

[C II] absorption and emission in the diffuse interstellar medium across the Galactic Plane[★]

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ABSTRACT

Aims. Ionized carbon is the main gas-phase reservoir of carbon in the neutral diffuse interstellar medium and its 158 μm fine structure transition [C II] is the most important cooling line of the diffuse interstellar medium (ISM). We combine [C II] absorption and emission spectroscopy to gain an improved understanding of physical conditions in the different phases of the ISM.

Methods. We present high resolution [C II] spectra obtained with the *Herschel*/HIFI instrument towards bright dust continuum sources regions in the Galactic plane, probing simultaneously the diffuse gas along the line of sight and the background high-mass star forming regions. These data are complemented by single pointings in the 492 and 809 GHz fine structure lines of atomic carbon and by medium spectral resolution spectral maps of the fine structure lines of atomic oxygen at 63 and 145 μm with *Herschel*/PACS.

Results. We show that the presence of foreground absorption may completely cancel the emission from the background source in medium spectral resolution PACS data and that high spectral resolution spectra are needed to interpret the [C II] and [O I] emission and the [C II]/FIR ratio. This phenomenon may explain part of the [C II]/FIR deficit seen in external luminous infrared galaxies where the bright emission from the nuclear regions may be partially canceled by absorption from diffuse gas in the foreground. The C⁺ and C excitation in the diffuse gas is consistent with a median pressure of $\sim 5900 \text{ K cm}^{-3}$ for a mean kinetic temperature of $\sim 100 \text{ K}$. A few higher pressure regions are detected along the lines of sight, as emission features in both fine structure lines of atomic carbon. The knowledge of the gas density allows us to determine the filling factor of the absorbing gas along the selected lines of sight. **The derived median value of the filling factor is 2.4 %, in good agreement with the properties of the Galactic Cold Neutral Medium.** The mean excitation temperature is used to derive the average cooling due to C⁺ in the Galactic plane : $9.5 \times 10^{-26} \text{ ergs}^{-1} \text{H}^{-1}$. Along the observed lines of sight, the gas phase carbon abundance does not exhibit a strong gradient as a function of Galacto-centric radius and has a weighted average of $C/H = 1.5 \pm 0.4 \times 10^{-4}$.

Key words. ISM : general; ISM : structure; ISM : individual objects : W28A, W31C, W33A, G34.3+0.15, W49N, W51, DR21(OH), W3-IRS5; ISM : lines and bands; Galaxy : disk; Infrared : ISM

1. Introduction

The diffuse interstellar medium presents a complex structure due to the coexistence of two stable phases of neutral gas, the Cold Neutral Medium (CNM) with moderate densities and temperatures ($n = n(\text{H I}) + n(\text{H}_2) \sim 30 \text{ cm}^{-3}$ and $T \sim 100 \text{ K}$) and the Warm Neutral Medium (WNM) with low densities and high temperatures ($n \sim 0.4 \text{ cm}^{-3}$ and $T \sim 8000 \text{ K}$), the two phases

being in pressure equilibrium (see Draine 2011). Higher pressure phases are associated either with HII regions created by massive stars, and with molecular clouds where the interplay of gravity and magnetic fields contributes to the creation of denser structures, in which thermal pressure does not balance other forces present. The classical tracer of both diffuse ISM phases is the hyperfine transition of neutral hydrogen at 21 cm (HI). Both phases contribute to the widespread emission, while the CNM can also be detected in absorption against either continuum sources or background HI emission. However it is still fairly difficult to estimate the relative contributions of these phases to the total

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amount of atomic hydrogen, and how this varies with the physical conditions in the Galaxy (Kalberla & Kerp 2009).

The $^2P_{3/2} - ^2P_{1/2}$ fine structure line of ionized carbon at $158 \mu\text{m}$, [C II], is the main cooling line of the CNM and as such is a complementary probe of this diffuse phase. The *Herschel* satellite has performed the first [C II] survey of the Galactic plane (Langer et al. 2010; Pineda et al. 2013; Langer et al. 2014). The measured [C II] emission has been compared with HI and CO data to identify the contributions from the different ISM phases, leading to the conclusion that the CNM gas represents $\sim 40\%$ of the atomic gas in the Galactic Plane. In addition to the molecular gas traced by CO emission, an additional 30% is traced by the [C II] emission. This ‘‘CO dark’’ molecular gas is also detected as an excess in thermal dust emission in the sub-millimeter regime as discussed by Planck Collaboration et al. (2011) using the *Planck Surveyor* data.

Detection of [C II] emission is limited by the sensitivity of the HIFI instrument. Absorption spectroscopy toward strong far infrared continuum sources can improve the ability to study low column density regions. This method has been pioneered for the fine structure lines of ionized carbon and neutral oxygen by Vastel et al. (2000, 2002) and by Lis et al. (2001) using the *Infrared Space Observatory* (ISO) long wavelength spectrograph (LWS) at moderate spectral resolution. The HIFI instrument on board the *Herschel* satellite has allowed for the first time high spectral resolution observations of both neutral and ionized carbon fine structure lines, with high sensitivity. It has also enabled the detection of ground state hydride lines from the diffuse interstellar medium (Gerin et al. 2012). These observations have revealed that neutral species like HF, CH and possibly water vapor, are reliable tracers of molecular hydrogen (Gerin et al. 2010; Neufeld et al. 2010; Flagey et al. 2013). Given the very large critical densities of the ground state hydride lines compared to the moderate densities in the diffuse ISM, the absorption data do not provide information on the physical conditions of the absorbing gas.

In this paper we present complementary data on the fine structure lines of ionized and neutral carbon aimed at better understanding the properties of the diffuse gas in the Galactic Plane. The high spectral resolution HIFI data are complemented by PACS spectral maps for the [C II] and [O I] fine structure lines, at medium spectral resolution. We study the density and pressure structure, as well as the the volume filling factor of the colder and denser phases (CNM and diffuse molecular gas) from the comparison of the emission and absorption data. Section 2 presents the observations. Section 3 discusses the analysis methods and gives the main results, while the conclusions are presented in Section 4 and summarized in Section 5. Appendices A and B give additional information on the observations and present the data in the complete set of sources observed.

2. Observations

2.1. Presentation of the HIFI observations

The source sample is based on the PRISMAS sources, as presented in e.g. Godard et al. (2010); Flagey et al. (2013). W3-IRS5 was targeted as an additional sight-line to probe the outer Galaxy as this source is located in the Perseus arm. The source positions, LSR velocities and distances are listed in Table 1 as well as the velocity intervals corresponding to the diffuse matter absorption and the associated atomic, molecular and total hydrogen column densities. Table A.1 in Appendix A summarizes the observing modes and provides the ObsID analyzed in this paper.

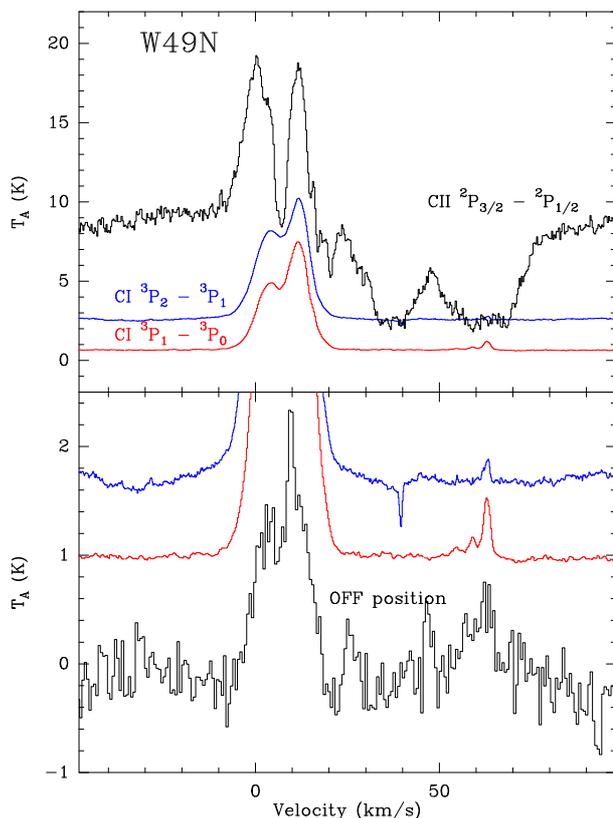


Fig. 1. Top : *Herschel*/HIFI spectra towards W49N. The red line shows the [C I] $^3P_1 - ^3P_0$ line at 492 GHz, the blue line shows the [C I] $^3P_2 - ^3P_1$ line at 809 GHz and the black line the [C II] $^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz. The horizontal axis is the LSR velocity in km s^{-1} and the vertical axis is the SSB antenna temperature in Kelvin. Bottom : zoom on the [C I] lines (red, blue as above) and the average [C II] spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

Analysis of absorption spectra requires accurate knowledge of the continuum level, which is best measured in DBS (Double Beam Switching) mode with heterodyne receivers like HIFI. However, since the [C II] emission is very extended in the Galaxy, this mode produces spectra that are contaminated by the presence of signals in the OFF beam. To avoid this problem, HIFI could be operated using the internal cold load (employed for calibration), as a reference for monitoring the system gain. This observing mode called ‘‘Load Chop’’ (LC) minimizes the contamination at the expense of a lower accuracy measurement of the continuum (de Graauw et al. 2010; Roelfsema et al. 2012). We therefore performed the [C II] and [C I] observations in two steps : we obtained a single pointing at the source position in DBS mode, followed by either a small map covering $\sim 50'' \times 50''$ (for [C II]) or a single pointing (for both [C I] lines) in LC mode. This combination provides the continuum at the central position and an accurate spectrum free of contamination, including both the background source emission and the foreground diffuse medium along the line of sight. As the continuum emission of the background source is semi-extended, we also used the PACS spectrometer (Poglitsch et al. 2010) to determine the $158 \mu\text{m}$ continuum over the area mapped with HIFI.

The spectroscopic parameters of the targeted fine structure lines are given in Table 2. At these frequencies, the angular resolution of the *Herschel* telescope is limited by diffraction, and the FWHM beam size is $44''$, $26''$ and $11''$, at 492, 809, and

Table 1. Summary of source properties

Name	RA	Dec	V_{LSR}	D	Ref. ^a	Velocity interval (km s ⁻¹)	$N_{tot}(\text{H})$ (10 ²¹ cm ⁻²)	$N(\text{HI})$ (10 ²¹ cm ⁻²)	$N(\text{H}_2)$ (10 ²¹ cm ⁻²)
	(J2000)	(J2000)	(km s ⁻¹)	(kpc)					
W28A(G5.89-0.4)	18:00:30.4	-24:04:00	10	1.3	1	[15,25]	6.5 ± 1.2	3.2 ^f	1.7 ± 0.4 ^e
W31C(G10.6-0.4)	18:10:28.7	-19:55:50	-2	5.0	2	[10,23]	15.6 ± 1.1	5.0 ^b	5.3 ± 0.5 ^c
						[23,34]	21.6 ± 2.0	5.4 ^b	8.1 ± 0.7 ^c
						[34,61]	19.3 ± 5.0	6.9 ^b	6.2 ± 0.6 ^c
W33A(G12.9-0.3)	18:14:39.4	-17:52:00	37	2.4	3	[14,25]	2.4 ± 0.5	2.0 ^c	0.2 ± 0.04 ^c
G34.3+0.15	18:53:18.7	+01:14:58	58	3.3	4	[8,20]	4.6 ± 0.5	2.2 ^b	1.2 ± 0.1 ^c
						[20,35]	7.6 ± 4.0	5.0 ^b	1.3 ± 0.1 ^c
W49N(G43.2-0.1)	19:10:13.2	+09:06:12	11	11.4	5	[25,50]	14.3 ± 0.7	7.0 ^b	3.7 ± 0.4 ^c
						[50,80]	22.6 ± 1.1	9.3 ^b	6.7 ± 0.7 ^c
W51(G49.5-0.4)	19:23:43.9	+14:30:30.5	57	5.4	6	[1,10]	3.3 ± 0.2	1.5 ^b	0.9 ± 0.1 ^c
						[10,16]	1.2 ± 0.1	1.0 ^b	0.12 ± 0.08 ^c
						[16,35]	2.5 ± 0.2	2.4 ^b	0.07 ± 0.02 ^c
DR21OH(G81.7+0.6)	20:39:00.7	+42:22:46.7	-2	1.5	7	[1,20]	12.8 ± 4.0	5.0 ^b	3.9 ± 0.4 ^c
W3-IRSS	02:25:40.6	+62:05:51	-37	2.0	8	[-25,-15]	1.6 ± 0.3	0.7 ^{b,d}	0.41 ± 0.1 ^e
						[-15,10]	1.7 ± 0.3	0.9 ^{b,d}	0.41 ± 0.1 ^e

^a References: (1) Motogi et al. (2011), (2) Sanna et al. (2014), (3) Immer et al. (2013), (4) Kuchar & Bania (1994), (5) Gwinn et al. (1992), (6) Sato et al. (2010), (7) Rygl et al. (2012), (8) Xu et al. (2006)

^b from (Winkel et al. 2014), ^c From Godard et al. (2012); ^d using data from Normandeau (1999); ^e this paper using $[\text{CH}]/[\text{H}_2] = 3.6 \times 10^{-8}$ (Sheffer et al. 2008) and $[\text{HF}]/[\text{H}_2] = 1.3 \times 10^{-8}$ (Sonnentrucker et al. 2010); ^f from Fish et al. (2003).

Table 2. Observed fine structure lines

Transition	ν (GHz)	λ (μm)	E_u (K)	g_u	A (s ⁻¹) ²⁰
C ⁺ (² P _{3/2} - ² P _{1/2})	1900.537	157.7	91.2	4	2.32 × 10 ⁻⁶
C (³ P ₁ - ³ P ₀)	492.160	609.1	23.6	3	7.93 × 10 ⁻⁸
C (³ P ₂ - ³ P ₁)	809.341	370.4	62.5	5	2.68 × 10 ⁻⁷
O (³ P ₁ - ³ P ₂)	4744.777	63.2	227.7	3	8.54 × 10 ⁻⁵
O (³ P ₀ - ³ P ₁)	2060.069	145.5	326.6	1	1.64 × 10 ⁻⁵

From CDMS (Müller et al. 2001, 2005)

1900 GHz, respectively. As we are focusing on the foreground absorption in this paper, we have not attempted to correct for the difference in angular resolution among the different lines.

2.2. HIFI Data processing

The observations have been analyzed with the Herschel Interactive Processing Environment (HIPE version 9¹). The HIFI data have been corrected for fringes using the *HIFIFitfringe* routine. The single pointing and maps have been exported to the CLASS package² for further processing. The HIFI mixers are double side band (DBS) and therefore sensitive to the emission in two frequency bands. Therefore the detected continuum levels are the combination of the continuum emission in these two side bands while the line emission and absorption is detected in one side band only. Correction for the relative gain of each mixer in the two side bands is included in the calibration pipeline. The HIFI calibration accuracy is better than 10%. As the foreground emission is extended relative to the HIFI beam at 1.9 THz, we present all data in this paper in units of antenna temperature denoted T_A . As an example of the resulting spectra and maps, the data toward W49N are shown in Fig. 1 and Fig. 2. The data for

¹ available at http://herschel.esac.esa.int/HIPE_download.shtml

² available at <http://www.iram.fr/IRAMFR/GILDAS/>

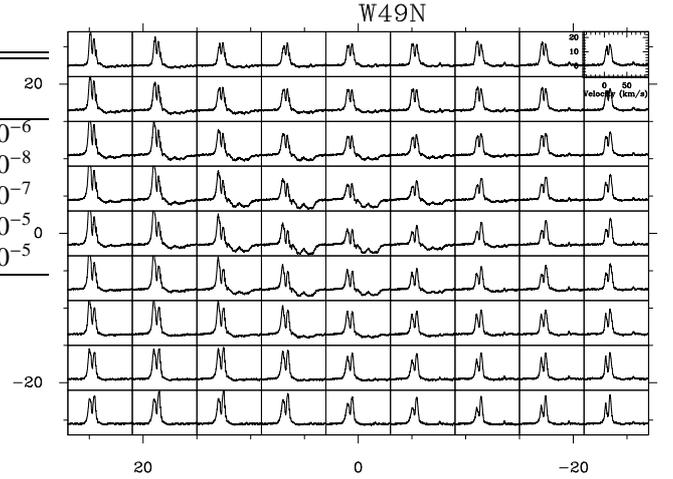


Fig. 2. Montage of [C II] HIFI spectra towards W49N. A baseline has been subtracted from each spectrum. In each box, the horizontal axis is the LSR velocity in km s⁻¹, which runs from -48 km s⁻¹ to 98 km s⁻¹ and the vertical axis is the antenna temperature in Kelvin, which runs from -8 to 24 K. The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

the other sources are presented in Appendix B. For each source, we display the spectrum observed in load chop mode toward the central position with the correct SSB continuum level, as well as the mean spectrum toward the OFF positions, obtained as the difference between the load chop and DBS spectra. This shows the level of contamination by emission signal in the OFF beam. A linear baseline has been removed from the spectra.

