



Supplementary Materials for

Black hole feedback in the luminous quasar PDS 456

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Materials and Methods

1. Observations and data reduction

PDS 456 was observed simultaneously with the *XMM–Newton* (28) and *NuSTAR* (29) X-ray satellites four times between August 27 and September 21, 2013 (hereafter Obs. 1–4), and once again on February 26–27, 2014 (Obs. 5). The details of each observation are summarized in Table S1.

The *XMM–Newton* data were reduced and calibrated according to the standard procedures, using the Scientific Analysis System version 13.0 and the latest version of the calibration database made available by the *XMM–Newton* Science Operations Centre. The observations were taken in the Large Window mode of the European Photon Imaging Camera (EPIC) pn CCD detector, with the medium optical blocking filter in place. We checked each light curve for particle background flares, removing any time intervals above the threshold of 0.5 count/s in Obs. 1–4, and 1 count/s in Obs. 5. The resultant cumulative net exposure was 483 ks. Source and background spectra were extracted from circular regions of 35'' and 60'' radius, respectively. The net count rates over the 0.45–10 keV energy band are in the range ~ 1.3 (Obs. 5) to 3.1 counts/s (Obs. 1). At these levels, the effects of photon pile-up on the spectra are negligible.

Although the *XMM–Newton* analysis was eventually performed on pn spectra only, we carried out a self-consistency check with the data from the two MOS detectors, whose effective areas drop quickly at the energies of the Fe-K features. A remarkable agreement was found in each observation down to ~ 0.45 keV, which was therefore assumed as the lower energy bound in our fits. The spectra were grouped to a minimum of 100 counts per energy bin in order to allow the use of χ^2 minimization during the spectral analysis, which was performed with the XSPEC fitting package version 12.8.1g.

We also considered the source photometry in the six wide-band *V*, *B*, *U*, *UVW1*, *UVM2* and *UVW2* filters of the *XMM–Newton* Optical Monitor (OM). We averaged the count rates from the small high-resolution window at the center of the OM field of view over the different observations, with a 10% uncertainty added in quadrature to account for systematic errors in the flux of the calibration stars. We applied an absorption correction based on the standard extinction law (30), with values of $A(V)$ and $E(B-V)$ appropriate for the source (31, 32). The five OM data sets were used to determine the shape of the optical–ultraviolet (UV) ionizing continuum.

The simultaneous *NuSTAR* level 1 data products were processed with the *NuSTAR* Data Analysis Software package version 1.3.1 included in the 6.15.1 HEASoft release. Event files were produced, calibrated, and cleaned using the standard filtering criteria and the latest files available in the *NuSTAR* calibration database (CALDB version 20131223). The total net exposure, after the exclusion of periods of source occultation by Earth and passages on the South Atlantic Anomaly, was around 299 ks for both *NuSTAR* co-aligned telescopes, with corresponding Focal Plane Modules A (FPMA) and B (FPMB). Extraction radii were 80'' for both source and background spectra, which were grouped to have at least 50 counts per bin.

2. Parameterization of the P-Cygni profile

The considerable curvature characterizing PDS 456 in the soft X-rays and the spectral variability on timescales as short as weeks (Fig. 1) have been usually attributed to reprocessing in multiple layers of absorbing gas along the line of sight. In the wake of this, we performed a preliminary inspection of the hard X-ray spectra between 2–30 keV using a partially covered continuum model. We jointly fitted all the simultaneous *XMM-Newton* and *NuSTAR* observations in the 2–10 and 3–30 keV energy intervals, respectively. Whilst not necessarily physically motivated at this stage, a good description of the overall continuum shape was obtained through a power law with photon index $\Gamma \sim 2.3$, partially covered (covering fraction $f_{\text{cov}} \sim 0.1\text{--}0.4$) by a neutral gas column of $N_{\text{H}} \sim 10^{23} \text{ cm}^{-2}$ (Table S2). Figure 2 shows the five ratio plots of the data against this model.

The P-Cygni-like feature in the Fe-K spectral region can be fitted with two Gaussian profiles (one in emission and the other in absorption) at $E_{\text{rest}} = 6.97 \text{ keV}$, corresponding to the rest-frame energy of the Fe XXVI Ly α resonance line. Additional residuals remain beyond $\sim 9 \text{ keV}$, indicating the possible presence of further atomic structures associated with the absorber. To investigate whether all these hard X-ray features arise from the same stream of outflowing gas, we included another Gaussian absorption line and a photoelectric edge into the baseline continuum model. The former was identified with the Ly β transition at $E_{\text{rest}} = 8.25 \text{ keV}$, while the latter was introduced to account for the K-shell absorption edge at $E_{\text{rest}} = 9.28 \text{ keV}$ (the minimum photon energy capable of stripping the last electron from Fe XXVI). We fixed all the rest-frame energies to their expected laboratory values, assuming a common broadening σ_{K} and allowing the relative outflow velocity between the emission line and the absorption components to vary.

Such an ensemble offers an excellent description of the Fe-K spectral range in each observation (Table S2). If attributed to Fe XXVI Ly α , the emission profile appears to be slightly blueshifted with respect to the systemic velocity of PDS 456, with $v_{\text{out}} = 0.029 \pm 0.017c$ (uncertainties are given at the 90% confidence level). On the other hand, a substantially larger outflow velocity $v_{\text{out}} \sim 0.24c$ is implied for the Fe XXVI Ly α , Ly β and K-shell edge absorption series, whose simultaneous and self-consistent detection rules out any identification with lower ionization species (e.g. Fe XXV). Both the emission and absorption profiles are well resolved, with $\sigma_{\text{K}} = 380^{+50}_{-30} \text{ eV}$. This corresponds to a velocity dispersion of $\sim 13,000 \text{ km/s}$ across the flow ($\sim 30,000 \text{ km/s}$ full width at half-maximum).

This model returns a very good fit of the entire 2–30 keV spectra ($\chi^2 = 2540.8$ for 2582 degrees of freedom). The equivalent width (EW) of the Ly α and Ly β absorption lines follows a general trend of being weaker at the start of the campaign and becoming systematically larger afterwards, while the overall count rate decreases accordingly. The Ly α EW varies from $-75 \pm 45 \text{ eV}$ (in Obs. 1) to $-350 \pm 60 \text{ eV}$ (in Obs. 3), and the Ly β exhibits a similar behavior (Table S2). The maximum optical depth of the absorption edge (τ_{max}) is in keeping with the observed EW trend, i.e. τ_{max} is higher when the Ly α and Ly β EWs are larger. This is consistent with a line variability driven by changes in the opacity of the absorbing gas.

The apparent Fe XXVI P-Cygni-like feature hints at a large covering factor of highly ionized matter in proximity of the central black hole. The Sobolev approximation with exact integration (SEI) is a commonly used method to calculate the P-Cygni line profiles

associated with the expansion of spherically symmetric stellar winds (33). A SEI-based model has been also developed to specifically reproduce Fe-K absorption features in AGN (10). We assumed a velocity field of the form $w = w_0 + (1-w_0)(1-1/x)^\gamma$, where $w = v/v_\infty$ is the wind velocity normalized to the terminal velocity, w_0 is the dimensionless velocity at the wind photosphere, $x = r/R_0$ is the radial distance in units of the photospheric radius R_0 , and γ governs how the velocity scales with distance. Since the exact values of γ and w_0 have little effect on the derived P-Cygni profiles, we chose $\gamma = 1$ and $w_0 = 0.01$. For this given velocity field, the line optical depth varies as $\tau(w) \propto \tau_{\text{tot}} w^{\alpha_1} (1-w)^{\alpha_2}$, where the quantities α_1 and α_2 determine the smoothness of the resultant P-Cygni profile, with larger values leading to more gradual shapes (34). We then fitted again all five observations simultaneously between 2–30 keV, replacing the pair of $K\alpha$ Gaussian lines with a self-consistent P-Cygni profile. For simplicity, we assumed $\alpha_1 = \alpha_2$ (in practice they cannot be constrained individually), and fixed the underlying continuum slope to the best-fit value from the previous model, to prevent any degeneracy between the spectral curvature and the P-Cygni parameters. The fit variables include the characteristic energy of the profile (E_c), corresponding to the onset of the absorption component, the terminal velocity of the wind (v_∞), and the total optical depth of the gas (τ_{tot}). While E_c and v_∞ were required to be the same over the five sequences, τ_{tot} was left free to vary independently to account for changes in the line intensity, and increases from $\tau_{\text{tot}} = 0.05 \pm 0.01$ (Obs. 1) to $\tau_{\text{tot}} \sim 0.10\text{--}0.15$ (Obs. 2–5), confirming that the opacity of the reprocessing gas is lower at the beginning of the campaign (Table S2). The characteristic energy of the P-Cygni line is $E_c \sim 7.4 \pm 0.1$ keV (rest-frame), and its terminal velocity is $v_\infty = 0.35 \pm 0.02c$. None of the higher order features are included in this model by default. Once the $K\beta$ absorption line and the K-shell edge are re-introduced, the fit statistic is $\chi^2/\nu = 2613.1/2604$, comparable to that achieved in the Gaussian case.

