

CONNECTING GRBs AND ULIRGs: A SENSITIVE, UNBIASED SURVEY FOR RADIO EMISSION FROM GAMMA-RAY BURST HOST GALAXIES AT $0 < z < 2.5$

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ABSTRACT

Luminous infrared galaxies and submillimeter galaxies contribute significantly to stellar mass assembly and provide an important test of the connection between the gamma-ray burst (GRB) rate and that of overall cosmic star formation. We present sensitive 3 GHz radio observations using the Karl G. Jansky Very Large Array of 32 uniformly selected GRB host galaxies spanning a redshift range from $0 < z < 2.5$, providing the first fully dust- and sample-unbiased measurement of the fraction of GRBs originating from the universe’s most bolometrically luminous galaxies. Four galaxies are detected, with inferred radio star formation rates (SFRs) ranging between 50 and $300 M_{\odot} \text{ yr}^{-1}$. Three of the four detections correspond to events consistent with being optically obscured “dark” bursts. Our overall detection fraction implies that between 9% and 23% of GRBs between $0.5 < z < 2.5$ occur in galaxies with $S_{3\text{GHz}} > 10 \mu\text{Jy}$, corresponding to $\text{SFR} > 50 M_{\odot} \text{ yr}^{-1}$ at $z \sim 1$ or $> 250 M_{\odot} \text{ yr}^{-1}$ at $z \sim 2$. Similar galaxies contribute approximately 10%–30% of all cosmic star formation, so our results are consistent with a GRB rate that is not strongly biased with respect to the total SFR of a galaxy. However, all four radio-detected hosts have stellar masses significantly lower than IR/submillimeter-selected field galaxies of similar luminosities. We suggest that the GRB rate may be suppressed in metal-rich environments but independently enhanced in intense starbursts, producing a strong efficiency dependence on mass but little net dependence on bulk galaxy SFR.

Key words: galaxies: starburst – gamma-ray burst: general – radio continuum: galaxies

1. INTRODUCTION

One of the primary appeals of the study of long-duration gamma-ray bursts (GRBs) lies in their potential to address broader questions of cosmology and galaxy evolution. As the explosions of massive stars at cosmological distances, GRBs are intimately connected with cosmic star formation, and the evolution of the cosmic GRB rate and the changing demographics of their host galaxies with time should reflect overall cosmological trends and inform our understanding of how galaxies and the stars within them form and evolve over the universe’s history (e.g., Hogg & Fruchter 1999; Blain & Natarajan 2000; Porciani & Madau 2001; Firmani et al. 2004; Natarajan et al. 2005; Kistler et al. 2008; Butler et al. 2010; Tanvir et al. 2012; Robertson & Ellis 2012; Salvaterra et al. 2013; Trenti et al. 2014).

Central to the utility of GRBs for this purpose is their extreme luminosity at all electromagnetic wavelengths, including wavelengths unaffected by absorption due to dust and gas. In particular, GRBs are first detected in hard X-rays and gamma-rays and are now routinely localized to $\sim 2''$ precision in soft X-rays using the *Swift* X-ray Telescope (XRT; Burrows et al. 2005), meaning that—with sufficient follow-up effort—their host-galaxy population and redshift distribution can be constructed independent

of the effects of obscuration (e.g., Hjorth et al. 2012; Jakobsson et al. 2012; Perley et al. 2013).

For this reason, one particular question in galaxy evolution that GRBs showed significant promise to help resolve is the relative importance of extremely luminous, dusty galaxies to cosmic star formation (e.g., Djorgovski et al. 2001; Ramirez-Ruiz et al. 2002; Berger et al. 2003). The UV/optical tracers by which galaxy and cosmic star formation rates (SFRs) are normally estimated are significantly impacted by interstellar dust. Most galaxies are predominately optically thin and the effects of dust can be corrected for via measurements of the UV spectral slope or Balmer decrement (Meurer et al. 1999; Smolčić et al. 2009). However, the most bolometrically luminous galaxies (such as submillimeter galaxies [SMGs] and local ultraluminous infrared galaxies [ULIRGs]) contain sufficient dust so as to be optically thick at UV and optical wavelengths, causing optical/UV-based tracers to inevitably underpredict the true SFRs of these galaxies even after dust correction (Blain et al. 2002; Goldader et al. 2002; Chapman et al. 2005). Instead, the SFRs of these objects (and therefore their contribution to the cosmic SFR) must be measured using long-wavelength tracers such as far-IR or radio continuum, but these methods have historically suffered from limited sensitivity, uncertain calibration, uncertain dust temperatures, and active galactic nucleus (AGN) contamination. Estimates of the importance of the most luminous and dusty galaxies to

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star formation have varied widely between different studies—in part due simply to varying definitions of what distinguishes a “luminous” and dusty galaxy from an “ordinary” one—but have ranged from estimates of 50% or more (Pérez-González et al. 2005; Michałowski et al. 2010; Magnelli et al. 2013) down to only about 10% (Rodighiero et al. 2011); see Casey et al. (2014) for a review.

In principle, the contribution of dusty star-forming galaxies (or any other galaxy population) to overall cosmic star formation could be determined simply by measuring the fraction of GRBs hosted within such galaxies: for example, if ULIRGs (galaxies with $L_{\text{IR}} > 10^{12} L_{\odot}$) represent 40% of all star formation at $1 < z < 2$, they should also produce 40% of all GRBs over the same redshift interval. However, this would strictly apply only if GRBs represent an unbiased tracer of overall star formation. A large volume of evidence now suggests they do not: GRB host galaxies at $z \lesssim 1.5$ have lower masses, lower metallicities, bluer colors, and more irregular morphologies compared to what would be expected from an unbiased SFR-tracing population or when compared to the core-collapse supernova host population (Fruchter et al. 2006; Modjaz et al. 2008; Levesque et al. 2010; Graham & Fruchter 2013; Kelly et al. 2014; Vergani et al. 2014), even when including dust-obscured “dark” bursts (Perley et al. 2013, although cf. Hunt et al. 2014).

Nevertheless, measuring the contribution of very luminous, dusty galaxies to the GRB rate remains important for understanding the overall link between GRBs and cosmic star formation, and the influence of dust on our view of each. What causes GRBs to favor certain environments over others is not well understood. A strong preference for low metallicity has been a long-favored explanation given the connection between line-driven winds and angular momentum loss (MacFadyen & Woosley 1999; Woosley & MacFadyen 1999; Woosley & Heger 2006), but a variable initial mass function (IMF) or other environmental factors (such as dynamical interaction of stars in dense stellar clusters that preferentially form in the most intense starbursting galaxies; van den Heuvel & Portegies Zwart 2013) may also help explain the observed characteristics of the host-galaxy population. As extrema of the star-forming galaxy population, dusty and luminous starbursts form an excellent test-bed for distinguishing different hypotheses: in particular, they are probably very metal-rich in their interiors (Nagao et al. 2012) but also very dense and intensely star-forming (Daddi et al. 2007).

A number of previous efforts to observe GRB hosts at mid-IR, submillimeter, and/or radio wavelengths have been performed over the past decade (Berger et al. 2003; Barnard et al. 2003; Tanvir et al. 2004; Le Floch et al. 2005; Stanway et al. 2010; Hatsukade et al. 2012; Michałowski et al. 2012b; Wang et al. 2012; Perley & Perley 2013; Hunt et al. 2014; Schady et al. 2014). However, all of these efforts have been either limited in scope (observing only small samples), limited in sensitivity (insensitive to even very luminous galaxies beyond $z > 1$), or subject to uncertain selection biases (most commonly a bias in favor of GRBs with afterglow-based redshift determinations that will systematically miss dust-obscured events, although other efforts have specifically targeted only the *most* heavily dust-obscured GRBs).

In this paper, we present results of the first GRB host-galaxy survey that is simultaneously unbiased in regard to target selection, sensitive enough to detect the long-wavelength emission from luminous star-forming galaxies even out to $z \sim 2.5$, and large enough to usefully statistically constrain the

fraction of GRBs that originate in such systems. Specifically, we survey a sample of 32 uniformly selected GRB host galaxies (a factor of ~ 2 larger than any previous long-wavelength survey) to a limiting radio flux density of approximately $10 \mu\text{Jy}$ at 3 GHz (a factor of 2–3 deeper than any other radio host survey)—detecting four in total and providing the first definitive measurement of the fraction of GRBs produced by the universe’s most luminous galaxies. Our sample selection, observations, and analysis are discussed in Section 2. We present our results and interpret our four detections as star formation-associated emission from the GRB host galaxies in Section 3. Further discussion of individual systems, including a detailed discussion of all detections and a few notable nondetections, is presented in Section 4 and Section 5. The overall statistical properties of our sample and its implications are discussed in Section 6, and conclusions are summarized in Section 7.

2. SAMPLE AND OBSERVATIONS

2.1. Sample Selection

Given the heterogeneous nature of previous long-wavelength efforts and the need to determine the fraction of luminous hosts even among both obscured and unobscured GRBs (a few dusty and luminous galaxies have been reported hosting even the latter), the primary consideration guiding the choice of targets was the need for a uniformly constructed sample. The production of unbiased subsets out of what is now a very large (but, often, poorly and non-uniformly observed) overall GRB catalog using simple observability cuts has become increasingly widespread in recent years, starting with the efforts of Jakobsson et al. (2006) and Fynbo et al. (2009) to select subsets of GRB afterglows for population analysis and, more recently, the host-galaxy focused The Optically Unbiased GRB Host (TOUGH; Hjorth et al. 2012) and BAT6 (Salvaterra et al. 2012) samples.

For this effort we chose a subset of GRBs from TOUGH, which was well suited to our goals for several reasons. First, the TOUGH sample size is large enough to be statistically informative but small enough to be feasible in a one-year Very Large Array (VLA) campaign. Second, deep optical host-galaxy imaging was available for all targets, needed in order to calculate accurate centroids and angular size constraints. Third, the survey boasts relatively high-redshift completeness ($\sim 85\%$, provided from a combination of afterglow spectroscopy available for most events and an extensive host-galaxy spectroscopic campaign; Jakobsson et al. 2012 and Krühler et al. 2012). Finally, all GRBs in the sample date to 2007 or earlier, ensuring a long temporal baseline between explosion and the present time such that any contribution from a radio afterglow is minimal.

