

Extremes of the jet–accretion power relation of blazars, as explored by *NuSTAR*

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ABSTRACT

Hard X-ray observations are crucial to study the non-thermal jet emission from high-redshift, powerful blazars. We observed two bright $z > 2$ flat-spectrum radio quasars (FSRQs) in hard X-rays to explore the details of their relativistic jets and their possible variability. S5 0014+81 (at $z = 3.366$) and B0222+185 (at $z = 2.690$) have been observed twice by the *Nuclear Spectroscopic Telescope Array (NuSTAR)* simultaneously with *Swift*/X-ray Telescope, showing different variability behaviours. We found that *NuSTAR* is instrumental to explore the variability of powerful high-redshift blazars, even when no γ -ray emission is detected. The two sources have proven to have respectively the most luminous accretion disc and the most powerful jet among known blazars. Thanks to these properties, they are located at the extreme end of the jet–accretion disc relation previously found for γ -ray detected blazars, to which they are consistent.

Key words: galaxies: active – quasars: general – quasars: individual: B0222+185 – quasars: individual: S5 0014+813 – X-rays: general.

1 INTRODUCTION

Blazars are active galactic nuclei (AGN) with their broad-band emission dominated by the relativistic jet, oriented close to our line of sight. The two humps that characterize their spectral energy distribution (SED) are the signature of this relativistically beamed emission. They are attributed to synchrotron (at low frequencies) and inverse Compton (at high frequencies) processes, and in the radio/sub-millimetre and X-/gamma-ray, respectively. The electron population involved in the inverse Compton emission is thought to interact either with the synchrotron photons involved in the low-

frequency emission, or with photons coming from structures external to the relativistic jet (synchrotron self-Compton, SSC, or external Compton, EC, emissions, respectively). The latter is likely the primary process in sources that present a pronounced dominance of the higher frequencies hump over the synchrotron one. This usually happens in the most powerful blazars, i.e. the flat-spectrum radio quasars (FSRQs). These sources are thought to have more sources of seed photons for an EC process (i.e. accretion disc, broad-line region, torus), compared to the BL Lacertae objects (BL Lacs) that have weak or absent broad lines and no accretion or torus emission (see e.g. Chiaberge, Capetti & Celotti 1999; Ghisellini et al. 2011; Sbarrato et al. 2012).

The most immediate signature of the blazar nature of an AGN is its emission in the γ -rays. The high-energy hump in very powerful

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blazar SED, in particular, usually peaks in the MeV–GeV range, and therefore it can be easily observed with γ -ray telescopes, such as the Large Area Telescope (LAT) onboard the *Fermi Gamma-Ray Space Telescope* (Atwood et al. 2009). The *Fermi/LAT* team built an all-sky γ -ray catalogue, providing a clear classification of all the sources included in the survey, through multifrequency studies. This provides a complete, all-sky blazar catalogue (Ackermann et al. 2015). Nevertheless, at higher redshifts, *Fermi/LAT* is less efficient in detecting blazars, even those with a very large bolometric luminosity. This is because the most powerful blazars have their high-energy peak at \sim MeV energies or below, and this peak is seen redshifted. This is the reason why the fraction of high-redshift blazars (i.e. at $z > 2$) detected in the hard X-rays by the Burst Alert Telescope (BAT) onboard the *Swift* satellite is much larger than for *Fermi/LAT* (see Ajello et al. 2009; Ghisellini et al. 2010).

Indeed, blazars observed so far show a trend: the humps in the SEDs of the more powerful ones peak at lower frequencies as compared to less powerful blazars. This trend is known as the ‘blazar sequence’ (Fossati et al. 1998). Although the original concept of ‘blazar sequence’ finds confirmation through *Fermi* blazar data (i.e. Ackermann et al. 2015; Ajello et al. 2015), there is some dispute about its reality (see e.g. the reviews by Padovani 2007 and Ghisellini & Tavecchio 2008). For instance, Giommi et al. (2012) proposed what they called a ‘simplified blazar scenario’, in which they postulate that the shape of the SED of blazars is uncorrelated with their luminosity. Then they assume a given probability for the different blazar shape: there are more blazar with low-frequency peaks than high-frequency ones. Taking into account the observation constraints and the limiting fluxes of the current blazar surveys, they can reproduce what is observed. This should be taken as a test that both the ‘blazar sequence’ and the ‘simplified scenario’ pass, not as a proof that the blazar sequence is wrong. Both frameworks can describe the considered existing data. On the other hand, the blazar sequence found an easy physical explanation in terms of radiative cooling (Ghisellini et al. 1998), while the simplified scenario is based on the assumed SED distribution, that has no physical explanation (yet). Whether the blazar sequence is intrinsic or only a selection effect, the most powerful and distant blazars are hardly detected by *Fermi/LAT*, their inverse Compton emission peak being shifted towards the (observed) MeV band (Ghisellini et al. 2010). Hard X-ray instruments, instead, like *Swift/BAT* and now the *Nuclear Spectroscopic Telescope Array* (*NuSTAR*; Harrison et al. 2013), are the most suitable instruments now available to investigate jet emission in the most powerful blazars at $z \sim 2$ –3.

