

The size and shape of the oblong dwarf planet Haumea

Alexandra C. Lockwood¹, Michael E. Brown¹ and John Stansberry²

alock@caltech.edu

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¹Division of Geological and Planetary Sciences, California Institute of Technology, Pasadena, CA 91125

²JWST/NIRCam Team, Space Telescope Science Institute, Baltimore, MD 21218

ABSTRACT

We use thermal radiometry and visible photometry to constrain the size, shape, and albedo of the large Kuiper belt object Haumea. The correlation between the visible and thermal photometry demonstrates that Haumea’s high amplitude and quickly varying optical light curve is indeed due to Haumea’s extreme shape, rather than large scale albedo variations. However, the well-sampled high precision visible data we present does require longitudinal surface heterogeneity to account for the shape of lightcurve. The thermal emission from Haumea is consistent with the expected Jacobi ellipsoid shape of a rapidly rotating body in hydrostatic equilibrium. The best Jacobi ellipsoid fit to the visible photometry implies a triaxial ellipsoid with axes of length 1920 x 1540 x 990 km and density 2.6 g cm^{-3} , as found by Lellouch et al. (2010). While the thermal and visible data cannot uniquely constrain the full non-spherical shape of Haumea, the match between the predicted and measured thermal flux for a dense Jacobi ellipsoid suggests that Haumea is indeed one of the densest objects in the Kuiper belt.

1. Introduction

Haumea, one of the largest bodies in the Kuiper Belt, is also one of the most intriguing objects in this distant population. Its rapid rotation rate, multiple satellites, and dynamically-related family members all suggest an early giant impact (Brown et al. 2007). Its surface spectrum reveals a nearly pure water ice surface (Barkume et al. 2006; Trujillo et al. 2007; Merlin et al. 2007; Dumas et al. 2011), with constraints on other organic compounds with upper bounds $< 8\%$ (Pinilla-Alonso et al. 2009). Lacerda et al. (2008) also find evidence for a dark spot on one side of the rotating body, make the surface albedo

non-uniform. Even more interesting is that shape modeling has suggested a density higher than nearly anything else known in the Kuiper belt and consistent with a body almost thoroughly dominated by rock (Rabinowitz et al. 2006; Lacerda et al. 2008; Lellouch et al. 2010). Lacerda & Jewitt (2007) concluded a density of 2.551 g cm^{-3} and follow up work found consistent values between 2.55 and 2.59 g cm^{-3} , depending on the model used. Such a rocky body with an icy exterior could be a product of initial differentiation before giant impact and subsequent removal of a significant amount of the icy mantle. Leinhardt et al. (2010) demonstrate that such an impact is possible in a graze and merge collision between equal size bodies, and are able to reproduce the properties of the Haumea family system.

Much of our attempt at understanding the history of Haumea relies on the estimate of the high density of the body. Haumea’s large-amplitude light curve and rapid rotation have been used to infer an elongated shape for the body. Assuming that the rotation axis lies in the same plane as the plane of the satellites, the amplitude of the light curve then gives a ratio of the surface areas along the major and minor axes of the bodies. If the further assumption is made that Haumea is large enough to be in hydrostatic equilibrium, the full shape can be uniquely inferred to be a Jacobi ellipsoid with fixed ratios of the three axes. From this shape and from the known rotation velocity, the density is precisely determined. Finally, with the mass of Haumea known from the dynamics of the two satellites (Ragozzine & Brown 2009), the full size and shape of Haumea is known (Rabinowitz et al. 2006; Lacerda et al. 2008; Lellouch et al. 2010).

The major assumption in this chain of reasoning is that Haumea is a figure of equilibrium. While it is true that a strengthless body can instantaneously have shapes very different from figures of equilibrium (Holsapple 2007), long-term deformation at the pressures obtained in a body this size should lead to essentially fluid behavior, at least at depth. Indeed, any non-spinning body large enough to become round due to self-gravity

has attained the appropriate figure of equilibrium. While the size at which this rounding occurs in the outer solar system is not well known, the asteroid Ceres, with a diameter of 900 km is essentially round (Millis et al. 1987; Thomas et al. 2005), while among the icy satellites, everything the size of Mimas and larger (~ 400 km) is essentially round. It thus seems reasonable to assume that a non-spinning Haumea, with a diameter of ~ 1240 km (Lellouch et al. 2013; Fornasier et al. 2013) would be round, thus a rapidly-spinning Haumea should likewise have a shape close to that of the Jacobi ellipsoid defined by Haumea’s density and spin rate.

