

THE ULTRAVIOLET RADIATION ENVIRONMENT AROUND M DWARF EXOPLANET HOST STARS*

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ABSTRACT

The spectral and temporal behavior of exoplanet host stars is a critical input to models of the chemistry and evolution of planetary atmospheres. Ultraviolet photons influence the atmospheric temperature profiles and production of potential biomarkers on Earth-like planets around these stars. At present, little observational or theoretical basis exists for understanding the ultraviolet spectra of M dwarfs, despite their critical importance to predicting and interpreting the spectra of potentially habitable planets as they are obtained in the coming decades. Using observations from the *Hubble Space Telescope*, we present a study of the UV radiation fields around nearby M dwarf planet hosts that covers both far-UV (FUV) and near-UV (NUV) wavelengths. The combined FUV+NUV spectra are publicly available in machine-readable format. We find that all six exoplanet host stars in our sample (GJ 581, GJ 876, GJ 436, GJ 832, GJ 667C, and GJ 1214) exhibit some level of chromospheric and transition region UV emission. No “UV-quiet” M dwarfs are observed. The bright stellar Ly α emission lines are reconstructed, and we find that the Ly α line fluxes comprise $\sim 37\%$ – 75% of the total 1150–3100 Å flux from most M dwarfs; $\gtrsim 10^3$ times the solar value. We develop an empirical scaling relation between Ly α and Mg II emission, to be used when interstellar H I attenuation precludes the direct observation of Ly α . The intrinsic unreddened flux ratio is $F(\text{Ly}\alpha)/F(\text{Mg II}) = 10 \pm 3$. The $F(\text{FUV})/F(\text{NUV})$ flux ratio, a driver for abiotic production of the suggested biomarkers O₂ and O₃, is shown to be ~ 0.5 – 3 for all M dwarfs in our sample, $> 10^3$ times the solar ratio. For the four stars with moderate signal-to-noise Cosmic Origins Spectrograph time-resolved spectra, we find UV emission line variability with amplitudes of 50%–500% on 10^2 – 10^3 s timescales. This effect should be taken into account in future UV transiting planet studies, including searches for O₃ on Earth-like planets. Finally, we observe relatively bright H₂ fluorescent emission from four of the M dwarf exoplanetary systems (GJ 581, GJ 876, GJ 436, and GJ 832). Additional modeling work is needed to differentiate between a stellar photospheric or possible exoplanetary origin for the hot ($T(\text{H}_2) \approx 2000$ – 4000 K) molecular gas observed in these objects.

Key words: planetary systems – stars: activity – stars: individual (GJ 581, GJ 876, GJ 436, GJ 832, GJ 667C, GJ 1214) – stars: low-mass – ultraviolet: stars

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1. INTRODUCTION

Stellar ultraviolet (UV) photons play a crucial role in the dynamics and chemistry of all types of planetary atmospheres. Extreme-UV (EUV; $200 \lesssim \lambda \lesssim 911$ Å) photons capable of ionizing hydrogen heat the atmosphere and are the primary drivers of atmospheric escape in both the solar system and exoplanetary systems. For short-period gas giant planets, EUV irradiation can heat the atmosphere to $\sim 10^4$ K (Yelle 2004; Murray-Clay et al. 2009), resulting in atmospheric mass loss (Sanz-Forcada et al. 2011) and the inflated exospheres observed in H I Ly α (Vidal-Madjar et al. 2003; Lecavelier des Etangs et al. 2012) and metal absorption line (Vidal-Madjar et al. 2004; Linsky et al. 2010; Fossati et al. 2010) transit spectra. For terrestrial atmospheres, an increase in the EUV flux can elevate the temperature of the thermosphere by a factor

of $\gtrsim 10$ (Tian et al. 2008), potentially causing significant and rapid atmospheric mass loss.

Far-UV (FUV; $912 \lesssim \lambda \lesssim 1700$ Å) and near-UV (NUV; $1700 \lesssim \lambda \lesssim 4000$ Å) photons from the host star are the drivers of atmospheric chemistry through their effects on the photoexcitation and photodissociation rates of many abundant molecular species. Solar chromospheric emission (specifically H I Ly β) excites H₂ fluorescence in gas giant planets, producing rovibrationally excited H₂ that catalyzes chemistry in the Jovian atmosphere (Cravens 1987; Kim & Fox 1994). In exo-Jovian atmospheres heated to $T > 1000$ K, Ly α fluorescence of both H₂ and CO may become an important excitation mechanism (Wolven et al. 1997); this process may be an observable diagnostic for gas giant atmospheres outside the solar system (Yelle 2004; France et al. 2010b). H₂O, CH₄, and CO₂ are sensitive to FUV radiation, in particular the bright H I Ly α line which has spectral coincidences with these species. UV radiation can be important for the long-term habitability of Earth-like planets (Buccino et al. 2006), and the combination of FUV and NUV photons can influence the molecular oxygen chemistry. For example, the O₂ abundance can be significantly

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enhanced through the photodissociation of CO_2 and H_2O . The subsequent formation of O_3 depends sensitively on the spectral and temporal behavior of the FUV and NUV radiation fields of the host star (Tian et al. 2012; Domagal-Goldman et al. 2012).

The UV (FUV + NUV) spectra of solar-type stars have been studied extensively, both observationally and theoretically (e.g., Woods et al. 2009; Fontenla et al. 2011; Linsky et al. 2012a). By studying solar analogs of various ages, we also have an understanding of the evolution of G star spectra (Ayes 1997; Ribas et al. 2005; Linsky et al. 2012b). By contrast, the observational and theoretical literature on the UV spectra of M dwarfs is sparse. The photospheric UV continuum of M dwarfs is very low relative to solar-type stars due to their lower effective temperature. Bright chromospheric and transition region emission lines dominate the UV spectrum of M dwarfs, comprising a much larger fraction of the stellar luminosity than for solar-type stars. Outside of a few well-studied flare stars (e.g., AU Mic, AD Leo, EV Lac, Proxima Cen), M dwarfs have largely been ignored by UV observers because most investigations were aimed at understanding the origin and nature of energetic events on low-mass stars. There are very few UV observations of the older, quiescent M dwarfs that are most likely to harbor potentially habitable planets, and at present no theoretical atmosphere models exist that self-consistently predict the UV emission from an M dwarf.

UV variability of M dwarf exoplanet host stars is also essentially unconstrained by observation at present. M dwarfs display variability on many timescales and flare activity occurs with greater frequency and amplitude than on main-sequence G and K dwarfs (Hawley et al. 1996; West et al. 2004; Welsh et al. 2007; Walkowicz et al. 2011). Because most flare activity is thought to be related to magnetic energy deposition in the corona, transition region, and chromosphere (Haisch et al. 1991), UV flare activity has an amplitude comparable to or greater than that observed in optical light curves. UV flare behavior on active M dwarfs is characterized by strong blue/NUV continuum emission and the enhancement of chromospheric and transition region emission lines (e.g., Hawley & Pettersen 1991; Hawley et al. 2003). FUV flares may also be signpost for enhanced soft X-ray, EUV, and energetic particle arrival rates on planets orbiting these stars, all of which can play an important role for the chemistry and evolution of exoplanetary atmospheres (Segura et al. 2010) and the long-term habitability of these worlds (Buccino et al. 2007).

Given the paucity of existing observational data and lack of theoretical models of the UV spectra of M dwarfs, the most reliable incident radiation field for photochemical models of extrasolar planetary atmospheres is one created from direct observations of the host star. Toward the goal of a comprehensive library of UV radiation fields for use as inputs in models of exoplanet atmospheres, we have carried out a pilot program to observe the spectral and temporal behavior of M dwarf exoplanet host stars, Measurements of the Ultraviolet Spectral Characteristics of Low-mass Exoplanet host Stars (MUSCLES).

In this paper, we present observational results from the MUSCLES pilot program, spectrally and temporally resolved UV data for six M dwarf exoplanet host stars (GJ 581, GJ 876, GJ 436, GJ 832, GJ 667C, and GJ 1214) observed with the *Hubble Space Telescope* (HST) Cosmic Origins Spectrograph (COS) and Space Telescope Imaging Spectrograph (STIS). We describe the targets and their planetary systems in Section 2. Sections 3 and 4 describe the observations, the spectral reconstruction of the important stellar $\text{Ly}\alpha$ emission line, and the creation of light curves for chromospheric and transition region

emission lines. In Sections 5 and 6, we describe the spectral and temporal characteristics, respectively, of our target sample. Section 7 places these results in context of models of terrestrial planetary atmospheres and discusses implications for UV transit studies. We summarize our results in Section 8.

2. TARGETS

In this section, we briefly summarize the stellar and planetary system properties of our M dwarf targets. Optical spectra of most of our targets have been published in the literature, and while we do not make a detailed comparison of the UV and visible emission lines here, it is valuable to place the MUSCLES targets in the context of traditional optical activity indicators. The stars in our sample would traditionally be considered “optically inactive,” based on their $\text{H}\alpha$ absorption spectra (Gizis et al. 2002). However, all of our stars with measured Ca II H and K profiles show weak but detectable emission (equivalent widths, $\text{EW}(\text{Ca II}) > 0$), indicating that at least a low level of chromospheric activity is present in these stars (Rauscher & Marcy 2006; Walkowicz & Hawley 2009). Our stars show $\text{H}\alpha$ in absorption with equivalent widths in the range $-0.4 < \text{EW}(\text{H}\alpha) < -0.2 \text{ \AA}$ and Ca II emission with equivalent widths in the range $0.2 < \text{EW}(\text{Ca II}) < 0.5 \text{ \AA}$. Adopting the M dwarf classification from Walkowicz & Hawley (2009), these stars would be referred to as possessing intermediate chromospheres, or weakly active M dwarfs.

GJ 581. GJ 581 is an M2.5 dwarf at a distance of 6.3 pc. It is estimated to have an age of 8 ± 1 Gyr (Selsis et al. 2007) and a somewhat subsolar metallicity, $[\text{Fe}/\text{H}] = -0.10$ to -0.02 (Johnson & Apps 2009; Rojas-Ayala et al. 2010). GJ 581 is not detected in X-ray surveys ($\log_{10} L_X < 26.89 \text{ erg s}^{-1}$; Poppenhaeger et al. 2010) and its optical spectrum displays $\text{H}\alpha$ in absorption, therefore chromospheric and coronal activity are thought to be low for this target. GJ 581 has one of the richest known planetary systems, with possibly up to six planets (four confirmed) including several of Earth/super-Earth mass (Mayor et al. 2009; Tuomi 2011). GJ 581d is a super-Earth ($M_P \approx 6 M_\oplus$) that resides on the outer edge of the habitable zone (HZ; $a_P = 0.22 \text{ AU}$; Wordsworth et al. 2011; von Braun et al. 2011).

