

NO EVIDENCE FOR EVOLUTION IN THE FAR-INFRARED–RADIO CORRELATION OUT TO $z \sim 2$ IN THE EXTENDED CHANDRA DEEP FIELD SOUTH

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ABSTRACT

We investigate the $70\ \mu\text{m}$ far-infrared–radio correlation (FRC) of star-forming galaxies in the Extended Chandra Deep Field South (ECDFS) out to $z > 2$. We use $70\ \mu\text{m}$ data from the Far-Infrared Deep Extragalactic Legacy Survey (FIDEL), which comprises the most sensitive ($\sim 0.8\ \text{mJy rms}$) and extensive far-infrared deep field observations using MIPS on the *Spitzer Space Telescope*, and $1.4\ \text{GHz}$ radio data ($\sim 8\ \mu\text{Jy beam}^{-1}\ \text{rms}$) from the Very Large Array. In order to quantify the evolution of the FRC, we use both survival analysis and stacking techniques, which we find give similar results. We also calculate the FRC using total infrared luminosity and rest-frame radio luminosity, q_{TIR} , and find that q_{TIR} is constant (within 0.22) over the redshift range 0–2. We see no evidence for evolution in the FRC at $70\ \mu\text{m}$, which is surprising given the many factors that are expected to change this ratio at high redshifts.

Key words: Galaxy: evolution – Galaxy: formation – infrared: galaxies – radio continuum: galaxies

Online-only material: color figures

1. INTRODUCTION

The correlation between the far-infrared (FIR) and radio emission for star-forming galaxies in the local universe was first observed by van der Kruit (1971, 1973) and is the tightest and most universal correlation known among global parameters of galaxies (Helou & Bicay 1993). The correlation is linear, spans five orders of magnitude of bolometric luminosity and has been shown to hold for a wide range of Hubble types (de Jong et al. 1985; Helou et al. 1985; Condon 1992; Yun et al. 2001).

The FIR–radio correlation (FRC) has been attributed to the presence of young, high-mass ($M > 8 M_{\odot}$) stars. The FIR emission arises from the absorption by dust of the UV emission and the subsequent reradiation of the energy at IR wavelengths. The radio emission is dominated by non-thermal synchrotron emission from cosmic ray electrons which are accelerated by supernovae shocks. There is also a thermal component which arises from free–free emission from ionized hydrogen in H II regions, but this contributes only $\sim 10\%$ of the radio emission at lower frequencies ($< 5\ \text{GHz}$; Condon 1992), and becomes more significant at higher frequencies.

The “calorimeter theory” (Voelk 1989) suggests that the FRC holds because galaxies are both electron calorimeters and UV calorimeters, hence the total radio and IR outputs remain proportional independent of variations within the galaxy. This theory, however, requires that galaxies are optically thick to UV light from the young high-mass stars and thus does not hold for optically thin galaxies. Helou & Bicay (1993) proposed the non-calorimetric “optically thin” scenario involving a correlation between the disk scale height and the escape scale length for cosmic ray electrons, while Bell (2003) concludes that the linearity of the FRC is a conspiracy as the star-formation rate is

underestimated for low-luminosity galaxies at both radio and infrared frequencies. However, these models typically leave out proton losses and non-synchrotron cooling (Lacki et al. 2010). Ultimately, the physical origin of the FRC is still not clear.

The far-reaching nature of the FRC has made it a valuable diagnostic. Some examples of its application include: using the FRC to identify radio-loud active galactic nucleus (AGN; Donley et al. 2005; Norris et al. 2006); using the FRC to define the radio luminosity/SFR relation (Bell 2003); and, at higher redshifts, using the FRC to estimate distances to submillimeter galaxies without optical counterparts (e.g., Carilli & Yun 1999). Consequently, it is of great importance to determine whether the FRC holds at high redshifts.

The FRC may fail at high redshifts for a number of reasons. Electrons are expected to lose energy by inverse Compton interactions with the cosmic microwave background, whose energy density scales as $(1+z)^4$, implying a lower level of radio emission at higher redshifts. Moreover, synchrotron emission is proportional to the magnetic field strength squared, so evolution of magnetic field strength should affect the FRC at higher redshifts (e.g., Murphy 2009). Changes in the spectral energy distributions (SEDs) may also be expected due to evolution in dust properties and metallicity (e.g., Amblard et al. 2010; Hwang et al. 2010; Chapman et al. 2010). Nonetheless, current studies show no firm evidence for evolution in the FRC (e.g., Garrett 2002; Appleton et al. 2004; Seymour et al. 2009; Bourne et al. 2011; Ivison et al. 2010a, 2010b; Sargent et al. 2010a, 2010b; Huynh et al. 2010).

Herschel was launched in 2009 May and can probe the FIR to submillimeter regime from $55\ \mu\text{m}$ to $671\ \mu\text{m}$, deeper than ever before (Pilbratt et al. 2010). Most recently, Jarvis et al. (2010) and Ivison et al. (2010b) used data from *Herschel* and found no

evidence for evolution in the FRC out to $z = 0.5$ and $z = 2$, respectively.

This paper studies the dependence of the FRC on redshift using deep $70\ \mu\text{m}$ data from the *Spitzer Space Telescope* and 1.4 GHz data from the Very Large Array (VLA). This work differs from previous studies as we are using FIDEL (Far-Infrared Deep Extragalactic Legacy Survey) $70\ \mu\text{m}$ data, which is the deepest $70\ \mu\text{m}$ data taken to date. FIDEL reaches a point source rms sensitivity of $0.8\ \text{mJy}$ at $70\ \mu\text{m}$, making it far more sensitive than previous studies of the FIR at $70\ \mu\text{m}$ such as Sargent et al. (2010b) whose data reached a point source rms sensitivity of $1.7\ \text{mJy}$.

We focus on the $70\ \mu\text{m}$ data as this band probes closer to the dust emission peak ($\sim 100\ \mu\text{m}$) than, for example, $24\ \mu\text{m}$. Furthermore, for $z < 3$, the $70\ \mu\text{m}$ band is not affected by emission from polycyclic aromatic hydrocarbons (PAHs; $7\text{--}12\ \mu\text{m}$). While Bourne et al. (2011) also study the FRC using FIDEL data, their work was entirely based on stacking analysis. This work is the first to use such deep $70\ \mu\text{m}$ data to study the evolution of the FRC based on individual sources.

The data are described in Section 2, Section 3 describes the data analysis, while Section 4 presents our results and analysis. This paper uses $H_0 = 71\ \text{km s}^{-1}\ \text{Mpc}^{-1}$, $\Omega_M = 0.27$, and $\Omega_\Lambda = 0.73$.

2. DATA

2.1. FIDEL

FIDEL, the Far-Infrared Deep Extragalactic Legacy Survey, is a legacy science program (PI: Dickinson) which comprises the most sensitive and extensive FIR deep field observations using the Multiband Imaging Photometer for SIRTf (MIPS) on the *Spitzer Space Telescope*. Characteristics of the FIDEL data are described in detail in Magnelli et al. (2009) and other papers. FIDEL observed three fields: ECDFS, EGS, and GOODS-North. Observations were taken at 3 bands: $24\ \mu\text{m}$, $70\ \mu\text{m}$, and $160\ \mu\text{m}$, focusing specifically on the $70\ \mu\text{m}$ data. This paper concentrates only on the $30' \times 30'$ ECDFS field centered on $03\ 32\ 00, -27\ 48\ 00$ (J2000). The observations were designed to achieve roughly uniform sensitivity at $70\ \mu\text{m}$ across most of the ECDFS, although the FIDEL data include a somewhat deeper central region with data from a GO program (PI: Frayer). The mean $70\ \mu\text{m}$ exposure time over the ECDFS is approximately $6600\ \text{s}$, yielding an rms point source sensitivity of approximately $0.8\ \text{mJy}$. The $24\ \mu\text{m}$ exposure time varies considerably more over the field, from 5000 to $35,000\ \text{s}$ over most of the ECDFS, with an average exposure time of approximately $16,000\ \text{s}$. The rms point source sensitivity at $24\ \mu\text{m}$ thus also varies, but is typically in the range $8\text{--}14\ \mu\text{Jy}$ over most of the field. FIDEL data were processed using the Mosaicking and Point-source Extraction (MOPEX; Makovoz & Marleau 2005) package to form the mosaicked images. The final $70\ \mu\text{m}$ mosaic has a pixel scale of $4''.0\ \text{pixel}^{-1}$ and a point response function (PRF) with an $18''$ FWHM. The final $24\ \mu\text{m}$ mosaic has a pixel scale of $1''.2\ \text{pixel}^{-1}$ and a PRF with a $5''.9$ FWHM.