The original selection criteria by which the TOUGH sample was chosen out of the broader catalog of *Swift* GRBs are described in detail in Hjorth et al. (2012), which together establish a list of 69 GRBs. We added two additional criteria for our VLA observations. First, to ensure observability from the VLA we required a declination of $\delta > -30^\circ$ (since the original TOUGH criteria establish a declination maximum of $\delta < +27^\circ$, this effectively limits the sample to $-30^\circ < \delta < 27^\circ$). In addition, we excluded any targets known to be at $z > 2.5$, since detecting even very luminous galaxies at higher redshift is extremely difficult at radio wavelengths. We did, however, include TOUGH GRBs in our declination range whose redshift is not yet known, to ensure that imposing a redshift cut did not produce any biases (by, for example, inadvertently excluding

Table 1
VLA Observations

| GRB | R.A. ^a | Decl. ^a | z^b | Config. ^c | Observation Date (UT) | t_{int}^d (min) | Beam Size ^e (") | rms Noise ^f ($\mu\text{Jy beam}^{-1}$) |
|----------------------|-------------------|--------------------|-----------|----------------------|--------------------------|-----------------------------|-------------------------------|--|
| 050416A | 12:33:54.64 | +21:03:26.8 | 0.654 | A | 2014 Jan 2 | 119.7 | 2.3×2.1 | 5.6 |
| 050525A | 18:32:32.67 | +26:20:21.6 | 0.606 | A, B | 2012 Dec 2, 2014 Jan 10 | 138.7 | 1.9×1.6 | 3.0 |
| 050714B | 11:18:47.71 | -15:32:49.0 | 2.438 | BnA | 2014 Jan 27 | 122.9 | 2.4×1.7 | 3.3 |
| 050801 | 13:36:35.32 | -21:55:42.7 | 1.560 | BnA | 2014 Jan 30 | 118.1 | 1.9×1.6 | 3.3 |
| 050803 | 23:22:37.85 | +05:47:08.5 | ? | A | 2012 Dec 24, 2013 Jan 5 | 141.5 | 0.7×0.6 | 3.5 |
| 050819 | 23:55:01.62 | +24:51:39.0 | 2.500 | B | 2013 Nov 1 | 121.3 | 2.5×2.1 | 3.6 |
| 050824 | 00:48:56.21 | +22:36:33.1 | 0.830 | B | 2013 Nov 23 | 119.7 | 2.2×2.0 | 3.5 |
| 050922B | 00:23:13.37 | -05:36:16.7 | ? | A, B ^g | 2012 Dec 3, 2013 Nov 22 | 161.1 | 0.8×0.6 | 7.8 |
| 050922C | 21:09:33.08 | -08:45:30.2 | 2.198 | A, B | 2012 Oct 8, 2013 Nov 17 | 145.5 | 2.3×1.6 | 3.4 |
| 051006 | 07:23:14.14 | +09:30:20.0 | 1.059 | A, B | 2013 Jan 5, 2013 Dec 7 | 134.0 | 1.7×1.4 | 3.2 |
| 051016B | 08:48:27.85 | +13:39:20.4 | 0.936 | B | 2013 Dec 6 | 118.1 | 2.6×2.0 | 3.9 |
| 051117B | 05:40:43.38 | -19:16:27.2 | 0.481 | A, BnA | 2012 Nov 1, 2014 Jan 27 | 140.6 | 1.5×1.2 | 5.4 |
| 060218 | 03:21:39.69 | +16:52:01.6 | 0.033 | A | 2012 Nov 30 | 74.5 | 0.6×0.6 | 3.9 |
| 060306 | 02:44:22.88 | -02:08:54.7 | 1.55 | A | 2012 Dec 15 | 70.7 | 0.7×0.5 | 6.2 |
| 060604 | 22:28:55.04 | -10:54:56.1 | 2.140 | A, B | 2012 Dec 30, 2013 Nov 23 | 137.8 | 2.3×1.5 | 3.2 |
| 060805A | 14:43:43.47 | +12:35:11.2 | 0.60/2.44 | A | 2012 Nov 11 | 109.8 | 0.7×0.6 | 3.3 |
| 060814 | 14:45:21.36 | +20:35:09.2 | 1.920 | B | 2013 Nov 17 | 119.7 | 2.2×2.0 | 3.2 |
| 060908 | 02:07:18.42 | +00:20:32.2 | 1.884 | B | 2013 Dec 10 | 121.3 | 2.2×1.9 | 9.5 |
| 060912A | 00:21:08.13 | +20:58:18.5 | 0.937 | A, B | 2013 Jan 1, 2013 Nov 3 | 134.0 | 1.8×1.6 | 3.1 |
| 060923A | 16:58:28.14 | +12:21:37.9 | 2.5 | B | 2013 Nov 23 | 118.1 | 2.3×2.1 | 6.1 |
| 060923C | 23:04:28.36 | +03:55:28.4 | ? | A, B | 2012 Oct 25, 2013 Nov 22 | 160.5 | 1.9×1.5 | 3.2 |
| 061021 | 09:40:36.12 | -21:57:05.2 | 0.346 | A, BnA | 2012 Nov 15, 2014 Jan 29 | 140.6 | 1.5×1.1 | 3.0 |
| 061110A | 22:25:09.89 | -02:15:30.4 | 0.758 | A, B | 2012 Dec 24, 2013 Oct 26 | 139.6 | 3.1×1.5 | 6.1 |
| 061121 | 09:48:54.59 | -13:11:42.1 | 1.314 | B | 2014 Jan 2 | 118.1 | 3.0×1.9 | 5.8 |
| 070129 | 02:28:00.98 | +11:41:03.4 | 2.340 | B | 2013 Dec 3 | 119.7 | 2.2×1.9 | 5.1 |
| 070224 | 11:56:06.57 | -13:19:48.8 | 1.992 | B | 2013 Nov 23 | 118.1 | 3.0×1.9 | 5.7 |
| 070306 | 09:52:23.29 | +10:28:55.5 | 1.496 | A | 2012 Nov 11 | 189.5 | 0.7×0.6 | 2.9 |
| 070506 | 23:08:52.31 | +10:43:20.8 | 2.310 | A, B | 2013 Jan 1, 2013 Nov 3 | 132.2 | 1.9×1.6 | 5.4 |
| 070611 | 00:07:58.12 | -29:45:20.4 | 2.040 | A, BnA | 2012 Nov 8, 2014 Feb 4 | 140.6 | 1.7×1.5 | 3.1 |
| 070621 | 21:35:10.08 | -24:49:03.1 | ? | A | 2012 Oct 31 | 127.9 | 1.0×0.6 | 6.2 |
| 070808 | 00:27:03.36 | +01:10:34.4 | ? | B | 2013 Nov 23 | 122.9 | 2.6×2.0 | 3.4 |
| 070810A ^h | 12:39:51.24 | +10:45:03.2 | 2.170 | A | 2012 Nov 11 | 69.8 | 0.7×0.6 | 4.6 |

Notes.^a Observation pointing center (J2000).^b Redshift of host or afterglow, generally from Hjorth et al. (2012) and sources quoted within or from Krühler et al. (2015). GRB 060805A has two possible host-galaxy candidates; we will generally assume $z = 2.44$ in our plots since it is closer to the *Swift* median redshift. The redshifts of GRB 060306 and GRB 060923A are from Perley et al. (2013); the redshift of GRB 060923A is photometric.^c VLA array configuration.^d Total time on-source, excluding overheads.^e Major and minor axis FWHM of the synthesized beam.^f Noise (1σ) estimated from the standard deviation of 1000 randomly chosen points in the final map.^g Only the A-configuration observations were usable due to severe RFI affecting the B-configuration data set.^h GRB 070810A was observed as part of the TOUGH campaign although it technically did not satisfy the sun-angle constraint of the survey. We include it as part of our survey here.

dust-obscured events at $z < 2.5$ whose redshifts may have been missed).

These additional cuts on declination and redshift produced a final sample of 32 GRBs. Of these, 27 are known to be at $z < 2.5$ (although one of these does not have a unique redshift/host identification) and 5 are at unknown redshift.

2.2. VLA Observations

All of the 32 targets described above were observed with the fully upgraded Karl G. Jansky VLA, using the S-band receivers (central frequency 3 GHz) and 8-bit samplers with a bandwidth of 2048 MHz. Observations were conducted during the A-configuration cycle of fall 2012, the B-configuration cycle of fall 2013, and the BnA-configuration of winter 2014 (project codes 12B-305 and 13B-316). Integration times for each field were approximately two hours on-target. Observations employed

the WIDAR correlator (Perley et al. 2011) using a spectral resolution of 2 MHz and a sampling time of 4 s. Local phase and amplitude calibration was established via observations of a nearby source approximately every six minutes, and the flux-density scale was calibrated with a single observation of 3C 48, 3C 138, or 3C 286 using the reference scale of Perley & Butler (2013) at the beginning or end of each observation sequence. A summary of all observations is presented in Table 1.

Phase stability was very good, with typical variations of less than 10° over the 6 minute calibration cycle. More troublesome were the receiver amplitude gain variations induced by transmission from geostationary satellites. As seen from the VLA's latitude, these satellites are located in a belt near a declination of $-4^\circ.5$, and examination of our data shows that observations of sources located within about 10° of this declination (i.e., $+5^\circ.5$ to $-14^\circ.5$) show significant gain variations due to variable power

from these satellites entering the analog signal path through the antennas’ sidelobes. As the antennas track the source, the input satellite power varies as the sidelobes sweep past the satellites. These gain variations are tracked by the antennas’ onboard switched power calibration system, and application of this monitoring system reduced the gain uncertainty to less than 5% for all affected sources. For those sources within $\sim 2^\circ$ – 3° of declination $-4:5$ (050922B, 060306, 061110A) the satellite input power occasionally saturated the analog receiver system. The data from those times were deleted.

Data reduction was carried out using the Astronomical Image Processing System. Following gain and amplitude calibration, the data were examined for sporadic radio frequency interference (RFI). Visibilities with amplitude values on the (2 MHz, 4 s) resolution scale exceeding 5σ were removed. In addition, the data within the transmission bands for the XM/Sirius satellite digital radio systems (2320–2350 MHz) were removed for all observations. Imaging/deconvolution was done using faceting over the primary beam to reduce non-coplanar distortions. For most targets, no attempt was made to account for the spectral index gradients of the background sources due to the wavelength-dependent primary beam size, as these background sources were generally far enough from the target position that the imaging errors were reduced to below the noise level at the position of the target. One exception is the field of GRB 061110A, which contains several strong sources; for this target we imaged each frequency window individually and averaged together the resulting maps. The other exception was for the field of GRB 070808, in which the target position was located along an image artifact associated with a single moderately strong quasar. Effective rms sensitivities of the final maps are between 3 and $4 \mu\text{Jy beam}^{-1}$ for most targets. A few fields are slightly noisier (4–9 μJy), in most cases due to the influence of geostationary-belt satellites. A mosaic of all fields is shown in Figure 1.

2.3. VLA Flux-density Measurement

We evaluate the significance of host detections/limits, and measure the host-galaxy fluxes, in two different ways. First, we take the position of the host centroid in the optical images and simply measure the flux-density value in the $0''.18$ cell nearest to this position in the reduced, deconvolved VLA map. To measure the uncertainty, we take the rms of the same final map, evaluated directly by a Monte Carlo sampling of the fluxes of 1000 randomly selected cells in the central region of the image. The positions and fluxes evaluated using this method are presented in Table 2.