Nonetheless, given the importance of understanding the interior structure of Haumea and the unusually high density inferred from these assumptions, we find it important to attempt independent size and density measurements of this object. Here we use unpublished photometric data from the Hubble Space Telescope (HST) to determine a best-fit Jacobi ellipsoid model. We then compare the predicted thermal flux from this best-fit Jacobi ellipsoid to thermal flux measured from the Spitzer Space Telescope (first presented in Stansberry et al. (2008)) and consider these constraints on the size and shape of the body.

2. Observations

Haumea was imaged on 2009 February 4 using the PC chip on the Wide Field/Planetary Camera 2 on HST. We obtained 68 100s exposures using the F606W filter, summarized in Table 1. The observations were obtained over 5 consecutive HST orbits, which provides a full sample of Haumea’s 3.9154 hr rotational lightcurve (Rabinowitz et al. 2006). With a pixel scale of ~ 5500 km at the distance of Haumea and semi-major axes of $\sim 50,000$ km and $\sim 25,000$ km for Hi’iaka and Namaka, respectively (Ragozzine & Brown 2009), this is the first published dataset where the object is resolved from its satellites, providing a pure lightcurve of the primary. During the observation both satellites are sufficiently spatially

separated from the primary that we are able to perform circular aperture photometry. Basic photometric calibrations are performed on the data including flat fielding, biasing, removing charge transfer efficiency effects, and identifying and removing hot pixels and cosmic rays¹. In five of our images the primary is contaminated by cosmic rays so we do not include these data. A 0.5" aperture is used to measure the object, and we apply an infinite aperture correction of 0.1 magnitudes (Holtzman et al. 1995). We present and model the data in the STMAG magnitude system, but a convolution of the F606W filter with the Johnson V filter shows a difference of approximately 0.1 magnitudes. This difference, along with the satellite flux contributions of 10% (Ragozzine & Brown 2009) included in previous photometry of the dwarf planet, indicate that the magnitude and amplitude of the lightcurve presented here ($\Delta m = 0.32$) is consistent with the findings of Rabinowitz et al. (2006) ($\Delta m = 0.28$) and Lacerda et al. (2008) ($\Delta m = 0.29$). The rotational period of Haumea is known sufficiently precisely that all of the observations are easily phased. We combine our data with that of Rabinowitz et al. (2006) and Lacerda et al. (2008) to get a 4-year baseline of observations and find a period of 3.91531 ± 0.00005 hours using phase dispersion minimization (Stellingwerf 1978). This period is consistent with that of Lellouch et al. (2010) ($P = 3.915341 \pm 0.000005$ h derived from a longer baseline), whose more precise solution we use to phase all observations, visible and infrared. With this period, there is a 60s uncertainty in the phasing over the 1.5 years between observations. We define a phase of 0 to be the point of absolute minimum brightness of Haumea and define a longitude system in which $\lambda = [(JD - 2454867.042)]$ modulo 360 degrees. This result is given in JD at Haumea, which is consistent with the phased data of Lacerda et al. (2008), and Lellouch et al. (2010) who instead quote their phasing in JD at the Earth. The

¹see the WFPC2 handbook at http://www.stsci.edu/instruments/wfpc2/Wfpc2_dhb/WFPC2_longdnhbcover.html

photometric results are shown in Fig. 1.

The thermal radiometry of Haumea was obtained 2007 July 13 – 19 using the 70 micron band of the MIPS instrument (Rieke et al. 2004) aboard the Spitzer Space Telescope (SST). Haumea’s lightcurve was unknown at the time, and the objective of the observations was simply to detect the thermal emission at a reasonably high signal to noise ratio (SNR). The data were collected as three 176 minute long observations, each of which is nearly as long as the (now known) lightcurve period. Stansberry et al. (2008) published the flux obtained by combining all three observations, as well as models indicating a diameter of about 1150 ± 175 km. Contemporaneously, Haumea’s lightcurve was published (Lacerda et al. 2008), a result that led us to consider re-analysing the Spitzer data to try and detect a thermal lightcurve. To that end, we split the original 176 minute exposures (each made up of many much shorter exposures) into 4 sub-observations, each 44 minutes long. The resulting data was processed using the MIPS Instrument Team pipeline (Gordon et al. 2005), resulting in flux calibrated mosaics for each of the 12 sub-observations. These individual observations are presented in Table 2. One of the points was obviously discrepant and was removed from the analysis. The data were also re-processed using improved (relative to the 2007 processing published in Stansberry et al. (2008)) knowledge of Haumea’s ephemeris. The reprocessing was undertaken as part of a project to reprocess all Spitzer/MIPS observations of TNOs (described in Mommert et al. (2012)) and is key to obtaining the highest SNR from the data. Had Haumea’s lightcurve been known at the time the Spitzer observations were planned, the observations probably would have been taken, for example, as a series of about 10 approximately 60 min exposures spaced about 4.1 rotations apart. The non-optimal observation plan may in part explain why the the uncertainties on the flux measurements appear to be somewhat optimistic, as discussed in more detail below.