2.2. Ancillary Data

2.2.1. Radio Data

The ECDFS has been observed at 1.4 GHz by both the VLA and the ATCA (Norris et al. 2006; Kellermann et al. 2008; Miller et al. 2008). Here, we use the Miller et al. (2008) radio data due to its high angular resolution and sensitivity over our field. The

radio data encompass a $34' \times 34'$ region, centered on $03\ 32\ 28.0, -27\ 48\ 30.0$ (J2000). The data have a typical rms sensitivity of $8\ \mu\text{Jy}$ per $2''.8 \times 1''.6$ beam. The catalog contains 464 sources above a 7σ cutoff.

2.2.2. Redshift Data

We obtained both spectroscopic and photometric redshift data from COMBO-17 (Wolf et al. 2004), MUSYC (Gawiser et al. 2006; Cardamone et al. 2010), GOODS (Balestra et al. 2010), and ATLAS (Mao et al. 2009).

COMBO-17 has photometric data in 17 passbands from $350\ \text{nm}$ to $930\ \text{nm}$ for 63501 objects in the ECDFS. The photometric redshifts are most reliable for sources with $R \leq 24$ (Wolf et al. 2004).

MUSYC (Multiwavelength Survey by Yale-Chile) has photometric data in the ECDFS. Cardamone et al. (2010) combined photometric data from the literature with new deep 18-medium-band photometry into a public catalog of $\sim 80,000$ galaxies in ECDFS, from which they computed the photometric redshifts. The photometric redshifts are most reliable for sources with $R \leq 25.5$. Cardamone et al. (2010) also compile a spectroscopic redshift catalog of 2551 galaxies from the literature (e.g., Balestra et al. 2010; Vanzella et al. 2008; Le Fèvre et al. 2004).

Balestra et al. (2010) observed the GOODS-South field (within the ECDFS) using VIMOS to obtain spectroscopic redshifts. Their campaign used two different grisms to cover different redshift ranges and used 20 VIMOS masks. They combined their resulting redshifts with those available in the literature to produce a catalog containing 7332 spectroscopic redshifts. Quality flags were provided for all the redshifts and we took only those redshifts that had quality flags of “secure” and “likely” yielding a catalog of 5528 spectroscopic redshifts. Although Cardamone et al. (2010) includes Balestra et al. (2010) data in their compilation of spectroscopic redshifts, they only include a subset of the data.

Mao et al. (2009) are undertaking a program of redshift determination and source classification of all ATLAS (Australia Telescope Large Area Survey) radio sources with AAOmega (Sharp et al. 2006) on the Anglo-Australian Telescope (AAT). Using redshifts from both the literature and their own campaign, they have a total of 261 spectroscopic redshifts in ECDFS and its surrounding region.

2.2.3. X-ray Data

We use the 2 Ms Chandra Deep Field South X-ray catalogs from Luo et al. (2008) to identify AGNs in our $70\ \mu\text{m}$ catalog (Section 4.1). This is one of the most sensitive X-ray surveys ever performed and detects 462 sources in $436\ \text{arcmin}^2$. While the X-ray data cover a smaller region than the FIDEL data, we use these data to eliminate some AGNs from our sample (Section 4.1).

3. DATA ANALYSIS

3.1. FIDEL

3.1.1. $70\ \mu\text{m}$

The $70\ \mu\text{m}$ catalog was produced with the Astronomical Point Source Extraction (APEX) module within the MOPEX package. APEX subtracts the local background by calculating the median in a region, which we set to 34×34 pixels, surrounding each pixel, and removing the 100 brightest pixels. Peak values with

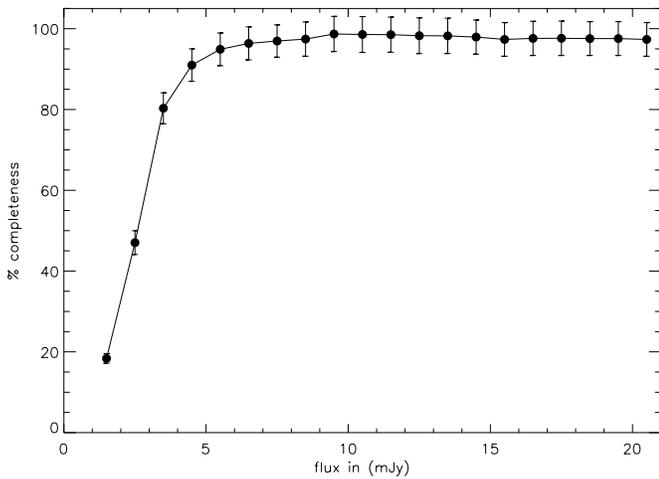


Figure 1. Percentage completeness plotted against flux density. Our catalog is almost 100% complete at 9 mJy, and 50% complete at ~ 2.5 mJy.

a signal-to-noise ratio greater than 3 were fitted using the PRF. The 3σ catalog extracted using APEX contained 515 sources.

3.1.2. $24\ \mu\text{m}$

A similar source detection and extraction process to the $70\ \mu\text{m}$ data were used for the $24\ \mu\text{m}$ data. However, a 4σ cutoff was used to reduce the number of spurious sources. Visual inspection was required to remove spurious sources due to artifacts surrounding bright objects. The final catalog of $24\ \mu\text{m}$ sources in the $\sim 30' \times 30'$ region of interest contained 5319 sources.

3.1.3. Final Catalog

To test the completeness of the $70\ \mu\text{m}$ catalog, we performed Monte Carlo simulations. A simulated source was injected at a random location in the $70\ \mu\text{m}$ image, and source extraction was performed using the same technique as for the production of the catalog. The input flux density of the simulated source varied over the range of fluxes in the “real” image. This process was repeated 10,000 times and we tracked the simulated source recovery rate, which enables us to estimate the overall completeness over the entire image. Figure 1 presents the completeness as a function of flux density. Our catalog is almost 100% complete at 9 mJy, and 50% complete at ~ 2.5 mJy. This high level of completeness at 2.5 mJy ($\sim 3\sigma$) is probably due to the $70\ \mu\text{m}$ image not being completely uniform. The 10×10 arcmin region in the center has an rms point source sensitivity of ~ 0.6 mJy, so 2.5 mJy is a $>4\sigma$ limit for $\sim 11\%$ of the image area. The completeness plot does not reach 100% due to the random positioning of the simulated source. If, by chance, the simulated source is injected upon a real source, the simulated source is not recovered as an individual source, and instead, the “real” source is recovered with a larger flux density.

The $70\ \mu\text{m}$ catalog of 515 sources obtained using APEX may contain confused sources, as well as spurious sources. We determined which sources required deblending by comparing it with the $24\ \mu\text{m}$ catalog which has better resolution. First, we matched the $70\ \mu\text{m}$ catalog to the $24\ \mu\text{m}$ catalog using a matching radius of $9''$ (the half-width at half-maximum of the PRF at $70\ \mu\text{m}$). We then calculated the ratio of the flux density of the brightest $24\ \mu\text{m}$ source to the total flux density of all $24\ \mu\text{m}$ sources within the $9''$ matching radius (see Figure 2). Those sources for which the ratio is 1, would not require deblending.

