

COSMOLOGICAL CONSTRAINTS FROM SUNYAEV-ZEL'DOVICH-SELECTED CLUSTERS WITH X-RAY OBSERVATIONS IN THE FIRST 178 SQUARE DEGREES OF THE SOUTH POLE TELESCOPE SURVEY

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ABSTRACT

We use measurements from the South Pole Telescope (SPT) Sunyaev Zel'dovich (SZ) cluster survey in combination with X-ray measurements to constrain cosmological parameters. We present a statistical method that fits for the scaling relations of the SZ and X-ray cluster observables with mass while jointly fitting for cosmology. The method is generalizable to multiple cluster observables, and self-consistently accounts for the effects of the cluster selection and uncertainties in cluster mass calibration on the derived cosmological constraints. We apply this method to a data set consisting of an SZ-selected catalog of 18 galaxy clusters at $z > 0.3$ from the first 178 deg² of the 2500 deg² SPT-SZ survey, with 14 clusters having X-ray observations from either *Chandra* or *XMM-Newton*. Assuming a spatially flat Λ CDM cosmological model, we find the SPT cluster sample constrains $\sigma_8(\Omega_m/0.25)^{0.30} = 0.785 \pm 0.037$. In combination with measurements of the CMB power spectrum from the SPT and the seven-year WMAP data, the SPT cluster sample constrains $\sigma_8 = 0.795 \pm 0.016$ and $\Omega_m = 0.255 \pm 0.016$, a factor of 1.5 improvement on each parameter over the CMB data alone. We consider several extensions beyond the Λ CDM model by including the following as free parameters: the dark energy equation of state (w), the sum of the neutrino masses (Σm_ν), the effective number of relativistic species (N_{eff}), and a primordial non-Gaussianity (f_{NL}). We find that adding the SPT cluster data significantly improves the constraints on w and Σm_ν beyond those found when using measurements of the CMB, supernovae, baryon acoustic oscillations, and the Hubble constant. Considering each extension independently, we best constrain $w = -0.973 \pm 0.063$ and the sum of neutrino masses $\Sigma m_\nu < 0.28$ eV at 95% confidence, a factor of 1.25 and 1.4 improvement, respectively, over the constraints without clusters. Assuming a Λ CDM model with a free N_{eff} and Σm_ν , we measure $N_{\text{eff}} = 3.91 \pm 0.42$ and constrain $\Sigma m_\nu < 0.63$ eV at 95% confidence. We also use the SPT cluster sample to constrain $f_{NL} = -192 \pm 310$, consistent with zero primordial non-Gaussianity. Finally, we discuss the current systematic limitations due to the cluster mass calibration, and future improvements for the recently completed 2500 deg² SPT-SZ survey. The survey has detected ~ 500 clusters with a median redshift of ~ 0.5 and a median mass of $\sim 2.3 \times 10^{14} M_\odot/h$ and, when combined with an improved cluster mass calibration and existing external cosmological data sets will significantly improve constraints on w .

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1. INTRODUCTION

Clusters of galaxies are the most massive collapsed objects in the universe. Their abundance is sensitive to multiple cosmological parameters, in particular the matter density, the amplitude of the matter power spectrum, and the dark energy equation of state (e.g., Wang & Steinhardt 1998; Haiman et al. 2001; Holder et al. 2001). Measurements of the cluster abundance that extend to higher redshifts become sensitive to dark energy through its effect on the growth of structure. This makes cluster abundance measurements an important systematic test of the standard dark energy paradigm, because they are affected by dark energy in a fundamentally different way than distance-redshift based tests, such as from type Ia supernovae and baryon acoustic oscillations. For the same reason, cluster abundance measurements also constrain different cosmological parameter combinations than distance-based tests, and their combination can break parameter degeneracies and achieve tighter constraints than either method alone (e.g., Linder & Jenkins 2003).

Recently there has been significant theoretical and experimental progress in efforts to use clusters as cosmological probes. Large-volume numerical simulations have calibrated a “universal” cluster mass function over a broad range of cosmologies at a level better than current experimental uncertainties (e.g., Jenkins et al. 2001; Warren et al. 2006; Tinker et al. 2008; Bhattacharya et al. 2011). Numerical simulations have also led to a better understanding of systematic biases in cluster mass estimates derived from a broad range of cluster observables (e.g., Kravtsov et al. 2006; Jeltema et al. 2008; Stanek et al. 2010; Becker & Kravtsov 2011). Measurements of the cluster abundance using optical, X-ray, and SZ selection methods have been used to place competitive

constraints on cosmology and dark energy parameters (e.g., Vikhlinin et al. 2009b; Mantz et al. 2010c; Rozo et al. 2010; Vanderlinde et al. 2010; Sehgal et al. 2010). Currently, the most precise dark energy constraints from clusters are derived from X-ray selected samples which use the X-ray emission from the hot intra-cluster gas as a tracer of the total mass in the cluster. X-ray observables, particularly the gas mass and inferred pressure, tend to correlate with cluster mass with low scatter, independent of the dynamical state of the cluster or the details of non-gravitational physics in clusters (e.g., Kravtsov et al. 2006).

Hot intra-cluster gas also causes a spectral distortion in the cosmic microwave background (CMB) in the direction of clusters from inverse Compton scattering, a phenomenon known as the Sunyaev-Zel’dovich (SZ) effect (Sunyaev & Zel’dovich 1972). The surface brightness of the SZ effect is redshift-independent and largest at mm-wavelengths. The integrated SZ effect from a cluster is effectively measuring the cluster pressure, and is an observable that is expected to have comparably low scatter with mass to the best X-ray observables (Nagai et al. 2007; Shaw et al. 2008; Stanek et al. 2010). Therefore, a mm-wavelength SZ survey with sufficient angular resolution is expected to provide clean, mass-limited catalogs out to high redshift, probing the regime where the cluster abundance is most sensitive to dark energy’s effect on the growth rate of structure (Carlstrom et al. 2002).

Recently, the first SZ cluster catalogs from three surveys have been released: the South Pole Telescope (SPT, Staniszewski et al. 2009; Vanderlinde et al. 2010; Williamson et al. 2011), the Atacama Cosmology Telescope (ACT, Marriage et al. 2011), and the *Planck* satellite (Planck Collaboration et al. 2011). However, even with only ~ 10 -20 clusters, the cosmological constraints from these surveys have been limited by the systematic uncertainty in the cluster mass calibration (Vanderlinde et al. 2010; Sehgal et al. 2011). X-ray surveys (Vikhlinin et al. 2009b; Mantz et al. 2010c) have achieved tighter constraints by adopting variations of the following calibration strategy: calibrating X-ray observable-mass relations using X-ray hydrostatic mass estimates of relaxed clusters, applying this calibration to a larger sample of relaxed and unrelaxed clusters, and verifying the overall mass calibration from other methods, particularly from weak lensing measurements. In this work, we apply a similar strategy to the SPT-SZ survey using the cluster sample from Vanderlinde et al. (2010) (hereafter V10), by incorporating an externally calibrated X-ray observable-mass relation and X-ray measurements of the V10 sample in order to present improved cosmological constraints.

This paper is organized as follows. In Section 2, we describe the relevant SZ, X-ray, and optical data, analysis methods, and the external cosmological data sets used in this work. In Section 3, we describe and implement a self-consistent cosmological analysis using SZ and X-ray observations of the SPT cluster sample that simultaneously constrains cosmology and the relevant SZ and X-ray cluster scaling relations while accounting for the SPT cluster selection function. In Section 4, we discuss the constraints on a Λ CDM cosmological model from the SPT cluster sample, and compare our results to the constraints from observations of the CMB power spectrum. In Section 5, we consider extensions to the Λ CDM model

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by including the following as free parameters: dark energy equation of state, the sum of the neutrino masses, the effective number of relativistic species, and a primordial non-Gaussianity. We report the relative improvements using the SPT data to constrain each extension. In Sections 6 and 7, we discuss the limiting systematics and implications for applying this method to the larger SPT cluster sample.

In this paper, unless otherwise specified, the cluster mass will refer to M_{500} , the mass enclosed within a spherical radius, r_{500} , where the cluster’s mean matter density is 500 times the critical density of the universe at the observed cluster redshift. The critical density is $\rho_{crit}(z) = 3H^2(z)/8\pi G$, where $H(z)$ is the Hubble parameter.

2. DATA AND OBSERVATIONS

2.1. Cluster Data and Observations

The cluster sample used in this work is a sub-sample of a SZ-selected catalog from the SPT that was described in V10. The V10 catalog consisted of 21 clusters selected by their SZ significance from 178 deg² of sky surveyed by the SPT in 2008. As in V10, we use only the 18 clusters at $z > 0.3$ for the cosmological results in this work. The optical and X-ray properties of this catalog have been described previously in High et al. (2010) and Andersson et al. (2011), hereafter H10 and A11, respectively. In this section, we summarize the V10, H10, and A11 data sets, analysis, and results used in this work. We also report additional spectroscopic redshift and X-ray measurements for several clusters.

2.1.1. SZ Observations and the Cluster Sample

The 10-meter diameter SPT is a mm-wavelength telescope designed to conduct a large-area survey with low noise and ~ 1 arcminute angular resolution. The SPT receiver consists of a 960 element bolometer array that is sensitive in three bands, at 95, 150, and 220 GHz. Details of the telescope and receiver can be found in Padin et al. (2008), Carlstrom et al. (2011), and Dobbs et al. (2011). The primary goal of the SPT survey is to search for clusters of galaxies via the SZ effect in a 2500 deg² survey that was completed in November 2011.

The first cosmological constraints from the SPT cluster survey were reported in V10, with an accompanying cluster catalog. These results were derived from SPT 150 GHz observations of 178 deg² observed in 2008, from two approximately equal area fields centered at right ascension (R.A.) 5^h30^m , declination (decl.) -55° and R.A. 23^h30^m , decl. -55° . Cluster candidates were identified in the SPT maps by using a matched spatial filter technique (Haehnelt & Tegmark 1996; Melin et al. 2006). In brief, the SPT maps are filtered in Fourier space to optimize the detection of cluster-like objects using a source template constructed from a β -model of variable angular size. This is done while accounting for the expected signals from the dominant sources of astrophysical contamination, instrumental and atmospheric noise, and the effects of the SPT beam and timestream filtering. Candidate galaxy clusters were assigned an SZ significance, ξ , defined as the highest signal-to-noise across all filter scales.

V10 used simulations to characterize the SZ selection

function and the scaling between ξ and cluster mass. Simulated SZ maps were generated from large-volume dark matter simulations (Shaw et al. 2009) using the semi-analytic gas model of Bode et al. (2007). The gas model was calibrated to match the observed X-ray scaling relations for low-redshift ($z < 0.25$) clusters. The cluster selection was characterized by applying the matched filter to multiple sky realizations that included the dominant astrophysical components (primary and lensed CMB, thermal SZ, and point sources), instrumental and atmospheric noise, and the SPT filtering. These simulations found that at $\xi > 5$, the SPT catalog was expected to be $\sim 95\%$ pure. This result is consistent with optical follow-up which found optical cluster counterparts to 21 of the 22 candidates above this threshold. The 21 optically confirmed clusters had a median redshift of $z = 0.74$, and the sample was predicted to be nearly 100% complete above a mass threshold of $M_{500} \sim 6 \times 10^{14} h^{-1} M_\odot$ at $z = 0.6$. The simulations were also used to put conservative priors on the ξ -mass relation in the V10 cosmological analysis. Even with only 18 clusters, the improvement in the cosmological constraints was limited by the assumed systematic uncertainty on the normalization of the ξ -mass relation.

The full cluster catalog used in this work is given in Table 1. For each cluster, we report the name, position, redshift, and the SZ and X-ray observables, where the latter assumes a default cosmology. We note that the only SZ product needed for the cosmological analysis described in Section 3 is the SZ observable ξ . In this work we improve the cosmological constraints relative to V10 by reducing the uncertainty on the ξ -mass relation through inclusion of X-ray observables which have an observable-mass relation that has been externally calibrated, as described in Vikhlinin et al. (2009a) and summarized in Section 3.1.1.

2.1.2. Optical Redshifts

Redshifts of the SPT clusters were measured through a combination of optical photometry and spectroscopy. The majority of the observations and data analyses are identical to those in H10, to which we refer the reader for a more detailed description. Relative to H10, we include spectroscopic redshift measurements for seven additional clusters, which we briefly describe here. All cluster redshifts are given in Table 1.

Optical counterparts and photometric redshifts were measured from a combination of imaging from the Blanco Cosmology Survey (BCS, see Ngeow et al. 2009) and targeted observations using the Magellan telescopes. Optical images were searched for red sequence galaxies within a $2'$ radius of the SPT candidate location. A cluster was identified through an excess of red sequence galaxies relative to the background, and the photometric redshift was estimated by fitting a red sequence model. The redshift uncertainty varies over the sample, however it is typically $\Delta z/(1+z) \sim 0.03$.