There is a possibility that the radio host centroid may be offset from the optical one, and there are a few fields where the exact host position is not known because the host is not detected in the available TOUGH imaging. To accommodate the possibility of an offset host galaxy, we also employ a second method in which we take the maximum value of any cell within a radius given by either the measured size of the host galaxy in the Very Large Telescope (VLT) optical images or (if not detected) by the size of the positional uncertainty region of the afterglow in the *Swift* XRT image. Uncertainties are estimated by a similar 1000 Monte Carlo sampling as for the fixed-position. Fluxes calculated via this method were (for all secure detections) consistent within uncertainties with the fixed-position values, and we conservatively adopt the lower fluxes of the fixed-position method for our later analysis. However, we do use the maximum-flux method to check the confidence of our detections: specifically, we consider as a “detection”

Table 2
Host-galaxy 3 GHz Flux-density Measurements

| GRB | F_ν^a (μJy) | $F_\nu^{\text{max} b}$ (μJy) | Conf. ^c |
|----------------------|---------------------------------|--|--------------------|
| 050416A | -7.87 ± 4.18 | -0.64 | 0.150 |
| 050525A | -1.25 ± 2.99 | 0.57 | 0.221 |
| 050714B | 4.91 ± 3.17 | 6.48 | 0.799 |
| 050801 | -4.79 ± 3.64 | 1.79 | 0.103 |
| 050803 | 1.46 ± 3.42 | 8.78 | 0.756 |
| 050819 | -0.92 ± 3.22 | 1.72 | 0.329 |
| 050824 | -0.11 ± 3.51 | 0.70 | 0.147 |
| 050922B | -2.72 ± 8.01 | 22.70 | 0.965 |
| 050922C | 8.75 ± 3.54 | 9.37 | 0.973 |
| 051006 | 9.08 ± 3.17 | 9.74 | 0.984 |
| 051016B | 0.97 ± 4.03 | 8.25 | 0.874 |
| 051117B | -4.87 ± 4.49 | 5.15 | 0.221 |
| 060218 | 5.52 ± 3.88 | 10.38 | 0.303 |
| 060306 | 7.03 ± 6.40 | 8.56 | 0.079 |
| 060604 | -3.50 ± 3.35 | 1.17 | 0.266 |
| 060805A ^d | -4.52 ± 3.45 | 8.43 | 0.619 |
| 060805A ^e | 2.25 ± 3.45 | 5.11 | 0.426 |
| 060814 | 11.34 ± 3.12 | 15.32 | 0.998 |
| 060908 | 4.53 ± 5.95 | 8.04 | 0.692 |
| 060912A | 4.54 ± 3.37 | 6.87 | 0.850 |
| 060923A | 0.05 ± 5.00 | 5.45 | 0.685 |
| 060923C | 2.87 ± 3.89 | 5.95 | 0.736 |
| 061021 | 0.82 ± 2.99 | 4.41 | 0.563 |
| 061110A | 14.20 ± 6.08 | 18.64 | 0.978 |
| 061121 | 17.07 ± 5.47 | 17.80 | 0.995 |
| 070129 | 4.92 ± 5.23 | 12.22 | 0.918 |
| 070224 | -1.35 ± 6.75 | 1.53 | 0.335 |
| 070306 | 11.31 ± 2.84 | 13.31 | 0.998 |
| 070506 | 3.69 ± 4.62 | 6.74 | 0.750 |
| 070611 | -9.62 ± 4.53 | 3.74 | 0.086 |
| 070621 | 11.35 ± 6.45 | 13.19 | 0.696 |
| 070808 | 9.78 ± 3.47 | 10.63 | 0.971 |
| 070810A | -1.45 ± 4.76 | 8.33 | 0.480 |

Notes. Boldface indicates the putative detections discussed in Sections 4.1–4.4.

^a Measured flux density (at 3 GHz) at the position of the optical/NIR host-galaxy centroid (or best-position afterglow centroid, if the host is not detected in optical/NIR observations).

^b Maximum flux density (at 3 GHz) in any $0''.18$ synthesized cell consistent with the position of the optical/NIR host-galaxy disk (or afterglow uncertainty region, if the host is undetected).

^c Significance of the detection, based on placing a large number of search apertures of identical size to the aperture used to calculate F_ν^{max} randomly across the image and calculating the maximum-flux density in each one.

^d Southern host candidate ($z = 0.60$).

^e Northern host candidate ($z = 2.44$).

any source with a measured fixed-position flux at least 2.5σ times the uncertainty and for which the flux measured from the maximum-flux method exceeds the flux measured in 98% of the maximum-flux Monte Carlo trials (rightmost column in Table 2). We also verify that any detections are pointlike (see Section 3).

Our flux measurement procedure assumes that the host galaxy is not extended at the resolution of our available imaging. This is generally a good assumption (most hosts are only marginally resolved even at the resolution of the VLT images, and the largest galaxies are only $\sim 2''$ in diameter, comparable to our beam size in most cases), although for a few fields with more extended host galaxies and/or compact beams (e.g., 060814, 060218, 051117B) our flux measurement/limit may be a modest underestimate of the true integrated value.

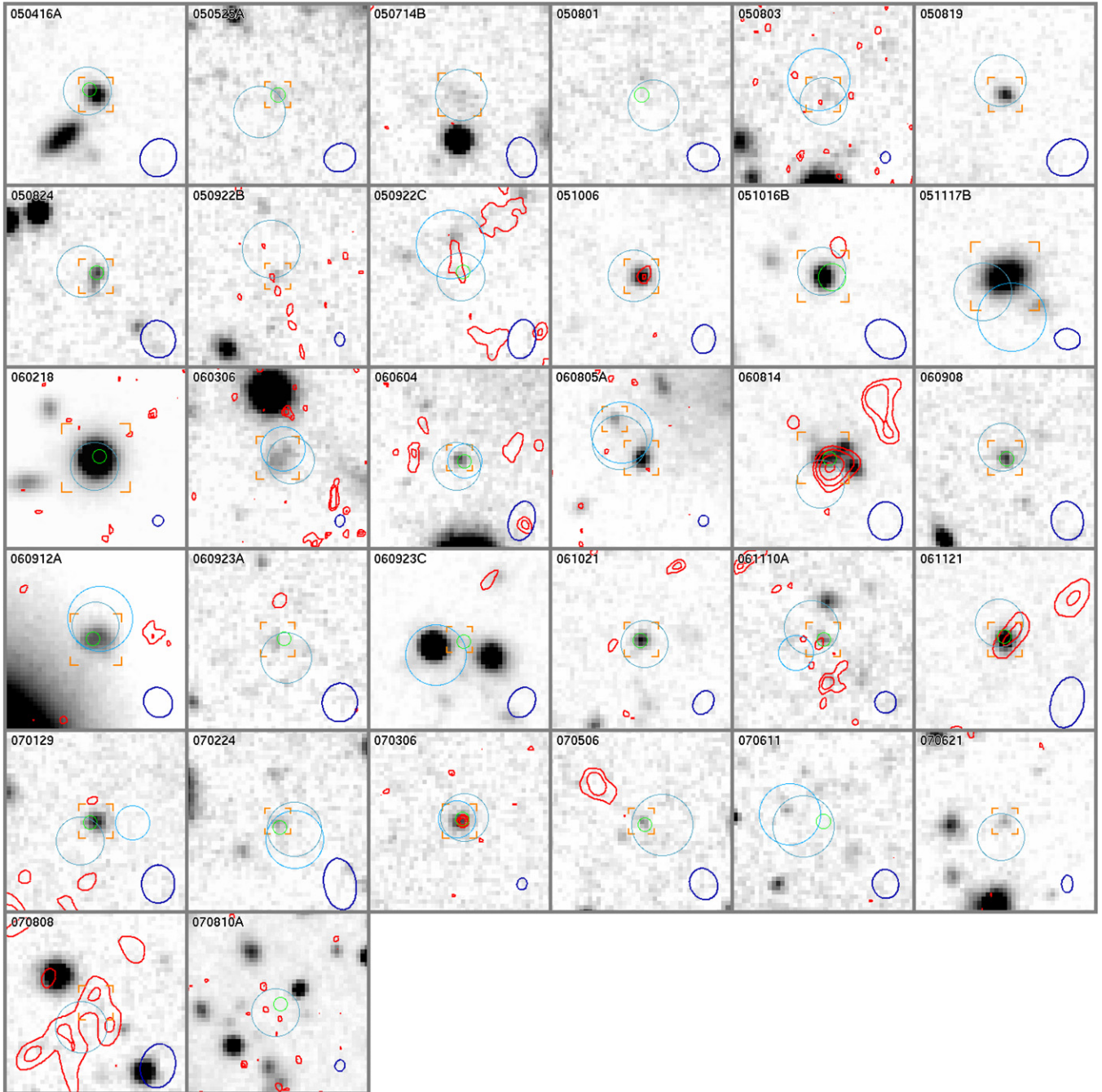


Figure 1. Mosaic of TOUGH *R*-band imaging of 32 GRB host-galaxy positions, with VLA 3 GHz flux-density contours overlaid in red. Contour levels are set to 2.3, 3.0, 4.0, 5.0, and 6.0 times the VLA image rms. (Several images have no flux above the 2.3σ contour within the $10''.5 \times 10''.5$ thumbnail.) The GRB afterglow location is represented by light blue circles (XRT positions from Butler 2007 or Evans et al. 2009) or by a green circle (optical/IR positions). The host galaxy, when detected in the optical images, is centered in the image and identified by an orange box reticle. The VLA synthesized beam is indicated by a dark blue ellipse at bottom right.

2.4. Optical/IR Observations

As discussed in the next section, four of our host-galaxy targets were detected (or are very likely detected). To investigate these systems in more detail, we also acquired deep optical, near-IR (NIR), and mid-IR imaging from a variety of sources in order to perform spectral energy distribution (SED) modeling and measure masses and UV-based SFRs. Many of these data points were taken from the literature (in particular Hjorth et al. 2012; Krühler et al. 2011; Perley et al. 2013), including all data points for the host galaxies of GRB 060814 and GRB 070306. New observations are briefly summarized below.

GRBs 051006 and 061121 were observed with the Low-Resolution Imaging Spectrometer (LRIS; Oke et al. 1995) at Keck Observatory on a number of occasions using several different filters between 2005 and 2014. Observations were reduced using the custom IDL pipeline `lpipe`.¹² Neither of these fields overlaps with the Sloan Digital Sky Survey (SDSS) footprint, so secondary standards within each field were calibrated by acquiring observations of each field with the Palomar 60 inch robotic telescope (P60; Cenko et al. 2006) in the *uBgriz* filters and calibrating relative to Landolt (2009) standards and to

¹² <http://www.astro.caltech.edu/~dperley/programs/lpipe.html>

Table 3
Host-galaxy Optical/NIR Photometry

| Filter | Mag ^a | Flux ^b (μ Jy) | Instrument |
|----------------------|------------------|----------------------------------|---------------------------|
| GRB 051006 | | | |
| <i>u</i> | 24.50 ± 0.20 | 0.75 ± 0.15 | Keck-I/LRIS |
| <i>B</i> | 24.03 ± 0.10 | 1.31 ± 0.13 | Keck-I/LRIS |
| <i>V</i> | 23.44 ± 0.10 | 1.89 ± 0.18 | Keck-I/LRIS |
| <i>R</i> | 22.99 ± 0.07 | 2.33 ± 0.15 | VLT-U1/FORS2 ^c |
| <i>i</i> | 22.51 ± 0.15 | 4.09 ± 0.61 | Keck-I/LRIS |
| <i>z</i> | 22.08 ± 0.10 | 6.00 ± 0.58 | Keck-I/LRIS |
| <i>J</i> | 20.76 ± 0.18 | 8.37 ± 1.51 | P200/WIRC |
| <i>H</i> | 19.82 ± 0.15 | 12.52 ± 1.85 | P200/WIRC |
| <i>K_s</i> | 19.04 ± 0.16 | 16.53 ± 2.62 | P200/WIRC |
| 3.6 | 17.72 ± 0.05 | 22.91 ± 1.08 | <i>Spitzer</i> /IRAC |
| GRB 061121 | | | |
| <i>u</i> | 23.08 ± 0.10 | 3.98 ± 0.38 | Keck-I/LRIS |
| <i>B</i> | 23.34 ± 0.05 | 3.33 ± 0.16 | Keck-I/LRIS |
| <i>g</i> | 22.95 ± 0.04 | 3.95 ± 0.15 | Keck-I/LRIS |
| <i>R</i> | 22.75 ± 0.04 | 3.49 ± 0.13 | VLT-U1/FORS2 ^c |
| <i>i</i> | 22.66 ± 0.10 | 4.12 ± 0.40 | Keck-I/LRIS |
| <i>z</i> | 22.33 ± 0.08 | 5.29 ± 0.40 | Keck-I/LRIS |
| <i>J</i> | 21.29 ± 0.28 | 5.49 ± 1.61 | P200/WIRC |
| <i>K</i> | 20.14 ± 0.19 | 6.16 ± 1.18 | VLT-U1/ISAAC ^c |
| 3.6 | 18.72 ± 0.10 | 9.12 ± 0.88 | <i>Spitzer</i> /IRAC |

Notes.