We derived the continuum intensities from the DBS spectra. We used the HIFI efficiencies given in Roelfsema et al. (2012) to convert the continuum values from Kelvin to Jy, namely 464, 469, and 506 Jy/K at 0.49, 0.81, and 1.9 THz, respectively. The continuum values are listed in Table 3. For the [C II] data at

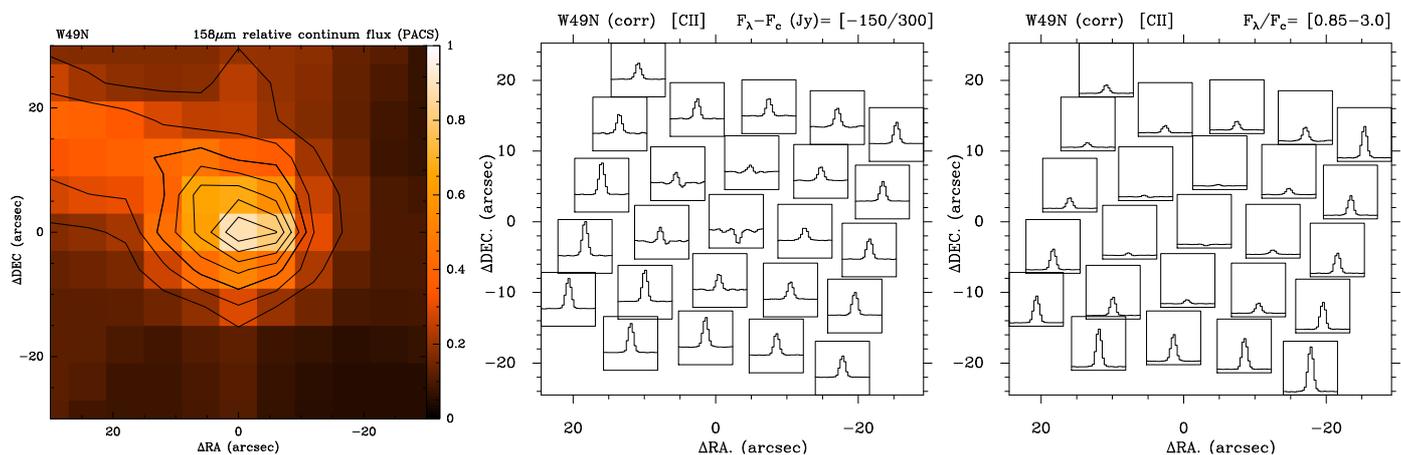


Fig. 3. PACS data towards W49N. For all maps, the offsets are given relative to the central position listed in Table 1. The data have been corrected from the weak emission in the OFF position. Left : Continuum emission at 158 μm . Contour levels are drawn at 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Middle : [C II] emission in the 25 PACS spaxels. The continuum emission has been subtracted. The vertical scale runs from -150 to 300 Jy. Right : Map of the line to continuum ratio in the 25 PACS spaxels. The vertical scale runs from 0.85 to 3.0. Note the very low contrast in the central spaxels.

1.9 THz, we provide a comparison with the continuum intensities measured in the central pixel of the PACS footprint.

2.3. PACS data processing

The PACS spectra presented in this work were obtained between 2010 and 2012 using the Range Spectroscopy SED mode. This mode is designed to cover the full PACS wavelength range in Nyquist sampling to get the complete far-IR Spectral Energy Distribution (SED). The observation ObsIDs are tabulated in Table A.1. The PACS spectrometer uses photoconductor detectors and provides 25 spectra over a $47'' \times 47''$ field-of-view (FoV) arranged as 5×5 spatial pixels (“spaxels”), each with an angular size of $\sim 9.4''$ on the sky (Poglitsch et al. 2010). The measured width of the PACS spectrometer point spread function (PSF) increases with wavelength for $\lambda \gtrsim 100 \mu\text{m}$. In particular, only $\sim 50\%$ ($\sim 70\%$) of a point source emission would fall in the central spaxel at $\sim 158 \mu\text{m}$ ($63 \mu\text{m}$). The resolving power in the R1 and B3A grating orders are $R \sim 1200$ ($\sim 250 \text{ km s}^{-1}$) at $158 \mu\text{m}$ and $R \sim 3300$ ($\sim 90 \text{ km s}^{-1}$) at $63 \mu\text{m}$. The total observing time was $\sim 1.4 \text{ h}$ per source in the R1 order and $\sim 1.3 \text{ h}$ in the B3A order.

Observations were carried out in the “chop-nod” mode with the largest chopper throw of ± 6 arc-min. This mode offers a reasonable measurement of the continuum level toward bright, compact star forming cores. In the observed regions, however, the reference nod positions are often contaminated by extended [C II] $158 \mu\text{m}$ and [O I] $63 \mu\text{m}$ line emission. Accurate calibration of the line intensities toward each source is thus not trivial and requires subtraction of the reference position emission. The continuum rms noise of the spectra toward each continuum peak position was typically $\sim 5\text{--}10 \text{ Jy/spaxel}$. Except for W33A (the faintest of the sample) the continuum peak in each source is a few thousand Jy/spaxel, close but still below the saturation levels of PACS detectors.

The PACS data were processed in HIPE 11 (Ott 2010). The ~ 158 , ~ 145 and $\sim 63 \mu\text{m}$ continuum levels in each observed position were extracted using the standard *ChopNodRangeScan* pipeline for chopped Range Spectroscopy observations. Table 3 shows a summary of the $158 \mu\text{m}$ continuum levels measured with PACS toward the continuum peak of each source. The ab-

solute flux calibration uncertainty is estimated to be $\sim 30\%^3$. Toward the peak positions, the continuum fluxes have been extracted using the PACS point source correction applied to the central spaxel fluxes. Owing to the relatively large distance to these sources, this is a good first approximation (specially for their bright and compact far-IR continuum peak). Indeed, the resulting fluxes are in very good agreement with those obtained from HIFI at the map center. Maps of the relative continuum emission with respect to the peak position were constructed for each observed source and are shown in Figure 3 and Figs B.3 to B.21. The PACS maps were re-centered with respect to the HIFI map ($0''$, $0''$) coordinates. In most sources, the ($0''$, $0''$) position agrees with the continuum peak.

In order to correct for the [C II] $158 \mu\text{m}$ and [O I] $63 \mu\text{m}$ line contamination in the reference nod positions we used the *ChopNodSplitOnOff* pipeline. For W49N, W51 and W31C, the [C II] $158 \mu\text{m}$ line emission in the A and B nod positions is almost the same ($\sim 10\text{--}30\%$ of the intrinsic ON source emission). The [O I] $63 \mu\text{m}$ line emission in the nod positions is typically $\sim 0\text{--}10\%$ of the ON source emission. For these sources we could correct for the extended emission by taking into account the line emission in each reference spaxel. The final spectral map thus contains the continuum emission plus the corrected [C II] $158 \mu\text{m}$ and [O I] $63 \mu\text{m}$ line emission/absorption. A polynomial baseline of order 3 was subtracted from each spectrum to obtain the continuum-subtracted data, and each spectrum was divided by the continuum to obtain the line/continuum spectra. Figure 3, and Figures B.3 to B.21 show the resulting [C II] maps for each source. We also display maps of the relative continuum emission at $158 \mu\text{m}$ (with respect to the peak position). Figure 4 shows the line to continuum maps for the [O I] $63 \mu\text{m}$ and $145 \mu\text{m}$ lines toward W49N.

3. Results

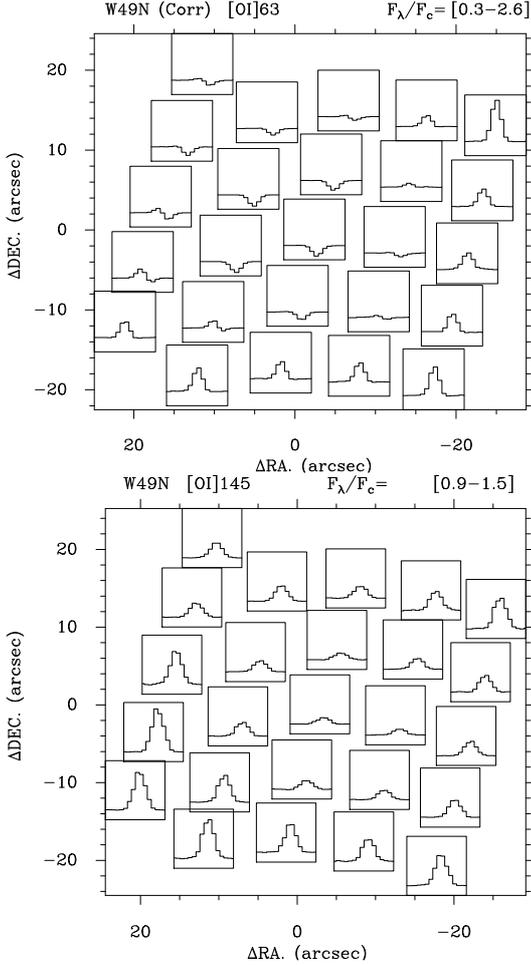
3.1. Spatial structure

We have made use of the information in both the PACS and HIFI data to investigate the spatial variations of the foreground absorption. The continuum emission was accurately measured at

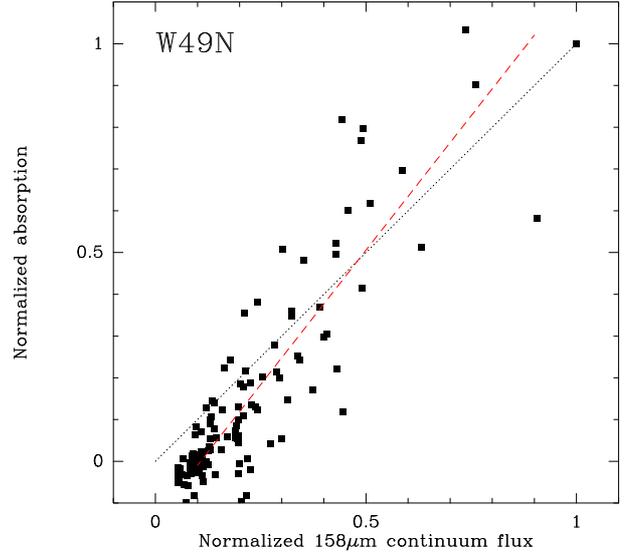
³ See PACS Spectroscopy performance and calibration,” PACS/ICC document ID P1CC-KL-TN-041, by B. Vandenbussche et al.

Table 3. Summary of continuum measurements

Source	T_{490} (K)	$S_{490}(\text{HIFI})^1$ (10^2 Jy/beam)	T_{809} (K)	$S_{809}(\text{HIFI})^2$ (10^2 Jy/beam)	T_{1900} (K)	$S_{1900}(\text{HIFI})^3$ (10^2 Jy/beam)	$S_{1900}(\text{PACS})^4$ (10^2 Jy/beam)
W28A	0.25 ± 0.02	1.16 ± 0.1	1.18 ± 0.1	5.50 ± 0.5	6.6 ± 0.5	33 ± 2.5	30 ± 4.5
W31C	0.39 ± 0.03	1.81 ± 0.14	1.96 ± 0.05	9.20 ± 0.2	6.7 ± 0.5	34 ± 2.5	33.8 ± 5.1
W33A	0.17 ± 0.01	0.79 ± 0.05	1.7 ± 0.2	8.6 ± 1.0	7.1 ± 1.1
G34.3+0.15	0.61 ± 0.02	2.83 ± 0.1	2.27 ± 0.1	10.6 ± 0.5	6.8 ± 0.5	34 ± 2.5	36.2 ± 5.4
W49N	0.7 ± 0.02	3.25 ± 0.1	2.45 ± 0.1	11.5 ± 0.5	8.6 ± 0.6	44 ± 3.0	39.4 ± 5.9
W51	0.8 ± 0.02	3.71 ± 0.1	2.9 ± 0.1	13.6 ± 0.5	7.6 ± 0.6	38 ± 2.5	39.2 ± 5.9
DR21(OH)	0.36 ± 0.02	1.67 ± 0.1	3.8 ± 0.5	19 ± 2.5	15.7 ± 2.4
W3-IRS5	0.2 ± 0.03	0.93 ± 0.1	4.0 ± 1.0	20.0 ± 5	23.7 ± 3.6

¹ using 464 Jy/K, ² using 469 Jy/K, ³ using 506 Jy/K, ⁴ using the point source calibration.

Fig. 4. Line to continuum ratio spectral maps obtained with PACS for the oxygen fine structure lines towards W49N. Top: The $^3P_1 - ^3P_2$ line at $63 \mu\text{m}$, Bottom: the $^3P_0 - ^3P_1$ line at $145 \mu\text{m}$. The vertical scale is indicated in each map.

the map center with HIFI. However, since the continuum information obtained in the HIFI map is not fully reliable, the HIFI map provides only the velocity integrated emission or absorption (in K km s^{-1}) at each position. We have therefore made use of the PACS spectral map data to determine the spatial variation of the far infrared continuum emission in each source. As the continuum emission is expected to vary smoothly at $158 \mu\text{m}$, we have interpolated the measured continuum fluxes in each PACS spaxel to obtain a far infrared continuum map of each source. Because this method is not expected to provide an accurate abso-


Fig. 5. Comparison of the velocity integrated absorption ($\int_{16.5}^{80} (T_A - T_C) dv$) measured in the HIFI map relative to the velocity integrated absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship.

lute calibration, we only make use of the spatial variation of the fluxes relative to the map center. Figure 3 (left panel) presents the resulting continuum map for W49N. Figure 5 presents the comparison of the velocity integrated absorption relative to the velocity integrated absorption at the map center for each pixel of the HIFI map. The relative integrated absorption and continuum emission (each normalized by value at the center of the map) show an excellent linear correlation, with a correlation coefficient of 0.88 for W49N and a slope close to unity. This behavior is expected for an extended and uniform foreground absorber. While small amplitude variations are not excluded given the limited signal to noise ratio, the diffuse gas along the line of sight behaves to first order as a uniform screen across the angular extent of the W49N background source. Therefore the molecular column densities derived from the ground state hydride absorption lines reported elsewhere with different angular resolutions are directly comparable to the present data. The same behavior is found for the other sources in the sample.

3.2. [C II] line to continuum ratio

All sources exhibit complex line profiles in the velocity ranges associated with the star forming regions, with multiple velocity

components, the possibility of self-absorption and in some cases broad line wings likely tracing UV-irradiated outflows, e.g. for W49N or for W28A (Fig. B.1, Gusdorf et al. (2014)). Nevertheless the velocity integrated [C II] signal from the massive star forming regions is always positive. W33A is the most extreme case since the [C II] emission almost vanishes at the HIFI map center as illustrated in Fig. B.8. The presence of diffuse material in the foreground leads to absorption of the background emission. At the PACS spectral resolution, the resulting [C II] signal in the central spaxel can appear either in absorption (towards W49N) or in emission (towards W51) depending on the relative strengths of the emission and absorbing signals. This shows the need for a sufficient spectral resolution to separate the different contributions to the [C II] signals in sources with a complex velocity structure.

The [C II] line is used as a tracer of the star formation rate in external galaxies, including distant systems at high redshifts (Stacey et al. 2010). We have used the present data to estimate the [C II] emission relative to the far infrared continuum in the three massive star forming regions W31C, W49N and W51. As explained above, the PACS [C II] and [O I] data could be corrected from the low level emission from the OFF position for these three sources only. The data are shown in Figure 6 in which each PACS spaxel is displayed as an independent point. For each source the PACS data where the contributions from the source and the foreground absorption are merged are shown as black squares, and the same quantity measured by HIFI as green stars. The HIFI data allow separation of the intrinsic source contribution (empty red squares) and the absorption from the foreground gas (filled blue squares with negative values). In all plots the horizontal scale is the product of the spaxel flux by the distance to the source, with the low values on the left side corresponding to the map edges and the high values on the right side to the continuum peak associated with the massive star forming region. The vertical scale is the relative [C II] emission $I_{rel}(\text{C II})$, defined as

$$I_{rel}(\text{C II}) = \frac{\int_{157.5}^{158.0} \frac{F(\lambda)}{F_c(\lambda)} d\lambda}{\int_{157.5}^{158.0} d\lambda} - 1. \quad (1)$$

To compute $I_{rel}(\text{C II})$ from the PACS observations we used a wavelength window of $\delta\lambda = 0.5 \mu\text{m}$ starting at $157.5 \mu\text{m}$ and ending at $158.0 \mu\text{m}$. This wavelength interval is broader than the intrinsic line widths and comparable to the full frequency bandpass of the HIFI band 7 receiver.

As both the line and continuum values are detected with the same instrument, $I_{rel}(\text{C II})$ is independent of calibration errors or uncertainties in the PSF. The ratio of the [C II] line luminosity $L([\text{C II}])$, to the far infrared luminosity L_{FIR} can be simply derived from $I_{rel}(\text{C II})$ and depends weakly on the shape of the far infrared spectral distribution of the source. Using infrared emission spectra computed with the Draine & Li dust model (Draine & Li 2007) and choosing the parameters appropriate to our sources, we determine

$$\frac{L([\text{C II}])}{L_{FIR}} \sim 3.1 \pm 0.35 \times 10^{-3} I_{rel}(\text{C II}). \quad (2)$$

The data for W31C, W49N and W51 show an excellent agreement between the unresolved data from PACS and the velocity resolved HIFI data. The dip in [C II] detected in the PACS maps towards W49N and W31C is fully explained by the foreground absorption. These plots also confirm the constancy of the absorption across the extent of the PACS map as the blue points fall nearly on a horizontal line. The slight decrease of the absorption

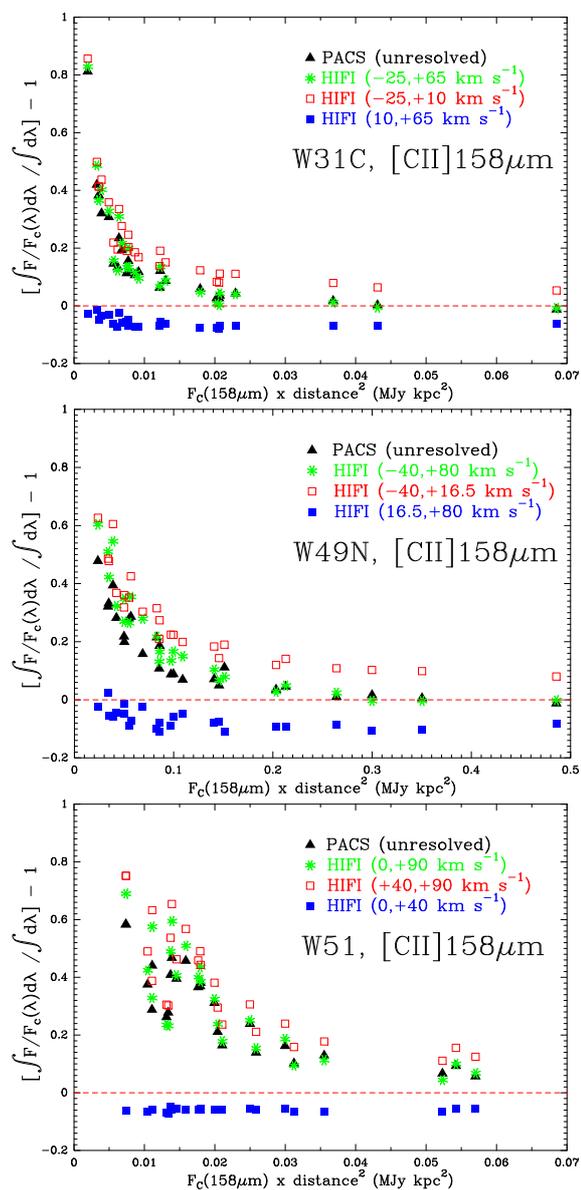


Fig. 6. Relative [C II] emission as a function of the $158\mu\text{m}$ continuum luminosity for W31C, W49N and W51. For each source, the black triangles show the spectrally unresolved PACS data, the blue squares the contribution of the foreground absorption measured by HIFI, the red squares the intrinsic source emission measured by HIFI, and the green stars the total emission and absorption measured by HIFI. The velocity intervals for the integration are given in each plot. In all plots the points on the left side correspond to the map edges while those on the right side correspond to the massive star forming regions associated with the dust continuum peak.

at lower continuum flux is likely to reflect the competition between the weak emission filling in the absorption at lower continuum fluxes, as expected from the physical conditions along these sight-lines (see below).

