

HAT-P-42b and HAT-P-43b

Two inflated transiting hot Jupiters from the HATNet Survey^{*,**}

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ABSTRACT

Aims. We announce the discovery of two new transiting planets, and provide their accurate initial characterization.

Methods. First identified from the HATNet wide-field photometric survey, these candidate transiting planets were then followed-up with a variety of photometric observations. Determining the planetary nature of the objects and characterizing the parameters of the systems were mainly done with the SOPHIE spectrograph at the 1.93 m telescope at OHP and the TRES spectrograph at the 1.5 m telescope at FLWO.

Results. HAT-P-42b and HAT-P-43b are typical hot Jupiters on circular orbits around early-G/late-F main sequence host stars, with periods of 4.641878 ± 0.000032 and 3.332687 ± 0.000015 days, masses of 1.044 ± 0.083 and $0.662 \pm 0.060 M_J$, and radii of 1.280 ± 0.153 and $1.281^{+0.062}_{-0.033} R_J$, respectively. These discoveries increase the sample of planets with measured mean densities, which are needed to constrain theories of planetary interiors and atmospheres. Moreover, their hosts are relatively bright ($V < 13.5$), which facilitates further follow-up studies.

Key words. planetary systems – techniques: radial velocities – techniques: photometric – stars: individual: HAT-P-42 – stars: individual: HAT-P-43

1. Introduction

Currently, more than 800 confirmed exoplanets are known, including about 200 well-characterized transiting planets (i.e. with planetary masses and radii measured to better than 10% accuracy). Transiting extrasolar planets are crucial for exploring the physical properties of exoplanets, such as their mass-radius relationships, atmospheric composition or orbital obliquity. They provide an important first step toward a valuable comparison of extrasolar planets. These objects will be the targets of upcoming space missions designed for atmospheric characterization.

Radial velocity (RV) surveys with well defined sample limits estimated an occurrence rate of hot Jupiters in the Solar neighborhood to be around one percent (Udry & Santos 2007; Wright et al. 2012). With a transit probability around ~10%, a huge number of stars have to be monitored to detect such rare objects. This is why the majority of the transiting planets were discovered by photometric surveys targeting tens of

* Full Table 2 is only available in electronic form at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/558/A86>

** The photometric/spectroscopic data presented in this paper are based in part on observations carried out by the Hungarian-made Automated Telescope Network, using telescopes operated at the Fred Lawrence Whipple Observatory (FLWO) of the Smithsonian Astrophysical Observatory (SAO), and at the Submillimeter Array (SMA) of SAO, by the Tillinghast Reflector 1.5 m telescope and the 1.2 m telescope, both operated by SAO at FLWO, by the SOPHIE spectrograph mounted on the 1.93 m telescope at Observatoire de Haute Provence, France (runs DDT-Dec. 2011), by the Nordic Optical Telescope, operated on the island of La Palma jointly by Denmark, Finland, Iceland, Norway, and Sweden, in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias, and by the facilities of the Las Cumbres Observatory Global Telescope.

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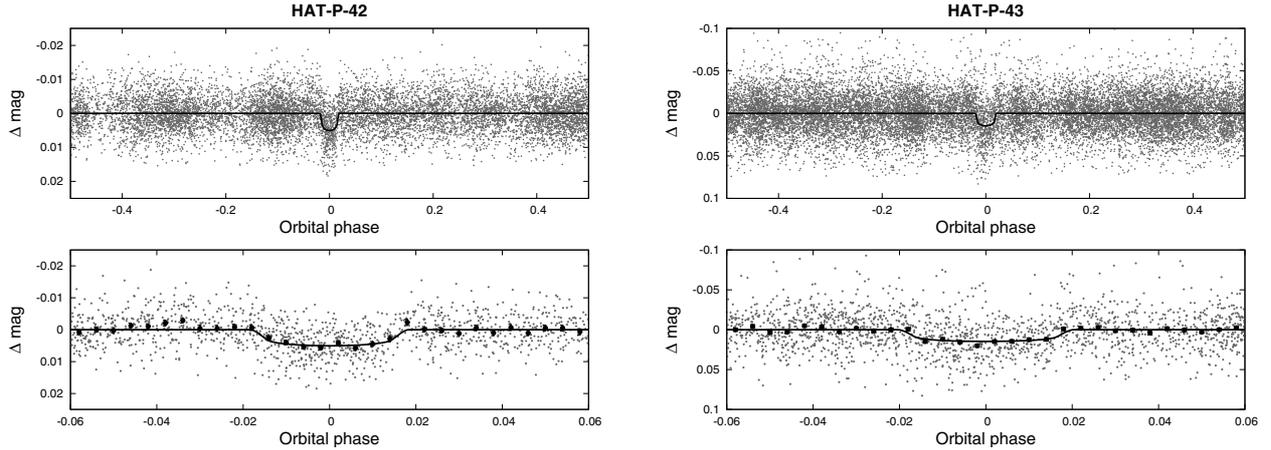


Fig. 1. HATNet light curves of HAT-P-42 (*left*) and HAT-P-43 (*right*) phase folded with the transit period. In both cases we show two panels: the *top* shows the unbinned light curve, while the *bottom* shows the region zoomed-in on the transit, with dark filled circles for the light curve binned in phase with a binsize of 0.002. The solid line shows the model fit to the light curve.

Table 1. Summary of photometric observations.

Instrument/Field ^a	Date(s)	Number of images	Approx. cadence (min)	Filter
HAT-P-42				
HAT-6/G366	2010 Nov.–2011 Apr.	5325	3.7	Sloan <i>r</i>
HAT-9/G366	2010 Nov.–2011 Apr.	4859	3.7	Sloan <i>r</i>
Keplercam	2011 Nov. 16	176	1.2	Sloan <i>i</i>
Keplercam	2012 Jan. 6	310	1.0	Sloan <i>i</i>
Keplercam	2012 Feb. 3	93	1.2	Sloan <i>i</i>
Keplercam	2012 Mar. 2	278	1.1	Sloan <i>i</i>
HAT-P-43				
HAT-5/G317	2010 Nov.–2011 Apr.	4174	3.7	Sloan <i>r</i>
HAT-8/G317	2010 Nov.–2011 Apr.	4300	3.7	Sloan <i>r</i>
HAT-7/G365	2010 Nov.–2011 May	8262	3.7	Sloan <i>r</i>
HAT-8/G365	2011 Apr.–2011 May	464	3.7	Sloan <i>r</i>
Keplercam	2011 Nov. 17	68	3.2	Sloan <i>i</i>
Keplercam	2012 Jan. 13	87	2.2	Sloan <i>i</i>
FTS	2012-Feb.-16	226	1.0	Sloan <i>i</i>
Keplercam	2012 Mar. 3	141	2.2	Sloan <i>i</i>
BOS	2012 Mar. 3	109	2.4	Sloan <i>r</i>
Keplercam	2012 Mar. 23	125	2.2	Sloan <i>i</i>

Notes. ^(a) HAT-5, -6, -7, and -10 are located at FLWO in AZ, USA. HAT-8 and -9 are located at Mauna Kea Observatory in HI, USA.

thousands of stars per night. Ground-based surveys are biased toward the detection of short period and relatively large planets, i.e. hot Jupiters (with some exceptions, e.g. HAT-P-11, Bakos et al. 2010). Complementary, space-based surveys (i.e. CoRoT and Kepler) have excelled in the detection of longer periods and smaller photometric transit depths, and hence smaller planet sizes. But, because the target stars are often relatively faint, radial-velocity follow-up to measure the planetary mass and confirm the planetary nature of the object is often not feasible for the space-based discoveries.

