

HYDROCARBON PHOTOCHEMISTRY AND LYMAN ALPHA ALBEDO OF JUPITER¹

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ABSTRACT

A combined study of hydrocarbon and atomic hydrogen photochemistry is made to calculate self-consistently the $L\alpha$ albedo of Jupiter. It is shown that the $L\alpha$ emissions observed by *Voyagers I* and *II* can be explained by resonance scattering of sunlight. Precipitation of energetic particles from the magnetosphere can provide the large required source of atomic hydrogen, although the contribution of direct particle excitation to the disk-averaged brightness is insignificant. The variability of the $L\alpha$ brightness inferred from many observations in recent years is examined. The large difference in the brightness of the He 584 Å resonance line observed by *Pioneer* and *Voyager* is briefly discussed. Driving the photochemistry by solar ultraviolet radiation alone yields a maximum mixing ratio of $C_2H_6 + C_2H_2$ at 10^{-2} atm of about 4×10^{-6} . The possibility of additional CH_4 dissociation from precipitation of magnetospheric particles is discussed. The photochemistry of C_2H_2 and C_2H_3 is sufficiently uncertain not to permit accurate calculations of their densities and the ratio C_2H_6/C_2H_2 .

Subject headings: molecular processes — planets: abundances — planets: Jupiter

1. INTRODUCTION

The $L\alpha$ brightness of Jupiter has been measured by a number of workers (Moos and Fastie 1969; Rottman, Mount, and Freer 1973; Giles, Moos, and McKinney 1976; Carlson and Judge 1974; Bertaux *et al.* 1979; Broadfoot *et al.* 1979; Clarke *et al.* 1979; see Table 1). All observations previous to 1976 were summarized and discussed by Giles, Moos, and McKinney (1976). More recent observations were compared by Bertaux *et al.* (1979). The most interesting question is why the brightness is so variable. Analysis by Carlson and Judge (1971) and Wallace and Hunten (1973) have shown that the $L\alpha$ brightness of Jupiter is primarily due to resonance scattering of incident sunlight by atomic hydrogen. The $L\alpha$ albedo depends on at least the following: (a) production of H atoms from H_2 and CH_4 dissociation, (b) vertical eddy mixing in the atmosphere, (c) photochemistry of the hydrocarbons, and (d) thermal structure of the atmosphere. Each of the above factors has been considered in some approximate manner in the previous models, but never all at once on a self-consistent basis.

The photochemistry of hydrocarbons in the atmosphere of Jupiter has been studied by one of the authors (Strobel 1969, 1973, 1974) whose prediction of abundant C_2H_6 and C_2H_2 concentrations has been

borne out by recent observations (Ridgway 1974; Combes *et al.* 1974; Tokunaga, Knacke, and Owen 1976; Orton and Aumann 1977; Hanel *et al.* 1979). A number of recent advances justify another examination of the hydrocarbon photochemistry in the Jovian atmosphere, in addition to the requirement of a self-consistent calculation of the $L\alpha$ albedo. Calculations of the chemical composition of the deep atmosphere by Barshay and Lewis (1978) enable us now to select a more appropriate boundary condition for C_2H_6 . There is strong evidence that the C/H ratio for Jupiter must be significantly higher than the solar value, which was used in previous photochemical calculations (Wallace and Hunten 1978; Sato and Hansen 1979). In addition, there have been important revisions in CH_4 and C_2H_6 photoabsorption cross sections (Mount, Warden, and Moos 1977; Mount and Moos 1978) and in the rate coefficients for three-body reactions $H + CH_3 + M \rightarrow CH_4 + M$ and $CH_3 + CH_3 + M \rightarrow C_2H_6 + M$ at low pressure (Van den Bergh 1976; Troe 1977). The detection of extensive polar emissions by Broadfoot *et al.* (1979) suggests that a large flux of energetic particles precipitates into the upper atmosphere. This magnetospheric interaction can provide a large source of atomic hydrogen and possibly hydrocarbons. In this paper we assess the possible importance of this interaction on the $L\alpha$ albedo and hydrocarbon abundances observed by *Voyager*.

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TABLE 1
OBSERVATION OF $L\alpha$ BRIGHTNESS OF JUPITER^a

Date Observed	$L\alpha$ Brightness ($4\pi I$ [kR])	πF ($10^{11} \text{ cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$)	$4I/F_J$ (\AA)	References
1967 Dec. 5.....	4.0	3.2	0.34	Moos and Fastie 1969
1972 Jan. 25.....	4.4	3.2	0.37	Rottman <i>et al.</i> 1973
1972 Sept. 1.....	2.2	3.0	0.20	Giles <i>et al.</i> 1976
1973 Dec. 1.....	0.4	2.4	0.045	Carlson and Judge 1974
1976 Jan. 5–Sept. 22.....	2.8–5.3	2.13	0.35–0.67	Bertaux <i>et al.</i> 1979
1979 Jan. 31.....	14	5.1	0.74	Broadfoot <i>et al.</i> 1979
1978 Dec. 1.....	13	5.1	0.69	Clarke <i>et al.</i> 1979

^a All brightness intensities refer to disk-averaged values in kR (10^9 photons $\text{cm}^{-2} \text{ s}^{-1}$). Respectively, πF and πF_J denote the solar $L\alpha$ fluxes at 1AU and 5.2AU. Bertaux *et al.* 1979 used simultaneous measurements of the $L\alpha$ emission of the geocorona to derive the albedo of the planet. For the observations of Broadfoot *et al.* 1979 and Clarke *et al.* 1979 we used the solar flux measured by Rottman and Mount 1980 on 1979 June 5. All other values of the solar flux in the table are deduced on the basis of the Sun's 10.7 cm flux and Zurich sunspot number, and application of Vidal-Madjar's 1975 formula.