It is also interesting to notice that the intrinsic emission from the massive star forming regions presents a strong spatial variation of their relative [C II] emission, from $I_{rel}(\text{C II}) \sim 0.5$ at the map edges down to $I_{rel}(\text{C II}) \sim 0.06$ at the peak of the far infrared emission, or from $L([\text{C II}])/L_{FIR} \sim 1.5 \times 10^{-3}$ down to $L([\text{C II}])/L_{FIR} \sim 2 \times 10^{-4}$. This behavior is reminiscent of the deficit in [C II] emission observed in compact and luminous infrared galaxies such as Arp 220 or NGC 4418

(Graciá-Carpio et al. 2011) with deeply embedded star forming regions and possible Active Galactic Nuclei. As for massive star forming regions, the heating sources are deeply embedded in large dust column densities. Therefore the emerging signal in the fine structure lines is heavily dependent on the geometry of the heating radiation relative to the location of the dense molecular gas, and the emergent intensity can be severely attenuated. The presence of absorption by a foreground layer, as in the PRISMAS sources, will also contribute to dimming the emergent [C II] signal, especially at low spectral resolution.

3.2.1. Diffuse ISM absorption toward NGC 4418

To illustrate the effect of foreground absorption, we have used the recent H I absorption observations towards the NGC 4418 galaxy nucleus (Costagliola et al. 2013) to evaluate the [C II] line opacity associated with the foreground absorption. The HI spectrum presents two components with total column densities of $5.8 \times 10^{21} \text{ cm}^{-2}$ and $1.5 \times 10^{22} \text{ cm}^{-2}$. Assuming that the gas producing the HI absorption is diffuse, it is expected to produce an absorption signal in [C II] as well. We can estimate the magnitude of this absorption from the HI column densities using a carbon to hydrogen abundance ratio of $x_C = 1.4 \times 10^{-4}$. These numbers lead to expected integrated line opacities of $\sim 5.8 \text{ km s}^{-1}$ and $\sim 15 \text{ km s}^{-1}$ for the [C II] line associated with each velocity component. To further estimate the impact of this absorption, we use the measured line widths of 130 and 160 km s^{-1} to derive the relative emission $I_{rel}(\text{C II})$ as defined in Eq.1, and obtain -0.04 and -0.11. Note that the figures are negative because they correspond to absorption. We therefore obtain comparable figures for the absorption caused by the diffuse gas along the line of sight to massive star forming regions, and for the diffuse gas in the foreground of NGC 4418 nucleus. This simple example shows that the interpretation of low [C II] emission, or even of the non detection of [C II], in luminous infrared galaxies is difficult. The presence of absorption lines, either from HI in the centimeter domain, or molecular lines in the far infrared domain should be used as a warning for the possibility of a [C II] absorption, which would affect the emission from the dense star forming regions. In that case the use of classical PDR modeling to interpret the detected signal should be done with caution. A detailed knowledge of the source geometry and its environment is needed to obtain meaningful results.

3.3. Atomic Oxygen fine structure lines

While the discussion is focused on [C II], it should be noted that the presence of foreground absorption is expected to affect the [O I] line at 63 μm even more severely than [C II], as this line is expected to have a larger opacity for the same total hydrogen column density. Indeed, using a carbon to hydrogen ratio of $x_C = 1.4 \times 10^{-4}$, an oxygen to hydrogen ratio of $x_O = 3.1 \times 10^{-4}$ (Sofia et al. 2004), the spectroscopic parameters listed in Table 2, and assuming low excitation as appropriate for diffuse gas, we obtain $\tau([\text{O I}_{63}]) = 1.6 \tau([\text{C II}])$.

Spectrally resolved line profiles of the [O I] 63 μm line are clearly needed to accurately evaluate the impact of foreground absorption, or self-absorption in the sources themselves.

Using the PACS data shown in Fig 4, we have computed the relative emission of the [O I] fine structure lines at 63 μm and 145 μm , as a function of the continuum luminosity at the same wavelength. The relative emission for both oxygen lines is de-

finied as in Eq.1. ⁴ In a similar manner as for [C II], the relative [O I] line emission can be related to the ratio of the the line luminosity to the far infrared luminosity ratios, but with different scalings. We obtain

$$\frac{L([\text{O I}_{63}])}{L_{FIR}} \sim 10^{-3} I_{rel}(\text{O I}_{63}) \quad (3)$$

and

$$\frac{L([\text{O I}_{145}])}{L_{FIR}} \sim 3.1 \times 10^{-3} I_{rel}(\text{O I}_{145}). \quad (4)$$

Figure 7 presents the relative emission of the [O I] lines for the W31C, W49N and W51 sources. In this figure the black triangles show the data for the ground state line at 63 μm and the gray symbols those for the excited line at 145 μm . As expected from the [C II] behavior, and the larger opacity of the [O I] 63 μm line, the oxygen maps show an even stronger effect of the absorption by the foreground gas than do the [C II] maps. To evaluate the effect of the foreground gas and restore the intrinsic emission from the background source, we have estimated the absorption caused by the foreground gas using the mean value of the [C II] absorption, scaled by 1.6, the ratio of the line opacities for the adopted carbon and oxygen abundances. It is shown as a blue line in Fig 7. The corrected points are shown as empty red squares. While the original PACS [O I] data at 63 μm present either a pure absorption (negative $I_{rel}(\text{O I}_{63})$), or a weak [O I] emission comparable to that of the [O I] 145 μm line, the corrected points are always positive, with $I_{rel}(\text{O I}_{63})$ stronger than $I_{rel}(\text{O I}_{145})$. For the massive sources, $I_{rel}(\text{O I}_{63})$ ranges from ~ 0.4 at the edge of the observed regions with moderate far infrared emission down to ~ 0.1 at the peak of the far infrared emission. $L(\text{O I}_{63})/L_{FIR}$ thus decreases from $\sim 4 \times 10^{-3}$ down to $\sim 10^{-4}$.

This shows that in a situation where a bright source is observed through a foreground of diffuse gas, the use of the [O I] 63 μm line alone for quantitative models is not recommended either. Line ratios comparing the two fine structure lines of oxygen at 63 and 145 μm , or [O I] with [C II] will need to be corrected for absorption before being of quantitative use. Luminous infrared galaxies showing a [C II] deficit are also expected to be deficient in the 63 μm [O I] line if the deficit is caused by a foreground layer of diffuse gas.

3.4. C⁺ Column densities

The spectroscopic parameters of the detected lines are summarized in Table 2. The C⁺ column density of the diffuse gas can be obtained from the velocity integrated absorption through the relationship

$$N(\text{C}^+) = \frac{8\pi\nu^3 Q(T_{ex})}{g_u A_{ul} c^3} \frac{1}{1 - \exp(-\frac{h\nu}{k_B T_{ex}})} \int \tau([\text{C II}]) dv, \quad (5)$$

where $Q(T_{ex})$ is the partition function at the excitation temperature T_{ex} ($Q(T_{ex}) = g_l + g_u e^{-h\nu/k_B T_{ex}}$), A_{ul} is the Einstein coefficient listed in Table 2, and ν is the line frequency. As discussed by Goldsmith et al. (2012), the excitation temperature of [C II] is significantly lower than the kinetic temperature in the diffuse

⁴ To compute I_{rel} from PACS observations we used a wavelength window of $\delta\lambda = 0.5 \mu\text{m}$ for [C II], 0.45 μm for [O I₁₄₅] and 0.2 μm for [O I₆₃], with the following lower and upper limits: 157.50 – 158.00 μm for [C II], 63.28 – 63.08 μm for [O I₆₃], and 145.30 – 145.75 μm for [O I₁₄₅]. These wavelength intervals are broader than the intrinsic line widths.

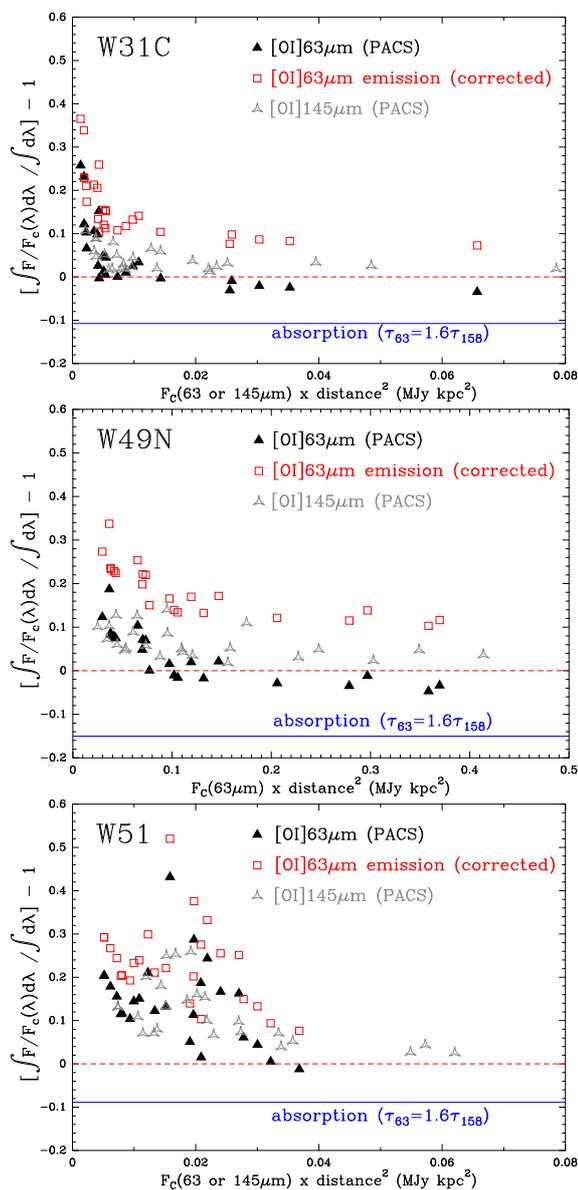


Fig. 7. As for Figure 6, but for the two [O I] fine structure transitions. The blue line shows the estimated level of absorption (negative relative emission) caused by the diffuse gas along the line of sight. It is obtained from the mean of the [C II] absorption scaled by a factor 1.6, the ratio of the [O I] 63 μm and [C II] line opacities for the adopted carbon and oxygen abundances.

gas because the densities are too low to produce a collisional excitation rate comparable to the spontaneous decay rate. For T_{ex} lower than ~ 25 K, the above expression can be written

$$N(\text{C}^+) (\text{cm}^{-2}) = 1.4 \times 10^{17} \times \int \tau([\text{C II}]) dv (\text{km s}^{-1}). \quad (6)$$

The above limit in the excitation temperature corresponds to a fraction of C^+ ions in the ground state larger than 95%. At higher excitation temperatures, equation 6 underestimates the C^+ column density.

In the diffuse interstellar gas, C^+ is the gas phase carbon reservoir. The total gas column density can therefore be derived, assuming that the gas phase carbon abundance relative to hydrogen is constant. Using the value derived by Sofia et al. (2004),

1.4×10^{-4} , we obtain :

$$N(\text{H}) (\text{cm}^{-2}) \sim 10^{21} \int \tau([\text{C II}]) dv (\text{km s}^{-1}). \quad (7)$$

This formula shows the excellent sensitivity of the absorption measurements that allow the detection of low column density regions, $\sim 10^{20} \text{cm}^{-2}$ corresponding to 0.05 magnitudes of visible extinction.

3.5. Atomic carbon column densities

For neutral carbon, most spectra show pure emission, the sole exception being the 40 km s^{-1} velocity feature in the $^3P_2 - ^3P_1$ spectrum towards W49N. In LTE the atomic carbon column densities can be derived from the expression appropriate for emission,

$$N(\text{C}) (\text{cm}^{-2}) = 1.4 \times 10^{16} \int T_{1-0} dv (\text{K km s}^{-1}), \quad (8)$$

where T_{1-0} is the main beam temperature of the ground state fine structure line at 492 GHz. This formula assumes that the emission is optically thin and the level populations can be described by a single excitation temperature, T_{ex} . It is accurate within 10% for $16 \text{ K} < T_{ex} < 120 \text{ K}$. The column density is underestimated for higher or lower values of T_{ex} , especially at the lower end where the populations of the excited levels become very small. This may be the case for instance in low density diffuse neutral gas. Therefore we performed statistical equilibrium calculations to derive the atomic carbon column densities, using the column density derived from the LTE formula as a lower limit.

The velocity feature at $\sim 40 \text{ km s}^{-1}$ towards W49N deserves a specific analysis because it shows up in absorption in the $^3P_2 - ^3P_1$ line at 809 GHz, but is not detected neither in emission nor in absorption in the $^3P_1 - ^3P_0$ line at 492 GHz. This behavior corresponds to the very special case where the excitation temperature of the ground state line is equal to the apparent continuum value of the background source. At 809 GHz, the rise of the thermal dust continuum emission results in a higher continuum temperature than the line excitation temperature, hence the appearance of an absorption line.

The radiative transfer equation for a plane parallel isothermal cloud located in front of a background continuum source producing a continuum antenna temperature T_c can be written

$$T = T_c e^{-\tau} + J(T_{ex})(1 - e^{-\tau}), \quad (9)$$

where $J(T_{ex}) = \frac{h\nu}{k_B} \frac{1}{e^{h\nu/k_B T_{ex}} - 1}$. In this formula the first term represents the absorption of the background radiation and the second one the emission of the plane parallel foreground cloud. These two terms can partially cancel each other, with the emission filling in the absorption. In the situation that $J(T_{ex}) = T_c$, equation 9 yields $T = T_c$ and the foreground cloud becomes invisible. For W49N, the continuum value at 492 GHz, expressed in main beam temperature units, is $0.9 \pm 0.03 \text{ K}$. This translates to an excitation temperature of $T_{ex} = 7 \pm 0.2 \text{ K}$. Such a low excitation temperature is consistent with the presence of cold molecular gas, in agreement with the detection of deep absorption in CN, NH_3 and H^{13}CO^+ (Godard et al. 2010; Persson et al. 2012). Indeed, as discussed by Tatematsu et al. (1999), the 492 GHz [C I] line is almost thermalized at the edges of dense molecular cores. The carbon column density for this particular feature is derived from the $^3P_2 - ^3P_1$ line using $T_{ex} = 7 \text{ K}$, $N(\text{C}) = (2.5 \pm 0.2) \times 10^{17} \text{ cm}^{-2}$. This large value is consistent with the strong H^{13}CO^+ absorption (Lucas & Liszt 2000) and confirms that this feature is associated to translucent material ($A_V \sim 3 - 5 \text{ mag}$) rather than diffuse gas ($A_V \sim 1 \text{ mag}$).

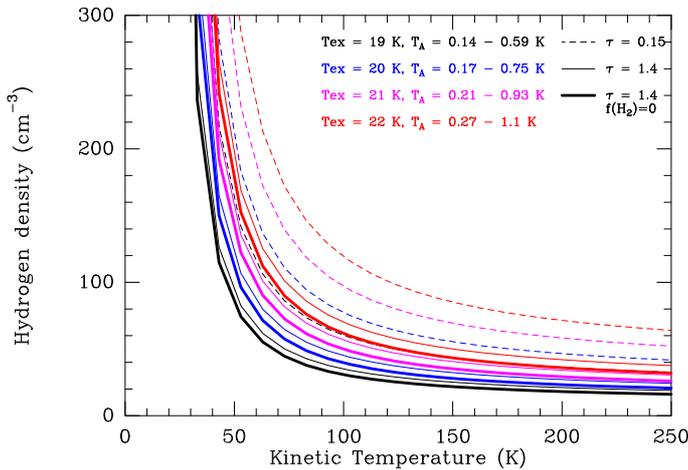


Fig. 8. Contours of constant excitation temperature of the [C II] transition in the diffuse neutral medium with values typical of those deduced here (indicated at the top of the figure). The dashed lines show the results for an optically thin line ($\tau = 0.15$), while the solid lines are for a mildly optically thick line ($\tau = 1.4$). For curves with normal weight lines, $n(\text{HI})=n(\text{H}_2)$, while the results for a pure atomic medium are given in bold lines; we see that the form of hydrogen makes little difference in terms of [C II] excitation.

3.6. Physical conditions from the [C II] data

The combination of the emission and absorption allows us to derive the excitation conditions of ionized carbon since it provides both the line opacity and peak temperature. Following Goldsmith et al. (2012) we have derived the excitation temperature for the velocity components detected both in absorption and in the OFF emission spectra. The well-known equations in the case that the foreground gas is uniform in terms of excitation temperature and optical depth are:

$$T_A(\text{ON}) = T_c e^{-\tau} + J(T_{ex})(1 - e^{-\tau}) \quad (10)$$

$$T_A(\text{OFF}) = J(T_{ex})(1 - e^{-\tau}) \quad (11)$$

$$\Delta T_A = T_A(\text{ON}) - T_A(\text{OFF}) = T_c e^{-\tau} \quad (12)$$

where T_c and $J(T_{ex})$ are defined above. ON - OFF spectrum provides the line opacity. This can then be used to analyze the OFF emission spectrum, which yields T_{ex} since the opacity is known. Given the large separation of the C^+ fine structure levels, 92 K, the modest density and column density of the diffuse medium produces only weak emission, typically at the 0.5 K level. The OFF position spectra displayed in the bottom panel of Fig. 1 and Fig B.1 to B.19 indeed show emission at this level in the same velocity range as the prominent absorption detected in the ON source spectra.

The derived optical depths and excitation temperatures are given in Table 4. Despite the range in [C II] opacities, the excitation temperatures cluster around the median value of $T_{ex} = 20$ K and a mean value of $T_{ex} = 20.4 \pm 1.2$ K. The weak feature towards W3-IRS5 has been omitted from this average as the S/N ratio is poor and the accuracy of the opacity determination is lowered by the presence of base line undulations.