In this paper, we report on observations of S5 0014+81 ($00^{\circ}17'08''.5 + 81^{\text{h}}35^{\text{m}}08^{\text{s}}$, $z = 3.366$) and B0222+185 ($02^{\circ}25'04''.7 + 18^{\text{h}}46^{\text{m}}49^{\text{s}}$, $z = 2.690$) by *NuSTAR*. These two blazars have been previously detected in the 3 yr all-sky survey of *Swift/BAT* (Ajello et al. 2009, see also Ajello et al. 2012 and Baumgartner et al. 2013), and are amongst the most powerful blazars ever observed. As with other powerful and high-redshift FSRQs, their optical flux shows contributions due to thermal emission from the accretion disc, particularly prominent in S5 0014+813, whose luminosity reaches $\sim 10^{48}$ erg s^{-1} (Ghisellini et al. 2010). For both sources, the *Swift/BAT* spectrum together with the *Fermi/LAT* upper limit already constrained the location of the high energy peak, but with a relatively large uncertainty given the poor spectral slope determination of *Swift/BAT*. This motivated the *NuSTAR* observations.

In this work, we adopt a flat cosmology with $H_0 = 68 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_{\text{M}} = 0.3$, as found by Planck Collaboration XIII (2015).

2 OBSERVATIONS AND DATA ANALYSIS

We performed simultaneous *NuSTAR* and *Swift* observations in two observing periods for each source. This section describes the data analysis performed on these new X-ray and optical-UV data. Along with the new data sets, we consider archival data for the overall SED modelling (see Sections 3 and 4). Specifically, S5 0014+813 radio and IR data are from Ghisellini et al. (2009), integrated with new IR photometry from the *Wide-field Infrared Explorer* (*WISE*;¹ Wright et al. 2010). In the case of B0222+185, archival data were all retrieved through the ASI Science Data Center (ASDC²). Both sources have not been detected by *Fermi/LAT*, but we obtained some information from this lack of detection with an ‘upper limit’ on their γ -ray fluxes. The sensitivity limit at 5σ estimated on 5 yr of observations gives a good constrain on the high-energy emission of S5 0014+813 and B0222+185.

2.1 NuSTAR observations

The *NuSTAR* satellite observed S5 0014+81 on 2014 December 21 (obsID 60001098002) and on 2015 January 23 (obsID 60001098004) for total net exposure times of 31.0 ks and 36.4 ks, respectively. B0222+185 was observed by *NuSTAR* on 2014 December 24 (obsID 60001101002) and on 2015 January 18 (obsID 60001101004). The total net exposure times were 32.0 ks and 37.4 ks, respectively.

The Focal Plane Module A (FPMA) and Focal Plane Module B (FPMB) data sets were first processed with the *NUSTARDAS* software package (v.1.4.1) jointly developed by the ASDC (Italy) and the California Institute of Technology (Caltech, USA). Event files were calibrated and cleaned with standard filtering criteria using the *nupipeline* task (version 20150316) of the *NuSTAR* CALDB.

The two sources were well detected in the *NuSTAR* 3–79 keV energy band. In both cases, the FPMA and FPMB spectra of the target were extracted from the cleaned event files using a circle of 20 pixel (~ 49 arcsec) radius, while the background was extracted from nearby circular regions of 40 pixel radius. The ancillary response files were generated with the *numkarf* task, applying corrections for the point spread function (PSF) losses, exposure maps and vignetting. All spectra were binned to ensure a minimum of 30 counts per bin.

2.2 Swift observations

The *Swift* satellite observed S5 0014+81 on 2014 December 21 (obsID 00080003001) and on 2015 January 23 (obsID 00080003002), while B0222+185 was observed on 2014 December 24 (obsID 00080243001) and on 2015 January 18 (obsID 00080243002). The total net exposure times were 6.5 ks (2014 December) and 6.6 ks (2015 January) for S5 0014+81 and 4.9 ks (2014 December) and 5.1 ks (2015 January) for B0222+185.

2.2.1 XRT observations

Swift/XRT (Burrows et al. 2005) observations were carried out using the photon counting CCD readout mode and in the four observations the sources were well detected in the 0.3–10 keV XRT energy band. The XRT data sets were first processed with the *XRTDAS*

¹ Data retrieved from <http://irsa.ipac.caltech.edu/>.

² <http://tools.asdc.asi.it/>

Table 1. Parameters of the X-ray spectral analysis, from the simultaneous fit of *NuSTAR* and *Swift*/XRT. The errors are at 90 per cent level of confidence for spectral index and break energy, at 68 per cent for the fluxes. Fluxes are corrected for galactic absorption.