II. PHOTOCHEMICAL MODEL

The details of the photochemical model were presented in Tables 2–5 in Strobel (1973), whose notation we follow. The standard model atmosphere, model A, is shown in Figure 1. (We also consider a colder isothermal atmosphere, model B, with temperature equal to 150 K, as described in Strobel 1973.) Altitude 0 km corresponds to a neutral number density of $1 \times 10^{19} \text{ cm}^{-3}$. The model assumes a bulk He/H₂ ratio equal to 0.11 by volume. We adopt a CH₄/H₂ mixing ratio of 1.7×10^{-3} , as required to interpret the visible and near-infrared CH₄ bands of Jupiter (Sato and Hansen 1979). The temperature at $z = 0$ km is 110 K and varies linearly with z to 385 K at $z = 500$ km. This choice is consistent with the *Voyager* IR results in the 0–150 km region (Hanel *et al.* 1979). Above this height, the temperature rises to an asymptotic value of 1200 K at ~ 1500 km. The vertical structure of the upper atmosphere is described by the simple analytic representation of Bates and Patterson (1961). In a one-dimensional photochemical model vertical transport for long-lived species is conveniently described by eddy

diffusion. The magnitude of the eddy diffusion coefficients $K(z)$ in the lower stratosphere is constrained by the requirement that it be compatible with the NH₃ and PH₃ photochemical destruction rate (Strobel 1977) and the CO abundance in the lower stratosphere (Strobel and Yung 1979). In the upper stratosphere, estimates of mixing rates can be derived from the $L\alpha$ albedo of the planet (Wallace and Hunten 1973; Carlson and Judge 1976), if solar radiation determines the production rate of H. However, in view of the large variability of the observed $L\alpha$ brightness, other factors may contribute. Eddy mixing in the upper atmosphere is considered an adjustable parameter in this work. The eddy diffusivity profile is expressed as

$$K(z) = \frac{K_1(z)K_2}{K_1(z) + K_2}, \quad (1)$$

where

$$K_1(z) = K_0 \left(\frac{n(0)}{n(z)} \right)^\gamma, \quad (2)$$

$n(z)$ is the number density of the atmosphere at z , and γ is a dimensionless number of order unity which measures the rate of increase of $K_1(z)$ with altitude. Realistic choices of the constants K_0 and K_2 for the Jovian atmosphere require $K_2 \gg K_0$. Hence, in the lower atmosphere $K(z) \approx K_1(z)$. In the upper atmosphere $K_1(z) \gg K_2$, and in this limit $K(z) \approx K_2$. The functional form of $K(z)$ was chosen to ensure continuous derivatives everywhere and to limit $K(z)$ as $z \rightarrow \infty$ to be much less than the molecular diffusion coefficients.

The major differences between the previous and the current models are summarized in Table 2; they represent the most important updates to the photochemical model since 1974. The adoption of the extremely slow reaction rate for R10 creates a problem as to the fate of C₂H₃ radicals in Jupiter's atmosphere. Since reaction R9 is much faster than R10 at all altitudes,

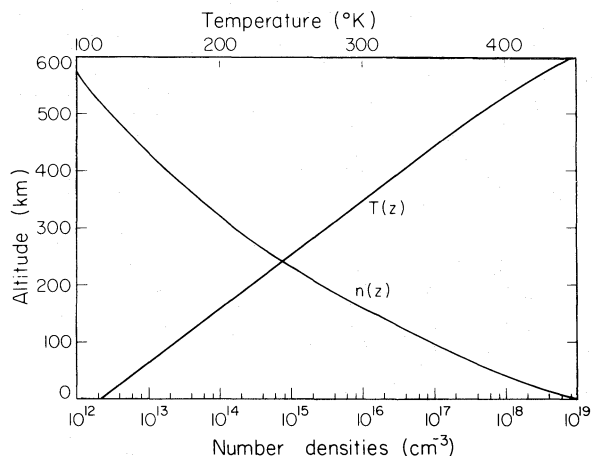
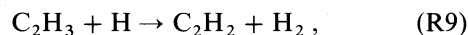


