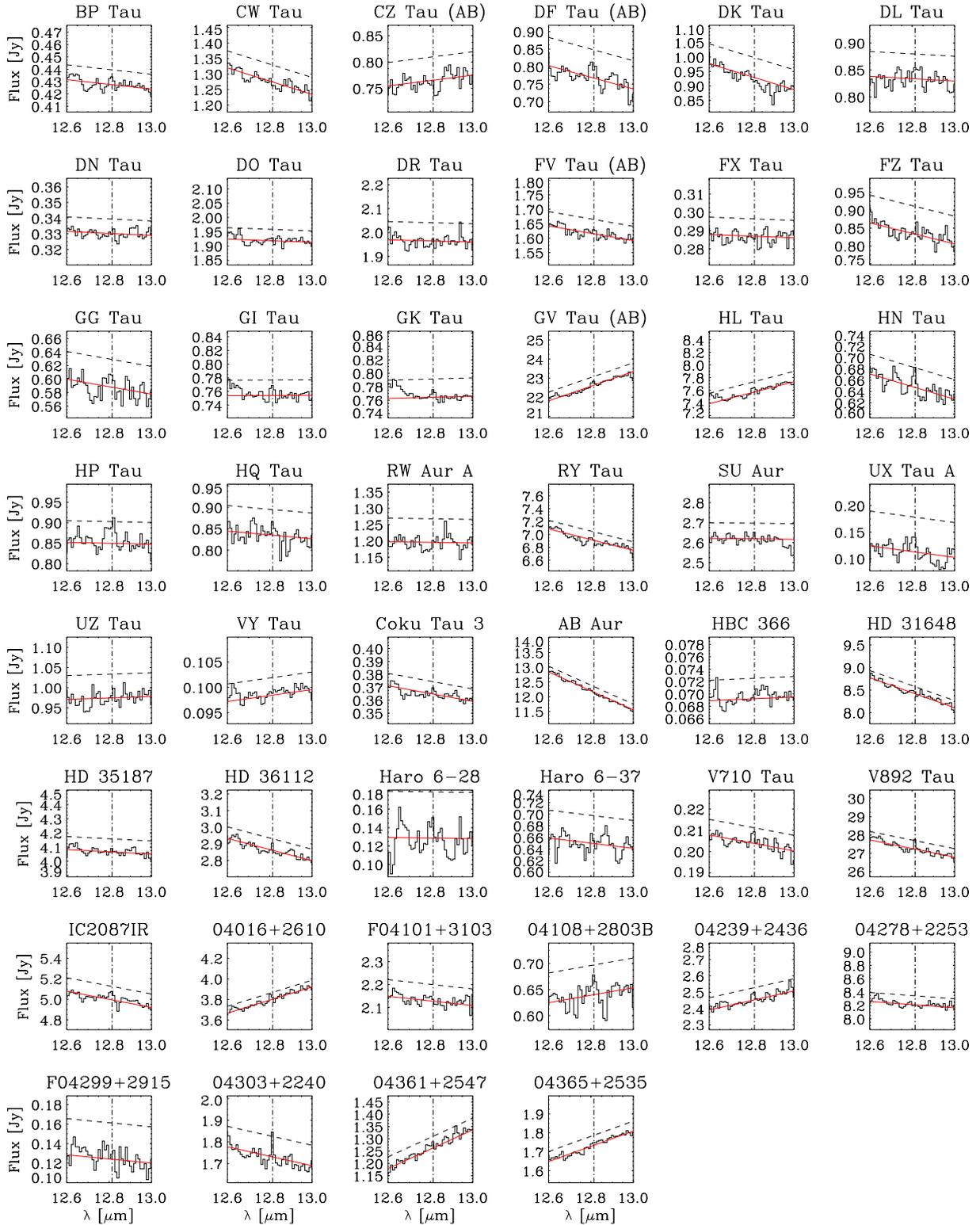


Fig. B.1. Spectra around the [Ne II] line of the targets in which we did not detect the line. The position of the line is plotted in dashed-dotted line and the 3σ noise level is plotted in dashed line. The name of each star is labeled in the top of each plot.