S5 0014+81							
Date	Γ_1	Γ_2	E_{break} (keV)	$F_{0.5-2\text{keV}}$ ($\text{erg cm}^{-2} \text{s}^{-1}$)	$F_{2-10\text{keV}}$ ($\text{erg cm}^{-2} \text{s}^{-1}$)	$F_{10-50\text{keV}}$ ($\text{erg cm}^{-2} \text{s}^{-1}$)	χ^2 / dof
2014 Dec 1	$1.18^{+0.22}_{-0.27}$	1.72 ± 0.05	$2.22^{+0.83}_{-0.58}$	$1.37^{+0.04}_{-0.1} \times 10^{-12}$	$3.83^{+0.15}_{-0.25} \times 10^{-12}$	$5.98^{+0.31}_{-0.71} \times 10^{-12}$	151.8/136
2015 Jan 23	$0.64^{+0.44}_{-0.61}$	1.65 ± 0.05	$1.61^{+0.61}_{-0.34}$	$1.31^{+0.04}_{-0.14} \times 10^{-12}$	$3.62^{+0.22}_{-0.44} \times 10^{-12}$	$6.36^{+0.28}_{-0.79} \times 10^{-12}$	133.2/147
B0222+185							
Date	Γ_1	Γ_2	E_{break} (keV)	$F_{0.5-2\text{keV}}$ ($\text{erg cm}^{-2} \text{s}^{-1}$)	$F_{2-10\text{keV}}$ ($\text{erg cm}^{-2} \text{s}^{-1}$)	$F_{10-50\text{keV}}$ ($\text{erg cm}^{-2} \text{s}^{-1}$)	χ^2 / dof
2014 Dec 24	$1.10^{+0.19}_{-0.10}$	$1.56^{+0.09}_{-0.03}$	$4.77^{+2.61}_{-0.55}$	$2.62^{+0.07}_{-0.12} \times 10^{-12}$	$1.22^{+0.01}_{-0.03} \times 10^{-11}$	$2.67^{+0.04}_{-0.07} \times 10^{-11}$	483.6/504
2015 Jan 18	1.06 ± 0.12	$1.69^{+0.05}_{-0.04}$	$4.69^{+0.83}_{-0.56}$	$1.86^{+0.09}_{-0.12} \times 10^{-12}$	$8.29^{+0.08}_{-0.18} \times 10^{-12}$	$1.51^{+0.03}_{-0.05} \times 10^{-11}$	406.8/400

software package (v.3.0.0) developed at the ASDC and distributed by HEASARC within the `HEASOFT` package (v. 6.16). In particular, event files were calibrated and cleaned with standard filtering criteria with the `xrtpipeline` task using the calibration files available in the version 20140709 of the `Swift`/XRT `CALDB`.

The energy spectra were then extracted from the calibrated and cleaned event files. Events for the source spectral analysis were selected within a circle of 20 pixel (~ 47 arcsec) radius, enclosing about 90 per cent of the PSF, while the background was extracted from a nearby circular region of 80 pixel radius. The ancillary response files were generated with the `xrtmkarf` task, applying corrections for the PSF losses and CCD defects using the cumulative exposure map. The source spectra were binned to ensure a minimum of 30 counts per bin.

2.2.2 *UVOT* observations

Swift/*UVOT* (Roming et al. 2005) observations were performed with all six optical and UV lenticular filters (namely *V*, *U*, *B*, *W1*, *M2*, *W2*). We performed aperture photometry for all filters in all the observations using the standard *UVOT* software distributed within the `HEASOFT` package (version 6.16) and the calibration included in the latest release of the `CALDB`. Counts were extracted from apertures of 5 arcsec radius for all filters and converted to fluxes using the standard zero-points (Poole et al. 2008). The fluxes were then de-reddened using the appropriate values of $E(B - V)$ taken from Schlegel, Finkbeiner & Davis (1998) and Schlafly & Finkbeiner (2011) with $A_\lambda / E(B - V)$ ratios calculated for *UVOT* filters using the mean Galactic interstellar extinction curve from Fitzpatrick (1999). No variability was detected within single exposures in any filter. The processing results were carefully verified, checking for possible contaminations from nearby objects within the source apertures and from objects falling within the background apertures.

2.3 X-ray spectral analysis

The spectral analysis of the 2014 December and 2015 January *NuSTAR* and *Swift*/XRT simultaneous observations of S5 0014+81 and B0222+185 were performed using the `XSPEC` package. In all four observations, a broken power-law model with an absorption hydrogen-equivalent column density fixed to the Galactic value of $1.35 \times 10^{21} \text{ cm}^{-2}$ (S5 0014+81) and $9.4 \times 10^{20} \text{ cm}^{-2}$ (B0222+185) (Kalberla et al. 2005) was found to provide a good description of the observed spectra in the 0.3–79 keV energy band. The inter-

calibration factors between the three instruments (*NuSTAR*/FPMA, *NuSTAR*/FPMB and *Swift*/XRT) were taken into account adding a multiplicative constant (kept to 1 for *NuSTAR*/FPMA) to the spectral model. We found values in the 2 per cent range for *NuSTAR*/FPMB and in the 10 per cent range for *Swift*/XRT which are consistent with the cross-calibration uncertainties for the instruments (Madsen et al. 2015). The results of the spectral fits are shown in Table 1, and Fig. 1 shows the X-ray spectra of the two sources.