FIG. 1.—Number density and temperature profile for model A

TABLE 2
MAJOR DIFFERENCES IN PHOTOCHEMICAL MODEL BETWEEN THIS STUDY AND STROBEL'S STUDY^a

Model	Major Changes	References
C/H ratio assumed for the planet	3.4×10^{-3}	Sato and Hansen 1979
Lower boundary conditions	$\frac{d}{dz} \ln [x] = \frac{\gamma}{H_{av}} (1 + \delta)$ $x = C_2H_2, C_2H_6$	Barshay and Lewis 1978
R1: $CH_4 + hv \rightarrow$ products		See text
R4: $C_2H_6 + hv \rightarrow$ products		Mount <i>et al.</i> 1977
R5: $CH_3 + H + M \rightarrow CH_4 + M$	$k_5 = \text{Min} (1 \times 10^{-10}, 1.3 \times 10^{-23} T^{-2} M)$	Mount and Moos 1978
R6: $CH_3 + CH_3 + M \rightarrow C_2H_6 + M$	$k_6 = \text{Min} (5.5 \times 10^{-11}, 1.7 \times 10^{-17} T^{-2.3} M)$	Teng and Jones 1972; Troe 1977
R8: $C_2H_2 + H + M \rightarrow C_2H_3 + M$	$k_8 = \text{Min} (9.2 \times 10^{-12} e^{-1205/T}, 10^{-30} e^{-770/T})$	Van den Bergh 1976; Troe 1977
R10: $C_2H_3 + H_2 \rightarrow C_2H_4 + H$	$k_{10} = 3 \times 10^{-13} e^{-5570/T}$	Payne and Stief 1976
R11: $C_2H_4 + H + M \rightarrow C_2H_5 + M$	$k_{11} = \text{Min} (3.7 \times 10^{-11} e^{-1040/T}, 3 \times 10^{-30} M)$	W. Tsang, private communication 1979
R12: $C_2H_5 + H \rightarrow CH_3 + CH_3$	$k_{12} = \text{Min} (1.8 \times 10^{-10} e^{-435/T}, 1 \times 10^{-20} T^{-3.3} M)$	Lee <i>et al.</i> 1978
R13: $^3CH_2 + CH_3 \rightarrow C_2H_4 + H$	$k_{13} = 7 \times 10^{-11}$	Teng and Jones 1972
		Pilling and Robertson 1975; Laufer and Bass 1975

^a The notation for reactions follows that in Tables 3 and 4 in Strobel 1973.

the catalytic cycle of R8 + R9 is essentially unbroken and no significant conversion of C_2H_2 to C_2H_4 results (cf. Strobel 1973). We did not include a reaction suggested by Prasad, Capone, and Schneck (1975),



The heats of formation of C_2H_3 and C_4H_4 are not accurately enough known to determine whether this reaction is exothermic or endothermic. Even if this reaction were exothermic, it is estimated to be slow at Jovian temperatures, since its preexponential factor is $\sim 10^{-12} \text{ cm}^{-3} \text{ s}^{-1}$ with an activation energy of $\sim 2 \text{ kcal mole}^{-1}$ (Benson and Haugen 1967). We should also point out that Prasad, Capone, and Schneck (1975) did not treat the photochemistry of C_4H_4 produced in reaction (3) to determine whether it is recycled back to C_2H_2 .

The photolysis of C_2H_2 is not completely understood, but Payne and Stief (1976) estimate that the quantum yield for C_2H is ~ 0.1 . The subsequent chemistry of C_2H has recently been investigated by Laufer and Bass (1979). From their laboratory results, C_2H should react preferentially with H_2 in the stratosphere to form C_2H_2 and H . The net result of C_2H_2 photolysis is thus destruction of H_2 . Our conclusion is that there is insufficient kinetic data available to adequately treat the chemistry of C_2H_2 and C_2H_3 in Jupiter's atmosphere. The absence of significant loss processes for C_2H_2 in our model leads to overestimates of its density when compared to observations. Rather than adopting a speculative loss process for the model, we chose to calculate the C_2H_2 density without any loss processes and caution the reader to interpret the resultant C_2H_2 densities as C_2H_2 plus other unknown species, e.g., C_4H_2 , C_4H_4 , additional C_2H_4 , or even Danielson dust.²

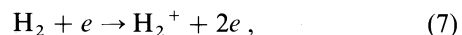
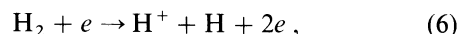
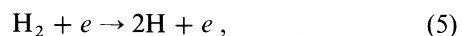
² At the suggestion of D. M. Hunten at the 1979 DPS meeting in St. Louis, we shall use the term Danielson dust in the place of Axel dust (Axel 1972).

The C_2H_6 mixing ratio in the deep Jovian atmosphere is predicted by Barshay and Lewis (1978) to be less than 10^{-8} . As a consequence, we anticipate a large flux of C_2H_6 from the stratosphere, where its mixing ratio could be as high as 10^{-5} , to the deep atmosphere where it undergoes pyrolysis. The appropriate lower boundary condition in our model is (cf. Strobel 1975)

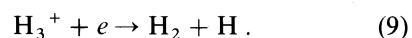
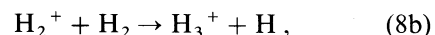
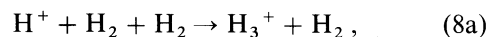
$$\frac{d}{dz} \ln [C_2H_6] = \frac{\gamma}{H_{av}} (1 + \delta), \quad (4)$$

where H_{av} is the scale height of the atmosphere, γ is a dimensionless parameter defined by equation (2), and δ is a small correction factor proportional to the vertical gradient of temperature. The same boundary condition also applies to C_2H_2 .