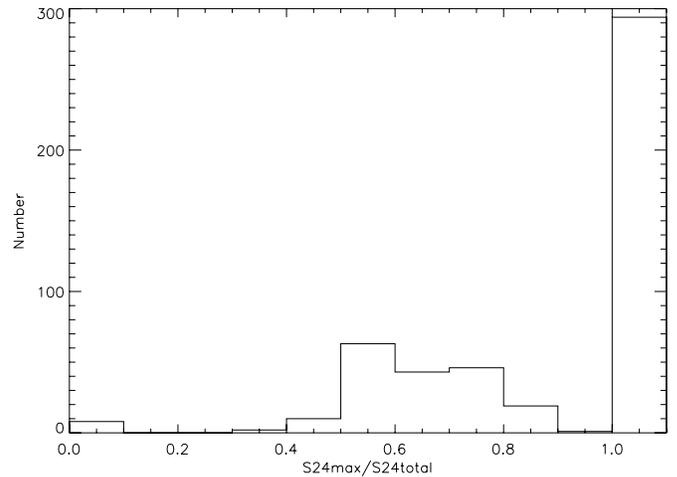


Figure 2. Histogram of the ratio of the maximum $24\ \mu\text{m}$ flux density over the total $24\ \mu\text{m}$ flux density within a $9''$ radius. Sources with a ratio less than 0.8 were deemed deblend candidates.

Using arguments similar to those of Pope et al. (2006), we assume that if a source is bright at $24\ \mu\text{m}$, it will also be bright at $70\ \mu\text{m}$. Consequently, if the ratio was greater than 0.8 (that is, if the brightest $24\ \mu\text{m}$ source contributed $>80\%$ of the total flux density in the matching radius), the $70\ \mu\text{m}$ emission was determined to be from the brightest $24\ \mu\text{m}$ source. If the ratio is less than 0.8, we considered the $70\ \mu\text{m}$ source as a candidate for deblending. Using these criteria we determined that 164 sources were deblend candidates. In principle, it is possible that for sources with a ratio of greater than 0.8, we are overestimating the $70\ \mu\text{m}$ flux density by up to 20%. An overestimation of the $70\ \mu\text{m}$ flux density may result in a higher FRC, but this effect is much less than $\Delta q_{70} = 0.1$.

The sources were deblended using double, triple or even quadruple Gaussian fits, constraining the centers of the Gaussians to the $24\ \mu\text{m}$ positions. We visually inspected all the deblend candidates and a small number of these clearly did not require deblending as the secondary $24\ \mu\text{m}$ source was right on the edge of the 9 arcsec radius with which we computed the deblend criteria, and were subsequently discarded as deblend candidates. In total, 143 sources were deblended, resulting in a total catalog of 658 putative sources. In the course of the visual inspection, we identified 41 spurious sources. 30 of these were sources that were faint ($<4\sigma$) in the $70\ \mu\text{m}$ catalog, and did not have $24\ \mu\text{m}$ counterparts, and 11 were clearly part of the Airy ring of an adjacent, bright source. This resulted in a final catalog containing 617 $70\ \mu\text{m}$ sources.

Assuming Gaussian statistics, noise spikes are expected to produce approximately 34 spurious sources above our 3σ cutoff in the $70\ \mu\text{m}$ data, and approximately 7 spurious sources above our 4σ cutoff in the $24\ \mu\text{m}$ data. The probability of any spurious $70\ \mu\text{m}$ source lying within a 9 arcsec radius of any spurious $24\ \mu\text{m}$ source is approximately 2%, and so we conclude that none of the 617 sources is likely to be spurious.

We note that after deblending the $70\ \mu\text{m}$ catalog we increased the number of faint sources, which changes the completeness levels. While we have improved our completeness at the fainter flux levels, this change cannot easily be quantified.

This catalog of 617 $70\ \mu\text{m}$ sources was matched to the $24\ \mu\text{m}$ catalog using a matching radius of $4''$. Following Huynh et al. (2008), the matching radius was determined by plotting the number of candidate matches against position offset (Figure 3). The matches from chance alone are determined from the source

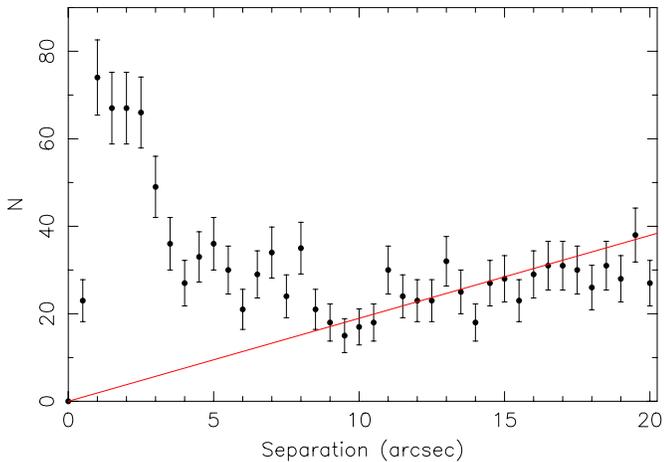


Figure 3. Number of candidate $24\ \mu\text{m}$ counterparts of $70\ \mu\text{m}$ sources as a function of position offset. The line shows the number of matches from chance alone.

(A color version of this figure is available in the online journal.)

densities of the catalogs. Although Figure 3 suggests that a greater number of matches would be obtained using a matching radius of $8''$, we choose the more conservative $4''$ as our matching radius to minimize spurious matches.

Given the relative sensitivities of the $70\ \mu\text{m}$ and $24\ \mu\text{m}$ data, it is expected that most if not all the $70\ \mu\text{m}$ sources will have a $24\ \mu\text{m}$ counterpart. A small number (10) of strong $70\ \mu\text{m}$ sources did not have $24\ \mu\text{m}$ counterparts in the catalogs, while the $24\ \mu\text{m}$ image showed faint detections. We therefore performed aperture photometry on these faint sources and successfully extracted eight. The final catalog has only two $70\ \mu\text{m}$ sources without $24\ \mu\text{m}$ counterparts but neither of these have extreme S_{70}/S_{24} ratios.

3.2. Radio, Optical, and X-ray Counterparts

3.2.1. Radio Counterparts

The $70\ \mu\text{m}$ catalog was matched to the radio catalog of Miller et al. (2008, hereafter M08), using a matching radius of $5''$. The matching radius was determined in a similar manner to that described in the previous section. Where available, the $24\ \mu\text{m}$ position was used because the higher resolution at $24\ \mu\text{m}$ allows for better positional accuracy. This resulted in 171 radio sources matched to the $70\ \mu\text{m}$ catalog.

The M08 radio catalog has a 7σ cutoff. In order to increase the number of radio counterparts, and given that we have the additional information of the $70\ \mu\text{m}$ sources, we extracted radio sources at the known $70\ \mu\text{m}$ source positions that had radio detections greater than 3σ . This was done by performing Gaussian fits twice, first with a fixed size ($2''.8 \times 1''.6$ —the beam of M08 data), and then without a fixed size. The relationship between peak and integrated flux densities was determined by

$$\frac{S_{\text{int}}}{S_{\text{peak}}} = \frac{\theta_{\text{maj}}\theta_{\text{min}}}{b_{\text{maj}}b_{\text{min}}}, \quad (1)$$

where S_{int} is the total integrated flux density, S_{peak} is the peak flux density, θ_{maj} is the semi-major axis of the free Gaussian, θ_{min} is the semi-minor axis of the free Gaussian, b_{maj} is the semi-major axis of the fixed Gaussian, and b_{min} is the semi-minor axis of the fixed Gaussian.

If $S_{\text{int}}/S_{\text{peak}}$ was less than 1.2, the source was determined to be unresolved and the peak flux density for the fixed Gaussian was

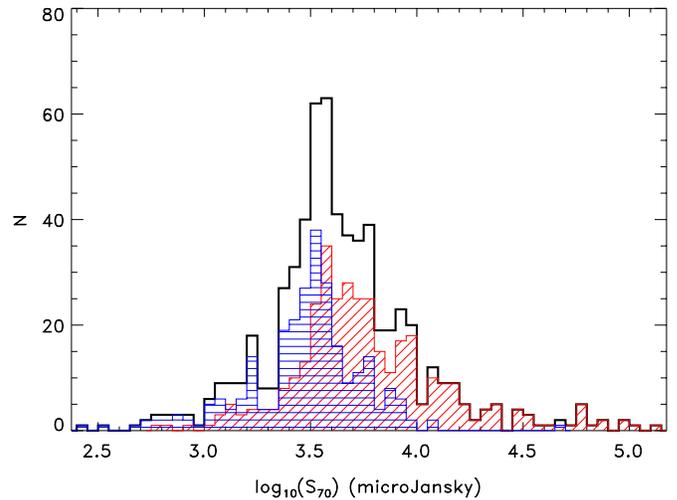


Figure 4. Histogram of $70\ \mu\text{m}$ fluxes. The black line is the histogram for all 617 sources, while the red diagonally shaded histogram shows the 353 sources that have a radio detection and the blue horizontally shaded histogram shows the 264 sources that do not have a radio detection.

