

The First AllWISE Proper Motion Discovery: WISEA J070720.50+170532.7

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ABSTRACT

While quality checking a new motion-aware co-addition of all 12.5 months of WISE data, the source WISE J070720.48+170533.0 was found to have moved $0.9''$ in 6 months. Backtracking this motion allowed us to identify this source as 2MASS J07071961+1705464 and with several entries in the USNO B catalog. An astrometric fit to these archival data gives a proper motion of $\mu = 1793 \pm 2$ mas/yr and a parallax of $\varpi = 35 \pm 42$ mas. Photometry from WISE, 2MASS and the POSS can be fit reasonably well by a blackbody with $T = 3658$ K and an angular radius of 4.36×10^{-11} radians. No clear evidence of H₂ collision-induced absorption is seen in the near-IR. An optical spectrum shows broad deep CaH bands at 638 & 690 nm, broad deep Na D at 598.2 nm, and weak or absent TiO, indicating that this source is an ultra-subdwarf M star with a radial velocity $v_{rad} \approx -21 \pm 18$ km/sec relative to the Sun. Given its apparent magnitude, the distance is about 39 ± 9 pc and the tangential velocity is probably ≈ 330 km/sec, but a more precise parallax is needed to be certain.

Subject headings: brown dwarfs – infrared:stars – solar neighborhood – stars:late-type – stars:low-mass – stars:individual

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1. Introduction

The Wide-field Infrared Survey Explorer (*WISE*) (Wright et al. 2010) mapped the entire sky between 14 Jan 2010 and 17 Jul 2010 in 4 thermal infrared bands centered at 3.4, 4.6, 12 & 22 μm , then continued on to map the entire sky again prior to 1 Feb 2011. On 7 Aug 2010 the outer cryogen tank ran out of solid hydrogen coolant, and the 3-band cryo phase began, where the 22 μm channel was lost and the 12 μm channel had much reduced sensitivity due to the background radiated by the optics which warmed to 45 K. On 30 Sep 2010 the inner cryogen tank ran out of solid hydrogen, and the two band post-cryo phase began. This phase lasted until 1 Feb 2011, with only the 3.4 and 4.6 μm channels operat-

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ing. This last phase of the mission, which lasted for 4 months, was never co-added or catalogued. Only single frame detections were found to search for Near Earth Objects as part of the NEOWISE project (Mainzer et al. 2011).

The AllWISE project is an effort to co-add all the WISE frames, find peaks in these coadds, and then fit for the properties of the source or sources responsible for each peak by doing profile fit photometry using the pixels on each frame that covers the position of the peak. While doing this fit, the AllWISE project is allowing for a motion in addition to the usual parameters of the position and the fluxes in each of the 4 bands.

Previous efforts to find nearby stars with the WISE data have concentrated on color-selected searches. Kirkpatrick et al. (2011) found over one hundred brown dwarfs using the color cut $W1-W2 > 2$, and many more have been studied by other papers such as Mace et al. (2013). These color-selected searches have found many high proper motion objects. However, several papers have searched for moving objects by using the 2MASS match ($3''$ radius) in the WISE catalog as a veto (Gizis et al. 2011; Liu et al. 2011; Thompson et al. 2013). With AllWISE we seek to find previously missed high motion objects using positive associations between observations over the full time span of WISE. This has been done for objects bright enough to appear in WISE single frame detections by Luhman (2013), leading to the discovery of a brown dwarf binary only 2 pc from the Sun. AllWISE is doing a similar search using coadded frames to reach the full sensitivity of WISE. Kirkpatrick et al. (in preparation) will discuss a large sample of AllWISE motion discoveries, but this brief note describes one of the most interesting AllWISE motion selected source whose colors are similar to hundreds of millions of WISE objects.

2. Observations

2.1. Astrometry & Photometry

2.1.1. AllWISE

The AllWISE catalog is formed by fitting the pixels in the individual frames with a model containing 8 parameters which are the position, motion and fluxes for a point source: $(\alpha, \delta, \mu_\alpha =$

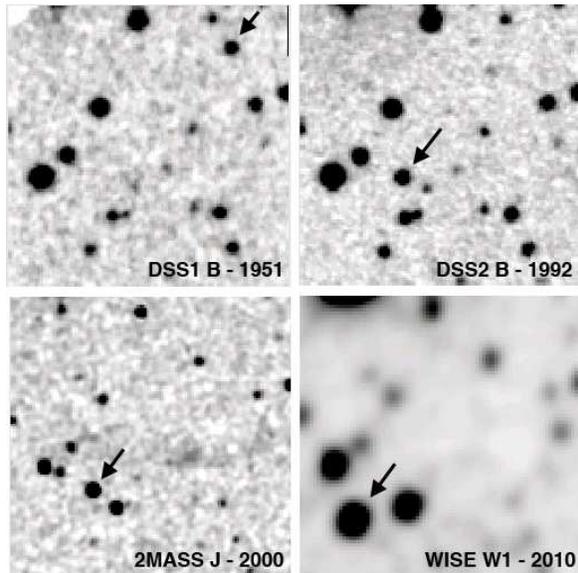


Fig. 1.— 2×2 arcmin images of the field near WISE 0707+1705. North is up and East is to the left. Arrows point to the high proper motion star at each epoch.