The primary sources of atomic hydrogen in the upper atmosphere are methane dissociation R1 and processes involving the EUV solar flux. The latter source has been estimated by Strobel (1973) to provide a mean flux of $\phi_0 = 7 \times 10^8 \text{ atoms cm}^{-2} \text{ s}^{-1}$ for solar activity consistent with Hinteregger's (1970) fluxes. There are additional sources of H , through reactions associated with energetic electrons of magnetospheric origin, such as

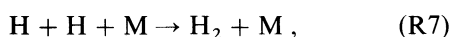


followed by

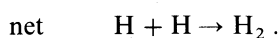
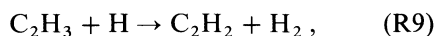


Detailed modeling (Gladstone and Yung 1979) of the auroral observations by Broadfoot *et al.* (1979) suggests that the globally averaged flux of hydrogen atoms produced in the auroras could greatly exceed ϕ_0 .

Hydrogen atoms are destroyed mainly by

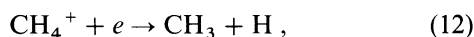
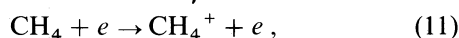
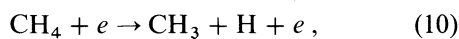


and the C_2H_2 catalytic cycle

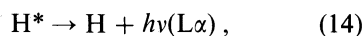


R5 is the most important reaction for destruction of hydrogen. R8 and R9 become important below the $\tau(L\alpha) = 1$ level for CH_4 . R7 limits the buildup of excessively large number densities of hydrogen.

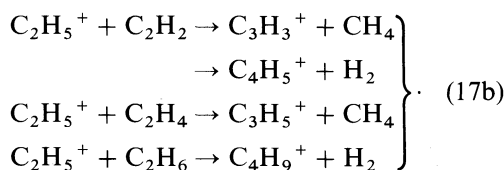
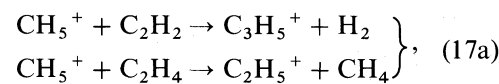
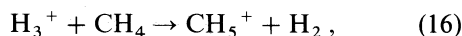
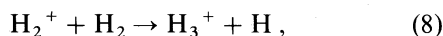
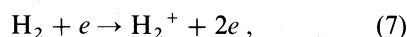
The effect of energetic particles on the hydrocarbons is less clear. If the energy of the incident electrons exceeds 10^5 eV, most of the particles are stopped below the homopause, where they can dissociate CH_4 . Direct reactions, such as



are not important since $\text{C}/\text{H} \sim 10^{-3}$. Secondary processes via a $L\alpha$ photon could become marginally important:



Probably most important is the following sequence of reactions based on Munson and Field (1969) and Huntress (1977):



Recombination of these complex hydrocarbon ions will lead to the formation of a variety of heavier hydrocarbons. Until further laboratory kinetics studies become available, we can only consider this scheme for additional destruction of CH_4 and subsequent production of C_2H_6 and C_2H_2 in a qualitative manner:



III. RADIATIVE TRANSFER MODEL

The basic atomic and molecular physics parameters that describe multiple scattering of the $L\alpha$ photons in the Jovian atmosphere are taken from Carlson and Judge (1971) and Wallace and Hunten (1973) and summarized in Table 4. The cross section for resonance scattering of $L\alpha$ by atomic hydrogen in standard notation is

$$\sigma(v - v_0) = \sigma_0 \sqrt{\pi} \Phi \left(\frac{v - v_0}{\Delta v_D} \right), \quad (18)$$

where

$$\Phi(x) = \frac{\alpha}{\pi^{3/2}} \int_{-\infty}^{\infty} du \frac{e^{-u^2}}{(x-u)^2 + \alpha^2}, \quad (19)$$

$$\sigma_0 = \frac{\pi e^2}{mc} \frac{f}{\sqrt{\pi} \Delta v_D}, \quad (20)$$

$$\alpha = \frac{\Delta v_N}{\Delta v_D}. \quad (21)$$

Note that

$$\lim_{x \rightarrow 0} \Phi(x) = \frac{1}{\sqrt{\pi}} e^{-x^2}, \quad (22)$$

$$\lim_{x \rightarrow \infty} \Phi(x) = \frac{\alpha}{\pi} \frac{1}{1 + x^2}, \quad (23)$$

and

$$\int_{-\infty}^{\infty} dv \sigma(v - v_0) = \frac{\pi e^2}{mc} f. \quad (24)$$

The most important absorber of a $L\alpha$ photon is methane, whose cross section is uniform over the width of the solar line.