3. Broadband modeling

The P-Cygni fit implies that the disk wind in PDS 456 is qualitatively consistent with a quasi-spherical flow. Our next goal is to quantify both the properties of the absorbing gas and the solid angle of the wind, by using the XSTAR photo-ionization code (35). The ionization balance is sensitive to the distribution of ionizing photons between 1–1,000 Rydberg (i.e., from 13.6 eV to 13.6 keV), as determined by the shape of the underlying UV to X-ray radiation field. By combining the EPIC pn and *NuSTAR* X-ray spectra with the data from the six *XMM-Newton* OM photometric bands, we thus obtained a proper input ionizing continuum for the generation of the relevant XSTAR absorption and emission models. The spectral energy distribution (SED) can be approximated from 2 eV to 30 keV by a three-segment broken power law. We adopted an X-ray photon index $\Gamma = 2.4$ to match the mean hard X-ray spectral slope found by *XMM-Newton* and *NuSTAR*, while the time-averaged OM data follow a $\Gamma = -0.7$ trend. The connecting UV to X-ray slope at 10–500 eV is $\Gamma \sim 3.3$. This model predicts a total ionizing luminosity L_{ion} of the order of 5×10^{46} erg/s in the 1–1,000 Rydberg bandpass. We also investigated a more realistic scenario, replacing the broken power law with multicolor blackbody emission from an accretion disk with photon temperature $kT \sim 5\text{--}10$ eV. This is similar to what is expected for a black hole with mass of $\sim 10^9 M_{\text{sun}}$ accreting at nearly Eddington rate (36),

and entails $L_{\text{ion}} \sim 10^{47}$ erg/s, as is likely the case for PDS 456 (see below).

The XSTAR code computes the radiative transfer through a spherically symmetric shell of absorbing gas, yielding two grids of absorption and emission spectra. Absorption tables have three main free parameters: the column density N_{H} of the gas, its ionization parameter ξ (defined as $\xi = L_{\text{ion}}/nR^2$, where n is the gas number density, and R the distance from the ionizing source), and the absorber's redshift z . Emission tables also take a normalization, which depends on the solid angle Ω subtended by the reprocessing gas shell to the illuminating source. With the information gathered from the SED analysis, we generated a pair of absorption/emission XSTAR grids with Solar abundances suitable for PDS 456, and covering the $5 \times 10^{21} < N_{\text{H}}/\text{cm}^{-2} < 2 \times 10^{24}$ and $1 < \log(\xi/\text{erg cm s}^{-1}) < 8$ parameter ranges. These components were used to model all of the effects imparted by the accretion disk wind on the observed spectra. Based on the observed width of the Fe-K features, we imposed a velocity smearing $v_{\text{K}} = 15,000$ km/s. Another grid with a modest broadening of 100 km/s, which is more typical of the soft X-ray warm absorbers on parsec scales (37, 38), was produced to account for a more distant layer of absorbing gas.

The ultimate interpretation of the observed spectral variability in the broadband X-ray emission of PDS 456 involves a combination of changes in the properties of the absorbing gas and subtle variations of the primary continuum. The mathematical structure of the model we applied can be represented as follows:

$$\text{galabs} \times \text{warmabs} \times \text{windabs} \times (\text{pcabs} \times \text{powerlaw} + \text{windem} + \text{softexc}). \quad (1)$$

The various components, identified above through self-evident tags, are illustrated in Fig. S2. The global fit, performed simultaneously over all the data sets, yields a statistic of $\chi^2 = 4372.5$ for 4180 degrees of freedom (while $\chi^2/\nu = 2651.3/2610$ above 2 keV, for a null hypothesis probability of 0.28) with no conspicuous residuals (Fig. S3). Below we proceed to a more detailed discussion of the model components.

- *galabs*: Galactic absorption was included in all our fits, with standard cross sections and Solar abundances (39). The hydrogen column density was fixed throughout to $N_{\text{H,Gal}} = 2.0 \times 10^{21} \text{ cm}^{-2}$, as appropriate for the coordinates of the source (40).
- *powerlaw*: The primary continuum was modeled using a power law, representing the intrinsic X-ray emission of PDS 456 before reprocessing. In order to fit the broadband spectra, some intrinsic variability is required, with the photon index Γ lying in the range ~ 2.3 – 2.7 and a factor of ~ 3 difference in the 7–30 keV flux across the observations (Table S3).
- *windabs*: Blueshifted K-shell absorption from highly ionized iron was reproduced by means of an XSTAR grid. Any changes in the depth of Fe-K absorption can be solely ascribed to ionization changes, with $\log(\xi/\text{erg cm s}^{-1})$ decreasing from $6.30^{+0.10}_{-0.16}$ in Obs. 1, when the feature is weakest, to a minimum of $5.75^{+0.06}_{-0.16}$ in Obs. 3, when it is strongest. Qualitatively, these variations trace the behavior of the continuum flux. Conversely, the column density is consistent within the errors over the entire campaign, and was therefore tied across the five data sets returning a best-fit value of $N_{\text{H}} = 6.9^{+0.6}_{-1.2} \times 10^{23} \text{ cm}^{-2}$. The blueshift of the high- ξ absorber turns out to be virtually constant (Table S3), implying an outflow velocity $v_{\text{out}} = 0.25 \pm 0.01c$, in agreement with the preliminary

Gaussian-line fit. Moreover, the trend of decreasing ξ (hence increasing opacity) confirms the findings of the P-Cygni analysis.

- *windem*: Emission associated with the disk wind was included in the model as an XSTAR grid. Indeed, photons absorbed within a quasi-spherical or wide-angle wind are subject to non-negligible re-emission along the line of sight, thereby leading to a P-Cygni-like profile. For this component, we assumed the same ionizing SED as per the absorption tables, with equal velocity smearing, ionization state and column density. The only free parameters of the emitter are its normalization and its redshift (common to the five epochs), which prefers a moderate outflow speed $v_{\text{out}} = 0.012 \pm 0.010c$.
- *pcabs*: Partial covering absorption has been established by previous studies as the origin of the pronounced spectral curvature below $\sim 3\text{--}4$ keV in PDS 456 (6–8). For consistency, we accounted for this component by using the same custom XSTAR absorption table generated for the high-ionization wind. For a best-fit ionization parameter of $\log(\xi/\text{erg cm s}^{-1}) = 2.9$, we found the overall spectral shape to be well explained by a variable column density, decreasing from $N_{\text{H}} \sim 2.5 \times 10^{23} \text{ cm}^{-2}$ in Obs. 1–2 to $N_{\text{H}} \sim 0.7 \times 10^{23} \text{ cm}^{-2}$ in Obs. 5, with small changes of the covering fraction in the range $f_{\text{cov}} \sim 0.3\text{--}0.5$ (Table S3). Most interestingly, the partially covering gas accepts some degree of blueshift relative to the local frame of PDS 456, with an outflow velocity commensurate with that of the Fe-K absorber ($v_{\text{out}} \sim 0.25c$). The fit is substantially worse if this component is forced to the systemic velocity of PDS 456 ($z = 0.184$), with $\Delta\chi^2 \sim +80$ for the same degrees of freedom. Given the lower ionization of this gas, and the slow variability over several weeks, it is then possible that partial covering arises from more distant clumps within the wind, as found in other AGN (11).
- *softexc*: Similarly to what had been previously reported (6, 41), the data seem to require a soft excess below ~ 1 keV. As the origin of this component (ubiquitous among AGN) is still largely ambiguous, here we modeled it through an ad hoc broad Gaussian emission line. Indeed, the exact parameterization of the soft excess does not affect in any way the fit of the high-ionization Fe-K features, nor our conclusions on the properties of the wind in PDS 456.
- *warmabs*: A fully covering absorber of low ionization $\log(\xi/\text{erg cm s}^{-1}) = 0.31 \pm 0.02$ and minor velocity broadening, possibly representing a distant warm absorber, also contributes to the observed spectral curvature. Its column $N_{\text{H}} = 4.3^{+0.2}_{-0.1} \times 10^{21} \text{ cm}^{-2}$ is essentially constant in Obs. 1–4, and appears to be lower in Obs. 5, reaching $N_{\text{H}} = 1.8 \pm 0.2 \times 10^{21} \text{ cm}^{-2}$. This suggests that the typical variability timescale for this component is a few months long. Due to the low column density value, this putative warm absorber is only manifest below ~ 2 keV, having no effect on the Fe-K region of the spectrum and the derived physical parameters.