The derivation of the excitation temperature relies on several hypotheses. In particular it assumes that the same component is responsible for the emission and absorption. It therefore ignores possible contribution of the warm ISM phases, the Warm Neutral Medium (WNM) and the Warm Ionized Medium (WIM) to the

[C II] signals. Although weaker than the CNM features, emission from the WIM has been detected by *Herschel* (Velusamy et al. 2012) at a level comparable to the features discussed here. The WNM is not expected to produce emission features given the low densities and electron fraction, but could contribute to the broad absorption features, and therefore lead to an overestimate of the line opacity. In this case, the excitation temperatures could be slightly underestimated.

The expected opacity from a WNM feature can be estimated from the average physical conditions of this phase as derived by Heiles & Troland (2003). The mean density $\sim 0.3 \text{ cm}^{-3}$, and the average HI column density per kpc of $N(\text{WNM}) = 8.7 \times 10^{20} \text{ cm}^{-2}$. With a C^+ fractional abundance of 1.4×10^{-4} , the corresponding C^+ column density is $N(\text{C}_{\text{WNM}}^+) \sim 1.2 \times 10^{17} \text{ cm}^{-2} \text{ kpc}^{-1}$, corresponding to an integrated line opacity of $\int \tau([\text{C II}]) dv \sim 0.9 \text{ km s}^{-1} \text{ kpc}^{-1}$. The line width is constrained by the typical velocity gradient along a sight-line through the Galactic Plane, about 10 km s^{-1} per kpc for the considered Galactic longitudes, and by the turbulent and thermal broadening, to be $\geq 10 \text{ km s}^{-1}$, hence the maximum opacity of each feature is $\tau_{\text{max}} \leq 0.09$.

A similar calculation holds for the WIM, with an expected higher C^+ abundance given the partial destruction of dust grains in hot gas. This simple calculation therefore shows that the warm phases make only a minor contribution to the absorption, except for long path-lengths in velocity-crowded regions. They may contribute to the weak and broad absorption features detected between the deep and narrow features. This minor contribution with $\tau \leq 0.1$ will not affect the determination of the excitation temperature except for the most optically thin features.

Figure 8 presents loci of constant excitation temperature of the [C II] transition in the diffuse ISM taking into account collisions with H, H_2 and electrons. In these calculations, the electron fraction x_e is set to 3×10^{-4} (Draine 2011), about twice the adopted gas-phase elemental carbon abundance and two hypothesis have been made for the relative fractions of atomic and molecular hydrogen : pure atomic hydrogen or equal amounts of H and H_2 . For a typical kinetic temperature of 100 K, the mean excitation temperature derived above (20.4 K) corresponds to gas densities between 40 and 80 cm^{-3} depending on the [C II] line opacity and the fraction of hydrogen in molecular form. These conditions clearly correspond to those of the diffuse ISM and confirm that the bulk of the detected [C II] emission and absorption is consistent with diffuse gas.

We have used the HI Galactic Plane survey (McClure-Griffiths et al. 2009; Kalberla et al. 2005, 2010) and new JVLA data (Winkel et al. 2014) to obtain the gas kinetic temperature and have derived the volume densities ($n = n(\text{HI}) + n(\text{H}_2)$) from the C^+ excitation temperature using the relationship

$$n = \frac{\beta A_{ul}}{\gamma_{ul}} \frac{e^{T_0/T_{kin}}}{e^{T_0/T_{ex}} - e^{T_0/T_{kin}}} \quad (13)$$

In this formula, $T_0 = \frac{h\nu}{k_B}$, β is the escape probability, which can be estimated from the line opacity τ ($\beta = (1 - e^{-\tau})/\tau$), and γ_{ul} is the collisional de-excitation rate coefficient so that $C_{ul} = n\gamma_{ul}$. We neglect the effect of any background radiation field on the excitation as the CMB and other contributions in the diffuse ISM are likely unimportant at the frequency of the [C II] line. The [C II] excitation is computed using the collisional cross sections for atomic and molecular hydrogen, helium, and for electrons. We have assumed that all carbon is ionized with a gas-phase carbon abundance of $x_C = 1.4 \times 10^{-4}$ (Sofia et al. 2004). This

Table 4. [C II] Excitation temperatures deduced from the emission in the OFF spectra^a

Source	V_{LSR} (km s ⁻¹)	T_{peak} (K)	Width (km s ⁻¹)	Area (K km s ⁻¹)	τ	T_{ex} (K)
W28A	[15;25]	0.53 ± 0.2	2 ± 0.6	1.1 ± 0.4	0.7 ± 0.10	20.8 ± 1.4
W31C	[15;30]	0.56 ± 0.1	11 ± 4	6.7 ± 3	1.3 ± 0.06	19.5 ± 1.7
	[30;50]	0.63 ± 0.1	14 ± 4	9.6 ± 3	1.4 ± 0.06	19.8 ± 1.0
G34.3+0.15	[10;20]	0.43 ± 0.1	1.2 ± 0.5	0.54 ± 0.2	0.65 ± 0.06	20.0 ± 0.9
	[20;40]	0.56 ± 0.1	6.6 ± 1.2	3.9 ± 0.7	0.83 ± 0.06	20.4 ± 0.5
W49N	[10;50]	0.53 ± 0.2	3.0 ± 0.9	1.6 ± 0.5	0.4 ± 0.1	22.6 ± 0.5
	[50;70]	0.63 ± 0.1	17 ± 3.8	11.3 ± 1.7	1.3 ± 0.06	19.8 ± 1.6
W51	[1;10]	0.33 ± 0.15	12.2 ± 2.4	4.3 ± 0.8	0.58 ± 0.06	19.3 ± 0.2
	[16;35]	0.35 ± 0.15	1.9 ± 0.6	0.71 ± 0.3	0.34 ± 0.06	21.4 ± 2.0
DR21(OH)	[5;25]	0.70 ± 0.3	20.3 ± 2.4	14.6 ± 1.5	1.5 ± 0.09	19.9 ± 1.0
W3-IRS5	[-25;-15]	0.71 ± 0.36	4.7 ± 1.9	3.6 ± 1.7	0.56 ± 0.09	23.1 ± 1.6
	[-15;15]	1.1 ± 0.4	18.2 ± 3	21.2 ± 3	0.11 ± 0.09	41 ± 3.0

^a The optical depths are determined from the ON and OFF difference spectra with assumption of uniform foreground gas; see text.

leads to an electron abundance $x_e = 3 \times 10^{-4}$ taking into account all sources of electrons (Draine 2011). We used the collisional cross sections summarized by Draine (2011). The results are displayed in Table 5. They apply to the absorption signals, for which the atomic and molecular column densities have been derived (Godard et al. 2012). Atomic and molecular hydrogen have only slightly different [C II] collision rate coefficients (Wiesenfeld & Goldsmith 2014), so the molecular H₂ fraction is not very important for determining the excitation of the [C II] fine structure line.

The median density derived from the [C II] data is $n \sim 60 \text{ cm}^{-3}$ and the mean value, ignoring the W3-IR5 sight-line is $n = 56 \pm 15 \text{ cm}^{-3}$. The thermal pressure ranges from $p/k_B = 3.8 \times 10^3 \text{ K cm}^{-3}$ to $2.2 \times 10^4 \text{ K cm}^{-3}$ with a median value of $5.9 \times 10^3 \text{ K cm}^{-3}$. In addition to the gas density, Table 5 also lists the ionized carbon column density $N_{tot}(C^+)$ and the kinetic temperature derived from the H I data.

3.7. Physical conditions from the [C I] data

To analyze the C I observations, we used the offline version of RADEX, a radiative transfer code based on the escape probability approximation (van der Tak et al. 2007). For each velocity component, we compared the observed line intensities to predictions obtained from a grid of RADEX models covering a pre-defined set of input parameters. The observed quantities include, for each velocity component, the integrated intensity of the C I ($^3P_1 - ^3P_0$) line, as well as that of the ($^3P_2 - ^3P_1$) transition, whenever observed. In the most favorable cases, we thus could use both the ($^3P_1 - ^3P_0$) integrated intensity value and the ($^3P_1 - ^3P_0$) to ($^3P_2 - ^3P_1$) integrated intensity ratio. In cases where the $^3P_2 - ^3P_1$ line was not detected, we used the $^3P_1 - ^3P_0$ integrated intensity and a lower limit to the integrated intensity ratio ($^3P_1 - ^3P_0$) to ($^3P_2 - ^3P_1$). In cases where the ($^3P_2 - ^3P_1$) line was not observed, we used only the ($^3P_1 - ^3P_0$) integrated intensity. We also determined the line width, Δv , and the peak temperature for each velocity component, from Gaussian fits over the observed spectral range.

The radiative transfer calculations were performed using collisional rates with ortho and para H₂, H, H⁺, He and electrons (Schröder et al. 1991; Launay & Roueff 1977; Roueff & Le Boulot 1990; Staemmler & Flower 1991; Johnson et al. 1987). For the model grid, the equivalent temperature of the background radiation field was set to the microwave

background field, $T_{CMB} = 2.73 \text{ K}$. The grid covers a range of gas densities n from 10 to 1000 cm^{-3} , a range of kinetic temperatures T_{kin} from 40 to 200 K with eight values for n and six for T_{kin} , corresponding to the parameters of the diffuse gas. We also sampled the effect of the composition of the gas by computing models for three values of the fraction of hydrogen in molecular form $f(H_2) = 2n(H_2)/(2n(H_2) + n(H I)) = 0.1, 0.4, 0.9$, and of the total column density of carbon using eight values of the C column density: 0.5, 0.75, 1.0, 3.0, 6.0, 9.0, 20., $50. \times 10^{16} \text{ cm}^{-2}$. As for the [C II] excitation calculations, we assumed that C⁺ is the main reservoir of carbon in these clouds, hence the electron density is derived from the total density using the same carbon abundance relative to hydrogen as above: $n(e) = 1.4 \times 10^{-4}n(H I) + 2.8 \times 10^{-4}n(H_2)$. The helium density is set to 10% of the hydrogen density and the H⁺ density is set to its minimum value $2 \times 10^{-3} \text{ cm}^{-3}$. We have checked that increasing the electron density to $n(e) = 3 \times 10^{-4}(n(H I) + 2n(H_2)) \text{ cm}^{-3}$ and the proton density to $n(H^+) = n(e) - n(C^+) \text{ cm}^{-3}$ does not change the model predictions for neutral carbon fine structure lines. For each velocity component, the observed quantities were compared to the appropriate model results. The fits are reported in Table 6 and illustrated in Fig. 9. They make use of the kinetic temperature derived from the H I data. Since collisional excitation rates with atomic hydrogen are larger than those with molecular hydrogen, the derived physical parameters show a clear dependence on the fraction of gas in molecular form. We have used the available information on the fraction of hydrogen in molecular form to select the best fit models. For about half of the velocity components, the fits are not well constrained and a broad interval of gas pressures is consistent with the data, while the range of best fit gas pressures is relatively narrow for a second half of the models. The pressure distribution is bimodal with about two thirds of the velocity components with pressures values ranging between 10^3 and $1.5 \times 10^4 \text{ K cm}^{-3}$, and the remaining components with significantly larger pressures. As the excitation analysis of the carbon lines does not return a single pressure value but rather an interval of acceptable values, we have derived the range of pressures using the median values of the minimum and maximum pressure values for each velocity component. We derive a pressure range of $2200 - 7800 \text{ K cm}^{-3}$, which is consistent with the median pressure value derived from the analysis of the [C II] data.

Table 5. Physical conditions derived from the [C II] spectra

Source	V_{LSR}	Width	$\int \tau \delta v$	$N_{tot}(C^+)^a$	T_{kin}^b	$f(H_2)^c$	n^d	Length ^e	L	Filling factor ^f
	(km s ⁻¹)	(km s ⁻¹)	(km s ⁻¹)	(10 ¹⁸ cm ⁻²)	(K)		(cm ⁻³)	(pc)	(kpc)	(%)
W28A	[15;25]	4.9 ± 0.7	3.7 ± 0.5	0.58 ± 0.1	100	0.51	67 ± 10	24	1.1	2
W31C	[10,23]	10.0 ± 2.4	13.5 ± 3.0	1.9 ± 0.3	105	0.68	39 ± 5	85	1.5	6
	[23,34]	7.9 ± 1.8	12.9 ± 5	1.8 ± 0.4	100	0.74	38 ± 5	116	0.9	13
	[34,61]	11.8 ± 2.5	17.7 ± 3.7	2.5 ± 0.5	90	0.45	45 ± 5	107	1.3	8
G34.3+0.15	[8,20]	7.2 ± 1.5	5.0 ± 0.8	0.70 ± 0.1	95	0.52	63 ± 10	18	0.85	2
	[20,35]	7.5 ± 1.0	6.6 ± 0.7	0.92 ± 0.1	85	0.34	68 ± 10	30	1.0	3
W49N	[25,50]	12.8 ± 0.8	20.0 ± 0.9	2.8 ± 0.5	90	0.51	80 ± 10	43	1.8	2
	[50,80]	19.0 ± 1.0	27.0 ± 1.5	3.8 ± 0.3	100	0.59	38 ± 5	134	2.6	5
W51	[1,10]	4.5 ± 1	2.8 ± 0.3	0.39 ± 0.05	90	0.55	58 ± 5	13	0.7	2
	[10,16]	5.0 ± 1.0	1.4 ± 0.2	0.2 ± 0.05	150	0.19	56 ± 10	6	0.5	1
	[16,35]	10.6 ± 1	3.8 ± 0.5	0.53 ± 0.07	130	0.06	70 ± 5	11	1.5	1
DR21OH	[0,20]	6.1 ± 1.0	12.9 ± 1.5	1.8 ± 0.3	95	0.61	36 ± 5	81	0.9	9
W3-IRS5	[-25;-15]	1.3 ± 1.	0.78 ± 0.5	0.11 ± 0.06	40	0.52	560 ± 50	0.7	0.9	0.07

^a $N_{tot}(C^+)$ is the C⁺ column density in the diffuse gas along the line of sight to the background sources, computed for the given LSR velocity interval. ^b HI data from the Galactic Plane survey <http://www.astro.uni-bonn.de/hisurvey/profile/index.php>, and from Winkel et al. (2014). ^c $f(H_2) = \frac{2n(H_2)}{n(H_1)+2n(H_2)}$ is the fraction of hydrogen in molecular form. ^d $n = n(H_1) + n(H_2)$ is the gas density. ^e $s = \frac{N_{tot}(H)}{n} (1 - f(H_2))/2$ is the total size of the CNM absorption features in the considered velocity interval. ^f $ff = s/L$ is the filling factor of the CNM absorption features in the considered velocity interval.

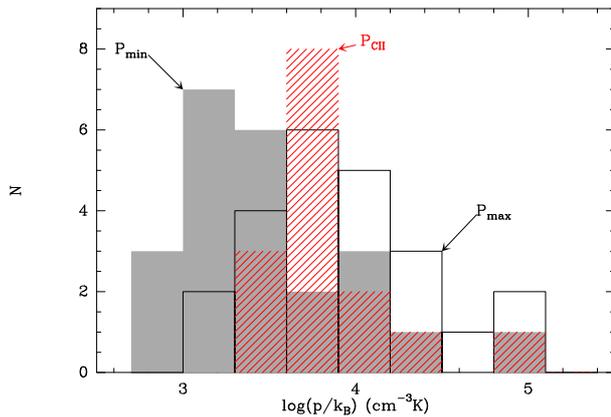


Fig. 9. Histograms of the pressure values derived from the analysis of the [C II] and [C I] data. The gray histogram represents the minimum values of the gas pressure derived from the [C I] data, the black lines show the maximum values also derived from [C I] and the dashed red histogram shows the pressure values derived from [C II].

4. Discussion

4.1. Gas pressure

Figure 9 presents histograms of the pressure derived from the [C II] and [C I] data. We show histograms of the minimum and maximum pressure compatible with the [C I] data given the wide range of acceptable solutions.

It is encouraging to see that the independent analysis of the C and C⁺ excitation lead to similar physical conditions for the diffuse interstellar gas in the Galactic plane. The median pressure derived from the velocity components detected in emission is $5.9 \times 10^3 \text{ K cm}^{-3}$ for [C II] and $2.2 - 7.8 \times 10^3 \text{ K cm}^{-3}$ for the [C I] velocity components. There is a slight tendency for n to decrease with increasing τ , as more opaque lines reach higher excitation temperatures for a given density. This shows that the emission data may be slightly biased towards the high pressure regions. In

addition, the regions with larger fractions of molecular gas tend to have slightly higher pressures than do the more atomic components. The pressures derived from the neutral carbon emission lines can be biased towards relatively dense regions because the ionization equilibrium is driven to higher C/C⁺ ratio at higher densities. Because of the more efficient collisional excitation of atoms and ions, higher pressure regions are also more likely to produce line emission than low pressure regions. The true median pressure may therefore be slightly lower than the median value, and closer to the lower bound of our pressure determination, $\sim 4.0 \times 10^3 \text{ K cm}^{-3}$.

The median density derived from the [C II] data is 60 cm^{-3} . These figures are very comparable to the physical conditions of the diffuse ISM in the solar neighborhood, determined from UV absorption measurements (Jenkins & Tripp 2011; Goldsmith 2013). The average pressure is also consistent with the theoretical predictions for a Galacto-centric radius of 5 kpc (Wolfire et al. 2003).

Interestingly these density and pressure conditions are also those of the low density gas responsible for the strong CH⁺ absorption features (Falgarone et al. 2010a; Godard et al. 2012). Models of the chemistry triggered by turbulent dissipation reproduce the observed CH⁺ column densities in the diffuse medium if the dissipation bursts occur in the CNM (Godard et al. 2009, 2014). The association with the CNM is further supported by the close similarity of the velocity profiles of the CH⁺ and [C II] absorption lines in the diffuse ISM (Falgarone et al. 2010b).

We conclude that the main phase responsible for the foreground absorption is the so-called cold neutral medium. The diffuse medium in this phase has a broad range of content of molecular gas, from nearly pure atomic gas up significant amounts of molecular gas (Godard et al. 2012; Winkel et al. 2014). The CNM is also responsible for the diffuse [C II] emission of the Galactic plane.