With the prescribed thermal profile and calculated densities of H and CH_4 from the photochemical model, we construct the scattering parameters, the optical depth τ , and the single scattering albedo $\tilde{\omega}_0(\tau)$. The incident solar line is divided uniformly into 156 steps to achieve adequate resolution. The $L\alpha$ albedo at each wavelength interval is calculated using a highly accurate (better than 10^{-3}) invariant embedding algorithm for inhomogeneous atmospheres (Sato, Kawabata, and Hansen 1977). For nearly conservative cases we choose the lower boundary at optical depth = 100, a choice that should simulate an infinite scattering atmosphere to better than 1%, based on comparisons with results obtained using H functions. We assume that scattering of $L\alpha$ photons can be approximated by the Rayleigh phase function³ without complete frequency redistribution. The neglect of polarization and frequency redistribution should not introduce more than 10% total error (Hansen and Travis 1974; Wallace 1971).

³ The phase function is nearly isotropic, but the difference between the Rayleigh and isotropic phase functions are small for albedo calculations.

TABLE 3
SUMMARY OF INPUT ASSUMPTIONS FOR MODEL RUNS^a

Model	Model Atmosphere	K_0	K_2	φ_H	Solar L α Flux
A11	A	1 (3)	1 (5)	1	1
A12	A	1 (3)	1 (6)	1	1
A13	A	1 (3)	1 (7)	1	1
A14	A	1 (3)	1 (8)	1	1
A15	A	1 (3)	1 (9)	1	1
A21	A	3 (2)	1 (5)	1	1
A22	A	3 (2)	1 (7)	1	1
A23	A	3 (2)	1 (9)	1	1
B11	B	1 (3)	1 (5)	1	1
B12	B	1 (3)	1 (6)	1	1
B13	B	1 (3)	1 (7)	1	1
B14	B	1 (3)	1 (8)	1	1
B15	B	1 (3)	1 (9)	1	1
C11	A	1 (3)	1 (5)	10	1
C12	A	1 (3)	1 (6)	10	1
C13	A	1 (3)	1 (7)	10	1
C14	A	1 (3)	1 (8)	10	1
C15	A	1 (9)	1 (9)	10	1
C21	A	1 (6)	1 (6)	30	1
C22	A	1 (7)	1 (7)	30	1
D11	A	1 (3)	1 (5)	1	3
D12	A	1 (3)	1 (6)	1	3
D13	A	1 (3)	1 (7)	1	3
D14	A	1 (3)	1 (8)	1	3
D15	A	1 (3)	1 (9)	1	3

^a K_0 and K_2 are as defined by equations (1) and (2) in units of $\text{cm}^2 \text{s}^{-1}$. The parameter $\gamma = 1$. φ_H is the total column production of hydrogen atoms in the units $7 \times 10^8 \text{ atoms cm}^{-2} \text{s}^{-1}$. The solar L α flux refers to the integrated flux of photons at 1 AU in units $3.5 \times 10^{11} \text{ photons cm}^{-2} \text{s}^{-1}$. Model A is a warm atmosphere as described in Figure 1. Model B is an isothermal atmosphere at 150 K as described in Strobel 1973. The numbers $a(b)$ read as $a \times 10^b$.

IV. DISCUSSION OF RESULTS

A fairly comprehensive set of models has been generated to understand the sensitivity of the major observable constituents, C₂H₆, C₂H₂, and H to the input parameters. The essential features of the models are summarized in Table 3. The A and B series are designed to explore the dependence on eddy diffusion coefficient and temperature when the production of H

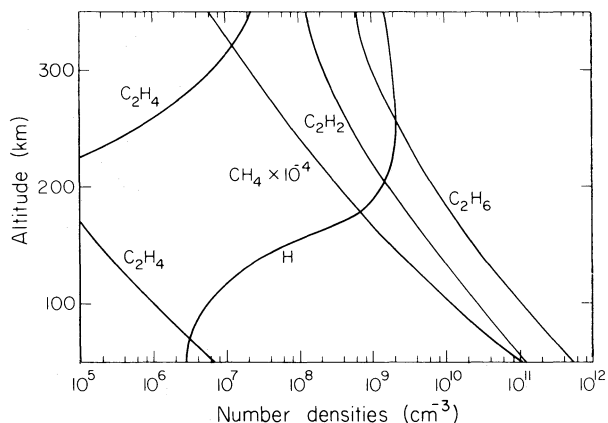


FIG. 2.—Number densities of major species computed with model A13. See Table 3 for input parameters used in model calculations for this and all subsequent figures.

is due solely to solar EUV radiation, estimated by Strobel (1973) to be $7 \times 10^8 \text{ atoms cm}^{-2} \text{s}^{-1}$. The C series explores the consequence of atomic hydrogen production associated with energetic particle precipitation observed by Broadfoot *et al.* (1979). The possible impact of reactions (17c) on the abundances of C₂H₆ and C₂H₂ is investigated in the D runs. The net result is an enhanced dissociation of CH₄, which is equivalent to an increased incident solar flux.