We present here two new transiting hot Jupiters, first identified from the Hungarian-made Automated Telescope Network wide-field photometric ground-based survey (HATNet; Bakos et al. 2004). Since 2006, HATNet has announced and published 41 transiting exoplanets, i.e. ~20% of the well-characterized sample. These two new typical hot Jupiters, HAT-P-42b and HAT-P-43b, were partially confirmed thanks to spectroscopic observations made with SOPHIE mounted at the 1.93 m telescope at the Observatoire de Haute-Provence (OHP) in France and TRES at the 1.5 m telescope at FLWO.

The layout of the paper is as follows. Section 2 presents the diverse photometric and spectroscopic observations that lead to the detection and characterization of these systems. In the third section, we derive the stellar parameters and the planetary orbits. The characteristics of these two hot Jupiters are discussed in the last section.

2. Observations

2.1. Photometric detection

HAT-P-42 and HAT-P-43 were identified as candidate transiting planets based on photometric observations conducted by the HATNet survey (Bakos et al. 2004). These observations are summarized, together with the follow-up photometric observations, in Table 1. The data were processed and searched for transits following the procedure of Bakos et al. (2010), see also Kovács et al. (2002, 2005). Figure 1 shows the folded HATNet light curves for both systems. We give cross-identifications, and catalog photometry on an absolute scale for each system in Table 5.

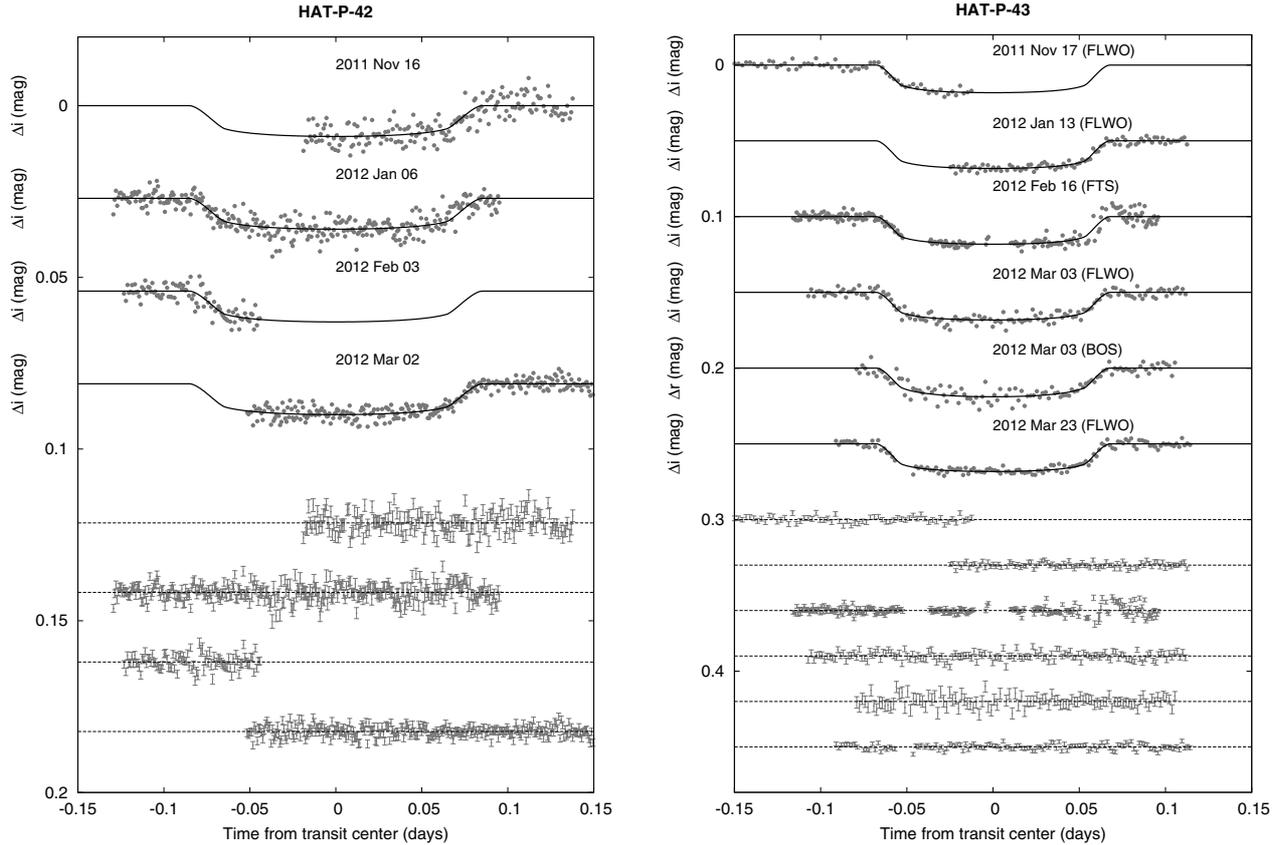


Fig. 2. Follow-up light curves for HAT-P-42 (*left*) and HAT-P-43 (*right*). For HAT-P-42 all observations were obtained with Keplercam on the FLWO 1.2 m telescope. For HAT-P-43 observations were obtained with Keplercam (indicated by FLWO in the figure), the Spectral CCD on the FTS 2.0 m (FTS), and CCD imager on the BOS 0.8 m (BOS). The Keplercam light curves have been corrected for trends during the modeling. The BOS and FTS light curves have also been corrected for trends, though for these light curves we only apply the external parameter decorrelation (EPD) procedure, and not the Trend Filtering Algorithm (TFA; see Bakos et al. 2010, for details). The dates of the events are indicated. Curves below the top one are displaced vertically for clarity. Our best fit from the global modeling is shown by the solid lines. Residuals from the fits are displayed at the bottom, in the same order as the top curves. The error bars represent the photon and background shot noise, plus the readout noise. They are only plotted on the residuals for readability.

2.2. Follow-up

We conducted follow-up spectroscopic and photometric observations of both systems to confirm their planetary natures and determine their physical properties. See Latham et al. (2009) for a more detailed description of our follow-up procedure. Below we provide specific details relevant to the discoveries of HAT-P-42b and HAT-P-43b.