^a Apparent magnitudes of the host galaxy, not corrected for foreground extinction. SDSS filters (lowercase) are reported in the SDSS system (nearly AB; Fukugita et al. 1996). Other filters are reported as Vega magnitudes.

^b Host-galaxy flux, corrected for foreground extinction.

^c From Hjorth et al. (2012).

various SDSS fields observed the same night during photometric weather. Magnitudes of each galaxy are measured in the reduced images using standard aperture photometry techniques.

GRB 061121 was also observed with the Wide Infrared Camera (WIRC) at Palomar Observatory in *J* band using the replacement engineering-grade detector on 2014 May 10 UT; we acquired 22 60 s exposures on the field. GRB 051006 was observed using the same instrument on 2014 October 11 UT (we acquired 72 15 s exposures in both *J* band and *K_s* bands, although only 42 exposures were included in the final stack due to suboptimal dithering) and on 2014 October 19 UT (we acquired 29 40 s exposures in *H* band) in all three broadband NIR filters. Data were reduced using custom IDL scripts. Magnitudes of the host galaxy were measured using aperture photometry with the calibration established via Two Micron Sky Survey stars in the field.

GRBs 051006 and 061121 were observed by the Infrared Array Camera (IRAC) on board the *Spitzer Space Telescope* during Cycle 9 (GO 90062, PI D. Perley). We used the PBCD reduced files from the *Spitzer* legacy archive and measure magnitudes using the standard zeropoints in the *Spitzer* handbook using an aperture radius of 1''.2.

All host-galaxy photometry for these two objects is presented in Table 3 and in Figure 2.

3. RESULTS

A mosaic of our fields is shown in Figure 1. The grayscale background image represents the TOUGH *R*-band optical imaging from Hjorth et al. (2012); overlaid contours are from our new VLA observations.

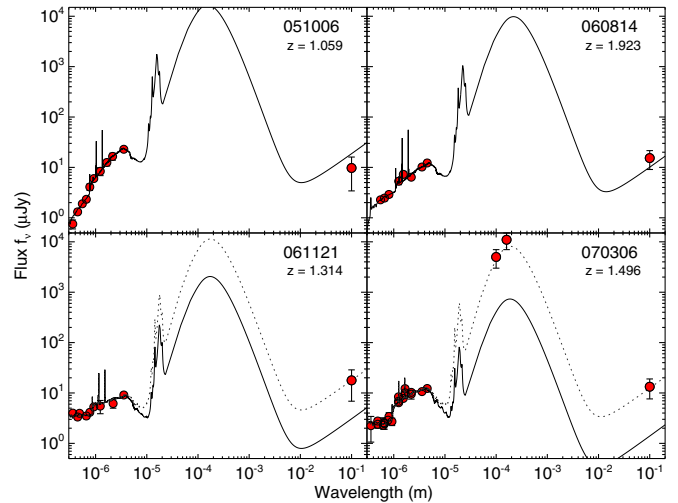


Figure 2. UV-to-radio SEDs of the host galaxies for all four radio-detected sources in our survey. The solid curve shows our multiwavelength SED model of each galaxy fit using the UV-optical-NIR data. In two cases this model accurately predicts the radio flux, indicating that the starburst is optically thin. For the remaining sources the UV-inferred star formation alone underpredicts the radio flux, and a heavily obscured nuclear starburst (dashed line) is also needed to explain the optical and radio measurements simultaneously.

The large majority of the hosts are not detected (the flux at or near the GRB location is consistent with the value expected for random locations in the same image; Table 2), which is no surprise—previous GRB host radio surveys have also produced low detection rates, and only a small fraction of cosmic star formation beyond $z > 0.1$ is in galaxies luminous enough to be individually detected at radio wavelengths, even at these deep levels (e.g., Karim et al. 2011; Dunne et al. 2009).

We do, however, clearly detect at least three (and very likely four) of our targets. Radio sources are detected at the locations of GRBs 060814, 061121, and 070306 at high significance (99% confidence or greater, based on our Monte Carlo analysis). A fourth radio source at the location of GRB 051006 has somewhat lower significance (98%), but the point-like nature of the object and its near-exact consistency with the optical host localization suggest that it is likely real as well.

A plot summarizing our detections and nondetections (along with a sample of measured or predicted fluxes from a sample of field galaxies; Section 6.1) is shown in Figure 3.

3.1. Star-formation, Afterglow, or AGN?

While it is impossible to determine with certainty whether or not any individual detection represents emission associated with star formation in the host galaxy or some other source, it is very likely that all of these detections do indeed correspond to the host-galaxy synchrotron continuum we are seeking.

There are three possible alternatives: emission from an unrelated foreground/background source, emission from an AGN within the host galaxy, or emission from the GRB afterglow. The first case can be ruled out on statistical grounds: only about 0.1% of the sky area in our fields contains a detected source, so the probability of even one chance intersection of a source anywhere in our sample with an unrelated object is low ($\sim 3\%$). An AGN origin within the host is also unlikely, for similar statistical reasons: most sources with a radio flux density close to 10μ Jy flux are star-forming galaxies, not AGN (Kimball et al. 2011; Condon et al. 2012). The optical spectroscopy of these systems, where available (e.g., Jaunsen

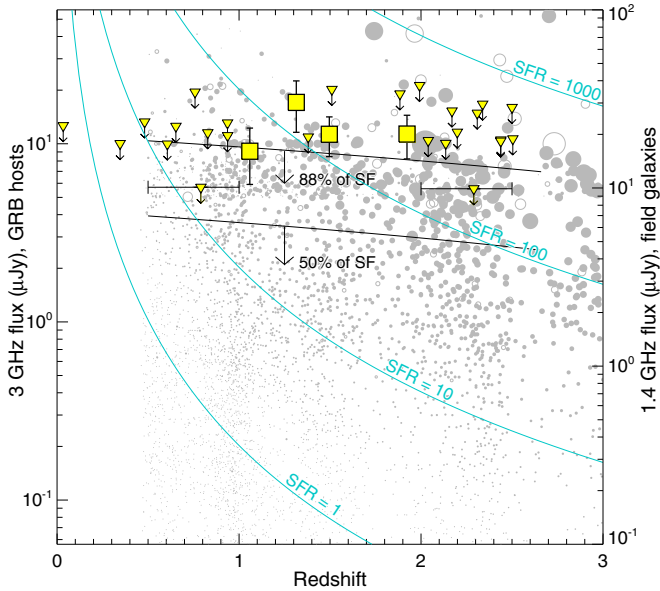


Figure 3. 3 GHz radio fluxes (or 3σ limits) of TOUGH galaxies from the VLA (yellow data points), compared to measured or predicted radio fluxes from galaxies inside GOODS-N (gray circles). To calculate the fluxes of the field galaxies, we use the reduced VLA map of this field presented by Morrison et al. (2010) to directly match against the galaxy catalog of Kajisawa et al. (2011), scaling the 1.4 GHz flux density from that survey to 3 GHz using an average spectral index of $\alpha = 0.75$. For sources that are not detected, we calculate the predicted radio flux density based on the star formation rate estimates in Kajisawa et al. (2011). Galaxies with hard X-ray detections are excluded (open circles). We also evaluate the fraction of cosmic star formation occurring in detected radio galaxies using this method by directly summing the star formation rates of all galaxies above and below $10 \mu\text{Jy}$ (see text for details). The solid lines show a linear fit to the redshift-dependent 50th and 88th percentiles. We also show the limits resulting from stacks of our nondetections at $0.5 < z < 1.0$ and $2.0 < z < 2.5$.

et al. 2008; Jakobsson et al. 2012), also shows no evidence for AGN features.

The afterglow possibility is most difficult to generically rule out: while all of our targets were observed long after the occurrence of the GRB (at least 5 yr and typically 7–8 yr), few GRBs have been followed to faint enough flux/luminosity limits to directly establish the “typical” distribution of afterglow fluxes on multi-year timescales. The discussion in Perley & Perley (2013) suggests that fewer than 10% of all GRB afterglows will be detectable at the flux levels and timescales of our study, but even 10% contamination would be sufficient to produce most of our detections, given the large overall sample size. However, it is possible to examine systems individually to rule out this possibility by other means, in particular, by the observation of any of the following: (1) lack of significant fading versus early-time measurements or limits, (2) physical extension of the radio source, or (3) confirmation of a star-forming galaxy-like SED from IR or submillimeter observations.

GRBs 061121 and 070306 were both observed at radio wavelengths a few days after the GRB and not detected to a flux-density limit of $\sim 100 \mu\text{Jy}$ (van der Horst et al. 2006a; Chandra & Frail 2007). The persistence of a 10–20 μJy radio afterglow at the time of our observations would require these events to have faded by a factor of only 10 or less between ~ 3 days and 2000 days, which is exceptionally unlikely given the behavior of all other known radio afterglows, which invariably show steep power-law decays beginning at ~ 100 days post-GRB or earlier (Chandra & Frail 2012). In the case of GRB 061121, even the possibility of an unusually persistent/late-fading source can be

ruled out directly since a second epoch showing clear fading was reported by van der Horst et al. (2006b). GRB 070306 also has sensitive FIR observations available (from the *Herschel* survey of Hunt et al. 2014) and the host is well detected in that data, consistent with its identification as a star-forming galaxy.

GRBs 060814 and 051006 have no observations available that would enable us to directly rule out an afterglow origin: the sources do not appear obviously extended in our VLA imaging (although the beam size of both observations is comparable to the optical host diameter) and neither was observed in the radio band before our observations. However, given that our optical observations (Section 4) indicate SFRs consistent with those inferred from our measured radio fluxes it is very likely that these detections represent host-galaxy emission as well.¹³

4. ANALYSIS OF DETECTED SOURCES

Since a very high SFR is required for a distant galaxy to be detectable in the radio band, there is good reason to expect that the general properties of the hosts detected in our observations will differ from the overall GRB host population and from “typical” star-forming galaxies; comparison of these galaxies to other radio/submillimeter-selected galaxy populations may also provide insight into the nature of the GRB-galaxy connection overall.