Table 6. Physical conditions from the [C I] data

Source	$V_{LSR}^{(1)}$ (km s ⁻¹)	Width ⁽¹⁾ (km s ⁻¹)	$\int T dv (1-0)^{(2)}$ (K km s ⁻¹)	$\int T dv (2-1)^{(2)}$ (K km s ⁻¹)	N(C) (10 ¹⁶ cm ⁻²)	n (cm ⁻³)	p/k_B (10 ³ K cm ⁻³)
W28A	20	4.0	4.40±1.20	< 3.30 ⁽³⁾	15 ± 5	20 - 200	0.7 - 11
	24	1.5	0.60±0.30	< 0.78 ⁽³⁾	1.8±1.3	40 - 1000	1.0 - 100
W31C	16.5	5.0	0.87±0.09	< 0.42 ⁽⁴⁾	2±1	40 - 400	2.2 - 34
	17.5	3.5	0.84±0.08	0.16±0.02	3 ± 0.5	40 - 70	2.2 - 3.4
	23	3.5	1.33±0.13	0.41±0.04	2 ± 0.5	200 - 400	11 - 20
	24.5	1.5	0.53±0.05	0.24±0.02	0.75 ± 0.5	400 - 1000	22 - 29
	27.5	2.5	2.70±0.27	0.63±0.06	6 ± 2	70 - 200	4.4 - 7.2
	30	2.5	2.30±0.23	0.59±0.06	4.5 ± 1	70 - 130	5.5 - 9.4
	34	5.0	2.00±0.20	0.43±0.04	10 ± 2	40 - 70	2.2 - 3.5
	37.5	4.0	1.59±0.16	< 0.48 ⁽⁴⁾	10±2	20 - 40	0.9 - 1.5
	41	2.5	0.83±0.08	< 0.12 ⁽⁴⁾	3±1	40 - 70	1.7 - 2.2
W33A	45	1.5	0.24±0.02	< 0.13 ⁽⁴⁾	0.75±0.25	40 - 200	1.7 - 8.4
	24	2.5	1.35±0.28	–	3 ± 1	40 - 200	2.7 - 7.8
G34.3	27	1.5	0.56±0.06	0.51±0.05	0.75±0.25	1000	90 - 100
	47.5	1.5	0.44±0.04	< 1.67 ⁽³⁾	0.75±0.25	130 - 200	9 - 13
W49N	51	1.5	0.07±0.02	< 0.07 ⁽⁴⁾	0.5±0.2	40 - 70	1.2 - 2.0
	54.5	3.5	0.35±0.04	0.06±0.01	1 ± 0.5	70 - 200	2.9 - 4.4
	59	2.5	0.51±0.05	0.80±0.03	2±1	40 - 200	1.5 - 7.5
	63	2.5	1.40±0.14	0.53±0.05	2±1	200 - 1000	12 - 26
	40 ⁽⁵⁾				30 ± 5	100 - 600	3.0 - 4.5
DR21(OH)	9	2.5	3.40±0.34	–	15 ± 5	20 - 130	0.65 - 4.5
	13.5	3.5	0.57±0.06	–	2.0±1	40 - 400	1.3 - 11
W3-IRS5	-21	1.0	0.11±0.01	–	0.5 ± 0.3	40 - 70	1.3 - 2.0

(¹) extracted from the (1–0) spectrum. (²) The error bars are dominated by the calibration uncertainty which is set to 10%. (³) upper limit due to blending and/or line confusion. (⁴) upper limit due to non detection. (⁵) velocity component detected in absorption in the ³P₂ – ³P₁ line (see section 3.5).

4.2. Filling factor

The knowledge of the median density and total gas column allows us to derive the average (1-dimensional) filling factor of this diffuse component along the sampled sight-lines. The length s and filling factor ff of the CNM absorption features are listed in Table 5. These figures have been derived using the total gas column $N_{tot}(\text{H})$, the mean fraction of gas in molecular form $f(\text{H}_2) = 2n(\text{H}_2)/[n(\text{H}_1) + 2n(\text{H}_2)]$, and the distance along the line of sight in a given velocity interval L as

$$s = \frac{N_{tot}(\text{H})}{n(\text{H}_1) + 2n(\text{H}_2)} = \frac{N_{tot}(\text{H})}{n} (1 - f(\text{H}_2)/2), \quad (14)$$

and

$$ff = \frac{s}{L}. \quad (15)$$

In these formulas, s represents the total length along the line of sight of each velocity feature and L is computed from the Galactic rotation curve $\Theta(R)$ assuming a flat rotation curve with $\Theta_0 = \Theta(R_0) = 220 \text{ km s}^{-1}$ and a distance to the Galactic Center, R_0 of 8.5 kpc. In the first Galactic quadrant and for $b = 0^\circ$, the distance along the line of sight can be written as

$$L = R_0 \cos(l) (-/+) \sqrt{R^2 - (R_0 \sin(l))^2}, \quad (16)$$

where the $-/+$ signs refer to the near and far distance branches and R is the distance to the Galactic center. R is determined from V_{LSR} assuming circular motions,

$$V_{LSR} = R_0 \sin(l) (\Theta(R)/R - \Theta_0/R_0). \quad (17)$$

For the studied lines of sight, the mean distance gradient produced by the rotation curve is about 100 pc per km s^{-1} . Since the velocity features we studied are broader than 1 km s^{-1} , they are likely not to be associated with a single and coherent entity but most likely result from the superposition of several individual ‘‘CNM cloudlets’’. The derived density and temperatures correspond to the average properties of these ‘‘CNM cloudlets’’. For all velocity intervals, the combined length along the line of sight of the ‘‘CNM cloudlets’’, s , is significantly smaller than the path-length along the sight-lines to the background sources L . The filling factors for the selected velocity intervals range from 0.07 to 13 % with higher filling factors in the largest opacity sight-lines, and a median value of 2.4%. These filling factors have been evaluated along paths of $\sim 1 - 2$ kpc length along the sight-lines to the background sources. The highest values may result from a selection bias as we chose sight-lines known for the presence of foreground absorption and adapted the velocity intervals to the line profiles. Note that the approach we followed to derive the filling factor leads to an upper limit as we assign all the absorbing gas to the CNM. Overall, these figures are consistent with the expected low volume filling factor of the CNM (Heiles & Troland 2003; Kalberla & Kerp 2009). These data therefore confirm that although most of the mass resides in the cold phases, these phases occupy a small fraction of the volume of the Galaxy.

The line profiles provide additional information on the spatial distribution of the CNM : for all sources, the [C II] line profiles show well defined velocity components, which correspond to localized regions in the Galactic plane. Therefore the diffuse medium is not uniformly distributed in space but spatially con-

centrated, most likely in the spiral arms as found in emission line surveys (e.g. Pineda et al. (2013) for [C II]).

4.3. Gas phase carbon abundance and C⁺ in the ISM phases

The C and C⁺ column densities listed in Tables 5 and 6 show that C⁺ dominates over atomic carbon in the diffuse gas, although a direct comparison between the velocity components present in emission and absorption is not straightforward. Nevertheless it is possible to compare the total C⁺ and C column densities along each sight-line. We find N(C)/N(C⁺) ranging from 1 to 28 %, with a median value of 6 %.

Because the gas-phase carbon abundance has been shown to spatially vary as a consequence of dust grain processing (Parvathi et al. 2012) we have investigated the gas-phase carbon abundance by comparing the derived C⁺ column densities with independent estimates of the hydrogen column densities derived from HI for the cold atomic gas and CH and HF for the diffuse molecular gas. Figure 10 shows the ratio of the C⁺ and total hydrogen column densities as a function of Galacto-centric radius. In this figure, the size of the symbol scales with the fraction of gas in molecular form $f(\text{H}_2)$. We do not find any strong effect with $f(\text{H}_2)$. The carbon abundance may be lower towards W3-IRS5 sampling the outer galaxy but the achieved S/N ratio is too low to draw a definitive conclusion. The median value is $\text{C}/\text{H} = 1.25 \times 10^{-4}$ and the weighted average $\text{C}/\text{H} = 1.5 \times 10^{-4}$. The right panel of Figure 10 illustrates how the *Herschel* values are fully consistent with the measurements derived from FUV absorption lines towards massive stars by Parvathi et al. (2012), especially comparing the same range of average hydrogen densities, $\langle n(\text{H}) \rangle = N_{\text{tot}}(\text{H})/L$, along the lines of sight, namely $\log_{10}(\langle n(\text{H}) \rangle) = [-0.3, 0.9]$.

There are several sources of uncertainties in this comparison, which contribute to the observed scatter :

- At large molecular hydrogen column densities, C and CO will make a significant contribution to the carbon budget. We do not expect this effect to be very strong because we do not see any significant effect with $f(\text{H}_2)$ and the total atomic carbon column density does not rise above 15% of the total C⁺ column density except towards W28A.
- The warm interstellar phases may be contributing to the [CII] absorption. Ionized carbon is expected to be present in the Warm Neutral Medium (WNM) and Warm Ionized Medium (WIM). Given the low densities, the excitation of the fine structure line will be moderate, leading to an absorption signal towards a strong continuum source. Additional information from other fine structure lines, [NII] at 205 μm and [OI] at 63 μm , is required to identify the relative contributions of these two warm phases. *Herschel* HIFI data of the [NII] at 205 μm line have been analyzed by Persson et al. (2014), and indeed show absorption towards W31C and W49N confirming the presence of the warm ionized medium along the same sight-lines. The WIM produces an excess [CII] absorption at the 7 - 10% level along both sight-lines.
- The estimations of the molecular hydrogen column densities are based on CH and HF measurements assuming a constant abundance relative to H₂. Optical and FUV measurements towards bright stars have shown that the relationship between CH and H₂ is linear and exhibits a significant scatter of 0.2dex (Sheffer et al. 2008), larger than the measurement uncertainties, and which probably reflects small variations in the physical conditions along the probed sight-lines. As shown by Levrier et al. (2012), averaging over long path-lengths tends to reduce the scatter, but local variations of the

CH abundance relative to H₂ cannot be excluded. The same can be true for HF as well.

Overall, the excellent agreement between the gas phase carbon abundance deduced from FUV and FIR data confirms the validity of the derivation of atomic and molecular gas column densities and supports the association of the [C II] absorption with CNM gas, with minor contributions from the WIM.

4.4. C⁺ Cooling

For an optically thin line, the cooling rate is proportional to the upper state population, and the cooling rate per C⁺ ion is given by $\Lambda(\text{C}^+) = hvA_{ul}g_u f_u$, where f_u is the fraction of the carbon ions in the upper level of the fine structure transition. The determination of the excitation temperature, T_{ex} , of the [C II] line allows to compute the C⁺ cooling of the diffuse medium. The low values from Table 4 indicate that only a few percent of the population is in the upper level. With the excitation temperature known, we can write the energy loss per H nucleon as

$$\Lambda(\text{C}^+) = \frac{hvA_{ul}g_u e^{-\frac{hv}{k_B T_{\text{ex}}}} N(\text{C}^+)}{Q(T_{\text{ex}}) N_{\text{tot}}(\text{H})}. \quad (18)$$

Taking the mean [C II] excitation temperature (Section 3.6), $T_{\text{ex}} = 20.4$ K, and assuming that all carbon is in ionized form so that the ratio of the C⁺ to H column densities is equal to x_C gives

$$\Lambda(\text{C}^+) = 9 \times 10^{-26} \frac{x_C}{1.4 \times 10^{-4}} \text{ erg s}^{-1} \text{ H}^{-1}. \quad (19)$$

Using a carbon abundance of 1.5×10^{-4} as derived from the present data, we obtain a cooling rate per H atom about three times larger than the mean value derived by Gry et al. (1992) for local lines of sight and five times larger than the cooling rate for low velocity clouds derived by Lehner et al. (2004) from a 43 object sample. The difference between the FUV and *Herschel* measurements can be understood as arising from the difference in the environments. The FUV data mostly target gas at high Galactic latitudes in the solar neighborhood while the far infrared *Herschel* data mainly probe the Galactic plane, where the star formation activity and the heating rate are expected to be larger. The *Herschel*-derived cooling rate is in very good agreement with the models for $n \sim 60 \text{ cm}^{-3}$ and Galacto-centric radii of 5 and 8.5 kpc presented by Wolfire et al. (2003), which include an increase of the heating rate in the inner Galaxy.

5. Conclusion

In this paper we have analyzed high-resolution spectra of the fine structure lines of ionized and neutral carbon, complemented by lower spectral resolution line maps of C⁺ and O, towards massive star forming regions in the Galactic plane.

- The high spectral resolution data exhibit complex line profiles from the background sources combined with absorption produced by the diffuse interstellar medium along the line of sight. We show that at low spectral resolution this combination leads to very low [C II] and [OI] signals, which can be wrongly interpreted. We suggest that part of the [C II] deficit in local luminous infrared galaxies is due to the presence of low density material in the foreground of the compact region producing the far-infrared emission. This diffuse gas component is expected to absorb part of the [C II] signal from the compact star forming region close to the nucleus. The effect

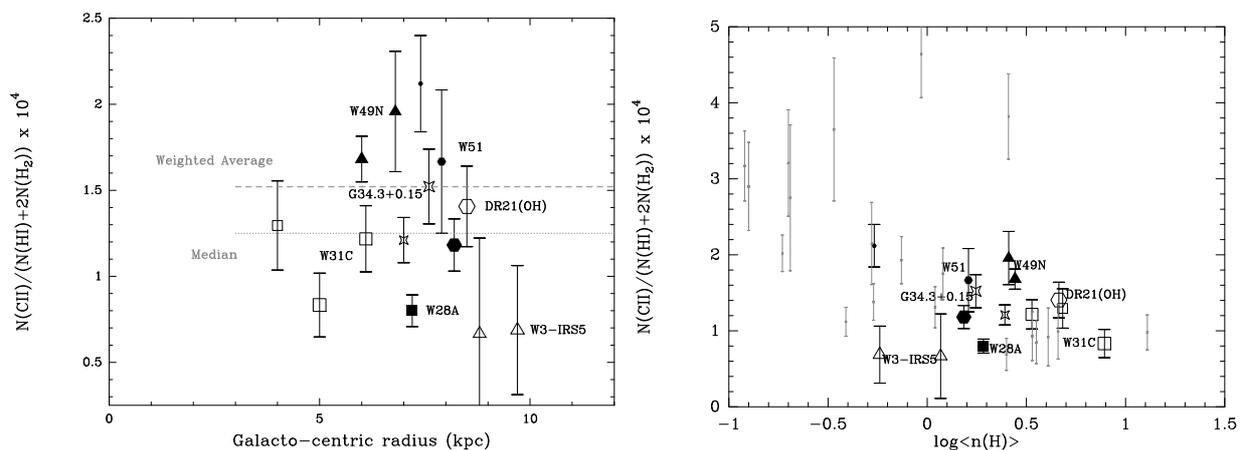


Fig. 10. Left : Ratio of the total C^+ column density to that of the total gas column measured by adding HI (from 21cm data) and H_2 (derived from CH and HF) as a function of Galacto-centric radius. The symbol size is proportional to the fraction of gas in molecular form ranging from ~ 0.1 to ~ 0.8 for the velocity intervals considered here, with each sight-line plotted with a different symbol. The weighted average (1.5×10^{-4}) is indicated with a dashed line and the median value (1.25×10^{-4}) with a dotted line. Right panel : comparison of the *Herschel* data (black symbols with source name) with the FUV absorption data from Parvathi et al. (2012) (gray dots with error bars). The gas-phase carbon abundance is plotted as a function of the logarithm of average gas density $\log_{10}(\langle n(H) \rangle) = \log_{10}(N_{tot}(H)/L)$.

is expected to be even larger for the $63 \mu\text{m}$ fine structure line of oxygen, which has a higher opacity for the same gas column density.

- The *Herschel* data have been combined with ground based HI data to derive the physical conditions in the diffuse gas. The median pressure is $p/k_B = 5900 \text{ K cm}^{-3}$ and the median density $n \sim 60 \text{ cm}^{-3}$, for a kinetic temperature of $\sim 100 \text{ K}$. These properties confirm that the absorbing gas can be associated with the so-called cold neutral medium (CNM). The fraction of hydrogen in molecular form in this gas is highly variable.
- The excitation conditions derived from the [C I] emission data are consistent with those deduced from [C II]. The atomic carbon column densities remain significantly smaller than the ionized carbon column densities, confirming that C^+ is the main carbon reservoir along the lines of sight in most cases.
- The line profiles show that the CNM is not uniformly distributed but follows the Galactic Spiral arms. The comparison of the gas densities and column densities has allowed us to determine the filling factor of the CNM along the lines of sight in the Galactic Plane. The derived median value of the filling factor is 2.4%. This value may be slightly biased upward by the source selection criteria.
- On average, the gas-phase abundance of carbon derived from the comparison of the [C II] absorption with the sum of the atomic and molecular hydrogen column densities, using CH and HF as tracers of H_2 , is in good agreement with previous measurements using FUV data, leading to $C/H \sim 1.5 \times 10^{-4}$. Some sight-lines may present a small excess [C II] absorption which indicates that two other ISM phases, the warm neutral medium and the warm ionized medium, may contribute to the [C II] absorption.
- The [C II] cooling per hydrogen nucleon in the Galactic plane is $\Lambda(C^+) \sim 9.5 \times 10^{-26} \text{ ergs}^{-1} \text{ H}^{-1}$, in good agreement with predictions from Wolfire et al. (2003).

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Appendix A: Presentation of the Herschel Observations

The individual *Herschel* observations analyzed in this paper are listed in Table A.1.