2.2.1. Reconnaissance spectroscopy

We obtained initial high-resolution, low signal-to-noise ratio (S/N) “reconnaissance” spectra of HAT-P-42 and HAT-P-43 using the Tillinghast Reflector Echelle Spectrograph (TRES; Fűrész 2008) on the 1.5 m Tillinghast Reflector at the Fred Lawrence Whipple Observatory (FLWO) in AZ, USA, and using the Fibre-fed Echelle Spectrograph (FIES) on the 2.5 m Nordic Optical Telescope (NOT) on the island of La Palma, Spain (Djupvik & Andersen 2010). These observations are summarized in Table 3. The data were analyzed via cross-correlation against synthetic templates as described by (Buchhave et al. 2010; Quinn et al. 2012). Both stars were found to be slowly rotating G dwarfs, with no evidence for composite spectra, and with low velocity variations indicating that neither object is an undiluted eclipsing binary star system.

2.2.2. Photometry

Photometric follow-up observations were obtained with Keplercam on the FLWO 1.2 m, the Spectral CCD on the 2.0 m Faulkes Telescope South (FTS) at Siding Springs Observatory in Australia, and with the CCD imager on the 0.8 m Byrne Observatory at Sedgwick (BOS) telescope at Sedgwick Reserve in the Santa Ynez Valley, CA, USA. These observations are summarized in Table 1. Both FTS and BOS are operated by the Las Cumbres Observatory Global Telescope (LCOGT¹; Shporer et al. 2011; Brown et al., in prep.). The Keplercam observations were reduced to light curves following the procedure of Bakos et al. (2010), while the FTS and BOS observations were reduced following the procedure of Fulton et al. (2011).

Figure 2 shows the photometric follow-up light curves for both systems together with our best-fit models after detrending as described in Sect. 3. HAT-P-42 and HAT-P-43 have photometric transit depths of 7.6 and 16.5 mmag, respectively. The data are provided in machine-readable format in Table 2.

2.2.3. Confirmation spectroscopy

We obtained high-resolution, high-S/N spectra of HAT-P-42 and HAT-P-43 using the SOPHIE spectrograph on the 1.93 m

¹ <http://lcogt.net>

Table 2. High-precision differential photometry of HAT-P-42 and of HAT-P-43.

Object	BJD ^a –2 455 000	Mag ^b	σ_{Mag}	Mag _{orig} ^c	Filter
HAT-P-42	882.87909	0.01318	0.00170	11.46340	<i>i</i>
	882.88049	0.01025	0.00191	11.46150	<i>i</i>
	882.88128	0.01064	0.00192	11.45990	<i>i</i>
	882.88208	0.00410	0.00192	11.45190	<i>i</i>
....
HAT-P-43	883.89634	0.00397	0.00167	12.40490	<i>i</i>
	883.89856	0.00002	0.00157	12.40130	<i>i</i>
	883.90082	–0.00162	0.00153	12.39990	<i>i</i>
	883.90306	0.00397	0.00150	12.40600	<i>i</i>
....

Notes. ^(a) Barycentric Julian Date calculated directly from UTC, *without* correction for leap seconds. ^(b) The out-of-transit level has been subtracted. These magnitudes have been subjected to the EPD and, in some cases, TFA procedures, carried out simultaneously with the transit fit. The TFA procedure has not been applied to the HAT-P-43 observations with $2\,455\,973 < \text{BJD} < 2\,455\,975$ or to the HAT-P-43 observations in the *r* filter with $2\,455\,990 < \text{BJD} < 2\,455\,991$. ^(c) Raw magnitude values without application of the EPD or TFA procedures. This table is available in a machine-readable at the CDS. A portion is shown here for guidance regarding its form and content.

telescope at OHP (Bouchy et al. 2009), and of HAT-P-42 using TRES on the FLWO 1.5 m telescope. These observations are summarized in Table 3. The TRES observations of HAT-P-42 were simply a continuation of the first two “reconnaissance” TRES observations of this target which were already of sufficient S/N to detect the $\sim 100 \text{ m s}^{-1}$ orbital variation. Of the 13 TRES measurements, 4 with excessively low S/N (i.e. precision $> 100 \text{ m s}^{-1}$) were excluded from the orbital analysis. While we have previously used SOPHIE to confirm HATNet planets (e.g. Bakos et al. 2007; Shporer et al. 2009), there have been significant changes to the instrument and reduction procedure. These significantly increased the RV accuracy (Perruchot et al. 2011; Bouchy et al. 2012), thus we provide a description of our observing and reduction procedures below.

We observed HAT-P-42 and HAT-P-43 with the *High Efficiency* fibers, yielding a resolving power of $R = \lambda/\Delta\lambda \approx 39\,000$ (at 550 nm). The spectra were correlated with a G2 numerical mask in order to calculate the cross-correlation function (CCF). This CCF is fitted by a Gaussian and its center gives the RV measurement (Baranne et al. 1996; Pepe et al. 2002). The bisector span values were calculated as described in Queloz et al. (2001). The error bars were derived from the CCF, following the procedure in Boisse et al. (2010), yielding a mean precision of $\sim 6 \text{ m s}^{-1}$ for HAT-P-42 and $\sim 10 \text{ m s}^{-1}$ for HAT-P-43. Since most of our spectra were polluted by moonlight, we corrected them following the method presented in Hébrard et al. (2008), and increased the error bars as determined empirically, leading to a mean precision of $\sim 40 \text{ m s}^{-1}$ for HAT-P-42 and $\sim 30 \text{ m s}^{-1}$ for HAT-P-43. A thorium-argon calibration was performed every two hours to monitor the instrumental drift mainly due to small variation of temperature and pressure in the spectrograph tank. SOPHIE is environmentally stabilized and this intrinsic drift is less than 2 m s^{-1} per hour. Each measurement was also corrected for this drift.

We computed the CCF using masks corresponding to different spectral types (F0, G2, K5, M5) and found no significant variation in the RV semi-amplitude. If a planet induces a RV shift, the amplitude of the variation should remain constant

regardless of the spectral lines used to measure the RV. In contrast, for the case of a blend consisting of a bright target and a faint binary having different spectral types, the relative contributions to the CCF from the different components will change when a different mask is used (Santos et al. 2002; Collier Cameron et al. 2007). This result thus favors the planetary hypothesis.

Figure 3 shows the high-precision RV curves for both systems, together with our best-fit circular-orbit models. Figure 4 shows the bisector spans as a function of the RV from the SOPHIE spectra. The BIS show no significant variability or trends relative to their error bars. The BIS values are not correlated to the RV variations in agreement with the planetary hypothesis. The data are provided in Table 4.

3. Analysis

To rule out the possibility that either object might actually be a blended stellar eclipsing binary system, we conducted a blend analysis similar to that done in Hartman et al. (2012). We find that in both cases we cannot rule out blend scenarios based on the photometric observations alone, however all scenarios which fit the photometric data predict significant RV and BIS variations (greater than 300 m s^{-1} for HAT-P-42 and greater than 1 km s^{-1} for HAT-P-43). Such variations of the RV and BIS are ruled out by our spectroscopic observations (Fig. 4). We conclude that both objects are transiting planet systems.