(A color version of this figure is available in the online journal.)

determined to be the final flux density. If the ratio was greater than 1.2, the source was determined to be extended, and the integrated flux was taken to be the final radio flux density. This resulted in 353 $70\ \mu\text{m}$ sources with radio counterparts.

In summary, our final $70\ \mu\text{m}$ catalog contains 617 sources, 615 of which have a $24\ \mu\text{m}$ counterpart. Three hundred and fifty three of these have a radio detection greater than 3σ from Miller et al. (2008). Figure 4 shows the histogram of the $70\ \mu\text{m}$ flux densities. We also overplot the population of sources with and without radio counterparts. Sources with no radio counterparts have a median $70\ \mu\text{m}$ flux density of 3.3 mJy while sources with radio counterparts have a median $70\ \mu\text{m}$ flux density of 5.2 mJy.

3.2.2. Redshift Data and Infrared Luminosities

The final $70\ \mu\text{m}$ catalog was matched to the various redshift catalogs using the matching radii shown in Table 1. Due to the better resolution of the radio data we used radio positions where available, followed by $24\ \mu\text{m}$ positions where available to match to the redshift catalogs.

Where possible, we matched spectroscopic redshifts to the sources. Where more than one spectroscopic redshift was available, we chose to prioritize ATLAS spectroscopic redshifts over GOODS spectroscopic data as we have access to the ATLAS spectra. The difference in redshift, Δz , was <0.0015 for the 12 sources in common.

Where no spectroscopic data were available we used photometric redshifts. Where both COMBO-17 and MUSYC photometric redshifts were available, we chose to prioritize MUSYC data over COMBO-17 because MUSYC's photometric redshifts were derived from 32 bands as opposed to COMBO-17, which uses 17. Furthermore, MUSYC has photometric data extending into the near-infrared (JHK), which improves the accuracy of the photometric redshift.

In summary, 562 of the 617 (91%) $70\ \mu\text{m}$ sources have redshift information, 206 (33%) of which are spectroscopic (Figure 5).

The total infrared luminosity (8–1000 μm) of the sources in our sample was estimated by fitting the $24\ \mu\text{m}$ and $70\ \mu\text{m}$ flux density to the SED templates of Chary & Elbaz (2001,

Table 1
Table of Redshift Catalogs

Catalog	Radius (arcsec)	Spec/Phot	Size	N_{matched}	N_{final}	Reference
COMBO-17	2	Phot	62337	555	122	Wolf et al. (2004)
MUSYC	1.5	Phot	59693	410	234	Cardamone et al. (2010)
MUSYC	1.5	Spec	2551	114	26	Cardamone et al. (2010)
GOODS	2	Spec	7332 (5528)	160	148	Balestra et al. (2010)
ATLAS	5	Spec	254	32	32	Mao et al. (2009)

Notes. Column 1 names the catalog while Column 2 gives the matching radius which was used to match the redshift data to the $70\ \mu\text{m}$ data. Columns 3 and 4 give the catalog size and whether the redshift data are photometric or spectroscopic and Columns 5 and 6 give the total number of sources that are matched to the $70\ \mu\text{m}$ data as well as the number that are in the final catalog. Column 7 gives the reference for the catalog. There are 562 sources with redshifts, of which 206 are spectroscopic.

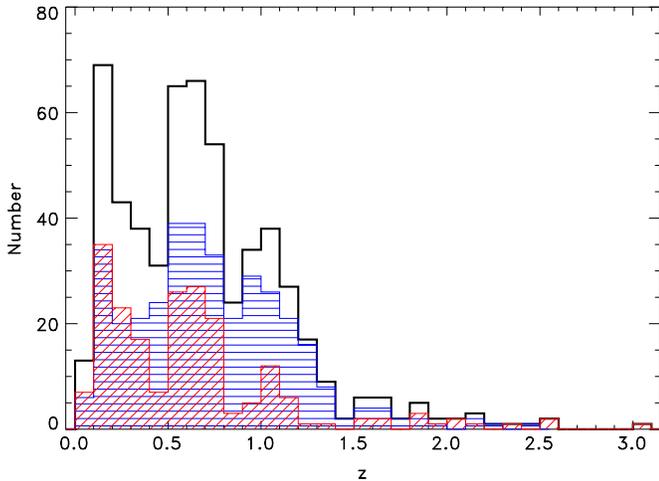


Figure 5. Histogram of all 562 sources that have redshift information. The black line is the histogram for all sources that have redshift information while the red diagonally shaded histogram shows sources with spectroscopic redshifts and the blue horizontally shaded histogram shows sources with photometric redshifts. (A color version of this figure is available in the online journal.)

hereafter CE01), while letting the templates scale in luminosity. The CE01 templates show observed $24\ \mu\text{m}/70\ \mu\text{m}$ flux density ratios, which are more representative of $z \sim 1$ galaxies than Dale & Helou (2002) or Lagache et al. (2003) templates. Magnelli et al. (2009) found that $70\ \mu\text{m}$ data seems to provide an estimate of L_{IR} that is nearly independent of the SED library used, but this is partly because $70\ \mu\text{m}$ is close to the peak emission, and therefore carries the largest fraction of FIR power.

We integrate the best fit template over 8–1000 μm to derive the total IR luminosity (L_{IR} ; Figure 6). We find luminous IR galaxies (LIRGs, $10^{11} L_{\odot} < L_{\text{IR}} < 10^{12} L_{\odot}$) are detected out to $z \sim 1.25$, while ultraluminous IR galaxies (ULIRGs, $L_{\text{IR}} > 10^{12} L_{\odot}$) are detected out to $z = 3$.

3.2.3. X-ray Counterparts

Using the 2Ms *Chandra* data (Luo et al. 2008), we find 55 of the $70\ \mu\text{m}$ sources are within 2 arcsec of an X-ray source. We use this data to calculate the hardness ratio so as to discriminate against AGNs (Section 4.1).

4. RESULTS AND ANALYSIS

We define q_{IR} as

$$q_{\text{IR}} = \log_{10} \left(\frac{S_{\text{IR}}}{S_{\text{radio}}} \right), \quad (2)$$

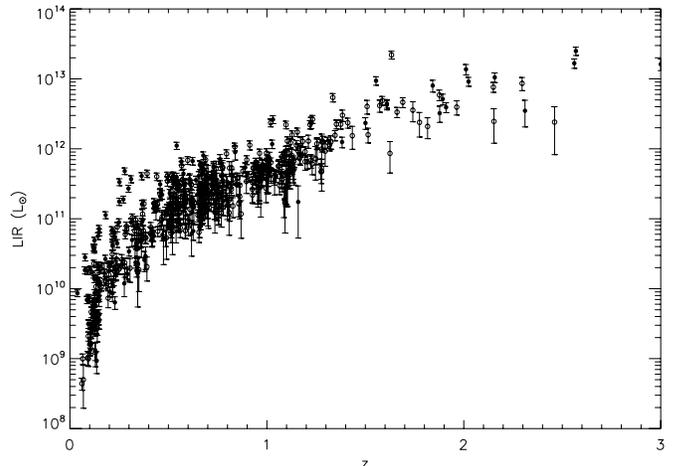


Figure 6. Total IR luminosity (8–1000 μm) as a function of redshift for all IR sources with redshift information (562/617). Sources with spectroscopic redshifts are shown with filled circles while sources with photometric redshifts are shown with open circles.

where S_{IR} is the observed infrared flux density at the specified IR wavelength (e.g., $70\ \mu\text{m}$), and S_{radio} is the observed flux density at 1.4 GHz.

4.1. AGN Identification

We wish to study the FRC of predominantly star-forming galaxies and so we removed galaxies from our sample if they satisfied any of the following four criteria which indicate AGNs.