Supplementary Text

1. Alternative models

The detection of blueshifted absorption lines in the hard X-ray spectra of AGN is usually associated with fast outflows. However, this interpretation is not unequivocal. It has also been proposed that some of these features could instead arise from resonant absorption in a layer of highly ionized, optically thick plasma located above the accretion disk, and co-

rotating with it (42). The main advantage of this scenario (referred to as ‘disk absorption’ in the following for ease of discussion) is that relativistic velocities are natural in proximity of the central black hole, and there is no need to invoke any wind launching mechanism. Here absorption is imprinted on the X-rays reflected off the disk surface, while any direct power-law continuum would be left unaffected (see Fig. 1 in (42) for a geometrical illustration). Thus, for prominent absorption lines to be produced, the observed 2–10 keV spectrum has to be dominated by the reflected component, rather than by the primary emission. In other words, the reflection strength RS (43) has to be much larger than one. In this framework, the Fe XXVI $K\alpha$ feature can be shifted from its rest-frame energy of 6.97 keV up to 9 keV and possibly beyond, depending on the exact geometry. This model has been successfully applied (12) to reproduce the absorption line observed in PG 1211+143, whose blueshift would correspond to an outflow velocity of $\sim 0.10\text{--}0.15c$ in the wind scenario. Fe-K emission from the disk, once blurred by relativistic effects, accounts for the remains of the claimed P-Cygni profile, which, however, emerged only after combining multiple observations taken several years apart (13).

We have tested this picture for PDS 456 on Obs. 3, during which the equivalent width of the absorption line appears to be larger. Energies below 3 keV were initially neglected to avoid the complications due to additional soft X-ray absorption and emission components. To mimic the presence of a hot atmosphere blanketing the disk, we defined a model with the following structure:

$$galabs \times (powerlaw + relblur \times diskabs \times diskref).$$

The standard power-law continuum also provides the input illumination for the disk reflection component (44), which is convolved with a blurring kernel (tagged as *relblur* above) that imparts the required relativistic distortions (45). The 9-keV absorption is included as an XSTAR table (*diskabs*), linked to the reflected spectrum only and subject to the same smearing. When varying freely, this model tries to account for the flux drop at the onset of the absorption trough as either the blue wing of a blurred Fe $K\alpha$ emission line or a K-shell edge, rejecting any effect from the hot-gas absorber through either a vanishing column density or an extremely high ionization. This pulls down the continuum, leaving obvious residuals at the upper end of the *XMM-Newton* bandpass, while at higher energies the shape of the *NuSTAR* spectra is gradually recovered by a modest Compton reflection hump (46).

For illustration purposes, we fixed the values of the dimensionless spin parameter $a^* = cJ/GM^2 = 0.7$ (J and M are the black hole’s angular momentum and mass) and of the disk inclination $i = 75$ degrees, as dictated by the desired blueshift of the Fe XXVI $K\alpha$ line (42). We also set the ionization of the absorbing layer to $\log(\xi/\text{erg cm s}^{-1}) = 6$. Figure S4 shows the fit residuals of this disk absorption model against those of the broadband wind model. The 9-keV feature is clearly missed, even if $10\times$ Solar Fe abundance is adopted. It is in fact the small reflection fraction of $RS \sim 0.6$ that prevents the absorption profile from reaching an adequate depth. It is still possible that more complex geometrical configurations than the one considered here can reproduce the Fe XXVI $K\alpha$ feature. Its apparent consistency over the different epochs, however, requires all the X-ray spectral states exhibited by PDS 456 in the course of this campaign to be reflection-dominated.

This seems quite unlikely, given the remarkable changes in both intensity and overall spectral shape.

In order to investigate this possibility, and to assess the reflection contribution in emission, we re-fitted the broadband 0.5–30 keV spectra of all observations after including in the wind model a self-consistent blurred reflection component originated by a lamppost source located above the disk (47). Any parsec-scale reflector can be disregarded due to the lack of a narrow Fe K line, while an independent soft excess component (again parameterized as a broad Gaussian line) is confirmed to be highly significant. The improvement achieved through this model is questionable ($\Delta\chi^2 \sim -27$ for eight degrees of freedom less), and disk reflection is not statistically required in four out of five observations (Obs. 1–4) according to an F -test. Even in Obs. 5, the height of the illuminating source is several tens of gravitational radii; this is not compatible with an extreme light bending regime and the consequent suppression of the direct continuum due to the focusing of the radiation towards the disk (48). Incidentally, the best-fit statistics cannot be recovered by replacing either the partial covering or the wind emission with disk reflection. Both cases were checked separately, and revealed a tendency to overestimate the *NuSTAR* spectra, which do not show any evidence of a reflection hump at ~ 20 keV (as opposed to many other local AGN). This is also clear from the comparison between Obs. 2 and the 2007 *Suzaku* observation (Fig. S5), where a hard excess indicative of a Compton hump was tentatively detected (7, 49).

In conclusion, we can definitely rule out the association of the 9-keV absorption trough in PDS 456 with a non-wind scenario, while the presence of a reflection component (either off the disk or within the wind itself) is not imposed by the data, and has been neglected.

2. Mass outflow rate

Let $\rho(r)$ and $v(r)$ be the local mass density and velocity of the wind, respectively. The mass outflow rate through a thin, spherical shell of radius r and solid angle Ω is given by $\dot{M}_{\text{out}} = \Omega r^2 \rho(r) v(r)$. For cosmic elemental abundances and full ionization $\rho(r) \sim 1.2 m_p n(r)$, where m_p is the proton mass and $n(r)$ is the ion number density. Neglecting the constant factor of order unity, and assuming that the wind has a constant velocity $v(r) = v_{\text{out}}$, we obtain $\dot{M}_{\text{out}} \sim \Omega r^2 m_p n(r) v_{\text{out}}$. The product $r^2 n(r)$ is constant, and can be associated to observables as follows:

$$N_{\text{H}} = \int_{R_{\text{in}}}^{\infty} n(r) dr \sim \int_{R_{\text{in}}}^{\infty} \left(\frac{\dot{M}_{\text{out}}}{\Omega m_p v_{\text{out}}} \right) r^{-2} dr = \frac{\dot{M}_{\text{out}}}{\Omega m_p v_{\text{out}} R_{\text{in}}}, \quad (2)$$

where the inner radius R_{in} is the starting point of the wind. The final expression brings out all the relevant dependencies for the mass outflow rate:

$$\dot{M}_{\text{out}} \sim \Omega N_{\text{H}} m_p v_{\text{out}} R_{\text{in}}. \quad (3)$$

We now discuss the measures of all the physical quantities involved.

2.1 Solid angle

Hydrodynamical models of the radiation transport through ionized biconical winds have been calculated to explain the P-Cygni profile of strong UV lines in bright Galactic sources such as cataclysmic variables (50, 51). The current quality and resolution of the

X-ray spectra from AGN only allows a qualitative comparison of the P-Cygni feature detected in PDS 456 with the atlases of synthetic P-Cygni profiles. Given the depth of the absorption trough, the line of sight has to lie inside the wind. Assuming a symmetric (Gaussian) shape and Doppler broadening, the continuous line profile maps the components of the velocity field projected along the direction to the observer (Fig. S1). From the red wing of the emission and the central energy of the absorption, we infer a line-of-sight velocity range of $-0.04c$ to $0.25c$, which implies an opening angle for a radial outflow of $\theta \sim \cos^{-1}(v_{\min}/v_{\max}) \sim 100$ degrees, corresponding to a solid angle $\Omega \sim 3\pi$. While any wind acceleration is neglected, this value is consistent with the outcome of the P-Cygni fit.

The two main arguments to evaluate the solid angle occupied by the wind are actually independent on both the exact geometrical structure and the local velocity field, as outlined below:

- The normalization κ_{obs} of the XSTAR wind emission table (*windem*) in the broadband analysis includes direct information on the fraction of solid angle covered by the reprocessing gas shell. The grids were generated assuming $\Omega = 4\pi$ and a 1–1,000 Rydberg luminosity $L_{\text{ion}} = 10^{47}$ erg/s. This provides a reference value against which the normalizations delivered by the best fit (which depend on the actual covering fraction and ionizing luminosity) can be interpreted. Following the adopted SED shape, we first rendered explicit the relation between L_{ion} and the 7–30 keV luminosity, which is a measurable quantity and refers to most critical energy range (i.e. above the Fe XXVI $K\alpha$ threshold) given the ionization state of the gas. We can then assess the solid angle through the following expression:

$$\frac{\Omega}{4\pi} = 3.15 \times 10^3 \kappa_{\text{obs}} \left(\frac{L_{7-30 \text{ keV}}}{10^{44} \text{ erg/s}} \right)^{-1}. \quad (4)$$

Note that, in terms of the spectral analysis, κ_{obs} is ultimately driven by the EW of the Fe XXVI $K\alpha$ emission line (Fig. S2). The resulting ratios for the five observations 1–5 are ~ 0.83 (> 0.49), 1.02 (> 0.59), 0.42 ± 0.18 , $0.80^{+0.16}_{-0.20}$, 0.99 (> 0.75). The weighted average is 0.79 ± 0.15 , thus $\Omega = 3.2 \pm 0.6\pi$.

