

Spitzer observations of SN 2014J and properties of mid-IR emission in Type Ia Supernovae

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ABSTRACT

SN 2014J in M82 is the closest Type Ia supernova (SN Ia) in decades. The proximity allows for detailed studies of supernova physics and provides insights into the circumstellar and interstellar environment. In this work we analyze *Spitzer* mid-IR data of SN 2014J in the 3.6 and 4.5 μm wavelength range, together with several other nearby and well-studied SNe Ia. We compile the first composite mid-IR light-curve templates from our sample of SNe Ia, spanning the range from before peak brightness well into the nebular phase. Our observations indicate that SNe Ia form a very homogeneous class of objects at these wavelengths. Using the low-reddening supernovae for comparison, we constrain possible thermal emission from circumstellar dust around the highly reddened SN 2014J. We also study SNe 2006X and 2007le, where the presence of matter in the circumstellar environment has been suggested. No significant mid-IR excess is detected, allowing us to place the most constraining upper limits to date on the amount of pre-existing dust in the circumstellar environment of any SN Ia. For SN 2014J, $M_{\text{dust}} \lesssim 10^{-5} M_{\odot}$ within $r_{\text{dust}} \sim 10^{17}$ cm, which is insufficient to account for the observed extinction. Similar limits are obtained for SNe 2006X and 2007le.

Key words: ISM: dust, extinction – supernovae: general - circumstellar matter – supernovae: individual: 2005df, 2006X, 2007af, 2007le, 2007sr, 2009ig, 2012cg, 2011fe, 2014J

1 INTRODUCTION

The use of Type Ia supernovae (SNe Ia) as distance indicators remains essential for the study of the expansion history of the Universe and for exploring the nature of dark energy (see e.g. review by Goobar & Leibundgut 2011). However, a lack of understanding of the progenitor systems and the requirement for empirically derived colour-brightness corrections represent severe challenges for precision cosmology. SN 2014J in the starburst galaxy M82 is the closest SN Ia in decades and offers a unique opportunity to study progenitor and explosion models, as well as the circumstellar (CS) and interstellar (IS) medium along the line-of-sight across an

unprecedented wavelength range, from gamma-rays to radio (Diehl et al. 2014; Pérez-Torres et al. 2014).

Our *Spitzer* observations of SN 2014J, along with observations of other well-studied nearby SNe Ia, allow us to compile the first mid-IR light-curve templates, presenting us with a new wavelength range to confront theoretical models of SNe Ia. For example, the study of the mid-IR light-curve around the near-IR secondary maximum, shown by Kasen (2006) to be a valuable diagnostic of the physical parameters governing SN Ia explosions, could be used to test the validity of the model predictions.

Furthermore, by searching for excess mid-IR emission towards SN 2014J, we set constraints on the amounts of pre-existing CS dust that could account for the non-standard reddening measured by Goobar et al. (2014a);

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Amanullah et al. (2014); Marion et al. (2014); Foley et al. (2014); Brown et al. (2014); Ashall et al. (2014), who otherwise find this SN to be a normal SN Ia, albeit with slightly higher than average expansion velocities, and showing signs of additional sources of luminosity in the first hours after the explosion (Goobar et al. 2014b). We also analyze other SNe Ia with peculiar reddening observed with *Spitzer*, SNe 2006X and 2007le, where CS material has been reported (Patat et al. 2007; Simon et al. 2009).

The outline of this paper is as follows: In Section 2 we present the observations, followed by a description in Section 3 of how the mid-IR light-curve templates are constructed. In Section 4 we add this new wavelength range along with optical and near-IR spectra, to present the spectral energy distribution (SED) in the full range 0.4 – 5.0 μm . The implications for emission from heated CS dust are discussed in Section 5, followed by concluding remarks in Section 6.

2 OBSERVATIONS

2.1 SN 2014J

SN 2014J was observed with *Spitzer* over 6 epochs between January 28 and March 3 and 4 epochs between May 31 and July 8 under the *Spitzer InfraRed Intensive Transients Survey* (SPIRITS) program (PI: M. Kasliwal). SPIRITS is an ongoing infrared survey that systematically searches two hundred nearby galaxies for all types of transients and variables within a volume of 20 Mpc. Data is promptly processed with subtractions relative to archival images in the Spitzer Heritage Archive. The SPIRITS team undertakes a large concomitant ground-based survey in the optical and near-IR to characterize the *Spitzer* findings. For additional details about the survey and first discoveries, see Kasliwal et al. *in prep.*

Aperture photometry was performed at the location of the SN on the aligned *Spitzer* Post-Basic Calibrated Data for the SN and pre-SN images. Throughout the paper we use the zero magnitude fluxes for the IRAC Channels 1 and 2 (CH1 and CH2, with central wavelengths of 3.6 and 4.5 μm , respectively) of $F_{\nu,0}^{\text{CH1}} = 280.9 \text{ Jy}$ and $F_{\nu,0}^{\text{CH2}} = 179.7 \text{ Jy}$.

We also present new optical and near-IR spectra (summarized in Table 2) and photometry (to be published in a future paper) of SN 2014J in Figs. 1 and 2.