The motion of Haumea was significant over the 6-day observing interval, so we were

able to make a clean image of the background sources (i.e. one without contamination from Haumea), and then subtract that sky image from our mosaics. The procedure used has been described previously, e.g. by Stansberry et al. (2008). We performed photometry on the sky-subtracted images, obtaining significantly smaller uncertainties than was possible using the original mosaics. The raw photometry was corrected for the size of the photometric aperture (15" radius). The signal-to-noise ratio of the resulting detections was about 7 in each of the 12 epochs. An additional calibration uncertainty of 6% should be systematically applied to the entire dataset. The thermal results are shown in Fig. 1. Where multiple observations are made at the same phase, these observations are averaged, and the uncertainty is taken from the standard deviation of the mean (or the full range if only two points go into the mean).

Though the uncertainties are large, the thermal light curve appears in-phase with the measured visible light curve. To robustly ascertain the detection of a thermal lightcurve, we look for a correlation between the thermal and visible datasets. We compare the mean visible flux during the phase of each 44 minute long thermal observation with the measured thermal flux during that observation (Fig. 2). A linear fit to the data suggests that a positive correlation between the optical and thermal brightness. As noted above, the deviation of the measurements from the model are larger than expected, so we assess the significance of the correlation between the optical and thermal data using the non-parametric Spearman rank correlation test Spearman (1904). With this test, we find that the two data sets are correlated at the 97% confidence level, and that the correlation is positive, that is: the observed optical and thermal flux increase and decrease in phase and there is only a 3% chance that this phase correlation is random. This positive correlation between the thermal and visible data sets indicates that we are viewing an elongated body and that the visible light curve must be caused – at least in part – by the geometric effects of this elongated body.

3. Photometric model

We begin with the assumption that Haumea is indeed a Jacobi ellipsoid whose shape is defined by its density and spin period. To find the Jacobi ellipsoid which best fits the photometric data, we model the expected surface reflection from an ellipsoid by creating a mesh of 4,000 triangular facets covering the triaxial ellipsoid and then determining the total visible light from the sum of the light reflected back toward the observer from each facet. Facets are approximately equal-sized equilateral triangles with length equal to 5 degrees of longitude along the largest circumference of the body. Mesh sizes a factor of 2 larger or smaller give identical results. For each facet we use a Hapke photometric model (Hapke 1993) to determine the reflectance as a function of emission angle. This model accounts for the effects of low phase angle observations, such as coherent backscattering and shadow hiding, and has been used to model the reflectance of many icy surfaces. For concreteness, we adopt parameters determined by Karkoschka (2001) for Ariel, a large satellite which exhibits deep water ice absorption and a high geometric albedo. We utilize the published values for the mean surface roughness ($\bar{\theta} = 23^\circ$), single-scattering albedo ($\varpi = 0.64$), asymmetry parameter ($g = -.28$), and magnitude ($S(0)$) and width (h) of the coherent backscattering and shadowing functions, $S(0)_{CB} = 4.0$, $h_{CB} = 0.001$, $S(0)_{SH} = 1.0$, and $h_{SH} = 0.025$ (Karkoschka 2001). While we have chosen Ariel because it is perhaps a good photometric analog to Haumea, we do note that within the range of Hapke parameters of icy objects throughout the solar system ($\varpi \sim 0.4 - 0.9$, $g = -.43 - -.17$, $\bar{\theta} = 10 - 36^\circ$), including the icy Galilean satellites and Triton (Buratti 1995; Hillier et al. 1990), the precise parameters chosen will affect only the geometric albedo and the beaming parameter, as discussed below.