These results describe intrinsic broken power laws, not consistent with absorption. We tested different spectral models, namely single power laws with an absorption hydrogen equivalent column density fixed to the Galactic value or left free to vary. These were possibilities explored by other authors also for these sources (e.g. Page et al. 2005; Piconcelli & Guainazzi 2005; Tavecchio et al. 2007; Eitan & Behar 2013). When we left N_{H} free to vary, the intercalibration factor between *Swift*/XRT and *NuSTAR*/FPMA-FMPB is no more consistent with the cross-calibration uncertainties for the instruments, with values that differ of more than 25 per cent. The χ^2 associated with this option is significantly higher than the broken power-law option, for all the observations. We performed a second test by fixing the absorption column density at the Galactic value, in the case of a single power law. The intercalibration factor in this case is even less consistent with the acceptable cross-calibration uncertainties (i.e. >35–40 per cent), along with higher χ^2 values. The results from these tests confirm the existence of an intrinsic spectral curvature within the observed 0.3–79 keV energy band.

3 BLACK HOLE MASS ESTIMATE

The black hole mass of a blazar is an important feature to characterize it. When optical or infrared spectra are not available, the virial mass estimate method cannot be applied. At high redshift, the available lines to apply such a method are also less reliable. With a good photometric coverage of the UV-optical-IR band, instead, the accretion disc emission can be fitted, and the black hole mass can be estimated from this fitting process (Calderone et al. 2013).

Fig. 2 shows the optical-UV band of the SEDs, along with *Swift*/*UVOT* photometric data, archival data from the ASDC³, that cover well the whole band. First, we need to point out that at frequencies higher than $\log(\nu/\text{Hz}) = 15.4$ rest frame a prominent absorption feature is usually present, i.e. the Ly α forest, due to intervening clouds absorbing hydrogen Ly α photons at wavelengths

³ www.asdc.asi.it

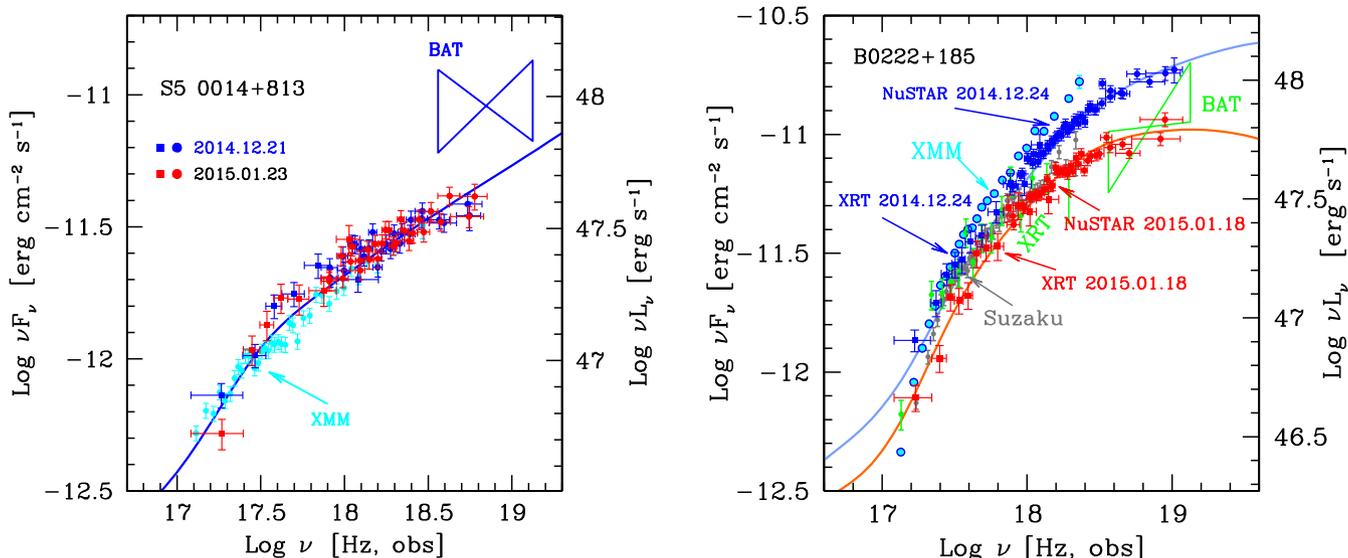


Figure 1. X-ray spectra of S5 0014+813 and B0222+185, along with the SED models described in the text with parameters as in Tables 2 and 3 (shown in solid blue and red lines). In both panels, new *Swift*/XRT and *NuSTAR* data are, respectively, filled squares and circles. The 2014 December observations are in blue, the 2015 January ones are in red. Left-hand panel: X-ray spectrum of S5 0014+813. Archival *Swift*/BAT data are shown in blue, while the cyan circles are archival *XMM* data, as labelled. These data were used and commented in detail in Ghisellini et al. (2009; 2010). Right-hand panel: X-ray spectrum of B0222+185. Archival data are shown in green (*Swift*/XRT and BAT), grey (*Suzaku*) and cyan (*XMM-Newton*). They were shown and described in detail in Ghisellini et al. (2009).