2.2 Comparison Supernovae

In order to analyze our mid-IR data on SN 2014J we need to compare with other well-studied SNe Ia. For this study, we include SNe that have multi-epoch *Spitzer* data and good optical/near-IR coverage. SN 2011fe was observed with *Spitzer* starting 145 days after *B*-band maximum. UV data (Mazzali et al. 2014) and optical to near-IR observations (Pereira et al. 2013; Matheson et al. 2012) together with high-resolution spectroscopy (Patat et al. 2013), show that SN 2011fe suffered little to no extinction, $E(B - V) = 0.026 \pm 0.036 \text{ mag}$, making it useful as a template of a pristine, normal SN Ia. We adopt a distance modulus to SN 2011fe of $\mu = 28.93 \pm 0.16 \text{ mag}$ ($D = 6.1 \pm 0.45 \text{ Mpc}$) based on near-IR light curves from Matheson et al. (2012) which is in good agreement with the Cepheid distance

in Mager et al. (2013). SN 2012cg was observed with *Spitzer* (PI: A. Goobar) starting 58 days after *B*-band maximum. The SN shows signs of modest host galaxy reddening, with a colour excess of $E(B - V) \approx 0.2 \text{ mag}$, derived from both optical photometry and high-resolution spectroscopy (Silverman et al. 2012; Munari et al. 2013, Amanullah et al. in prep.). Using optical light curves, Munari et al. (2013) put SN 2012cg at a distance $\mu = 30.95 \text{ mag}$, close to the Tully-Fisher estimate in Cortés et al. (2008). By accounting for the reddening and scaling the NIR photometry in Amanullah et al. (in prep.) to match SN 2011fe, we adopt a distance modulus of $\mu = 30.70 \pm 0.16 \text{ mag}$ ($D = 13.8 \pm 1.0 \text{ Mpc}$).

To compare our limits on CS dust for SN 2014J we also include archival data of the reddened SNe 2006X and 2007le. SN 2006X in M 100 was observed by *Spitzer* (PIs: P. Meikle and R. Kotak) starting 136 days after *B*-band maximum. Similar to SN 2014J, 2006X showed signs of non-standard reddening, $E(B - V) \sim 1.4 \text{ mag}$ with $R_V \sim 1.5$ (Wang et al. 2008b; Folatelli et al. 2010). SN 2007le suffered less extinction than SN 2006X, $E(B - V) \sim 0.39 \text{ mag}$, but also had a low $R_V \sim 1.5$ (Burns et al. 2014). The detection of time varying Na I D absorption for both these SNe has been interpreted as being due to CSM at distances $\sim 10^{17} \text{ cm}$ from the SN (Patat et al. 2007; Simon et al. 2009). Since the SNe are reddened, it has been speculated that dust in the CS environment could play an important role. We also include data of SN 2005df and SN 2009ig previously presented in McClelland et al. (2013) and Gerardy et al. (2007), adding two epochs for SN 2009ig at -3 and +36 days from peak brightness, serendipitously observed with *Spitzer* (PI: K. Sheth). SN 2007af and SN 2007sr were observed with *Spitzer* (PI: R. Kotak) and have well measured optical/near-IR light-curves and precise Cepheid distance estimates (Riess et al. 2011).

3 MID-IR LIGHT CURVES

SN 2014J is the best object to date to build mid-IR light-curve templates, capturing the full range from before peak brightness to the nebular phase. To fill in the gaps, caused by limited visibility windows and scheduling constraints, we make use of archival data of the SNe Ia described in Sect. 2.2. The composite light-curves shown in Fig. 1 are compiled by shifting each SN by their estimated distance modulus and correcting for host galaxy and Milky Way extinction (see Table 1). We do not find evidence for variations in the mid-IR light-curve shapes corresponding to the different optical decline rates in our sample, $\Delta m_{15}^B = 0.9$ to 1.3, suggesting that SNe Ia are a very homogeneous class of objects at longer wavelengths. A larger sample of SNe Ia with multi-epoch mid-IR coverage is needed to make a more quantitative study. Three different decline time scales can be recognized in the CH1 and CH2 light-curves. Although our first detections are before the optical and near-IR maximum brightness, we can not fully measure the mid-IR light-curve shapes at these epochs, i.e. fitting the time of maximum in CH1 and CH2 is impossible. However, fitting the early epochs (-5 d to 15 d from *B*-band maximum) gives linear decline rates of 0.081 and 0.135 mag d^{-1} in CH1 and CH2 respectively, similar to the decline rates in optical bands.

A break in the mid-IR light-curves occurs ~ 15 days

after B -band maximum. This roughly coincides with the onset of the secondary maximum in the near-IR bands, although data from +15 to +30 days show no signs of a secondary maximum in the mid-IR. From 15 days after B -band maximum and onwards, the decline rate changes to 1.67 and 1.93 mag/100d in CH1 and CH2, respectively. After ~ 150 days past B -band maximum the linear decline rates of SN 2011fe are 1.48 mag/100d and 0.78 mag/100d in CH1 and CH2, respectively. The decline rate in CH1 is similar to what is observed at optical wavelengths, while the decline in CH2 is slower, which can also be seen in the CH1–CH2 color panel in Fig. 1 as a change towards redder colors. McClelland et al. (2013) analyze the mid-IR late-time (> 200 days) decline rates for four SNe (SN 2011fe, SN 2009ig, SN 2008Q and SN 2005df) and argue that the different decline rates are a result of doubly ionized elements dominating the bluer CH1 band ($3.6 \mu\text{m}$) while singly ionized iron-peak species are controlling the redder CH2 band ($4.5 \mu\text{m}$). They also suggest that the interpolated color at +230 days correlates with the light-curve decline rate at maximum brightness, i.e. with Δm_{15}^B .

4 SPECTRAL ENERGY DISTRIBUTION

Lacking spectra in the wavelength range probed by our *Spitzer* observations, we examine how the shape of the SED, as measured through optical/near-IR spectroscopy together with the *Spitzer* broad-band observations, matches model spectra from numerical simulations. In Fig. 2 we show observed optical and near-IR spectra (red lines) together with optical, near-IR and mid-IR broad-band photometry (red circles) of SN 2014J at three epochs (5, 31 and 126 days after B -band maximum).