$\cos \delta d\alpha/dt, \mu_\delta, F_{W1}, F_{W2}, F_{W3}, \& F_{W4}$). This model is convolved with the point spread function of WISE, compared to the pixel data, and then adjusted to get the best fit. The position and motion parameters, which have a non-linear effect on the pixel values, are often hard to fit in a robust and stable way, so the fitting procedure first finds the best fit with $\mu_\alpha = \mu_\delta = 0$, which just repeats the procedure used for the previous WISE catalogs, and then uses this as a starting point for an iterative adjustment to find the best fitting motion. In addition, the first $\mu = 0$ solution is used to generate the source designation and when extracting fluxes from individual frames that go into the multi-epoch photometry database (see §2.2).

The AllWISE catalog entry for W0707+1705 has the designation WISEA J070720.50+170532.7. The first position solution, for $\mu_\alpha = \mu_\delta = 0$ gives $\alpha = 106.8354461^\circ \pm 0.0446''$, $\delta = 17.0924311^\circ \pm 0.0431''$. The magnitudes obtained in the fixed position solution are $W1234 = 12.646(24), 12.502(25), > 11.701 \& > 8.719$. The contamination and confusion flags are all zero, indicating that the detection is not affected by known artifacts. This

Table 1: Astrometric Data [J2000]

MJD	α [°]	σ_α [mas]	δ [°]	σ_δ [mas]	Source
33957.456	106.814542	195	17.113712	79	USNO B1.0
48986.792	106.829188	333	17.098708	333	POSS2 B FITS WCS
49277.000	106.829498	114	17.098514	25	USNO B1.0
50405.411	106.830604	333	17.097333	333	POSS2 R FITS WCS
50923.000	106.831139	999	17.096873	999	USNO B1.0
51496.920	106.831671	333	17.096272	333	POSS2 IR FITS WCS
51579.657	106.831722	70	17.096224	70	2MASS
51902.000	106.832112	122	17.095854	122	UCAC4
52064.000	106.832222	88	17.095716	63	CMC14
55287.577	106.835361	90	17.092520	90	WISE allsky
55479.133	106.835568	90	17.092344	90	WISE postcryo
56545.350	106.836599	50	17.091278	50	FanCam

position is then used to initialize the simultaneous solution for the position at epoch MJD=55400, the proper motion, and the four WISE magnitudes, giving $\alpha = 106.8354766^\circ \pm 0.0459''$, $\delta = 17.0924049^\circ \pm 0.0446''$, $\mu_\alpha = 1355 \pm 59$ mas/yr, $\mu_\delta = -1256 \pm 62$ mas/yr, and W1234 = 12.660(23), 12.512(25), 12.430(530) & > 8.702 . Note that the W3 flux is essentially identical in both solutions but one is slightly more than 2σ while the other is slightly less and thus quoted as an upper limit. The quoted motion assumes that the parallax is zero, and a star with a non-zero parallax will have a motion that differs from the true proper motion by as much as $\pm 4\varpi$ since the parallactic movement in 6 months is $\pm 2\varpi$ and the time interval is one half year.

Note that the error on the motion of W0707+1705 is considerably smaller than the error on the position. This happens because the reference frame error terms cancel out in the position difference. For the optimal range of magnitudes, $8 < W1 < 10$, with high SNR but not saturated, the motion errors can be as small as 35 mas/yr, 1σ , 1-axis. Thus the smallest motions that can reach 5σ are about 175 mas/yr. But for fainter sources the noise terms dominate the motion error, so for $W1 \approx 15$ the errors are $10\times$ larger, giving a 5σ threshold of about 1750 mas/yr. The upper limit on motions that can be detected by AllWISE is still being investigated and will be discussed further in Kirkpatrick et al. (in preparation). But Barnard’s star and Kapteyn’s star are recovered as moving sources, even though they are quite saturated.

2.1.2. Separated WISE epochs

In order to combine the WISE astrometry with archival data from earlier surveys, we use the position from the WISE Allsky catalog, $\alpha = 106.8353609^\circ \pm 0.0898''$, $\delta = 17.0925197^\circ \pm 0.0889''$ at a mean MJD of 55287.577.

Since W0707+1705 is bright enough to be easily detected on individual frames, the mid-average of the positions found on 12 post-cryogenic detections is used, giving $\alpha = 106.835568^\circ \pm 0.09''$, $\delta = 17.092344^\circ \pm 0.09''$ at a mean MJD of 55479.133.