The visible flux reflected from the body then becomes:

$$F_{vis} = p_{vis} \frac{F_{\odot,606}}{R_{AU}^2} A \frac{H \cos e}{\pi \Delta^2} \quad (1)$$

where p_{vis} is the visual albedo, $F_{\odot,606}$ is the solar luminosity over the bandwidth of the F606W filter, A is the projected surface area, H is the Hapke reflectance function, e is the angle of incidence, Δ is the geocentric distance to the body, and R is the heliocentric distance.

Our modeled body is rotated about the pole perpendicular to our line of sight – consistent with the hypothesis that the rotation pole is similar to the orbital pole of the satellites – and the photometric light curve is predicted. For such a model, the peak and trough of the visible lightcurve correspond to the largest and smallest cross-sectional areas of the body, and the ratio of the length of the largest non-rotational axis (a) to the smallest non-rotation axis (b) controls the magnitude of the photometric variation. With a uniform albedo across the surface of Haumea, however, no triaxial ellipsoid can fit the asymmetric observed lightcurve. We confirm the assertion of Lellouch et al. (2010) that the photometric variations of Haumea are caused primarily by shape and that surface albedo variations add a only minor modulation. In this approximation, the brightest peak and brightest trough of the data are assumed to be from essentially uniform albedo surfaces and are modeled to determine the ratio of the axes of the body. For plausible values of the dimension of the rotational axis (c), the measured peak-to-trough amplitude in the lightcurve of $\Delta m = 0.32 \pm 0.006$ is best modeled with an axis ratio of $b/a = 0.80 \pm 0.01$. Using the brightest trough and darkest peak instead, the axial ratio would be $b/a = 0.83$, but the surface heterogeneity necessary for this assertion is less likely than the single darker spot proposed here and observed by others (Lacerda et al. 2008).

Our measured value of 0.80 differs from previous measurements of b/a of .78 Rabinowitz et al. (2006) and .87 Lacerda et al. (2008) for a number of reasons. Our visible

dataset resolves Haumea from its satellites which results in a slightly deeper lightcurve. More to the point, including a realistic surface reflectance model changes that estimated shape significantly. Rabinowitz et al. (2006) do not actually model the shape of the body, while Lacerda et al. (2008) assume the surface to be uniformly smooth, giving a value for b/a that is too high. More recently, Lellouch et al. (2010) confirm a more elongated body ($b/a = .80$), after having tested two different models, including that of Lacerda et al. (2008).

With the ratio of the axes fixed, we now find the simplest surface normal albedo model consistent with the data. We divide the surface longitudinally into 8 slices and allow the albedo to vary independently between sections to account for the possible hemispherical variation apparent in our data and explored by others (Lacerda et al. 2008; Lacerda 2009). Dividing the surface further does not significantly improve the fit. The middle panel of Figure 1 shows how the geometric albedo varies across the surface of the body.

Assuming that Haumea is indeed a Jacobi ellipsoid, the ratio $b/a = 0.80$ combined with the rotation period uniquely defines $c/a = 0.517$ and a density of 2.6 g cm^{-3} . Combining these parameters with the known mass of Haumea from Ragozzine & Brown (2009) implies $a = 960 \text{ km}$, $b = 770 \text{ km}$, and $c = 495 \text{ km}$. These radii agree with the ones obtained by Lellouch et al. (2010) using a Lommel-Seelinger reflectance function.

4. Thermal model

To calculate the thermal emission from our shape and albedo model we first determine the temperature of each facet of the body. Due to Haumea’s rapid rotation, the temperature of any given face of the surface does not have time to equilibrate with the instantaneous incoming insolation. Instead we calculate the average amount of sunlight received by a facet during a full rotation, which is only dependent on the angle between rotational pole

and the facet normal. Although we implement a standard thermal model, the rapidly spinning object gives rise to surface temperatures indicative of an isothermal latitude model (Stansberry et al. 2008), which we precisely calculate for the non-spherical geometry of the object. Indeed the agreement between the visible and infrared datasets as seen in Figure 2 supports the hypothesis of a body with negligible thermal conductivity.

We model the average amount of sunlight absorbed for each facet by multiplying the geometric albedo of each facet by an effective phase integral q and averaging over a full rotation. The average facet phase integral is a mild function of the shape of the body, but for simplicity we simply adopt a value of $q = 0.8$, as used by Stansberry et al. (2008) for large, bright KBOs. In fact, the precise value used has little impact on our final results.