GJ 876. GJ 876 is an M4 dwarf at 4.7 pc, and is the only planet-hosting M dwarf with a well-characterized UV spectrum prior to the present work (Walkowicz et al. 2008; France et al. 2012a). GJ 876 has super-solar metallicity ($[\text{Fe}/\text{H}] = 0.37\text{--}0.43$; Johnson & Apps 2009; Rojas-Ayala et al. 2010), and differing estimates on the stellar rotation period ($40 \text{ days} \leq P_* \lesssim 97 \text{ days}$) result in large uncertainties in the age estimate for this system, 0.1–5 Gyr (Rivera et al. 2005, 2010; Correia et al. 2010). While the star would be characterized as weakly active based on its $\text{H}\alpha$ absorption spectrum, UV and X-ray observations have shown the presence of an active upper atmosphere (Walkowicz et al. 2008; France et al. 2012a; $\log_{10} L_X = 26.48 \text{ erg s}^{-1}$; Poppenhaeger et al. 2010). GJ 876 has a rich planetary system, with four planets ranging from a super-Earth (GJ 876d, $M_P \approx 6.6 M_\oplus$) in a short-period orbit ($a_P = 0.02 \text{ AU}$; Rivera et al. 2010) to two Jovian-mass planets in the HZ (GJ 876b, $M_P \approx 2.27 M_{\text{Jup}}$, $a_P = 0.21 \text{ AU}$; GJ 876c, $M_P \approx 0.72 M_{\text{Jup}}$, $a_P = 0.13 \text{ AU}$; Rivera et al. 2010).

GJ 436. GJ 436 is an M3 dwarf star located at a distance of 10.3 pc. It has a 45 day rotation period, a relatively old age ($\sim 6_{-5}^{+4}$ Gyr; Torres 2007), and may have a super-solar metallicity ($[\text{Fe}/\text{H}] = 0.00\text{--}0.25$; Johnson & Apps 2009; Rojas-Ayala et al. 2010). GJ 436 does show signs of an active corona ($\log_{10} L_X = 27.16 \text{ erg s}^{-1}$; Poppenhaeger et al. 2010), and its chromospheric

$\text{Ly}\alpha$ emission has been observed by Ehrenreich et al. (2011). GJ 436 is notable for its well-studied transiting Neptune-mass planet (Butler et al. 2004; Pont et al. 2009), orbiting at a semimajor axis of ≈ 0.03 AU, interior to its HZ (0.16–0.31 AU; von Braun et al. 2012). Additional low-mass planets may also be present in this system (Stevenson et al. 2012).

GJ 832. GJ 832 is an M1 dwarf at $d = 4.9$ pc. GJ 832 is not as well characterized as other targets in our sample; an age determination for this star is not available. Coronal X-rays have been detected from GJ 832, $\log_{10} L_X = 26.77 \text{ erg s}^{-1}$ (Poppenhaeger et al. 2010). This subsolar metallicity star ($[\text{Fe}/\text{H}] = -0.12$; Johnson & Apps 2009) hosts a $0.64 M_{\text{Jup}}$ mass planet in a 9.4 year orbit ($a_p = 3.4$ AU; Bailey et al. 2009).

GJ 667C. GJ 667C (M1.5V) is a member of a triple star system (GJ 667AB is a K3V + K5V binary) at a distance of 6.9 pc. This 2–10 Gyr M dwarf (Anglada-Escudé et al. 2012) is metal-poor ($[\text{Fe}/\text{H}] = -0.59 \pm 0.10$, based on an analysis of GJ 667AB; Perrin et al. 1988) and may host as many as three planets, including a super-Earth mass planet (GJ 667Cc, $M_p \approx 4.5 M_{\oplus}$, $a_p = 0.12$ AU) orbiting in the HZ (0.11–0.23 AU; Anglada-Escudé et al. 2012).

GJ 1214. GJ 1214 is a late M dwarf (M6V) at 13 pc, making it the coolest and most distant of the targets in the MUSCLES pilot study. It has an age of 6 ± 3 Gyr (Charbonneau et al. 2009), a super-solar metallicity ($[\text{Fe}/\text{H}] = +0.39 \pm 0.15$; Berta et al. 2011; Rojas-Ayala et al. 2010), and shows signs of optical flare activity (Kundurthy et al. 2011). GJ 1214b is a transiting super-Earth ($M_p \approx 6.5 M_{\oplus}$, $a_p = 0.014$ AU; Charbonneau et al. 2009), possibly harboring a dense, water-rich atmosphere (Bean et al. 2010; Désert et al. 2011).

3. OBSERVATIONS

The MUSCLES observing plan uses the two primary ultraviolet spectrographs on *HST* to create quasi-continuous M dwarf exoplanet host star data from 1150 to 3140 Å. Data for the MUSCLES program were obtained as part of *HST* GTO and GO programs 12034, 12035, and 12464, acquired between 2011 June and 2012 August. COS has a factor of >10 times more effective area than the medium resolution modes of STIS in the FUV, and its low detector background and grating scatter make it factors of ~ 50 times more efficient for the study of faint far-UV chromospheric and transition region emission from M dwarfs (Green et al. 2012). We used COS, in its time-tag (TTAG) acquisition mode, to observe the FUV spectra of our target stars and study time variability on scales from minutes to hours (Section 4.2 and Section 6). Multiple central wavelengths and focal-plane offset (FP-POSS) settings with the COS G130M and G160M modes were used to create a continuous FUV spectrum from 1145 to 1795 Å. These modes provide a point-source resolution of $\Delta v \approx 17 \text{ km s}^{-1}$ with 7 pixels per resolution element (Osterman et al. 2011).

Because COS is a slitless spectrograph, observations at H I $\text{Ly}\alpha$ are heavily contaminated by geocoronal emission. We therefore observed the stellar $\text{Ly}\alpha$ profile with STIS, using either the G140M/cenwave 1222 mode ($\Delta v \approx 30 \text{ km s}^{-1}$) through the $52'' \times 0'.1$ slit, or the E140M mode ($\Delta v \approx 7.5 \text{ km s}^{-1}$) through the $0'.2 \times 0'.2$ slit for the brightest targets. In order to include the important $\text{Ly}\alpha$ emission line for GJ 436 and GJ 1214, we downloaded STIS G140M observations of these stars from the MAST archive (program IDs 11817 and 12165, respectively). The $\text{Ly}\alpha$ profile of GJ 436 has been described in Ehrenreich et al. (2011). In order to maximize the combination of wavelength coverage and sensitivity, near-UV spectra were

Table 1
MUSCLES *HST* Observations

Target	Date	Instrument, Mode	PID	T_{exp} (s)
GJ 581	2012 Apr 18	STIS, G140M	12034	1107
GJ 581	2012 Apr 18	STIS, G230L	12034	265
GJ 581	2011 Jul 20	COS, G130M	12034	1478
GJ 581	2011 Jul 20	COS, G160M	12034	900
GJ 876	2011 Nov 12	STIS, G140M	12464	1138
GJ 876	2011 Nov 12	STIS, G230L	12464	267
GJ 876	2012 Jan 5	COS, G130M	12464	2017
GJ 876	2012 Jan 5	COS, G160M	12464	2779
GJ 436	2010 Jan 5	STIS, G140M	11817	1762
GJ 436	2012 May 10	STIS, G230L	12464	1665
GJ 436	2012 Jun 23	COS, G130M	12464	3372
GJ 436	2012 Jun 23	COS, G160M	12464	4413
GJ 832	2011 Jun 9	STIS, E140M	12035	2572
GJ 832	2012 Jun 10	STIS, E230H	12035	3135
GJ 832	2012 Apr 10	STIS, G230L	12464	917
GJ 832	2012 Jul 28	COS, G130M	12464	2163
GJ 832	2012 Jul 28	COS, G160M	12464	2925
GJ 667C	2011 Sep 4	STIS, E140M	12035	2396
GJ 667C	2011 Sep 4	STIS, E230H	12035	3023
GJ 1214	2011 Apr 27	STIS, G140M	12165	7282
GJ 1214	2012 Aug 12	STIS, G230L	12464	1620
GJ 1214	2012 Aug 4	COS, G130M	12464	3289
GJ 1214	2012 Aug 4	COS, G160M	12464	4368

observed with the STIS G230L/cenwave 2376 mode ($\Delta v \approx 600 \text{ km s}^{-1}$) through the $52'' \times 0'.1$ slit. In order to constrain the effects of interstellar absorption on the stellar Mg II emission profile, we also acquired Mg II spectra with the STIS E230H mode ($\Delta v \approx 2.6 \text{ km s}^{-1}$) through the $0'.2 \times 0'.2$ slit for GJ 832 and GJ 667C.

Exposure times were generally short due to the pilot-study nature of this project; considerably longer integrations will be essential for future observations to measure UV flare frequencies in exoplanet host stars and to produce high-quality line profiles of important chromospheric and transition region tracers such as C II $\lambda 1335$, N V $\lambda 1240$, and Si IV $\lambda 1400$. A complete list of the M dwarf observations used in this work is presented in Table 1, and complete spectra for the four targets in which $\text{Ly}\alpha$, FUV, and NUV emission is observed are displayed in Figure 1.

In order to place the M dwarf exoplanet host stars in context with other well-studied cool stars, we assembled archival spectra of AD Leo (M3.5Ve), HD 189733 (K1V), and the Sun (G2V). AD Leo is one of the best-studied flare stars, and despite its extreme activity levels it has been the object of choice for models of extrasolar planets orbiting M dwarfs, mainly because of the lack of reasonable alternatives. We created a complete UV spectrum of AD Leo by combining STIS E140M monitoring observations (Hawley et al. 2003) with an “average” NUV spectrum from *IUE*. The *IUE* spectrum is the average of flare and quiescent states, scaled to the FUV flux from STIS. HD189733 is an active K dwarf hosting the prototypical transiting hot Jupiter. While high-quality COS spectra of HD 189733 exist in the archive (e.g., Linsky et al. 2012b; Haswell et al. 2012), we elected to use the combined NUV+FUV (STIS E140M + E230M) spectra of a surrogate star taken from the STARCAT library (Ayres 2010) to eliminate uncertainties about temporal variability between different observations. We chose HD37394 (K1V variable star; $V = 6.23$) because it is a good match in stellar mass and activity level to HD 189733, and because the FUV spectra of the two stars are qualitatively similar. The UV

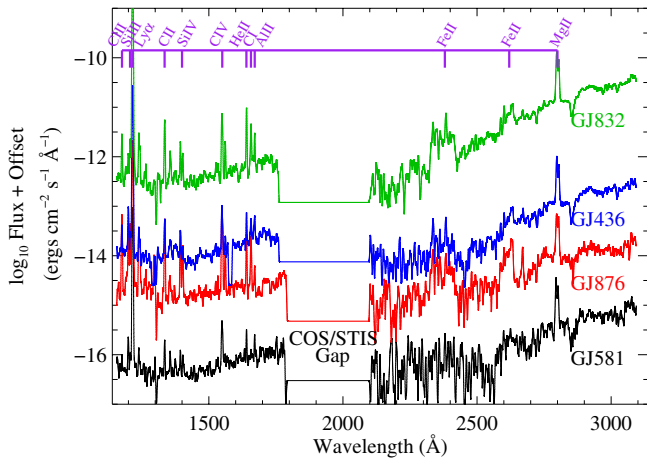


Figure 1. Observed 1150–3140 Å fluxes from the M dwarf exoplanet host stars in the MUSCLES sample with complete (Ly α + FUV + NUV) data sets. The 1760–2100 Å emission has been omitted because the flux levels are below the STIS G230L detection level. For display purposes, the data have been convolved with a 3 Å FWHM Gaussian kernel and offset as follows: GJ 581, no offset; GJ 876, $\log_{10} F_{\lambda} + 1.2$; GJ 436, $\log_{10} F_{\lambda} + 2.4$; GJ 832, $\log_{10} F_{\lambda} + 3.6$.

(A color version of this figure is available in the online journal.)

spectrum of HD37394 was then scaled to the V -magnitude of HD 189733 ($V = 7.77$). The spectrum of the quiet Sun was taken from Woods et al. (2009), and serves as the prototype main-sequence G-type star.