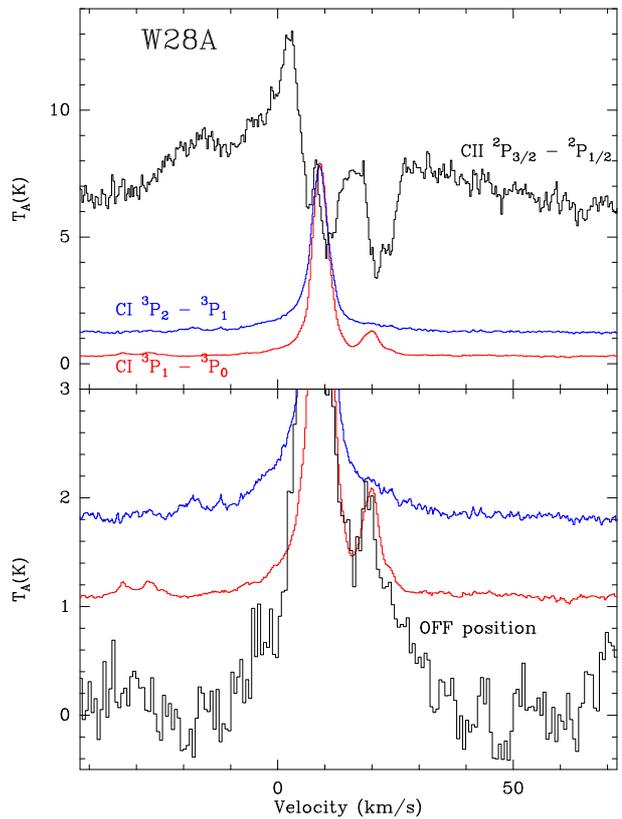


Fig. B.1. Top : Herschel/HIFI spectra towards W28A. The red line shows the $[C\text{I}]^3P_1 - ^3P_0$ line at 492 GHz, the blue line shows the $[C\text{I}]^3P_2 - ^3P_1$ line at 809 GHz and the black line the $[C\text{II}]^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz. The horizontal axis is the LSR velocity in km s^{-1} , and the vertical axis the antenna temperature in Kelvins. The continuum level for $[C\text{II}]$ corresponds to the SSB continuum level. Bottom : zoom on the $[C\text{I}]$ lines (red, blue as above), and average $[C\text{II}]$ spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

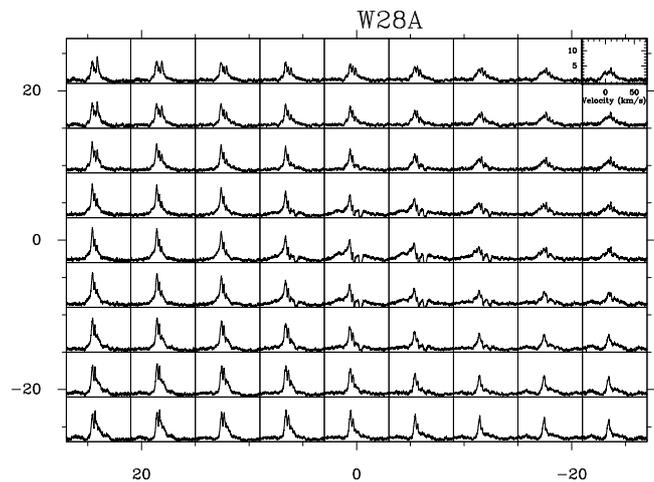


Fig. B.2. Montage of $[C\text{II}]$ spectra towards W28A. A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -42 km s^{-1} to 72 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -1 K to 14 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1

Table A.1. Summary of observations

Source	Line	Obs mode	ObsID
W28A	[C II]	DBS	1342243682; 1342243683
	[C II]	LC-OTF	1342243684
	[C I] $^3P_1 - ^3P_0$	DBS	1342216303; 1342216304; 1342216305
	[C I] $^3P_1 - ^3P_0$	LC	1342218938
	[C I] $^3P_2 - ^3P_1$	DBS	1342242856; 1342242857
	[C I] $^3P_2 - ^3P_1$	LC	1342242858
	[C II]	SED-R1	1342217940
	[O I] 63	SED-B3A	1342217941
W31C	[C II]	DBS	1342244100; 1342244101
	[C II]	LC-OTF	1342244102
	[C I] $^3P_1 - ^3P_0$	DBS	1342206657; 1342206658; 1342206659
	[C I] $^3P_1 - ^3P_0$	LC	1342251659
	[C I] $^3P_2 - ^3P_1$	DBS	1342244106; 1342244107; 1342244108
	[C I] $^3P_2 - ^3P_1$	LC	1342244109
	[C II]	SED-R1	1342217945
	[O I] 63	SED-B3A	1342228531
W33A	[C II]	DBS	1342229846; 1342229847
	[C II]	LC-OTF	1342229848
	[C I] $^3P_1 - ^3P_0$	DBS	1342208038; 1342208039; 1342208040
	[C I] $^3P_1 - ^3P_0$	LC	1342218943
	[C II]	SED-R1	1342239714
	[O I] 63	SED-B3A	1342239713
G34.3+0.15	[C II]	DBS	1342244578; 1342244579
	[C II]	LC-OTF	1342244580
	[C I] $^3P_1 - ^3P_0$	DBS	1342219189; 1342219190; 1342219191
	[C I] $^3P_1 - ^3P_0$	LC	1342231776
	[C I] $^3P_2 - ^3P_1$	DBS	1342244110; 1342244111; 1342244112
	[C I] $^3P_2 - ^3P_1$	LC	1342244113
	[C II]	SED-R1	1342209734
	[O I] 63	SED-B3A	1342209733
W49N	[C II]	DBS	1342231449; 1342231450
	[C II]	LC-OTF	1342231451
	[C I] $^3P_1 - ^3P_0$	DBS	1342206661; 1342206662; 1342206663
	[C I] $^3P_1 - ^3P_0$	LC	1342244802
	[C I] $^3P_2 - ^3P_1$	DBS	1342230253; 1342230254; 1342230255
	[C I] $^3P_2 - ^3P_1$	LC	1342230256
	[C II]	SED-R1	1342207774
	[O I] 63	SED-B3A	1342207775
W51	[C II]	DBS	1342231446; 1342231447
	[C II]	LC-OTF	1342231448
	[C I] $^3P_1 - ^3P_0$	DBS	1342219179; 1342219180; 1342219181
	[C I] $^3P_1 - ^3P_0$	LC	1342232813
	[C I] $^3P_2 - ^3P_1$	DBS	1342230243; 1342230244; 1342230245
	[C I] $^3P_2 - ^3P_1$	LC	1342230246
	[C II]	SED-R1	1342193698
	[O I] 63	SED-B3A	134219697
DR21(OH)	[C II]	DBS	1342220468; 1342220469
	[C II]	LC-OTF	1342220470
	[C I] $^3P_1 - ^3P_0$	DBS	1342198320; 1342198321; 1342198322
	[C I] $^3P_1 - ^3P_0$	LC	1342219171
	[C II]	SED-R1	1342209401
	[O I] 63	SED-B3A	1342209400
W3-IRS5	[C II]	DBS	1342248401
	[C II]	LC-OTF	1342248400
	[C I] $^3P_1 - ^3P_0$	LC	1342247845
	[C II]	SED-R1	1342229092
	[C II]	SED-B3A	1342229093

DBS stands for Double Beam Switch, LC for Load chop and LC-OTF for On The Fly map in Load Chop mode. SED-R1 refers to the SED mode for the red channel of the PACS spectrometer, SED-B3A refers to the blue channel of the PACS spectrometer.

Appendix B: Pointed data and spectral maps towards the individual sources

Appendix B.1: W28A

W28A (G5.89-0.4) is a massive star formation region at a distance of 1.3 kpc (Motogi et al. 2011). The HIFI spectra are shown in Figure B.1 and B.2. W28A hosts a powerful outflow with prominent velocity wings extending from -40 to 60 km s⁻¹ (Gusdorf et al. 2014). The foreground gas is associated with the absorption feature at ~ 20 km s⁻¹ in the [C II] spectrum, and with the emission line at the same velocity in the [C I] $^3P_1 - ^3P_0$ spectrum. This gas is probably local. The PACS spectral maps data are presented in Figure B.3.

Appendix B.2: W31C

W31C(G10.6-0.4) is a massive star forming region located in the Norma arm, at a distance of 5 kpc (Sanna et al. 2014). The foreground material samples the Sagittarius and Scutum-Centaurus arms, in addition to local material. The source velocity is about -2 km s⁻¹ while the foreground material is detected between ~ 10 and ~ 50 km s⁻¹. The HIFI spectra are shown in Figures B.4 and B.5. The PACS spectral maps near $158\mu\text{m}$ are presented in Figure B.6.

Appendix B.3: W33A

W33A(G12.9-0.3) is a highly embedded massive star at a distance of 2.4 kpc (Immer et al. 2013), with a massive cold envelope. A weak absorption from local foreground gas has been detected in CH⁺ by Godard et al. (2012) but the sensitivity of the present data does not allow to detect the associated [C II] feature. The HIFI spectra are shown in Figure B.7 and B.8. The [C II] emission from W33A is weak. It is barely detected at the continuum peak but increases as one moves away from the continuum peak. Broad line wings are detected in [C II], probably associated with a molecular outflow. The [C I] $^3P_1 - ^3P_0$ data show multiple weak emission lines from line of sight diffuse gas, including features located further away from W33A. The PACS spectral maps near $158\mu\text{m}$ are presented in Figure B.9.

Appendix B.4: G34.3+0.15

G34.3+0.15 is an ultra compact HII region located at about 3.3 kpc (Kuchar & Bania 1994). The line of sight crosses the Sagittarius arm at velocities between ~ 20 and ~ 35 km s⁻¹, and also samples the local interstellar medium for LSR velocities lower than 20 km s⁻¹. The HIFI spectra are shown in Fig. B.10 and B.11. The [C II] line is strongly self-absorbed towards G34.3+0.15 and presents an asymmetric spatial distribution. The [C I] $^3P_1 - ^3P_0$ emission is weak in the foreground gas. The tiny absorption at ~ 15 km s⁻¹ is due to a low level contamination in the reference position used in the Load Chop observations. The emission features in the [C I] $^3P_2 - ^3P_1$ spectra are excited molecular lines associated with the G34.3+0.15 hot cores. The PACS spectral maps near $158\mu\text{m}$ are presented in Fig B.12.

Appendix B.5: W49N

W49N is one of the most luminous star forming region in the Galaxy. It is located in the Perseus arm at a distance of 11.4 kpc (Gwinn et al. 1992). The line of sight crosses the Sagittarius arm twice along this very long line of sight.

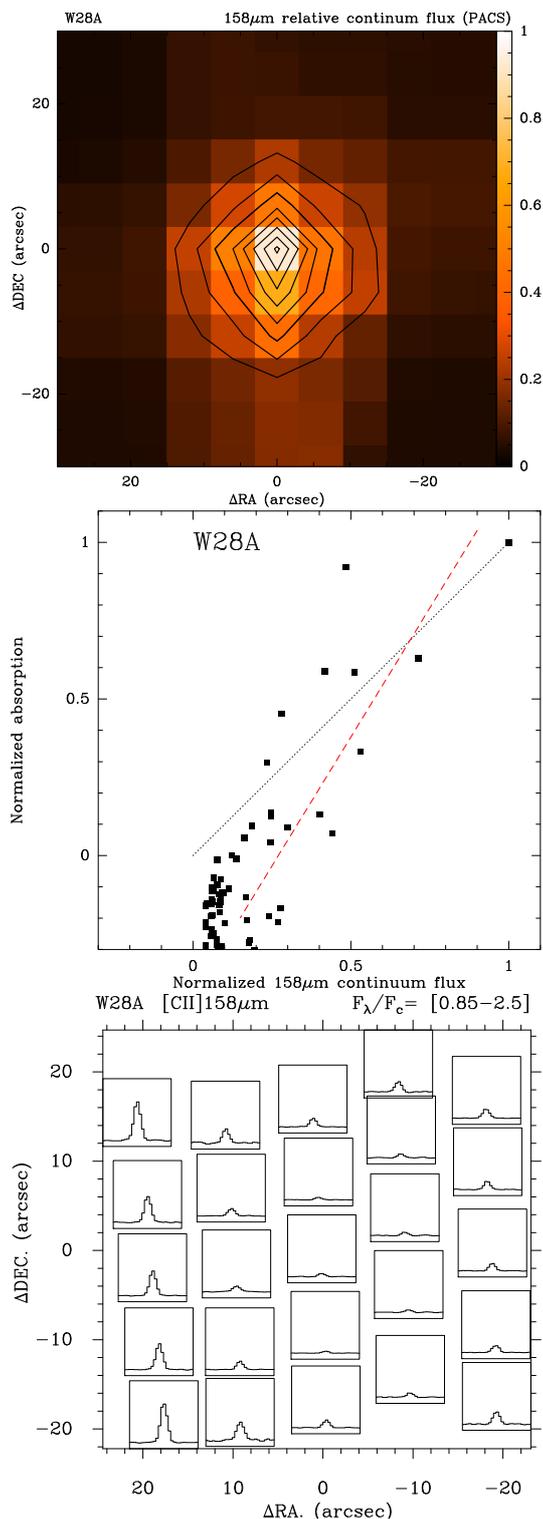


Fig. B.3. PACS data towards W28A at 158 μm . For all maps, the offsets are given relative to the central position listed in Table 1. Top: Continuum emission at 158 μm . Contour levels are 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Middle: Comparison of the integrated absorption measured in the HIFI map relative to the absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship. Bottom: Map of the line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 2.5. The data have not been corrected for the possible contamination in the OFF position.

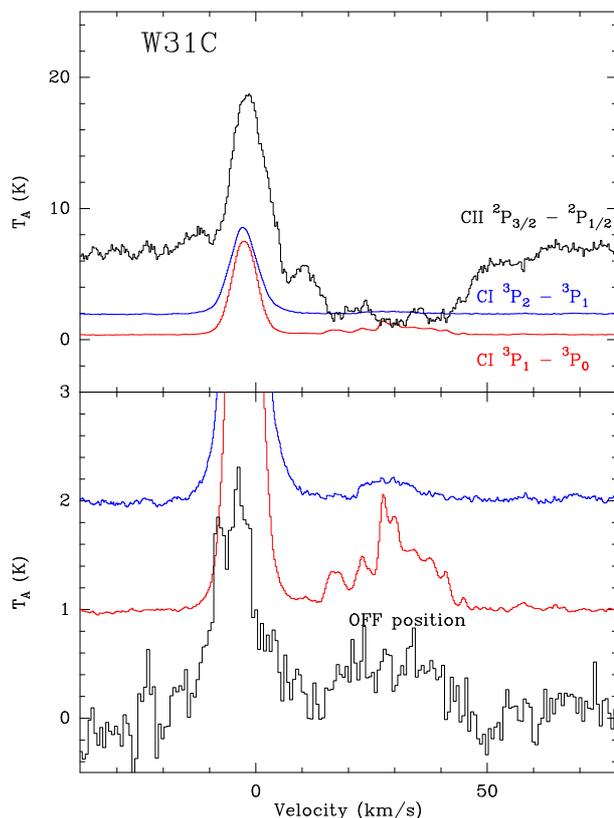


Fig. B.4. Top : Herschel/HIFI spectra towards W31C. The red line shows the $[\text{C}\text{I}]^3P_1 - ^3P_0$ line at 492 GHz, the blue line shows the $[\text{C}\text{I}]^3P_2 - ^3P_1$ line at 809 GHz, and the black line the $[\text{C}\text{II}]^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz. The horizontal axis is the LSR velocity in km s^{-1} and the vertical axis the antenna temperature in Kelvins. The continuum level for $[\text{C}\text{II}]$ corresponds to the SSB continuum levels. Bottom : zoom on the $[\text{C}\text{I}]$ lines (red and blue as above), and average $[\text{C}\text{II}]$ spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

Appendix B.6: W51

W51 is a very massive star formation complex in the Sagittarius arm at a distance of ~ 5.4 kpc (Sato et al. 2010). The line of sight crosses the local interstellar medium, and also probes relatively diffuse material in the foreground of W51 at velocities ~ 45 and ~ 65 km s^{-1} . The HIFI spectra are shown in Fig. B.13 and B.14. The $[\text{C}\text{II}]$ spectra show extended wings associated with the CO molecular outflow. The spectra are self reversed over a significant fraction of the map. Despite the high sensitivity reached in these data, we could not detect $[\text{C}\text{I}]$ either in absorption nor in emission associated with the prominent features at ~ 6 and ~ 20 km s^{-1} . The weak emission features in the $[\text{C}\text{I}]^3P_1 - ^3P_0$ and $^3P_2 - ^3P_1$ spectra are unrelated molecular lines associated with the W51 E1/E2 hot cores. The PACS spectral maps near 158 μm are presented in Figure B.15.