Having confirmed the planetary nature of both systems, we proceeded with their analysis following the methods described in Bakos et al. (2010), with some modifications as described in Hartman et al. (2012). Briefly, this consists of: (1) inferring the stellar atmospheric parameters from the available high-resolution spectra (we used the TRES spectra of HAT-P-42, and the TRES and FIES spectra of HAT-P-43, together with the stellar parameter classification (SPC) method of Buchhave et al. 2012); (2) conducting a global Markov-chain Monte Carlo (MCMC)-based modeling of the available photometric light curves and RVs (we fix the limb darkening coefficients using the tables in Claret 2004); (3) using the spectroscopically inferred stellar effective temperatures and metallicities, together with the stellar densities determined from the light curve modeling, and the Yonsei-Yale theoretical stellar evolution models (Yi et al. 2001), to determine the stellar masses, radii and ages, as well as the planetary parameters (e.g. mass and radius) which depend on these values (Fig. 5); (4) re-analyzing the high-resolution spectra fixing the stellar surface gravities to the values found in (3), and then re-iterating steps (2) and (3).

In fitting the photometric follow-up light curves, we apply the external parameter decorrelation (EPD) and trend-filtering algorithms (TFA) to correct for possible systematic errors in the photometry. As of Bakos et al. (2010) this has been done as a matter of course in analyzing all HATNet discoveries. This amounts to modeling the observed magnitude m_{ij} , which is the j th point in light curve i , as:

$$m_{ij} = \bar{m}_i + \sum_{k=1}^{N_{\text{EPD}}} u_{ik} \text{EPD}_{ijk} + \sum_{k=1}^{N_{\text{TFA}}} v_k \text{TFA}_{jk} + MA(t_{ij}, \theta) \quad (1)$$

where \bar{m}_i is the out-of-transit magnitude of light curve i , MA is the Mandel & Agol (2002) transit model, and sums including the EPD_{ijk} and TFA_{jk} terms are the EPD and TFA filters, respectively. The MA model takes as input the time of

Table 3. Summary of spectroscopic observations.

Instrument	Date(s)	Number of observations	Resolution $\lambda/\Delta\lambda$	Wavelength coverage [Å]	Velocity ^a precision
HAT-P-42					
FLWO 1.5 m/TRES	2011 Oct.–2012 Jan.	13	44 000	3900–8900	$\sim(30 + 35)^b$ m s ⁻¹
OHP 1.93 m/SOPHIE	2011 Dec. 7–13	5	39 000	3900–6900	$\sim(6 + 35)^c$ m s ⁻¹
HAT-P-43					
NOT 2.5 m/FIES	2011 Oct. 24–27	2	46 000	3700–7300	~ 100 m s ⁻¹
FLWO 1.5 m/TRES	2011 Nov.–2012 Jan.	2	44 000	3900–8900	~ 100 m s ⁻¹
OHP 1.93 m/SOPHIE	2011 Dec. 6–13	6	39 000	3900–6900	$\sim(10 + 25)^c$ m s ⁻¹

Notes. ^(a) We give the median velocity precision of the observations, and not the best precision that can be attained with the given instrument. The FIES and TRES observations of HAT-P-43 were done in reconnaissance mode (lower S/N) which so they yield velocities that are less precise than the SOPHIE observations. ^(b) A precision of 30 m s⁻¹, plus 35 m s⁻¹ added in quadrature to achieve a χ^2 per degree of freedom of unity for the Keplerian fit of HAT-P-42. ^(c) Values within parentheses include the sum of the theoretical noise due to the instrument (6 and 10 m s⁻¹, respectively), plus empirical errors taking into account the correction due to moonlight in the spectra.

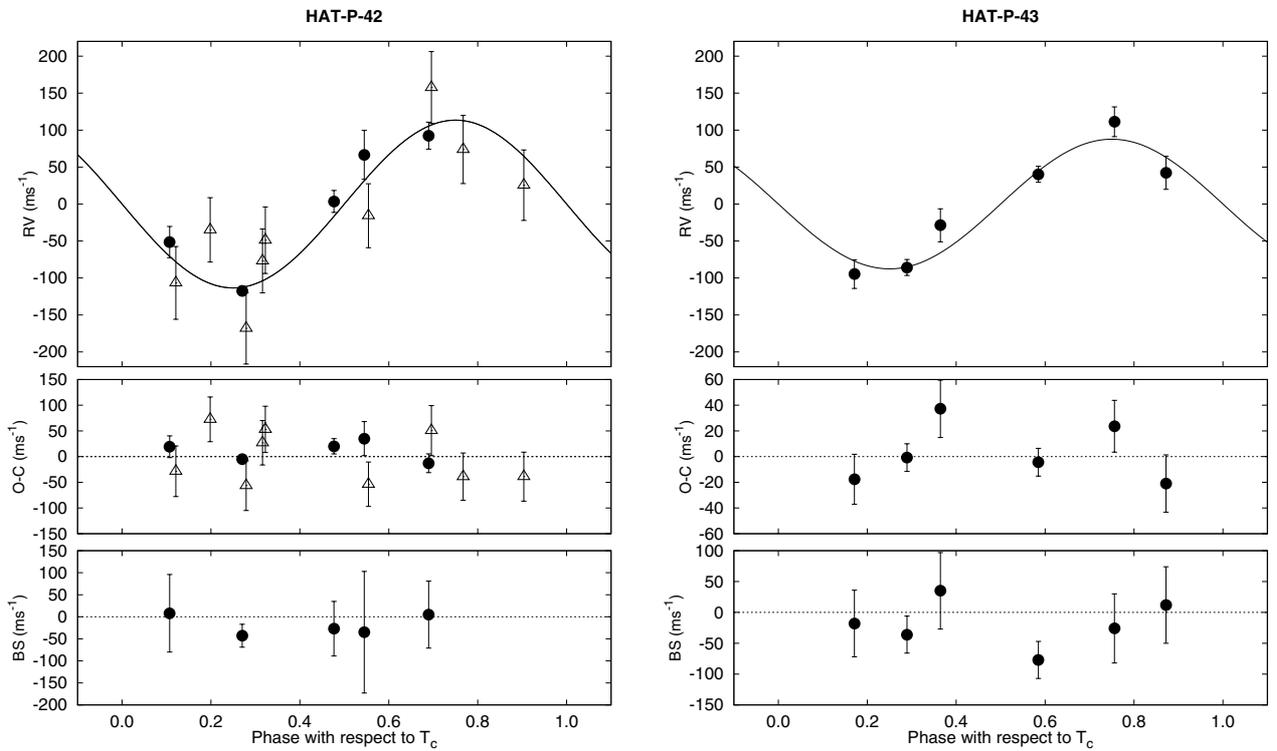


Fig. 3. High precision RV measurements and bisector spans (BIS) for HAT-P-42 (*left*) and HAT-P-43 (*right*). Filled circles show OHP 1.93 m/SOPHIE observations, while open triangles show FLWO 1.5 m/TRES observations. The *top panels* show the RV measurements as a function of orbital phase, along with our best-fit circular orbit model. Zero phase corresponds to the time of mid-transit. The center-of-mass velocity has been subtracted for each system. The *middle panels* show velocity O–C residuals from the best fit circular orbit model. The *bottom panels* show the BIS values, with the mean values subtracted. Note the different vertical scales of the panels.