If we assume the surface is in thermal equilibrium, the temperature of each facet, is determined by balancing this absorbed sunlight with the emitted thermal radiation. We choose a typical thermal emissivity of 0.9 and invoke a beaming parameter, η , which is a simple correction to the total amount of energy radiated in the sunward direction, usually assumed to be caused by surface roughness, but which can be taken as a generic correction factor to the assumed temperature distribution. For asteroids of known sizes, Lebofsky et al. (1986) found η to be approximately 0.756, a correction which agrees well with measurements of icy satellites in the outer solar system (Brown et al. 1982a,b). The beaming parameter value range for Trans-Neptunian objects is fully described in Lellouch et al. (2010), who found η of 1.15-1.35 for Haumea, with hemispherical variations consistent with a much lower value ($\eta \sim .4-.5$). The MIPS and PACS fluxes values presented in Lellouch et al. (2010) were updated and presented in Fornasier et al. (2013), who, also using the SPIRE data, found a beaming factor of $0.95^{+0.33}_{-0.26}$ with a NEATM model. We leave this as a free parameter in our modeling but consider an inclusive range.

For each rotational angle of our model, we predict the total thermal flux by calculating

the blackbody spectrum from each visible facet and integrating this flux in the full 70 μm band pass of MIPS, according to

$$F_{IR} = \frac{A \cos e}{\pi^2 \Delta^2} \epsilon \int B_\lambda(T(\theta, \phi)) \sin \theta d\theta d\phi \quad (2)$$

where A is the cross-sectional surface area, $\epsilon = 0.9$ is the emissivity, B is the Planck function and T is the temperature at each piece of the surface. Assuming a solar flux at 70 μm of S at the distance of Haumea, R , then T is calculated for an edge-on rotating body and

$$T = \left[\frac{S * (1 - qp_{vis})}{\epsilon \sigma \eta R^2} \right]^{1/4} \quad (3)$$

Figure 1 shows the measured Spitzer flux along with the flux predicted from a model for the theoretical Jacobi ellipsoid with $a = 960$ km and $a : b : c = 1.00 : 0.80 : 0.52$ and a thermal beaming parameter of $\eta = 0.76$. The best fit is obtained by assuming $\eta = 0.89$, but values of η between 0.82 and 0.97 are within the $1\text{-}\sigma$ error limits. The larger values measured by Lellouch et al. (2010) and quoted for the majority of Kuiper Belt objects (~ 1.2 by Stansberry et al. (2008) and $\sim 1\text{-}2.5$ by Lellouch et al. (2013)) are consistent with our result if we consider the difference of thermal models employed. Stansberry et al. (2008) explain that the surface temperature difference between an isothermal latitude model (that is used in this work) and a standard thermal model (used by Lacerda (2009) and Lellouch et al. (2010) is simply a factor of $\pi^{-\frac{1}{4}}$. If this value is incorporated into the beaming factor, the inconsistency between the models and resulting beaming parameters is resolved. Remarkably, the shape and albedo model constructed from only photometric observations and the assumption of fluid equilibrium provides an acceptable prediction of the total thermal flux at 70 μm and its rotational variation.

5. Discussion

The thermal and photometric light curves of Haumea are consistent with the assumption that Haumea is a fluidly relaxed, rapidly rotating Jacobi ellipsoid with a density of 2.6 g cm^{-3} and minor albedo variation across its surface. Although we allow the albedo to change longitudinally, we have demonstrated that the reason for the double-peaked lightcurve of Haumea is in fact a shape effect. This work presents new data and an informed Hapke model that agree with the findings of previous authors (Lellouch et al. 2010; Lacerda et al. 2008; Rabinowitz et al. 2006).

The lightcurve in several colors (Lacerda et al. 2008; Lacerda 2009) indicates that the albedo variation is concentrated in a large spot on one side of the body, an argument supported by our allowed albedo variation. The precise geometric albedo presented here is less than that of Fornasier et al. (2013), but with a small variation in either the single-scattering albedo or the asymmetry parameter, we easily find agreement between the two values. We do not focus on the absolute value of the albedo here, as it is closely tied to the unknown Hapke parameters for the surface. The only affect this has on the thermal fit is to change the best fit beaming parameter, which does not differ by more than $1\text{-}\sigma$ from the reported value. The important point here, however, is that our precise visible data set agrees with the dark spot proposed by Lacerda et al. (2008).