2.1.3. Cross Identifications

The AllWISE motion backtracked ten years to the 2MASS epoch predicts a position NW of the WISE position. Indeed, an easily seen source in the 2MASS (Skrutskie et al. 2006) image NW of the WISE position is not seen in the WISE image, while the WISE source is not seen in the 2MASS image. This source is 2MASS J07071961+1705464 with $\alpha = 106.831722^\circ \pm 0.07''$, $\delta = 17.096224^\circ \pm 0.07''$, and magnitudes $JHK_s = 13.473(25)$, $13.045(29)$ & $12.923(17)$ on MJD = 51579.657.

W0707+1705 appears in the Carlsberg Meridian Catalogue 14 (Evans et al. 2002, CMC14). This observation also gives an AB magnitude of 15.811 ± 0.17 in the SDSS r band, where the quoted error is the external uncertainty for sources with $r_{AB} \approx 16$ (Evans et al. 2002). The UCAC4 (Zacharias et al. 2013) contains W0707+1705 as entry 536-039034. No proper motion is given. The 579-642 nm magnitude is 15.881 ± 0.30 on

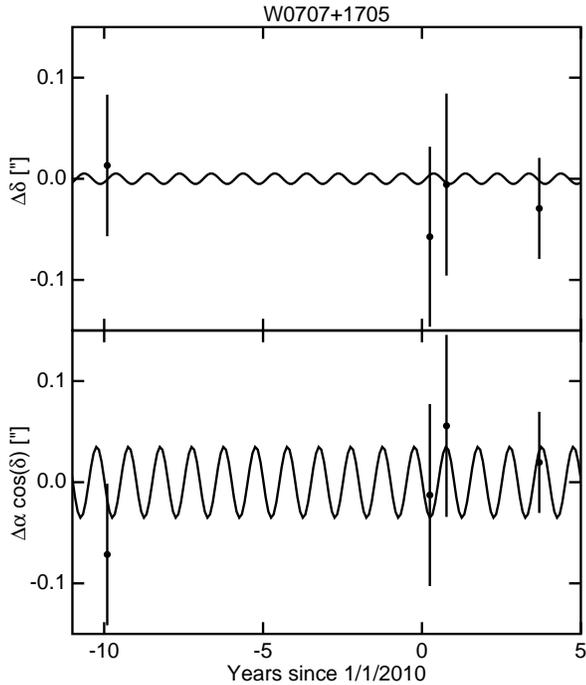


Fig. 2.— Astrometric residuals after a proper motion only fit is removed. Only the more precise data from 2MASS, WISE & FanCam have been plotted to reduce clutter.

the Vega system. These CCD-based surveys give better photometry than photographic plates, but W0707+1705 is close to the survey limit in both cases.

The USNO B1.0 catalog (Monet et al. 2003) contains several entries that appear to refer to W0707+1705. Entries 1071-0154281, 1070-0147856 & 1070-0147865 have positions and epochs that are consistent with the movement of W0707+1705, while the entry 1070-0147855 appears to be W0707+1705 in the second epoch blue plate associated with an incorrect identification on the first epoch blue plate. In addition to the astrometric data, we take photometric values of $B = 17.66, 17.70$ & 18.33 , $R = 15.46$ and $I = 14.97$ from these entries. Some of these entries have large position errors (999 max) but all are reasonably consistent with the astrometric solution.

Magnitudes from WISE, 2MASS and the USNO-B are on the Vega system.

In addition to the WISE, 2MASS & USNO

B astrometry, data were also obtained by fitting parabolae to the data around the peak of the source on FITS postage stamp images from the Digitized Sky Survey maintained at the Space Telescope Science Institute. These values rely on the World Coordinate System in the FITS header and the pixel positions reported by the SAOImage DS9 FITS viewer (Joye & Mandel 2003). For these values the error is taken to be 1/3 of a pixel in the FITS image, typically 333 mas.

2.1.4. Fan Mountain Observatory/FanCam

YJH&K_s photometry and astrometry of W0707+1705 were obtained on 5 Sep 2013 with FanCam, a HAWAII-1 based near-infrared imager operating at the University of Virginia’s Fan Mountain 31-inch telescope (Kanneganti et al. 2009). Results are given in Tables 1 and 2.

All of the astrometric data used in this paper are listed in Table 1. Figure 1 shows images of W0707+1705 from 1952 to 2010.

2.1.5. Ultraviolet

W0707+1705 was not detected by GALEX (Martin et al. 1999) but its position was observed during the Allsky Imaging Survey. Examining the expected position of W0707+1705 on the near UV FITS image of the field shows it is at least 17 times fainter than the nearby source at $\alpha = 106.837313^\circ$, $\delta = 17.099230^\circ$ which has a near UV AB magnitude of 18.32, so W0707+1705 is fainter than 21.4 or $< 10^{-5}$ Jy at $0.24 \mu\text{m}$.