For 15 of the 18 clusters, we have also measured spectroscopic redshifts, which we use for the cluster’s redshift when measured. For eight of the clusters we use the spectroscopic redshifts as reported in H10, which were measured using the Low Dispersion Survey Spectrograph (LDSS3) on the Magellan Clay 6.5-m telescope. For SPT-CL J0546-5345, we use the redshift reported in Brodwin et al. (2010), measured using multi-slit spec-

TABLE 1
THE SPT 178 deg² CLUSTER CATALOG AND OBSERVABLES

Object Name	R.A. (deg)	decl. (deg)	Photo-z	Spec-z	ξ	Y_X (10 ¹⁴ M_\odot keV)
SPT CL J0509-5342	77.336	-53.705	0.47(4)	0.463	6.61	4.3 ± 0.8
SPT-CL J0511-5154	77.920	-51.904	0.74(5)	-	5.63	-
SPT-CL J0521-5104	80.298	-51.081	0.72(5)	-	5.45	-
SPT-CL J0528-5300	82.017	-53.000	0.75(5)	0.765	5.45	1.6 ± 0.5 ^b
SPT-CL J0533-5005	83.398	-50.092	0.83(5)	0.881	5.59	1.0 ± 0.4 ^b
SPT-CL J0539-5744	85.000	-57.743	0.77(5)	-	5.12	-
SPT-CL J0546-5345	86.654	-53.761	1.16(6)	1.067 ^a	7.69	4.8 ± 0.8 ^b
SPT-CL J0551-5709	87.902	-57.156	0.41(4)	0.423	6.13	1.9 ± 0.4 ^b
SPT-CL J0559-5249	89.925	-52.826	0.66(4)	0.611	9.28	6.4 ± 0.8
SPT-CL J2301-5546	345.469	-55.776	0.78(5)	0.748 ^a	5.19	-
SPT-CL J2331-5051	352.958	-50.864	0.55(4)	0.571	8.04	3.5 ± 0.6
SPT-CL J2332-5358	353.104	-53.973	0.32(3)	0.403 ^a	7.30	6.1 ± 0.8 ^b
SPT-CL J2337-5942	354.354	-59.705	0.77(5)	0.781	14.94	8.5 ± 1.7
SPT-CL J2341-5119	355.299	-51.333	1.03(5)	0.998	9.65	4.7 ± 1.0
SPT-CL J2342-5411	355.690	-54.189	1.08(6)	1.074 ^a	6.18	1.4 ± 0.3 ^b
SPT-CL J2355-5056	358.955	-50.937	0.35(4)	0.320 ^a	5.89	2.2 ± 0.4 ^b
SPT-CL J2359-5009	359.921	-50.160	0.76(5)	0.774 ^a	6.35	1.8 ± 0.4 ^b
SPT-CL J0000-5748	0.250	-57.807	0.74(5)	0.701 ^a	5.48	4.2 ± 1.6 ^b

NOTE. — ξ is the maximum signal-to-noise of the SPT-detection obtained over the set of filter scales for each cluster. The cluster positions in R.A. and decl. are given in degrees and refer to the center of the SZ brightness in the SPT map filtered at the preferred scale to maximize the signal-to-noise. We give the estimated photometric redshift and spectroscopic redshifts, where available. To be consistent with A11, Y_X is calculated assuming a preferred Λ CDM cosmology using WMAP7+BAO+ H_0 data with $\Omega_M = 0.272$, $\Omega_\Lambda = 0.728$ and $H_0 = 70.2$ km s⁻¹ Mpc⁻¹ (Komatsu et al. 2011). In §4 and §5, Y_X is recalculated as a function of cosmology and scaling relations for each step in the Markov chain.

^aNew spectroscopic redshift since V10.

^bUpdated Y_X since A11.

troscopy with the Inamori Magellan Areal Camera and Spectrograph (IMACS) on the Magellan Baade 6.5-m telescope. Finally, there are six clusters that have new spectroscopic redshifts, which we report in this work for the first time in Table 1. These redshifts were measured with a combination of IMACS and GMOS on Gemini South, and the details of the data and analysis will be described in Bazin et al. (2012, in prep.).

2.1.3. X-ray Observations

X-ray observations were obtained using *Chandra* and *XMM-Newton* for 14 of the clusters in Table 1. The majority of the X-ray observations, data reduction, and analyses are the same as described by A11, to which we refer the reader for a more detailed description. Relative to A11, we include new *Chandra* observations for five clusters, and re-run the X-ray analysis for the five clusters with new optical spectroscopic redshifts, one of which also had new *Chandra* observations. In this section, we summarize the X-ray observations and results, and describe additional analyses required to incorporate the X-ray measurements in the cosmological analysis.

Summarizing A11, 15 of the 16 highest ξ clusters from V10 were targeted for X-ray observations, however, in this work, we use only the 14 clusters at $z > 0.3$. Of these, twelve were observed with *Chandra* and four clusters were observed with *XMM-Newton*. Two clusters were observed by both *Chandra* and *XMM-Newton*, and for these clusters only the *Chandra* data was included in the analysis. From the data, the X-ray observables, M_g , T_X , and Y_X , were measured in a manner identical

to Vikhlinin et al. (2009a), where M_g is the gas mass within r_{500} , T_X is the core-excised X-ray temperature in an annulus between $0.15 - 1.0 \times r_{500}$, and $Y_X \equiv M_g T_X$. We solved for each observable and r_{500} iteratively, to maintain consistency with their respective observable-mass relations. Since A11, five of the clusters have new spectroscopic redshift measurements. For these clusters we repeat the A11 reduction and analysis using the new redshifts, and give the updated results in Table 7.

Five of the clusters from A11 have had additional *Chandra* observations, which we include in this work. In Table 8, we list these clusters, the *Chandra* observation IDs, and the improvement in exposure time and cluster source counts adding the new observations. We repeat the A11 reduction and analysis to derive new constraints on the X-ray observables, which are given in Table 7. For these results, relative to A11, we use more recent *Chandra* analysis software (CIAO 4.3) and calibration files (CALDB 4.3.3). We find that the new *Chandra* calibration files typically change Y_X by $< 5\%$. This is at a level below the assumed mass-normalization uncertainty that we assign in Section 3.1.1.

For the cosmological analysis in this work, described in Section 3, we need to calculate the X-ray observables as a function of cosmology and scaling relation parameters. To do this, we derive density and temperature profiles for all 14 clusters with X-ray data. We calculate $T_X(r)$ and $M_g(r)$ (for the calculation of $Y_X(r)$) from the X-ray observations of each cluster assuming a reference cosmology, where r corresponds to a physical radius in

the cluster and the profiles are defined to return the cluster observable within r . The reference cosmology is chosen to match A11; a preferred Λ CDM cosmology using WMAP7+BAO+ H_0 data with $\Omega_M = 0.272$, $\Omega_\Lambda = 0.728$ and $H_0 = 70.2 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (Komatsu et al. 2011). For three of the clusters with the lowest X-ray photon counts, the $T_X(r)$ profiles have jumps which appear unphysical. For this reason, we have assumed a functional form of $T_X(r)$ that, in combination with the measured $M_g(r)$, matches the pressure profile from Arnaud et al. (2010) and is normalized to give the measured Y_X assuming the reference cosmology. When considering the eleven clusters with well-behaved temperature profiles, we find that our cosmological results in Section 4 negligibly change when assuming either the functional form of $T_X(r)$, or the profile derived from the data. Therefore, we consider this approximation valid for this work.

2.2. External Cosmological Data Sets

In addition to the SPT cluster data set, we incorporate several external cosmological data sets, including measurements of the CMB power spectrum (CMB), the Hubble constant (H_0), baryon acoustic oscillations (BAO), type Ia supernova (SNe), and big bang nucleosynthesis (BBN). We will use these abbreviations when referring to these data sets, and will use several different combinations of them in our analysis and results. Below we give references and a brief description of each external data set. Also, when discussing our results in Section 4 and onward, we will define the SPT_{CL} data set as the combination of the SPT-SZ data, optical redshift, and X-ray measurements described in Section 2.1.

We use measurements of the CMB power spectrum from the seven-year WMAP data release (WMAP7, Larson et al. 2011) and 790 deg² of sky observed with the SPT (Keisler et al. 2011). Following Keisler et al. (2011),¹ we fit the CMB data to a model including primary CMB anisotropy plus three nuisance parameters that model “foreground” signals detectable in the SPT data. We use low-redshift measurements of H_0 from the Hubble Space Telescope (Riess et al. 2011), which we include as a Gaussian prior of $H_0 = 73.8 \pm 2.4 \text{ km s}^{-1} \text{ Mpc}^{-1}$. We use measurements of the BAO feature using SDSS and 2dFGRS data (Percival et al. 2010). The BAO constraints have been applied as a measurement of $r_s/D_V(z = 0.2) = 0.1905 \pm 0.0061$ and $r_s/D_V(z = 0.35) = 0.1097 \pm 0.0036$; where r_s is the comoving sound horizon size at the baryon drag epoch, $D_V(z) \equiv [(1+z)^2 D_A^2(z) cz/H(z)]^{1/3}$, $D_A(z)$ is the angular diameter distance, and $H(z)$ is the Hubble parameter. The inverse covariance matrix given in Eq. 5 of Percival et al. (2010) is used for the BAO measurements. We use measurements of the luminosity distances of Type Ia supernovae (SNe) from the Union2 compilation of 557 SNe (Amanullah et al. 2010), and include their treatment of systematic uncertainties. Finally, we use a BBN prior from measurements of the abundances of ⁴He and Deuterium (Kirkman et al. 2003), which we include as a Gaussian prior of $\Omega_b h^2 = 0.022 \pm 0.002$.

3. COSMOLOGICAL ANALYSIS

In this section, we outline the cosmological analysis method for the SPT cluster data set, including the calculation of the cosmological likelihood and the assumed parameterization for the cluster mass-observable relations. This implementation allows for self-consistent constraints on cosmology and the cluster scaling relations, i.e., the cluster mass calibration, by simultaneously varying the cluster-mass observable relations and cosmological parameters using a Markov Chain Monte Carlo (MCMC) technique. The method is generalizable in way that can include additional cluster observables from other data sets (e.g., weak lensing shear, optical velocity dispersions). We have incorporated our calculation of the SPT cluster likelihood into the CosmoMC code² of Lewis & Bridle (2002) to compute its joint likelihood with the external cosmological data sets.

3.1. Scaling Relation Parameterization

3.1.1. X-ray: $Y_X - M_{500}$

Following Vikhlinin et al. (2009b), we use Y_X as an X-ray proxy for cluster mass, M_{500} . We assume a $Y_X - M_{500}$ relation of the form

$$\frac{M_{500}}{10^{14} M_\odot/h} = \left(A_X h^{3/2} \right) \left(\frac{Y_X}{3 \times 10^{14} M_\odot \text{ keV}} \right)^{B_X} E(z)^{C_X}, \quad (1)$$

parameterized by the normalization A_X , the slope B_X , the redshift evolution C_X , where $E(z) \equiv H(z)/H_0$, and a log-normal scatter D_X on Y_X . Relative to the form of this equation in Vikhlinin et al. (2009b), we have multiplied the right-hand side by an extra factor of h , so that the cluster mass M_{500} , is in units of M_\odot/h to match the $\zeta - M_{500}$ relation in Section 3.1.2. For our cosmological analysis, we assume Gaussian priors on the scaling relation parameters, which we list in Table 2. The priors are motivated by constraints from X-ray measurements by Vikhlinin et al. (2009a) and simulations, which we describe below.

Vikhlinin et al. (2009a) constrained the $Y_X - M_{500}$ relation using X-ray observations of a low-redshift ($z < 0.3$) sample of 17 relaxed clusters to estimate the hydrostatic total mass and Y_X . From simulations, Kravtsov et al. (2006) put an upper limit on the systematic offset in the $Y_X - M_{500}$ relation between relaxed and unrelaxed clusters of 4%. Simulations also expect that biases in hydrostatic mass estimates are less for relaxed clusters and typically $\lesssim 15\%$ (Nagai et al. 2007). Therefore, a $Y_X - M_{500}$ relation calibrated from hydrostatic mass estimates of a relaxed cluster sample should have minimal biases and be applicable to a larger cluster sample of both relaxed and unrelaxed clusters.

From the above measurements, Vikhlinin et al. (2009a) obtained best-fit values of $A_X = 5.77 \pm 0.20$ and $B_X = 0.57 \pm 0.03$, where the uncertainties are statistical only. They estimated the systematic uncertainty in the A_X calibration by comparing to weak lensing mass estimates from Hoekstra (2007) for a sample of 10 low-redshift clusters. From this analysis, they estimated a 1σ uncertainty of 9% on the *Chandra* mass scale calibration, which we add in quadrature with their quoted statistical uncertainty on A_X . We have therefore assumed Gaussian pri-

¹ <http://pole.uchicago.edu/public/data/keisler11>

² <http://cosmologist.info/cosmomc/>

ors of $A_X = 5.77 \pm 0.56$ and $B_X = 0.57 \pm 0.03$.

We assume a Gaussian prior of $C_X = -0.4 \pm 0.2$, consistent with self-similar evolution and a 50% uncertainty. There are relatively few observational constraints on the normalization of the $Y_X - M_{500}$ relation at $z > 0.3$. However, this level of C_X uncertainty was chosen to match the prior assumed in Vikhlinin et al. (2009b), which was in turn motivated by constraints from the simulations of Kravtsov et al. (2006). The 1σ prior on C_X would correspond to a 6% difference in the mass calibration between $z = 0.0$ and 0.6 . For the highest redshift cluster in this work, $z = 1.074$, this would correspond to an additional 11% uncertainty in the mass calibration, and a 15% total uncertainty when also considering the prior on A_X .

We assume a Gaussian prior of $D_X = 0.12 \pm 0.08$, which we truncate below 0.02 and where $D_X = 0.12$ corresponds to a 12% log-normal scatter in Y_X for a given mass. This scatter has been measured to have values ranging from 0-12% (Vikhlinin et al. 2009a; Mantz et al. 2010a). Analogous to Vikhlinin et al. (2009b), for our cosmological analysis we have chosen a prior centered on a value which is consistent with simulations (e.g., Kravtsov et al. 2006). Although we have assumed a larger uncertainty on the scatter than the range typically found in simulations, this uncertainty is negligible for our cosmological constraints, see Section 6.

3.1.2. SZ: $\zeta - M_{500}$

As in V10, we use the detection significance ξ as an SZ mass proxy. However, since the relation between ξ and halo mass is complicated by the comparable effects of intrinsic scatter and instrumental noise, we introduce the unbiased significance, ζ : the average detection signal-to-noise of a simulated cluster, measured across many noise realizations, evaluated at the preferred position and filter scale of that cluster as determined by fitting the cluster in the absence of noise.

We relate unbiased significance ζ to the detection significance ξ as follows. Firstly, ζ is related to $\langle \xi \rangle$ through the relation

$$\zeta = \sqrt{\langle \xi \rangle^2 - 3} \quad (2)$$

at $\xi > 2$. This maximization bias comes from having maximized ξ across possible cluster positions and filters scales, effectively adding three degrees of freedom to the fit with ξ analogous to a χ^2 . Additionally, $\langle \xi \rangle$ relates to ξ by a Gaussian scatter of unit width. Simulations have been used to verify that these approximations introduce negligible bias or scatter compared to the Poisson noise of the sample. For further details we refer the reader to V10.

We assume a $\zeta - M_{500}$ relation of the form

$$\zeta = A_{SZ} \left(\frac{M_{500}}{3 \times 10^{14} M_\odot h^{-1}} \right)^{B_{SZ}} \left(\frac{E(z)}{E(0.6)} \right)^{C_{SZ}}, \quad (3)$$

parameterized by the normalization A_{SZ} , the slope B_{SZ} , the redshift evolution C_{SZ} , and a log-normal scatter, D_{SZ} , on ζ . V10 motivated the form of this relation based on physical arguments, and the expected range of these parameters based on self-similar arguments. In V10, the cluster mass was defined within a spherical region in which the density is equal to 200 times the mean

matter density at the cluster redshift. In this work, to be consistent with the $Y_X - M_{500}$ relation, we are defining the cluster mass as M_{500} , the mass in a spherical radius, r_{500} , within which the density is equal to 500 times the critical density of the universe at the cluster redshift. This change has motivated a change in the redshift evolution term from $(1+z)$ to $E(z)$, because of the expected self-similar scaling between Y_{SZ} and M_{500} (e.g., Kravtsov et al. 2006). In addition, we allow for a correlated scatter between ζ and Y_X with a correlation coefficient ρ , which we allow to uniformly vary between 0.02 and 0.98, but away from 0 and 1 for numerical reasons.