Appendix B.7: DR21(OH)

DR21(OH) is one of the most massive cores in the Cygnus molecular cloud complex. At a distance of ~ 1.5 kpc (Rygl et al. 2012), the line of sight mostly probes the material in the Cygnus region. While the material directly associated with DR21(OH) appears at slightly negative LSR velocities, the foreground gas is seen near 10 km s^{-1} and could be associated with W75N a few arc-minutes North East. The HIFI spectra are shown in Fig. B.16

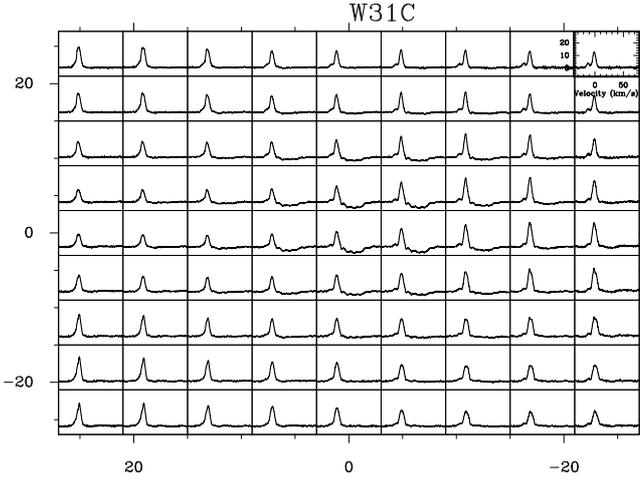


Fig. B.5. Montage of [C II] spectra towards W31C. A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -38 km s^{-1} to 78 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -7 K to 29 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

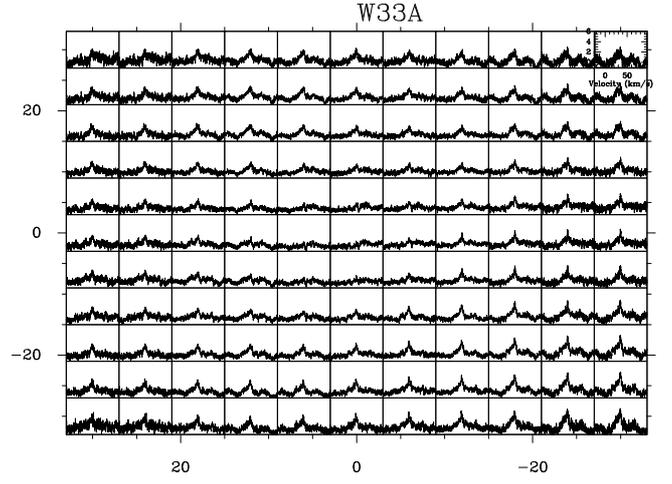


Fig. B.8. Montage of [C II] spectra towards W33A. A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -23 km s^{-1} to 95 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -1 to 6 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

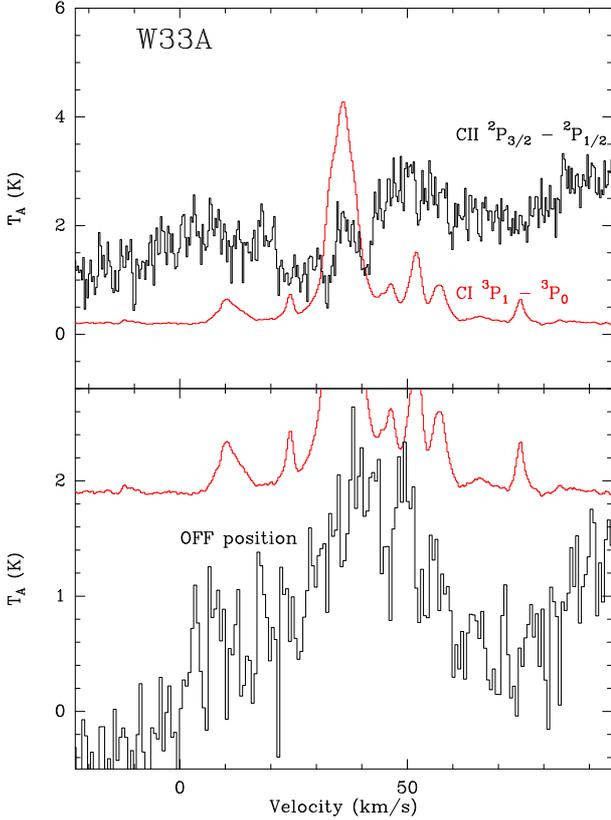


Fig. B.7. Top : Herschel/HIFI spectra towards W33A. The red line shows the $[\text{C I}]^3P_1 - ^3P_0$ line at 492 GHz and the black line the $[\text{C II}]^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz . The horizontal axis is the LSR velocity in km s^{-1} , and the vertical axis the antenna temperature in Kelvins. The continuum level for [C II] corresponds to the SSB continuum levels. Bottom : zoom on the [C I] lines (red, as above), and average [C II] spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

As DR21(OH) is a very weak continuum source at 1.4 GHz , the HI data also make use of the absorption towards the stronger continuum source DR 21, 3 arc-minutes South.

Appendix B.8: W3-IRS5

W3-IRS5 is a massive protostar in the W3 star forming complex at a distance of $\sim 2 \text{ kpc}$ (Xu et al. 2006). Located in the second Galactic quadrant, the line of sight towards W3-IRS5 is probing the outer Galaxy. The HIFI spectra are shown in Figs B.19 and B.20. They include the ground state transition of HF at 1.2 THz . A deep and narrow velocity component is detected at $\sim -20 \text{ km s}^{-1}$, and a broader one near 0 km s^{-1} . The H_2 column densities have been derived from the HF data using $[\text{HF}]/[\text{H}_2] = 1.3 \times 10^{-8}$ (Sonnentrucker et al. 2010).

The PACS spectral maps near $158 \mu\text{m}$ are presented in Fig B.21. The line of sight towards W3 has been studied in HI Normandeau (1999). Herschel-HIFI spectra exhibiting absorption from both the W3 complex and foreground material have been presented by Benz et al. (2010).

and B.17. The PACS spectral maps near $158 \mu\text{m}$ are presented in Figure B.18.

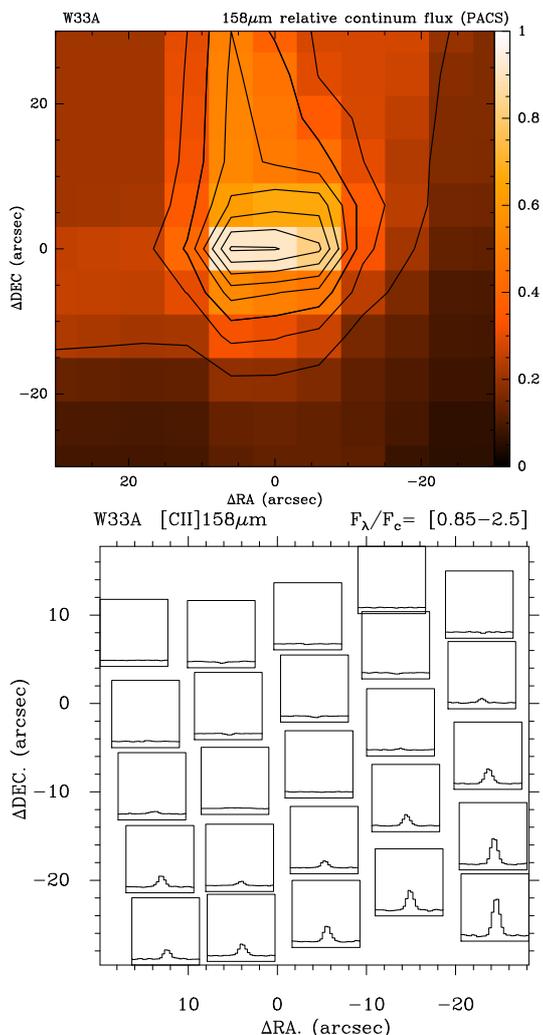


Fig. B.9. PACS data towards W33A. The offsets are given relative to the central position listed in Table 1. Top: Continuum emission at $158\mu\text{m}$. Contour levels are 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Bottom: Map of the line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 2.5. Note the very low contrast in the central spaxels. The data have not been corrected for the contamination in the OFF beam.

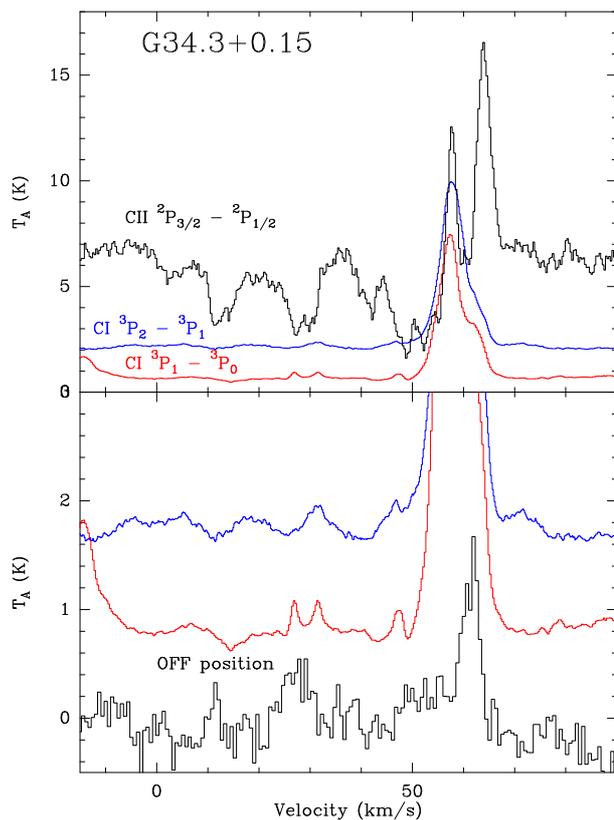


Fig. B.10. Top : Herschel/HIFI spectra towards G34.3+0.15. The red line shows the $[\text{C I}]^3P_1 - ^3P_0$ line at 492 GHz, the blue line shows the $[\text{C I}]^3P_2 - ^3P_1$ line at 809 GHz and the black line the $[\text{C II}]^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz. The horizontal axis is the LSR velocity in km s^{-1} and the vertical axis is the antenna temperature in Kelvins. The continuum level for $[\text{C II}]$ corresponds to the SSB continuum level. Bottom : zoom on the $[\text{C I}]$ lines (red, blue as above), and average $[\text{C II}]$ spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

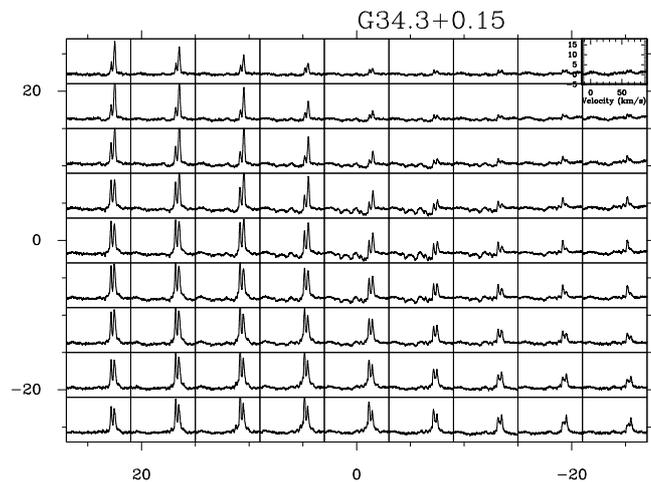


Fig. B.11. Montage of $[\text{C II}]$ spectra towards G34.3+0.15. A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -15 km s^{-1} to 90 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -5 K to 18 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

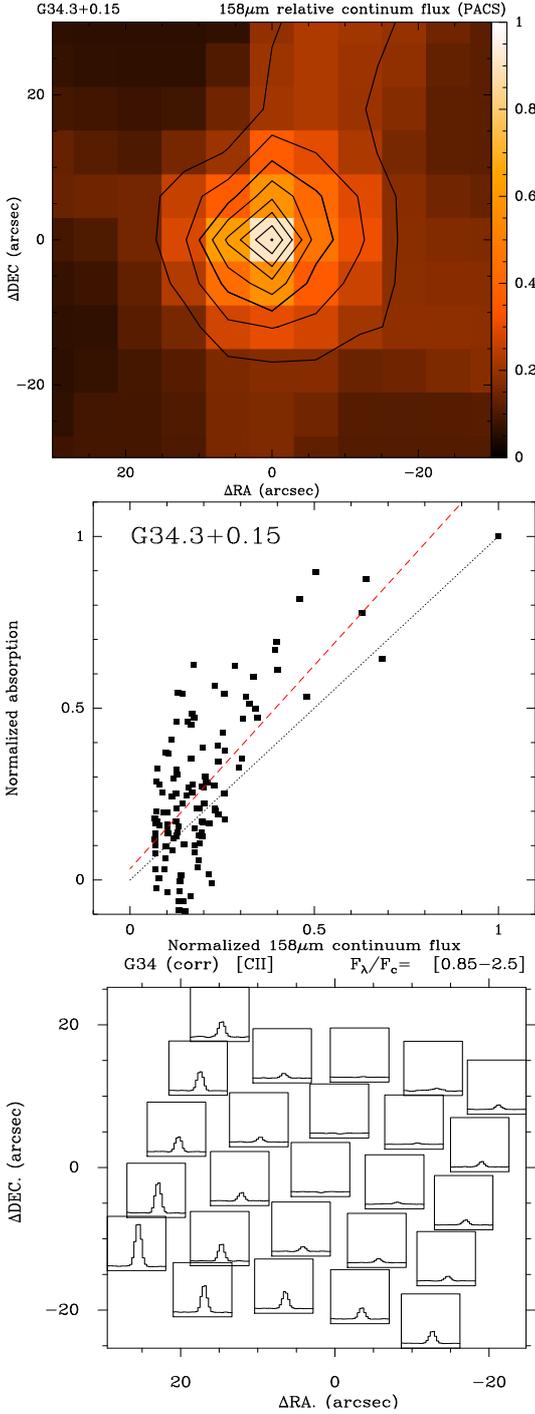


Fig. B.12. PACS data towards G34.3+0.15. For all maps, the offsets are given relative to the central position listed in Tab 1. Top: Continuum emission at $158\mu\text{m}$. Contour levels are at 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Middle: Comparison of the integrated absorption measured in the HIFI map relative to the absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship. Bottom: [C II] Map of the line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 2.5. The data have not been corrected for possible contamination in the OFF position.

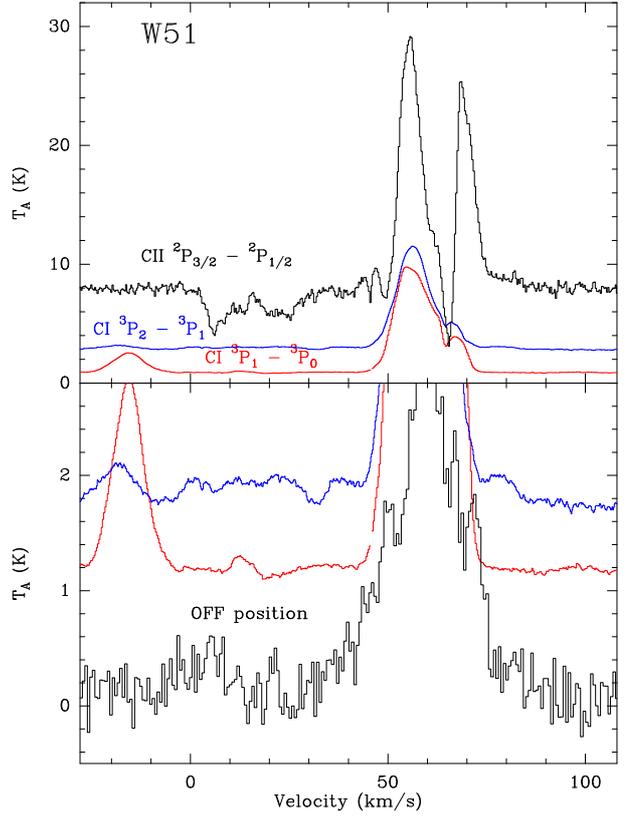


Fig. B.13. Top : Herschel/HIFI spectra towards W51. The red line shows the [C I] $^3P_1 - ^3P_0$ line at 492 GHz, the blue line shows the [C I] $^3P_2 - ^3P_1$ line at 809 GHz and the black line the [C II] $^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz. The horizontal axis is the LSR velocity in km s^{-1} and the vertical axis the antenna temperature in Kelvins. The continuum level for [C II] corresponds to the SSB continuum level. Bottom : zoom on the [C I] lines (red, blue as above), and average [C II] spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

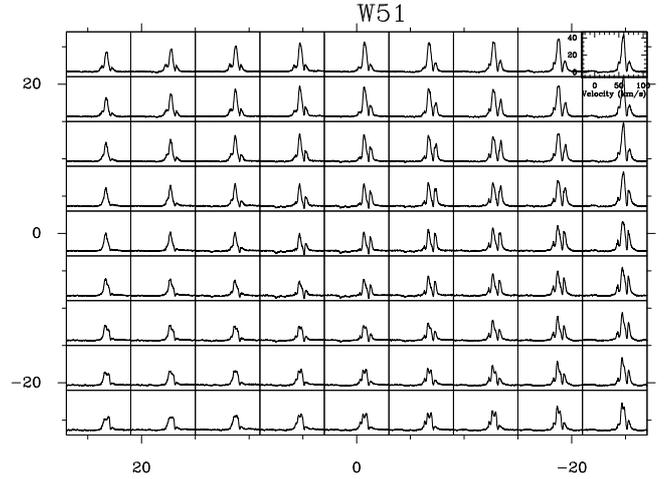


Fig. B.14. Montage of [C II] spectra towards W51. A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -28 km s^{-1} to 108 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -6 K to 47 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

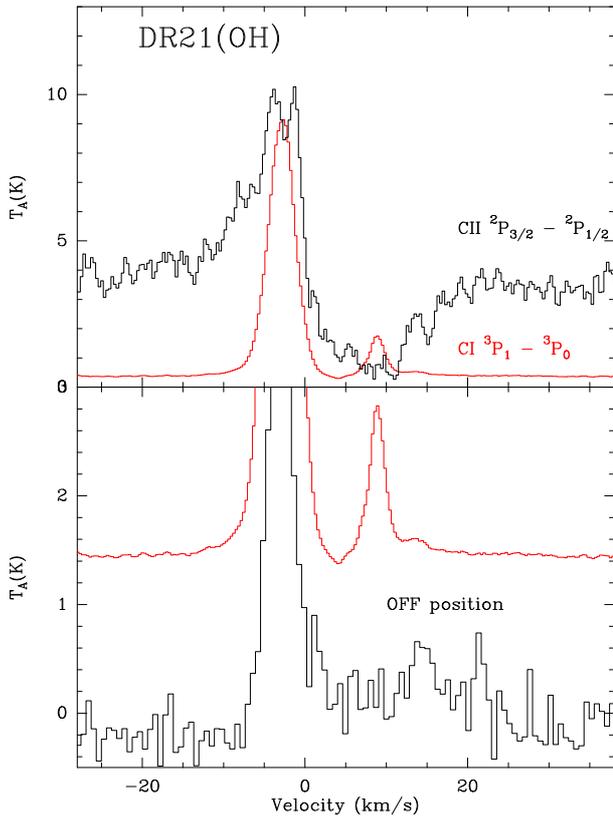


Fig. B.16. Top : Herschel/HIFI spectra towards DR21(OH). The red line shows the $[C\text{ I}]^3P_1 - ^3P_0$ line at 492 GHz, the blue line shows the $[C\text{ I}]^3P_2 - ^3P_1$ line at 809 GHz and the black line the $[C\text{ II}]^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz. The horizontal axis is the LSR velocity in km s^{-1} and the vertical axis the antenna temperature in Kelvins. The continuum level for $[C\text{ II}]$ corresponds to the SSB continuum level. Bottom : zoom on the $[C\text{ I}]$ lines (red, blue as above), and average $[C\text{ II}]$ spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.