There are still a number of questions remaining regarding this KBO. The temperature of the dark spot is still uncertain and our data are not precise enough to constrain it, although the lower albedo we use to match the visible data agree well with a warmer region. Another assumption used here that could be disputed is the rotation axis of the body as perpendicular to our line of sight. However, assuming that the majority of the depth of the light curve is from the major axes of the body, the pole position will only affect the size of the third dimension of the body and the thermal beaming factor, which become somewhat

degenerate when fitting the thermal data anyway. If the body were not mostly edge-on, it would be difficult to explain the particular surface patterning needed to recreate the visible lightcurve.

While it is encouraging that a dense Jacobi ellipsoid fits both data sets, it is important to point out that we can only obtain a unique solution under the assumption that the object has a prescribed shape. If Haumea is modeled as an arbitrary triaxial ellipsoid rather than a Jacobi ellipsoid, large families of solutions are possible. To first order, the photometric data constrain the ratio b/a , while the thermal flux is roughly proportional to the emitting surface area which is proportional to ac and bc . As long as the ratio of b/a is kept constant, however, equally good fits can be obtained from very elongated ellipsoids with a very short rotation axis, or from only moderately elongated ellipsoids with a large rotation axis, as long as the value ac is approximately constant. For this unconstrained problem, densities anywhere between 1 g cm^{-3} (for very elongated objects) and 3 g cm^{-3} for compact objects are compatible with both data sets. While such large deviations from an equilibrium shape appear implausible, the thermal data alone cannot rule them out.

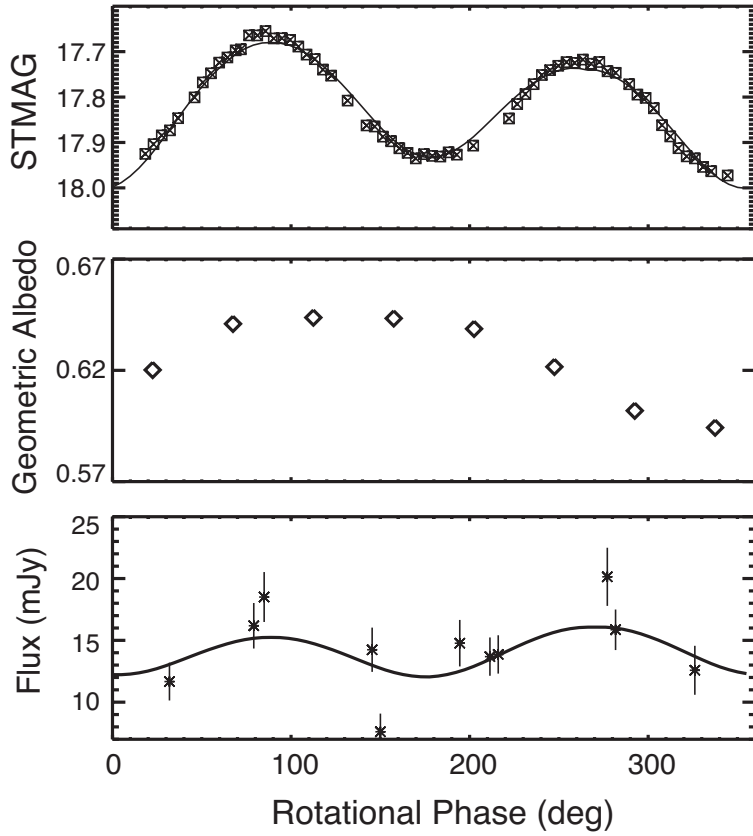


Fig. 1.— The visible lightcurve, geometric albedo, and thermal lightcurve plotted over one rotation. The error bars in the top two panels are smaller than the size of the plotted point. The visible photometry are fit with a Jacobi ellipsoid of dimensions 1920 x 1540 x 990 km with the modest longitudinal variation in reflectance shown. This ellipsoid model provides a quite good fit to the 70 μm thermal data from Spitzer.