Thanks to a combination of photometry from the literature and our own additional observations, we have excellent optical and infrared photometric data for all four radio-detected systems in our sample (although not yet for the nondetections). After correcting for foreground extinction (Schlegel et al. 1998), we analyzed the optical–NIR SEDs of these galaxies using a similar procedure as previously employed in Perley et al. (2013), using a custom SED-fitting code based on the stellar population-synthesis templates of Bruzual & Charlot (2003) assuming a Chabrier (2003) IMF and metallicity of 0.5 solar.¹⁴ The star formation history of each galaxy is fit as a two-population model with the maximum age of the older population fixed to the age of the universe at the redshift of the host, and the maximum age of the younger population free (but required to be at least 20 Myr). Both populations are assumed to have a continuous star formation history from the maximum age until today, so the overall star formation history is constant with an abrupt increase at an arbitrary time t_{burst} , the age of the current ongoing starburst. Results from these fits are shown in Table 4 and Figure 2.

Radio SFRs are calculated from our observed VLA fluxes using Equation (17) of Murphy et al. (2011) and assuming a synchrotron spectral index ($F_\nu \propto \nu^{-\alpha}$) of $\alpha = 0.75$; the contribution of free-free radiation is assumed to be negligible at this frequency. We note that the calibration of radio SFRs is subject to some systematic uncertainty at about the level of a factor of ~ 2 , and the use of other relations would produce values offset by a common factor; in particular, the measured SFRs and limits would drop by about 40% if we instead used Equation (15) of Yun & Carilli 2002.

For two of our galaxies (GRB 051008 and GRB 060814), the SFRs inferred from the SED fitting procedure and from our radio observations are consistent; that is, an SED model

¹³ Even if these detections were partially due to afterglow, the overall conclusion of this work—that the radio measurements rule out any significant excess star formation beyond that inferred optically for these objects—would be unaltered.

¹⁴ We also attempted models of 1.0 and 0.2 solar metallicity and found the derived parameters to be generally consistent with the 0.5 solar model, but in all cases with significantly higher χ^2 residuals.

Table 4
Physical Properties of Radio-detected Host Galaxies

| GRB | z | M_* (M_\odot) | t_{burst}^a (Myr) | A_V (mag) | $\text{SFR}_{\text{SED}}^b$ ($M_\odot \text{ yr}^{-1}$) | $\text{SFR}_{\text{radio}}^c$ ($M_\odot \text{ yr}^{-1}$) |
|--------|-------|------------------------------------|-------------------------------|------------------------|--|--|
| 051006 | 1.059 | $1.3_{-0.1}^{+0.1} \times 10^{10}$ | 128 | $1.73_{-0.01}^{+0.03}$ | 98_{-1}^{+2} | 51_{-18}^{+22} |
| 060814 | 1.923 | $1.6_{-0.6}^{+1.4} \times 10^{10}$ | 209 | $1.17_{-0.17}^{+0.05}$ | 209_{-53}^{+27} | 256_{-70}^{+160} |
| 061121 | 1.314 | $1.5_{-0.6}^{+0.6} \times 10^{10}$ | 179 | $0.45_{-0.26}^{+0.16}$ | 27_{-6}^{+27} | 160_{-51}^{+58} |
| 070306 | 1.496 | $5.0_{-0.2}^{+0.1} \times 10^{10}$ | 20 | $0.21_{-0.10}^{+0.13}$ | 17_{-5}^{+7} | 143_{-35}^{+61} |

Notes.

^a Age of the young (starburst) component in the stellar population-synthesis model fit.

^b Star-formation rate derived from fitting the UV–optical–IR SED of the galaxy.

^c Star-formation rate derived from the measured radio flux density at the optical host centroid, using the conversion from Murphy et al. (2011).

based only on the optical/NIR photometry accurately predicts (within 2σ) the observed radio flux. For the remaining two sources (GRB 070306 and, marginally, GRB 061121), the radio detection is significantly in excess of what would be expected from the SED fit, indicating that an additional, heavily obscured star-forming component is present in these galaxies. We implement this in Figure 2 by simply adding an additional component with parameters of $A_V = 15$ mag, $t_{\text{burst}} = 10$ Myr, and $\text{SFR} = \text{SFR}_{\text{radio}} - \text{SFR}_{\text{OIRSEDfit}}$; the combined model with this additional component added is shown as a dashed line. (Note that because this component is, by definition, completely obscured in the optical bands, it does not contribute significantly to our photometry outside the radio bandpass and its contribution can be calculated analytically without re-running the SED fit.)

Since luminous and star-forming galaxies are also often very dust-obscured, the afterglow properties of these events in relation to other GRBs (i.e., whether these bursts were optically dark or bright) are also of relevance. We generally rely on previous analysis of the afterglows of these events from the literature, supplemented by our own reanalysis where necessary.

4.1. GRB 051006

GRB 051006 was a relatively weak *Swift* burst with a faint X-ray afterglow. Ground-based follow-up was limited and none of it was conducted less than six hours after the burst occurred. Three optical sources near the burst location were initially identified as afterglow candidates by de Ugarte Postigo et al. (2005), none of which were reported to fade or vary (Cenko et al. 2005; Holland et al. 2005; Ruyantsev et al. 2005). This event may be a dark burst, although unfortunately without deeper, earlier, or redder-wavelength imaging it is impossible to know if this event was actually dust-extinguished to a significant extent. While we classify it as a dark burst for the purpose of subsequent discussions, in this case the assignment is not definitive.

The refined XRT position (Evans et al. 2009) includes only source “B” of de Ugarte Postigo et al. (2005), which has therefore been identified as the probable host galaxy (Hjorth et al. 2012), an association strengthened by its characterization as distant ($z = 1.059$) and highly star-forming (Jakobsson et al. 2012). The galaxy is easily detected in every filter in which we have observed it from the near-UV to $4.5 \mu\text{m}$. Modeling of the SED shows a moderately red optical (rest-frame UV) continuum indicative of a luminous, moderately extinguished galaxy. We derive a best-fit extinction column of $A_V \sim 1.7$ mag and an SFR of $\sim 100 M_\odot \text{ yr}^{-1}$. (While a very low uncertainty on this value is inferred from our SED modeling, the exact

SFR value depends on the assumed star formation history; post-starburst models with lower instantaneous SFRs also give a good fit to the data but were disallowed.) The mass of the older stellar population converged to zero in our fits, and the stellar population is dominated by the current episode with an intermediate age of about 100 Myr.

The UV/optical SFR is comparable to, and possibly even in excess of, what we infer from the radio observations ($\sim 50 M_\odot \text{ yr}^{-1}$), indicating that this system is a luminous, star-forming galaxy but not an extreme SMG and it does not harbor large amounts of heavily obscured star formation.

4.2. GRB 060814

GRB 060814 is a prominent dark burst, with clear detections of an afterglow in *K* band but not in any bluer filters (e.g., Levan et al. 2006; Campana et al. 2007), unambiguously indicating an origin from an obscured region (but not an *extremely* obscured region, or else the *K*-band afterglow would also be undetectable: a reasonable bracket on the dust column is $1.8 < A_V < 8$ mag).

The host galaxy of GRB 060814 is discussed in some detail in Perley et al. (2013), as well as in Hjorth et al. (2012) and Jakobsson et al. (2012). It is well detected in every optical/NIR band and is also one of the physically largest GRB hosts known, showing a complex, distorted morphology with an angular extent of $\sim 2''.5$ (21 kpc at the spectroscopic redshift of $z = 1.923$) that makes it easily resolved even in ground-based optical observations (Figure 1; note that, as noted by Jakobsson et al. 2012, the southwest source visible in the VLT image is a foreground object at $z = 0.84$ but the remainder of the emission is from the host). Our SED modeling indicates a very large SFR of approximately $200 M_\odot \text{ yr}^{-1}$ after correction for dust extinction.

The 3 GHz flux density of $\sim 15 \mu\text{Jy}$ measured for this source translates to an SFR of $260 M_\odot \text{ yr}^{-1}$ at this redshift. This is clearly within the range of ULIRGs and SMGs—although, again, the fact that this is in agreement with the optical SFR indicates that the galaxy is not actually optically thick but simply at the high-luminosity end of the population of ordinary, moderately obscured galaxies. However, this galaxy is clearly in a starburst phase, given its modest mass ($\sim 1.5 \times 10^{10} M_\odot$) and highly distorted morphology.

A second source is also (marginally) detected in the VLA image a few arcseconds to the northwest of the host galaxy. While there is no counterpart at this position in the *R*-band image shown in Figure 1, a very red source is clearly detected there in the NIR imaging (*Hubble Space Telescope* [*HST*] WFC3-IR, *K* band, and IRAC). No redshift is available for this object, but given the relative rarity of very red objects and of faint radio sources, it is possible that it may also be associated with the host system in some way.

4.3. GRB 061121

Unlike the other events in this list with radio detections, GRB 061121 was definitely not a dark burst. Its afterglow was very bright in the optical and ultraviolet at early times, was detected in all UVOT filters including UVW2 (which partially overlaps the Lyman limit at the host redshift), and its optical spectral index was relatively normal (Page et al. 2007), consistent with an unextinguished event. These are all typical properties of the optically bright GRB afterglows commonly studied (and targeted for host follow-up) in the pre-*Swift* era.

Nevertheless, the event is hosted within a very luminous, radio-detected galaxy. The $17 \mu\text{Jy}$ source detected in the VLA

image indicates an SFR of $160 M_{\odot} \text{ yr}^{-1}$ (though with a substantial statistical uncertainty, since the rms noise level of $5.5 \mu\text{Jy}$ is higher than average). The host galaxy of this event is well detected in all optical/NIR bands; the color is fairly blue and the inferred optical extinction is modest ($A_V \sim 0.4$ mag). The UV/optical SFR is $30 M_{\odot} \text{ yr}^{-1}$.

A secure interpretation of this system is somewhat complicated by the weak detection: nominally, the radio and optical SFRs are inconsistent by a factor of ~ 5 , but the radio measurement itself is uncertain by a factor of ~ 2 simply due to the marginal detection and additional systematic uncertainties underlie the SFR derivations from UV and radio fluxes. Nevertheless, taking our current measurement at face value indicates that a (modest) additional optically thick starburst is present in the galaxy. However, as the GRB itself shows no evidence of significant obscuration, GRB 061121 must have occurred in the relatively unobscured regions of the galaxy seen in our optical observations and *not* from the optically thick component.

A second source with similar flux level similar that of the host galaxy is present approximately $4''.5$ to the northwest. However, there is no optical counterpart at this location and the probability of chance occurrence of a noise peak at this flux level within a $4''.5$ radius search region is significant, so we do not currently regard this object as real.

4.4. GRB 070306

GRB 070306 was a heavily obscured GRB (Jaunsen et al. 2008); its afterglow was detected only in *K* and *H* bands and its color between these filters was very red (consistent with $A_V \sim 5.5$ mag of extinction). The optical/IR host-galaxy observations of this target are all the same as previously presented in Perley et al. 2013 (which in turn incorporated data from the earlier studies of Jaunsen et al. 2008 and Krühler et al. 2011), and despite somewhat different modeling assumptions, we derive nearly identical parameters, including a very modest UV/optical SFR of $\sim 15 M_{\odot} \text{ yr}^{-1}$. The radio (and *Herschel*-) inferred dust-unbiased SFR for this source is unambiguously much larger: approximately $\sim 140 M_{\odot} \text{ yr}^{-1}$, indicating the presence of substantial heavily obscured star formation. The large column toward the GRB is consistent with an interpretation in which the burst itself actually occurred in the dominant, heavily obscured part of the host.