2.1.6. Astrometric Fit

A parallax and proper motion fit to the data in Table 1 gives $\mu_\alpha = 1228 \pm 3$ mas/yr, $\mu_\delta = -1307 \pm 1$ mas/yr, and a parallax of $\varpi = 35 \pm 42$ mas. The residuals give a χ^2 per degree of freedom = 8.72/19, so the errors are reasonable if perhaps conservative. This best fit solution to a long time baseline is consistent with the six month baseline AllWISE catalog motion within 2σ : $\Delta\mu_\alpha = -127 \pm 59$ and $\Delta\mu_\delta = -51 \pm 62$ mas/yr. Figure 2 shows the residuals after a position and proper motion fit is removed.

2.1.7. Photometric Fit

The magnitudes from WISE, 2MASS and the USNO B were fit to a blackbody using a least sum

Table 2: Photometric Observations of W0707+1705.

Filter	Magnitude	Instrument
B	17.9 ± 0.25	USNO B
579-642 nm	15.881 ± 0.3	UCAC4
r_{AB}	15.811 ± 0.17	CMC14
R	15.46 ± 0.33	USNO B
I	14.97 ± 0.33	USNO B
Y(MKO)	13.90 ± 0.01	FanCam
J	13.47 ± 0.01	FanCam
J	13.473 ± 0.025	2MASS
H	13.08 ± 0.01	FanCam
H	13.045 ± 0.029	2MASS
K_s	12.91 ± 0.01	FanCam
K_s	12.923 ± 0.017	2MASS
W1	12.660 ± 0.023	WISEA
W2	12.512 ± 0.025	WISEA
W3	12.430 ± 0.530	WISEA
W4	> 8.702	WISEA

NOTE.—All magnitudes except r_{AB} are Vega magnitudes and the JHK_s bands use 2MASS filters. The Y-band calibration uses the Hamuy et al. (2006) transformation of Y-K_s vs. J-K_s.

of absolute value of error criterion. Monet et al. (2003) state that the photometric calibration of the USNO B catalog is of marginal quality due to the lack of an all-sky network of faint photometric calibration stars, and that the standard deviation of the fit to the existing calibration network is 0.25 mag. Thus we used assumed uncertainties of 0.33 mag for the USNO B except in the blue where there are three consistent estimates so a 0.25 mag uncertainty was used. For WISE and 2MASS the actual magnitude uncertainties were replaced by 0.1 mag to allow for the crudeness of the blackbody model. The error for the CMC14 was rounded up to 0.2 mag for the same reason. The photometric points and the best fit model with $T = 3658$ K and a solid angle of $10^{-20.224}$ sr are shown in Figure 3. This solid angle corresponds to the Sun at a distance of 517 pc with an angular radius of $9 \mu\text{as}$.

A more pertinent comparison is the moderately metal poor Barnard’s star (GJ 699) which has a measured radius of $0.1867 \pm 0.0012 R_{\odot}$

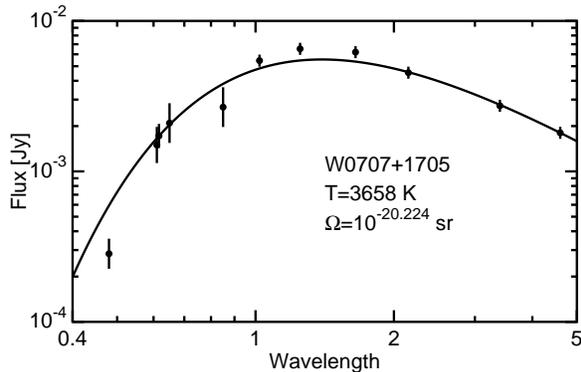


Fig. 3.— Photometric points and a blackbody fit.

(Boyajian et al. 2012). This would put W0707+1705 at a distance of 97 pc. A blackbody fit to the fluxes of Barnard’s star, similar to the one we did for W0707+1705, gives a radius of 587 mas. The distance of Barnard’s star is 1.8 pc, so if W0707+1705 were like Barnard’s star its distance would be 117 pc.

The near-infrared fluxes are clearly not suppressed by H₂ collision-induced absorption, as would be expected for a cool white dwarf with a hydrogen atmosphere like LHS 3250 (Harris et al. 1999). In fact, the slight excess over the fit at J and H might indicate the presence of H⁻ absorption.