1. The source has a radio morphology that displays the classic double-lobed AGN morphology. Figure 7 shows the three sources that were identified and subsequently removed from our sample.
2. The source has $\log(S_{70}/S_{24}) < 0.5$, because AGNs are expected to have low S_{70}/S_{24} ratios (Frayer et al. 2006). Twenty-two sources were identified as AGNs in this way. Figure 8 shows the log of the ratio of $70\ \mu\text{m}$ flux density over $24\ \mu\text{m}$ flux density, plotted against redshift.
3. The source has a hardness ratio > 0.2 based on the 2Ms *Chandra* data (Luo et al. 2008), where hardness ratio is defined as

$$\text{HR} = \frac{S_{\text{hard}} - S_{\text{soft}}}{S_{\text{hard}} + S_{\text{soft}}}, \quad (3)$$

where HR is the hardness ratio, S_{hard} is the flux density of the 2–8 keV band, S_{soft} is the flux density of the 0.5–2.0 keV band (Rosati et al. 2002). Fourteen sources were identified as AGNs in this way.

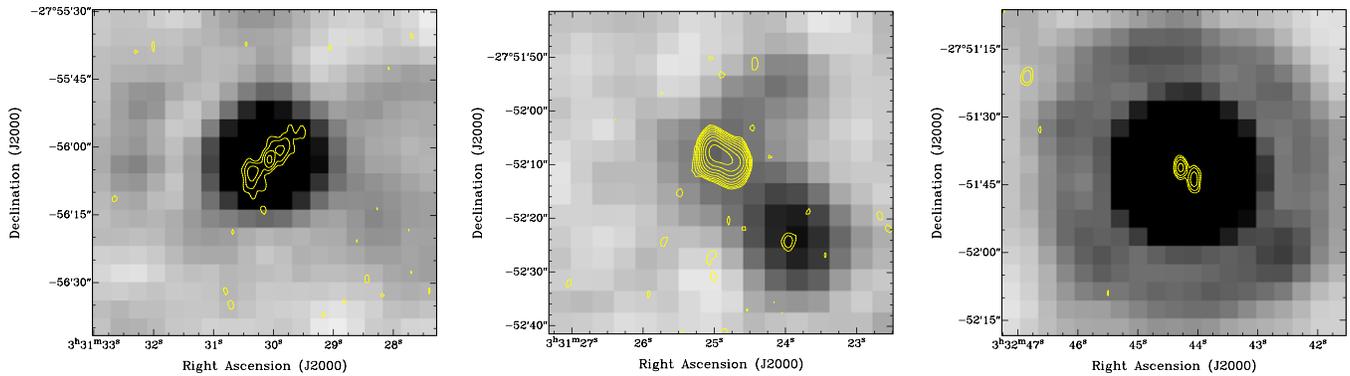


Figure 7. Sources which we classified as AGNs based on radio morphology. The grayscale image is the $70\ \mu\text{m}$ image and contours are the 1.4 GHz image starting at $24\ \mu\text{Jy}$ (~ 3 times the rms) and increasing by factors of two.

(A color version of this figure is available in the online journal.)

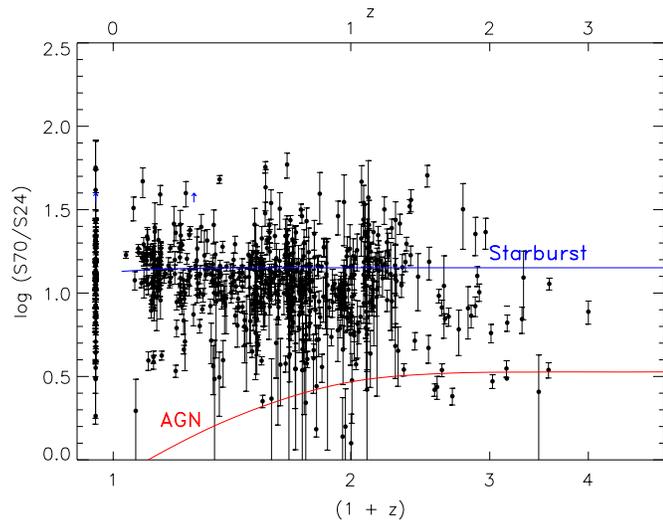


Figure 8. Ratio of $70\ \mu\text{m}$ flux density over $24\ \mu\text{m}$ flux density plotted against redshift. Sources with no redshift information are shown at an artificial redshift of -0.05 . The blue line represents a simple modified blackbody SED model for starbursts with a dust temperature of 30K with a mid-infrared slope of $\alpha = 2.4$, while the red line represents a model for AGNs with a dust temperature of 90K with a mid-infrared slope of 1.1 (Frayser et al. 2006). Error bars are derived by combining in quadrature the standard errors in radio and IR fluxes.

(A color version of this figure is available in the online journal.)

4. The source has a soft X-ray ($0.5\text{--}2\ \text{keV}$) to R -band flux density ratio of greater than 1, as used by Luo et al. (2008) to classify AGNs. Fourteen sources were identified as AGNs in this way.

Low q_{70} and q_{24} values may also be used to discriminate against AGNs (e.g., Middelberg et al. 2008). However, doing so could potentially bias our results against low values of q . Applying these criteria removed only a further five AGNs from our sample and we found made a negligible difference to our results, so we do not use these criteria here.

In total, 44 ($\sim 7\%$) sources were classified as AGNs and subsequently removed from further analysis, leaving us with a final catalog of 573 $70\ \mu\text{m}$ sources. Most of the sources were identified as AGNs based on only one of the above four diagnostics, highlighting the need for multiple AGN diagnostics. There were only six sources that were identified as AGNs by both the X-ray diagnostics. 41/44 of the sources identified as AGNs had redshift information. The median redshift of the AGNs is 0.969, slightly higher than the median redshift of the

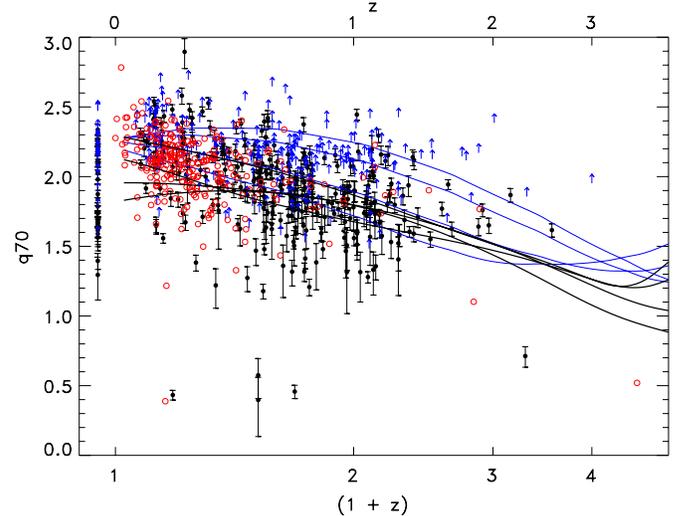


Figure 9. $70\ \mu\text{m}$ FIR–radio correlation plotted against redshift. The black points are sources that have both $70\ \mu\text{m}$ and 1.4 GHz detections while the blue arrows show the lower limit of q_{70} for sources with no radio detection. The red open circles are data from xFLS (Appleton et al. 2004). The black lines are expected q_{70} tracks derived from SED templates for galaxies with total infrared luminosities of 10^9 (normal galaxies), 10^{11} , 10^{12} , and $10^{13} L_{\odot}$ (ULIRGs) going from bottom to top, from Chary & Elbaz (2001) and the blue lines are q_{70} tracks derived from SED templates from Dale & Helou (2002). Sources with no redshift information are shown at an artificial redshift of -0.05 . Error bars are derived by combining in quadrature the standard errors in radio and IR fluxes.

(A color version of this figure is available in the online journal.)

entire sample (0.655). The median L_{IR} of the AGNs is $5.10 \times 10^{11} L_{\odot}$, also higher than the median L_{IR} of the entire sample ($2.07 \times 10^{11} L_{\odot}$). The AGNs tend to be at higher redshifts and higher L_{IR} due to Malmquist bias.