observation t_{ij} and a set of physical parameters θ , which includes: the radii ratio R_p/R_* , the linear and quadratic limb darkening coefficients c_1 and c_2 , the reciprocal of the half duration of the transit ζ/R_* , which is related to a/R_* by the expression $\zeta/R_* = a/R_*(2\pi(1 + e \sin \omega))/(P\sqrt{1 - b^2}\sqrt{1 - e^2})$ (Bakos et al. 2010), the minimum normalized impact parameter squared b_{\min}^2 , two reference times of mid-transit T_1 and T_2 (we assume a strictly periodic ephemeris), and parameters characterizing the eccentricity of the orbit $\sqrt{e} \sin \omega$ and $\sqrt{e} \cos \omega$. The coefficients u_{jk} and v_k are free parameters giving the contribution from the k th EPD vector for light curve i and the contribution from the k th TFA vector, while EPD_{ijk} and TFA_{jk} are the components of the EPD and TFA vectors, respectively. For the Keplercam follow-up light curves we use 5 EPD vectors per light curve,

including the hour angle of the observation, the square of the hour angle, and three parameters describing the shape of the point spread function for that observation; for the FTS and BOS follow-up light curves we use only one, rather than three, parameters to describe the point spread function shape. We use 20 TFA vectors, corresponding to the light curves for 20 non-variable stars observed simultaneously with each object. Due to the limited field of view, which restricts the number of potential comparison stars, we do not include the TFA terms in the models for the BOS or FTS light curves. Note that the number of free parameters in this model is much less than the number of data points in our light curves. Nonetheless by including a number of empirical detrending filters, there is a risk of overfitting the data. Because the filtering is done simultaneously with the fitting in

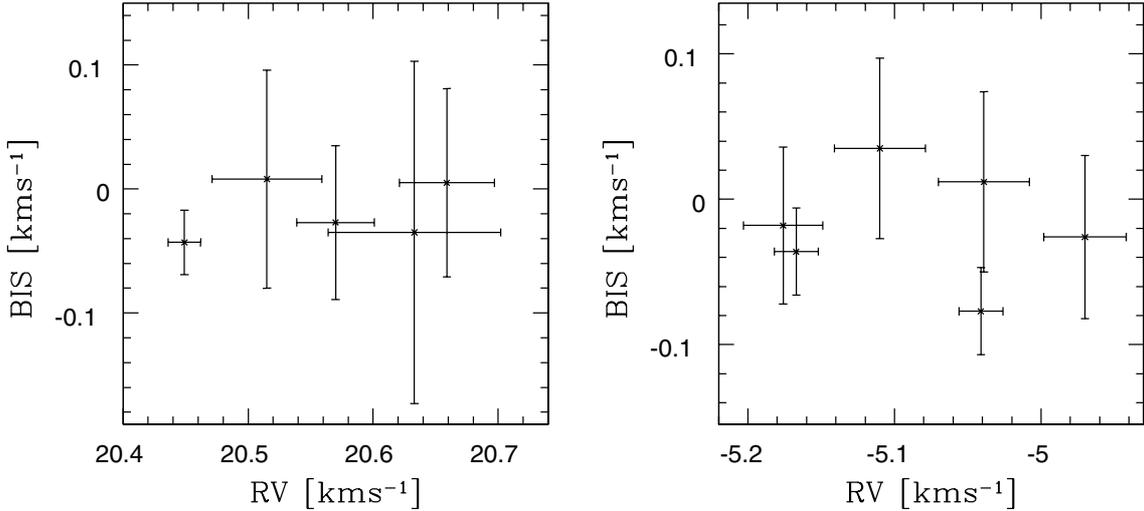


Fig. 4. Bisector spans (BIS) as a function of the high precision RV measurements for HAT-P-42 (*left*) and HAT-P-43 (*right*) from the OHP 1.93 m/SOPHIE observations. In both panels the scale is the same in the x and y axes. The mean values of the BIS have been subtracted for each system.

Table 4. Relative radial velocities and bisector spans of HAT-P-42 and HAT-P-43.

Object	BJD ^a	RV ^b	σ_{RV} ^c	BIS ^d	Phase
	-2455 000	m s ⁻¹	m s ⁻¹	m s ⁻¹	
HAT-P-42	889.03539	-47.4	28.4	...	0.322
	900.05380	159.1	33.8	...	0.696
	901.01864	27.1	32.4	...	0.904
	902.02682	-105.1	34.6	...	0.121
	902.71983	-112.1	13.0	-43.0	0.270
	902.93220	-75.4	25.4	...	0.316
	903.67805	8.9	31.0	-27.0	0.477
	904.03850	-14.2	25.4	...	0.554
	904.66696	97.9	38.0	5.0	0.690
	905.02718	75.3	29.9	...	0.767
	906.60475	-46.1	44.0	8.0	0.107
	907.02711	-33.4	25.9	...	0.198
	908.63608	71.9	69.0	-35.0	0.545
	939.89586	-166.5	33.4	...	0.279
HAT-P-43	901.68768	-85.8	15.0	-36.0	0.289
	902.67245	40.3	15.0	-77.0	0.585
	903.62875	42.3	31.0	12.0	0.872
	904.62601	-94.8	27.0	-18.0	0.171
	906.57509	111.3	28.0	-26.0	0.756
	908.60322	-28.8	31.0	35.0	0.364

Notes. Observations without a BIS value were obtained with the TRES instrument, observations with a BIS value were obtained with SOPHIE. ^(a) Barycentric Julian Date calculated directly from UTC, *without* correction for leap seconds. ^(b) A zero-level γ velocity, fitted independently for each instrument, has been subtracted from these measurements. For reference, for HAT-P-42, the SOPHIE γ velocity from the fit was 20.561 ± 0.014 km s⁻¹, which is an estimate of the true recession velocity of the center of mass of the system. For TRES the value is 0.014 ± 0.015 km s⁻¹. As an artifact of the reduction procedure the TRES RVs were measured relative to an arbitrary zero-point, the fitted γ velocity in this case has no physical meaning. For HAT-P-43, the SOPHIE γ velocity from the fit was -5.082 ± 0.009 km s⁻¹. ^(c) Internal errors excluding any component of astrophysical jitter. ^(d) We use the relation $\sigma_{BIS} = 2\sigma_{RV}$ to estimate the BIS uncertainties.

our MCMC, overfitting may inflate the resulting parameter uncertainties, by allowing bad models to fit the data, but will not prevent the correct model from fitting the data, and so should not bias the results away from the true values.

In conducting the analysis we inflated the TRES RV errors for HAT-P-42 by adding in quadrature a value of 35 m s⁻¹ to the internal errors. This “jitter” is needed to achieve a χ^2 per degree of freedom of unity for the TRES RVs in the best-fit model. On the other hand, the SOPHIE RVs of HAT-P-42 or HAT-P-43 show in both cases a χ^2 per degree of freedom less than one, indicating that the SOPHIE formal uncertainties may be overestimated. We completed the analysis by rescaling the error bars to obtain a χ^2 per degree of freedom equal to one.