2.2. Variability

A comparison of the 2MASS and FanCAM magnitudes shows that W0707+1705 has had a constant brightness over 13 years. But a new product in the AllWISE data release allows one to check for variability over hourly or six month intervals. In addition to checking for motion of sources when generating the catalog, AllWISE is also producing a very large database of multi-epoch photometry. This MEP database contains the flux estimates in each WISE band for each frame that covers a source. These individual frame flux estimates have been generated in past releases, and used to set the variability flag, but not stored. However, for AllWISE the single frame values are kept in a searchable database available at irsa.ipac.caltech.edu. These flux estimates are forced photometry at the position of the source assuming it is celestially fixed, not the moving po-

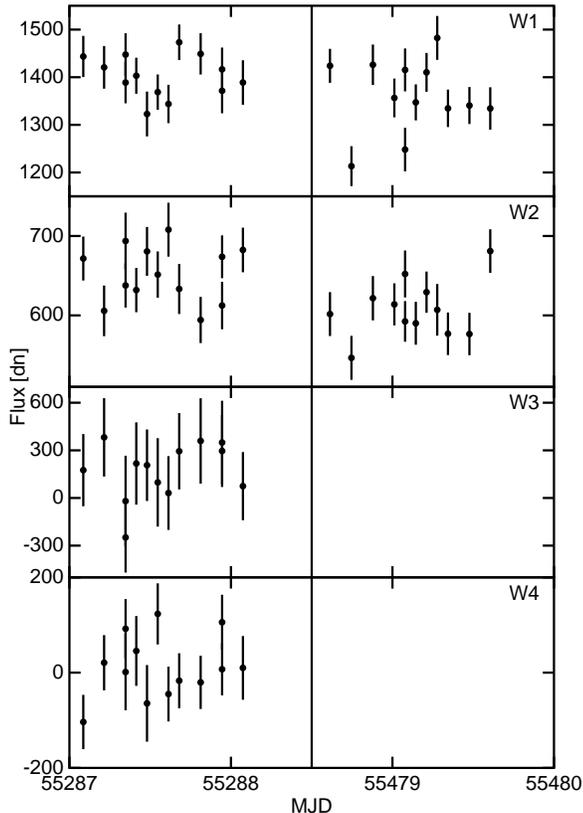


Fig. 4.— Plot of multi-epoch photometry data for W0707+1705. In each band, the left and right panels cover the first and second epochs of observation, separated by 6 months. Fluxes are in data numbers. 1 dn corresponds to magnitudes $W1234 = 20.5, 19.5, 18$ & 13 .

sition from the proper motion fit. Since the All-WISE catalog and reject table combined contain more than 10^9 sources, and each source is observed on an average of 36 frames, the MEP database has 42,759,337,365 rows and thus should not be used for large area blind surveys for variable objects. But it is easy to collect the individual frame data for a source or a list of sources.

The MEP database differs from the previously released single frame detection list in one important regard: the single frame detections had to be $> 5\sigma$ in at least one band in a single frame. For sources that are just over 5σ in W1 in the coadd of 30 frames, which includes a large fraction of catalogued WISE sources, the individual frames will usually only yield upper limits. The single

frame detection list will only have the few frames with upward noise fluctuations leading to a biased view of the source flux. Because most fluxes in the MEP will be insignificant, it is a good idea to use the flux and flux sigma columns for analysis. The fluxes are quoted in data numbers, which are close to the actual quantization in the downloaded images, but the data numbers have been scaled to correct for the different pixel and frame sensitivities, so 1 dn always corresponds to a consistent flux. The zeropoint magnitudes that correspond to a flux of 1 dn are $W1234 = 20.5, 19.5, 18$ & 13 , the same as the zeropoints for the allsky release given in Table 1 of §2.3.f of the Explanatory Supplement.

Figure 4 shows the flux values and uncertainties in the 4 WISE bands for the 25 frames covering W0707+1705. Fitting a constant flux to each band gives 1383.10 ± 8.38 dn with a χ^2 of 54.8 for 24 degrees of freedom in W1, 627.93 ± 5.72 dn with $\chi^2/\text{dof} = 49.9/24$ in W2, 162.7 ± 67.3 dn with $\chi^2/\text{dof} = 7.0/12$ in W3, and 11.24 ± 17.29 dn with $\chi^2/\text{dof} = 14.2/12$ in W4. These uncertainties are statistical only and do not include systematic terms such as the uncertainty in the zeropoint. The W3 value corresponds to $W3 = 12.47$ with an SNR of 2.42:1, providing no evidence for a warm dust disk around W0707+1705.

Moderate improvements in the χ^2 for W1 and W2 can be made by fitting the epochs separately. $\Delta\chi^2 = 6.4$ in W1 with the later epoch $\Delta m = 0.033$ mag fainter. In W2, $\Delta\chi^2$ is 14.9 with the second epoch 0.076 mag fainter. Given all the changes in the detector operating conditions as the telescope warmed up, these results are only weak evidence for a small amount of long term variability over a 6 month time span. But the comparison of 2MASS and FanCam over a 13 year span indicates that W0707+1705 is probably not varying.