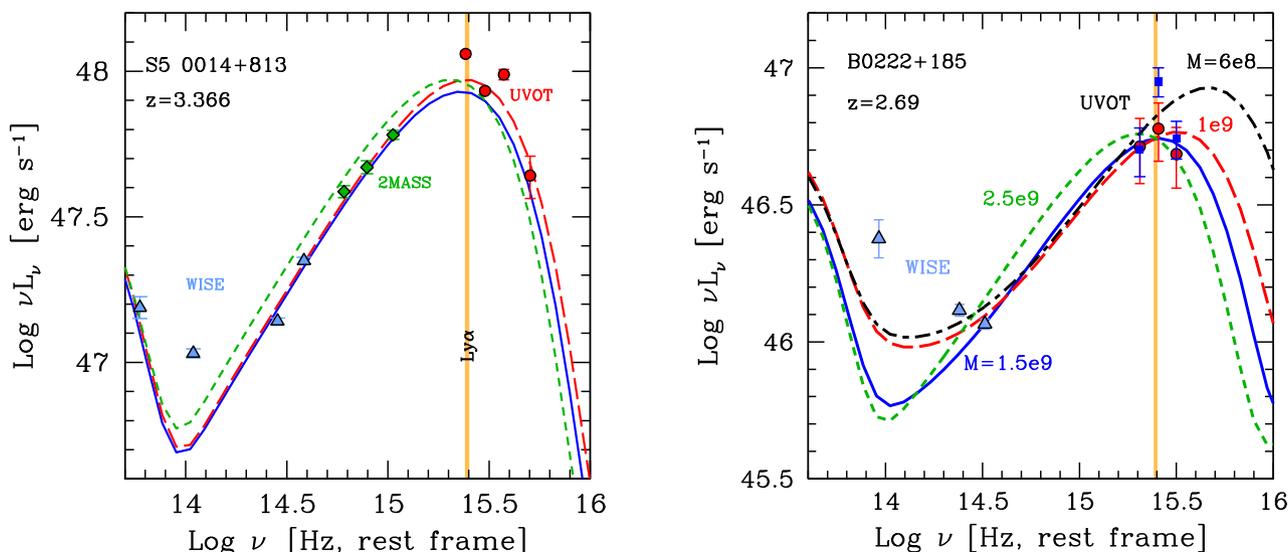


Figure 2. IR-optical-UV SEDs of S5 0014+813 and B0222+185, along with models used to derive L_d and M_{BH} . In both panels, the yellow vertical line highlight the Ly α line position. At frequencies larger than this line, the *Swift*/uvot data points have been corrected for the absorption by intervening clouds. The cloud distribution, though, is random and strongly dependent on the line-of-sight. Hence the absorption correction is only statistical, with little reliability on single sources. We do not consider this wavelength range for our modelling. Left-hand panel: S5 0014+815 IR-optical-UV SED. The red circles are new *Swift*/uvot data. Green diamonds are 2MASS and blue triangles are *WISE* data points. The solid blue line corresponds to our best model, with $M_{BH} = 7.5 \times 10^9 M_\odot$ and $L_d = 0.85 L_{Edd} = 8.3 \times 10^{47} \text{ erg s}^{-1}$. The green short dashed line shows a model with $M_{BH} = 10^{10} M_\odot$ and $L_d = 0.7 L_{Edd} = 9.1 \times 10^{47} \text{ erg s}^{-1}$, while the red long dashed line with $M_{BH} = 7 \times 10^9 M_\odot$ and $L_d = L_{Edd} = 9.1 \times 10^{47} \text{ erg s}^{-1}$. The last two represent the ‘confidence range’ of M_{BH} and L_d . Right-hand panel: B0222+183 IR-optical-UV SED. Red circles are *Swift*/uvot data from the January observation, blue squares from the December observation. Blue triangles are archival *WISE* data points. The solid blue line is the best model, with $M_{BH} = 1.5 \times 10^9 M_\odot$ and $L_d = 0.27 L_{Edd} = 5.3 \times 10^{46} \text{ erg s}^{-1}$. The red long dashed line corresponds to $M_{BH} = 1 \times 10^9 M_\odot$ and $L_d = 0.42 L_{Edd} = 5.5 \times 10^{46} \text{ erg s}^{-1}$, while the green short dashed to $M_{BH} = 2.5 \times 10^9 M_\odot$ and $L_d = 0.17 L_{Edd} = 5.5 \times 10^{46} \text{ erg s}^{-1}$. The black dot-dashed line represents instead a tentative model obtained by trying to fit the uvot December data at $\nu = 10^{15.4} \text{ Hz}$. The last two represent the ‘confidence range’ of M_{BH} and L_d .