Analogous to V10 and summarized in Section 2.1.1, we have used simulated SZ maps to characterize the scaling between ζ and cluster mass. We have repeated this exercise to match the form of the scaling given in equation 3, and we give the Gaussian priors in Table 2. The fractional uncertainty on each parameter matches V10, except for the log-normal scatter, for which we allow a larger uncertainty in this work. However, this uncertainty remains negligible for these cosmological constraints, see Section 6.

3.2. Likelihood Model

The analysis method employed in this work closely mirrors the one presented by V10 with extensions to incorporate the X-ray data. In V10, the parameter space was explored through importance sampling of pre-existing WMAP MCMC chains. In this work, we have elected to utilize a full MCMC algorithm. This is accomplished through the use of the CosmoMC analysis package, where we have included the cluster abundance likelihood as an additional module in the CosmoMC likelihood calculation. Among the numerous advantages to this approach is the ability to enforce quantitative convergence criteria as well as the optional inclusion of supplemental data sets.

Each step in the Markov chain selects a new point in the joint cosmological and scaling relation parameter space. Prior to passing these variables to the cluster likelihood evaluation, we use the Code for Anisotropies in the Microwave Background (CAMB) (Lewis et al. 2000) to compute the matter power spectrum at 20 logarithmically spaced redshifts between $0 < z < 2.5$. The matter power spectra, as well as the proposed scaling relation and relevant cosmological parameters, are the inputs to the cluster likelihood function.

At this point, the analysis follows a similar path to that laid out by V10. First, the matter power spectra and cosmology are used to calculate a mass function based upon the Tinker et al. (2008) prescription, which we calculate for an over-density of $\Delta = 500 \Omega_m(z)$, to match our cluster mass definition in Section 3.1. As noted in Tinker et al. (2008), this function predicts the halo abundance as a function of input cosmology across a mass range of $10^{11} h^{-1} M_\odot \leq M \leq 10^{15} h^{-1} M_\odot$ and a redshift range of $0 \leq z \leq 2.5$. Tinker et al. (2008) claim an overall calibration of their mass function to simulations of $\lesssim 5\%$. Stanek et al. (2010) found that the inclusion of non-gravitational physics can shift the normalization of the mass function by $\sim 10\%$ along the mass direction. However, this effect is approximately degenerate with an uncertainty between intra-cluster gas observables and mass, which we account for explicitly in our scaling

relation uncertainty through equations 1 and 3.

As in V10, the next step in the analysis is to move the theoretically predicted cluster abundances from their native M_{500} mass space into the observable space for this analysis. V10 define this space by the SZ detection significance, ξ , and the optically derived redshift, z . This resulted in a two-dimensional surface of predicted cluster abundances in the observable space. In this analysis, we perform a similar transformation, this time including a third dimension, the X-ray parameter Y_X . This results in a three-dimensional volume of predicted cluster abundances, now as a function of ξ , Y_X , and z .

Using the scaling relations discussed in §3.1, the halo mass function is recast as a predicted number density in terms of ξ , Y_X and z , which we write as

$$\frac{dN(\xi, Y_X, z|\vec{p})}{d\xi dY_X dz} = \int dM P(\xi, Y_X|M, z, \vec{p}) P(M, z|\vec{p}) \Theta(\xi-5) \quad (4)$$

where \vec{p} is the set of cosmological and scaling relation parameters, and Θ is the Heaviside step function. The likelihood function is then given by the Poisson probability:

$$\ln \mathcal{L}(\vec{p}) = \sum_i \ln \frac{dN(\xi_i, Y_{X_i}, z_i, |\vec{p})}{d\xi dY_X dz} - \int \frac{dN(\xi, Y_X, z, |\vec{p})}{d\xi dY_X dz} d\xi dY_X dz, \quad (5)$$

where the sum over the i index runs over the SPT cluster catalog. Note that we have neglected a global offset to the log-likelihood.

We compute Equation 4 on a three-dimensional grid that is 200 by 200 by 30 in the ζ , Y_X , and z dimensions, respectively. For each value of Y_X and z we then convert to the ξ basis by using the ζ - ξ relation defined in Equation 2, where we also convolve with a unit-width Gaussian in ξ to account for the noise in the SPT measurement.

For each step in the MCMC, we recalculate Y_X for each cluster given its $T_X(r)$ and $M_g(r)$ profiles from Section 2.1.3, so that its calculated Y_X is consistent with the $Y_X - M_{500}$ relation and r_{500} at that step. To account for this in the cosmological likelihood, we modify the likelihood by adding $\sum_i \ln Y_{X_i}$ to the right hand side of Equation 5. For a detailed explanation, see Appendix B. For each cluster, we account for finite measurement errors or missing data in z and Y_X by modifying the first term in Equation 5 by marginalizing over the relevant parameter, weighted by either a Gaussian likelihood determined from its uncertainty or a uniform distribution over the allowed range.

From this calculation we obtain a value for the cluster likelihood corresponding to this particular set of cosmological and scaling relation parameters. This value is then returned to CosmoMC where it may be combined with other likelihood calculations from supplemental data sets and is used in the MCMC step acceptance/rejection computation.

4. Λ CDM RESULTS

We first consider the SPT_{CL} data constraints for a spatially flat Λ CDM cosmological model. For this model, we fit 15 parameters: the nine scaling relation parameters

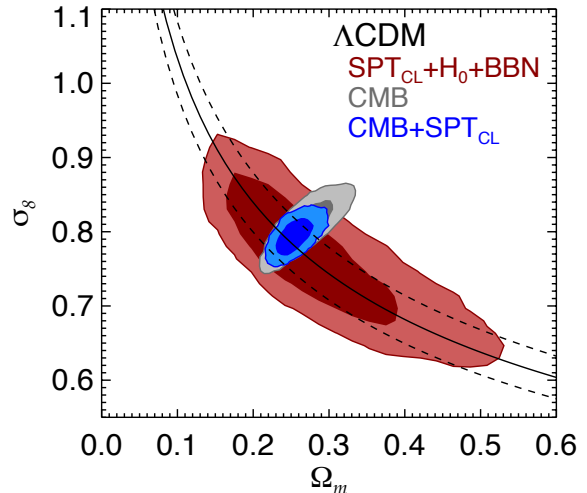


FIG. 1.— Assuming a Λ CDM cosmology, the two-dimensional marginalized constraints on σ_8 and Ω_m . Contours show the 68% and 95% confidence regions for the SPT_{CL}+ H_0 +BBN (red), CMB (gray), and CMB+SPT_{CL} (blue) data sets. The black lines are the best-fit constraint (solid) and 68% confidence region (dashed) for the combination of parameters that the SPT_{CL}+ H_0 +BBN data set best constrains: $\sigma_8(\Omega_m/0.25)^{0.30} = 0.785 \pm 0.037$.

and six primary cosmology parameters listed in Table 3. For constraints on any individual parameter, we always quote the mean of the likelihood distribution and the 68% confidence interval about the mean. The confidence interval reflects uncertainties after marginalizing over all other parameters, and includes systematic uncertainties in the cluster scaling relations and mass calibration, as described in Section 3.1. In this analysis, we use the SPT_{CL} and external cosmological data sets as described in Section 2.

4.1. Cosmological Constraints

The SPT_{CL} data are not sensitive to all six Λ CDM cosmology parameters. Here and in Section 5.1, when considering the SPT_{CL} cosmological constraints without CMB data, we always include BBN and H_0 priors, as indicated. For the SPT_{CL}+ H_0 +BBN data set, we also fix the optical depth of reionization, τ , and allow the scalar tilt, n_s , to vary uniformly between 0.944 and 0.989, the 95% confidence range from Keisler et al. (2011) assuming a Λ CDM model. However, we note that the SPT_{CL} cosmological constraints vary negligibly over this range of n_s . As noted in Section 2.2, whenever we refer to the SPT_{CL} data, we are implicitly referring to the combined SPT-SZ data, optical redshift, and X-ray measurements described in Section 2.1.

In Figure 1, we show the constraints on the σ_8 and Ω_m parameters for the individual and combined SPT_{CL} and CMB data sets. In Table 3, we give the marginalized constraints for the cosmological and scaling relation parameters. The latter will be discussed further in Sections 4.2 and 6. In a Λ CDM cosmology, the SPT_{CL} data is most sensitive to σ_8 and Ω_m . The number of clusters increases with either parameter, so the cluster abundance data effectively constrain a product of the two. We find that the SPT_{CL}+ H_0 +BBN constraints are well approximated as $\sigma_8(\Omega_m/0.25)^{0.30} = 0.785 \pm 0.037$, which we show in

TABLE 2
PARAMETER TABLE

Type	Symbol	Meaning	Gaussian Prior
Scaling	A_{SZ}	ζ -mass normalization	5.58 ± 1.67
Relation	B_{SZ}	ζ -mass slope	1.32 ± 0.26
Parameters	C_{SZ}	ζ -mass redshift evolution	0.87 ± 0.44
	D_{SZ}	Log-normal scatter in ζ	0.24 ± 0.16
	A_X	Y_X -mass normalization	5.77 ± 0.56
	B_X	Y_X -mass slope	0.57 ± 0.03
	C_X	Y_X -mass redshift evolution	-0.40 ± 0.20
	D_X	Log-normal scatter in Y_X	0.12 ± 0.08
	ρ	Correlated scatter between ζ and Y_X	Uniform:(0.02, 0.98)
Primary	$\Omega_c h^2$	Dark matter density	
Cosmology	$\Omega_b h^2$	Baryon density	
Parameters	$100\Theta_s$	Angular scale of the sound horizon at last scattering	
	n_s	Scalar tilt of power spectrum	
	$10^9 \Delta_R^2$	Scalar amplitude of power spectrum	
	τ	Optical depth to reionization	
Extension	w	Dark energy equation of state	
Cosmology	f_ν	Fraction of dark matter in the form of neutrinos, $\Sigma m_\nu = 94\text{eV}(f_\nu \Omega_c h^2)$	
Parameters	N_{eff}	The effective number of relativistic species	
	f_{NL}	Primordial non-Gaussianity parameter	
Derived	σ_8	Matter fluctuations on 8 Mpc scales at $z = 0$	
Cosmology	Ω_m	Total matter density	
Parameters	h	$h \equiv H_0/100 \text{ km s}^{-1} \text{ Mpc}^{-1}$, where H_0 is the Hubble constant at $z = 0$	

TABLE 3
 Λ CDM CONSTRAINTS

Parameter	Prior	SPT _{CL} + H_0 +BBN	CMB	CMB+SPT _{CL}
A_{SZ}	5.58 ± 1.67	5.31 ± 0.98	-	4.91 ± 0.71
B_{SZ}	1.32 ± 0.26	1.39 ± 0.15	-	1.40 ± 0.15
C_{SZ}	0.87 ± 0.44	0.90 ± 0.34	-	0.83 ± 0.30
D_{SZ}	0.24 ± 0.16	0.21 ± 0.10	-	0.21 ± 0.09
A_X	5.77 ± 0.56	5.69 ± 0.51	-	5.82 ± 0.48
B_X	0.57 ± 0.03	0.564 ± 0.029	-	0.563 ± 0.029
C_X	-0.40 ± 0.20	-0.37 ± 0.16	-	-0.35 ± 0.16
D_X	0.12 ± 0.08	0.14 ± 0.07	-	0.14 ± 0.07
ρ	(0.02, 0.98)	0.52 ± 0.27	-	0.52 ± 0.27
$\Omega_c h^2$	-	0.133 ± 0.045	0.111 ± 0.0048	0.109 ± 0.0032
$\Omega_b h^2$	-	0.0221 ± 0.0020	0.0222 ± 0.0004	0.0223 ± 0.0004
$100\Theta_s$	-	1.065 ± 0.041	1.041 ± 0.0016	1.041 ± 0.0016
n_s	(0.944, 0.989)	0.966 ± 0.013	0.965 ± 0.011	0.967 ± 0.010
$10^9 \Delta_R^2$	-	2.16 ± 1.30	2.44 ± 0.10	2.40 ± 0.08
τ	-	(0.090)	0.086 ± 0.014	0.087 ± 0.014
σ_8	-	0.766 ± 0.062	0.808 ± 0.024	0.795 ± 0.016
Ω_m	-	0.285 ± 0.083	0.268 ± 0.025	0.255 ± 0.016
h	-	0.739 ± 0.024	0.707 ± 0.022	0.717 ± 0.016

NOTE. — The marginalized constraints on the scaling relation and primary cosmology parameters from Table 2, where we report the mean of the likelihood distribution and the 68% confidence interval about the mean. The priors are Gaussian, except for ρ and n_s , which are uniform over the range given. The n_s prior is only used for the SPT_{CL}+ H_0 +BBN data set.

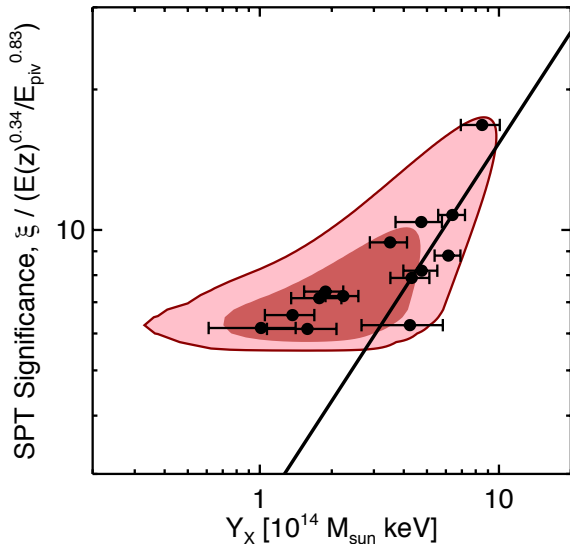


FIG. 2.— A plot of the SZ-significance, ξ , versus the X-ray observable Y_X for the 14 SPT clusters with X-ray measurements. From the CMB+SPT_{CL} data set fit to a Λ CDM cosmology, we use the best-fit $\zeta - M_{500}$ and $Y_X - M_{500}$ scaling relations to calculate the expected form and redshift evolution of the ξ - Y_X relation (solid-line), where $E_{\text{piv}} \equiv E(z = 0.6)$. With the best-fit cosmology parameters, we also predict the effective 68 and 95% confidence intervals for the expected distribution of clusters in the ξ - Y_X plane (red contours). The measured and predicted cluster distribution show qualitatively good agreement.