4.2. q_{70}

To explore how the FRC, as probed by the observed $70\ \mu\text{m}$ band, changes with redshift, we plot q_{70} against redshift in Figure 9. Our data are plotted in black and lower limits of q_{70} from the radio non-detections are plotted as blue arrows. The lower limits are calculated by choosing the upper limit to the radio flux density for the radio non-detections to be $3 \times \text{rms}$. Data from the Spitzer Extragalactic First Look Survey (xFLS; Appleton et al. 2004) are plotted as red open circles. It is immediately evident that our data contain many more high redshift ($z > 1$) sources than the xFLS data. Appleton et al.

Table 2Median q_{70} for Different Redshift Bins for Only Sources with Both $70\ \mu\text{m}$ and Radio Detections, and are Not Classified as AGNs

z	Median z	N_{sources}	Median q_{70}
$0 \leq z < 0.25$	0.148	53	2.20 ± 0.05
$0.25 \leq z < 0.50$	0.370	46	2.07 ± 0.04
$0.50 \leq z < 0.75$	0.626	93	1.87 ± 0.04
$0.75 \leq z < 1.00$	0.890	43	1.81 ± 0.03
$1.00 \leq z < 1.50$	1.13	52	1.80 ± 0.04
$z \geq 1.50$	1.91	10	1.65 ± 0.11
No redshift	...	26	1.79 ± 0.06

Note. There are a total of 323 sources.

(2004) found a q_{70} value of 2.15 ± 0.16 for their sample. The data points that are at low q_{70} values are likely to be AGNs.

We also overplot the expected q_{70} ratio as a function of redshift from IR SED templates of Dale & Helou (2002, DH02) and Chary & Elbaz (2001, CE01). The CE01 SED templates are obtained from low-redshift galaxies with appropriate luminosities. We use four CE01 templates which range in L_{IR} from 10^9 (normal galaxy) to $10^{13} L_{\odot}$ (ULIRGs). The DH02 SED templates are derived from combining theoretical SEDs from dust emission of individual regions within a galaxy. We use four DH02 templates with $\alpha = 1, 1.5, 2,$ and 2.5 , where α represents the relative contributions of the different dust emission SEDs. The overplotted q_{70} tracks are derived using the IR SED templates, while the radio flux is derived from the infrared flux assuming both thermal and non-thermal components and a constant value of q .

In order to quantify the evolution of q_{70} , we binned all data with both $70\ \mu\text{m}$ and radio detections in redshift and determined the median q_{70} in each redshift bin. The results are shown in Table 2 as well as in the black points of Figure 10. We calculate a median q_{70} of 2.16 ± 0.03 out to $z = 0.5$, which is in agreement with Appleyton et al. (2004), as well as other previous studies (e.g., Seymour et al. 2009). Sources with no redshift information have a mean q_{70} ratio similar to $z \sim 1$ sources.

Figure 10 shows no evidence for evolution in the FRC because the data points are consistent with the overplotted q_{70} tracks. At $z > 1.25$ we are detecting only ULIRGs (see Figure 6), and this is reflected in the higher median q_{70} values, which are better traced by the q_{70} track derived from the ULIRG SED template.

4.3. Survival Analysis

Our $70\ \mu\text{m}$ selected sample has sources with both $70\ \mu\text{m}$ and 1.4 GHz detections, as well as sources that do not have 1.4 GHz

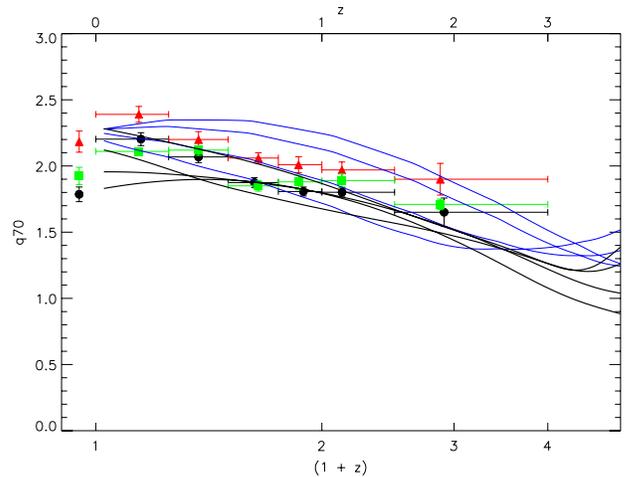


Figure 10. Median q values for different redshift bins. The black circles show the median q_{70} values for sources with both $70\ \mu\text{m}$ and 1.4 GHz detections, the red triangles show the median q_{70} value for all sources taking into account the lower limits using survival analysis and the green squares show the median q_{70} value after stacking. The tracks derived from the SED templates are the same as in Figure 9. Vertical error bars denote standard errors and horizontal error bars indicate the range of the redshift bin.

(A color version of this figure is available in the online journal.)

detections. To study the full IR sample, we have to include radio non-detections. As we know the rms of the 1.4 GHz data we can estimate an upper limit to the radio flux density, which translates to a lower limit for q_{70} . We choose the upper limit to the radio flux density for the radio non-detections to be $3 \times \text{rms}$. One method of including limits from non-detections is using a branch of statistics called survival analysis. Survival analysis is an extensive field of statistics and was first applied to astronomy by Feigelson & Nelson (1985). We perform survival analysis on our data using the ASURV (Astronomical SURVival analysis) package developed by Lavalley et al. (1992).

We took the 50th percentile of the Kaplan–Meier estimator to obtain a median q_{70} value for each redshift bin. The results are shown in Table 3, as well as in Figure 10. The two highest redshift bins appear to have higher median q_{70} values, but this is because at $z > 1.25$ we are detecting only ULIRGs, and hence these two points lie along the most actively star-forming, or ULIRG track.

Our q_{70} values appear to agree, within the errors, to the q_{70} tracks derived from the empirical SED templates from the local universe. This is evidence for little, if any evolution in q_{70} , which implies that galaxies at high redshifts ($z \sim 2.5$) share many of the same properties as galaxies in the local universe.

Table 3Median q_{70} for Different Redshift Bins, Calculated Using Survival Analysis, for All $70\ \mu\text{m}$ Sources that are Not Classified as AGNs

z	Median z	N_{sources}	Median q_{70} (3σ)	q_{70} (DH02)	q_{70} (CE01)
$0 \leq z < 0.25$	0.141	107	2.39 ± 0.06	2.11–2.32	1.87–2.22
$0.25 \leq z < 0.50$	0.370	80	2.20 ± 0.06	2.00–2.35	1.90–2.13
$0.50 \leq z < 0.75$	0.646	154	2.06 ± 0.04	1.88–2.33	1.88–2.01
$0.75 \leq z < 1.00$	0.862	75	2.01 ± 0.06	1.77–2.28	1.83–1.92
$1.00 \leq z < 1.50$	1.13	85	1.97 ± 0.06	1.66–2.20	1.75–1.81
$z \geq 1.50$	1.88	20	1.90 ± 0.12	1.41–1.92	1.48–1.55
No redshift	...	52	2.18 ± 0.08		

Notes. Survival analysis was performed for radio flux density limits of $3 \times \text{rms}$. The range of values given for the q_{70} values derived from model SED templates of DH02 are for $\alpha = 2.5$ (first value) and $\alpha = 1$. The range of values given for the model q_{70} values derived from SED templates of CE01 are for normal galaxies and ULIRGs. There are a total of 573 sources.

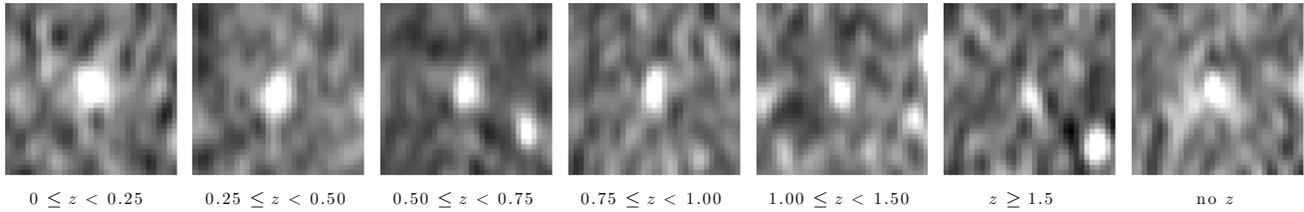


Figure 11. Postage stamps of the radio stacks of the IR sources, for the redshift bins as shown. The stacks are 16×16 arcsec in size. There is significant radio flux at the center of the stack in all of the redshift bins. The measured flux densities are reported in Table 4.