For both systems we conducted the analysis fixing the eccentricities to zero, as well as allowing the eccentricities to vary. For each system we find that the eccentricity is consistent with zero (the 95% confidence upper limits on the eccentricity are $e < 0.2$ for HAT-P-42b, and $e < 0.29$ for HAT-P-43b). Following the suggestion of Anderson et al. (2012) we adopt the parameter values associated with the fixed circular orbits. The adopted stellar parameters are given in Table 5, while the adopted planetary parameters are given in Table 7. For completeness we also provide the parameters which result when the eccentricities are allowed to be non-zero in Tables 6 and 8. Note that all the eccentric parameters are within 1σ of the circular orbit values.

We searched for sinusoidal signals and additional transits in the light curves. Neither present a significant peak in the Lomb-Scargle periodogram (false alarm probability of 50% or more). After removing the detected in-transit data, neither light curve shows a transit signal that would pass our automated or by-eye selections.

4. Discussion

We have presented the discovery, confirmation and characterization of two new transiting planets. HAT-P-42b and HAT-P-43b are inflated hot Jupiters with $P = 4.641878 \pm 0.000032$ and 3.332687 ± 0.000015 days, $M_p = 1.044 \pm 0.083$ and $0.662 \pm 0.060 M_J$, and $R_p = 1.280 \pm 0.153$ and $1.281^{+0.062}_{-0.033} R_J$, respectively. Assuming zero albedo and full heat redistribution, and using the stellar and planetary parameters determined by our analysis, their equilibrium temperatures are in the same regime, i.e. $T_{eq} = 1428 \pm 60$ and 1361 ± 25 K, respectively. HAT-P-42b has larger uncertainties on its derived parameters as the transit depth is shallower due to the larger radius of the slightly

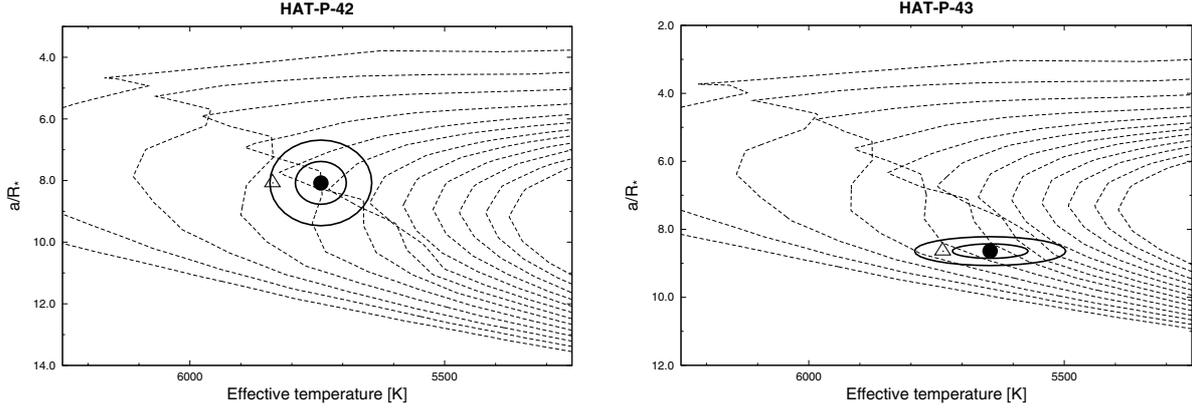


Fig. 5. Model isochrones from Yi et al. (2001) for the metallicities of HAT-P-42 (left) and HAT-P-43 (right) and ages of 1 to 13 Gyr in 1 Gyr increments (left to right). The adopted values of effective temperature, $T_{\text{eff}\star}$, and a/R_{\star} are shown together with their 1σ and 2σ confidence ellipsoids. In each plot the initial values of $T_{\text{eff}\star}$ and a/R_{\star} from the first SPC (see Sect. 3) and light curve analyses are represented with a triangle.

Table 5. Adopted stellar parameters for HAT-P-42 and HAT-P-43 assuming circular orbits.

Parameter	HAT-P-42 Value	HAT-P-43 Value	Source
<i>Astrometric properties</i>			
GSC ID	GSC 0232-01451	GSC 0801-00608	
2MASS-ID	2MASS 09012265+0605500	2MASS 08354217+1012239	
RA (J2000)	09 ^h 01 ^m 22.66s	08 ^h 35 ^m 42.18s	2MASS
Dec (J2000)	+06°05′50.0″	+10°12′24.0″	2MASS
μ_{RA} (mas yr ⁻¹)	-6.2 ± 1.9	-10.3 ± 2.6	UCAC4
μ_{Dec} (mas yr ⁻¹)	-29.3 ± 2.0	-16.0 ± 3.2	UCAC4
<i>Spectroscopic properties</i>			
$T_{\text{eff}\star}$ (K)	5743 ± 50	5645 ± 74	SPC ^a
[Fe/H]	0.27 ± 0.08	0.23 ± 0.08	SPC
$v \sin i$ (km s ⁻¹)	3.5 ± 0.5	2.4 ± 0.5	SPC
γ_{RV} (km s ⁻¹)	20.26 ± 0.15 ^b	-4.86 ± 0.15 ^b	TRES
<i>Photometric properties</i>			
B (mag)	12.827 ± 0.020	14.120 ± 0.060	APASS
V (mag)	12.168 ± 0.030	13.356 ± 0.030	APASS
J (mag)	10.960 ± 0.023	12.146 ± 0.021	2MASS
H (mag)	10.677 ± 0.024	11.809 ± 0.029	2MASS
K_s (mag)	10.626 ± 0.026	11.764 ± 0.023	2MASS
<i>Derived properties</i>			
M_{\star} (M_{\odot})	1.178 ± 0.068	1.048 ^{+0.031} _{-0.042}	YY+ a/R_{\star} +SPC ^c
R_{\star} (R_{\odot})	1.530 ± 0.140	1.103 ^{+0.041} _{-0.021}	YY+ a/R_{\star} +SPC
$\log g_{\star}$ (cgs)	4.14 ± 0.07	4.37 ± 0.02	YY+ a/R_{\star} +SPC
L_{\star} (L_{\odot})	2.27 ± 0.43	1.11 ± 0.10	YY+ a/R_{\star} +SPC
M_V (mag)	3.94 ± 0.21	4.73 ± 0.11	YY+ a/R_{\star} +SPC
M_K (mag,ESO)	2.38 ± 0.20	3.11 ± 0.07	YY+ a/R_{\star} +SPC
Age (Gyr)	5.2 ^{+1.8} _{-0.7}	5.7 ^{+1.9} _{-1.1}	YY+ a/R_{\star} +SPC
A_V (mag) ^d	0.000 ± 0.020	0.000 ± 0.019	YY+ a/R_{\star} +SPC
Distance (pc)	447 ± 41	542 ⁺²² ₋₁₆	YY+ a/R_{\star} +SPC