Our discussion of the numerical results will concentrate on observable constituents: C₂H₆, C₂H₂, and H. Figure 2 illustrates altitude profiles for the major hydrocarbon species computed in model A13 with solar UV radiation as the only source and an “average” L α reflectivity (cf. Fig. 5). The results would have been a plausible prediction for the *Voyager* encounter. The eddy diffusion coefficients were $K_0 = 10^3$ and $K_2 = 10^7 \text{ cm}^2 \text{s}^{-1}$. Figure 3 shows the dependence of C₂H₆ on the choice of various input parameters. The C₂H₆ concentration increases with higher K in the upper atmosphere and lower K at the lower boundary, with lower temperatures in the inversion region, and enhanced production by particle precipitation. The C₂H₆ dissociation rate is not neglig-

TABLE 4
VALUES OF IMPORTANT PARAMETERS IN L α BRIGHTNESS CALCULATIONS^a

Parameter	Value
Natural line width Δv_N	$4.986 \times 10^7 \text{ s}^{-1}$
Doppler line width Δv_D	$1.061 \times 10^9 \sqrt{T} \text{ s}^{-1}$
Cross section at line center $\sigma_0 = \frac{\pi e^2 f}{mc \pi \Delta v_D}$	$5.876 \times 10^{-12} \frac{1}{\sqrt{T}} \text{ cm}^2$
Methane absorption	$1.9 \times 10^{-17} \text{ cm}^2$
Solar L α flux at 1 AU used in photochemical calculations ^b	$3.5 \times 10^{11} \text{ photons cm}^{-2} \text{s}^{-1}$
Solar L α flux at 1 AU used in brightness emissions calculations. (assume rectangular profile with full-width = 1 Å)	$5.1 \times 10^{11} \text{ photons cm}^{-2} \text{s}^{-1} \text{ \AA}^{-1}$

^a T is temperature in kelvins. The numerical values are from Carlson and Judge 1971. The solar L α flux is from Rottman and Mount 1980.

^b This assumes about 70% of the incident solar L α is absorbed by CH₄. The remaining 30% is reflected.

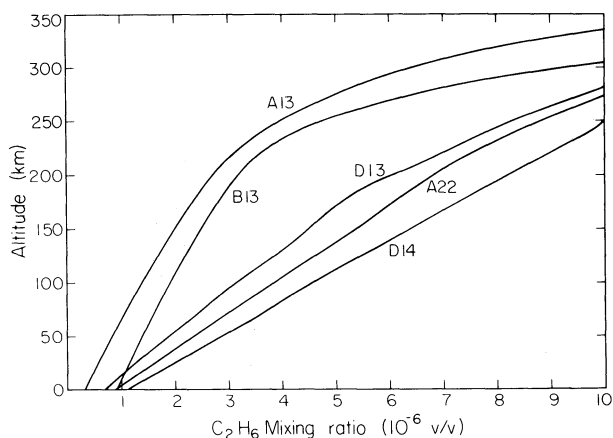


FIG. 3.— C_2H_6 mixing ratio by volume computed with models A13, A22, B13, D13, and D14.

ible in comparison to the mixing rate for small K . With the CH_4 absorption cross sections based on Mount, Warden, and Moos (1977), we find that CH_4 does not effectively shield C_2H_6 from dissociation as in previous calculations (Strobel 1973, 1974). Consequently the C_2H_6 density is sensitive to $K(z)$ in the lower stratosphere. Also, at the lower boundary, downward mixing of C_2H_6 from its source region is a major sink. Our calculations suggest an upper limit on the C_2H_6 mixing ratio of 3×10^{-6} at the $p = 10^{-2}$ atm level. We emphasize that our lower mixing ratio (compared to the large mixing ratio of Strobel 1974 with $\phi = 0$ lower boundary condition) results from the boundary condition (5) and the larger C_2H_6 dissociation rate. According to Tokunaga, Knacke, and Owen (1976) their observations would require an inversion layer at 190 K if the C_2H_6 mixing ratio were only $\sim 5 \times 10^{-7}$. However, an inversion layer at 150 K requires a mixing ratio of 10^{-5} to be in accord with their observations. Indeed, Orton and Aumann (1977) prefer a warm inversion layer with $T > 155$ K and increasing with height to ~ 200 K at $p = 10^{-4}$ atm. The results of Figure 3 in conjunction with mixing ratios of C_2H_6 inferred from *Voyager* IR data may indicate sources of C_2H_6 and C_2H_2 in addition to solar UV dissociation (Hanel 1979). It is important to recognize the time constants for various processes in the upper atmosphere. The lifetime associated with the column integrated hydrocarbon abundance ($C_2H_6 + C_2H_2$) exceeds 100 yr and thus should not vary significantly in response to changes in auroral activity and the column integrated atomic hydrogen abundance. The lifetime associated with the column integrated atomic hydrogen abundance is of order 100 days. Steady state calculations of C_2H_6 and C_2H_2 densities in the lower stratosphere for *Voyager* conditions may be underestimates as a consequence of the large H atom concentration which suppresses recombination of CH_3 radicals to C_2H_6 .