4. INTRINSIC Ly α PROFILE RECONSTRUCTION AND TIME VARIABILITY ANALYSIS

4.1. Ly α Reconstruction

The stellar Ly α emission line dominates the UV output of M dwarfs, containing approximately as much energy as the rest of the FUV+NUV spectrum combined (France et al. 2012a). Therefore, measurements of Ly α emission provide an important constraint on the source term for calculations of the heating and photochemistry of exoplanetary atmospheres (Linsky et al. 2012a). Due to resonant scattering of neutral hydrogen and deuterium in the interstellar medium (ISM), the intrinsic Ly α radiation field cannot be directly measured, even in the nearest stars ($\log_{10} N(\text{H I}) \lesssim 18.5$). In order to produce an accurate estimate of the local Ly α flux incident on the planets around our target stars, the Ly α profiles observed by STIS must be reconstituted. We do this using a technique that simultaneously fits the observed Ly α emission line profile and the ISM component, using the MPFIT routine to minimize χ^2 between the fit and data (Markwardt 2009).

The intrinsic Ly α emission line is approximated as a two-component Gaussian, as suggested for transition region emission lines from late-type stars by Wood et al. (1997). This approximation is further justified by the reconstructed Ly α profiles for M dwarfs, which show approximately Gaussian line shapes and little evidence for self-reversal compared to F, G, and K dwarfs (Wood et al. 2005). The Ly α emission components are characterized by an amplitude, FWHM, and velocity centroid. Interstellar absorption is parameterized by the column density of neutral hydrogen ($N(\text{H I})$), Doppler b parameter, D/H ratio, and the velocity of the interstellar absorbers. We assumed a fixed D/H ratio ($\text{D}/\text{H} = 1.5 \times 10^{-5}$; Linsky et al. 1995, 2006). Degeneracy between the two emission components makes errors difficult to define for the individual fit parameters. We estimate the uncertainty on the integrated Ly α fluxes from our technique

by comparing the reconstructed Ly α flux for a range of initial guesses on the parameters. The uncertainty on the integrated intrinsic flux is $\approx 10\%$ – 20% for the stars observed with E140M (GJ 832 and GJ 667C) and $\approx 15\%$ – 30% for the stars observed with G140M (GJ 581, GJ 876, GJ 436). In order to test our methodology, we fitted the STIS E140M Ly α profile of AU Mic. Our total integrated flux agrees with that presented in Wood et al. (2005) to $\approx 5\%$, with an identical derivation of the interstellar atomic hydrogen column on the line of sight ($\log_{10} N(\text{H I}) = 18.36 \pm 0.01$).

Employing our iterative least-squares reconstruction technique, we also derived the intrinsic Ly α spectrum of HD 189733. Using the 2010 April 6 STIS G140M spectra (Lecavelier des Etangs et al. 2012), we find an intrinsic flux of $F(\text{Ly}\alpha) = 7.5 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ and an interstellar hydrogen column density $\log_{10} N(\text{H I}) = 18.45$. We also refitted the intrinsic Ly α emission from AD Leo, finding $F(\text{Ly}\alpha) = 7.5 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ and an interstellar hydrogen column density $\log_{10} N(\text{H I}) = 18.47 \pm 0.02$. This value agrees with the Ly α reconstruction from Wood et al. (2005) to $\approx 30\%$, which gives us confidence in the iterative technique, even for sightlines with multiple interstellar velocity components.

We present the individual Ly α spectra in Figure 2, and narrow/broad emission component ratios ($F_{\text{narrow}}/F_{\text{broad}}$), spectral line widths, H I column densities, and Doppler b -values in Table 2. We find $F_{\text{narrow}}/F_{\text{broad}}$ to be in the range 5–12 for the MUSCLES stars, again independent of the resolution of the observations on which the reconstructions are based. For comparison, we find $F_{\text{narrow}}/F_{\text{broad}} \sim 3$ for the more active M dwarf (AD Leo) and more massive star (HD 189733). Reconstructed Ly α fluxes are given in Tables 3, 4, and 5. GJ 1214 has no Ly α emission detected, and we discuss this in greater detail in Section 5.1.1.

4.2. Timing Analysis

In our initial study of GJ 876, we described the large chromospheric/transition region flare observed in several UV emission lines during our COS observations (France et al. 2012a). We have carried out a similar analysis on the entire MUSCLES sample to constrain the heretofore unknown level of UV variability in optically quiet M dwarf exoplanet host stars. Light curves in several chromospheric and transition region lines were extracted from the calibrated three-dimensional data by exploiting the time-tag capability of the COS microchannel plate detector (France et al. 2010a). We extract a $[\lambda_i, y_i, t_i]$ photon list (where λ is the wavelength of the photon, y is the cross-dispersion location, and t is the photon arrival time) from each exposure i and combine these to create a master $[\lambda, y, t]$ photon list. The total number of counts in a $[\Delta\lambda, \Delta y]$ box is integrated over a time step Δt . We use a flux-dependent time step of $\Delta t = 40$ – 120 s for the MUSCLES targets. The instrument background level is computed in a similar manner, with the background integrated over the same wavelength interval as the emission lines, but physically offset below the science region in the cross-dispersion direction. The signal-to-noise ratio (S/N) was not high enough to temporally resolve the GJ 1214 data.

5. STELLAR UV EMISSION LINES AND THE FUV/NUV RATIO

5.1. Chromospheric and Transition Region Emission

Chromospheric and transition region emission lines are observed in all of the MUSCLES spectra, suggesting that

Table 2
Stellar Ly α Emission Line Widths and Interstellar Atomic Hydrogen^a

Target	$F_{\text{narrow}}/F_{\text{broad}}$	FWHM _{narrow} (km s ⁻¹)	FWHM _{broad} (km s ⁻¹)	log ₁₀ $N(\text{H I})$ (cm ⁻²)	$b_{\text{H I}}$ (km s ⁻¹)
GJ 581 ^b	8.4	110	228	18.35 ± 0.06	10.1
GJ 876 ^b	5.0	132	303	18.06 ± 0.03	8.0
GJ 436 ^b	11.8	109	294	18.19 ± 0.04	7.4
GJ 832 ^c	7.6	96	163	18.47 ± 0.02	9.7
GJ 667C ^c	8.9	85	233	18.07 ± 0.03	12.7
AD Leo ^c	2.9	166	412	18.47 ± 0.01	9.0
HD 189733 ^{b,d}	3.5	170	480	18.45	10

Notes.^a Ly α profiles reconstructed using the iterative least-squares technique described in Section 4.1.^b STIS G140M observations.^c STIS E140M observations.^d HD 189733 reconstruction failed to converge, the presented profile parameters provide a fit to the observed line profile that is qualitatively similar to the formal fits for the M dwarf spectra.

Table 3
M dwarf Broadband UV and Reconstructed Ly α Fluxes (erg cm⁻² s⁻¹)^a

Target	d (pc)	$F(\text{Ly}\alpha)^{\text{b}}$ ($\Delta\lambda^{\text{d}} = 1210\text{--}1222 \text{ \AA}$)	$F(\text{FUV})^{\text{c}}$ ($\Delta\lambda = 1160\text{--}1690 \text{ \AA}$)	$F(\text{NUV})^{\text{c}}$ ($\Delta\lambda = 2300\text{--}3050 \text{ \AA}$)
GJ 581	6.3	3.0×10^{-13}	3.6×10^{-14}	2.3×10^{-13}
GJ 876	4.7	4.4×10^{-13}	1.7×10^{-13}	3.9×10^{-13}
GJ 436	10.2	3.5×10^{-13}	3.3×10^{-14}	2.8×10^{-13}
GJ 832	4.9	5.0×10^{-12}	9.8×10^{-14}	2.2×10^{-12}
GJ 667C	6.9	7.6×10^{-13}
GJ 1214	13.0	$<2.4 \times 10^{-15}$	1.0×10^{-14}	2.0×10^{-14}
AD Leo	4.7	7.5×10^{-12}	3.4×10^{-12}	9.3×10^{-12}
HD 189733 ^e	19.5	7.5×10^{-13}	4.1×10^{-13}	1.0×10^{-10}
Sun	1 AU	5.9×10^0	5.9×10^0	1.6×10^4

Notes.^a Flux measurements are averaged over all exposure times.^b Uncertainty on the reconstructed Ly α flux is estimated to be between 10%–20% for the stars observed with E140M (GJ 832, GJ 667C, and AD Leo) and 15%–30% for stars observed with G140M (GJ 581, GJ 876, GJ 436, and HD 189733).^c $F(\text{FUV})$ refers to the integrated 1160–1690 Å flux *excluding* Ly α . Flux uncertainties for the broad band measurements are dominated by instrumental calibrations and target acquisition errors for faint sources. The flux errors for $F(\text{FUV})$ and $F(\text{NUV})$ are estimated to be $\leq 10\%$ and $\leq 5\%$, respectively.^d $\Delta\lambda$ is the bandpass over which the flux is integrated.^e HD 189733 Ly α reconstruction made from direct STIS G140M observations. $F(\text{FUV})$ and $F(\text{NUV})$ were measured using a proxy star (HD 37394) scaled to the V -magnitude of HD 189733.

Table 4
M dwarf Broadband UV Luminosity (erg s⁻¹)^a

Target	d (pc)	log ₁₀ $L(\text{Ly}\alpha)$ ($\Delta\lambda^{\text{c}} = 1210\text{--}1222 \text{ \AA}$)	log ₁₀ $L(\text{FUV})^{\text{b}}$ ($\Delta\lambda = 1160\text{--}1690 \text{ \AA}$)	log ₁₀ $L(\text{NUV})$ ($\Delta\lambda = 2300\text{--}3050 \text{ \AA}$)
GJ 581	6.3	27.16	26.23	27.03
GJ 876	4.7	27.07	26.65	27.01
GJ 436	10.2	27.65	26.62	27.55
GJ 832	4.9	28.15	26.44	27.80
GJ 1214	13.0	<25.69	26.31	26.61
		log ₁₀ $L(\text{Ly}\alpha)/L_{\text{bol}}^{\text{d}}$	log ₁₀ $L(\text{FUV})/L_{\text{bol}}$	log ₁₀ $L(\text{NUV})/L_{\text{bol}}$
GJ 581	6.3	-4.11	-5.04	-4.24
GJ 876	4.7	-4.26	-4.68	-4.32
GJ 436	10.2	-4.07	-5.10	-4.17
GJ 832	4.9	-3.50	-5.21	-3.85
GJ 1214	13.0	<-6.24	-5.62	-5.31

Notes.^a Flux measurements are averaged over all exposure times.^b $L(\text{FUV})$ refers to the integrated 1160–1690 Å flux *excluding* Ly α .^c $\Delta\lambda$ is the bandpass over which the flux is integrated.^d L_{bol} is the integral of the X-ray through IR flux of each star (Section 5.2).