The best-fit stellar mass of this galaxy is the largest of the four radio-detected hosts in our sample ($\sim 5 \times 10^{10} M_{\odot}$). *HST* WFC3-IR imaging is available for this source (Perley et al. 2013); while the galaxy is resolved there are no obvious signs of an ongoing major merger at the resolution of the IR camera, though the galaxy does show some asymmetry (an extension to the northwest but not the southeast) and two much fainter companions (possibly, extended spiral arms) are present nearby.

5. NOTABLE NON-DETECTIONS

Since the radio SFRs inferred for two of our detections are close to the values inferred from UV/optical observations of the same galaxies (ruling out significant optically thick star formation), we also examined our sample to determine if any of our *nondetections* imply radio luminosities close to our measured limits, which would also rule out additional optically thick star formation. The available TOUGH data only provide two filters (*R* and *K_s* bands), which are not adequate to evaluate the dust-corrected UV SFRs of these sources (dust corrections are critical even for reasonably optically thin galaxies). However,

in two cases additional data are available that suggest SFRs comparable to our radio limits, discussed below.

5.1. GRB 060306

The host galaxy of GRB 060306 was studied in Perley et al. (2013). Its optical/IR SED indicates a very luminous and dust-obscured galaxy; we previously derived an estimated SFR of $245^{+130}_{-67} M_{\odot} \text{ yr}^{-1}$ at the favored redshift of $z = 1.55$ (Perley et al. 2013). Our limiting radio SFR of $\text{SFR} < 260 M_{\odot} \text{ yr}^{-1}$ is not inconsistent with this figure, but indicates that the galaxy—while probably a ULIRG, based on the bolometric luminosity implied by our fit—contains minimal optically thick star formation. This conclusion is similar to the one reached for this system from our previous study of dark GRB hosts with the VLA using *C*-band observations (Perley & Perley 2013) but significantly more constraining. This target is at a declination near the satellite belt and the observation was strongly impacted by satellite RFI; had we achieved similar rms sensitivities during our VLA observation of this target as were achieved for most of the other fields, we predict that the host would have been detected.

5.2. GRB 060218

GRB 060218 is by far the lowest-redshift event in the sample and represents something of a special case: it is extremely optically underluminous compared to almost any class of galaxy, is physically very small (~ 1 kpc), and clearly has little in common with the LIRG-like systems that our survey otherwise targets. Nevertheless, given the close proximity of this galaxy (redshift $z = 0.033$) the radio limit reaches very low SFRs. This source was observed only in A-configuration and the beam ($0''.6$ FWHM) is significantly smaller than the full optical extent of the host, but the core of the galaxy (in which the large majority of the star formation is occurring) is very compact: Kelly et al. (2014) estimate a half-light radius of only $0''.31$ from *HST* imaging, so at least half of the galaxy’s star formation is contained within one VLA beam. Conservatively allowing for a factor of ~ 2 loss in sensitivity for this reason, our observations place a limit on the radio-derived SFR of $< 0.04 M_{\odot} \text{ yr}^{-1}$, similar to the optically derived value of $0.03\text{--}0.04 M_{\odot} \text{ yr}^{-1}$ (Wiersema et al. 2007; Levesque et al. 2010). Radio observations of very young, dwarf starburst galaxies often underpredict the optically derived SFRs because the ongoing starburst has not had sufficient time to produce enough supernova remnants to accelerate the electrons that produce the radio continuum (Roussel et al. 2003), so a mild discrepancy would not be surprising for this system.

5.3. Stacked Observations

To provide a limit on the *average* flux density of the host galaxies in our sample, we summed all 28 images in which no significant host emission was detected, using the host position from the VLT images as the alignment point. Both a direct sum and a sum weighted using the inverse of the rms were performed. In neither case is significant emission detected at the position of the stacked host galaxies. We measure an averaged flux density of $1.77 \pm 0.85 \mu\text{Jy}$ in the unweighted stack, or $1.47 \pm 0.77 \mu\text{Jy}$ in the weighted stack. At the mean redshift of the stacked sample of $z = 1.54$, this corresponds to a 2σ limiting average SFR of $41 M_{\odot} \text{ yr}^{-1}$, although we caution that the flux–SFR conversion is a strong function of redshift and

significant evolution is expected in the population across this period.

Somewhat more informative limits can be determined by stacking in limited redshift intervals. Using the same procedure as in the previous paragraph, we stacked the images of 6 fields at $0.5 < z < 1.0$ and 9 fields at $2.0 < z < 2.5$; in the weighted stacks we do not detect a source in either case with nominal fluxes of $1.05 \pm 1.55 \mu\text{Jy}$ and $1.60 \pm 1.33 \mu\text{Jy}$, respectively. Equivalent 2σ limiting SFRs are $\text{SFR} < 11.7 M_{\odot} \text{yr}^{-1}$ at $z = 0.79$ and $< 143 M_{\odot} \text{yr}^{-1}$ at $z = 2.29$. These binned limits are shown in Figure 3.

6. POPULATION ANALYSIS AND DISCUSSION

6.1. Detection Rate and Connection to Cosmic Star-formation

The successful detection of four objects with SFRs ranging up to $\sim 300 M_{\odot} \text{yr}^{-1}$ clearly indicates that GRBs can form in extremely luminous galaxies in at least some circumstances, while at the same time the low overall detection rate (approximately 15%: 4 out of 24 low-rms fields at known $z < 2.5$, or 4 out of 29 including unknown- z) indicates that most GRB hosts are not systems of this type. While this qualitative observation obviously rules out the most extreme models (e.g., in which GRBs are never produced by very luminous galaxies, or are produced only within very luminous galaxies), we wish to be more quantitative: is the observed detection rate above, below, or consistent with the detection rate *expected* if GRBs neither prefer nor avoid the universe’s most luminous galaxies?

We followed two different approaches to predict the expected numbers under this SFR-tracing null hypothesis. First, we numerically integrated long-wavelength luminosity functions from both *Spitzer* (the $24 \mu\text{m}$ luminosity functions of Pérez-González et al. 2005, converted to total infrared luminosity) and *Herschel* (the bolometric IR luminosity functions of Gruppioni et al. 2013) to determine the quantity of star formation occurring in galaxies with SFRs large enough to be detected in our survey at each redshift, using $10 \mu\text{Jy}$ as a characteristic flux limit of our survey. Specifically, we calculate the quantity

$$f_{\text{det}}(z) = \frac{\int_{L_{\text{limit}}(z)}^{\infty} L \times \phi_L(L, z) dL}{\int_0^{\infty} L \times \phi_L(L, z) dL}.$$

Here, L_{limit} is the total IR luminosity of a star-forming galaxy with $F_{3\text{GHz}} = 10 \mu\text{Jy}$ (our radio flux limit) and $\phi_L(L, z)$ is the IR luminosity function as taken from the literature.¹⁵ For the *Herschel*-based luminosity functions of Gruppioni et al. (2013) the radio-detectable fraction drops gradually from 40% to 25% over our redshift range of interest ($0.5 < z < 2.5$). The *Spitzer*-based fraction is flat at 20% between $0.5 < z < 1.5$; it then rises sharply above $z > 1.5$ (Figure 4) although $24 \mu\text{m}$ based luminosities are highly uncertain at these redshifts (e.g., Papovich et al. 2007; Elbaz et al. 2010; Nordon et al. 2010).

Taking a more direct approach, we also calculated a sum of the actual inferred SFRs of all radio-detected galaxies from GOODS-N, using the redshift and SFR catalogs of Kajisawa

¹⁵ Since it only examines dust-radiated light, this procedure is insensitive to *unobscured* star formation and likely overestimates the fraction of star formation in very luminous galaxies, but as most of the UV light in even “normal” galaxies is obscured at high-redshift (e.g., Meurer et al. 1999) the correction factor is likely to be small. Based on the estimates of Oesch et al. (2010), the dust-*uncorrected* UV star formation rate density is not more than 20% of the FIR-based star formation rate density at any redshift relevant to this study.

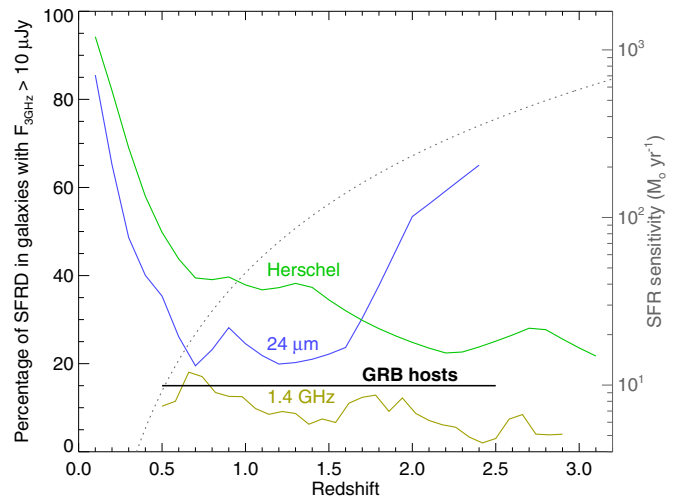


Figure 4. Fraction of star formation detectable to the approximate average sensitivity of our EVLA survey ($F_{3\text{GHz}} > 10 \mu\text{Jy}$) as a function of redshift, as calculated by three different means: integrating the *Herschel*-PACS luminosity functions of Gruppioni et al. (2013), integrating the *Spitzer*-MIPS luminosity functions of Pérez-González et al. (2005), and directly matching sources detected at 1.4 GHz in Morrison et al. (2010) to the catalogs of Kajisawa et al. (2011). Between approximately $z = 0.5$ and $z = 2.5$, radio observations probe about the same fraction of cosmic SFR with little variation with redshift, even though the *absolute* star formation rate sensitivity limit (right axis and dotted line) varies strongly with redshift. Our GRB detection fraction is shown as a single bin, since our sample size is not yet large enough to provide redshift-resolved constraints.

et al. (2011) matched with detected sources in the 1.4 GHz radio map of Morrison et al. (2010), which has similar resolution and sensitivity as our *S*-band observations. We independently measured 1.4 GHz flux densities of sources in this map using the same procedure as for our GRB hosts, and scaled these values to 3 GHz using an average spectral index of $\alpha = 0.75$; galaxies with hard X-ray detections were excluded as AGN-dominated. (The scaled, predicted fluxes from this method are shown as gray circles in 3.) We then calculate the detectable fraction:

$$f_{\text{det}}(z_{\text{min}}, z_{\text{max}}) = \frac{\sum_{\text{galaxies with } F > 10\mu\text{Jy}} \text{SFR}}{\sum_{\text{all galaxies}} \text{SFR}},$$

where both sums are restricted to galaxies over the range $z_{\text{min}} < z < z_{\text{max}}$. SFR estimates are taken from the “IR+UV” column of Kajisawa et al. (2011), which uses the sum of UV and $24 \mu\text{m}$ SFRs for MIPS-detected galaxies or the dust-corrected UV SFRs otherwise. Since the GOODS-N field is small, this method is more affected by small-number statistics than the *Herschel/Spitzer* maps, but the overall result is similar, between 5% and 15% on average between $z = 0.5$ and 2.2.

Note that despite the fact that there is rapid cosmic evolution with redshift over this range (the characteristic luminosity of galaxies increases by over an order of magnitude) and strong variation in sensitivity with redshift in the radio bands, these two effects largely cancel out, and our survey is sensitive to about the same *proportion* of cosmic star formation at every redshift: specifically, at *any* redshift between approximately $z = 0.5$ and $z = 2.5$, the galaxy population at or above our sensitivity limit is responsible for $\sim 5\%$ – 30% of the cosmic SFR rate density at that time. Therefore, regardless of the actual redshift distribution of the sample (as long as it is within this range), for an unbiased star formation tracer we expect a detection rate of 5% – 30% .