- Alternatively, we can compute $\Omega/4\pi$ as the fraction of the absorbed continuum luminosity that is re-emitted within the wind, $L_{\text{em}}/L_{\text{abs}}$. We derived these values over the 0.5–30 keV range, thanks to the simultaneous, broadband spectral coverage. Since both L_{em} and L_{abs} depend in a complex way from several parameters (Γ , N_{H} , ξ and flux normalizations), the errors were estimated through a bootstrap re-sampling technique over the parameters involved. As the distributions are not perfectly symmetric, the uncertainties correspond to the 16th and 84th percentiles, and are comparable to the 1.645 SD (90 per cent confidence level) error bars in the Gaussian approximation. We obtain the following measures of Ω for the individual observations: $1.9 \pm 0.7\pi$ (Obs. 1), $1.4 \pm 0.6\pi$ (Obs. 2), $0.7^{+0.3}_{-0.2}\pi$ (Obs. 3), $1.6 \pm 0.4\pi$ (Obs. 4), $1.8^{+0.6}_{-0.4}\pi$ (Obs. 5). The weighted value against the average absorbed and re-emitted luminosities is $\Omega = 1.4 \pm 0.2\pi$. However, this does not encompass the fluorescent yield of the gas, i.e. the probability that the creation of a vacancy in an inner shell is followed by radiative de-excitation rather than auto-ionization (i.e. the ejection of a second electron, known as the Auger effect). For Fe XXV

and Fe XXVI no electron is available for the Auger mechanism, but the Ly α transition is the most likely result of the recombination cascade, so that the effective fluorescent yield is ~ 0.5 and 0.7 , respectively (52). With such a correction, the solid angle of the wind is again $\Omega > 2\pi$.

2.2 Column density

The column density $N_{\text{H}} = 6.9^{+0.6}_{-1.2} \times 10^{23} \text{ cm}^{-2}$ of the highly ionized absorbing (and re-emitting) gas is a direct output of the spectral model in the broadband analysis. Since ion abundances slightly depend on the resolution of XSTAR grids, a lower sampling over the ionization parameter range might lead to underestimated opacities. The generation and implementation of finely sampled grids over a large area of the parameter space is rather time-consuming. Having identified the region the wind spectral components likely fall in, we produced another set of finer XSTAR tables with 15 linear steps at $5 < \log(\xi/\text{erg cm s}^{-1}) < 8$. With a negligible improvement in the overall fit statistic ($\Delta\chi^2 < -2$), the column density increases to $N_{\text{H}} = 8.8^{+0.4}_{-0.5} \times 10^{23} \text{ cm}^{-2}$. The values of $\log(\xi/\text{erg cm s}^{-1})$ grow by ~ 0.1 , remaining consistent with those given in Table S3, and none of the other variables are affected.

Another estimate of N_{H} is based on the optical depth τ_{tot} from the P-Cygni fit, which is related to the ionic column density of reprocessing gas N_i through the expression $\tau_{\text{tot}} = (\pi e^2/m_e c) f \lambda_0 N_i/v_\infty$, where e and m_e are the electron charge (in statcoulomb) and mass, f is the oscillator strength of a given transition, and λ_0 (in cm) is the laboratory wavelength of the line (33). For an identification with Fe XXVI Ly α , which has a mean $f = 0.21$ for the resonant line doublet (53), the average $\tau_{\text{tot}} \sim 0.12$ for Obs. 2–5 implies $N_i \sim 1.3 \pm 0.2 \times 10^{19} \text{ cm}^{-2}$. At the measured ionization state of the absorbing gas, the Fe XXVI Ly α ionic fraction is close to 1 (54). Thus, for standard Solar iron abundance (55), this suggests an equivalent hydrogen column density $N_{\text{H}} = 4.0 \pm 0.6 \times 10^{23} \text{ cm}^{-2}$, comparable to that found through the XSTAR photo-ionization code.

2.3 Outflow velocity

The terminal velocity of the wind is one of the fit parameters in the custom P-Cygni model. The best value is $v_\infty = 0.35 \pm 0.02c$, and is consistent with the velocity corresponding to the full width at zero intensity of the absorption trough in the Gaussian-line fit. While this should be regarded as the actual speed of the outflowing gas, in keeping with the standard practice for narrow, unresolved lines, we adopt for the computation of the wind's energetics the more conservative line-of-sight bulk velocity $v_{\text{out}} = 0.25 \pm 0.01c$, averaged over the viewing angle through the absorbing gas and commensurate with the energy of maximum depth in the absorption profile.

2.4 Launch radius

The large uncertainty on the location of the absorbing gas usually prevents an accurate determination of the wind's energetics. The best way to overcome this limitation is through a timing approach, in order to measure how absorption varies over timescales of days/weeks. Since this is possible only for a small number of well-studied sources, a single-epoch estimate of the distance of the absorber can be derived from the spectral analysis, based on the ionization parameter of the gas $\xi = L_{\text{ion}}/nR^2$. The thickness Δr of the gas layer responsible for the high-ionization absorption is lower than or equal to its

maximum distance R_{\max} from the central source. Since $N_{\text{H}} \sim n\Delta r$, this condition becomes $R_{\max} < L_{\text{ion}}/N_{\text{H}}\xi$. With an average $\xi \sim 10^6 \text{ erg cm s}^{-1}$, and $L_{\text{ion}} \sim 5 \times 10^{46} \text{ erg/s}$ (obtained from the SED fit), we get $R_{\max} < 7 \times 10^{16} \text{ cm}$ or $\sim 500 r_{\text{g}}$. This is compatible with a recombination time of nearly one week for an inferred electron density of $\sim 10^7 \text{ cm}^{-3}$.

The present campaign on PDS 456 allowed us to witness the considerable spectral variability of the source on relevant timescales to constrain R_{in} . The most dramatic variations occur from Obs. 1 to Obs. 2, when the hard X-ray power-law flux measured by *NuSTAR* decreases by a factor of three over the span of about ten days, and the Fe XXVI $K\alpha$ line flux also drops significantly (Table S2). In general, the wind emission follows the same variability pattern of the high-energy continuum in the five epochs, suggesting that the gas is in photo-ionization equilibrium. In fact, the variations of the ionization parameter ξ correlate fairly well with those of the overall intensity. The light-crossing time between Obs. 1 and Obs. 2 corresponds to a distance of $\sim 2.6 \times 10^{16} \text{ cm}$ ($170 r_{\text{g}}$), and sets an upper limit to R_{in} . Even if we cannot perform a proper reverberation study, R_{in} is unlikely to be much smaller than this value. The order-of-magnitude estimate of a few hundreds of r_{g} is further supported by the sizable changes in the depth of the absorption line over a week, observed just a few months earlier during the *Suzaku* monitoring (18). In principle, R_{in} can be as low as the escape radius, which is the innermost starting point of the wind to be physically acceptable and lies between $\sim 16 r_{\text{g}}$ (for $v_{\infty} = 0.35c$) and $32 r_{\text{g}}$ (for $v_{\text{out}} = 0.25c$). Even in this case, the corresponding mass loss would still be a sufficient fraction of the Eddington accretion rate to produce significant feedback (see below). However, such a small spatial extent would likely give rise to spectral variability on timescales of ~ 1 – 2 days (i.e. detectable within the single observations), for which no clear evidence was found.

3. Wind energetics

There is no direct measurement for the mass of the central black hole in PDS 456. An estimate can be obtained from the scaling relations between the virial radius of the broad-line region (BLR) and the black hole mass, based on reverberation studies, $M_{\text{BH}} =$

$G^{-1}R_{\text{BLR}}V_{\text{BLR}}^2$, where V_{BLR} is the Keplerian velocity of the BLR clouds. With $\lambda L_{\lambda}(5100 \text{ \AA}) \sim 2 \times 10^{46} \text{ erg/s}$ (56) and $\text{FWHM}(\text{H}\beta) = 3974 \pm 764 \text{ km/s}$ (30), we compute R_{BLR} using the latest $R_{\text{BLR}}-\lambda L_{\lambda}(5100 \text{ \AA})$ relationship (57), and assume $V_{\text{BLR}} \sim \text{FWHM}(\text{H}\beta)$ neglecting any correction factor of order unity to account for the BLR geometry. This results in $\log(M_{\text{BH}}/M_{\text{sun}}) = 9.24 \pm 0.17$. Given the uncertainties, and keeping in mind that all the subordinate quantities below are accurate within a factor of ~ 2 , we have assumed for simplicity a value of $M_{\text{BH}} = 10^9 M_{\text{sun}}$ throughout this work. This implies for PDS 456 an Eddington luminosity $L_{\text{Edd}} \sim 1.3 \times 10^{47} \text{ erg/s}$, and an Eddington mass accretion rate $\dot{M}_{\text{Edd}} \sim 1.4 \times 10^{27} \text{ g/s}$ or $22 M_{\text{sun}}/\text{yr}$ (for a radiative efficiency $\eta \sim 0.1$).