2.3. Spectroscopy

We obtained an optical spectrum on UT 2013 September 3 with the Double Spectrograph at Palomar Mountain Observatory through partial cloud coverage during morning twilight. This spectrum is plotted in Figure 5. No calibration star could be observed so the sensitivity calibration is based on calibrations over the previous month, and atmospheric features have not been divided out. The 3680 K blackbody fit to the pho-

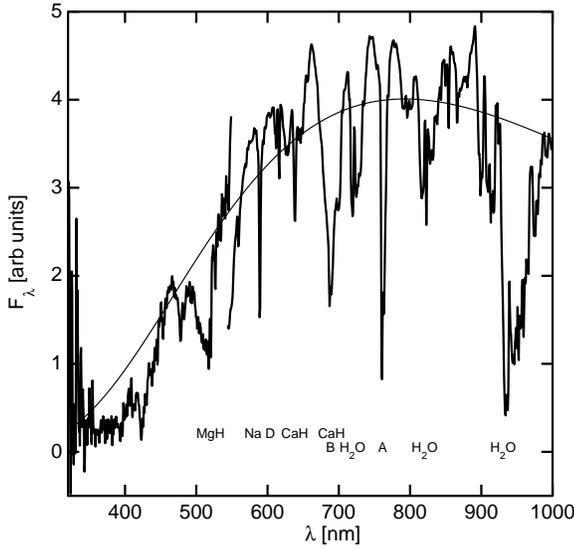


Fig. 5.— Palomar Double Spectrograph spectrum of W0707+1705 is shown as the heavy dark line. The discontinuity near 550 nm is due to the dichroic separating the red and blue sides of the spectrograph. This spectrum has been smoothed to the resolution of the plot with a 1.5 nm FWHM smoothing kernel. The lower line of spectral identifications is due to telluric absorption bands. An arbitrarily scaled 3680 K blackbody gives the thin solid curve.

ometric data is a reasonable match to the trend of the spectrum, but the flux scaling is arbitrary.

The sodium D doublet of W0707+1705 is deep and broad with an equivalent width of ≈ 1.2 nm and a FWHM for the merged lines of ≈ 1.4 nm. It shows a radial velocity < 100 km/s, as seen in Figure 6. The Ca II infrared triplet indicates $v_{rad} = -45 \pm 18$ /km/sec with respect to the Earth, or -21 km/sec with respect to the Sun. Figure 7 gives an expanded view of the Ca II infrared triplet.

The TiO band at 715 nm is much weaker than the 690 nm CaH band, and the telluric water band at 718.5 nm may provide all the structure seen in this region of the spectrum. Weak TiO to CaH is an indicator of an extreme or ultra-subdwarf, according to Lépine et al. (2007). Figure 8 compares the spectrum of W0707+1705 to the spectral standards in Lépine et al. (2007). In this region W0707+1705 resembles an usdM2 or usdM3.5 star. The usdM3.5 spectral standard is LHS 325, while the usdM2 standard is LHS 1691. Their resemblance to W0707+1705 includes similar spectral energy distributions: blackbody fits to the USNO B, 2MASS and WISE data give 3499 K for LHS 325 and 3184 for LHS 1691, which are similar to the fit for W0707+1705.

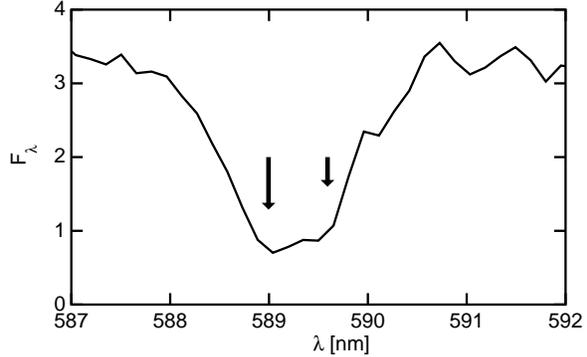


Fig. 6.— Full resolution blowup of the Na D region of the W0707+1705 spectrum. The arrows indicate the rest wavelengths of the doublet.

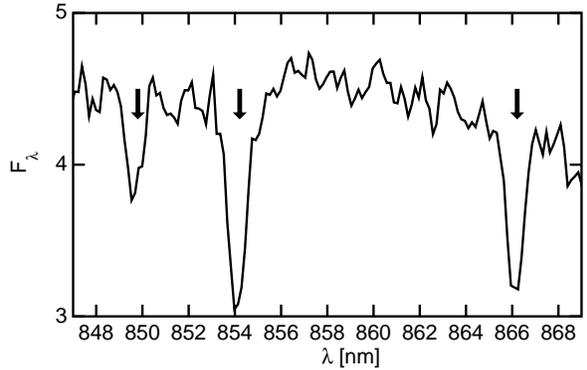


Fig. 7.— Expanded view of the Ca II IR triplet part of the spectrum.

The sequence from ordinary red dwarf (dM), through subdwarf (sdM), extreme subdwarf (esdM) to ultra subdwarf (usdM) is one of weakening metal oxide bands relative to metal hydride bands and atomic lines. Thus this is a sequence a decreasing metallicity, since to make a TiO molecule in a very low metallicity star requires two very rare atoms.