Figure 4 shows the distribution of H in a number of models. The abundance of H above the methane

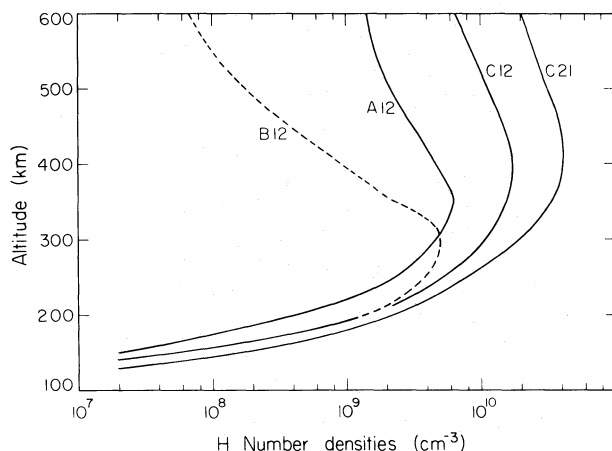


FIG. 4.—Atomic H number densities computed with models A12, B12, C12, and C21.

absorption layer is a sensitive function of the eddy diffusion coefficient in the upper atmosphere K_2 and the flux of hydrogen atoms produced from particle precipitation. The reflectivity $4I/F_J$, as defined by Wallace and Hunten (1973), is calculated for the various models and summarized in Figure 5. The curves A and B roughly reproduce the results previously obtained by Wallace and Hunten (1973), but

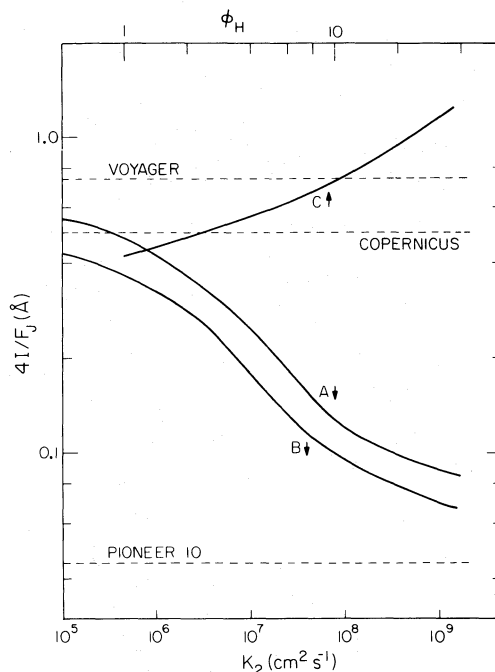


FIG. 5.—Disk-averaged reflectivity $4I/F_J$ at $L\alpha$ in Å units. The curves A and B were computed with models A_{ij} and B_{ij} , respectively. The abscissa is the eddy diffusion coefficient in the upper atmosphere K_2 . The curve C was computed using model atmosphere A with $K_2 = 10^6$ cm^2 s^{-1} . The abscissa is the flux of hydrogen atoms at the upper boundary in units of 7×10^8 atoms cm^{-2} s^{-1} . The dashed lines show the mean values of $4I/F_J$ as measured by *Pioneer 10*, *Copernicus*, and *Voyager* (see Table 1).

are consistently higher by $\sim 30\%$. This may be partly due to the use of a Rayleigh phase function rather than an isotropic phase function for scattering of $L\alpha$ photons by H atoms. We conclude that the $L\alpha$ albedo increases slightly with temperature, and the exceedingly high temperature (~ 1200 K) in the thermosphere of model A has practically no effect. The reason is that most of the photons are scattered deep in the atmosphere from the Lorentz wings of the $L\alpha$ line. The curve C was computed to test the sensitivity of $4I/F_J$ to φ_H . In the standard model, we used model atmosphere A, $K_0 = 10^3$, $K_2 = 10^6 \text{ cm}^2 \text{ s}^{-1}$ and $\varphi_H = \varphi_0 = 7 \times 10^8 \text{ atoms cm}^{-2} \text{ s}^{-1}$. The value of φ_H was then varied between φ_0 and $30 \varphi_0$.

Broadfoot *et al.* (1979) reported a disk-averaged $L\alpha$ brightness of 14 kR. Although no measurement of the solar $L\alpha$ intensity was made during the *Voyager* encounter, a solar $L\alpha$ flux of $3.75 \times 10^{11} \text{ photons cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ is estimated on the basis of the solar 10.4 cm flux (Vidal-Madjar 1975). A slightly higher flux of $5.1 \times 10^{11} \text{ photons cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ was measured by Rottman and Mount (1980) on 1979 June 5. Only a small fraction of the 14 kR can be due to direct excitation by energetic particles. Our reasons are twofold. First, the night-side $L\alpha$ intensity is less than 1 kR (Sandel *et al.* 1979). Second, the center-to-limb variation of the observed $L\alpha$ intensities is in good agreement with that predicted for resonance scattering, but not with that due to particle excitation as shown in Figure 6 (Canahan and Zipf 1977; Gladstone and Yung 1979). The center-to-limb variation of $L\alpha$ intensities could offer additional confirmation of the scattering in the Lorentz wing. The dashed line in Figure 6 was computed by arbitrarily restricting the scattering to within 6 Doppler units from the core and is not in good agreement with observation. However, we need higher-quality data before we can settle this issue.