Table 1: HST F606W flux of Haumea

JD	phase	F606W	uncertainty	JD	phase	F606W	uncertainty
-2450000	(deg)	mag		-2450000	(deg)	mag	
4866.943802	142.27	17.863	0.004	4867.0898	104.08	17.689	0.003
4866.945802	146.87	17.865	0.004	4867.0918	108.68	17.707	0.003
4866.948202	151.47	17.887	0.004	4867.0937	113.28	17.717	0.003
4866.950202	156.06	17.897	0.004	4867.0962	117.88	17.740	0.003
4866.952102	160.66	17.913	0.004	4867.0981	122.47	17.753	0.003
4866.954602	165.26	17.923	0.004	4867.1025	131.67	17.808	0.004
4866.956502	169.86	17.935	0.004	4867.1435	222.08	17.847	0.004
4866.958502	174.45	17.926	0.004	4867.1455	226.68	17.815	0.004
4866.960402	179.05	17.930	0.004	4867.1474	231.28	17.793	0.004
4866.962902	183.65	17.931	0.004	4867.1494	235.87	17.772	0.004
4866.964802	188.24	17.922	0.004	4867.1518	240.47	17.751	0.003
4866.966802	192.84	17.927	0.004	4867.1538	245.07	17.741	0.003
4866.971202	202.04	17.907	0.004	4867.1557	249.66	17.731	0.003
4867.010702	289.39	17.772	0.004	4867.1577	254.26	17.723	0.003
4867.012702	293.98	17.795	0.004	4867.1601	258.86	17.725	0.003
4867.014602	298.58	17.802	0.004	4867.1621	263.46	17.718	0.003
4867.017102	303.18	17.825	0.004	4867.1640	268.05	17.727	0.003
4867.019002	307.77	17.862	0.004	4867.1660	272.65	17.723	0.003
4867.021002	312.37	17.887	0.004	4867.1684	277.25	17.743	0.003
4867.022902	316.97	17.913	0.004	4867.1704	281.85	17.747	0.003
4867.025402	321.57	17.931	0.004	4867.2143	18.39	17.925	0.004
4867.027302	326.16	17.935	0.004	4867.2163	22.99	17.904	0.004
4867.029302	330.76	17.954	0.004	4867.2182	27.58	17.884	0.004
4867.031202	335.36	17.963	0.004	4867.2202	32.18	17.873	0.004
4867.035602	344.55	17.973	0.004	4867.2226	36.78	17.847	0.004
4867.075202	71.90	17.695	0.003	4867.2265	45.97	17.801	0.004
4867.077102	76.50	17.664	0.003	4867.2285	50.57	17.768	0.004

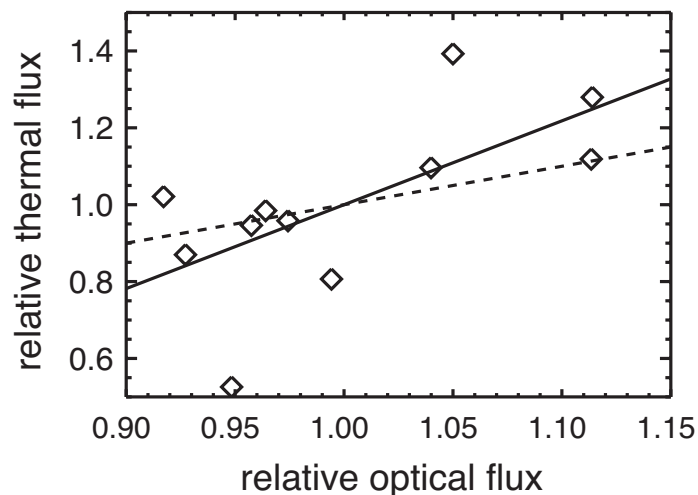


Fig. 2.— The correlation of the relative optical and thermal flux. The dashed line shows a one-to-one correlation while the solid line shows the best fit. A rank correlation test shows that the two distributions are correlated at the 97% confidence level. The in-phase thermal light curve of Haumea demonstrates that it is an elongated body.

Table 2: Thermal Flux of Haumea

JD - 2450000	phase (deg)	MIPS 70 μm flux (mJy)	uncertainty (mJy)
4294.6770	194.5	14.77	1.88
4294.7367	326.1	12.58	1.98
4294.7664	31.9	11.67	1.53
4297.4004	84.9	18.51	2.02
4297.4302	150.0	7.61	1.47
4297.4601	216.0	13.86	1.55
4297.4899	281.8	15.86	1.64
4300.4980	79.1	16.18	1.84
4300.5277	145.4	14.24	1.79
4300.5576	211.4	13.69	1.55
4300.5874	277.1	20.14	2.35

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