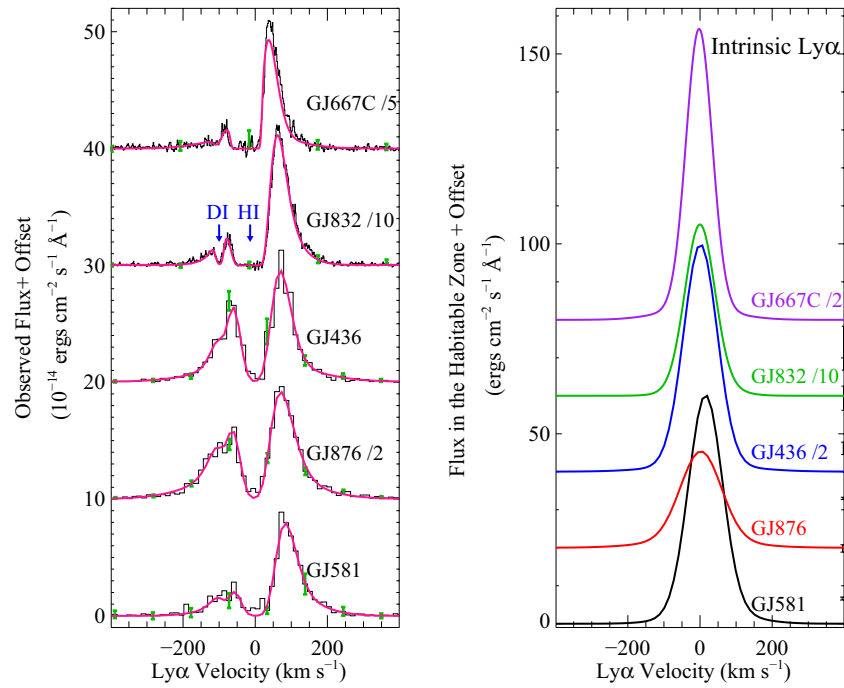


Figure 2. Left: *HST*/STIS E140M and G140M Ly α profiles (black histograms). The two-component Gaussian emission line + interstellar absorption model is overplotted in pink. The neutral hydrogen and deuterium absorption components are marked with blue arrows. The spectra are scaled as marked and offset in $+10 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ steps; photometric uncertainties are displayed as green error bars. Right: intrinsic two-component Gaussian model to the stellar Ly α emission line, scaled to the flux at 0.16 AU, in the habitable zone of these M dwarf exoplanet host stars. The spectra are scaled as marked and offset in $+20 \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ steps.

(A color version of this figure is available in the online journal.)

Table 5
Brightest Chromospheric and Transition Region M dwarf Emission Line Fluxes ($\text{erg cm}^{-2} \text{ s}^{-1}$)^a

Target	d (pc)	F^b (Ly α)	F^c (C IV)	F^c (Mg II)
GJ 581	6.3	3.0×10^{-13}	$1.80 (\pm 0.16) \times 10^{-15}$	$2.13 (\pm 0.13) \times 10^{-14}$
GJ 876	4.7	4.4×10^{-13}	$1.99 (\pm 0.17) \times 10^{-14}$	$3.37 (\pm 0.23) \times 10^{-14}$
GJ 436	10.2	3.5×10^{-13}	$1.38 (\pm 0.11) \times 10^{-15}$	$2.30 (\pm 0.13) \times 10^{-14}$
GJ 832	4.9	5.0×10^{-12}	$7.30 (\pm 0.48) \times 10^{-15}$	$2.98 (\pm 0.18) \times 10^{-13}$
GJ 667C	6.9	7.6×10^{-13}	...	$5.59 (\pm 0.47) \times 10^{-14}$
GJ 1214	13.0	$<2.4 \times 10^{-15}$	$2.62 (\pm 0.47) \times 10^{-16}$	$2.21 (\pm 0.15) \times 10^{-15}$

Notes.

^a Flux measurements are averaged over all exposure times.

^b Uncertainty on the reconstructed Ly α flux is estimated to be between 10%–20% for the stars observed with E140M (GJ 832 and GJ 667C) and 15%–30% for stars observed with G140M (GJ 581, GJ 876, and GJ 436).

^c C IV and Mg II fluxes are Gaussian fits to both lines of the doublets.

many M dwarf exoplanet host stars have UV-active atmospheres (Figure 1). The chromosphere is observed in H I Ly α , C II $\lambda\lambda 1334, 1335$, Al II $\lambda 1671$, Fe II multiplets near 2400 and 2600 \AA , and Mg II $\lambda\lambda 2796, 2803$; lines that are formed from $\sim 1\text{--}3 \times 10^4$ K. The transition region is traced by lines with formation temperatures from $\sim 0.6\text{--}1.6 \times 10^5$ K, including the C III 1175 multiplet, Si III $\lambda 1206$, O V $\lambda 1218$, N V $\lambda\lambda 1239, 1243$, Si IV $\lambda\lambda 1394, 1403$, C IV $\lambda\lambda 1548, 1550$, and He II $\lambda 1640$. C II and C IV line fluxes for GJ 581, GJ 876, GJ 436, and GJ 832 have been published in Linsky et al. (2012a), and a comprehensive set of stellar emission line measurements will be presented in a future work. In Table 5, we present the reconstructed Ly α fluxes and the observed fluxes of C IV and Mg II.

Figure 3 displays the C IV profiles for GJ 581, GJ 876, GJ 436, and GJ 832 scaled to the flux level at 0.16 AU from the parent star, a value near the center of the HZ for the MUSCLES

targets.⁸ While we conclude that most M dwarf exoplanet hosts have a non-zero local UV radiation field, the amplitude of this field is quite variable. The HZ C IV fluxes from these four stars varies by a factor of 10 from the weakest (GJ 581) to the strongest (GJ 832) C IV emitter. This is particularly interesting in light of the established correlation between X-ray and C IV luminosity for M dwarf emission line (dMe) flare stars (Byrne & Doyle 1989). With improved atmospheric models for M dwarfs or simultaneous UV and X-ray observations of a larger sample of M dwarf planet hosts, the C IV flux in the HZ of low-mass stars could be calibrated as a proxy for the EUV and soft X-ray emission incident on the planetary atmospheres (Linsky et al. 2012a). C IV-based EUV irradiance estimates could provide a

⁸ von Braun et al. (2012) compute the HZ of GJ 436 to extend 0.16–0.31 AU, so 0.16 AU may represent the inner edge of the HZ for this system.

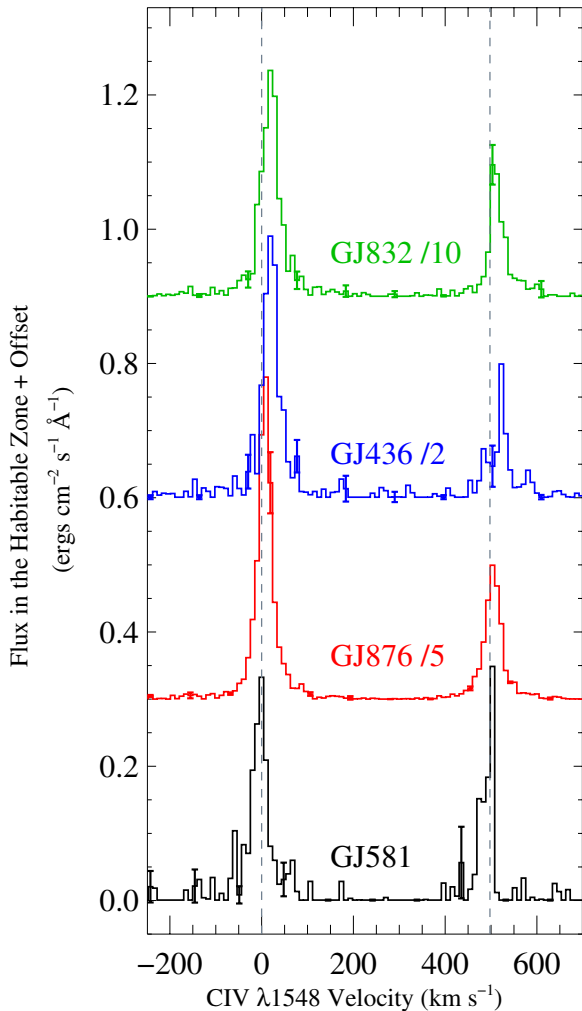


Figure 3. CIV $\lambda\lambda 1548, 1550$ resonance doublets for the four stars with Ly α detections and FUV spectra, plotted in the heliocentric velocity frame. The emission lines have been scaled to the flux at 0.16 AU, in the habitable zone of these M dwarf exoplanet host stars, and offset in $0.3 \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ steps. The dashed lines are the rest velocities of the 1548 \AA and 1550 \AA lines, respectively. (A color version of this figure is available in the online journal.)

complementary constraint to those derived from coronal X-ray measurements (Sanz-Forcada et al. 2011).

Ly α is the brightest line in the UV spectrum of low-mass stars, although resonant scattering in the ISM makes direct line profiles inaccessible, even for the nearest stars (Wood et al. 2000, 2005). As described above, Ly α reconstruction can be a powerful technique for determining the intrinsic stellar Ly α line flux, but this technique is not possible for more distant targets where the ISM completely removes stellar Ly α photons from our line of sight ($N(\text{H I}) \gtrsim 10^{20} \text{ cm}^{-2}$). In this case, one requires a proxy formed in a similar region of the chromosphere to estimate the local Ly α emission flux. For our limited sample, we have found that Ly α and Mg II are well correlated (see also Linsky et al. 2012a), and we present a scaling relation to infer the Ly α flux from an Mg II observation.

Figure 4 shows the reconstructed Ly α flux versus the observed Mg II flux, scaled to 0.16 AU from the host star for GJ 581, GJ 876, GJ 436, GJ 832, and GJ 667C (black filled circles). A linear fit to the correlation ($F(\text{Ly}\alpha) = a + bF(\text{Mg II})$), where $a = -2.3 \pm 2.2$ and $b = +14.8 \pm 1.2$), is also displayed in Figure 4. The observed $F(\text{Ly}\alpha)/F(\text{Mg II})$ ratio is the

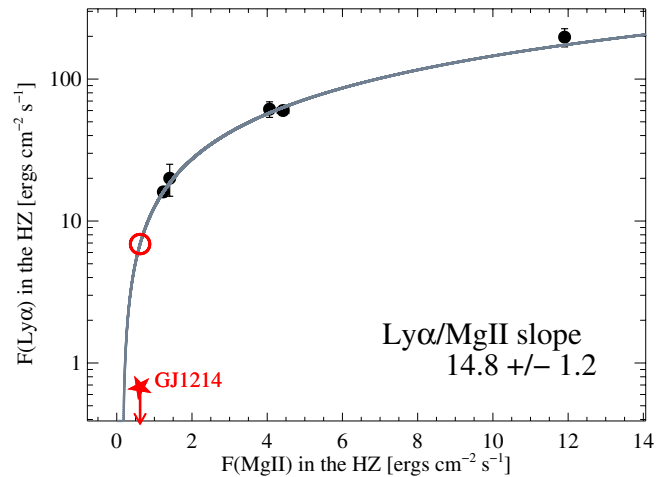


Figure 4. Intrinsic, reconstructed Ly α emission fluxes compared with the observed Mg II ($h + k$) flux, both scaled to the flux at 0.16 AU from the star. The Ly α and Mg II fluxes are observed to be well correlated. The slope of the relation is 14.8 ± 1.2 . After correcting for the effects of interstellar absorption of the Mg II lines, the $F(\text{Ly}\alpha)/F(\text{Mg II})$ ratio for weakly active M dwarf exoplanet host stars is 10 ± 3 . The expected Ly α emission line strength of GJ 1214 is shown as the open red circle. The observed GJ 1214 Ly α emission line flux upper limit is shown as the filled red star.