To this end, we have calculated synthetic spectra for the hydrodynamic explosion model W7 (Nomoto et al. 1984). W7 is known to reproduce the observed characteristics of normal SNe Ia at optical (e.g. Branch et al. 1985) and near-IR wavelengths (e.g. Gall et al. 2012). To obtain synthetic spectra at 5 and 31 days past B -band maximum we performed radiative transfer simulations using the ARTIS code (Kromer & Sim 2009) and the atomic data set described in Gall et al. (2012). For the latest epoch (+126 days past B -band maximum) we used a nebular code (described in Jerkstrand et al. 2011, 2014). The flux in the model spectra between B -band maximum and +10 days stems from a mixture of Co II, Co III and Fe II features and a tiny contribution from intermediate-mass elements. Starting at 10 days after maximum, the flux in the 2.8 to 3.5 micron region is dominated by singly-ionized iron-group elements. The lines in the W7 model possibly contributing to the flux in CH1 can be attributed to Ni II (at 2.85, 2.95, 3.11, 3.29 and $3.54 \mu\text{m}$) and Fe II ($3.08 \mu\text{m}$).

In the late-time model spectrum (+126 days after maximum), the fluxes in the optical and near-IR agree with the observations while the flux levels in the mid-IR are underpredicted. The dominant lines in CH1 and CH2 are [Fe II], with little contribution from [Co II]. In CH1 there is also [Fe III] (at 2.90 and $3.04 \mu\text{m}$) and [Co III] (at $3.48 \mu\text{m}$) emission.

Our observations, shown in Figs. 1 and 2, bridge the gap between the late-time near-IR spectra in Friesen et al. (2014) and mid-IR spectra of SN 2014J in the 8–13 μm wave-

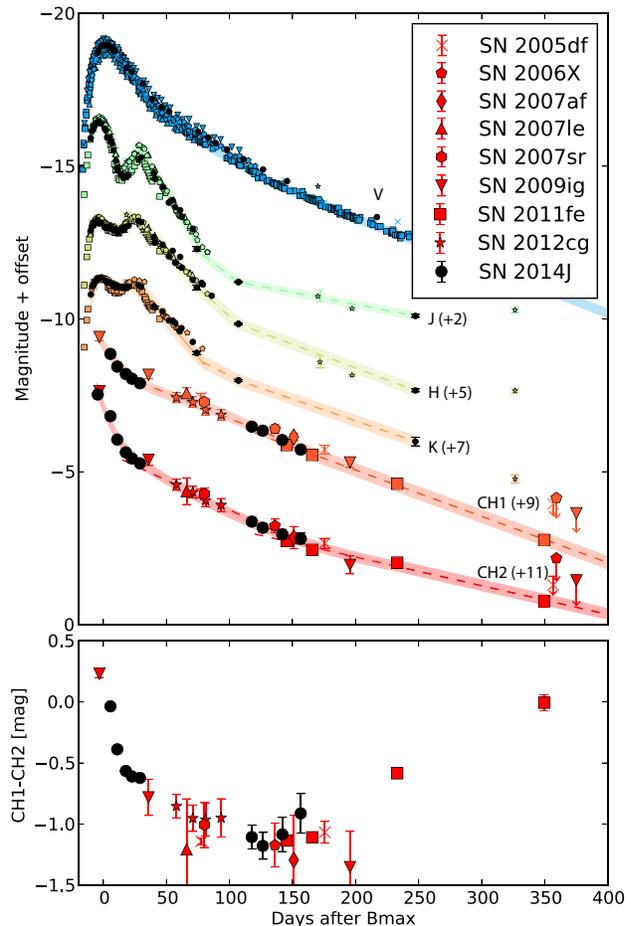


Figure 1. Upper panel: Absolute magnitude V -band, near-IR (J , H and K -band) and mid-IR light-curves of the Type Ia SNe used in this study. The magnitudes have been de-reddened using the best fit extinction values and shifted by the distance moduli (typically known to ~ 0.2 mag accuracy) listed in Tab. 1. The lower panel shows the *Spitzer* CH1–CH2 colour evolution.

length range in Telesco et al. (2014). Friesen et al. (2014) present near-IR spectra of SN 2014J at 70 days past maximum and find that [Ni II] fits the emission feature near $1.98 \mu\text{m}$, suggesting that a substantial mass of ^{58}Ni exists near the center of the ejecta, arising from nuclear burning at high density. A tentative identification of Mn II at $1.15 \mu\text{m}$ may support this conclusion as well. Telesco et al. (2014) compare their observed mid-IR spectra to a delayed detonation model with $\sim 0.6 M_{\odot}$ of ^{56}Ni and claim that the model is consistent with observations. Recent multi-dimensional hydrodynamical simulations of Chandrasekhar-mass explosions, however, struggle in producing a concentration of stable iron-group elements near the center of the ejecta (e.g. Seitenzahl et al. 2013).