The near infrared CN and CO bands are not seen in a spectrum obtained on UT 2013 September 15 with the Triple-Spec instrument on the Apache Point Observatory 3.5 m telescope. The flux in the spectrum has been normalized using the 2MASS photometry. Figure 9 shows the spectrum divided by the blackbody fit in Figure 3. This is compared to a NASA IRTF SpeX prism spectrum of LHS 3768, which is tentatively classified as usdM3 based on the spectrum in Kirkpatrick et al.

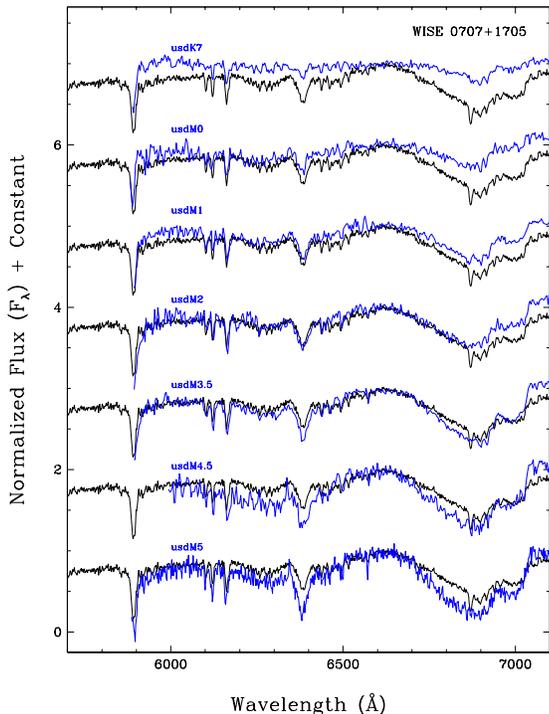


Fig. 8.— Spectrum of W0707+1705 compared to spectral standards from Lépine et al. (2007). All spectra are normalized at 660 nm.

(1995). These spectra are very similar, and neither shows any strong molecular bands. The CN band at $1.1 \mu\text{m}$ in the spectrum of the carbon dwarf G77-61 shown by Joyce (1998) is at least ten times stronger than in W0707+1705. The iron abundance of G77-61 is very low, 10^{-4} of solar, but both carbon and nitrogen are enhanced relative to iron by a factor of 400 (Plez & Cohen 2005). The weakness of any bands from carbon containing molecules in W0707+1705 suggests that carbon is depleted along with the metals in this source.

3. Discussion

The spectral standards LHS 325 and LHS 1691 do not have published precise parallaxes so it is difficult to estimate the distance and absolute magnitude of W0707+1705. But Bochanski et al. (2013) gives estimated absolute magnitudes for red dwarfs and subdwarfs seen by the SDSS (York et al. 2000). This paper defines a standard

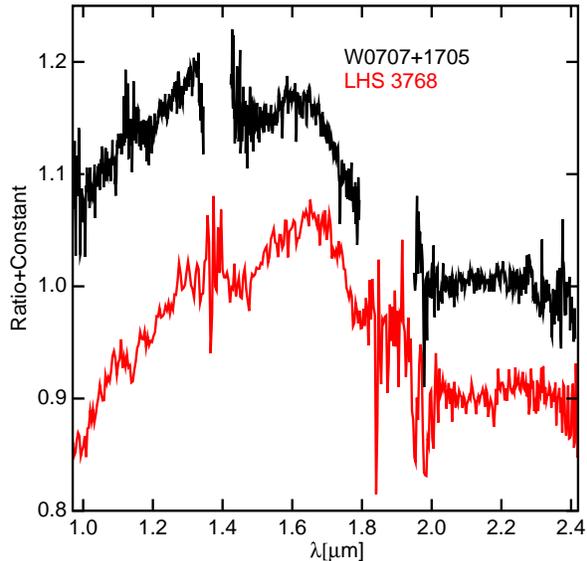


Fig. 9.— The Triple-Spec spectrum divided by the blackbody fit, compared to a SpeX spectrum of LHS 3768. The spectra are normalized at $2.2 \mu\text{m}$, and LHS 3768 is displaced downward by 0.1 for clarity.