(A color version of this figure is available in the online journal.)

slope of this fit. Chromospheric Mg II emission lines of G and K dwarfs are known to show self-reversed profiles due to non-LTE effects in their upper atmospheres (e.g., Linsky & Wood 1996; Wood & Linsky 1998; Redfield & Linsky 2002); however, the situation is less clear for M dwarfs. For example, Redfield et al. (2002) present GHRS spectra of AU Mic, where no appreciable self-reversal of the reconstructed Mg II profile is observed. We therefore assume a simple Gaussian line shape for the M dwarf Mg II emission lines. Interstellar Mg II absorption will produce the appearance of a self-reversed emission line when the stellar velocity is close to the velocity of the Mg⁺-bearing interstellar cloud (e.g., Wood et al. 2000); therefore a correction for interstellar Mg II absorption needs to be made in order to give an unbiased estimate of the $F(\text{Ly}\alpha)/F(\text{Mg II})$ ratio for the MUSCLES stars.

We cannot analyze the interstellar Mg II absorption in detail for most of our targets because the stellar and interstellar features are unresolved in STIS G230L observations. However, utilizing our STIS E230H high-resolution ($\Delta v < 3 \text{ km s}^{-1}$) spectra of GJ 832 and GJ 667C, the correction for interstellar Mg II absorption can be estimated. The profile of the Mg II k line ($\lambda_o = 2795.528 \text{ \AA}$) of GJ 832 is shown in Figure 5. For GJ 832 and GJ 667C, we find that the interstellar absorber removes 30%–35% of the intrinsic Gaussian emission line flux. Therefore, we estimate that the scaling of the ISM-corrected Mg II emission line flux to the intrinsic Ly α emission line flux is ≈ 10 . Factoring in uncertainties on the interstellar correction and the Ly α reconstruction, we estimate that this scaling relation is good to $\sim 30\%$, or the ISM-corrected intrinsic $F(\text{Ly}\alpha)/F(\text{Mg II})$ ratio for weakly active M dwarf exoplanet host stars is 10 ± 3 . Linsky et al. (2012a) find that the $F(\text{Ly}\alpha)/F(\text{Mg II})$ ratios for the more active M dwarfs AU Mic and AD Leo are 2.5 and 4.0, respectively. This suggests that the ratio may decrease with activity.

5.1.1. Non-detection of Ly α in GJ 1214

GJ 1214 is the coolest (M6V, $T_{\text{eff}} 2949 \pm 30 \text{ K}$; Kundurthy et al. 2011) and most distant source in the MUSCLES pilot

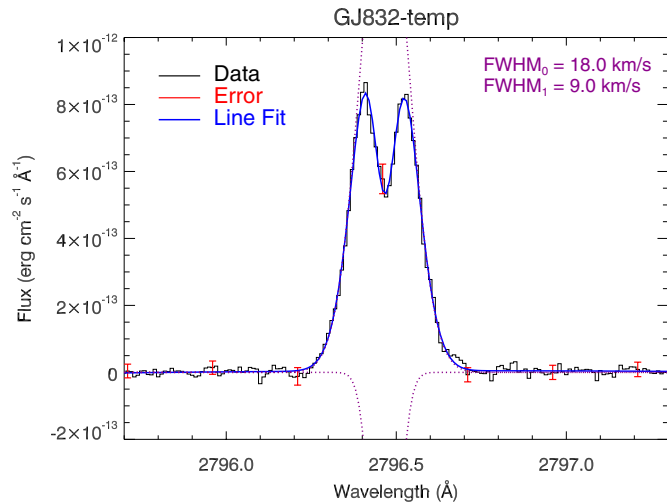


Figure 5. STIS E230H profile of the Mg II k line ($\lambda_o = 2795.528 \text{ \AA}$) of GJ 832. The dotted magenta lines are the stellar emission and interstellar absorption components of the fit, respectively. Fitting the high-resolution echelle spectra of the Mg II k and h lines of GJ 832 and GJ 667 C, we estimate that the interstellar correction for the Mg II flux is $\approx 30\%$ – 35% .

(A color version of this figure is available in the online journal.)

study, and the only star where H I Ly α is not detected. In the top panel of Figure 6, we display the full UV spectrum of GJ 1214. The only clearly detected stellar emission lines are those of C IV and Mg II. In the lower panel, we focus on the immediate spectral region around the lines of interest. The upper limit from the co-added STIS G140M spectra of GJ 1214 is $F(\text{Ly}\alpha) < 2.4 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$. The short-wavelength component of the C IV doublet is shown in red, and we show a possible detection of N V in blue. Taking the observed Mg II flux from GJ 1214 ($2.21 (\pm 0.15) \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$), we can use the $F(\text{Ly}\alpha)/F(\text{Mg II})$ relation described above to compare the expected flux from Ly α with the upper limit (red circle and red star in Figure 4). The upper limit on the GJ 1214 Ly α flux is an order of magnitude below the predicted point based on the reconstructed Ly α emission lines and we have no information on which to base a reconstruction in GJ 1214.

The presence of the UV metal line emission demonstrates that GJ 1214 maintains some basal UV flux level. However, the non-detection of Ly α is surprising in light of the results presented above. We suggest three possible explanations that might account for the non-detection. First, the 2012 COS and STIS G230L observations may have occurred during a period of stellar activity, whereas the star may have been UV-quiet during the 2011 STIS G140M observations. This type of “binary” UV activity behavior has been observed previously in the M8 flare star VB 10 (Linsky et al. 1995). Second, if the GJ 1214 Ly α emission line was significantly narrower than the Ly α lines in the other M dwarfs, then it might be possible that the somewhat longer interstellar sightline removed the signal. However, as shown by Wood et al. (2005), neutral interstellar hydrogen columns rarely exceed $\log_{10} N(\text{H I}) = 18.5 \text{ cm}^{-2}$ for stars nearer than 50 pc. Therefore, this second option is really that the stellar Ly α emission line from GJ 1214 is unusually narrow. There is not enough information from C IV or Mg II to make a meaningful comparison of GJ 1214 line widths to the other stars in our sample. The third, more speculative, possibility is that for the low effective temperature of GJ 1214, a significant fraction of the available neutral hydrogen in the stellar atmosphere is locked

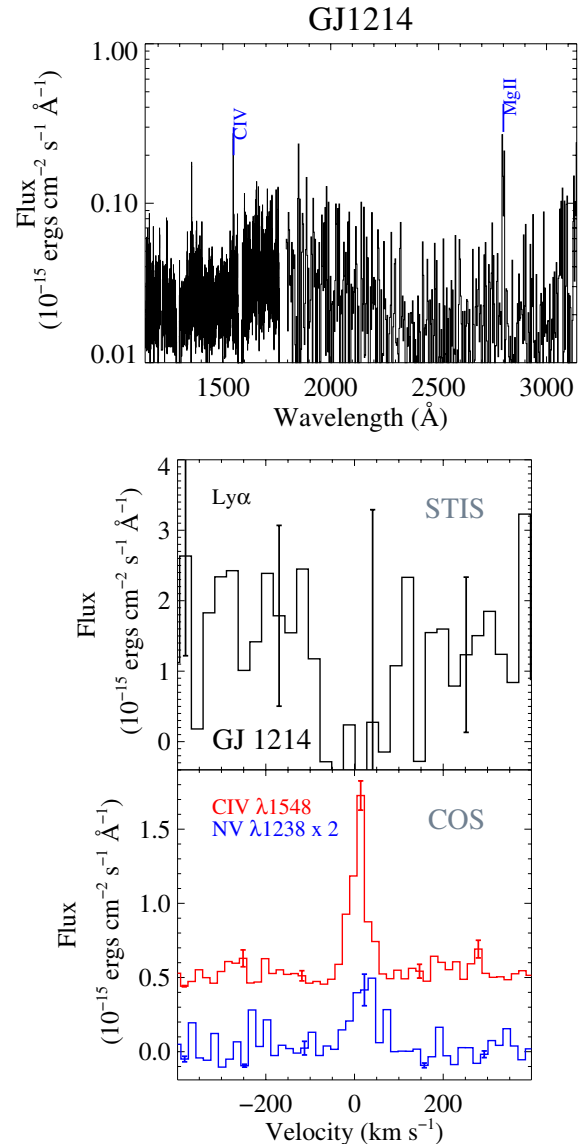


Figure 6. Top: *HST*/COS + STIS 1140–3140 Å spectrum of GJ 1214. Emissions from C IV and Mg II are the only strongly detected features. At bottom, we show the $\pm 400 \text{ km s}^{-1}$ spectrum around H I Ly α (upper panel, black histogram), C IV $\lambda 1548$ (lower panel, red histogram), and N V $\lambda 1239$ (lower panel, blue histogram) scaled up by a factor of two for display. The Mg II and C IV lines demonstrate the presence of an active atmosphere; several possibilities for the non-detection of Ly α are presented in Section 5.1.1.

(A color version of this figure is available in the online journal.)

up in molecules (mainly H₂). As we will discuss in Section 5.3, we detect H₂ emission in all of the other M dwarfs observed with COS. We do not detect H₂ fluorescence in GJ 1214, but this could be the consequence of a feedback process driven by increasing molecular fraction. Ly α is the energy source for the observed fluorescence, so as fewer atoms are available to produce Ly α , the H₂ fluorescence process will become unobservable in FUV observations even if more molecules are present in the stellar atmosphere.

5.2. FUV/NUV Flux Ratios and the Ly α Fraction

In earlier work describing the UV spectrum of GJ 876, it was shown that the FUV/NUV flux ratio is $> 10^3$ times the solar value (France et al. 2012a). This is a consequence of an enhanced Ly α line strength and a much lower NUV flux owing

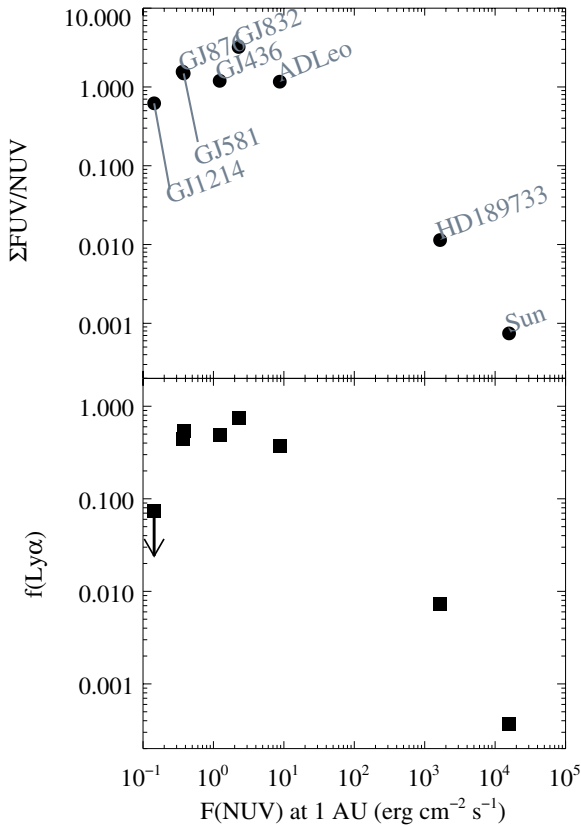


Figure 7. Top: the total FUV (including the reconstructed Ly α emission line) to NUV flux ratio is ~ 0.5 –3 for M dwarfs and $\lesssim 10^{-2}$ for K dwarfs and earlier. This ratio has important implications for the oxygen photochemistry on Earth-like planets orbiting low-mass stars. Bottom: the fraction of the total UV energy in the Ly α line is $\gtrsim 35\%$ for most of the M dwarfs (GJ 581, GJ 876, GJ 436, GJ 832, and AD Leo), $< 8\%$ for GJ 1214, and $\lesssim 1\%$ for earlier-type stars.

to the cooler effective temperature of GJ 876. In Figure 7, we show that this is a generic property of M dwarfs. Here, we define the FUV flux, $F(\text{FUV})$, as the total stellar flux integrated over the 1160–1690 Å bandpass, excluding Ly α . Therefore, the total FUV emission is $\Sigma\text{FUV} = (F(\text{Ly}\alpha) + F(\text{FUV}))$, the total FUV/NUV ratio is

$$\Sigma\text{FUV}/\text{NUV} = (F(\text{Ly}\alpha) + F(\text{FUV}))/F(\text{NUV}), \quad (1)$$

and the Ly α fraction is defined as

$$f(\text{Ly}\alpha) = F(\text{Ly}\alpha)/(F(\text{Ly}\alpha) + F(\text{FUV}) + F(\text{NUV})), \quad (2)$$

where $F(\text{Ly}\alpha)$ is integrated over 1210–1222 Å and $F(\text{NUV})$ is integrated over 2300–3050 Å.