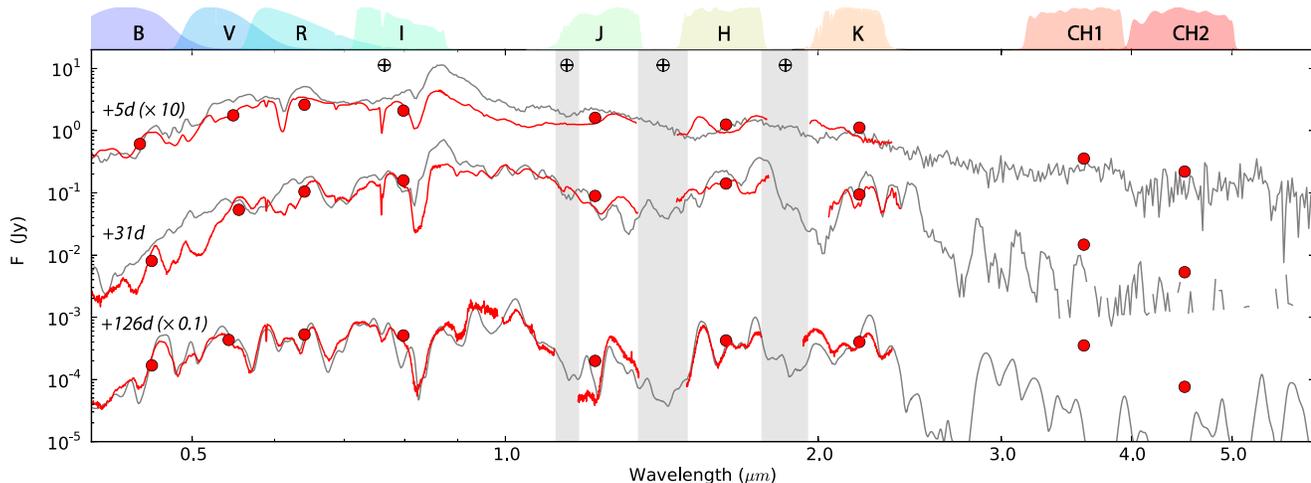


Figure 2. Observed SEDs at +5, +31 and +126 days past *B*-band maximum from optical/NIR spectroscopy (red lines) and broad band *BVR IJHK* and *Spitzer* CH1 and CH2 photometry (red circles). For comparison, synthetic spectra of the W7 model (gray lines) at similar epochs are shown. Vertical, gray bars indicate regions of low atmospheric transmission.

5 CONSTRAINTS ON EMISSION FROM CS DUST

The existence of CS material around individual nearby SNe Ia has been suggested by studies of sodium absorption lines (e.g. SNe 1999cl, 2006X, 2007le, and PTF11kx Patat et al. 2007; Blondin et al. 2009; Simon et al. 2009; Dilday et al. 2012). High-resolution spectra reveal the presence of time-variable and blueshifted Na I D features, possibly originating from CSM within the progenitor system. Studies of large samples of SNe Ia (Sternberg et al. 2011) find that half of all SNe Ia with detectable Na I D absorption at the host-galaxy redshift have Na I D line profiles with significant blueshifted absorption relative to the strongest absorption component. This indicates that the absorption occurs in the vicinity of the progenitor systems rather than in the ISM. For SN 2014J, high-resolution spectra show that the line of sight is very rich in absorbing material, but reveal no signs of time-variable Na I D absorption (Goobar et al. 2014a; Foley et al. 2014; Welty et al. 2014; Ritchey et al. 2014). The wavelength dependent extinction towards SN 2014J has been measured with great accuracy using a very wide wavelength range, 0.2 – 2.2 μm , through a combination of *Hubble Space Telescope* and ground-based observations. Amanullah et al. (2014) found that the reddening can be described with either a MW-type extinction law with a low value of the total-to-selective extinction, $R_V = 1.4 \pm 0.1$, i.e., corresponding to non-standard dust grains in the ISM, or by invoking the effective extinction law from Goobar (2008). The latter arises from multiple scattering of photons on “normal” dust grains, but surrounding the supernova, i.e., in the CS medium, as also discussed by Wang (2005).

Non-standard reddening has been noted in studies of individual and large samples of SNe Ia. For example, the extinction of SN 2006X was studied in Folatelli et al. (2010), showing that the reddening is incompatible with the average extinction law of the Milky Way. Their findings augmented the large body of evidence indicating that the red-

dening of many SNe Ia show a steeper wavelength dependence ($R_V < 3.1$) than that which is typically observed for stars in our Galaxy. Previously, Nobili & Goobar (2008) derived $R_V = 1.75 \pm 0.27$ from a statistical study of 80 low-redshift SNe Ia. Similarly, when the colour-brightness relation is fitted jointly with cosmological parameters in the SN Ia Hubble diagram, using a wide range of SN Ia redshifts, low values of R_V are obtained (see e.g. Betoule et al. 2014, for a recent compilation).

Amanullah & Goobar (2011) simulated the impact of thin CS dust shells located at radii $r_d \sim 10^{16} - 10^{19}$ cm ($\sim 0.003 - 3$ pc) from the SN and found that this scenario would also perturb the optical light-curve shapes and introduce a time dependent color excess, $\Delta E(B - V) \sim 0.05 - 0.1$ mag. Foley et al. (2014) claim to have detected a time variable color excess for SN 2014J, which led them to conclude that dimming by CS dust accounts for about half of the extinction. However, this interpretation has been challenged by Brown et al. (2014). By exploring the mid-IR wavelength range, we have a unique way to test if dust in the CSM plays a significant role in explaining the non-standard reddening towards SN 2014J and other highly reddened normal SNe Ia.

5.1 Dust models

If pre-existing CS dust is the source of non-standard reddening, it will be radiatively heated by absorption of UV/optical photons from the SN or collisionally heated by the SN shock. Thermal emission at IR wavelengths could therefore be the “smoking gun” for detecting or ruling out the presence of CS dust. To model the emission from pre-existing CS dust we consider the idealized case (described in Hildebrand 1983; Dwek 1985; Fox et al. 2010) of an optically thin (at mid-IR wavelengths) dust cloud of mass M_d with dust particles of radius a , emitting thermally at a single equilibrium temperature T_d . The expected flux at a distance D is,

$$F_\nu = M_d \frac{\kappa_\nu(a) B_\nu(T_d)}{D^2}, \quad (1)$$

where $B_\nu(T_d)$ is the Planck blackbody function and the dust mass emissivity coefficient, $\kappa_\nu(a)$, is

$$\kappa_\nu(a) = \left(\frac{3}{4\pi\rho a^3} \right) \pi a^2 Q_\nu(a) = \frac{3Q_\nu(a)}{4a\rho}. \quad (2)$$

$Q_\nu(a)$ is the absorption efficiency and the dust bulk (volume) density, $\rho \approx 2 - 3 \text{ g cm}^{-3}$ depending on grain composition. Since we do not know the nature of the SN Ia progenitor systems and their potential dust production mechanisms, we will consider separate scenarios of either silicate or graphite grains of radius $a = 0.1 \mu\text{m}$ and a mixture of silicate and graphitic grains of different sizes (MW3.1) that reproduce the standard $R_V = 3.1$ Milky Way dust properties (described in Draine & Lee 1984; Laor & Draine 1993; Weingartner & Draine 2001).