Table 4
Summary of Radio Stacking Results of IR Sources

z	Median z	N_{det}	N_{stack}	Stacked $S_{1.4\text{GHz}}$ (μJy)	Average q_{70}
$0 \leq z < 0.25$	0.141	53	54	27.7 ± 2.8	$2.11^{+0.02}_{-0.03}$
$0.25 \leq z < 0.50$	0.370	46	34	25.4 ± 1.9	$2.12^{+0.03}_{-0.04}$
$0.50 \leq z < 0.75$	0.646	93	61	19.7 ± 1.5	$1.85^{+0.04}_{-0.04}$
$0.75 \leq z < 1.00$	0.862	43	32	18.9 ± 1.6	$1.88^{+0.05}_{-0.06}$
$1.00 \leq z < 1.50$	1.13	52	33	13.6 ± 0.9	$1.89^{+0.05}_{-0.05}$
$z \geq 1.50$	1.88	10	10	19.5 ± 2.5	$1.71^{+0.04}_{-0.04}$
No redshift	...	26	26	46.8 ± 9.4	$1.93^{+0.06}_{-0.07}$

Notes. N_{det} is the number of IR sources in each bin, which are detected in the radio image. N_{stack} is the number of IR sources not detected in the radio image and therefore the number stacked in this analysis, resulting in the stacked $S_{1.4\text{GHz}}$ listed in column 5. Average q_{70} is the FIR–radio flux density ratio for all IR sources in the redshift bin, obtained by combining the detections with the stacked results.

4.4. Stacking

To extend the results, we performed a stacking analysis of the IR sources by stacking the radio sources in bins of redshift. We started by stacking the M08 radio data at the positions of all sources not detected in the radio image, using a noise-weighted mean, for which the noise was determined from the radio image in the vicinity of the IR source. The stacked radio flux density was then obtained from a Gaussian fit to the signal at the center of the radio stack, leaving all Gaussian parameters unconstrained. The resulting stacks are shown in Figure 11, and each has a significant detection.

To derive an average q_{70} for all IR sources in the various redshift bins (green points in Figure 10), we calculated a mean by adding the flux densities of the detected sources to the stacked flux density. Each was added with a weight = $1/n$ (where n is the total number of measurements, including both stacked and detected images) and the stacked flux was weighted by m/n where m is the number of stacked images. The results are summarized in Table 4. The higher redshift bins have higher q values when compared to the survival analysis results because stacking analysis imposes lower limits on the radio flux densities. As before, the FIR–radio flux density ratios appear to follow the ratios expected for local SEDs and hence we find little evolution in the FRC as a function of time.

4.5. q_{TIR}

We computed the ratio of total infrared luminosity over rest-frame radio luminosity to confirm our results from Sections 4.2–4.4. Furthermore, this allows us to make direct comparisons with other studies that compute the FRC in this way.

We computed q_{TIR} for all 521 sources that had redshift information (and were not identified as AGNs) using

Table 5
Median q_{TIR} for Different Redshift Bins for Only Sources with Both $70 \mu\text{m}$ and Radio Detections, and are Not Classified as AGNs

z	Median z	N_{sources}	Median q_{TIR}
$0 \leq z < 0.25$	0.148	53	2.54 ± 0.05
$0.25 \leq z < 0.50$	0.370	46	2.52 ± 0.04
$0.50 \leq z < 0.75$	0.626	93	2.41 ± 0.03
$0.75 \leq z < 1.00$	0.890	43	2.35 ± 0.04
$1.00 \leq z < 1.50$	1.13	52	2.33 ± 0.03
$z \geq 1.50$	1.91	10	2.52 ± 0.11

Note. There are a total of 297 sources.

Table 6
Median q_{TIR} for Different Redshift Bins, Calculated Using Survival Analysis, for All $70 \mu\text{m}$ Sources that are Not Classified as AGNs

z	Median z	N_{sources}	Median q_{TIR} (3σ)
$0 \leq z < 0.25$	0.141	107	2.74 ± 0.06
$0.25 \leq z < 0.50$	0.370	80	2.70 ± 0.06
$0.50 \leq z < 0.75$	0.646	154	2.60 ± 0.04
$0.75 \leq z < 1.00$	0.862	75	2.53 ± 0.06
$1.00 \leq z < 1.50$	1.13	85	2.55 ± 0.06
$z \geq 1.50$	1.88	20	2.75 ± 0.12

Notes. Survival analysis was performed using $L_{1.4\text{GHz}}$ calculated from radio flux density limits of $3 \times \text{rms}$. There are a total of 521 sources.

$$q_{\text{TIR}} = \log \left(\frac{L_{\text{IR}}}{3.75 \times 10^{12} \text{W}} \right) - \log \left(\frac{L_{1.4\text{GHz}}}{\text{WHz}^{-1}} \right), \quad (4)$$

where L_{IR} is the total infrared luminosity (Section 3.2.2), and $L_{1.4\text{GHz}}$ is the rest-frame 1.4 GHz luminosity. In order to account for k -correction in the radio, we set the spectral index, α , to 0.8,⁹ following the work of Ibar et al. (2010) and Sargent et al. (2010b).

Figure 12 shows the distribution of q_{TIR} as a function of redshift. For sources without a radio detection, we set an upper limit to the radio flux of $3 \times \text{rms}$. Similar to Figure 9, sources without a radio detection (i.e., lower limits) tend to have higher q_{TIR} when compared to sources with a radio detection. In order to take into account the effect of the radio non-detections, we bin our data in redshift and perform survival analysis (Tables 5 and 6). Figure 13 shows the median q_{TIR} for only the detected sources (black circles), and for all sources by using survival analysis (red triangles). We have included results from previous studies in Figure 13. The blue star at $z = 0$ is at $q_{\text{TIR}} = 2.64$, the mean q_{TIR} found by Bell (2003). The open upside-down triangles are q_{TIR} derived from stacking analyses by Bourne et al. (2011) and the open squares are the “uncorrected” q_{TIR} for star-forming galaxies from Sargent et al. (2010b) (see Section 4.5.1 for a discussion on the Kellermann Correction).

⁹ $S_{\nu} \propto \nu^{-\alpha}$, where S_{ν} is the flux density at frequency ν .

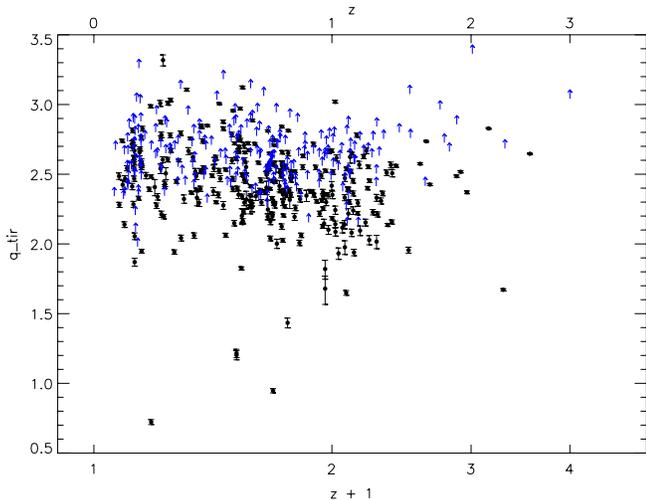


Figure 12. Total infrared luminosity FRC plotted against redshift. The same symbol and color scheme is used as for Figure 9.

(A color version of this figure is available in the online journal.)

Our q_{TIR} values are all within ~ 0.22 of each other. Our values agree, within the errors, with the work of Sargent et al. (2010b), and also agree with the work of Bourne et al. (2011), with the exception of the redshift bins $0.75 < z < 1$ and $z > 1.5$, where our values differ by $< 2\sigma$ compared to the values given by Bourne et al. (2011) for similar redshift ranges. Sargent et al. (2010b) use $3 \times \text{rms}$ to determine the upper limit of the radio flux density for their data set. Unlike the work by Bourne et al. (2011) and Sargent et al. (2010b), our q_{TIR} values do not consistently increase or decrease over the redshift range studied, although we do see marginally significant evidence for a decreasing q_{TIR} over the redshift range 0–1.5, followed by an increase from 1.5 to 3. However, Figure 13 shows that the data points that use survival analysis (which are more reliable than those that do not) are statistically consistent with a horizontal line at about $q_{\text{TIR}} = 2.62$, and so we conclude that these apparent variations of q_{TIR} with redshift may not be statistically significant.