Figure 1 by the solid and dashed lines. Combining the SPT_{CL} and CMB data, we constrain $\sigma_8 = 0.795 \pm 0.016$ and $\Omega_m = 0.255 \pm 0.016$, a factor of 1.5 improvement on each over the constraints from the CMB alone.

The SPT_{CL} constraints are consistent with results using optical and X-ray selected cluster samples. Recently, Rozo et al. (2010) compared the cluster constraints from several different methods, and found generally good agreement and comparable constraints. It is typical for cluster based constraints to be quoted in terms of the product of σ_8 and Ω_m to an exponent which varies depending on the mass scale of the cluster sample. One example for comparison is Vikhlinin et al. (2009b), who constrained $\sigma_8(\Omega_m/0.25)^{0.47} = 0.813 \pm 0.027$. For typical Λ CDM model constraints of $\Omega_m \sim 0.25 - 0.30$, this agrees well with our result.

4.2. Scaling Relation Constraints

In Figure 2, we show the relationship between ξ and Y_X for the 14 clusters with X-ray observations, over-plotted with the expected distribution of clusters and the best-fit relation determined using the CMB+SPT_{CL} data from Section 4.1. The combination of the steep mass function and SZ selection yields a distribution of clusters visibly offset from the best-fit scaling relation, an effect often referred to as Eddington bias. We note that our cosmological analysis method described in Section 3.2, explicitly accounts for the SZ selection and therefore Eddington bias. We also predict the expected distribution of clusters in the ξ - Y_X plane assuming the best-fit cosmology and scaling relation parameters, and applying a comparable selection as was used for the SPT X-ray follow-up ($z > 0.3$ and $\xi > 5.45$). The predicted 14.2 clusters is

consistent with the 14 detected. In Figure 2, we over-plot the effective 68 and 95% confidence region in the ξ - Y_X plane where we would expect to find these clusters. Qualitatively we find good agreement between the observed and predicted cluster distribution.

In Table 3, we give the constraints on the $Y_X - M_{500}$ and $\zeta - M_{500}$ scaling relations using the SPT_{CL} and CMB+SPT_{CL} data sets. Because the $Y_X - M_{500}$ relation has significantly tighter priors than the $\zeta - M_{500}$ relation, we will not give the $Y_X - M_{500}$ constraints for the modified cosmologies presented in Section 5. Similarly, for the parameter ρ , the correlated scatter between ζ and Y_X , we have virtually no constraining power. In all cases, ρ moves nearly uniformly across the entire allowed range, and has a negligible effect on the cosmological constraints. In Appendix C, we give posterior mass estimates for each cluster using a similar method as described in V10 and briefly reviewed in the appendix.

If there were a significant discrepancy between the simulation-based prior on the $\zeta - M_{500}$ relation and the observational prior on the $Y_X - M_{500}$ relation, we would observe it as an offset between the central value of the $\zeta - M_{500}$ prior and its best-fit value. From the CMB+SPT_{CL} constraints, the largest offset is for A_{SZ} , with a best-fit value of 4.91 ± 0.71 compared to the simulation prior of 5.58 ± 1.67 . An offset in this direction would be consistent with the SZ simulation prior under-estimating the mass of a cluster by a factor of $\sim ((4.91 \pm 0.71)/5.58)^{1/1.4} = 0.91 \pm 0.09$. This result is consistent with preliminary estimates from A11, who estimated this factor to be 0.78 ± 0.06 . We note that A11 did not marginalize over uncertainties in either the X-ray scaling relation or cosmological parameters, both of which affect this result. The derived offset is also a function of the assumed cosmology. For example, if we assume a Λ CDM cosmology with a non-zero neutrino mass, as in Section 5.2, we find a value closer to the simulation prior, $A_{SZ} = 5.39 \pm 0.79$, using the CMB+H₀+SPT_{CL} data set. Therefore, we find no significant inconsistency between the simulation-based prior on the $\zeta - M_{500}$ relation and the observational prior on the $Y_X - M_{500}$ relation.

5. EXTENSIONS TO Λ CDM

In this section, we consider extensions to a spatially flat Λ CDM cosmology. For each extension, we also fit the nine scaling relation parameters and six primary cosmology parameters listed in Table 2. We consider four extension cosmologies where we include the following as free parameters: the dark energy equation of state (w), the sum of the neutrino masses (Σm_ν), the sum of neutrino masses and the effective number of relativistic species (N_{eff}), and a primordial non-Gaussianity (f_{NL}). For constraints on any individual parameter, we always quote the mean of the likelihood distribution and the 68% confidence interval about the mean. The confidence interval will include uncertainties after marginalizing over all other parameters, which includes systematic uncertainties in the cluster scaling relations and mass calibration, as described in Section 3.1. In this analysis, we use the SPT_{CL} and external cosmological data sets as described in Section 2.

5.1. w CDM

TABLE 4
 w CDM CONSTRAINTS

		CMB	H_0 +BBN +SPT _{CL}	CMB+BAO+SNe	CMB+BAO+SNe +SPT _{CL}
Scaling	A_{SZ}	-	5.12 ± 1.36	-	4.75 ± 0.79
Parameters	B_{SZ}	-	1.40 ± 0.15	-	1.41 ± 0.15
	C_{SZ}	-	0.92 ± 0.36	-	0.85 ± 0.29
	D_{SZ}	-	0.22 ± 0.10	-	0.21 ± 0.10
Cosmology	σ_8	0.864 ± 0.120	0.773 ± 0.088	0.823 ± 0.040	0.793 ± 0.028
Parameters	Ω_m	0.244 ± 0.089	0.293 ± 0.113	0.279 ± 0.016	0.273 ± 0.015
	h	0.775 ± 0.128	0.740 ± 0.025	0.698 ± 0.018	0.697 ± 0.018
	w	-1.19 ± 0.37	-1.09 ± 0.36	-1.014 ± 0.078	-0.973 ± 0.063

NOTE. — The marginalized constraints on a subset of the scaling relation and cosmology parameters from Table 2. Scaling relation and primary cosmology parameters not given are still varied in the MCMC and marginalized over for these constraints. We report the mean of the likelihood distribution and the 68% confidence interval about the mean.

TABLE 5
 Λ CDM + Σm_ν + N_{eff} CONSTRAINTS

		CMB+H ₀ +BAO	CMB+H ₀ +BAO +SPT _{CL}	CMB+H ₀ +SPT _{CL}	CMB+H ₀ +BAO +SPT _{CL}
Scaling	A_{SZ}	-	5.26 ± 0.79	5.39 ± 0.79	5.01 ± 0.85
Parameters	B_{SZ}	-	1.39 ± 0.14	1.39 ± 0.14	1.41 ± 0.15
	C_{SZ}	-	0.89 ± 0.30	0.89 ± 0.30	0.89 ± 0.30
	D_{SZ}	-	0.20 ± 0.09	0.20 ± 0.10	0.21 ± 0.10
	Cosmology	σ_8	0.761 ± 0.043	0.770 ± 0.026	0.771 ± 0.023
Parameters	Ω_m	0.275 ± 0.016	0.272 ± 0.015	0.260 ± 0.018	0.284 ± 0.018
	h	0.698 ± 0.014	0.701 ± 0.013	0.712 ± 0.017	0.727 ± 0.020
	Σm_ν (eV)	0.19 ± 0.14	0.15 ± 0.10	0.12 ± 0.09	0.34 ± 0.17
	Σm_ν (eV), 95% CL	< 0.45	< 0.33	< 0.28	< 0.63
	N_{eff}	(3.046)	(3.046)	(3.046)	3.91 ± 0.42

NOTE. — The marginalized constraints on a subset of the scaling relation and cosmology parameters from Table 2. Scaling relation and primary cosmology parameters not given are still varied in the MCMC and marginalized over for these constraints. We report the mean of the likelihood distribution and the 68% confidence interval about the mean.

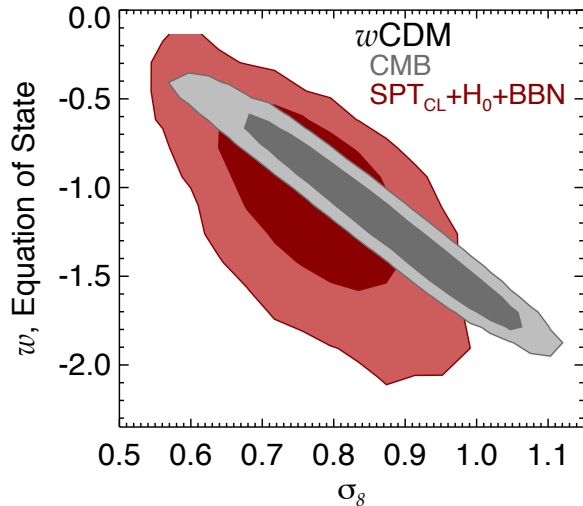


FIG. 3.— Assuming a w CDM cosmology, the two-dimensional marginalized constraints on w and σ_8 . Contours show the 68% and 95% confidence regions for the $\text{SPT}_{\text{CL}}+H_0+\text{BBN}$ (red) and CMB (gray) data sets.

The first extension we consider is a w CDM cosmology, a model in which the equation of state of dark energy is a constant w . The cluster abundance and the shape of the mass function depend on w through its effect on the growth of structure, or equivalently the redshift evolution of σ_8 . The CMB data measure structure at $z \sim 1100$, and therefore require significant extrapolation to predict the cluster abundance in the redshift range of the SPT sample ($0.3 < z < 1.1$). Therefore, consistency between the implied w from both data sets is already an important systematic test of dark energy.

In Figure 3, we show the constraints on w and σ_8 using the CMB and $\text{SPT}_{\text{CL}}+H_0+\text{BBN}$ data sets. The likelihood contours have significant overlap, implying the data are in good agreement. Relative to the CMB, the SPT_{CL} data tend to disfavor cosmologies with large σ_8 and more negative w . In Table 4, we give marginalized constraints for several cosmological and scaling relation parameters. The SPT_{CL} data constrain $w = -1.09 \pm 0.36$, and have similar constraining power to the CMB data, for which the constraints have a significant degeneracy between w and σ_8 . The SPT_{CL} data simultaneously constrain $\sigma_8 = 0.773 \pm 0.088$. This constraint has a factor of ~ 1.4 lower uncertainty than that from the CMB data.

5.1.1. w CDM with BAO and SNe data sets

In this section, we consider the improvement in w CDM cosmological constraints when adding the SPT_{CL} data to the CMB, BAO, and SNe data sets. In Figure 4, we show the constraints of the combined CMB+BAO+SNe data set, before and after including the SPT_{CL} data. The SPT_{CL} data most significantly improve the constraints on σ_8 and w ; reducing the allowed two-dimensional likelihood area by a factor of ~ 1.8 . In Table 4, we give the marginalized constraints for several parameters before and after the inclusion of the SPT_{CL} data. The combined constraints are $w = -0.973 \pm 0.063$ and $\sigma_8 = 0.793 \pm 0.028$, a factor of 1.25 and 1.4 improvement, respectively, over the constraints without clusters. The

combined data set also constrains $\Omega_m = 0.273 \pm 0.015$ and $h = 0.697 \pm 0.018$. These constraints are consistent with previous cluster-based results (Vikhlinin et al. 2009b; Mantz et al. 2010c; Rozo et al. 2010), which used X-ray and optically selected samples of typically lower redshift clusters. The sensitivity of the SPT_{CL} cluster data to the amplitude of structure, σ_8 , is primarily what gives it the ability to break degeneracies with the distance-relation based constraints from the BAO and SNe data sets. We note the slight tension with the H_0 constraints from Riess et al. (2011) of $h = 0.738 \pm 0.024$. While this tension is not significant, it helps to intuitively explain some constraints on neutrino mass in Section 5.2.

5.1.2. $\zeta - M_{500}$ Constraints

Given the work of V10 and other cluster results (e.g., Vikhlinin et al. 2009b; Mantz et al. 2010c; Rozo et al. 2010), we expect the cluster mass-calibration to be the dominant systematic uncertainty limiting our results. In Figure 5, we show the constraints on A_{SZ} and σ_8 . The $\text{SPT}_{\text{CL}}+H_0+\text{BBN}$ data set has a significant degeneracy between its constraints on A_{SZ} and σ_8 . From this data, we constrain the fractional uncertainty, $\delta A_{\text{SZ}}/A_{\text{SZ}}$, to be 27%, which is effectively constrained only by the 14 clusters that have both X-ray and SZ measurements. This constraint is not significantly better than the uncertainty in the simulation based prior of 30%. With enough X-ray observations, we expect the $\zeta - M_{500}$ calibration to be limited by the uncertainty of the $Y_X - M_{500}$ relation, because the latter is currently better observationally constrained. In this limit, we would expect a fractional uncertainty on A_{SZ} of $B_{\text{SZ}} \times (\delta A_X/A_X) \sim 14\%$. The above would suggest that for a w CDM cosmology we would need X-ray observations of ~ 50 clusters, i.e., ~ 14 clusters $\times (27\%/14\%)^2$, to calibrate A_{SZ} in terms of mass so that its not the dominant source of uncertainty. When adding the BAO and SNe data sets, we improve the constraints on A_{SZ} to an accuracy of $\sim 16\%$. However, these data sets are not sensitive to either A_{SZ} or σ_8 , and cannot completely break their degeneracy. In Section 6, we will discuss the systematic uncertainties from this degeneracy on our cosmological constraints in more detail.

5.2. Λ CDM with Massive Neutrinos

We next consider a Λ CDM cosmology with non-zero neutrino masses. Cosmological measurements are primarily sensitive to the neutrino masses through their effect on structure formation. A massive neutrino additionally affects the CMB power spectrum if it was non-relativistic at the redshift of recombination. For example, if the heaviest neutrino had a mass $\lesssim 0.6$ eV, it would be relativistic at recombination and therefore would not significantly affect the structure in the CMB (Komatsu et al. 2009). However, as the universe expanded and cooled neutrinos would transition to non-relativistic, and would contribute to Ω_m but not to structure formation below their free streaming scale, implying a lower σ_8 at $z = 0$ and fewer clusters. This implies that measurements of the CMB power spectrum alone cannot constrain the neutrino mass to significantly less than 0.6 eV per species (i.e., $\Sigma m_\nu \lesssim 1.8$ eV), and the constraints will be significantly degenerate with σ_8 . Local measurements of structure break this degeneracy, and significantly improve the neutrino mass constraints.