In what follows, we treat the low-reddening SNe 2011fe and 2012cg as mid-IR templates of a SN Ia, to put limits on any possible excess emission from CS dust around the highly-reddened SNe 2014J, 2006X and 2007le. Taking both instrumental noise and distance modulus uncertainties into account, the differences between the reddened and un-reddened SNe are not statistically significant. Hence, we compute 3σ upper limits on the dust temperature and mass for SN 2014J (red contours in Fig. 3) as well as for other SNe (red contours in Fig. 4). We complement this with a similar analysis using K -band data of SNe 2014J and 2006X (green contours in Fig. 4).

5.2 Expected emission from heated dust in the thin shell approximation

Although the geometric distribution of CS dust could be complex, we adopt the simple thin shell approximation to provide an estimate of the expected mid-IR emission from heated dust, to be compared with the models for reddening proposed in Goobar (2008); Amanullah & Goobar (2011). This allows us to estimate the expected temperature as a function of shell radius, shown as the right-hand side vertical axis of Figs. 3 and 4. We estimate the minimal dust shell radius, $r_d = ct/2$, that could give rise to a detectable IR echo at each observed epoch, t . Dust at this radius will be heated to

$$T_{d,\text{exp}} \sim 4.0 \left(\frac{L}{a} \right)^{\frac{1}{6}} r_d^{-\frac{1}{3}}, \quad (3)$$

where we assume that the peak SN bolometric luminosity of $\sim 3 \times 10^9 L_\odot$ is heating a spherical dust shell with grain sizes of $a = 0.1 \mu\text{m}$ (Kruegel 2003). The upper bound on the dust temperature is set by the evaporation temperature of the dust grains ($T \lesssim 2000 \text{ K}$), corresponding to a minimal dust survival radius $r_{\text{evap}} \sim 10^{16} \text{ cm}$.

Furthermore, in order for a thin dust shell at r_d to have significant opacity in the optical V -band, $\tau(V) \sim 1$, the required dust mass can be estimated from the absorption cross-sections, $\sigma_{\text{abs}}(\lambda = 0.55 \mu\text{m})/m_{\text{dust}} = \kappa_V$, as

$$M_{d,\text{exp}} \sim 4\pi r_d^2 \frac{\tau(V)}{\kappa(V)}, \quad (4)$$

where $\kappa(V) \sim 5 \cdot 10^4 \text{ cm}^2 \text{ g}^{-1}$ for graphitic grains, $\kappa(V) \sim 2 \cdot 10^3 \text{ cm}^2 \text{ g}^{-1}$ for silicate grains of size $a = 0.1 \mu\text{m}$ and $\kappa(V) \sim 2 \cdot 10^3 \text{ cm}^2 \text{ g}^{-1}$ for the MW3.1 mixture. Thus,

for a thin spherical dust shell at r_d the total dust mass, $M_{d,\text{exp}} \sim 10^{-4} - 10^{-5} (r_d/10^{16} \text{ cm})^2 M_\odot$, depending on dust grain composition, is needed to explain the observed reddening if it mainly arises in the CS environment. This corresponds to the blue lines in Figs. 3 and 4. The black horizontal lines in Fig. 3 indicate the expected dust temperature (and corresponding shell radius) at two of the epochs (29 and 157 days from B -band maximum) from where the limits are derived for SN 2014J. Similarly, the symbols in Fig. 4 indicate the limits derived at different epochs for an extended sample of SNe observed with *Spitzer*.

Using K -band data for SNe 2014J and 2006X, to explore possible emission from hot CS dust close to the exploding star ($T_d \gtrsim 1200 \text{ K}$ for $r_d < 5 \cdot 10^{16} \text{ cm}$), we are able to rule out $M_d > 10^{-5} M_\odot$. By adding the *Spitzer* data we can significantly reduce the allowed parameter space, excluding $M_d > 10^{-6} M_\odot$ ($T_d > 700 \text{ K}$, for $r_d < 2 \cdot 10^{17} \text{ cm}$).

Our combined limits for SNe 2014J correspond to a very severe upper limit on the possible amount of pre-existing dust surrounding the SN progenitor, $M_d \lesssim 10^{-5} M_\odot$ within $r_d \sim 2 \cdot 10^{17} \text{ cm}$, depending on the assumed grain size and composition. However, regardless of the specific dust grain composition, the derived upper limits on the CS dust mass are significantly lower than what would be required to explain the observed reddening of SN 2014J, in apparent contradiction to the claims in Foley et al. (2014). Similar limits are obtained for the reddened SNe 2006X and 2007le (see Fig. 4). For reference, we also show detections of dust around a subset of peculiar SNe Ia interacting strongly with a dense CSM (SNe Ia-CSM) in Fig. 4. *Spitzer* observations of SNe Ia-CSM 2002ic, 2005gj and 2008J (Fox et al. 2011; Fox & Filippenko 2013; Taddia et al. 2012) show evidence for late-time (> 500 days after maximum brightness) mid-IR emission from warm ($T_d \sim 500\text{--}800 \text{ K}$) dust. Assuming a simple dust population of a single size ($a = 0.1 \mu\text{m}$) graphitic grains yield dust masses of $0.006 - 0.022 M_\odot$. The dust parameters are most consistent with a pre-existing dust shell that lies beyond the forward-shock radius, most likely radiatively heated by optical and X-ray emission continuously generated by late-time CSM interaction.