This range agrees quite well with our measured fraction of 15%, so our radio data alone are consistent with the hypothesis that GRBs represent unbiased tracers of star formation. However, note that the prediction is systematically uncertain to a factor of ~ 2 – 3 depending on the comparison study, and our own measurement is statistically uncertain to within a factor of ~ 2 – 3 (Section 6.2), so this is not yet a highly constraining test on its own. Fortunately, we know much more about these galaxies than simply their radio fluxes and SFRs and can apply additional, more constraining tests using additional parameters; we will return to this issue in Section 6.3.

We note that our detections cover a wide redshift range ($1.06 < z < 1.92$) and show no correlation between the observed flux and redshift, in support of our assertion that the detection rates and radio fluxes for the most luminous star-forming galaxies should be approximately constant with redshift. We do not detect any galaxies at $z < 1$ or $2 < z < 2.5$ within our sample, but given the small sample size this is not likely significant; indeed, previous targeted surveys have detected hosts at radio wavelengths at both higher and lower redshifts.

6.2. Intrinsic Flux and Luminosity Distributions

The convenient scaling between sensitivity and characteristic luminosity discussed above means that the detection fraction alone provides perhaps the most informative constraint on the degree to which GRBs prefer or avoid the most luminous galaxies, as long as it is understood that “most luminous” is treated as a *relative* statement (versus other galaxies at the same redshift) as opposed to an absolute one (e.g., above some fixed luminosity threshold).¹⁶ Still, it is informative to produce formal constraints on the intrinsic flux distribution inferred from our survey and on the luminosity distribution as a function of redshift.

While we have treated our detection fraction as a single value above a single flux point, in reality our sensitivity limits do vary from field to field (from $7.1 \mu\text{Jy}$ for 070306 to $18.6 \mu\text{Jy}$ for 060306 and 070621, at 2.5σ). Translating our results into a statistical constraint on the actual fraction with flux density above a specific value therefore requires a prior on the shape of the flux-density distribution of the underlying population. We assume it follows a single flux-weighted Schechter function ($dN/dS \propto S \times (S/S_*)^\alpha e^{-S/S_*}$, with $\alpha = -1.1$ and S_* a free parameter)—a reasonable assumption if GRBs approximately trace star formation given the redshift-independence of the characteristic flux as described in the previous paragraph—although the results are not particularly sensitive to the choice of model. Based on a simple Monte Carlo analysis using the distribution of 2.5σ sensitivity limits achieved by our survey and including the effects of measurement noise, we calculate that, intrinsically, between 9% and 23% of all $z < 2.5$ GRB hosts have $S_{3\text{GHz}} > 10 \mu\text{Jy}$ at 90% confidence.

We do not have enough detections to constrain the flux-density function shape in a meaningful way, but it is worth noting that we detect *no* host galaxies with a flux density greater than $20 \mu\text{Jy}$ at any redshift, so the luminosity function is at

¹⁶ There are physical reasons why the relative definition may indeed be more relevant than an absolute one: the SFRs of all types of star-forming galaxies appear to scale up with redshift in similar ways, such that luminous LIRGs and ULIRGs at $z \sim 0$ are primarily mergers (Sanders & Mirabel 1996; Veilleux et al. 2002), whereas galaxies of similar luminosity at $z \sim 1$ – 2 are primarily normal disks; major mergers instead produce extreme ULIRGs and HyLIRGs at these redshifts (Chapman et al. 2003; Fadda et al. 2010; Kartaltepe et al. 2010, 2012).

least moderately steep (as indeed is expected for the high-luminosity exponential tail of a Schechter-like population). This observation produces some tension with the pre-*Swift* study of Berger et al. (2003), which reported a detection rate of 20% at a 2.5σ radio sensitivity level of 20 – $30 \mu\text{Jy}$ in *X* band (equivalent to 40 – $60 \mu\text{Jy}$ in *S* band for $\alpha = 0.75$). The distribution of redshifts for the targets in their survey, and for their reported detections, is similar to what is seen in our own. We suspect that some of their reported detections may have either been noise fluctuations or affected by late-time contribution from the GRB afterglow, and/or that their detection of several intrinsically rare, luminous objects in such a small sample was a statistical fluke.

Again assuming that the ratio of host luminosity sensitivity and characteristic star-forming galaxy luminosity can be assumed to be constant over our redshift range, the flux-distribution constraint can be translated to an equivalent SFR-distribution constraint at any redshift within a factor of ~ 2 simply by calculating the equivalent limiting SFR for our $10 \mu\text{Jy}$ limiting flux at each redshift. For example, we estimate that 9%–23% of GRB hosts at $z \sim 0.6$ have $\text{SFR} > 14 M_\odot$, 9%–23% of hosts at $z \sim 1.0$ have $\text{SFR} > 50 M_\odot$, 9%–23% of hosts at $z \sim 1.5$ have $\text{SFR} > 125 M_\odot$, and 9%–23% of hosts at $z \sim 2.0$ have $\text{SFR} > 250 M_\odot$.

6.3. The Stellar Masses of GRB-selected Submillimeter Galaxies

In the study of Perley et al. (2013) we examined the influence of several parameters on the rate at which a galaxy produces GRBs relative to stars in general (the GRB production “efficiency,” $\epsilon \equiv R_{\text{GRB}}/\text{SFR}$). The total SFR of a galaxy had no discernible effect on its efficiency for GRB production over the redshift range we examined ($z \sim 0.5$ – 3). The consistency of the number of radio detections in our TOUGH-VLA sample with the expected number for an unbiased tracer (Section 6.2) can be seen as an extension of this conclusion up to the highest SFRs. On the other hand, we previously found that mass had a strong effect (GRBs in the most massive galaxies are much rarer than expected given the contributions of such galaxies to overall cosmic star formation), so it is relevant to investigate the masses of the luminous galaxies detected in this effort as well.

Our optical/NIR SED modeling provides firm estimates of the stellar masses of all four-radio detected galaxies (Section 4),¹⁷ enabling us to evaluate whether or not their mass distribution is consistent with that which would be expected for the most luminous members of a star formation-sampled galaxy population. Results are plotted in Figure 5; note that due to the strong cosmic evolution in the mass–SFR relation over the redshift interval spanned by our sample, we have applied a redshift scaling factor of $(z/1.5)^{-1.8}$ to all SFRs¹⁸ to remove this effect from our analysis. The masses of our radio-detected hosts (large yellow squares) are well below the masses of almost all high-luminosity (scaled $\text{SFR} > 100 M_\odot \text{yr}^{-1}$) galaxies in the

¹⁷ Of course, stellar masses in this and all other studies are subject to systematic uncertainties related to the input assumptions. As discussed earlier, our modeling assumes a Chabrier IMF and a two-population star formation history, which are thought to be the most realistic assumptions for fitting galaxies of this type (Michałowski et al. 2012a, 2014a). However, even when employing alternative star formation histories we derived consistent results, and in this analysis we scale all comparison studies to the same IMF.

¹⁸ The power-law index of -1.8 is derived by comparing the mass–SFR relation at $z \sim 1$ from Elbaz et al. (2007) to that at $z \sim 2$ from Daddi et al. (2007). After this adjustment these two relations lie along the same solid green line, shown on the plot.

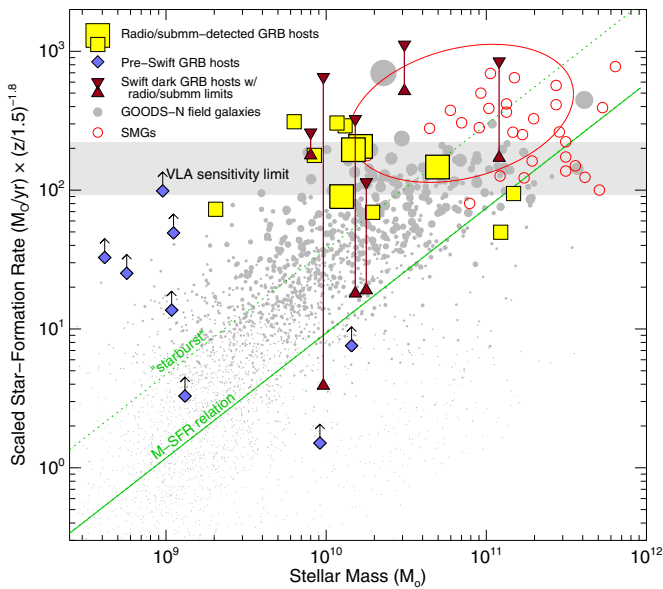


Figure 5. Stellar masses and star formation rates of GRB host galaxies and field galaxies from our sample and from the literature at $0.8 < z < 2.2$. In addition to the four radio-detected targets presented here (large squares), we plot GRB hosts with FIR/submillimeter/radio detections from Michałowski et al. (2008) and Hunt et al. (2014) (small squares), as well as events from Perley & Perley (2013) that were *not* detected at radio wavelengths, showing the range between the UV/optical SFR as a minimum and the radio SFR as a maximum in the form of vertical lines. Pre-*Swift* GRB hosts are also shown, plotted as lower limits since sensitive radio/submillimeter limits are generally not available. For the GOODS-N field galaxies (Kajisawa et al. 2011) we take $24 \mu\text{m}$ SFRs when available or dust-corrected UV SFRs otherwise, correcting by an IMF adjustment factor of 1.6 to convert from Salpeter to Chabrier/Kroupa. To minimize the impact of redshift evolution and varying radio luminosity sensitivity over this broad interval, we have scaled all SFRs by a factor of $(z/1.5)^{-1.8}$. The galaxy $M - \text{SFR}$ relation at $z = 1-2$ (Elbaz et al. 2007; Daddi et al. 2007) is also shown; galaxies with sSFRs a factor of five above this relation are defined as “starbursts” (Rodighiero et al. 2011). Compared to field galaxies of comparable luminosities and compared to SMGs (small circles indicate individual systems from Michałowski et al. 2010; the large red ellipse represents the region of parameter space inhabited by SMGs in Gruppioni et al. 2013), the most luminous GRB hosts have unusually low stellar masses.

GOODS-N field as estimated by Kajisawa et al. (2011), with the exception of GRB 070306 whose stellar mass is fairly typical for a luminous galaxy.

Because the GOODS-N field is small, galaxies with the highest masses and SFRs are not well sampled, and this comparison may be biased by cosmic variance. However, wide-area FIR and submillimeter surveys select galaxies out of the underlying field population in a manner similar to the way in which our four detections were selected out of the parent TOUGH sample, and provide a comparison sample over a much larger volume suitable to evaluate the GRB-SFR connection within luminous galaxies in more detail. In Figure 5 we plot both a sample of submillimeter-selected SMGs from Michałowski et al. (2010) and a broad ellipse that encloses nearly every SMG in the *Herschel* studies of Rodighiero et al. (2011, their Figure 1) and Gruppioni et al. (2013, their Figure 15).