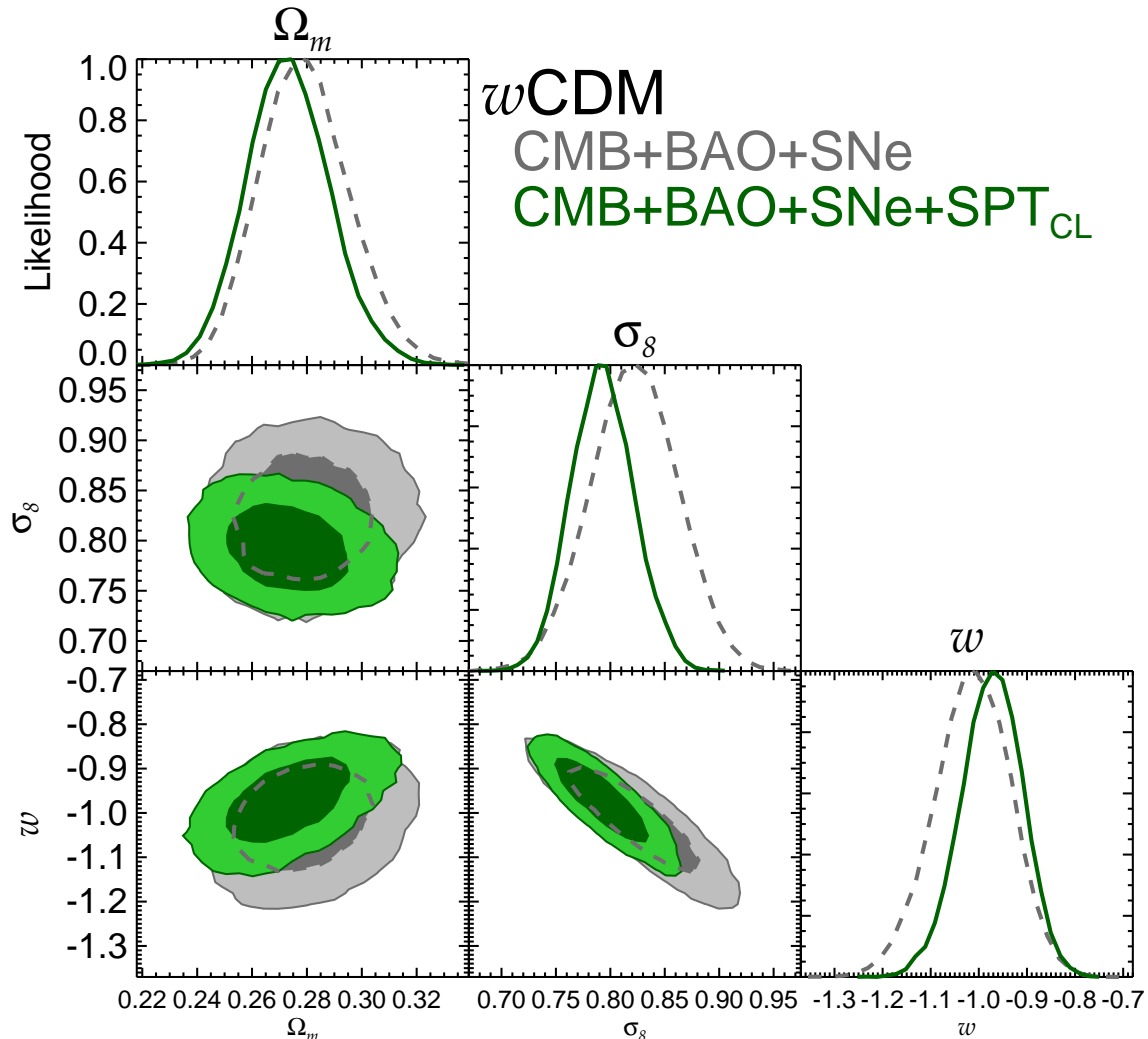


FIG. 4.— Assuming a w CDM cosmology, the constraints on Ω_m , σ_8 , and w . The plots along the diagonal are the one-dimensional marginalized likelihood. The off-diagonal plots are the two-dimensional marginalized constraints showing the 68% and 95% confidence regions. We show the constraints for the CMB+BAO+SNe (gray, dashed), and CMB+BAO+SNe+SPT_{CL} (green, solid) data sets. The SPT_{CL} data improves the constraints on σ_8 and w , by factors of 1.4 and 1.25, respectively.

We also note the significant degeneracy between the CMB power spectrum constraints on H_0 and Σm_ν . Massive neutrinos affect the amplitude of the early integrated Sachs-Wolfe effect causing a shift of the first peak of the CMB power spectrum towards larger angular scales that can be absorbed by a lower value of H_0 (Ichikawa et al. 2005). Komatsu et al. (2011) used a combination of WMAP7+ H_0 +BAO data to set a limit of $\Sigma m_\nu < 0.58$ at a 95% confidence limit (CL). Following Komatsu et al. (2011), we consider the same combination of data sets to add to the CMB power spectrum measurements, which were chosen because of their insensitivity to systematic errors and their ability to maximally constrain Σm_ν by breaking the degeneracy with H_0 . We define $\Sigma m_\nu = 94 \text{ eV} (f_\nu \Omega_c h^2)$, where f_ν is the fraction of dark matter in the form of massive neutrinos.

In Figure 6, we show the constraints on σ_8 and Σm_ν , using the CMB+ H_0 +BAO data set, before and after including the SPT_{CL} data. In Table 5, we give the marginalized constraints on each parameter. Using the

CMB+ H_0 +BAO+SPT_{CL} data set, we constrain $\Sigma m_\nu < 0.33 \text{ eV}$ at a 95% CL, a factor of 1.4 improvement over the constraints without the SPT_{CL} data. This improvement is primarily due to the tighter constraints on σ_8 for which the uncertainty decreased by a factor 1.8. The constraint is lower by excluding the BAO data; using only the CMB+ H_0 +SPT_{CL} data set we constrain $\Sigma m_\nu < 0.28 \text{ eV}$ at a 95% CL. These improved constraints can be understood from the H_0 measurements, as also noted in Section 5.1. The results of Riess et al. (2011) favor a marginally higher H_0 value than the CMB+BAO data. Because of the degeneracy between Σm_ν and H_0 in the CMB constraints, a higher value of H_0 tends to favor lower values of Σm_ν . The constraints presented here are comparable to other recent results using optically and X-ray selected cluster samples with similar cosmological data sets (Reid et al. 2010; Mantz et al. 2010b).

5.2.1. Number of Relativistic Species

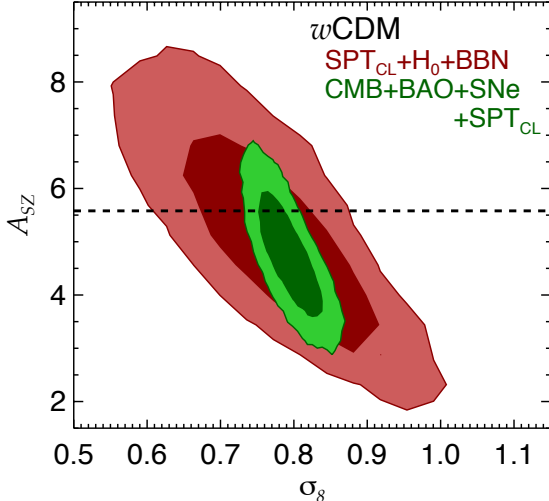


FIG. 5.— Assuming a w CDM cosmology, the two-dimensional marginalized constraints on A_{SZ} and σ_8 . Contours show the 68% and 95% confidence regions for the $SPT_{CL}+H_0+BBN$ (red) and $CMB+BAO+SNe+SPT_{CL}$ (green) data sets. The horizontal black dashed line is the center of the theory prior on A_{SZ} .

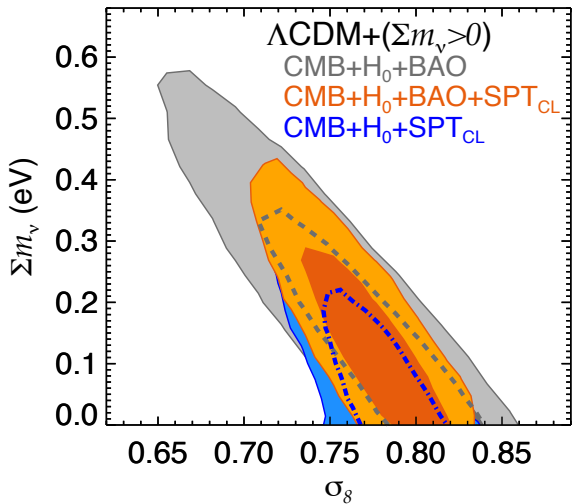


FIG. 6.— Assuming a Λ CDM cosmology ($\Sigma m_\nu > 0$), the two-dimensional marginalized constraints on Σm_ν and σ_8 . Contours show the 68% and 95% confidence regions for the $CMB+H_0+BAO$ (gray, dashed), $CMB+H_0+BAO+SPT_{CL}$ (orange, solid), and $CMB+H_0+SPT_{CL}$ (blue, dot-dashed) data sets. The SPT_{CL} data improves the constraints on σ_8 and Σm_ν , by factors of 1.8 and 1.4, respectively.

Recent measurements have shown a $\sim 2\sigma$ preference for increased damping in the tail of the CMB power spectrum (Dunkley et al. 2011; Keisler et al. 2011). This damping could be caused by several different physical mechanisms, such as a high primordial helium abundance, a running of the scalar spectral index, or additional relativistic species. This last explanation is particularly timely because of recent measurements from atmospheric (Aguilar-Arevalo et al. 2010) and nuclear reactor (Mention et al. 2011) neutrino oscillation experiments

that find some evidence for a sterile neutrino species. It has been pointed out that these measurements are most consistent with two sterile neutrinos and $\Sigma m_\nu \gtrsim 1.7$ eV (Kopp et al. 2011). Therefore, we consider the joint cosmological constraints on N_{eff} and Σm_ν to compare with these terrestrial results.

With only three neutrino species, we would expect $N_{\text{eff}} = 3.046$, a value slightly larger than three because of energy injection from electron-positron annihilation at the end of neutrino freeze-out (Dicus et al. 1982; Lopez et al. 1999; Mangano et al. 2005). As N_{eff} increases, the contribution to the gravitational potential of the additional neutrino perturbations boosts the early growth of dark matter perturbations (Bashinsky & Seljak 2004), which also increases σ_8 (Hou et al. 2011). As explained in Section 5.2, adding neutrino mass at the levels considered here only affects the low-redshift universe, suppressing structure formation, and lowering σ_8 at $z = 0$. Therefore, increasing N_{eff} will also allow an increasing Σm_ν . Keisler et al. (2011) used a combination of $CMB+H_0+BAO$ data to constrain $\Sigma m_\nu < 0.69$ eV at a 95% CL, $\sigma_8 = 0.803 \pm 0.056$, and $N_{\text{eff}} = 3.98 \pm 0.43$.

In Figure 7, we show the constraints on N_{eff} , Σm_ν , and σ_8 , using the $CMB+H_0+BAO$ data set, before and after including the SPT_{CL} data. In Table, 5 we give the marginalized constraints. When varying N_{eff} we assume consistency with BBN for our constraints. Using the $CMB+H_0+BAO+SPT_{CL}$ data set, we constrain $\Sigma m_\nu < 0.63$ eV at a 95% CL, $\sigma_8 = 0.777 \pm 0.031$, and $N_{\text{eff}} = 3.91 \pm 0.42$. Relative to Keisler et al. (2011), the addition of the SPT_{CL} data improves the constraints on σ_8 by a factor of 1.8, and reduces the upper limit on Σm_ν by a factor of 1.1. However, the addition of the SPT_{CL} data does noticeably sharpen the peak in the marginalized one-dimensional likelihood for Σm_ν , such that the maximum likelihood constraint peaks away from zero, $\Sigma m_\nu = 0.34 \pm 0.17$ eV.

As noted in Keisler et al. (2011), models of the CMB power spectrum that include increased damping are favored at the 1.6-1.9 σ level. However, even if one accepts the need for an extra parameter to explain the damping, its physical origin is unclear. Regardless, considering the N_{eff} model extension is instructive to help understand the model dependency of the neutrino mass constraints. Keisler et al. (2011) considered three models to explain the excess damping and found that the N_{eff} model had the most significant effect on σ_8 . The inclusion of N_{eff} also weakens the constraints on Σm_ν , because of the degeneracies between N_{eff} , σ_8 , and Σm_ν . In the combined cosmological data set, the SPT_{CL} data mainly constrains σ_8 , which helps to break this degeneracy and indirectly improve the neutrino mass constraints. Therefore, the Σm_ν constraint from the N_{eff} model can be considered a conservative upper limit on Σm_ν regardless of the physical mechanism for the increased damping.