Based on the discussion in the previous section, we adopt $\Omega = 2\pi$, $N_{\text{H}} = 6 \times 10^{23} \text{ cm}^{-2}$, $v_{\text{out}} = 0.25 \pm 0.01c$, and $R_{\text{in}} = 1.5 \times 10^{16} \text{ cm}$ in Eq. 4, and calculate a mass outflow rate:

$$\dot{M}_{\text{out}} \sim 7 \times 10^{26} \text{ g/s} \sim 11 M_{\text{sun}}/\text{yr} \sim 0.5 \dot{M}_{\text{Edd}}. \quad (6)$$

Under the same assumptions, the kinetic power of the wind is therefore:

$$P_{\text{kin}} = \frac{1}{2} \dot{M} v_{\text{out}}^2 \sim 2 \times 10^{46} \text{ erg/s} \sim 0.15 L_{\text{Edd}}. \quad (7)$$

While the above fractions of Eddington limits depend on the actual black hole mass, the absolute values of the mass loss and of the associated kinetic power do not. We can then determine the total mechanical power E_{wind} released by the AGN through the wind over a typical lifetime of 10^8 yr (58). We take a duty cycle of 0.1. In other words, the wind is in action for $\sim 10^7$ yr. The fraction of AGN with detected X-ray signatures of fast outflows (about 40%) can be interpreted in two different ways: for a duty cycle of 1, this is a measure of the covering fraction $\Omega/4\pi$, whereas it would indicate the wind lifetime in case of full covering. Unless the wide-angle, high-column, high-velocity phase of the wind is extremely short-lived, our guess is therefore rather conservative, and leads to $E_{\text{wind}} \sim 6 \times 10^{60}$ erg. In order to understand the possible impact in terms of feedback, this has to be compared to the binding energy of the galactic bulge in a system like PDS 456, whose order of magnitude has been again evaluated from the scaling relations as $E_{\text{grav}} \sim M_{\text{bulge}} \sigma_*^2$, where M_{bulge} is the dynamical mass of the bulge and σ_* is the stellar velocity dispersion. For a $10^9 M_{\text{sun}}$ black hole, $M_{\text{bulge}} \sim 10^{12} M_{\text{sun}}$ (59, 60), while $\sigma_* \sim 330$ km/s (61, 62). Hence $E_{\text{grav}} \sim 2 \times 10^{60}$ erg, and the condition $E_{\text{wind}} \gtrsim E_{\text{grav}}$ safely holds for $\log(M_{\text{BH}}/M_{\text{sun}}) < 9.5$. Since the deposition of a few per cent of the wind power in the interstellar matter can be enough to trigger a snowplow effect on galactic scales, even the constraints on both the actual launch radius (which can be closer to the escape radius if the matter is too ionized to produce any detectable effect) and the wind lifetime can be significantly relaxed. Indeed, even a modest coupling (with $\sim 10\%$ efficiency) between the disk wind and any energy-conserving, large-scale outflow would culminate in the usual values of mass-loss rate ($\sim 10^3 M_{\text{sun}}/\text{yr}$) and velocity ($\sim 10^3$ km/s) observed in the molecular gas component of several local AGN (22).

With respect to the wind driving mechanism in PDS 456, no persistent outflow can arise from resonant line absorption due to the extreme ionization state of the gas. For sources accreting at about the Eddington limit, however, the wind is expected to be optically thick to electron scattering, so that continuum radiation pressure becomes effective (63). Here the best-fit column density of the highly ionized absorber implies a Thomson optical depth of $\tau \sim 0.5$, which is in practice a lower limit, since additional, fully ionized gas would remain undetected. Assuming $\tau = 1$, conservation of momentum requires that $\dot{M}_{\text{out}} v_{\text{out}} \sim L_{\text{bol}}/c$. For the derived wind parameters, we find that the momentum of the gas exceeds the momentum of the photons by a factor of ~ 1.5 . Given the uncertainties on \dot{M}_{out} and L_{bol} , this is a reasonable agreement with momentum conservation. Moreover, it is possible that some other form of driving (e.g. magnetic) is also present, further alleviating any tension.

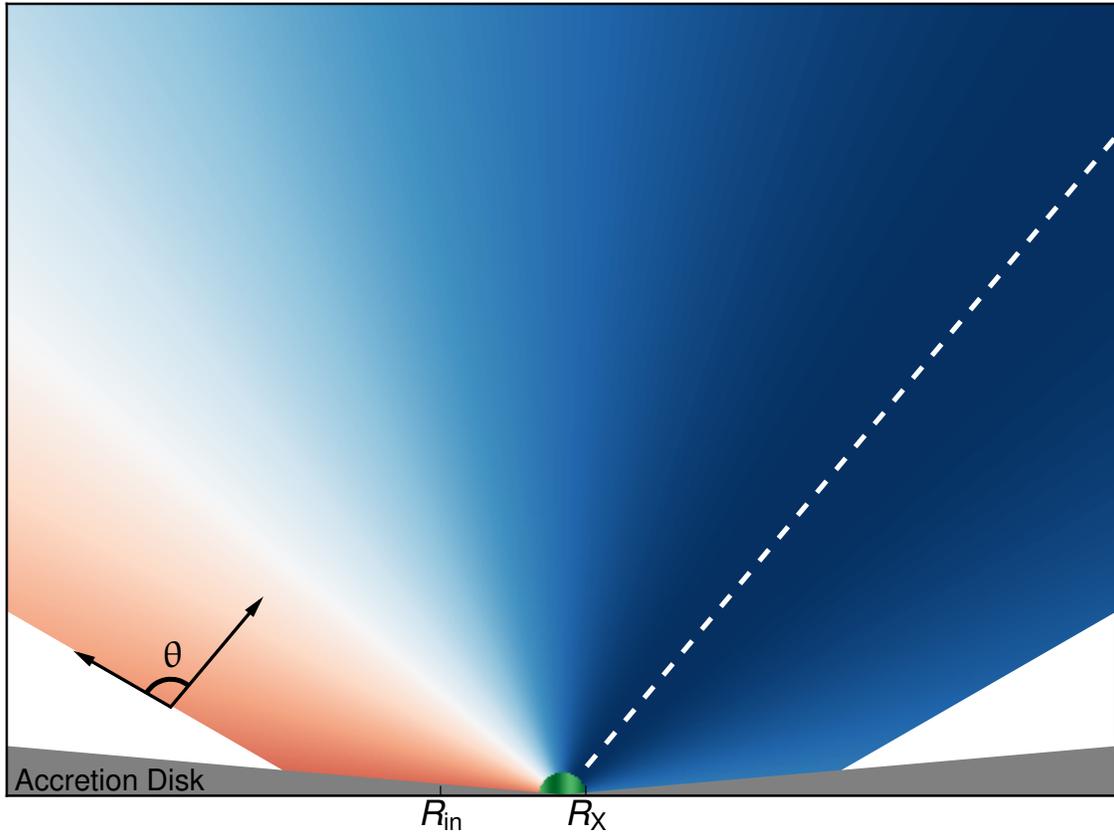


Fig. S1.

Possible origin of the P-Cygni line profile. For a preliminary, qualitative evaluation of the wind's opening angle θ , we can consider a radial outflow with uniform velocity. Due to the Doppler effect, the different energies in the P-Cygni profile correspond to different velocity components projected along the line of sight (marked by the dashed line). In this diagram (not to scale), the angular color code indicates the degree of blueshift/redshift of the Fe XXVI $K\alpha$ line for various directions within the outflow. In this oversimplified framework, the observed range of velocities implies that $\theta \sim \cos^{-1}(v_{min}/v_{max}) \sim 100$ degrees (hence $\Omega \sim 3\pi$ sr). It is worth noting that the overall geometrical structure (e.g. polar versus equatorial) of the outflow is actually unknown, and no assumptions have been made in the derivation of its solid angle. This notwithstanding, the flux variability timescale in PDS 456 suggests a size of the X-ray source of the order of $R_X \sim 10 r_g$, while the starting point of the wind is most likely around $R_{in} \sim 100 r_g$.