The observed Ly α , FUV, and NUV fluxes are presented in Table 3. We put these values in context in Table 4 by presenting the Ly α , FUV, and NUV luminosities as the fraction of the bolometric luminosities. We calculate the bolometric luminosity for each star by creating a continuous spectrum from 0.115 to 2.5 μm , using the appropriate M dwarf model atmosphere (Pickles 1998) to cover the optical and IR portions of the spectral energy distribution. The model atmosphere flux is scaled to the STIS G230L data from 3050–3080 Å. The use of the Pickles models will underestimate the bolometric luminosity somewhat (by not including the entire Rayleigh–Jeans tail into the mid-IR); however, we expect that this effect is similar to or smaller than the uncertainty introduced by scaling to the NUV data, where the relative contribution of the photosphere is not clear in most cases.

$\Sigma\text{FUV}/\text{NUV}$ is in the range ≈ 1 –3 for all M dwarfs, except GJ 1214. Mostly due to the lack of Ly α emission (Section 5.1.1), $\Sigma\text{FUV}/\text{NUV}$ is in the range 0.50–0.62 for GJ 1214. While the $\Sigma\text{FUV}/\text{NUV}$ ratio is lower for GJ 1214 than for the rest of the M dwarfs in our sample, it is still significantly higher than for earlier type stars. For HD 189733, this ratio has fallen to $\sim 10^{-2}$, and it is $< 10^{-3}$ for the Sun. Similarly, $f(\text{Ly}\alpha)$ spans 37%–75% for most M dwarfs, is $< 7.4\%$ for GJ 1214, is $\approx 0.7\%$ for HD 189733, and is $\approx 0.04\%$ for the Sun. AD Leo has an absolute NUV flux that is 4–40 times larger than for the less active M dwarfs in this survey, and the $\Sigma\text{FUV}/\text{NUV}$ ratio could decrease dramatically during flares when strong Balmer continuum emission dominates the spectrum (Hawley & Pettersen 1991). The implications of the large FUV/NUV flux ratio in M dwarfs are discussed in more detail below (Section 7.1).

5.3. The Ubiquity of H₂ Fluorescence: an Exoplanetary Origin?

We observe fluorescent H₂ emission lines in GJ 581, GJ 876, GJ 436, and GJ 832. These features are photoexcited (“pumped”) by Ly α photons out of the $v = 2$ level of the ground electronic state (Shull 1978). Under the assumption of thermalized rovibrational populations, this process is a signpost for 2000–4000 K molecular gas. We detect between four and eight fluorescent emission lines in each of these targets, pumped by Ly α through the (1–2) $R(6)$ $\lambda 1215.73$ Å and (1–2) $P(5)$ $\lambda 1216.07$ Å transitions. The S/N of any of the individual lines is not sufficiently high for spectral fitting, so we co-added four to six of the brightest lines.⁹ The co-added line profiles were normalized to unity and fitted by assuming a Gaussian emission line shape that has been convolved with the 1500 Å *HST*+COS line spread function (Kriss 2011). The co-added fluorescence lines and fits are shown in Figure 8. The H₂ emission lines are consistent with the stellar radial velocities to within half of the 15 km s^{−1} wavelength solution uncertainty of the COS G160M mode, [$v_{\text{H}_2}(581) = -5.5$ km s^{−1}, $v_{\text{H}_2}(876) = +9.0$ km s^{−1}, $v_{\text{H}_2}(436) = +17.5$ km s^{−1}, $v_{\text{H}_2}(832) = +19.7$ km s^{−1}] compared with [$v_{\text{rad}}(581) = -9.5$ km s^{−1}, $v_{\text{rad}}(876) = +8.7$ km s^{−1}, $v_{\text{rad}}(436) = +10$ km s^{−1}, $v_{\text{rad}}(832) = +18.0$ km s^{−1}]. Except in the case of GJ 581 (the lowest S/N H₂ spectrum), the H₂ emission features are spectrally unresolved. In this subsection, we describe three possible physical origins for the observed fluorescence: (1) the stellar photospheres, (2) the atmospheres of the planets orbiting these stars, or (3) circumstellar gas disks populated by atmospheric mass loss from the planets in these systems.

Ly α -pumped H₂ emission lines were first detected in sounding rocket spectra of sunspots (Jordan et al. 1977), and have been modeled assuming a thermalized H₂ population at a temperature of 3200 K (Shull 1978). H₂ fluorescence lines are also powerful diagnostics of the molecular surface layers of protoplanetary disks over a similar range of temperatures (Herczeg et al. 2004; France et al. 2012b). These lines are not clearly detected in previous STIS observations of M dwarfs *without* planets (AD Leo, Proxima Cen, EV Lac, and AU Mic¹⁰), although this could be related to much larger instrumental backgrounds

⁹ The co-added H₂ line profiles include at least four of the following emission lines: (1–7) $R(6)$ $\lambda 1500.24$ Å, (1–7) $P(8)$ $\lambda 1524.65$ Å, (1–6) $R(3)$ $\lambda 1431.01$ Å, (1–6) $P(5)$ $\lambda 1446.12$ Å, (1–7) $R(3)$ $\lambda 1489.57$ Å, and (1–7) $P(5)$ $\lambda 1504.76$ Å.

¹⁰ AU Mic displays H₂ fluorescence from lower temperature molecular gas, but this gas is consistent with an origin in a remnant protoplanetary gas disk (France et al. 2007).

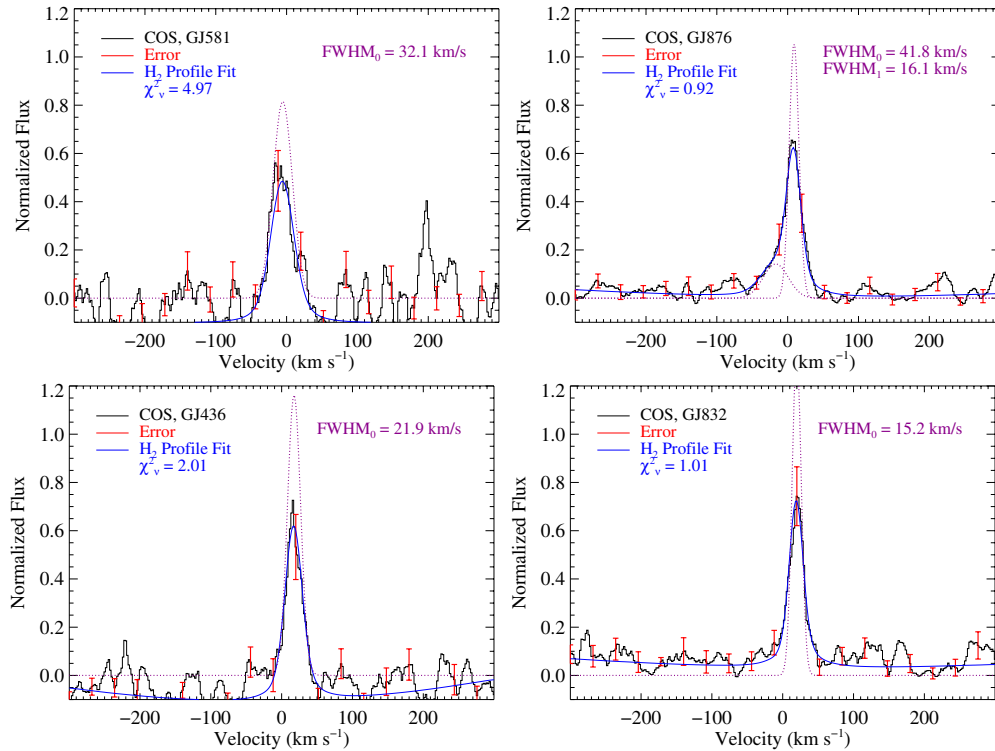


Figure 8. Co-added H_2 emission spectra detected in the MUSCLES pilot program. Each velocity profile is the co-addition of four to six fluorescent emission lines. The dotted line is the intrinsic Gaussian line profile which has been convolved with the COS line spread function (solid blue line) for comparison with the data (black histogram). The red error bars are representative of the uncertainty on the normalized line profiles. The origin of these lines is not yet clear; more detailed modeling is required to differentiate between a photospheric and exoplanetary origin for this hot, photoexcited molecular gas.

(A color version of this figure is available in the online journal.)

associated with the STIS echelle modes. In any event, the 3200 K sunspot temperature is characteristic of the photospheric effective temperature of the MUSCLES stars, and simple energy budget calculations (e.g., France et al. 2012a) suggest that the fluorescence may be consistent with an origin in a hot molecular stellar surface layer.

However, it is interesting that these lines are either not detected or are marginally detected with much lower H_2 /atomic line ratios in the non-planet-hosting stars. H_2 is the primary constituent of gas giant planets. After $\text{H I Ly}\alpha$, the electronically excited emission spectrum of H_2 dominates the FUV spectrum of giant planets (Broadfoot et al. 1979; Wolven & Feldman 1998; Gustin et al. 2010). Fluorescence driven by solar $\text{Ly}\beta$ is the main constituent of the dayglow (non-auroral illuminated disk) spectrum of Jupiter (Feldman et al. 1993) and in areas where the Jovian surface temperature is elevated (e.g., the Shoemaker–Levy 9 impact site; Wolven et al. 1997), $\text{Ly}\alpha$ -pumped fluorescence becomes strong.

All four of the stars showing strong H_2 fluorescence host planets of Neptune to super-Jupiter mass. Our analysis cannot rule out the possibility that the ubiquitous $\text{Ly}\alpha$ -pumped H_2 fluorescence observed toward M dwarf planet hosts is the result of photoexcited dayglow emission in planetary atmospheres. In light of the recent tentative detections of FUV H_2 emission from the planetary systems around HD 209458 (France et al. 2010b) and brown dwarf 2MASS J12073346–3932539 (France et al. 2010a), it is tempting to attribute an exoplanetary origin to these lines. We caution that further study is required to differentiate between these two scenarios. If the H_2 in the GJ 436 system originates in the atmosphere of GJ 436b, we would have expected to see a significant velocity shift of the

fluorescent lines relative to the stellar radial velocity ($v_{436b} = 110\text{--}120\text{ km s}^{-1}$ at orbital phase ≈ 0.8). However, the H_2 lines are consistent with the radial velocity of the star, indicating little direct dayglow and/or auroral contribution from GJ 436b. The upper limit on the H_2 (1–7) R (3) emission line flux at $+110\text{ km s}^{-1}$ (assuming unresolved emission lines) is $\approx 6 \times 10^{-17}\text{ erg cm}^{-2}\text{ s}^{-1}$ in the GJ 436 spectrum. Modeling of both the H_2 fluorescence spectrum from M dwarf photospheres and the expected UV emission spectrum from nearby M dwarf planetary systems would be extremely valuable.