Fig. B.17. Montage of $[C\text{ II}]$ spectra towards DR21(OH). A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -28 km s^{-1} to 38 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -4 K to 14 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

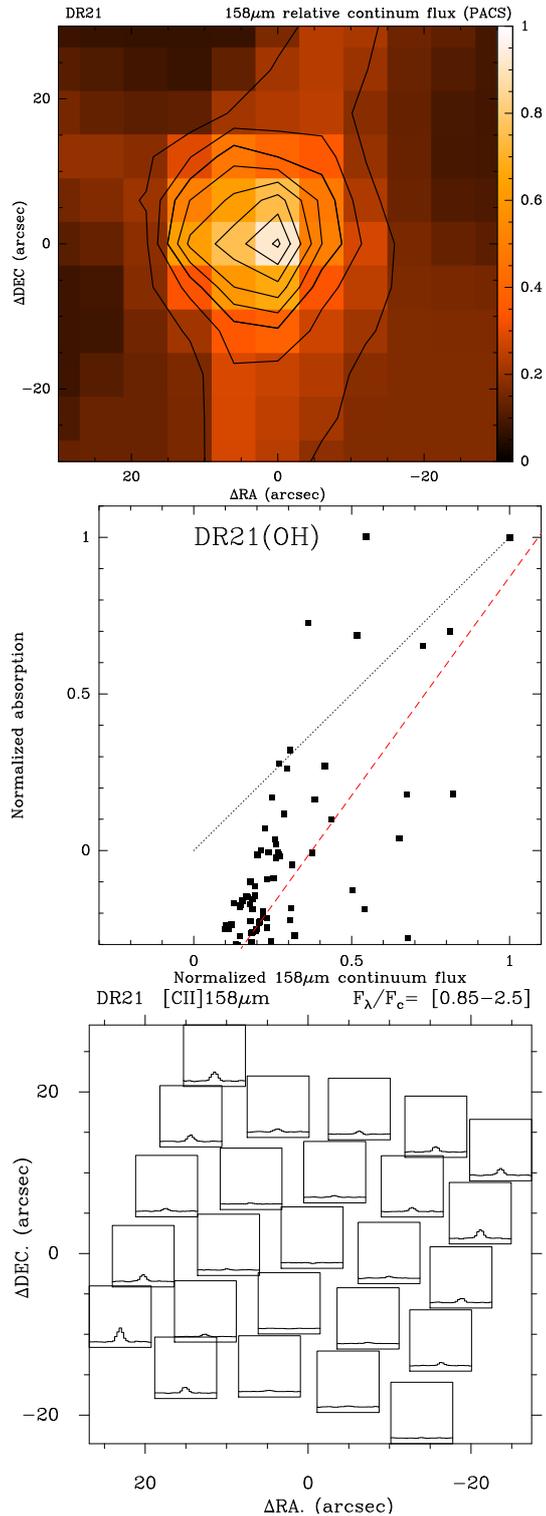


Fig. B.18. PACS data towards DR21(OH). For all maps, the offsets are given relative to the central position listed in Table 1. Top: Continuum emission at $158\mu\text{m}$. Contour levels are 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Middle: Comparison of the integrated absorption measured in the HIFI map relative to the absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship. Bottom: Map of the line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 2.5. The data have not been corrected for contamination in the OFF position.

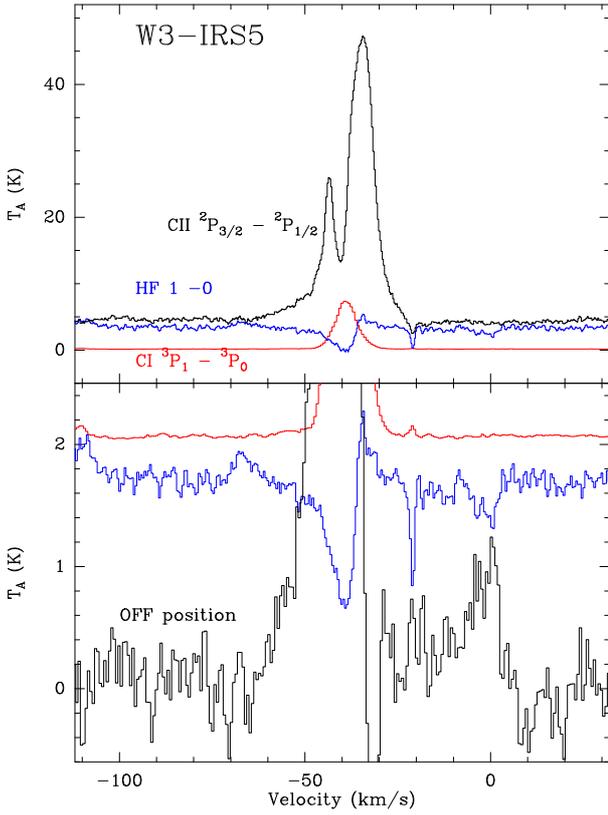


Fig. B.19. Top : Herschel/HIFI spectra towards W3-IRS5. The black line shows the [C II] $^2P_{3/2} - ^2P_{1/2}$ line at 1.9 THz, the red line shows the [C I] $^3P_1 - ^3P_0$ line at 492 GHz, and the blue line the ground state transition of HF at 1.2 THz. The horizontal axis is the LSR velocity in km s^{-1} and the vertical axis the antenna temperature in Kelvins. The continuum level for [C II] corresponds to the SSB continuum level. Bottom : zoom on the [C I] and HF lines (red and blue as above), and average [C II] spectrum of the OFF positions (black). The continuum levels have been shifted for clarity in the bottom panel.

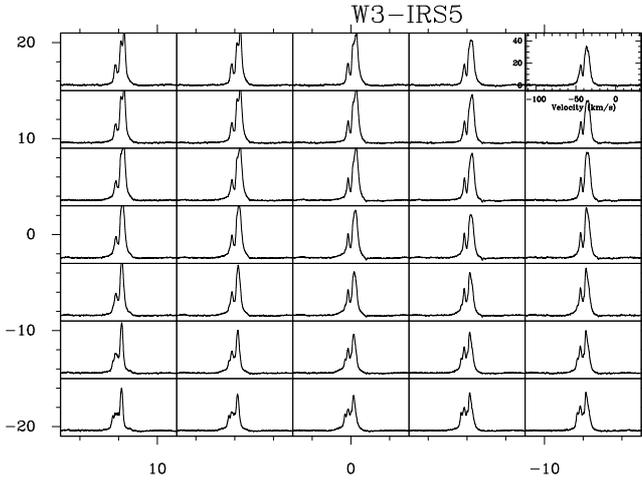


Fig. B.20. Montage of [C II] spectra towards W3-IRS5. A baseline has been subtracted from all spectra. The horizontal axis is the LSR velocity in km s^{-1} , which runs from -112 km s^{-1} to 32 km s^{-1} , and the vertical axis is the antenna temperature in Kelvins, which runs from -5 K to 47 K . The x-axis shows the right ascension offset in arc-sec and the y-axis the declination offset in arc-sec, relative to the source position given in Table 1.

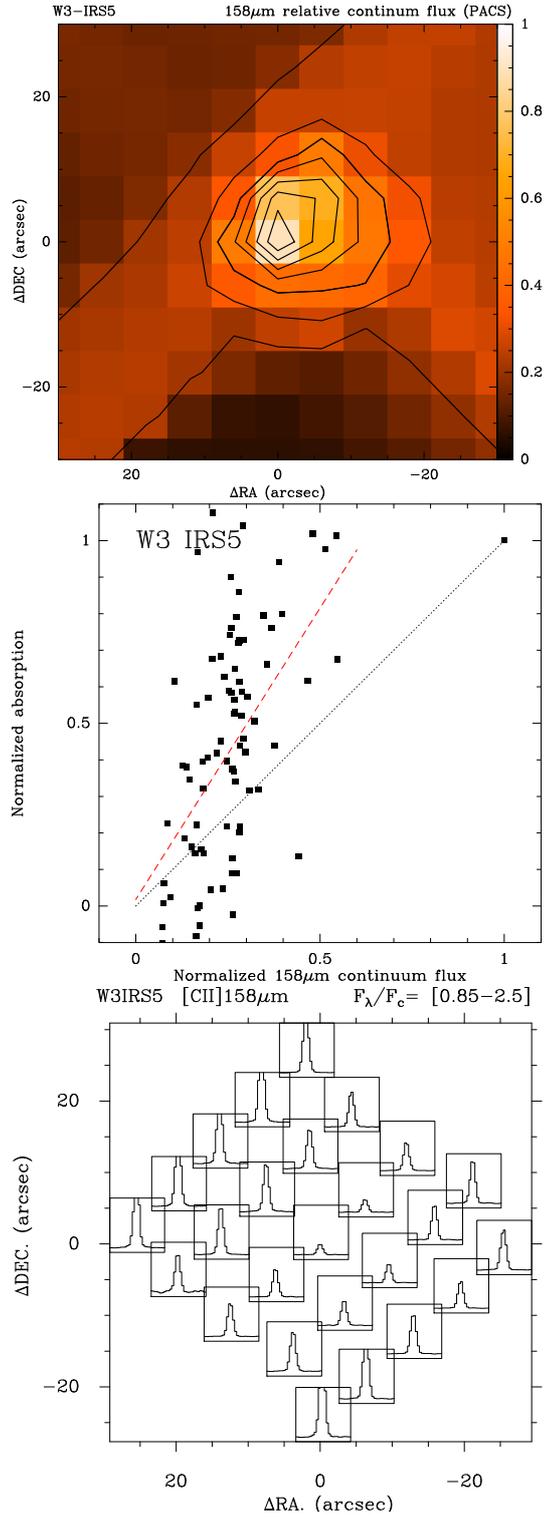


Fig. B.21. PACS data towards W3-IRS5. For all maps, the offsets are given relative to the central position listed in Table 1. Top: Continuum emission at $158 \mu\text{m}$. Contour levels are 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Middle: Comparison of the integrated absorption measured in the HIFI map relative to the absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship. Bottom: Map of the line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 2.5. The data have not been corrected for the contamination in the OFF position.

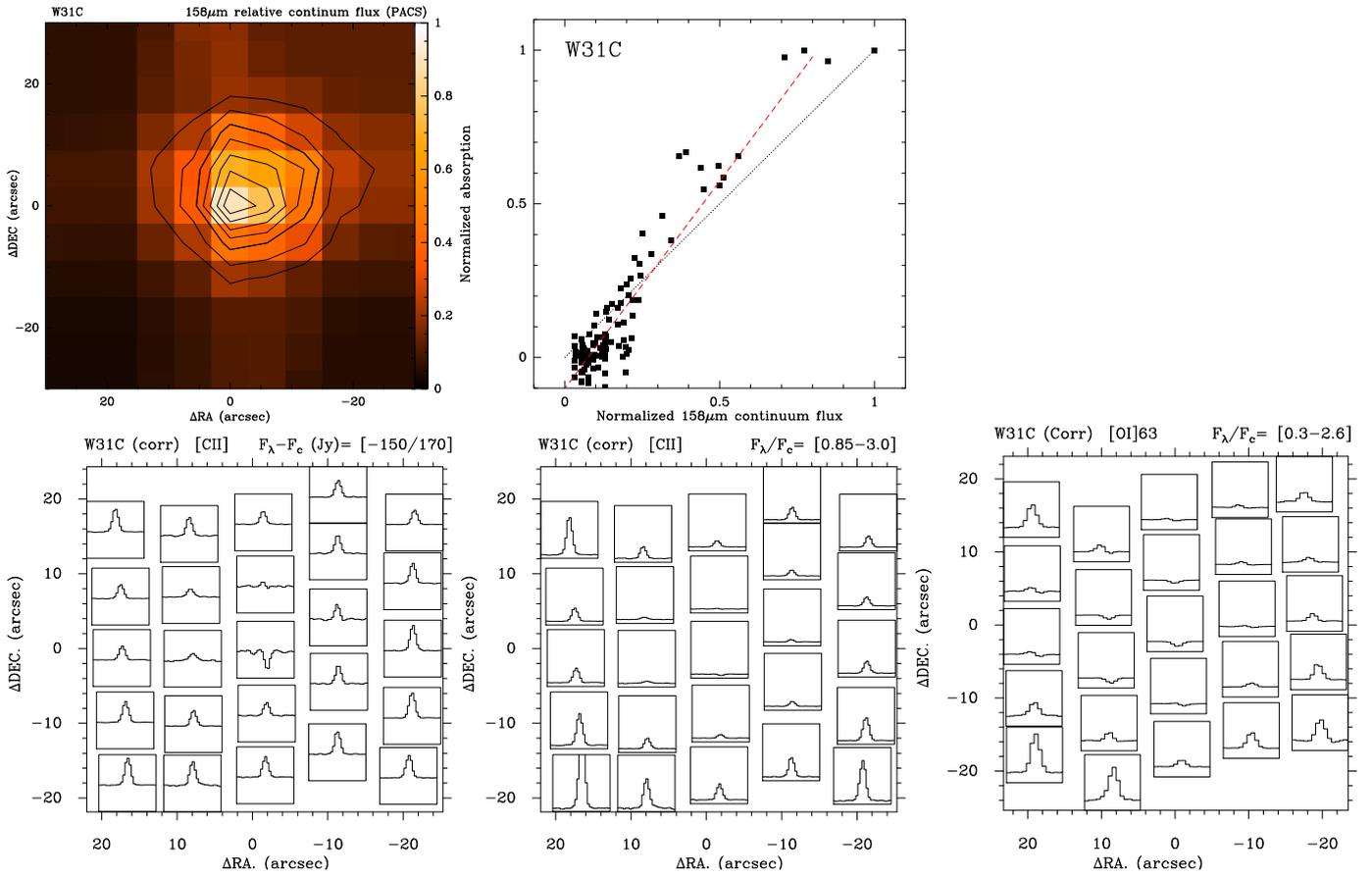


Fig. B.6. PACS data towards W31C. For all maps, the offsets are given relative to the central position listed in Table 1. Top left: Continuum emission at 158 μm . Contour levels are drawn at 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Top right: Comparison of the integrated absorption measured in the HIFI map relative to the absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship. Bottom left: [C II] emission in the 25 PACS spaxels. The continuum emission has been subtracted. The vertical scale runs from -150 to 170 Jy. Bottom middle: Map of the [C II] line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 3.0. Note the very low contrast in the central spaxels. Bottom right: Map of the [O I] 63 μm line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.3 to 2.6.

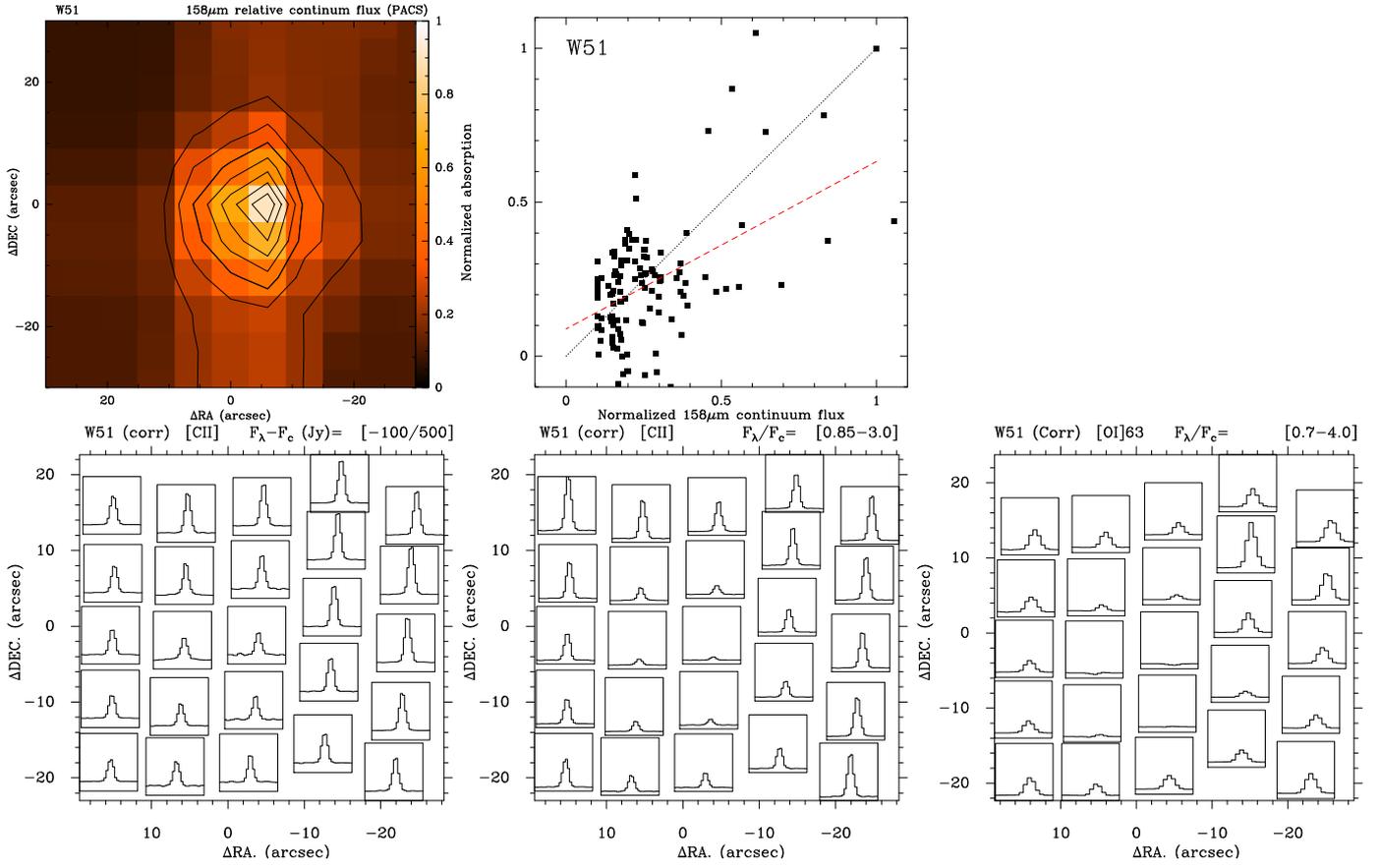


Fig. B.15. PACS data towards W51. For all maps, the offsets are given relative to the central position listed in Tab 1. Top left: Continuum emission at $158\mu\text{m}$. Contour levels are drawn at 0.2, 0.3, 0.4 ... 0.9 relative to the maximum. Top right: Comparison of the integrated absorption measured in the HIFI map relative to the absorption at the map center, with the continuum flux measured in the PACS map relative to the map center. The dashed red line shows the linear regression line and the dotted black line a 1:1 relationship. Bottom left: [C II] emission in the 25 PACS spaxels. The continuum emission has been subtracted. The vertical scale runs from -100 to 500 Jy. Bottom middle: Map of the [C II] line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.85 to 3. Bottom right: Map of the [O I] $63\mu\text{m}$ line to continuum emission/absorption in the 25 PACS spaxels. The vertical scale runs from 0.7 to 4.