6 SUMMARY AND CONCLUSIONS

We have analyzed the mid-IR light curves of SN 2014J and several other SNe Ia observed with *Spitzer*, spanning the range from before peak brightness well into the nebular phase. We have characterized, for the first time, the SN Ia light-curve evolution at 3.6 and $4.5 \mu\text{m}$. Our observations indicate that SNe Ia form a very homogeneous class of objects at these wavelengths, possibly without any light-curve shape variations. In particular, the mid-IR light curves do not show any evidence for a secondary maximum, as opposed to the case in the near-IR. The latter was investigated by Kasen (2006) to explore the physics of the exploding system. Extending these studies to the now probed longer wavelengths should provide critical tests for SN Ia models.

The *Spitzer* observations provide a completely new way to test models for the circumstellar environment of SNe Ia and may help understanding the non-standard reddening measured both for individual SNe Ia and in the large SN Ia samples used to derive cosmological parameters.

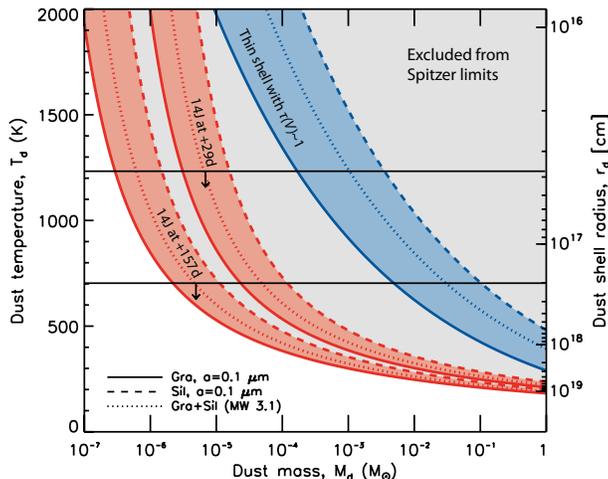


Figure 3. 3σ upper limits on dust emission around SN 2014J at 29 and 157 days after maximum brightness. The solid, dashed and dotted contours indicate limits using the graphitic, silicate and MW3.1 dust models. Assuming the thin shell approximation, we can estimate the expected dust temperatures and shell radii probed at these epochs (black lines) and the expected dust mass from Eq. 4 (blue lines).

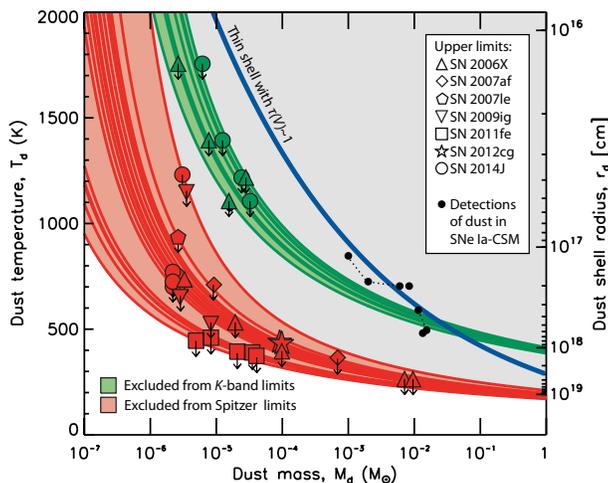


Figure 4. 3σ upper limits on CS dust emission around the reddened SNe 2014J, 2006X and 2007le at different epochs from K -band data (green contours and symbols) and *Spitzer* mid-IR data (red contours and symbols), assuming graphitic dust grains of radius $0.1 \mu\text{m}$. Also shown are the *Spitzer* limits for SNe 2007af and 2009ig along with the late time non-detections of SNe 2011fe and 2012cg. The blue solid line indicates the expected dust mass from Eq. 4. Black dots are detections of warm dust around Ia-CSM SNe.

By comparing the measured mid-IR fluxes at different epochs for the reddened SNe 2014J, 2006X and 2007le to unreddened SNe, we can place strong constraints on the emission from heated dust within $\sim 10^{18}$ cm from the exploding stars. This allows us to exclude the possibility that the bulk of the observed extinction towards these highly-reddened SNe Ia is due to CS dust. Foley et al. (2014) claim that half of the extinction ($A_V \sim 1$ mag) towards SN 2014J can be attributed to CS dust, while the other half is due to interstellar dust in M82. However, our limits on dust

emission imply that at most $\tau(V) \approx A_V \lesssim 0.1$ mag of extinction can be accounted for by CS dust. We emphasize that the derived limits are relatively insensitive to the assumed dust models (as illustrated in Fig. 3). Thus, our findings reaffirm the conclusions from polarization studies of SNe 2014J and 2006X, which indicate that the dust in the line of sight towards these objects is most likely of interstellar nature (Kawabata 2014; Patat et al. 2014). Furthermore, the lack of heated material in the CS environment of SN 2014J is compatible with the non-detection in X-rays and radio (Margutti et al. 2014; Pérez-Torres et al. 2014). The only comparable previous study of searches for emission from heated circumstellar dust to date was carried out by Johansson et al. (2013), presenting far-IR non-detections of both SNe 2011fe and 2012cg using the *Herschel* PACS $70 \mu\text{m}$ instrument. These non-detections exclude CS dust masses $M_d \gtrsim 7 \times 10^{-3} M_\odot$ for dust temperatures $T_d \sim 500$ K at a 3σ -level for SN 2011fe, while the upper limits are one order of magnitude weaker for SN 2012cg, excluding dust masses $M_d \gtrsim 10^{-1} M_\odot$. Thus, our *Spitzer* study is more than two orders of magnitude more sensitive than previous attempts.