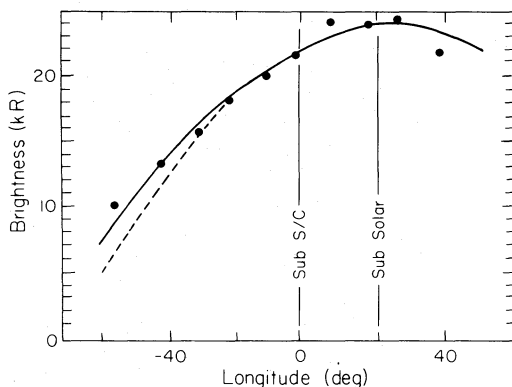


FIG. 6.— $L\alpha$ brightness as a function of longitude computed with model C12 for the same viewing geometry as described by Broadfoot *et al.* 1979. The intensity at 20° has been normalized to 24 kR. The dashed line results from arbitrarily cutting off the emission line profile beyond about 6 Doppler units and renormalizing the intensity.

The observed variability of the $L\alpha$ brightness (cf. Table 1 and Bertaux *et al.* 1979) suggests that the lifetime of atomic hydrogen in the Jovian thermosphere is less than a year. This implies an eddy diffusion coefficient at the homopause of more than $\sim 10^6 \text{ cm}^2 \text{ s}^{-1}$. From the results of Figure 5 we are forced to conclude that the *Voyager* $L\alpha$ brightness requires energetic particle dissociation of H_2 rather than a very low eddy diffusion coefficient. The principal problem in understanding the variability of Jupiter's upper atmosphere is the *Pioneer 10* UV photometer results (Carlson and Judge 1974). From Table 1 we note the reflectivity varies by a factor of 2 about a central value of 0.4 \AA , with the exception of the *Pioneer 10* results. From Figure 5 we infer that modest changes in H production, temperature, and eddy diffusion coefficient would easily explain this factor of 2 variability. However, the very large increase in K_2 and possibly temperature decrease required to understand *Pioneer 10* UV results are puzzling. It should be noted that the *Pioneer* results are at least internally self-consistent. Assuming a He/ H_2 ratio equal to 0.11 (Hanel *et al.* 1979), a very high eddy diffusion coefficient ($K_2 \gtrsim 10^8 \text{ cm}^2 \text{ s}^{-1}$) is required to account for the *Pioneer* observation of 5.2 R emission in the short-wavelength channel. Broadfoot *et al.* (1979) reported an upper limit of 0.1 R for the 584 \AA resonance emission of He. If confirmed, this would imply a much lower $K_2 \approx 10^5 \text{ cm}^2 \text{ s}^{-1}$. It is conceivable that the atmosphere had undergone orders of magnitude of change in K_2 in the last few years. But until further evidence becomes available on the variability of the upper atmosphere, the *Pioneer 10* results should be taken with reservation.

V. CONCLUSIONS

The high disk-averaged $L\alpha$ emission observed by Broadfoot *et al.* (1979) can be explained by resonance scattering of sunlight by hydrogen atoms. These observations require a larger source of H than solar EUV radiation. We suggest precipitation of magnetospheric particles to dissociate H_2 and heat the upper atmosphere and estimate an average column production rate of at least $7 \times 10^9 \text{ atoms cm}^{-2} \text{ s}^{-1}$, about 10 times more than solar EUV can produce. This is consistent with auroral emissions observed by *Voyager*. A low eddy diffusion coefficient alone cannot account for the observed $L\alpha$ brightness; our calculations suggest $K_2 \approx 10^6 - 10^7 \text{ cm}^2 \text{ s}^{-1}$. A value of $K_2 \approx 10^6$ requires an H column production rate of 7×10^9 , whereas, $K_2 \approx 10^7$ requires a production rate of $\sim 7 \times 10^{10}$ to obtain the *Voyager* observed albedo. For the average reflectivity, $4I/F_J \approx 0.4 \text{ \AA}$ in Table 1, we infer $K_2 \approx 10^7 \text{ cm}^2 \text{ s}^{-1}$ for solar production of H only.

To account for the inferred abundances of the hydrocarbons ($\text{C}_2\text{H}_6 + \text{C}_2\text{H}_2$) our updated model

requires a low eddy coefficient at the tropopause (10^{-1} atm) $K_0 \approx 10^3 \text{ cm}^2 \text{ s}^{-1}$ and a fairly high eddy coefficient at the homopause $K_2 \approx 10^7 \text{ cm}^2 \text{ s}^{-1}$. Since the lifetime for total (column integrated) C_2 compounds in the stratosphere exceeds 100 yr, the values for K_0 and K_2 deduced from our calculations are an average over this period of time. This is consistent with average conditions inferred from $L\alpha$ albedo observations. We also suggest that magnetospheric particle precipitation may produce CH_4 dissociation rates in

excess of solar UV radiation during a certain portion of this period when solar activity is high.

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