A final possibility is that the observed H_2 fluorescence originates in a circumstellar gas envelope that is being replenished by atmospheric mass loss from escaping planetary atmospheres. The narrow H_2 line width in GJ 436 rules out the possibility that the fluorescing gas is confined to the semimajor axis of the planetary orbit ($a = 0.03\text{ AU}$ for GJ 436b); however, over the $\sim 5\text{ Gyr}$ lifetimes of these systems, this material could disperse into a low-density “second generation” circumstellar disk. While detailed modeling of such a disk is beyond the scope of this work, we can make a rough estimate of the planetary mass-loss rate required to sustain the observed H_2 emission.

The balance equation of H_2 photodissociation and replenishment is $\beta_{\text{diss}}n(\text{H}_2) = \Phi_{\text{loss}}$. Due to the negligible 912–1120 Å continuum flux from inactive M-type stars, we assume that the interstellar radiation field (ISRF) dominates H_2 dissociation. The photodissociation rate is $\beta_{\text{diss}} = B_{\text{diss}}\beta_o$. B_{diss} is the dissociation fraction of molecules excited into the Lyman and Werner bands following absorption of a 912–1300 Å photon, ≈ 0.2 for a hot H_2 gas. β_o is the standard H_2 photoabsorption rate from the ISRF, $\beta_o = 5 \times 10^{-10}\text{ s}^{-1}$ (Jura 1975). We assume an H_2 column density of $N(\text{H}_2) \approx 10^{17}\text{ cm}^{-2}$ (France et al.

2012a) and a fiducial cloud diameter of $d_{\text{H}_2} = 1$ AU such that $n(\text{H}_2) = N(\text{H}_2)/d_{\text{H}_2} \approx 7 \times 10^3 \text{ cm}^{-3}$. The replenishment term from the escaping atmospheres of our short-period planets is $\Phi_{\text{loss}} = \dot{M}_{\text{loss}} m_{\text{H}_2}^{-1} F(\text{H}_2/\text{H}_1) V_{\text{emit}}^{-1}$, where \dot{M}_{loss} is the planetary atmosphere mass-loss rate, $F(\text{H}_2/\text{H}_1)$ is the fraction of the total cloud in molecular form, and V_{emit} is the volume of the emitting area ($V_{\text{emit}} = \pi r_{\text{H}_2}^2 z_{\text{H}_2}$). Models of escaping hot Jupiter atmospheres find that most of the escaping material is atomic (e.g., Tremblin & Chiang 2013), so we assume a low molecular fraction $F(\text{H}_2/\text{H}_1) = 10^{-3}$. Any disk populated by an orbiting planet will most likely be geometrically thin, and we take $z_{\text{H}_2} = 0.01 d_{\text{H}_2}$ as the vertical height. Keeping the crudeness of this approach in mind, we find that the H_2 disk can be sustained for mass-loss rates $\dot{M}_{\text{loss}} \gtrsim 6 \times 10^{10} \text{ g s}^{-1}$, well within the range of mass-loss estimates of short-period planets around G-, K-, and M-type stars (Vidal-Madjar et al. 2008; Linsky et al. 2010; Lecavelier des Etangs et al. 2012; Ehrenreich et al. 2011; Ehrenreich & Désert 2011). Therefore, the H_2 fluorescence observed in our MUSCLES sample appears compatible with either a photospheric or circumstellar origin.

6. TIME VARIABILITY OF M DWARF UV RADIATION FIELDS

We show examples of the UV light curves of GJ 581, GJ 876, GJ 436, and GJ 832 in Figure 9. The gaps in the light curves shown in Figure 9 are due to Earth occultations during the orbit of *HST*. We observe time variability in the UV spectra of all of our M dwarf exoplanet host stars observed with COS (GJ 1214 has insufficient S/N for timing analysis), even on weakly active stars (Walkowicz & Hawley 2009). In Figure 9, we show example light curves from all four targets in chromospheric and transition region tracers Si III $\lambda 1206$, N V $\lambda \lambda 1239, 1243$, C II $\lambda \lambda 1334, 1335$, Si IV $\lambda \lambda 1394, 1403$, and C IV $\lambda \lambda 1548, 1550$, and the detector background, evaluated in $\Delta t = 40$ s or 120 s time steps. Our observations were designed to provide a first estimate of the flux levels and temporal variability in low-mass exoplanet host stars, so the observations do not span a sufficient baseline to determine the UV flare frequency. However, we can qualitatively identify three flare behaviors in our spectra.

1. *The impulsive flare.* The GJ 876 and GJ 436 UV flares exhibit a qualitatively similar behavior to optical light flares on M dwarfs (e.g., Figure 5 of Kowalski et al. 2009). We observed a flare peak/quiescent flux ratio of ≥ 10 on GJ 876, with an approximately exponential decay ($\tau_{876} \approx 400$ s). The UV flare behavior on GJ 876 is similar to the factor of ~ 5 – 20 flux increase observed on AD Leo (Hawley et al. 2003). GJ 876 also displays X-ray flare activity (Poppenhaeger et al. 2010). A similar but lower amplitude C IV flare was observed on GJ 436. The C IV flux approximately doubled in 240 s, then decayed to pre-flare levels on a timescale $\tau_{436} \approx 1200$ s.
2. *The slow rise flare.* The Si IV light curve of GJ 832, binned to $\Delta t = 120$ s time steps, is shown in the lower left-hand panel of Figure 9. Compared with the abrupt flux increase seen in GJ 876 and GJ 436, the GJ 832 Si IV flux increases by a factor of 1.5–2.0 over the course of the first 1200 s exposure. The peak and turnover are lost to Earth occultation between COS exposures for GJ 832, but the flare is observed to slowly decay during the last 960 s of G130M observation.
3. *The short flare.* GJ 581 displays a shorter duration flare event, shown in the lower right panel in Figure 9. The

C IV light curve shows a flux increase of a factor of ~ 2 – 3 at $T_{\text{exp}} \approx 750$ s. The signal to noise in the GJ 581 UV emission lines is low, but one can resolve the flare using $\Delta t = 40$ s time steps. This 200 s long C IV emission enhancement comes during the first *HST* orbit, when the basal transition region flux was approximately a factor of two larger than during the second *HST* orbit following Earth occultation.

The conclusion from this section is that even with a small sample of relatively low S/N spectra obtained over limited temporal baselines, we have identified variability with a range of amplitudes and durations. Continuous *HST* observing campaigns, spanning 6–10 orbits each, of a larger sample of exoplanet host stars will be essential to better constraining the frequency, duration, and amplitude of the variable UV radiation environment of planets around M dwarfs. Alternatively, long-duration NUV observations could be made with balloon-borne experiments. Time-domain studies at optical wavelengths are also of interest for characterizing the flare frequency of the largest flares that are panchromatic events (e.g., Hawley et al. 2003). This is particularly relevant in light of the projected paucity of UV-capable observatories. A more complete understanding of UV variability on weakly active and inactive M dwarfs, in combination with better constrained flare frequencies for stars of a variety of activity strengths, would allow one to use the wealth of data from current and upcoming large optical surveys to inform planetary atmosphere models.

7. DISCUSSION

7.1. UV Radiation Fields for Atmospheric Photochemistry Models

Owing to the scarcity of M dwarf exoplanet host spectra previously available in the literature, authors modeling the atmospheres of Earth-like planets around M dwarfs have typically been driven to one of two extremes: assuming no UV flux (Segura et al. 2005; Rauer et al. 2011; Kaltenegger et al. 2011) or assuming the spectrum of one of the most active M dwarfs known, AD Leo (Segura et al. 2005; Wordsworth et al. 2010). In the case of no UV flux, the short-wavelength spectral cutoff is determined by the photosphere of a theoretical star with $T_{\text{eff}} \lesssim 3500$ K, typically having negligible flux below 2500 Å and no chromospheric emission features. Walkowicz et al. (2008) demonstrated that even M dwarfs with modest X-ray emission and no chromospheric H α emission display NUV spectra that are qualitatively similar to more optically active stars. Following this work, Kaltenegger et al. (2011) noted that an active UV atmosphere is most likely present on almost all M dwarfs, which is consistent with the observations presented here. Additional observations are required to understand the UV flux from M dwarfs as a function of stellar age and metallicity, for spectral types later than M6, and for a larger sample of M dwarfs without X-ray emission or H α activity.

The MUSCLES observations show that the UV spectral behavior of four of our stars (GJ 581, GJ 876, GJ 436, and GJ 832) is similar to AD Leo in both the spectral and the temporal sense. The MUSCLES stars (except GJ 1214) showed similar a $\Sigma\text{FUV}/\text{NUV}$ flux ratio and $f(\text{Ly}\alpha)$ to that of AD Leo, and they also displayed time variability with UV emission line amplitudes changing by factors of $\gtrsim 2$ during our brief observing snapshots.

While AD Leo appears to be a more appropriate choice for exoplanetary modeling input, a primary difference is the absolute flux of the NUV radiation field from AD Leo. AD Leo has

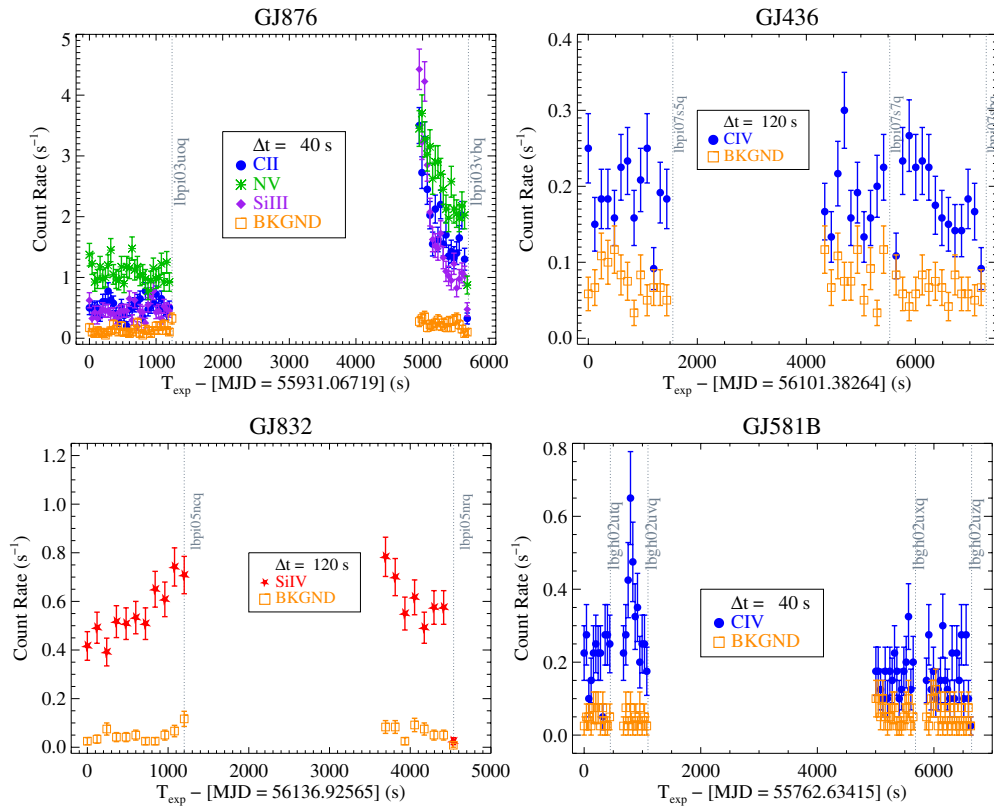


Figure 9. We have detected a range of temporal variability scales in the chromospheric and transition region emission lines from weakly active M dwarf exoplanet host stars. We separate the time variability into three qualitative groups: (1) the impulsive flare, shown in several chromospheric and transition region tracers (C II, Si III, C IV, and N V) in GJ 876 and GJ 436 (top row). (2) The slow rise flare, shown in Si IV in GJ 832 (lower left). (3) The short flare, shown in C IV in GJ 581. This demonstrates that time-variable, active atmospheres are a generic property of M dwarf exoplanet host stars. The gray vertical dotted lines denote the end of an individual *HST*/COS exposure. Exposure IDs are included.