5.3. Λ CDM with f_{NL}

Finally, we consider a Λ CDM cosmology with primordial non-Gaussianity. Standard inflationary cosmology predicts that density fluctuations in the universe were seeded by random Gaussian fluctuations. However, inflationary models can be constructed that predict significant levels of non-Gaussianity (e.g., Bartolo et al. 2004). The leading order non-Gaussian term is typi-

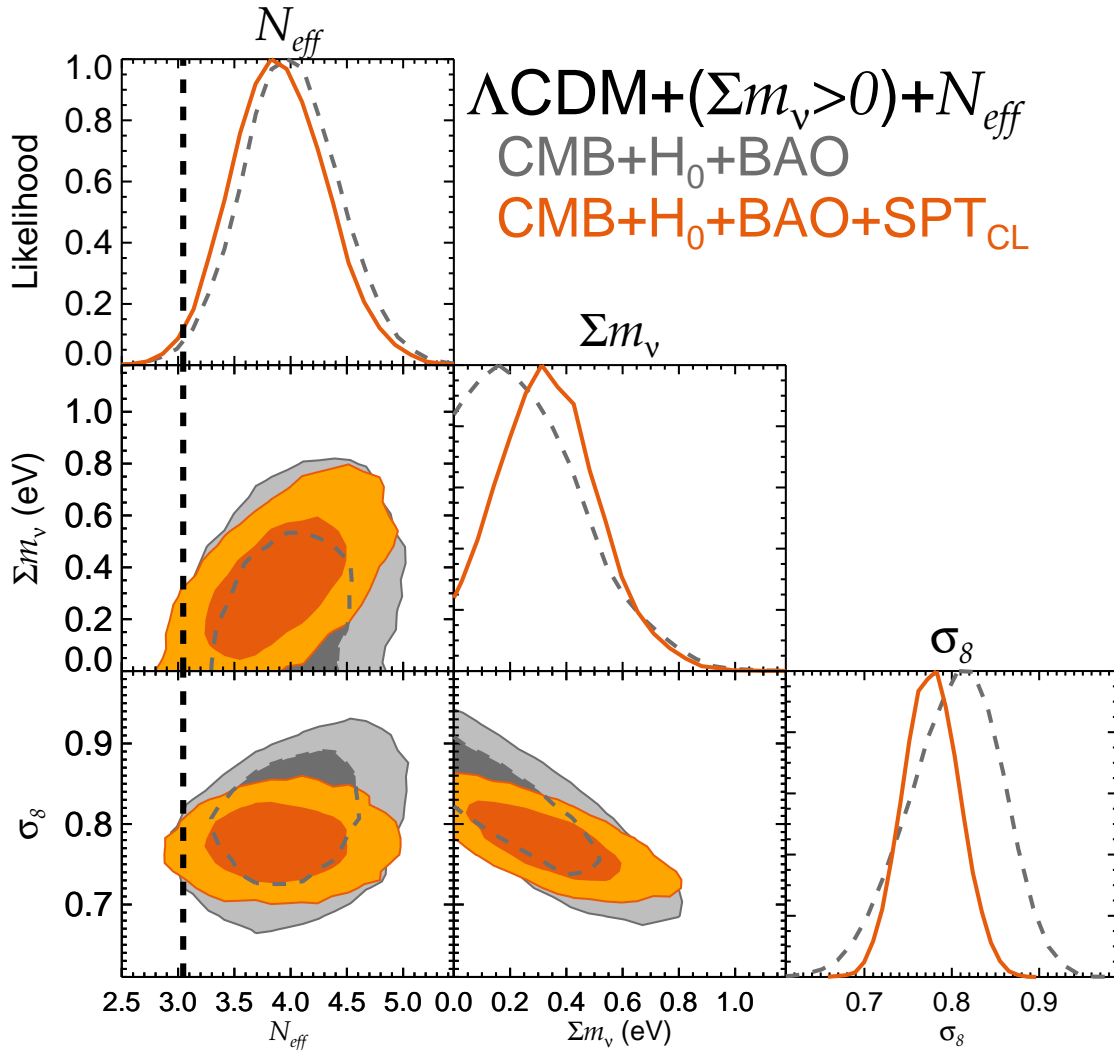


FIG. 7.— Assuming a Λ CDM cosmology with massive neutrinos and a free number of relativistic species, N_{eff} , the constraints on σ_8 , Σm_ν , and N_{eff} . The plots along the diagonal are the one-dimensional marginalized likelihood. The off-diagonal plots are the two-dimensional marginalized constraints showing the 68% and 95% confidence regions. We show the constraints for the CMB+ H_0 +BAO (gray, dashed), and CMB+ H_0 +BAO+SPT_{CL} (orange, solid) data sets. The black vertical line shows $N_{\text{eff}} = 3.046$, the expected value for three neutrino species. Using the CMB+ H_0 +BAO+SPT_{CL} data set, the 68% confidence marginalized constraints are $\sigma_8 = 0.777 \pm 0.031$, $N_{\text{eff}} = 3.91 \pm 0.42$, and $\Sigma m_\nu = 0.34 \pm 0.17$ eV, with a 95% CL of $\Sigma m_\nu < 0.63$ eV.

cally described by the parameter f_{NL} . Using a measurement of the CMB power spectrum from WMAP7 data, Komatsu et al. (2011) measured a 95% CL of $-10 < f_{NL} < 74$. Primordial non-Gaussianity can also manifest itself through the abundance of massive galaxy clusters. While the constraints from an SPT-like SZ survey are not expected to be competitive with current CMB constraints (Dalal et al. 2008), they constitute a constraint independent from the CMB results which is sensitive to very different physical scales. In principle, even a single massive high-redshift cluster can falsify standard Λ CDM cosmology (Mortonson et al. 2011), though currently the most massive cluster known at $z > 1$ is not in significant tension (Foley et al. 2011). In our analysis, we incorporate f_{NL} as a modification of the cluster mass function following the prescription of Dalal et al. (2008), and only consider the effect of f_{NL} on the SPT_{CL} data set.

In Figure 8, we show the constraints on σ_8 and f_{NL}

using the CMB+SPT_{CL} data set. The marginalized constraints are $f_{NL} = -192 \pm 310$, $\sigma_8 = 0.803 \pm 0.022$, and $A_{SZ} = 5.27 \pm 0.89$. The best-fit value of f_{NL} is slightly negative, generally implying fewer massive high-redshift clusters. In Section 4, we found that for a Λ CDM cosmology using the CMB+SPT_{CL} data, our constraint for A_{SZ} was 0.9σ lower than the simulation prior. This corresponds to a higher mass for a cluster of a given ξ , which can also be thought of as the SPT survey having fewer clusters than expected given the simulation prior. When f_{NL} is added as a parameter, A_{SZ} moves back towards its simulation prior, and the deficit of clusters can be maintained by a more negative f_{NL} . For the model to match the number of clusters measured by SPT, a more negative f_{NL} can be balanced with either a larger σ_8 or A_{SZ} , creating a degeneracy in this direction. Regardless, any deficit of clusters is not significant relative to the uncertainty on either σ_8 or A_{SZ} , even in a Λ CDM cosmology for which they are best constrained.

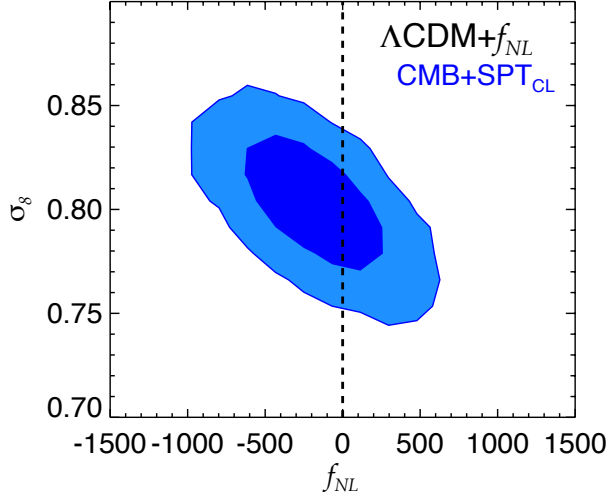


FIG. 8.— Assuming a Λ CDM cosmology with a primordial non-Gaussianity characterized by the parameter f_{NL} , the two-dimensional marginalized constraints on σ_8 and f_{NL} using the CMB+SPT_{CL} data set (blue). Contours show the 68% and 95% confidence regions. We only consider the affect of f_{NL} on the SPT_{CL} data set. We measure $f_{NL} = -192 \pm 310$, consistent with zero non-Gaussianity.

Our results are consistent with Williamson et al. (2011) which used the 26 most massive clusters in the full 2500 deg² SPT-SZ survey to constrain $f_{NL} = 20 \pm 450$. Our work differs from Williamson et al. (2011) in that we use a much smaller area of the SPT-SZ survey, we select clusters down to a lower ξ threshold, and we use an improved mass calibration.

6. SOURCES OF UNCERTAINTY

Previous SPT cluster survey results, namely V10 and Williamson et al. (2011), found cosmological constraints that were limited most significantly by the cluster mass calibration, or equivalently the fractional uncertainty in A_{SZ} . In this work, we have reduced this uncertainty by incorporating the external mass calibration from the $Y_X - M_{500}$ relation using X-ray observations of the SPT clusters. We can directly estimate the impact of the uncertainties in the X-ray and SZ scaling relations by importance sampling the MCMC chains, where we post-process the chains by imposing a narrow prior on each scaling relation parameter centered around the best-fit value. The resulting increase in precision on the cosmological parameters allows a measure of the impact from the uncertainty in the scaling relations. In this way, we effectively “fix” the X-ray and scaling relation parameters, a process which we will implicitly be referring to throughout this section. For a w CDM cosmology, we also consider the impact of the SNe systematic uncertainty on the cosmological results presented here.

With enough SZ and X-ray observations, we expect the $\zeta - M_{500}$ calibration to be limited by the calibration of the $Y_X - M_{500}$ relation because the latter has tighter external priors. In practice, there will be an additional uncertainty in the $\zeta - M_{500}$ calibration from the limited number of SZ and X-ray observations for cross-calibration, and this uncertainty will also degrade the cosmological constraints. We wish to separate this ef-

fect, which we will refer to as the SZ- Y_X scaling uncertainty, from the additional systematic uncertainty from the $Y_X - M_{500}$ calibration, which we will refer to as the X-ray scaling uncertainty, and the statistical uncertainty from the cluster sample size. By fixing the X-ray and SZ scaling relation parameters, as described above, we can measure the impact of the SZ- Y_X scaling uncertainty, X-ray scaling uncertainty, and statistical uncertainty on our cosmological constraints.

6.1. Λ CDM Cosmology: Scaling Relation Uncertainty

We first consider the Λ CDM constraints using the SPT_{CL}+ H_0 +BBN data set, the results of which were described in Section 4. This data best constrained the combination of $\sigma_8(\Omega_m/0.25)^{0.30} = 0.785 \pm 0.037$. The sources of uncertainty for this result are summarized in Table 6, and are discussed below.

For the X-ray scaling relation parameters, only the uncertainty in A_X and C_X , the normalization and redshift evolution parameters, contribute significantly to the uncertainty on $\sigma_8(\Omega_m/0.25)^{0.30}$. Fixing each parameter separately implies that they contribute an uncertainty on $\sigma_8(\Omega_m/0.25)^{0.30}$ of ± 0.022 and ± 0.015 , respectively. It is not surprising that the normalization of the mass calibration significantly affects the constraints, and the redshift evolution can be understood for similar reasons. For a cluster at the median redshift of the SPT sample, $z = 0.74$, the prior on the C_X value effectively contributes an additional 7% to the cluster mass calibration. This can be compared to the 10% mass calibration uncertainty from the prior on A_X . Fixing all X-ray parameters simultaneously, implies that they contribute an uncertainty on $\sigma_8(\Omega_m/0.25)^{0.30}$ of ± 0.028 .

For the SZ scaling relation parameters, only the uncertainty in A_{SZ} contributes significantly to the uncertainty on $\sigma_8(\Omega_m/0.25)^{0.30}$. Fixing all the SZ scaling parameters, we measure an uncertainty on $\sigma_8(\Omega_m/0.25)^{0.30}$ of ± 0.023 from statistical uncertainty, ± 0.010 from the SZ- Y_X scaling uncertainty, and ± 0.028 due to X-ray scaling uncertainty (as discussed above). The relatively low contribution from the SZ- Y_X scaling uncertainty is not surprising considering the constraints on the fractional uncertainty of A_{SZ} , which was near the systematic limit of 14% imposed by the $Y_X - M_{500}$ calibration.

Therefore, the Λ CDM constraints are nearly at the systematic limit from the calibration of the $Y_X - M_{500}$ relation. For our constraint of $\sigma_8(\Omega_m/0.25)^{0.30} = 0.785 \pm 0.037$, the X-ray scaling and statistical uncertainty contribute almost equal amounts of ± 0.028 and ± 0.023 , respectively. By only increasing the cluster sample size we could reduce the uncertainty by up to ~ 1.3 ($\sim 0.037/0.028$). Further improvements would require a more accurate cluster mass calibration.

6.2. w CDM Cosmology: Scaling Relation and SNe Uncertainty

We next consider the sources of uncertainty for the w CDM cosmology discussed in Section 5.1. We will concentrate on using the SPT_{CL}+ H_0 +BBN and CMB+BAO+SNe+SPT_{CL} data sets, which produce constraints of $w = -1.09 \pm 0.36$ and $w = -0.973 \pm 0.063$, respectively. The sources of uncertainty for this result are summarized in Table 6, and are discussed below.

TABLE 6
ERROR BUDGET

	Λ CDM	w CDM		w CDM	
	SPT _{CL} + H_0 +BBN $\sigma_8(\Omega_m/0.25)^{0.30}$	SPT _{CL} + H_0 +BBN		CMB+BAO+SNe+SPT _{CL}	
		w	σ_8	w	σ_8
Baseline, Section 4.5.1	0.785 ± 0.037	-1.09 ± 0.36	0.773 ± 0.088	-0.973 ± 0.063	0.793 ± 0.028
SNe Systematic	—	—	—	± 0.026	± 0.005
SZ- Y_X Scaling	± 0.010	± 0.19	± 0.066	± 0.013	± 0.013
X-ray Scaling Systematic	± 0.028	± 0.15	± 0.036	± 0.019	± 0.014
Statistical	± 0.023	± 0.27	± 0.046	± 0.033	± 0.013

NOTE. — We give the mean and 68% confidence intervals for a subset of the cosmological parameters reported in Sections 4 and 5.1. The last four rows give the 1σ error in each cosmological parameter due to the stated uncertainty.

We first consider the SPT_{CL}+ H_0 +BBN data set. For the X-ray scaling relation parameters, we again find that the uncertainty in A_X and C_X contribute the largest uncertainty on w . Fixing each parameter independently implies they contribute an uncertainty of $\delta w = \pm 0.10$ and ± 0.11 , respectively, and a total uncertainty of ± 0.15 . For the SZ scaling relation parameters, only the uncertainty in A_{SZ} contributes significantly to the uncertainty on w . Fixing all the SZ scaling parameters, we measure an uncertainty on w of: ± 0.27 from statistical uncertainty, ± 0.19 from the SZ- Y_X scaling uncertainty, and ± 0.15 due to X-ray scaling uncertainty. Therefore, unlike the Λ CDM case, we find that our constraints on w would be significantly improved by adding more clusters and additional Y_X measurements. A similar conclusion is reached repeating the above analysis for σ_8 . In principle, adding more clusters and Y_X measurements would reduce the uncertainty on w and σ_8 to values limited by the X-ray scaling uncertainty. In this limit we should measure w and σ_8 with an uncertainty of $\delta w = \pm 0.15$ and $\delta\sigma_8 = \pm 0.036$, or ~ 2.5 times better than our current constraints.

When considering the CMB+BAO+SNe+SPT_{CL} data set, we reach qualitatively similar conclusions, however the total uncertainty is significantly lower because of the parameter degeneracies that are broken from the additional data sets. We first re-run the MCMC chains without SNe systematic uncertainty. Fixing all the SZ scaling parameters, we measure a statistical uncertainty of $\delta w = \pm 0.033$, a factor of two improvement relative to the constraints including all systematic uncertainties. Comparing this uncertainty to that with no fixed parameters, we estimate an uncertainty on w of ± 0.019 from X-ray scaling uncertainty and ± 0.013 from SZ- Y_X scaling uncertainty. The addition of the SPT_{CL} data also significantly reduces the systematic uncertainty from SNe. Running a CMB+BAO+SNe MCMC chain with and without SNe systematics, we measure $w = -1.014 \pm 0.078$ and $w = -1.017 \pm 0.050$, respectively. This implies that SNe systematics are contributing an uncertainty of $\delta w = \pm 0.060$. After adding the SPT_{CL} data the uncertainty from SNe systematics is reduced to $\delta w = \pm 0.026$, a factor of ~ 2.3 improvement.