Bourne et al. (2011) find their q_{TIR} values are systematically higher than the median q_{TIR} found by Bell (2003) of 2.64 ± 0.02 , except at $z > 1$ where it appears to decline. They suggest the apparent decline may be attributed to the assumptions they make about spectral indices. Sargent et al. (2010b) find that their q_{TIR} values are constant with redshift when they apply a correction (see Section 4.5.1), otherwise ~ 0.3 dex of evolution is found. While Bourne et al. (2011) also uses FIDEL data in the ECDFS, their sample is IRAC selected whereas our sample is $70 \mu\text{m}$ selected. Sargent et al. (2010b) use a $24 \mu\text{m}$ selected sample to study the q_{TIR} in the COSMOS field.

There are a number of factors that may contribute to uncertainties in q_{TIR} . The predominant source of uncertainty is in the estimation of L_{IR} as we are fitting the $24 \mu\text{m}$ and $70 \mu\text{m}$ fluxes to SED templates. The uncertainty from the choice of model SEDs can add 0.2–0.3 dex (Magnelli et al. 2009; Le Floch et al. 2005). Uncertainties also arise from the rest-frame radio luminosity where we are assuming $\alpha = 0.8$.

Jarvis et al. (2010) suggest that the slight upturn in q_{IR} seen at low redshifts ($z < 0.5$) in high redshift studies of the FRC (e.g., Bourne et al. 2011) may be attributed to the resolving of extended structure, which means radio flux would be missed and hence increase the FRC. Our radio data has a $2''.8 \times 1''.6$ beam, so if the Jarvis hypothesis is correct, we would expect to see this effect at low redshifts.

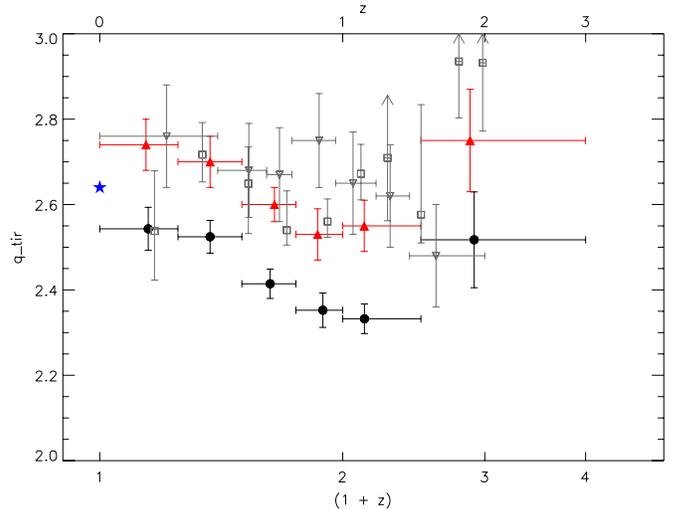


Figure 13. Median q_{TIR} values for different redshift bins. The black filled circles show the median q_{TIR} values for sources with both $70 \mu\text{m}$ and 1.4 GHz detections and the red triangles show the median q_{TIR} value for all sources taking into account the lower limits using survival analysis. The gray upside-down open triangles show the median q_{TIR} values derived by Bourne et al. (2011) and the gray open squares show the median q_{TIR} values derived by Sargent et al. (2010b) for star-forming galaxies. None of the data shown include the Kellermann correction (Section 4.5.1). Vertical error bars are standard errors for both our data set and Bourne’s data, but the vertical error bars for Sargent’s data are upper and lower 95% confidence levels. The blue star at $(1+z) = 1$ represents the median $q_{\text{TIR}} = 2.64 \pm 0.02$ from Bell (2003).

(A color version of this figure is available in the online journal.)

Although we see a slight upturn in q_{TIR} from $z = 1$ to $z = 0$, this increase has a roughly constant gradient, implying that resolution effects are important as high as $z = 1$. However, at $z = 1$ a typical star-forming galaxy of diameter 10 kpc is unresolved by our beam. Thus, the increasing q_{TIR} toward low redshifts is probably not due to this effect.

4.5.1. The Kellermann Correction

Kellermann (1964) showed that flux ratios (such as q) or spectral indices of a flux-limited sample are biased by a factor which depends sensitively on the scatter in the flux ratio and on the source intensity distribution, and in particular on the power law index β of the differential source counts (i.e., $dN/dS \propto S^{-\beta}$). Other formulations of the same effect are given by Condon (1984), Francis (1993), and Lauer et al. (2007) and were first noted as being relevant to the evolution of q by Sargent et al. (2010a).

Sargent et al. (2010b) assume a Euclidean ($\beta = 2.5$) source intensity distribution at $z < 1.4$ and a sub-Euclidean ($\beta \sim 1.5$) at $1.4 < z < 2$ to derive a correction of 0.22 in their value of q at $z > 1.4$. However, we note that (1) the expression assumes a common value of β for both the radio and infrared source counts, and (2) the value of β is very uncertain at low flux densities. For example, the median radio flux density for our detected sources at $z < 1.4$ is $\sim 72 \mu\text{Jy}$, at which the radio source count power law index lies in the range ~ 1.5 – 2.5 (e.g., Huynh et al. 2005), leading to a correction factor between 0.22 (the adopted value of Sargent, which gives no evolution of q with z) and 0.51 (which implies that q decreases with z with marginal significance).

Because of this uncertainty in the value of the Kellermann correction for faint sources such as those studied here and for consistency with other authors (e.g., Appleton et al. 2004; Seymour et al. 2009; Bourne et al. 2011), we do not apply it to

our data but note that it is responsible for a further uncertainty in the slope of q as a function of redshift. The range of possible Kellermann corrections (which depend strongly on the assumed value of β) to the variation of q over our redshift range is roughly centered on zero, with a total spread of about ± 0.3 .

5. SUMMARY AND CONCLUSIONS

We have studied the FRC out to $z > 2$ of ULIRGs in ECDFS. Our results for q_{70} showed that they could be broadly described by q_{70} tracks derived from the SED templates of Chary & Elbaz (2001) and Dale & Helou (2002). To quantify the evolution of q_{70} , we binned our data in redshift and determined the median q_{70} using both survival analysis and a stacking analysis. Both survival analysis and the stacking analysis gave similar results. We see no clear evidence for evolution in the FRC at $70 \mu\text{m}$.

We also calculate the FRC using L_{IR} and $L_{1.4\text{GHz}}$ and find that evolution in q_{TIR} is constrained within 0.22. Our calculated q_{TIR} appears slightly lower than previous studies, but this may merely be due to uncertainties involved in calculating q_{TIR} . We also acknowledge the importance of the Kellermann correction, but due to uncertainties in the value we do not apply it to our data.

A lack of evolution in the FRC is surprising because it implies that the myriad of effects on which the FRC rely must either also not evolve, or must evolve in such a way so as to preserve the FRC. Effects such as inverse Compton cooling of the electrons (Murphy 2009), evolution of the magnetic field strength, and evolution in the SEDs due to dust and metallicity changes should all affect the FRC.

We therefore conclude that either all these factors are insignificant at $z \sim 2$ or there is a complex interplay between these factors conspiring in the preservation of the FRC at high redshifts.

Early science results from *Herschel* have already hinted that the FRC shows no evidence for evolution to $z = 2$ (Jarvis et al. 2010; Ivison et al. 2010b). More results on the FRC from *Herschel* are expected over the next few years as observations are completed and the data analyzed. *Herschel* will measure the far-infrared properties of normal galaxies to $z \sim 1$ and ULIRGs out to $z \sim 4$. Consequently, we will be able to study the FRC out to higher redshifts and hence gain a better understanding of the evolution of star-forming galaxies, especially in the high redshift universe.

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