(A color version of this figure is available in the online journal.)

factors of ~ 4 – 40 stronger absolute flux levels at its effective HZ distance than do the rest of the MUSCLES targets. Furthermore, the strong Balmer continuum associated with its frequent flare activity will periodically drive the $\Sigma\text{FUV}/\text{NUV}$ ratio significantly below the level (~ 1 – 3) found for the less active M dwarfs. This could lead to significant differences between the atmospheric chemistry of a planet orbiting an AD Leo-like star and one orbiting a more typical M dwarf. It has been recently shown that the strong FUV flux and high $\Sigma\text{FUV}/\text{NUV}$ flux ratio incident on a terrestrial planet orbiting in the HZ of an M dwarf can lead to the buildup of a large, abiotic atmospheric O_2 abundance through the photodissociation of CO_2 (Tian et al. 2012). The creation of O_3 via O_2 photolysis depends sensitively on the strength of the 1700–2400 Å radiation field and the balance of the ocean–atmosphere system on these planets. In some models, this process can lead to detectable levels of O_3 without the presence of an active biosphere (Domagal-Goldman et al. 2012).

Spectral observations of O_2 and O_3 , along with CH_4 and CO_2 , are expected to be the most important signatures of biological activity on planets with Earth-like atmospheres (Des Marais et al. 2002; Seager et al. 2009), which may be the case for such planets orbiting solar-type stars. However, the above work has shown that when realistic inputs for the UV radiation field are included, it may be possible to produce O_2 and O_3 with observable abundances on terrestrial planets in the HZ of M dwarfs *without the presence of biological life*. Therefore, a detailed knowledge of the stellar spectrum will be critical to interpreting the observations of potential biomarkers on these

worlds when they are detected in the coming decades. Until such time that stellar models of M-type stars can reliably predict the observed UV-through-IR spectrum, a direct UV observation will be the most reliable means for determining the radiation field incident on planets orbiting M dwarfs.

7.2. UV Transit Spectroscopy of M dwarf Planets

In Section 6, we described the range of time variability observed in the UV chromospheric and transition region emission lines from our M dwarf exoplanet host stars. We observe emission line flux variability of 50%–500% on timescales shorter than the typical transit duration of short-period exoplanets (few hours). For comparison, the typical depth of a resolved UV transit observation is $\sim 5\%$ – 15% (Vidal-Madjar et al. 2003; Ben-Jaffel 2007; Linsky et al. 2010; Lecavelier des Etangs et al. 2012). The importance of stellar variability for UV transit studies of G and K dwarfs is just now being explored (Haswell et al. 2012). While an older solar-type transiting planet host star such as HD 209458 has been shown to have very steady transition region emission (observed in Si IV; Linsky et al. 2010), UV variability on younger, less-massive transiting planet hosts like HD 189733 may significantly hinder the ability to detect transits and characterize their depth and ingress/egress profiles. GJ 436 and GJ 1214 both host transiting planets (Neptune and super-Earth mass planets, respectively), and our results suggest that transit studies using UV chromospheric emission lines as background sources to search for signs of atmospheric escape may be complicated by stellar activity. NUV searches for O_3 in the

atmospheres of Earth-like planets will also likely be impacted by M dwarf stellar activity, even for weakly active stars. Long-baseline observations over multiple transits will be required to mitigate the effects of stellar activity in future M dwarf exoplanet transit observations. We will consider this issue in more detail in a future work.

7.3. The MUSCLES Spectral Database

As part of the MUSCLES pilot program, we have started building a library of medium/high-resolution UV spectra of M dwarf exoplanet host stars. These data are available to download for use in photochemical models of Earth-like, ice giant, and gas giant planets.¹¹

The high resolution of the MUSCLES sample is important for detailed molecular emission/absorption modeling. As most of the UV flux from low-mass stars arises in discrete emission lines, molecular species in the atmospheres of HZ planets will be “selectively pumped”: only species that have spectral coincidences with stellar emission lines will be subject to large energy input from the host star. H₂O, CH₄, and CO₂ in particular are highly sensitive to far-UV radiation, and provide an illustration of the utility of spectrally resolved data. CH₄ absorbs in a broad band at $\lambda < 1500$ Å, and the photoabsorption rates are relatively insensitive to the spectral resolution. CO₂ displays a banded structure over a broad absorption region from ~ 1200 – 1600 Å, and high-resolution data show a spectral coincidence between the narrow CO₂ bands and the Si III $\lambda 1298$ + O I $\lambda 1304$ multiplet and the C II $\lambda 1335$ stellar emission lines. H₂O dissociation is influenced by far-UV photons as the molecules absorb in a broad continuum from 1400 – 1900 Å, $B - X$ absorption bands near 1280 Å, and sharp Rydberg transitions at $\lambda < 1240$ Å that overlap with the Ly α emission line (Yi et al. 2007). The spectral overlap between the photo-absorption cross-sections of CO₂ and H₂O and the chromospheric and transition region emission lines from an M dwarf are shown in Figure 10. Low-resolution spectra or molecular absorption cross-sections will fail to preserve the specific energy absorption in these narrow ($\Delta\lambda < 10$ Å) wavelength coincidences and could lead to an underestimation of the dissociation rates of important atmospheric species.

8. SUMMARY

We have presented the first study of the UV radiation fields around nearby M dwarf planet hosts that covers both FUV and NUV wavelengths. Using a combination of data from the COS and STIS instruments on *HST*, we present 1150 – 3100 Å spectra for GJ 581, GJ 876, GJ 436, GJ 832, GJ 1214, and Ly α and Mg II observations for GJ 667C. Using an iterative least-squares approach, we reconstruct the intrinsic Ly α radiation field strengths assuming a two-component Gaussian emission line. Using the time-tagged capability of the COS FUV detector, we calculated emission line light curves in several chromospheric and transition region diagnostics. The main results of this work are as follows.

1. All six of the stars in our sample display both FUV and NUV emission line flux, even though these stars do not display H α emission (however, they do show some degree of H α absorption and Ca II emission lines). This suggests

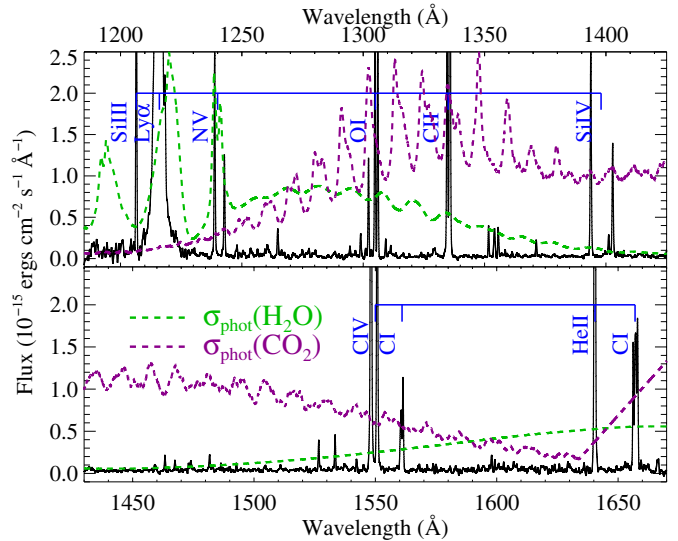


Figure 10. Photoabsorption cross-sections of H₂O (green dashed line; Mota et al. 2005) and CO₂ (purple dashed line; Yoshino et al. 1996) showing the spectral overlap with chromospheric and transition region emission lines from an M dwarf (black histogram). Bright chromospheric and transition region emission lines are labeled.

(A color version of this figure is available in the online journal.)

that an active stellar UV atmosphere should be included in calculations of atmospheric chemistry on planets orbiting most M dwarfs.

2. The Ly α emission lines are well-fit by a two-component Gaussian emission line model with narrow ($\text{FWHM}_{\text{narrow}} \sim 90$ – 130 km s⁻¹) and broad ($\text{FWHM}_{\text{broad}} \sim 160$ – 300 km s⁻¹) components. The $F_{\text{narrow}}/F_{\text{broad}}$ flux ratio ranges from ~ 5 – 12 . Interstellar H I column densities are derived; $18.06 \leq \log_{10} N(\text{H I}) \leq 18.47$ cm⁻² for our targets ($d = 4.7$ – 10.3 pc).
3. Ly α emission was not detected toward the M6 super-Earth host star GJ 1214, although emission from Mg II and C IV indicate that an active atmosphere is present, at least during certain times.
4. The reconstructed Ly α line flux correlates with the observed Mg II line flux. The ISM-corrected intrinsic $F(\text{Ly}\alpha)/F(\text{Mg II})$ ratio for weakly active M dwarf exoplanet host stars is 10 ± 3 .
5. The total FUV/NUV flux ratio is in the range ≈ 0.5 – 3 for all M dwarfs, including AD Leo. This ratio is $\lesssim 10^{-2}$ for K1V dwarfs and $< 10^{-3}$ for the Sun. The percentage of total UV energy from the star in the Ly α emission line is: (1) 37%–75% for GJ 581, GJ 876, GJ 436, GJ 832, and AD Leo, (2) $< 7.4\%$ for GJ 1214, (3) $\approx 0.7\%$ for K1V dwarfs, and (4) $\approx 0.04\%$ for the Sun. The FUV/NUV flux ratio can have a profound effect on the production of O₂ and O₃ in the atmospheres of Earth-like planets orbiting M dwarfs.
6. Fluorescent emission from hot H₂ ($T(\text{H}_2) \approx 2000$ – 4000 K) is detected toward all four stars with moderate S/N COS spectra (GJ 581, GJ 876, GJ 436, and GJ 832). A first-order calculation suggests that this emission originates in the stellar photosphere, but this emission could also arise on the illuminated dayside of the planets in these systems or in a circumstellar envelope replenished by mass loss from the planetary atmosphere.
7. Time variability is detected in the chromospheric and transition region emission lines observed from all four

¹¹ The combined 1150 – 3100 Å spectra (0.05 Å pixel⁻¹ binning) and the full-resolution, co-added COS spectra of the MUSCLES targets are available for download at <http://cos.colorado.edu/~kevinf/muscles.html>.

stars with moderate S/N in the COS spectra (GJ 581, GJ 876, GJ 436, and GJ 832). Variability with amplitudes 50%–500% is observed on timescales of 10^2 – 10^3 s. This will most likely make UV transit studies of these systems challenging.

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