The mid-IR non-detections for SNe 2006X, 2007le, 2011fe and 2012cg at late epochs (> 600 days after max) can constrain the possible presence of dust at very large radii ($> 10^{18}$ cm). E.g., Soker (2014) proposed that the sodium responsible for the time variable absorption seen in SNe 2006X and 2007le is released from dust grains. In this model, the absorbing shell/ring, $\sim 1 - 10 M_\odot$, resides at distances $10^{17} - 10^{19}$ cm away from the SN, most likely concentrated in an equatorial plane. This range, probed by our observations shown in Fig. 4, is larger than what a regular AGB wind can supply, but is compatible with a planetary nebula shell or a planetary nebula shell that entrained some interstellar medium (ISM). Light echoes have been detected for reddened SNe 2014J and 2006X (Crotts 2014; Wang et al. 2008a), and reportedly from the low-reddening SNe 2007af and 2009ig (Drozdov et al. 2014; Garnavich et al. 2013). Our observations can provide limits on the allowed distance at which the reflecting dust responsible for these echoes can exist. Furthermore, Nozawa et al. (2011) investigate the possibility of forming new dust grains in SN Ia ejecta. They compare their dust-formation models with *Spitzer* photometry for SN 2005df (Gerardy et al. 2007) and derive an upper limit of $0.1 M_\odot$ of newly formed dust after $\sim 100 - 300$ days. Our limits are stricter by at least an order of magnitude.

To summarize, this work significantly expands on previous efforts to study the wide wavelength range of the SED of SNe Ia and provides the first statistical sample of SNe Ia in the mid-IR and a detailed study of the highly reddened object SN 2014J and its CS environment. The non-detection of thermal emission from heated dust in the CS environment of SN 2014J, as well as for SNe 2006X and 2007le where the detection of circumstellar matter at 10^{17} cm has been claimed, argues against the proposed explanation for the non-standard reddening of SNe Ia invoking multiple-scattering on CS dust (Wang 2005; Goobar 2008). Thus, the non-standard reddening may be unrelated to the SN site and originate from the host galaxy ISM being different than what has been derived for the Milky Way. This could have a serious impact for our understanding of the properties of dust grains in distant galaxies, with profound implications for essentially all areas of extragalactic astronomy.

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Table 1. Type Ia supernovae observed in the mid-IR

Target	$t_{B,\max}$ MJD	Host galaxy	Distance modulus (mag)	Distance (Mpc)	$E(B - V)$ (mag)	R_V	Δm_{15}^B (mag)	Reference
2005df	53598.4	NGC 1559	31.26 (0.14)	15.7 (2.3)	0.03	3.1	1.20	1, 2, 3
2006X	53785.6	NGC 4321	30.91 (0.14)	15.2 (1.0)	1.24	1.6	1.29	4, 5
2007af	54173.8	NGC 5584	31.72 (0.07)	22.0 (0.7)	0.04	3.0	1.20	6, 7, 8
2007le	54398.9	NGC 7721	32.24 (0.16)	28.1 (2.1)	0.39	1.5	1.10	9
2007sr	54448.3	NGC 4038	31.66 (0.08)	21.5 (0.8)	0.13	1.6	0.97	6, 10
2009ig	55080.9	NGC 1015	32.82 (0.09)	36.6 (1.5)	<0.05	-	0.89	2, 11, 12
2011fe	55815.8	NGC 5457	28.93 (0.16)	6.1 (0.5)	0.03	3.1	1.11	2, 13, 14
2012cg	56082.0	NGC 4424	30.70 (0.16)	13.8 (1.0)	0.18	2.7	0.89	14, 15, 16
2014J	56689.2	M 82	27.60 (0.16)	3.3 (0.15)	1.37	1.4	0.95	17, 18, 19

(1): Gerardy et al. (2007), (2): McClelland et al. (2013), (3): Milne et al. (2010), (4): Wang et al. (2008b), (5): Folatelli et al. (2010) (6): Riess et al. (2011), (7): Simon et al. (2007), (8): Hicken et al. (2009), (9): Burns et al. (2014), (10): Schweizer et al. (2008), (11): Foley et al. (2012), (12): Marion et al. (2013), (13): Matheson et al. (2012), (14): Munari et al. (2013), (15): Silverman et al. (2012), (16): Amanullah et al., *in prep.*, (17): Amanullah et al. (2014), (18): Foley et al. (2014), (19): Marion et al. (2014)

Table 2. Log of spectroscopic observations of SN 2014J

UT date	MJD	Days from B_{\max}	Observatory / Instrument	Wavelength range (μm)	Exp. time (s)	Airmass
2014-02-06 ^a	56694.9	+5.1	NOT/ALFOSC	0.35 – 0.90	180	1.61
2014-02-06 ^a	56695.0	+5.2	Mt Abu	0.85 – 2.35	1200	1.59
2014-03-03	56718.9	+29.1	Mt Abu	0.85 – 2.35	1080	1.48
2014-03-05	56721.1	+31.3	NOT/ALFOSC	0.35 – 0.90	180	1.35
2014-06-07	56815.3	+125.5	Keck/MOSFIRE	1.15 – 2.35	480	2.19
2014-06-09	56817.2	+127.4	ARC/DIS	0.35 – 0.95	120	1.73

^a Published in Marion et al. (2014).

Table 3. Spitzer observations and photometry of Type Ia supernovae in the mid-IR.