color $(g-r)_{dM}$ for a red dwarf as a function of the $(r-z)$ color, and then uses the color excess $\delta(g-r) = (g-r) - (g-r)_{dM}$ as an indicator of subdwarfism. Bochanski et al. (2013) then gives fits for the absolute magnitudes $M_{r,i,z}$ as a function of $(r-z)$ and $\delta(g-r)$. Unfortunately we do not have SDSS *griz* photometry on W0707+1705, but we do have SDSS data on LHS 325, with $ugriz = 21.59, 17.96, 16.08, 15.26$ & 14.81 in DR10 (Ahn et al. 2013). These data imply $\delta(g-r) = 0.49$, appropriate for an ultra-subdwarf, and an absolute magnitude of $M_r = 12.883 \pm 0.41$ based on the polynomial fits in Bochanski et al. (2013). If W0707+1705 has the same absolute magnitude, then the apparent r magnitude (AB system) of $\approx 15.811 \pm 0.17$ from the CMC14 data implies a distance of 39 ± 9 pc and a tangential velocity of $v_{tan} = 330$ km/sec. This seems unlikely, given the small radial velocity, but it is not ruled out. The parallax determination is consistent with $d = 39$ pc, but has such a large uncertainty that it does not improve our knowledge of the distance, except to impose a lower limit, $d > 6$ pc.

Note that the proper motion of W0707+1705 is primarily an increasing galactic longitude

($\cos(b)dl/dt = 1705$ mas/yr, $db/dt = 555$ mas/yr), as would be expected for a source not rotating with the Galaxy, since its location is near the anti-center $[(l, b) = (199.29^\circ, 11.13^\circ)]$. The radial velocity of -21 km/sec with respect to the Sun gives UVW components $(19, 7, -4)$ km/sec, and the proper motion gives $(3.15d, -7.46d, +2.58d)$ /km/sec with d in pc, while the Sun's velocity with respect to the Local Standard of Rest (LSR) is $(9, 12, 7)$ km/sec (Binney & Tremaine 1987), so the total velocity with respect to the LSR is $(28 + 3.15d, 19 - 7.46d, 3 + 2.58d)$ km/sec.

If W0707+1705 is a member of a non-rotating singular isothermal sphere halo population, with $\sigma = v_c/\sqrt{2}$ where $v_c \approx 220$ km/sec is the local circular velocity, then the velocity distribution function is

$$\begin{aligned} f(U, V, W) &\propto \exp[-(U^2 + W^2 - (V + v_c)^2)/v_c^2] \\ &\propto \exp[-(1.2 - 0.070d + 0.00147d^2)]. \end{aligned} \quad (1)$$

The argument of the exponential is only -0.715 for $d = 39$ pc. If W0707+1705 is 1.1 mag less luminous than we have assumed, the argument of the exponential reaches its maximum value of -0.371 at $d = 24$ pc. This is equivalent to a Gaussian in d with a mean of 24 pc and a standard deviation of 18 pc.

But a singular isothermal sphere velocity distribution is an extreme assumption with an infinite escape velocity. Bochanski et al. (2013) also derive an ellipsoidal velocity distribution for usdM stars. For this distribution, the peak is at $d = 15$ pc and σ is 15 pc. This case is more in favor of W0707+1705 being 2 magnitudes more sub-luminous than an ultra-subdwarf star, but even the $d = 39$ pc case with $v_{tan} = 330$ km/sec is only 1.7σ out in the tail of this distribution.

For $d = 39$ pc W0707+1705 is in a retrograde orbit around the Milky Way, and its speed with respect to the center of mass of the galaxy is no more than the local rotational speed. Thus there is no indication that W0707+1705 is escaping from the Milky Way even if $v_{tan} \approx 330$ km/sec.

An accurate parallax would certainly be useful to distinguish between these scenarios for the luminosity and distance of W0707+1705. Improved optical photometry and spectroscopy would help to put W0707+1705 into the context of the many

faint red stars seen by large surveys such as the SDSS.

4. Conclusions

W0707+1705 shows the power of the upcoming AllWISE data release to reveal odd denizens of the Solar neighborhood. With an extremely short interval between epochs, the source association problems that have plagued previous searches for large proper motion objects are eliminated. Some examples based on discoveries within the last decade are: Teegarden's star (Teegarden et al. 2003) is recovered with AllWISE proper motion values of $\mu_\alpha = 3471 \pm 56$ & $\mu_\delta = -3666 \pm 25$ mas/yr; WD 1339-340 (Lépine et al. 2005) is recovered with $\mu_\alpha = -1916 \pm 78$ & $\mu_\delta = 999 \pm 84$ mas/yr; and Luhman's star (Luhman 2013) is recovered with $\mu_\alpha = -2646 \pm 26$ & $\mu_\delta = 408 \pm 25$ mas/yr. These stars are all easily visible on optical sky survey plates but images at different epochs were not correctly associated. Previously known higher proper motion stars like Barnard's star and Kapteyn's star are also easily recovered. See §II.6.d of the AllWISE Explanatory Supplement for more information on the AllWISE motion limits.

Based on sampling in test regions, we estimate that 2000-3000 real new proper motion detections will be found in the whole sky (Kirkpatrick et al. in preparation). In addition to proper motions, AllWISE will provide a deeper survey of the sky at 3.4 & 4.6 μm , and new tools to study the variability of sources all over the sky.

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