Notes. ^(a) SPC = “stellar parameter classification” method, described by Buchhave et al. (2012), which derives stellar atmospheric parameters from high-resolution spectra. These parameters rely primarily on SPC, but have a small dependence also on the iterative analysis incorporating the isochrone search and global modeling of the data, as described in the text. ^(b) These velocities corresponds to an absolute scale references thanks to nightly observations of HD 182488 and correction of the gravitational redshift of the Sun. ^(c) YY+ a/R_{\star} +SPC = Based on the YY isochrones (Yi et al. 2001), a/R_{\star} as a luminosity indicator, and the SPC results. ^(d) V band extinction determined by comparing the measured 2MASS and APASS photometry for the star to the expected magnitudes from the YY+ a/R_{\star} +SPC model for the star. We use the Cardelli et al. (1989) extinction law.

more evolved host. Within 1σ , HAT-P-42b and HAT-P-43b have the same radius and T_{eq} . They mainly differ by their periods (semi-major axes) and by their masses (densities). HAT-P-43b is very similar to HAT-P-4b ($M_p = 0.672 M_J$, $P = 3.05$ d, $R_p = 1.27 R_J$, Kovács et al. 2007) with HAT-P-4 slightly more massive, hotter and more evolved star, with similar metallicity ([Fe/H] = 0.24), and to OGLE-Tr-10b ($M_p = 0.63 M_J$,

$P = 3.1$ d, $R_p = 1.25 R_J$, Torres et al. 2008). HAT-P-42b is similar to CoRoT-19b ($M_p = 1.107 M_J$, $P = 3.9$ d, $R_p = 1.29 R_J$, Guenther et al. 2012). HAT-P-13b is the other closest analog to HAT-P-42b and HAT-P-43b but in the intermediate mass-regime ($M_p = 0.851 M_J$, $P = 2.9$ d, $R_p = 1.28 R_J$, Winn et al. 2010b; Bakos et al. 2009). The stellar hosts of all the previously quoted planets are slow rotators, with late F/early G spectral type, and

Table 6. Derived stellar parameters for HAT-P-42 and HAT-P-43 allowing nonzero eccentricity.

Parameter	HAT-P-42 Value	HAT-P-43 Value
$M_\star (M_\odot)$	1.165 ± 0.064	1.043 ± 0.048
$R_\star (R_\odot)$	1.411 ± 0.173	$1.158^{+0.136}_{-0.072}$
$\log g_\star$ (cgs)	4.21 ± 0.09	4.33 ± 0.07
$L_\star (L_\odot)$	$1.95^{+0.60}_{-0.39}$	$1.20^{+0.33}_{-0.16}$
M_V (mag)	4.10 ± 0.27	4.65 ± 0.21
M_K (mag, ESO)	2.55 ± 0.26	3.01 ± 0.19
Age (Gyr)	$4.9^{+1.5}_{-0.8}$	$6.6^{+2.5}_{-1.5}$
A_V (mag) ^c	0.000 ± 0.023	0.000 ± 0.014
Distance (pc)	414 ± 51	566^{+67}_{-37}

Notes. Quantities and abbreviations are as in Table 5, which gives our adopted values, determined assuming circular orbits. We do not list parameters that are independent of the eccentricity.

Table 7. Adopted orbital and planetary parameters for HAT-P-42b and HAT-P-43b assuming circular orbits.

Parameter	HAT-P-42 Value	HAT-P-43 Value
<i>Light curve parameters</i>		
P (days)	4.641878 ± 0.000032	3.332687 ± 0.000015
T_c (BJD) ^a	$2\,455\,952.52603 \pm 0.00075$	$2\,455\,997.37106 \pm 0.00032$
T_{14} (days) ^a	0.1681 ± 0.0039	0.1354 ± 0.0011
$T_{12} = T_{34}$ (days) ^a	0.0193 ± 0.0037	0.0149 ± 0.0009
a/R_\star	$8.08^{+0.82}_{-0.54}$	$8.64^{+0.12}_{-0.28}$
ζ/R_\star ^b	13.40 ± 0.14	16.60 ± 0.08
R_p/R_\star	0.0860 ± 0.0033	0.1193 ± 0.0018
b^2	$0.333^{+0.085}_{-0.141}$	$0.036^{+0.061}_{-0.024}$
$b \equiv a \cos i/R_\star$	$0.577^{+0.068}_{-0.162}$	$0.190^{+0.111}_{-0.096}$
i (deg)	$85.9^{+1.3}_{-0.8}$	88.7 ± 0.7
<i>Limb-darkening coefficients^c</i>		
c_1, i (linear term)	0.2863	0.3003
c_2, i (quadratic term)	0.3274	0.3180
<i>RV parameters</i>		
K (m s ⁻¹)	113.5 ± 7.9	87.7 ± 7.7
e	0 (fixed)	0 (fixed)
<i>Planetary parameters</i>		
M_p (M_J)	1.044 ± 0.083	0.662 ± 0.060
R_p (R_J)	1.280 ± 0.153	$1.281^{+0.062}_{-0.033}$
$C(M_p, R_p)$ ^d	0.23	0.09
ρ_p (g cm ⁻³)	$0.61^{+0.34}_{-0.15}$	0.39 ± 0.05
$\log g_p$ (cgs)	3.20 ± 0.10	3.00 ± 0.05
a (AU)	0.0575 ± 0.0011	$0.0443^{+0.0004}_{-0.0006}$
T_{eq} (K) ^e	1428 ± 60	1361 ± 25
Θ ^f	$0.079^{+0.014}_{-0.009}$	0.044 ± 0.004
$\langle F \rangle$ (10 ⁸ erg s ⁻¹ cm ⁻²) ^g	9.38 ± 1.56	$7.74^{+0.65}_{-0.47}$

Notes. ^(a) Reported times are in Barycentric Julian Date calculated directly from UTC, *without* correction for leap seconds. T_c : reference epoch of mid transit that minimizes the correlation with the orbital period. T_{14} : total transit duration, time between first to last contact; $T_{12} = T_{34}$: ingress/egress time, time between first and second, or third and fourth contact. ^(b) Reciprocal of the half duration of the transit used as a jump parameter in our MCMC analysis in place of a/R_\star . It is related to a/R_\star by the expression $\zeta/R_\star = a/R_\star(2\pi(1 + e \sin \omega))/(P\sqrt{1 - b^2}\sqrt{1 - e^2})$ (Bakos et al. 2010). ^(c) Values for a quadratic law, adopted from the tabulations by Claret (2004) according to the spectroscopic (SPC) parameters listed in Table 5. ^(d) Correlation coefficient between the planetary mass M_p and radius R_p . ^(e) Planet equilibrium temperature averaged over the orbit, calculated assuming a Bond albedo of zero, and that flux is re-radiated from the full planet surface. ^(f) The Safronov number is given by $\Theta = \frac{1}{2}(V_{\text{esc}}/V_{\text{orb}})^2 = (a/R_p)(M_p/M_\star)$ (see Hansen & Barman 2007). ^(g) Incoming flux per unit surface area, averaged over the orbit.

with enhanced metallicity with respect to the Sun (except for CoRoT-19, which is solar).