SN	MJD	Epoch (days)	F_{ν}^{CH1} (μJy)	F_{ν}^{CH2} (μJy)	CH1 (mag)	CH2 (mag)	Sub? (y/n)
2005df	53676.0	77.6	327.0 (6.7)	73.2 (5.7)	14.835 (0.022)	15.976 (0.082)	y
2005df	53774.0	175.6	68.3 (5.6)	16.4 (5.3)	16.535 (0.085)	17.602 (0.307)	y
2005df	53955.1	356.7	13.3 (5.2)	4.6 (5.3)	>18.136	>17.642	y
2006X	53922.2	136.1	157.9 (8.0)	35.3 (6.0)	15.626 (0.054)	16.768 (0.171)	y
2006X	54145.4	359.3	19.7 (8.0)	13.3 (6.0)	>17.750	>17.572	y
2006X	54662.7	876.6	-0.7 (7.6)	1.8 (5.9)	>17.731	>17.520	y
2006X	56722.8	2936.7	-6.9 (8.3)	-1.3 (6.5)	>17.633	>17.415	y
2006X	56751.9	2965.8	1.6 (8.6)	0.2 (6.3)	>17.586	>17.447	y
2007af	54324.9	151.1	62.2 (4.6)	12.8 (5.0)	16.594 (0.081)	17.865 (0.357)	y
2007af	55280.4	1106.6	2.4 (5.0)	2.2 (5.0)	>18.181	>17.696	y
2007le	54464.7	66.0	152.0 (5.0)	32.0 (5.0)	15.667 (0.059)	16.873 (0.312)	y
2007sr	54528.1	80.1	220.0 (8.0)	55.7 (10.0)	15.265 (0.146)	16.271 (0.146)	y
2007sr	55616.4	1168.4	3.3 (8.0)	19.2 (8.0)	>17.671	>17.186	y
2009ig	55076.9	-3.1	580.7 (10.0)	459.1 (10.0)	14.211 (0.019)	13.982 (0.023)	n
2009ig	55115.9	35.8	186.1 (8.0)	57.9 (8.0)	15.447 (0.046)	16.229 (0.140)	n
2009ig	55275.7	195.6	13.3 (1.4)	2.4 (0.7)	18.316 (0.110)	19.668 (0.270)	n
2009ig	55455.2	375.2	2.9 (1.1)	1.5 (0.8)	>19.876	>19.672	n
2011fe	55960.7	145.4	670.6 (8.5)	150.9 (6.4)	14.055 (0.014)	15.190 (0.045)	y
2011fe	55981.0	165.6	498.2 (8.1)	114.9 (6.3)	14.378 (0.018)	15.486 (0.058)	y
2011fe	56048.4	233.0	209.2 (7.0)	78.1 (5.7)	15.320 (0.036)	15.905 (0.076)	y
2011fe	56165.0	349.7	38.3 (5.7)	24.4 (5.4)	17.163 (0.151)	17.168 (0.217)	y
2011fe	56337.1	521.8	3.3 (5.6)	14.1 (6.1)	>18.058	17.763 (0.390)	y
2011fe	56348.1	532.8	2.5 (5.8)	15.6 (5.0)	>18.020	17.654 (0.302)	y
2011fe	56393.8	578.5	4.4 (5.1)	7.8 (7.7)	>18.160	>17.227	y
2011fe	56452.6	637.3	3.9 (6.0)	-2.6 (4.6)	>17.983	>17.787	y
2011fe	56742.8	927.5	-	2.4 (5.7)	-	>17.554	y
2011fe	56771.8	956.5	-	1.9 (5.8)	-	>17.534	y
2011fe	56902.0	1086.7	3.1 (7.4)	-8.9 (11.2)	>17.755	>16.820	y
2012cg	56139.9	57.9	549.0 (20.3)	160.0 (13.3)	14.272 (0.039)	15.126 (0.087)	y
2012cg	56152.9	70.9	481.0 (19.0)	128.0 (12.3)	14.416 (0.042)	15.368 (0.100)	y
2012cg	56163.1	81.1	380.0 (20.4)	100.0 (11.9)	14.672 (0.057)	15.636 (0.122)	y
2012cg	56175.4	93.4	326.0 (20.6)	87.0 (12.2)	14.838 (0.067)	15.788 (0.142)	y
2012cg	56723.8	641.8	11.9 (17.2)	6.2 (12.6)	>16.837	>16.694	y
2012cg	56751.9	669.9	4.1 (18.4)	2.2 (12.7)	>16.764	>16.682	y
2014J	56685.4	-4.5	-	42273.0 (200.0)	-	9.071 (0.005)	y
2014J	56695.4	5.6	35895.9 (154.9)	22130.1 (97.1)	9.734 (0.005)	9.774 (0.005)	y
2014J	56700.8	11.0	24508.9 (149.2)	10554.5 (95.9)	10.148 (0.007)	10.578 (0.010)	y
2014J	56707.6	17.8	18295.8 (117.4)	7130.0 (73.1)	10.466 (0.007)	11.004 (0.011)	y
2014J	56712.4	22.6	16491.0 (121.1)	5932.9 (75.2)	10.578 (0.008)	11.203 (0.014)	y
2014J	56718.8	29.0	14325.6 (129.4)	5145.3 (80.5)	10.731 (0.010)	11.358 (0.017)	y
2014J	56807.5	117.7	3812.2 (127.3)	876.4 (80.3)	12.168 (0.036)	13.280 (0.095)	y
2014J	56816.2	126.4	3273.6 (122.5)	766.0 (75.8)	12.334 (0.040)	13.426 (0.102)	y
2014J	56831.7	141.9	2338.2 (121.3)	600.4 (79.9)	12.699 (0.055)	13.690 (0.136)	y
2014J	56846.3	156.5	1922.3 (126.5)	448.6 (79.9)	12.912 (0.069)	14.007 (0.178)	y