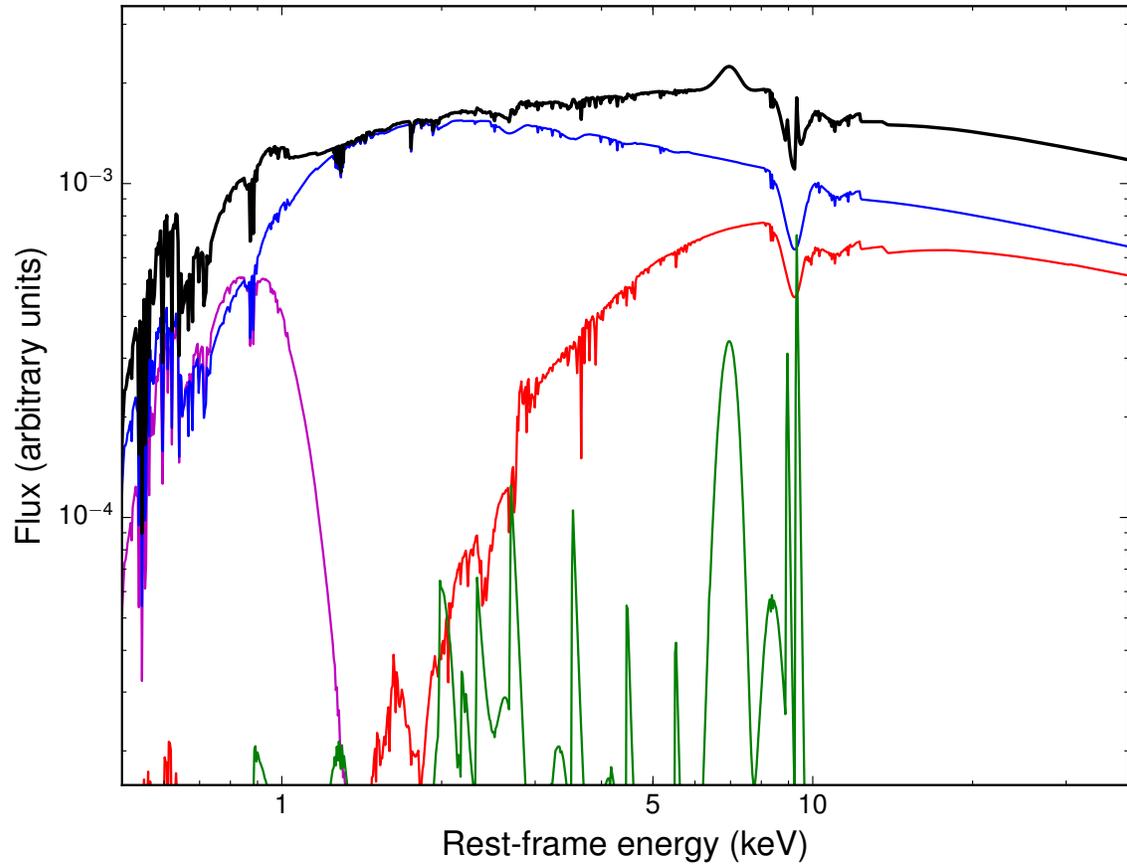


Fig. S2

Illustration of the wind model components. Template of the self-consistent broadband model (black curve; based on Obs. 4) with all the emission components disentangled: primary power-law continuum, either direct (blue) or transmitted through the partial covering layer (red), wind emission (green) and soft excess (magenta). The absorption features imprinted by the highly ionized gas are clearly appreciable. Note, however, that CCD detectors have a resolving power of only ~ 20 – 50 , for an energy resolution of around 80 eV at 1 keV, and 150 eV at 6 keV.

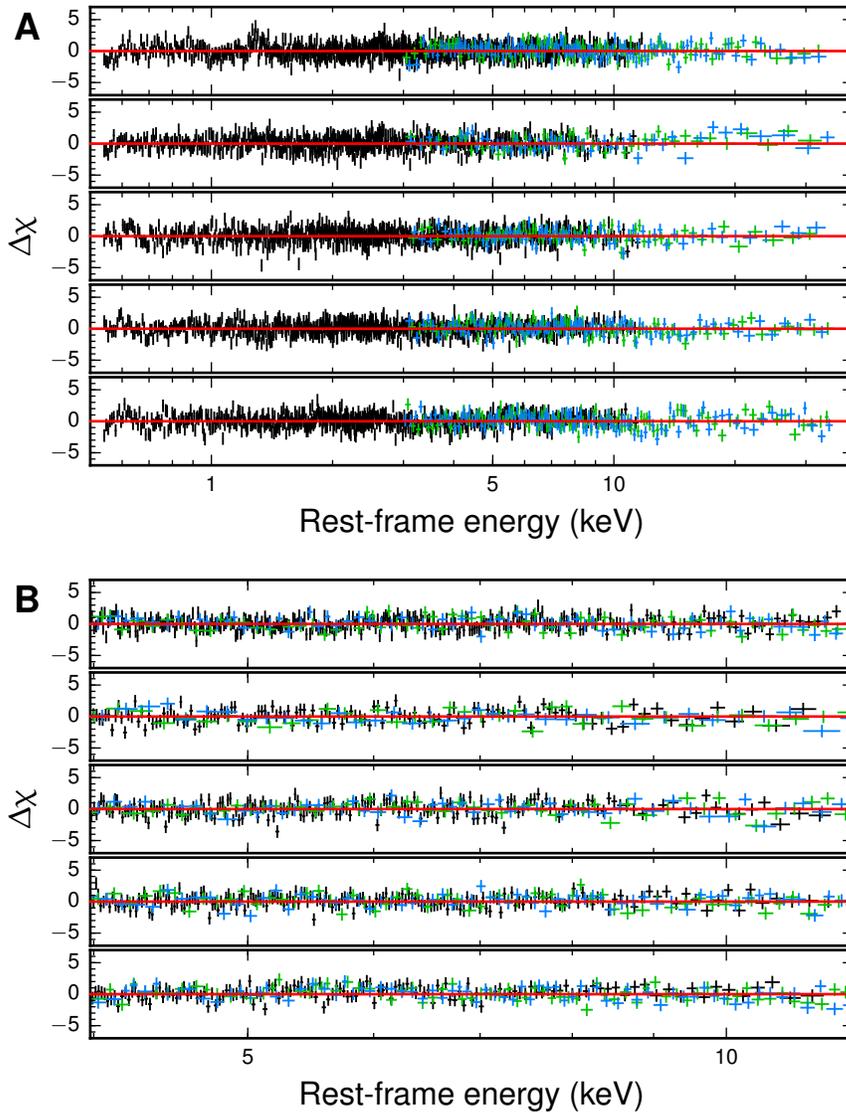


Fig. S3

Spectral residuals of the broadband fit. (A) Residuals with respect to the best-fitting wind model ($\chi^2/\nu = 4372.5/4180$) for each of the five observations 1–5 (in descending order), plotted in units of SD with error bars of size one. (B) Zoom on the 4–12 keV energy range (rest-frame), showing that no emission/absorption structure is left unmodeled in the Fe-K band ($\chi^2/\nu = 1465/1460$).