shortward of 1216 Å. We corrected for the absorption following Ghisellini et al. (2010). However, the distribution of intervening clouds varies randomly along every line of sight. Only an average correction can be done, and it is not sufficiently reliable when ap-

plied on single sources. For this reason, we do not consider the data point at $\log(\nu/\text{Hz}) > 15.4$ in our modelling. The bluer uvot bands fall in this frequency range in the case of S5 0014+318 and B0222+185. At frequencies lower than this prominent absorption

feature, though, a peak in the SEDs is clearly visible. Below this peak, the optical flux decreases with frequency, suggesting a power-law trend, especially in S5 0014+813. This is the clear signature of an accretion disc, which can be fitted with a simple Shakura–Sunyaev model (Shakura & Sunyaev 1973; Calderone et al. 2013; Sbarrato et al. 2013).

At lower frequencies, another feature is evident from the IR-optical-UV SEDs of these two sources: the *WISE* IR bands show an increase of the flux, that breaks the power-law-like trend in the optical. This is likely the signature of the IR emission from a dusty torus around the nucleus. Such a steep IR spectrum, in fact, could not be produced by synchrotron emission: either a self-absorption frequency larger than $\sim 10^{13}$ Hz or a steep thin synchrotron spectrum ending with an exponential cut would be needed to justify such a spectral profile. Both options would show up with prominent signatures in the high-energy emission of the sources, that we do not observe. Fig. 2 shows that our models do not perfectly reproduce far-IR data. In fact our model oversimplifies the torus emission: we describe it as a blackbody emission, while there is evidence (see e.g. Calderone, Sbarrato & Ghisellini 2012) that it is best represented as a multitemperature structure. The hottest part, closer to the accretion disc, has likely a temperature ~ 2000 K, i.e. of the order of the dust sublimation temperature. Our far-IR data, in fact, show an increase in flux at frequencies $\sim 1.6 \times 10^{14}$ Hz, that roughly corresponds to these range of temperatures.

With these premises, a reliable way to estimate the black hole mass of S5 0014+318 and B0222+185 is to fit their IR-optical-UV SEDs with a simple model of accretion disc emission. We applied the radiatively efficient, geometrically thin, optically thick Shakura & Sunyaev (1973) model. Assuming a standard radiative efficiency $\eta = 0.08$, only two free parameters are left to be fitted: the accretion rate \dot{M} , that can be traced by the intrinsic disc luminosity $L_d = \eta \dot{M} c^2$, and the black hole mass M_{BH} itself. In the case of S5 0014+318, we can constrain the overall disc luminosity thanks to the visibility of the peak of the disc emission, with some consideration regarding its anisotropic properties (as thoroughly explained by Calderone et al. 2013):

(i) according to the Shakura–Sunyaev model, the peak luminosity $v_p L_{v_p}$ corresponds to half the total observed luminosity $L_{\text{obs}} = 2v_p L_{v_p}$;

(ii) the observed luminosity depends on the viewing angle of the accretion disc:

$$L_{\text{obs}} = 2 \cos \theta_v L_d, \quad (1)$$

where L_d is the intrinsic total luminosity emitted by the accretion disc. In the case of a blazar $L_{\text{obs}} \simeq 2 L_d$ since we see the accretion disc face-on.

We can therefore derive the intrinsic total luminosity from the peak luminosity of our sources:

$$L_d = \frac{v_p L_{v_p}}{\cos \theta_v} \simeq v_p L_{v_p}. \quad (2)$$

This means that in the case of S5 0014+318, L_d is constrained by observations (i.e. 2MASS and UVOT data), and only M_{BH} is left as a free parameter to be derived with the IR-optical-UV SED fitting.

We find that both sources have large black hole masses and are fast accreting, even if not super-Eddington. We derive $M_{\text{BH}} = 1.5 \times 10^9 M_{\odot}$ and $L_d = 5.3 \times 10^{46} \text{ erg s}^{-1}$ for B0222+185 and $M_{\text{BH}} = 7.5 \times 10^9 M_{\odot}$ and $L_d = 8.3 \times 10^{47} \text{ erg s}^{-1}$ for S5 0014+813. These values are significantly smaller than what was derived in Ghisellini et al. (2009 and 2010). For S5 0014+813, the reason is due to (i)

the better coverage of the IR band achieved with *WISE* data and (ii) neglecting the optical data taken from Bechtold et al. (1994). We now prefer to discard those data because the derivation of flux and luminosities are not sufficiently clear in that paper. The confidence range of S5 0014+813 mass ($7 \times 10^9 - 10^{10} M_{\odot}$) is indicated by the dashed lines in the left-hand panel of Fig. 2. Note that it strongly depends on data quality. In this case, the range is rather narrow because of very good data. More precise data would lead to even more refined estimates. A lower limit on the mass is anyway fixed by the strong constraint given by the Eddington limit.