6.3. Point Source Contamination

In V10, it was argued that point source contamination contributed a negligible level of uncertainty relative to the statistical precision of the cluster sample. Since we are using the same cluster sample, we expect the same conclusion to hold, though we briefly summarize their

arguments here. From Poisson distributed sources, the probability of a chance superposition of a bright point source ($\gtrsim 6$ mJy) with a cluster is negligible, given the sky density of sources at 150 GHz ($\sim 1 \text{ deg}^{-2}$, Vieira et al. (2010)). Furthermore, our cosmological analysis in Section 3 explicitly accounts for a Poisson distributed background of sources, and the X-ray measurements gives an additional systematic check on an offset in the SZ measurements. Correlated emission from cluster members could potentially fill in cluster decrements. However, correlated radio emission has previously been calculated to be negligible at 150 GHz for clusters of the typical SPT mass scale and redshift range (Lin et al. 2009; Sehgal et al. 2010). In V10, it was also argued that the level of correlated dusty emission is negligible, from the known quenching of star formation in massive clusters (Hashimoto et al. 1998), and the sub-millimeter luminosity function (Pascale et al. 2009). In addition, more recent *Spitzer* infrared observations of a sample of X-ray selected groups and low-mass clusters, found that correlated dusty emission is insignificant compared to the SZ signal (George et al. 2011). These arguments apply to the clusters in the SPT sample, which span a similar redshift range and are of higher mass than the objects considered in that work.

7. DISCUSSION

7.1. Improvement Relative to V10

In this work, the cluster sample is the same as used in V10 for their cosmological analysis. However, we have improved the cosmological constraints relative to V10 by including X-ray measurements in order to reduce the cluster mass calibration uncertainty. It is not straightforward to quantify the improvement for two main reasons. First, V10 used pre-existing WMAP7 MCMC chains from Komatsu et al. (2011) that they importance sampled by re-weighting the chains by the likelihood of the SPT cluster catalog given each set of parameters. In this work, we generated new MCMC chains while simultaneously fitting both data sets. Second, each result uses somewhat different external data sets, in particular for the CMB power spectrum measurements, where the results in this work also include CMB measurements from Keisler et al. (2011).

Without explicitly correcting for these differences, we can approximate the improvement from including the X-ray measurements by considering the relative improvements of adding the SPT_{CL} data to the CMB data used in either result. For a Λ CDM cosmology, the WMAP7 data constrains $\sigma_8 = 0.801 \pm 0.030$. In V10, the ad-

dition of the SPT data, this constraint improved to $\sigma_8 = 0.791 \pm 0.027$, a factor of 1.1 improvement. In this work, for a Λ CDM cosmology using the SPT_{CL} and CMB data, we constrained $\sigma_8 = 0.795 \pm 0.016$, a factor of 1.5 improvement over the constraints from the CMB alone. Therefore, the addition of the X-ray measurements improved the Λ CDM constraints on σ_8 by a factor of ~ 1.4 . A comparison of the w CDM cosmological constraints is more complicated because of the somewhat different handling of the external data sets. Regardless, the significant improvement in the constraints from the X-ray measurements is clear.

7.2. Prospects for Further Improvement

The results in this paper were derived using 18 clusters from 178 deg² of the 2500 deg² SPT-SZ survey. The full survey will significantly increase both the area and overall depth of the SZ maps. Reichardt et al. (2012, in prep.) will present a catalog of ~ 200 clusters from the first 800 deg² of the SPT-SZ survey, with a median redshift of ~ 0.5 and a median mass of $M_{500} \sim 2.3 \times 10^{14} M_\odot/h$. This sample is representative of the cluster yield and properties for the full survey, which was completed in November 2011 and has detected ~ 500 clusters. Therefore it is useful to consider how the method used in this work will be applied to the full survey, and what level of improvement we can expect on the cosmological constraints.

Using the SPT_{CL}+ H_0 +BBN data set, we found that our constraints are currently limited by both statistical uncertainty and the SZ- Y_X scaling uncertainty. Both uncertainties would be improved by adding more SPT clusters with additional Y_X measurements. Recently the SPT collaboration was awarded a *Chandra* X-ray Visionary Project (XVP) to complete X-ray observations of the 80 most significant clusters at $z > 0.4$ detected in the first 2000 deg² of the SPT-SZ survey. As argued in Section 5.1.2, we would need $\gtrsim 50$ clusters with Y_X measurements for the $\zeta - M_{500}$ calibration to be limited by the $Y_X - M_{500}$ uncertainty. With this many clusters, the statistical uncertainty on w should decrease to a level below the systematic uncertainty from the X-ray scaling relation, $\delta w = \pm 0.15$. Combining the full 2500 deg² SPT-SZ survey with the *Chandra* XVP observations, we would be limited to this constraint from the current calibration of the $Y_X - M_{500}$ relation.

To reduce the systematic uncertainty further, we would need more accurate cluster mass estimates than currently exist from X-ray measurements alone. In Section 6.2, we found that the X-ray scaling systematics were currently limited by the uncertainty in A_X and C_X , whose fractional uncertainty was 10% and 50%, respectively. Reducing their uncertainty by a factor of two, would reduce their contribution to the systematic uncertainty to $\delta w = \pm 0.037$ and ± 0.074 , respectively. This would effectively correspond to an overall mass calibration uncertainty of 5% with an additional 6% uncertainty in the evolution of the mass calibration between $z = 0.0 - 1.1$.

This level of mass calibration should be achievable by incorporating additional data sets, such as optical velocity dispersion (White et al. 2010) or weak lensing measurements (Hoekstra 2007; Becker & Kravtsov 2011). For example, in massive clusters, the scatter in weak lensing mass estimates is expected to be $\sim 20\%$ (Becker &

Kravtsov 2011). Therefore, with weak lensing observations of two sets of ~ 15 -20 clusters at low and high redshift, this level of accuracy should be achievable. Towards this goal, the SPT collaboration has been approved for weak lensing observations of ~ 35 SPT-detected clusters spanning $0.30 < z < 1.3$ using the Magellan and *Hubble* telescopes. Additionally, the SPT collaboration has been approved for optical velocity dispersion observations of ~ 100 SPT-detected clusters using the Very Large Telescope (VLT) and a large NOAO program on Gemini South. With these data sets, we expect to achieve the factor of two improvement in mass calibration, as discussed above.

Applying this calibration to the full 2500 deg² SPT_{CL}+ H_0 +BBN data set, we should constrain w with an accuracy of $\sim 8\%$, or a factor of ~ 4.5 tighter than the current SPT_{CL}+ H_0 +BBN constraints. This improved constraint would be comparable to the current constraints from the CMB+BAO+SNe data, and would be an independent systematic test of the standard dark energy paradigm by measuring the effect of dark energy on the growth of structure. Combining the existing CMB+BAO+SNe data with the 2500 deg² SPT cluster sample, the uncertainty from the SZ- Y_X scaling and the cluster sample size is expected to be negligible compared to the uncertainty contributed by the improved cluster mass calibration. In this case, the SPT cluster data would contribute an uncertainty of only $\sim 1\%$ to the significantly improved constraint on w from the combined data set.

8. CONCLUSIONS

We use measurements from the SPT-SZ cluster survey in combination with X-ray measurements to constrain cosmological parameters. We have described and implemented a method that simultaneously fits for cosmological parameters and the scaling of the SZ and X-ray observables with cluster mass. The method is generalizable to multiple cluster observables, and self-consistently accounts for the effects of cluster selection and uncertainties in cluster mass calibration on the derived cosmological constraints. We apply this method to a SZ-selected catalog of 18 galaxy clusters identified in 178 deg² of the 2500 deg² SPT-SZ survey. This is the first analysis of an SZ survey to directly incorporate X-ray observations, which has reduced the uncertainty on both the cluster mass calibration and the cosmological constraints.

For a Λ CDM cosmology, we find that the SPT_{CL}+ H_0 +BBN data best constrain $\sigma_8(\Omega_m/0.25)^{0.30} = 0.785 \pm 0.037$, where the total uncertainty consists of an approximately equal amount of statistical and systematic uncertainty. These constraints are consistent, and comparable, with other constraints using X-ray-selected (Vikhlinin et al. 2009b; Mantz et al. 2010c) and optically-selected (Roza et al. 2010) cluster samples. In combination with measurements of the CMB power spectrum from the SPT data and the seven-year WMAP data, the SPT cluster data constrain $\sigma_8 = 0.795 \pm 0.016$ and $\Omega_m = 0.255 \pm 0.016$, a factor of 1.5 improvement on each parameter over the constraints from the CMB data alone.

We consider several extensions beyond a Λ CDM cosmological model by including the following as free parameters: the dark energy equation of state (w), the sum of

the neutrino masses (Σm_ν), the effective number of relativistic species (N_{eff}), and a primordial non-Gaussianity (f_{NL}).

For a w CDM cosmology, the SPT_{CL}+ H_0 +BBN data constrain $w = -1.09 \pm 0.36$ and $\sigma_8 = 0.773 \pm 0.088$, consistent with dark energy being due to a cosmological constant, and with comparable uncertainties to constraints from the CMB data alone. Using the CMB+BAO+SNe+SPT_{CL} data set, we constrain $w = -0.973 \pm 0.063$ and $\sigma_8 = 0.793 \pm 0.028$, a factor of 1.25 and 1.4 improvement, respectively, over the constraints without SPT cluster data. The uncertainty on w consists of approximately equal contributions from statistical uncertainty, systematic uncertainty from SNe, and systematic uncertainty from cluster scaling relations, with the latter contributing an uncertainty of $\delta w = \pm 0.023$.

We next consider a Λ CDM cosmology with a non-zero neutrino mass. Using a CMB+ H_0 +BAO+SPT_{CL} data set, we constrain the sum of the neutrino masses Σm_ν to be < 0.33 eV at 95% confidence, a factor of 1.4 improvement over the constraints without SPT cluster data. We find even tighter constraints when we exclude the BAO data set, which tend to favor a lower value of H_0 and therefore a higher neutrino mass. Using a CMB+ H_0 +SPT_{CL} data set, we constrain $\Sigma m_\nu < 0.28$ eV at 95% confidence. We also consider a model with a free effective number of relativistic species, N_{eff} , to explain the increased damping that is observed in the the CMB power spectrum. Using a CMB+ H_0 +BAO+SPT_{CL} data set, we jointly measure $N_{\text{eff}} = 3.91 \pm 0.42$ and $\Sigma m_\nu = 0.34 \pm 0.17$ eV, while constraining $\Sigma m_\nu < 0.63$ eV at 95% confidence.

Finally, we consider a Λ CDM cosmology where we allow the number of observed clusters to be affected by non-Gaussian density fluctuations characterized by the parameter f_{NL} . Using a CMB+SPT_{CL} data set, we measure $f_{NL} = -192 \pm 310$, consistent with zero non-Gaussianity.

The results presented in this paper use 18 clusters from 178 deg² of the 2500 deg² SPT-SZ survey, and are limited by the combination of the cluster sample size and mass calibration. The SPT-SZ survey was completed in November 2011, and has detected ~ 500 clusters with a median redshift of ~ 0.5 and a median mass of $M_{500} \sim 2.3 \times 10^{14} M_\odot/h$. Ongoing X-ray, weak lensing, and optical velocity dispersion observations of SPT-SZ-selected clusters will be used to produce an improved cluster mass calibration of the sample. The full SPT-SZ survey and improved mass calibration will produce constraints on w comparable to current constraints from the combination of CMB+BAO+SNe data, and would represent an independent systematic test of the standard dark energy paradigm by measuring the effect of dark energy on the growth of structure. The combination of CMB+BAO+SNe data with the SPT cluster sample will

break degeneracies between the data sets resulting in significantly tighter constraints on dark energy.

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REFERENCES

- Aguilar-Arevalo, A. A., et al. 2010, Phys. Rev. Lett., 105, 181801
Amanullah, R., et al. 2010, ApJ, 716, 712
Andersson, K., et al. 2011, ApJ, 738, 48
Arnaud, M., Pratt, G. W., Piffaretti, R., Böhringer, H., Croston, J. H., & Pointecouteau, E. 2010, A&A, 517, A92+
Bartolo, N., Komatsu, E., Matarrese, S., & Riotto, A. 2004, Phys. Rep., 402, 103
Bashinsky, S., & Seljak, U. 2004, Phys. Rev. D, 69, 083002
Bazin et al. 2012, In Preparation
Becker, M. R., & Kravtsov, A. V. 2011, ApJ, 740, 25
Bhattacharya, S., Heitmann, K., White, M., Lukić, Z., Wagner, C., & Habib, S. 2011, ApJ, 732, 122
Bode, P., Ostriker, J. P., Weller, J., & Shaw, L. 2007, ApJ, 663, 139
Brodwin, M., et al. 2010, ApJ, 721, 90
Carlstrom, J. E., et al. 2011, PASP, 123, 568