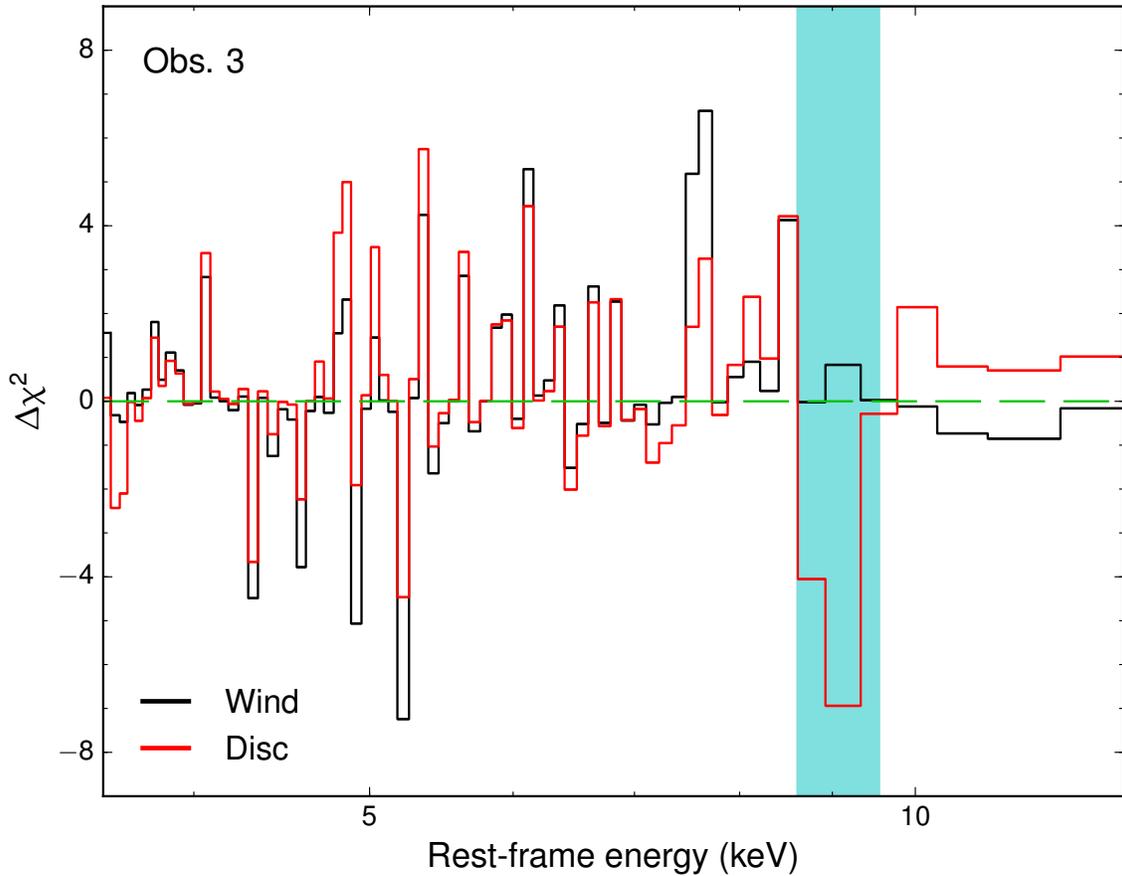


Fig. S4

Disk versus wind Fe $K\alpha$ absorption. Residuals over the best fit for Obs. 3 for both the disk absorption (in red) and wind absorption (in black) models. The former scenario is not capable of accounting for the depth of the absorption trough around 9 keV even if $10\times$ Solar Fe abundance is adopted. This is due to the fact that the observed continuum is not dominated by a disk reflection component (see the text for thorough discussion). The shaded area indicates the energy range corresponding to the full width at half maximum of the absorption feature when fitted with a simple Gaussian profile. Only *XMM-Newton* data (re-binned) are plotted for clarity.

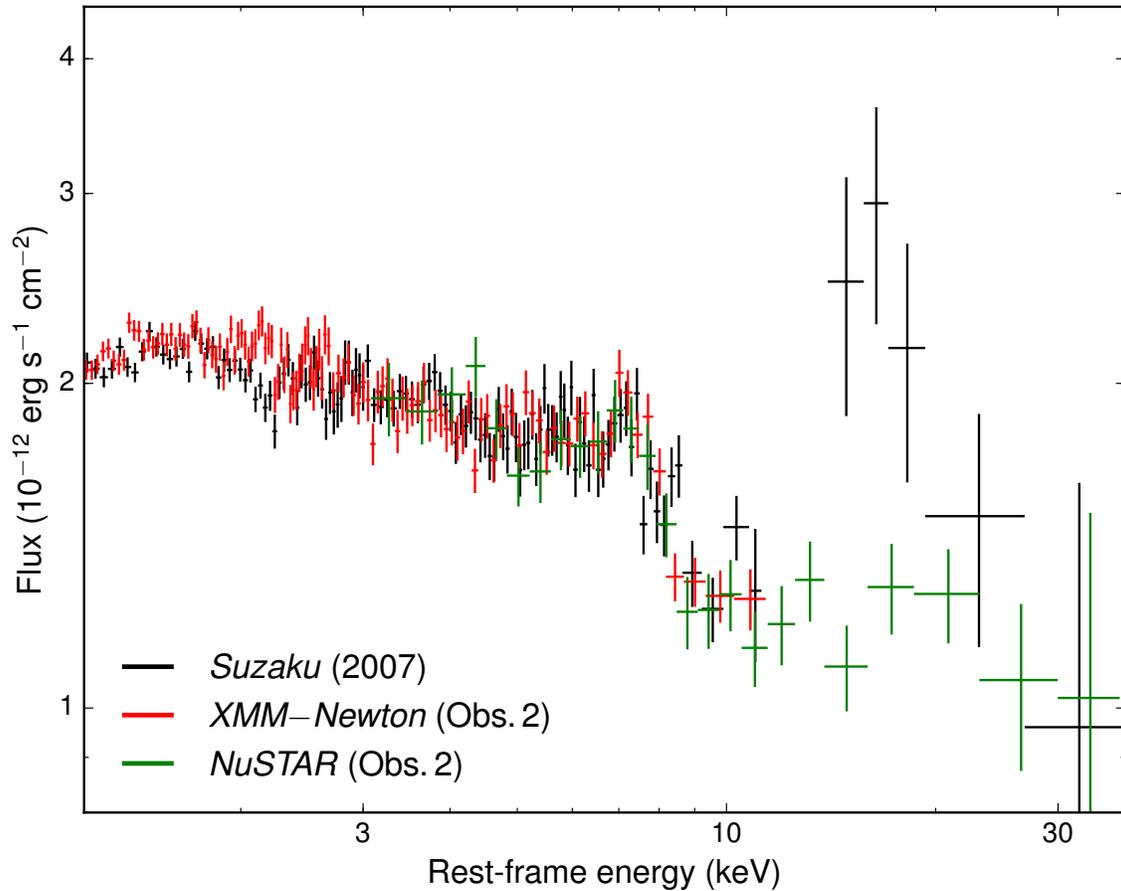


Fig. S5

Lack of any Compton reflection hump. Comparison between Obs. 2 of the current *XMM-Newton* and *NuSTAR* campaign and the 2007 *Suzaku* observation of PDS 456 (the latter scaled down by a factor of ~ 0.75). While the hard excess tentatively detected by *Suzaku* is consistent with a standard ($RS \sim 1$) contribution from disk reflection to the observed spectrum, a similar Compton hump is not present in any of the much more sensitive *NuSTAR* data sets. This implies that any reflection is quite faint and hard to constrain, possibly due to a different configuration of the X-ray illuminating source.

Table S1.

***XMM-Newton* and *NuSTAR* observation log.** T_{tot} : total elapsed time. T_{net} : net exposure after standard screening. Background subtracted count rates refer to the 0.45–10 keV (*XMM-Newton*) and 2.5–30 keV (*NuSTAR*) bands.

Obs.	Satellite	Obs. Start (UT)	Obs. End (UT)	T_{tot} (ks)	Detector	T_{net} (ks)	Count rate (s^{-1})
1	<i>XMM-Newton</i>	2013-08-27, 04:41	2013-08-28, 11:13	110.0	pn	95.7	3.083 ± 0.006
					MOS1	105.8	0.730 ± 0.003
					MOS2	107.7	0.932 ± 0.003
	<i>NuSTAR</i>	2013-08-27, 03:41	2013-08-28, 11:41	113.9	FPMA	43.8	0.165 ± 0.002
					FPMB	43.8	0.156 ± 0.002
2	<i>XMM-Newton</i>	2013-09-06, 03:24	2013-09-07, 10:36	112.3	pn	92.2	1.990 ± 0.005
					MOS1	106.3	0.545 ± 0.002
					MOS2	106.3	0.566 ± 0.002
	<i>NuSTAR</i>	2013-09-06, 02:56	2013-09-07, 10:51	113.9	FPMA	42.9	0.064 ± 0.001
					FPMB	43.0	0.061 ± 0.001
3	<i>XMM-Newton</i>	2013-09-15, 18:47	2013-09-17, 03:57	119.4	pn	102.0	1.883 ± 0.004
					MOS1	111.0	0.501 ± 0.002
					MOS2	111.4	0.574 ± 0.002
	<i>NuSTAR</i>	2013-09-15, 17:56	2013-09-17, 04:01	119.1	FPMA	44.0	0.089 ± 0.002
					FPMB	44.0	0.083 ± 0.002
4	<i>XMM-Newton</i>	2013-09-20, 02:47	2013-09-21, 09:37	111.0	pn	93.0	1.854 ± 0.004
					MOS1	107.8	0.519 ± 0.002
					MOS2	107.7	0.574 ± 0.002
	<i>NuSTAR</i>	2013-09-20, 03:06	2013-09-21, 11:11	113.8	FPMA	58.5	0.095 ± 0.001
					FPMB	58.6	0.089 ± 0.001
5	<i>XMM-Newton</i>	2014-02-26, 08:03	2014-02-27, 22:51	139.7	pn	100.3	1.323 ± 0.004
					MOS1	119.1	0.415 ± 0.002
					MOS2	110.7	0.411 ± 0.002
	<i>NuSTAR</i>	2014-02-26, 08:16	2014-02-28, 22:56	224.3	FPMA	109.7	0.060 ± 0.001
					FPMB	109.3	0.059 ± 0.001

Table S2.

Preliminary parameterization of the Fe XXVI K-shell features. Γ : photon index. f_{cov} : covering fraction. σ : Gaussian line width. τ_{max} : optical depth of the absorption edge. $\Delta\chi^2$: change in the overall statistics (for an extra degree of freedom) after feature removal and re-fitting. E_c : characteristic energy of the P-Cygni profile (in keV). v_∞ : terminal outflow velocity. τ_{tot} and $\alpha_{1,2}$ control the dependence of the optical depth of the reprocessing gas on the velocity field (10, 32). (f) and (t) are frozen and tied parameters, respectively. Energies are given in the rest frame, while continuum fluxes are intrinsic at 2–30 keV. All errors correspond to 90% confidence levels for the single parameters.