For B0222+185, the smaller values of M_{BH} and L_d are due to the new infrared data (not available in the previous work), that now help in roughly constraining the peak frequency of the disc emission. The data coverage in this case is not enough to constrain the peak frequency luminosity, therefore the estimate on the black hole mass is less constrained. This should be taken as an indication of M_{BH} , not as a best fit.

A side result of M_{BH} and L_d studies is an estimate of the broad-line region covering factor with respect to the accretion disc f_{BLR} . The BLR is thought to reprocess a fraction $f_{\text{BLR}} \sim 0.05 - 0.2$ of the radiation emitted from the disc, thus usually a standard value $f_{\text{BLR}} \simeq 0.1$ is used. When BLR and disc luminosities are obtained independently, the BLR covering factor can be derived, and this is the case. Cao & Jiang (1999) derived $L_{\text{BLR}} = 4.348 \times 10^{46} \text{ erg s}^{-1}$ for S5 0014+813. By comparing it with our result, we obtain a covering factor $f_{\text{BLR}} = 0.05$.

4 MODELLING THE BROAD-BAND SED

Fig. 3 shows that both S5 0014+318 and B0222+185 have overall SEDs characterized by a prominent high-energy component, that along with the characteristic flat and intense radio luminosity is attributed to non-thermal emission from a relativistic jet. In the IR-optical-UV range of both sources, the SEDs are dominated by thermal emission attributed to the accretion disc, as discussed in Section 3.

Not being detected in the γ -rays by *Fermi*/LAT, X-ray data are necessary to study the non-thermal high-energy emission of S5 0014+318 and B0222+185. Specifically, *NuSTAR* data are crucial for understanding the X-ray spectral profile and possible variability in this kind of high-redshift source, as can be seen in Fig. 1. X-ray data contribute significantly to the modelling of the broad-band SEDs of the two sources (Fig. 3).

To interpret the SEDs of the two sources, we used a leptonic one-zone emitting model, fully described in Ghisellini & Tavecchio (2009). We refer to the original paper for details, providing here only a very brief description of the most important features of the models. The emitting source is assumed to be a spherical region in which relativistic electrons emit by synchrotron and inverse Compton processes. This homogeneous spherical blob is assumed to be located at a distance R_{diss} from the central black hole, moving with a bulk Lorentz factor Γ at an angle θ_v from our line of sight. Relativistic electrons are injected throughout the source, with a power P'_1 as measured in the comoving frame. The energy distribution $Q(\gamma)$ of the injected electrons is a smoothly broken power law with slopes s_1 and s_2 (defined as $Q(\gamma) \propto \gamma^{-s}$) below and above the random Lorentz factor γ_b . Note that, even if $Q(\gamma)$ is a broken power law, the particle energy distribution N_γ derived through the continuity equation maintains a break, albeit smoother than the injected broken power law. This produces a gently curved spectrum, as shown in Fig. 1. The broad-line region is located at a distance $R_{\text{BLR}} = 10^{17} L_{d,45}^{1/2} \text{ cm}$ from the black hole, while the infrared