HAT-P-42b and HAT-P-43b are inflated hot Jupiters compared to models of coreless giant planets. The models of Fortney et al. (2007) predict a maximal radius of $\sim 1.1 R_J$ for both planet (see their Table 4 and Fig. 7) considering the system age, planet

distance and mass (maximal in the meaning of a core free planet of pure H-He). Laughlin et al. (2011) derived a simple fitting relation from the Bodenheimer et al. (2003) models to infer the expected radius of a H-He composition planets for different masses and insolation. From the Laughlin et al. (2011) relations, we derived a predicted radius of 1.19 and 1.17 R_J , respectively for

Table 8. Orbital and planetary parameters for HAT-P-42b and HAT-P-43b allowing eccentric orbits.

Parameter	HAT-P-42 Value	HAT-P-43 Value
<i>Light curve parameters</i>		
a/R_\star	8.74 ± 0.96	$8.20^{+0.54}_{-0.74}$
ζ/R_\star	13.41 ± 0.14	16.62 ± 0.09
i (deg)	86.6 ± 1.1	$88.5^{+0.7}_{-1.0}$
<i>RV parameters</i>		
K (m s ⁻¹)	108.2 ± 8.9	89.4 ± 8.6
$\sqrt{e} \cos \omega$	$-0.143^{+0.119}_{-0.086}$	$-0.134^{+0.182}_{-0.132}$
$\sqrt{e} \sin \omega$	$-0.191^{+0.238}_{-0.157}$	$0.164^{+0.162}_{-0.220}$
$e \cos \omega$	-0.042 ± 0.034	$-0.035^{+0.047}_{-0.066}$
$e \sin \omega$	$-0.050^{+0.063}_{-0.091}$	$0.041^{+0.094}_{-0.058}$
e	0.084 ± 0.063	0.084 ± 0.070
ω (deg)	231 ± 62	132 ± 74
<i>Secondary eclipse parameters</i>		
T_s (BJD)	$2\,455\,954.723 \pm 0.101$	$2\,455\,995.631 \pm 0.122$
$T_{s,14}$	0.1568 ± 0.0177	0.1465 ± 0.0240
$T_{s,12}$	0.0155 ± 0.0051	0.0165 ± 0.0038
<i>Planetary parameters</i>		
M_p (M_J)	0.981 ± 0.097	0.672 ± 0.068
R_p (R_J)	$1.175^{+0.193}_{-0.146}$	$1.348^{+0.163}_{-0.089}$
$C(M_p, R_p)$	0.48	0.41
ρ_p (g cm ⁻³)	$0.75^{+0.40}_{-0.22}$	0.34 ± 0.08
$\log g_p$ (cgs)	3.24 ± 0.11	$2.96^{+0.06}_{-0.08}$
a (AU)	0.0573 ± 0.0010	0.0443 ± 0.0007
T_{eq} (K)	1378 ± 75	1390^{+79}_{-47}
Θ	$0.081^{+0.014}_{-0.010}$	0.042 ± 0.005
$\langle F \rangle$ (10^8 erg s ⁻¹ cm ⁻²)	$8.14^{+2.15}_{-1.47}$	$8.43^{+2.27}_{-1.05}$

Notes. Quantities and definitions are as in Table 7, which gives our adopted values, determined assuming circular orbits. Here we do not list parameters that are effectively independent of the eccentricity.

HAT-P-42b and HAT-P-43b, still at the 1- σ and 3- σ level from our observations.

Enoch et al. (2012) determined an empirical relation between R_p , T_{eq} , and the semi-major axis a (see their Eq. (9)). This leads to an estimated radius of $1.325 \pm 0.054 R_J$ for HAT-P-42b, matching the observed radius within the error bars. The same relations predict a radius of $1.180^{+0.021}_{-0.023} R_J$ for HAT-P-43b, 3- σ below our measured value. Considering also the results observed for the highly inflated hot Jupiters presented by Hartman et al. (2012), it is possible that parameters other than T_{eq} and a play a role in the planetary radius in the mass domain where the equation was derived for $0.5 < M_p < 2 M_J$. Independently, Béky et al. (2011) derived a relation to determine the radius for the planets with $0.3 < M_p < 0.8 M_J$ that we could apply to HAT-P-43b. Their equation leads to an inferred radius of $1.133 \pm 0.044 R_J$. Note that the predicted radius is smaller than the measured value.

Another caveat in using the above relations is the assumption of zero Bond albedo and full redistribution of the heat when calculating T_{eq} , whereas in reality these assumptions are unlikely to be true for all of the planets.

Few measurements per orbit lead to poor constraints on a small eccentricity. We followed the recommendation of Anderson et al. (2012), and adopted circular orbits, after a significance test on the eccentricity measurement. We note that observing the secondary transit detection is a very efficient way to characterize the shape of the orbit, that would need several tens of RV values to obtain an equivalent precision (e.g. Boisse et al. 2009; Husnoo et al. 2011). Another method to constrain a small

eccentricity is the measurement of the RVs of planetary lines during the transit, but that would currently need brighter hosts to be significantly detected (Snellen et al. 2010; Montalto et al. 2011).

The Rossiter-McLaughlin effect measured via spectroscopy during the transit allows the measurement of the sky-plane projected angle between the stellar spin-axis and the planetary orbital momentum vector. With the hypothesis that the projected angle is close to zero, the expected semi-amplitude of the effect is of 26 and 34 m s⁻¹ for HAT-P-42 and HAT-P-43, respectively.

We note that the transit of HAT-P-42b is far from equatorial ($b = 0.577^{+0.068}_{-0.162}$), which facilitates the measurement of the projected angle, since there is a smaller correlation between the stellar $v \sin i$ and the projected spin-orbit angle. According to Winn et al. (2010a), planets around stars with the $T_{\text{eff}\star} < 6250$ K are preferentially aligned (except for the systems with longest timescales for obliquity damping, e.g. HD 80606). Additionally, the (old) age of the host stars as well as the (low) mass of the planets also increase the chances that both planetary systems are well-aligned (Triaud 2011; Hébrard et al. 2011). We applied Eq. (2) of Albrecht et al. (2012) where they estimated the tidal timescale to align a planet with the stellar equator from a calibration of binary studies considering that our targets have convective envelopes (CE). These results also favor alignment for both HAT-P-42 and HAT-P-43 (the calculated timescale are equal respectively to $\tau_{\text{CE}} = 247$ and 558 to be reported in their Fig. 24).

To conclude, hot gas giant planets are surprisingly diverse, and the physical reasons behind the diversity and the observed

correlations are not well understood. It is likely that the physical properties of such planets depend on a large number of parameters. Hence, there is a continued need to broaden the sample of well-characterized transiting planets in order to develop a comprehensive view of exoplanet formation, structure and evolution.

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