The stellar masses of SMGs are subject to some systematic uncertainty, with the values derived varying by up to a factor of almost 10 depending on the assumptions considered in modeling the stellar populations (in particular the SFR history, choice of stellar templates, and IMF; see Michałowski et al. 2012a, 2014a for an in-depth discussion of this issue). The offset of the Michałowski et al. (2012a) SMGs and the *Herschel* samples in Figure 5 is largely a reflection of this. In our own modeling we

have assumed a two-component stellar population, which tends to lead to *higher* masses when applied to SMGs.¹⁹ However, even when compared to SMG mass catalogs derived using assumptions leading to lower masses, very few such galaxies are of masses comparable to the GRB hosts in our sample (three out of four of the GRB hosts in our sample have a mass below $2 \times 10^{10} M_{\odot}$, and SMGs below this value are extremely uncommon).²⁰

Our sample size is small (four hosts), but can be expanded by considering additional submillimeter/radio-detected GRB hosts from previous work (although these were not selected out of an unbiased GRB sample, and one of which was detected in radio but undetected at submillimeter wavelengths and could conceivably be afterglow-contaminated; Section 6.1). Michałowski et al. (2008) analyzed the SEDs of four submillimeter- or radio-detected GRB hosts (all at $z = 1-2$) and derived stellar masses less than $2 \times 10^{10} M_{\odot}$ (after IMF adjustment) for all four cases. Furthermore, a few pre-*Swift* host galaxies with even lower stellar masses ($\sim 10^9 M_{\odot}$) have measured UV dust-corrected SFRs indicating predicted radio fluxes similar to our detection level (Figure 5). This result is fully consistent with the distribution observed within our own sample and provides additional support to the notion that, even within SMGs and at $z \sim 1.5$, GRBs exhibit a strong aversion against the most massive galaxies.

GRBs in more massive submillimeter-like galaxies do exist, particularly among dark GRB hosts: combining the dark GRB samples of Perley & Perley (2013) and Hunt et al. (2014), four galaxies with masses of $> 5 \times 10^{10} M_{\odot}$ have been detected by *Herschel* or VLA. However, it is important to keep in mind that these are not unbiased samples, but are selected out of the much larger parent pool of the hundreds of *Swift* and pre-*Swift* GRBs and subject to very specific selection criteria, explicitly favoring heavily reddened afterglows and IRAC-bright (massive) galaxies.

Further long-wavelength observations of unbiased samples will be necessary to confirm this interpretation, but these results offer support to the notion that GRBs preferentially explode in low-mass galaxies (albeit without entirely excluding massive galaxies) across the entirety of the host luminosity range.

6.4. GRBs from Optically Thick Super-starbursts?

Traditional SMGs are extremely heterogeneous systems, with dense, intensely star-forming cores that dominate the bolometric output (and overall SFR) as well as less-obscured outer components that dominate the UV/optical luminosity despite their minor contributions to the total energy budget. Since the degree of obscuration to a GRB can be determined from afterglow observations, when a GRB is localized to an SMG it is often possible to crudely discern where within the host the GRB actually occurred: a GRB that is extremely obscured

¹⁹ We verified this explicitly by refitting all of the SMGs with UV-through-NIR host detections from Michałowski et al. (2008) using the same procedure employed for our GRB hosts and derived masses consistent with the higher values that work.

²⁰ An additional systematic issue concerns the completeness and purity of the submillimeter and *Herschel* samples discussed here. While recent *Herschel* papers claim very high redshift completeness (80%–100%; Gruppioni et al. 2013), it is not inconceivable that some *Herschel* sources are incorrectly identified with bright counterparts in the *Herschel* beam when in fact they should be associated with fainter, lower-mass sources, leading to a bias in favor of larger masses and lower redshifts. ALMA follow-up will be necessary to clarify this issue unambiguously, but since our survey targets a relatively nearby redshift range ($z = 1-2.5$) at which counterpart matching is fairly secure we will treat the published *Herschel*/James Clerk Maxwell telescope results at face value.

($A_V > 5$ mag) likely occurred within the central, optically thick starburst, whereas a GRB showing little or no obscuration must have occurred in the optically thin regions.

The sites of GRBs within heterogeneous SMGs represent an important indicator of the connection between GRB production and star formation, since the dusty inner regions of these galaxies are likely to be much more metal-rich than the unobscured regions but are also much more intensely star-forming. Different physical hypotheses for the origin of the variations in the GRB rate have opposing predictions for the rate of GRBs in dusty central starbursts: if the GRB efficiency is governed primarily by metallicity we would expect GRBs to avoid the obscured regions; if the GRB efficiency is governed primarily by star formation intensity then we would expect GRBs to be abundant in obscured regions.

Only two of the galaxies in our sample show evidence of strong dust heterogeneity: the hosts of GRB 061121 and GRB 070306. These two sources offer a split verdict on the location of the burst within these hosts. The bright, blue afterglow of GRB 061121 is clearly not heavily obscured: Page et al. (2007) estimate $A_V = 0.4$ – 1.6 mag depending on the assumed dust law and intrinsic spectral index. This is consistent with the mean obscuration inferred from the host UV/optical observations in our SED modeling and clearly situates the event outside the heavily obscured star-forming region that our data (tentatively) indicate is also present. On the other hand, GRB 070306 was extremely obscured and this event is quite consistent with having occurred within the dusty starburst region of an SMG.

Two events are not sufficient to give a definitive answer as to whether or not GRBs preferentially prefer or avoid any particular area of this type of galaxy. Again, however, comparison to other literature results can offer a hint. Among the four pre-*Swift* submillimeter host galaxies discussed in Michałowski et al. (2008)—all of which showed evidence for dust heterogeneity as discussed above—three were optically bright and therefore unobscured; only one, GRB 000210, was consistent with an obscured origin. While it is conceivable that this is a selection effect (few optically obscured dark bursts were sufficiently well localized in the pre-*Swift* era to perform a host-galaxy search), dedicated studies of dark GRBs have failed to turn up many counterexamples: all of the submillimeter or radio-detected galaxies of dark GRBs in the studies of Hunt et al. (2014) and Perley & Perley (2013) have corresponded to optically reddened galaxies in which the gap between the extinction-corrected UV and the FIR/radio SFRs is small. This may be a selection effect on its own (in Perley & Perley 2013 we explicitly favored events with redder optical colors and higher inferred dust-corrected SFRs; Hunt et al. 2014 observed only hosts previously targeted with, and detected by, *Spitzer*), but even so, additional analogs of GRBs 070306 and 000210 have proven remarkably elusive considering that the bolometric luminosity of obscured regions of typical SMGs can exceed the luminosities of the unobscured regions by huge factors.

Larger, preferentially unbiased campaigns to observe GRB host galaxies at submillimeter/radio wavelengths will be required to resolve this question unambiguously. However, the small numbers of GRBs found to date within heavily obscured regions of very luminous, blue galaxies (two) in relation to the much larger number found in unobscured regions of similar galaxies (as many as five, although cf. our discussion in Section 6.1) may point toward nonuniform efficiency for GRB production across the different parts of these hosts: preference

for outer, unobscured regions and for blue, moderately extinguished ULIRGs, and an aversion against the central, dusty starburst.

7. CONCLUSIONS

We have presented sensitive radio observations to search for host-galaxy synchrotron emission at the locations of 31 GRBs selected from the TOUGH host-galaxy survey. Among 25 sources at known redshift, we securely detect three (and probably detect a fourth), indicating that approximately 15% of all GRBs occur in the most luminous galaxies (those detectable at radio wavelengths).

As about 15% of all cosmic star formation occurs in galaxies of similar radio luminosity over this redshift range, our results suggest that GRBs are not particularly biased tracers of cosmic star formation with regard to bulk galaxy SFR, neither markedly preferring nor avoiding galaxies with the highest luminosities. However, the stellar masses of the systems we detect with the VLA are about an order of magnitude lower than those of field-selected star-forming galaxies of similar luminosities and redshifts. Both of these results are consistent with the trends seen in Perley et al. (2013), in which the GRB host population had significantly lower mass than an SFR-weighted field sample but matched the expected distribution of UV-based SFRs.

Why would GRBs appear to preferentially avoid massive galaxies, yet not avoid the most actively star-forming galaxies, when it is well established that the most massive galaxies are, on average, the most-star-forming? Our results suggest that the GRB efficiency, while depressed in massive galaxies, must also be *enhanced* in galaxies with high *specific* SFRs, allowing moderate-mass starburst galaxies to compensate for the underabundance of GRBs in non-starbursting high-mass galaxies. This effect probably operates at the low-mass end as well, given the remarkable abundance of GRBs within low-mass, modest total-SFR, extremely high-specific star formation rate (sSFR) galaxies (blue points in Figure 5, or Figure 13 from Perley et al. 2013; see also Christensen et al. 2004, Castro Cerón et al. 2006, 2010, and Svensson et al. 2010), which contribute almost negligibly to overall cosmic star formation except at the highest redshifts.

The traditional hypothesis for the fundamental factor controlling the GRB rate is a metallicity effect, since such an effect is predicted theoretically in many models (e.g., MacFadyen & Woosley 1999; Yoon & Langer 2005) and several observational studies have found a strong trend for the GRB population to favor low-metallicity galaxies (Modjaz et al. 2008; Levesque et al. 2010; Graham & Fruchter 2013). Since mass and metallicity are strongly correlated (e.g., Tremonti et al. 2004), this naturally explains why GRB hosts also seem to favor low-mass galaxies. Evidence for an additional dependence on specific SFR (independent of that on mass) does not necessarily nullify this conclusion, since high-sSFR galaxies may be more metal-poor than low-sSFR galaxies of the same mass (Mannucci et al. 2010; Lara-López et al. 2010)—but the sSFR–metallicity trend is not particularly strong, so it remains to be seen whether the distribution of GRB host properties is truly consistent with a model in which the GRB efficiency depends only on metallicity and no other factors. Direct measurements of the metallicities of the systems studied here (and of other low-metallicity galaxies) should provide some insight: are these galaxies actually significantly less chemically enriched than others of the same mass, or simply more active? In the latter case, this would provide

evidence for alternative scenarios in which another parameter instead of or in addition to metallicity governs the GRB efficiency, such as a variable IMF or binary interactions in dense clusters (van den Heuvel & Portegies Zwart 2013). Some independent evidence for such a scenario is provided by the spatially resolved characteristics of GRB host galaxies: *HST* studies have shown that GRB hosts are more compact than core-collapse supernova hosts and that the GRB itself preferentially occurs in or near the brightest, densest regions of each galaxy (Fruchter et al. 2006; Svensson et al. 2010; Kelly et al. 2014; Michałowski et al. 2014b).

While only four radio-luminous systems have been identified in an unbiased manner so far, the path to identifying more is now well established: it is now possible to efficiently survey large numbers of fields with the VLA, and an extension of this technique could easily identify many new submillimeter/radio-luminous host galaxies in the next few years. Our survey demonstrates the depths necessary to detect an appreciable fraction of GRB hosts out of an unbiased sample ($10 \mu\text{Jy}$ at 3 GHz to identify $\sim 15\%$, independent of the redshift of the host population being targeted as long as it is $z \lesssim 2.5$), a limit achievable in only ~ 2 hr of observation. ALMA observations will be even more sensitive, reaching an order of magnitude fainter and probing GRB hosts beyond $z > 3$, as has already been demonstrated by Wang et al. (2012) for two objects. A complete understanding of the entire GRB host population, from the least to the most luminous galaxies and from the nearby universe out to high redshifts, is clearly within reach.

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