- Carlstrom, J. E., Holder, G. P., & Reese, E. D. 2002, *ARA&A*, 40, 643
- Dalal, N., Doré, O., Huterer, D., & Shirokov, A. 2008, *Phys. Rev. D*, 77, 123514
- Dicus, D. A., Kolb, E. W., Gleeson, A. M., Sudarshan, E. C. G., Teplitz, V. L., & Turner, M. S. 1982, *Phys. Rev. D*, 26, 2694
- Dobbs, M. A., et al. 2011, *ArXiv e-prints*, 1112.4215
- Dunkley, J., et al. 2011, *ApJ*, 739, 52
- Foley, R. J., et al. 2011, *ApJ*, 731, 86
- George, M. R., et al. 2011, *ArXiv e-prints*, 1109.6040
- Haehnelt, M. G., & Tegmark, M. 1996, *MNRAS*, 279, 545+
- Haiman, Z., Mohr, J. J., & Holder, G. P. 2001, *ApJ*, 553, 545
- Hashimoto, Y., Oemler, Jr., A., Lin, H., & Tucker, D. L. 1998, *ApJ*, 499, 589
- High, F. W., et al. 2010, *ApJ*, 723, 1736
- Hoekstra, H. 2007, *MNRAS*, 379, 317
- Holder, G., Haiman, Z., & Mohr, J. J. 2001, *ApJ*, 560, L111
- Hou, Z., Keisler, R., Knox, L., Millea, M., & Reichardt, C. 2011, *ArXiv:1104.2333*
- Ichikawa, K., Fukugita, M., & Kawasaki, M. 2005, *Phys. Rev. D*, 71, 043001
- Jeltema, T. E., Hallman, E. J., Burns, J. O., & Motl, P. M. 2008, *ApJ*, 681, 167
- Jenkins, A., Frenk, C. S., White, S. D. M., Colberg, J. M., Cole, S., Evrard, A. E., Couchman, H. M. P., & Yoshida, N. 2001, *MNRAS*, 321, 372
- Keisler, R., et al. 2011, *ApJ*, 743, 28
- Kirkman, D., Tytler, D., Suzuki, N., O’Meara, J. M., & Lubin, D. 2003, *ApJS*, 149, 1
- Komatsu, E., et al. 2009, *ApJS*, 180, 330
- . 2011, *ApJS*, 192, 18
- Kopp, J., Maltoni, M., & Schwetz, T. 2011, *Physical Review Letters*, 107, 091801
- Kravtsov, A. V., Vikhlinin, A., & Nagai, D. 2006, *ApJ*, 650, 128
- Larson, D., et al. 2011, *ApJS*, 192, 16
- Lewis, A., & Bridle, S. 2002, *Phys. Rev. D*, 66, 103511
- Lewis, A., Challinor, A., & Lasenby, A. 2000, *ApJ*, 538, 473
- Lin, Y., Partridge, B., Pober, J. C., Boucheffry, K. E., Burke, S., Klein, J. N., Coish, J. W., & Hufferberger, K. M. 2009, *ApJ*, 694, 992
- Linder, E. V., & Jenkins, A. 2003, *MNRAS*, 346, 573
- Lopez, R. E., Dodelson, S., Heckler, A., & Turner, M. S. 1999, *Physical Review Letters*, 82, 3952
- Mangano, G., Miele, G., Pastor, S., Pinto, T., Pisanti, O., & Serpico, P. D. 2005, *Nuclear Physics B*, 729, 221
- Mantz, A., Allen, S. W., Ebeling, H., & Rapetti, D. 2008, *MNRAS*, 387, 1179
- Mantz, A., Allen, S. W., Ebeling, H., Rapetti, D., & Drlica-Wagner, A. 2010a, *MNRAS*, 406, 1773
- Mantz, A., Allen, S. W., & Rapetti, D. 2010b, *MNRAS*, 406, 1805
- Mantz, A., Allen, S. W., Rapetti, D., & Ebeling, H. 2010c, *MNRAS*, 406, 1759
- Marriage, T. A., et al. 2011, *ApJ*, 731, 100
- Melin, J.-B., Bartlett, J. G., & Delabrouille, J. 2006, *A&A*, 459, 341
- Mention, G., Fechner, M., Lasserre, T., Mueller, T. A., Lhuillier, D., Cribier, M., & Letourneau, A. 2011, *ArXiv:1101.2755*
- Mortonson, M. J., Hu, W., & Huterer, D. 2011, *Phys. Rev. D*, 83, 023015
- Nagai, D., Kravtsov, A. V., & Vikhlinin, A. 2007, *ApJ*, 668, 1
- Ngeow, C. C., Mohr, J., Zenteno, A., Data Management, D., BCS, & SPT Collaborations. 2009, in *Bulletin of the American Astronomical Society*, Vol. 41, *Bulletin of the American Astronomical Society*, 335–+
- Padin, S., et al. 2008, *Appl. Opt.*, 47, 4418
- Pascale, E., et al. 2009, *ApJ*, 707, 1740
- Percival, W. J., et al. 2010, *MNRAS*, 401, 2148
- Planck Collaboration, et al. 2011, *A&A*, 536, A8
- Reichardt et al. 2012, *In Preparation*
- Reid, B. A., Verde, L., Jimenez, R., & Mena, O. 2010, *Journal of Cosmology and Astroparticle Physics*, 1, 3
- Riess, A. G., et al. 2011, *ApJ*, 730, 119
- Rozo, E., et al. 2010, *ApJ*, 708, 645
- Sehgal, N., Bode, P., Das, S., Hernandez-Monteagudo, C., Hufferberger, K., Lin, Y., Ostriker, J. P., & Trac, H. 2010, *ApJ*, 709, 920
- Sehgal, N., et al. 2011, *ApJ*, 732, 44
- Shaw, L. D., Holder, G. P., & Bode, P. 2008, *ApJ*, 686, 206
- Shaw, L. D., Zahn, O., Holder, G. P., & Doré, O. 2009, *ApJ*, 702, 368
- Stanek, R., Rasia, E., Evrard, A. E., Pearce, F., & Gazzola, L. 2010, *ApJ*, 715, 1508
- Staniszewski, Z., et al. 2009, *ApJ*, 701, 32
- Sunyaev, R. A., & Zel’dovich, Y. B. 1972, *Comments on Astrophysics and Space Physics*, 4, 173
- Tinker, J., Kravtsov, A. V., Klypin, A., Abazajian, K., Warren, M., Yepes, G., Gottlöber, S., & Holz, D. E. 2008, *ApJ*, 688, 709
- Vanderlinde, K., et al. 2010, *ApJ*, 722, 1180
- Vieira, J. D., et al. 2010, *ApJ*, 719, 763
- Vikhlinin, A., et al. 2009a, *ApJ*, 692, 1033
- . 2009b, *ApJ*, 692, 1060
- Wang, L., & Steinhardt, P. J. 1998, *ApJ*, 508, 483
- Warren, M. S., Abazajian, K., Holz, D. E., & Teodoro, L. 2006, *ApJ*, 646, 881
- White, M., Cohn, J. D., & Smit, R. 2010, *MNRAS*, 408, 1818
- Williamson, R., et al. 2011, *ApJ*, 738, 139

APPENDIX

X-RAY OBSERVATIONS AND RESULTS

In Table 7, we give updated X-ray observables for the clusters used in this work, as discussed in Section 2.1.3. For the five clusters without new measurements, we give the results from A11 directly, in order to provide a complete listing for the cluster sample. In Table 8, we give the complete list of *Chandra* observation identifications (ObsIDs) used for clusters with new *Chandra* observations

TABLE 7
X-RAY OBSERVABLES FOR SPT CLUSTERS

Name	z	r_{500} (kpc)	T_X (keV)	M_g ($10^{13} M_\odot$)	Y_X ($10^{14} M_\odot \text{ keV}$)
SPT-CL J0509-5342 ^a	0.463	1062 ± 39	7.0 ^{+1.4} _{-1.1}	5.6 ^{+0.2} _{-0.2}	4.3 ± 0.8
SPT-CL J0528-5300 ^b	0.765	765 ± 47	5.2 ^{+1.9} _{-1.2}	2.8 ^{+0.3} _{-0.3}	1.6 ± 0.5
SPT-CL J0533-5005 ^b	0.881	666 ± 51	3.9 ^{+1.6} _{-1.1}	2.3 ^{+0.5} _{-0.4}	1.0 ± 0.4
SPT-CL J0546-5345 ^b	1.067	823 ± 27	6.8 ^{+1.2} _{-0.9}	7.4 ^{+0.4} _{-0.3}	4.8 ± 0.8
SPT-CL J0551-5709 ^b	0.423	923 ± 34	4.0 ^{+0.6} _{-0.6}	5.1 ^{+0.6} _{-0.6}	1.9 ± 0.4
SPT-CL J0559-5249 ^a	0.611	1071 ± 30	7.7 ^{+1.1} _{-0.8}	8.3 ^{+0.3} _{-0.2}	6.4 ± 0.8
SPT-CL J2331-5051 ^a	0.571	972 ± 34	5.9 ^{+1.3} _{-0.8}	5.7 ^{+0.2} _{-0.2}	3.5 ± 0.6
SPT-CL J2332-5358 ^c	0.403	1166 ± 31	7.8 ^{+1.0} _{-0.9}	7.6 ^{+0.2} _{-0.3}	6.1 ± 0.8
SPT-CL J2337-5942 ^a	0.781	1046 ± 39	8.9 ^{+2.0} _{-1.4}	9.5 ^{+0.4} _{-0.6}	8.5 ± 1.7
SPT-CL J2341-5119 ^a	0.998	847 ± 37	8.0 ^{+1.9} _{-1.6}	5.6 ^{+0.2} _{-0.2}	4.7 ± 1.0
SPT-CL J2342-5411 ^c	1.074	648 ± 29	5.0 ^{+0.9} _{-0.8}	2.6 ^{+0.3} _{-0.3}	1.4 ± 0.3
SPT-CL J2355-5056 ^c	0.320	997 ± 31	5.3 ^{+0.9} _{-0.7}	3.9 ^{+0.2} _{-0.1}	2.2 ± 0.4
SPT-CL J2359-5009 ^{b,c}	0.774	778 ± 36	5.2 ^{+1.3} _{-0.9}	3.1 ^{+0.3} _{-0.3}	1.8 ± 0.4
SPT-CL J0000-5748 ^c	0.701	950 ± 68	8.3 ^{+3.6} _{-2.2}	4.4 ^{+0.5} _{-0.5}	4.2 ± 1.6

NOTE. — X-ray observables for clusters with *Chandra* or *XMM-Newton* observations. For clusters with new spectroscopic redshifts or new *Chandra* observations, we have recalculated their X-ray observables, as described in Section 2.1.3. To maintain consistency with A11, all X-ray observables are calculated assuming a preferred Λ CDM cosmology using WMAP7+BAO+ H_0 data with $\Omega_M = 0.272$, $\Omega_\Lambda = 0.728$ and $H_0 = 70.2 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (Komatsu et al. 2011).

^aX-ray observables taken from A11.

^bUpdated for new *Chandra* observations.

^cUpdated for new spectroscopic redshift.

TABLE 8
CLUSTERS WITH NEW *Chandra* X-RAY OBSERVATIONS

Name	z	A11		This Work		ObsIDs
		Exposure [ks]	Source Counts	Exposure [ks]	Source Counts	
SPT-CL J0528-5300	0.765	36.5	356	115.9	1732	9341, 10862, 11996, 11747 , 11874 , 12092 , 13126
SPT-CL J0533-5005	0.881	41.5	201	67.7	344	11748, 12001, 12002
SPT-CL J0546-5345	1.067	55.6	1304	67.8	1512	9332, 9336, 10851, 10864, 11739
SPT-CL J0551-5709	0.423	19.8	876	33.2	1212	11871, 11743
SPT-CL J2359-5009	0.774	57.9	713	122.4	1522	9334, 11742, 11864, 11997

NOTE. — The ObsIDs refer to all the *Chandra* observations used in this work. The ObsIDs that are new, relative to A11, are highlighted in **bold**. For clusters not listed here, we use the same *Chandra* and *XMM-Newton* observations as listed by A11.

LIKELIHOOD MODIFICATION TO ACCOUNT FOR COSMOLOGICAL DEPENDENCE OF Y_X

In Section 3.2 we presented a procedure for translating the theoretical mass function, $dN/dMdz$, into observable space, $dN/d\xi dY_X dz$. Under the assumption that this transformation is independent of the cosmological and scaling relation parameters \vec{p} we are ultimately trying to recover, this procedure modifies the log likelihood by a constant offset.

However, in the case of Y_X , this assumption ceases to hold true as Y_X is a derived quantity, calculated explicitly for each new value of \vec{p} (i.e., at each likelihood evaluation in the MCMC). This added subtlety can be addressed in the following fashion: Let us define Y_X^* as the value of Y_X when evaluated at some reference point in parameter space \vec{p}^* . In order for the probability contained in a differential volume to be independent of a change of variables, we need to

multiply by the Jacobian of the transformation, as follows

$$P(\vec{z}, \vec{\xi}, \vec{Y}_x^* | \vec{p}) = \left| \frac{\partial(\vec{Y}_x)}{\partial(\vec{Y}_x^*)} \right| P(\vec{z}, \vec{\xi}, \vec{Y}_x | \vec{p}) = \left| \prod_i f_i(\vec{p}) \right| P(\vec{z}, \vec{\xi}, \vec{Y}_x | \vec{p}) \quad (\text{B1})$$

where $f_i(\vec{p})$ is the ratio of the Y_X value of the i^{th} cluster at \vec{p} to its value at \vec{p}^* . Expressing this in terms of log probability, and ignoring constant offsets, we obtain

$$\ln P(\vec{z}, \vec{\xi}, \vec{Y}_x | \vec{p}) = \ln P(\vec{z}, \vec{\xi}, \vec{Y}_x^* | \vec{p}) + \sum_i \ln Y_{x,i} \quad (\text{B2})$$

This results in the straightforward prescription of adding $\sum \ln Y_X$ to the likelihood at each step in the MCMC, a process very similar to that suggested and performed in Mantz et al. (2008) and Vikhlinin et al. (2009a).

MASS ESTIMATES

We present posterior mass estimates for all 18 clusters considered in this work in Table 9. Where applicable, these are joint X-ray and SZ posterior mass estimates, for clusters without X-ray data we use the SZ posterior mass estimate. We calculate a probability density function on a mass grid at each point in the Λ CDM chain that was calculated using the CMB+SPT_{CL} data, from Section 4.1. The probability density functions are combined to obtain a mass estimate that has been fully marginalized over all cosmological and scaling relation parameters. We report the mean and the 68% confidence interval for the mass estimate.

TABLE 9
MASS ESTIMATES FOR THE SPT CLUSTER CATALOG

Object Name	ξ	z	$M_{500}(\rho_{crit})(10^{14} M_{\odot} h_{70}^{-1})$
SPT-CL J0509-5342	6.61	0.463	5.11 ± 0.68
SPT-CL J0511-5154 ^a	5.63	0.74	3.36 ± 0.86
SPT-CL J0521-5104 ^a	5.45	0.72	3.21 ± 0.86
SPT-CL J0528-5259	5.45	0.765	2.96 ± 0.54
SPT-CL J0533-5005	5.59	0.881	2.54 ± 0.54
SPT-CL J0539-5744 ^a	5.12	0.77	2.93 ± 0.86
SPT-CL J0546-5345	7.69	1.067	4.79 ± 0.64
SPT-CL J0551-5709	6.13	0.423	3.61 ± 0.54
SPT-CL J0559-5249	9.28	0.611	6.36 ± 0.79
SPT-CL J2301-5546 ^a	5.19	0.748	3.00 ± 0.86
SPT-CL J2331-5051	8.04	0.572	4.89 ± 0.68
SPT-CL J2332-5358	7.30	0.403	6.21 ± 0.79
SPT-CL J2337-5942	14.94	0.781	7.68 ± 1.04
SPT-CL J2341-5119	9.65	0.998	5.14 ± 0.71
SPT-CL J2342-5411	6.18	1.074	2.75 ± 0.46
SPT-CL J2355-5056	5.89	0.320	3.96 ± 0.54
SPT-CL J2359-5009	6.35	0.774	3.32 ± 0.54
SPT-CL J0000-5748	5.48	0.701	4.04 ± 0.68

^aThese clusters have only SZ data, and no X-ray observations.