Component Parameter	Observation				
	1	2	3	4	5
Continuum					
Power law					
Γ	$2.29^{+0.04}_{-0.03}$	(t)	(t)	(t)	(t)
F (10^{-12} erg s $^{-1}$ cm $^{-2}$)	$11.6^{+0.9}_{-0.8}$	$4.3^{+0.4}_{-0.2}$	$6.7^{+0.5}_{-0.4}$	$7.3^{+0.6}_{-0.5}$	$4.8^{+0.4}_{-0.3}$
Partial covering					
N_{H} (10^{22} cm $^{-2}$)	12 ± 1	(t)	(t)	(t)	(t)
f_{cov}	$0.32^{+0.04}_{-0.03}$	$0.06^{+0.06}_{-0.05}$	$0.39^{+0.02}_{-0.04}$	$0.40^{+0.02}_{-0.03}$	0.33 ± 0.04
Fe K emission					
Fe XXVI Ly α					
Energy (keV)	6.97(f)	(t)	(t)	(t)	(t)
σ_{K} (eV)	380^{+50}_{-30}	(t)	(t)	(t)	(t)
F (10^{-14} erg s $^{-1}$ cm $^{-2}$)	7.2 ± 0.3	4.0 ± 1.3	2.6 ± 1.6	5.1 ± 1.8	1.5 ± 1.3
EW (eV)	100 ± 30	145 ± 100	< 140	110 ± 80	< 135
v_{out}/c	0.029 ± 0.017	(t)	(t)	(t)	(t)
Fe K absorption					
Fe XXVI Ly α					
Energy (keV)	6.97(f)	(t)	(t)	(t)	(t)
σ_{K} (eV)	380(t)	(t)	(t)	(t)	(t)
F (10^{-14} erg s $^{-1}$ cm $^{-2}$)	$-4.3^{+2.1}_{-2.5}$	$-5.2^{+1.2}_{-1.5}$	$-11.1^{+1.5}_{-1.9}$	$-9.7^{+1.5}_{-2.1}$	$-7.9^{+1.2}_{-1.5}$
EW (eV)	-75 ± 45	-250 ± 70	-350 ± 60	-280 ± 50	-345 ± 55
v_{out}/c	$0.239^{+0.004}_{-0.005}$	(t)	(t)	(t)	(t)
Fe XXVI Ly β					
Energy (keV)	8.25(f)	(t)	(t)	(t)	(t)
σ_{K} (eV)	380(t)	(t)	(t)	(t)	(t)
F (10^{-14} erg s $^{-1}$ cm $^{-2}$)	> -3.9	-4.1 ± 1.5	-7.4 ± 1.8	$-4.3^{+1.8}_{-2.0}$	$-5.1^{+1.2}_{-1.4}$
EW (eV)	> -90	-250 ± 90	-290 ± 70	-160 ± 70	-280 ± 70
v_{out}/c	0.239(t)	(t)	(t)	(t)	(t)
$\Delta\chi^2$	0.7	20.5	44.1	14.5	43.6
Fe XXVI K edge					
Energy (keV)	9.28(f)	(t)	(t)	(t)	(t)
τ_{max}	< 0.06	$0.18^{+0.11}_{-0.14}$	$0.31^{+0.12}_{-0.11}$	0.29 ± 0.10	$0.41^{+0.11}_{-0.09}$
v_{out}/c	0.239(t)	(t)	(t)	(t)	(t)
$\Delta\chi^2$	–	4.3	21.1	23.6	52.9
P-Cygni profile					
E_c (keV)	$7.42^{+0.09}_{-0.08}$	(t)	(t)	(t)	(t)
v_∞/c	0.35 ± 0.02	(t)	(t)	(t)	(t)
τ_{tot}	0.05 ± 0.01	0.12 ± 0.03	0.13 ± 0.02	0.11 ± 0.02	0.12 ± 0.02
$\alpha_{1,2}$	$6.4^{+2.4}_{-1.6}$	(t)	(t)	(t)	(t)

Table S3.

Best-fit parameters of the broadband wind model. v_K : velocity smearing of the photo-ionization spectral grids. ξ : ionization parameter of the gas. κ_{obs} : normalization of the wind emission component. F_K : intrinsic ionizing flux for Fe XXVI (7–30 keV rest-frame). All the other symbols are the same as defined in the legend of Table S2, except for the continuum and soft excess fluxes that now refer to the 0.45–30 keV band.

Component Parameter	Observation				
	1	2	3	4	5
Galactic absorption					
$N_{\text{H,Gal}} (10^{22} \text{ cm}^{-2})$	0.2(f)	(t)	(t)	(t)	(t)
Primary continuum					
Γ	$2.28^{+0.01}_{-0.02}$	2.68 ± 0.02	2.43 ± 0.02	2.35 ± 0.02	$2.39^{+0.01}_{-0.02}$
$F (10^{-12} \text{ erg s}^{-1} \text{ cm}^{-2})$	25 ± 3	16 ± 2	16 ± 2	16 ± 2	10 ± 1
Wind absorption ($v_K = 15,000 \text{ km/s}$)					
$N_{\text{H}} (10^{22} \text{ cm}^{-2})$	69^{+6}_{-12}	(t)	(t)	(t)	(t)
$\log(\xi/\text{erg cm s}^{-1})$	$6.30^{+0.10}_{-0.16}$	$5.99^{+0.14}_{-0.13}$	$5.75^{+0.06}_{-0.16}$	$5.80^{+0.07}_{-0.15}$	$5.84^{+0.07}_{-0.12}$
v_{out}/c	$0.254^{+0.009}_{-0.008}$	$0.244^{+0.010}_{-0.014}$	$0.251^{+0.004}_{-0.003}$	$0.254^{+0.006}_{-0.003}$	$0.240^{+0.006}_{-0.005}$
Wind emission ($v_K = 15,000 \text{ km/s}$)					
$N_{\text{H}} (10^{22} \text{ cm}^{-2})$	69(t)	(t)	(t)	(t)	(t)
$\log(\xi/\text{erg cm s}^{-1})$	6.30(t)	5.99(t)	5.75(t)	5.80(t)	5.84(t)
$\kappa_{\text{obs}} (\times 10^{-3})$	$1.47^{+0.58}_{-0.60}$	$0.57^{+0.25}_{-0.24}$	0.38 ± 0.16	$0.82^{+0.16}_{-0.20}$	$0.59^{+0.16}_{-0.14}$
v_{out}/c	0.012 ± 0.010	(t)	(t)	(t)	(t)
Partial covering ($v_K = 15,000 \text{ km/s}$)					
$N_{\text{H}} (10^{22} \text{ cm}^{-2})$	25^{+2}_{-3}	25^{+2}_{-3}	18^{+2}_{-1}	17 ± 1	6.8 ± 0.4
$\log(\xi/\text{erg cm s}^{-1})$	2.91 ± 0.02	(t)	(t)	(t)	(t)
f_{cov}	$0.33^{+0.04}_{-0.05}$	0.44 ± 0.06	0.48 ± 0.05	$0.45^{+0.05}_{-0.06}$	0.50 ± 0.03
v_{out}/c	0.254(t)	0.244(t)	0.251(t)	0.254(t)	0.240(t)
Soft excess					
$E_s (\text{keV})$	0.63 ± 0.01	(t)	(t)	(t)	0.80 ± 0.01
$\sigma_s (\text{eV})$	196^{+6}_{-3}	(t)	(t)	(t)	105^{+9}_{-10}
$F_s (10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2})$	22^{+3}_{-1}	23^{+3}_{-1}	27^{+4}_{-2}	22^{+3}_{-1}	4.2 ± 0.3
Warm absorption ($v_K = 100 \text{ km/s}$)					
$N_{\text{H}} (10^{22} \text{ cm}^{-2})$	$0.43^{+0.03}_{-0.01}$	(t)	(t)	(t)	$0.18^{+0.01}_{-0.02}$
$\log(\xi/\text{erg cm s}^{-1})$	0.31 ± 0.02	(t)	(t)	(t)	(t)
v_{out}/c	0(f)	(t)	(t)	(t)	(t)
Cross norm.					
FPMA/pn	1.03 ± 0.02	$1.05^{+0.04}_{-0.03}$	$1.07^{+0.04}_{-0.03}$	$1.02^{+0.02}_{-0.03}$	$1.01^{+0.03}_{-0.02}$
FPMB/pn	$1.02^{+0.02}_{-0.03}$	1.05 ± 0.04	1.05 ± 0.03	$1.00^{+0.02}_{-0.03}$	$1.08^{+0.02}_{-0.03}$
7–30 keV flux					
$F_K (10^{-12} \text{ erg s}^{-1} \text{ cm}^{-2})$	5.6	1.8	2.8	3.2	1.9

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