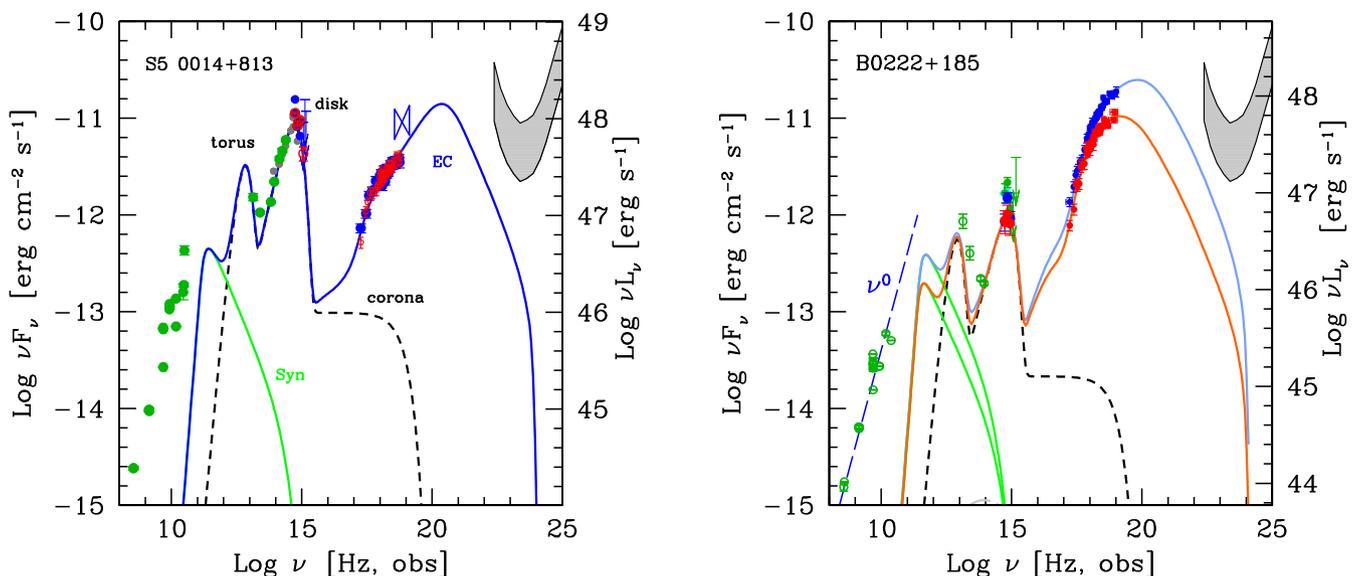


Figure 3. Broad-band SEDs of S5 0014+813 and B0222+185 along with the models discussed in the text and parametrized as in Tables 2 and 3. In both panels, the grey stripe is the 5σ *Fermi*/LAT sensitivity limit, calculated for 5 yr (lower edge) and 1 yr of operations (upper edge). New *Swift*/UVOT, *Swift*/XRT and *NuSTAR* data for the two observation periods are red (2015 January) and blue (2014 December) circles. Left-hand panel: S5 0014+813 SED with its broad-band model (blue solid line). The green solid line is the self-absorbed synchrotron emission, while the thermal emission from accretion disc, torus and X-ray corona is shown with the dashed black line. Green data points in radio and IR are from the literature (for details see Ghisellini et al. 2009). Archival *Swift*/BAT data are shown in blue. Right-hand panel: SED of B0222+185 with the two models corresponding to the low state (orange solid line) and the high state (blue solid line). Synchrotron emission in the two states are both shown with green solid lines. The thermal emission from the nuclear region (black dashed line) does not vary between the two states. Green open circles are archival data (from ASDC).

Table 2. Input parameters used to model the SED. Col. (1): source name. H indicates the higher state, L the lower; Col. (2): redshift; Col. (3): black hole mass in solar mass units (see Fig. 2 and Section 3 for the confidence range); Col. (4): distance of the blob from the black hole in units of 10^{15} cm. The size of the emitting region is defined as $R_{\text{blob}} = \psi R_{\text{diss}}$, where $\psi = 0.1$ rad is the jet aperture angle; Col. (5): radius of the BLR in units of 10^{15} cm; Col. (6): radius of the torus in units of 10^{15} cm; Col. (7): disc luminosity in units of 10^{45} erg s^{-1} (see Fig. 2 and Section 3 for the confidence range); Col. (8): disc luminosity in units of the Eddington luminosity; Col. (9): bulk Lorentz factor; Col. (10): viewing angle (degrees); Col. (11): power injected in the blob calculated in the comoving frame, in units of 10^{45} erg s^{-1} ; Col. (12): magnetic field in Gauss; Col. (13) and (14): break and maximum random Lorentz factors of the injected electrons; Col. (15): random Lorentz factors of the electrons cooling in R/c ; Col. (16) and (17): slopes of the injected electron distribution ($Q(\gamma)$) below and above γ_b . The spectral shape of the corona is assumed to be $\propto \nu^{-1} \exp(-h\nu/150 \text{ keV})$.

Source (1)	z (2)	M_{BH} (3)	R_{diss} (4)	R_{BLR} (5)	R_{T} (6)	L_{d} (7)	$L_{\text{d}}/L_{\text{Edd}}$ (8)	Γ (9)	θ_{v} (10)	P'_{i} (11)	B (12)	γ_{b} (13)	γ_{max} (14)	γ_{cool} (15)	s_1 (16)	s_2 (17)
S5 0014+813	3.336	7.5e9	1350	2878	7.2e4	829	0.85	14	3	0.042	3.5	45	2e3	1	-1	3.3
B0222+185 H	2.690	1.5e9	360	725	1.8e4	52.7	0.27	14	3	0.065	3.3	35	2e3	3.2	1	3.2
B0222+185 L	2.690	1.5e9	360	725	1.8e4	52.7	0.27	14	3	0.032	4.7	17	2e3	3	1.2	3.2

emitting torus is at $R_{\text{torus}} = 2.5 \times 10^{18} L_{\text{d},45}^{1/2}$ cm. $L_{\text{d},45}$ is the accretion disc luminosity in units of 10^{45} erg s^{-1} , and it is derived as in Section 3, together with the central black hole mass. The values of the parameters adopted for the models are reported in Table 2. Note that the model we apply is very sensitive to changes in the derived parameters. The emission profile and intensity reproduced by the SED fitting change significantly even after small parameter variations, as shown for small variations in the viewing angle in fig. 3 of Sbarrato et al. (2015).

Table 3 reports the different forms of the power carried by the jet: the power P_{r} spent in producing the radiation we observe, the Poynting flux P_{B} , the power associated with the bulk motion of relativistic electrons (P_{e}) and cold protons (P_{p}), assuming one proton per relativistic electron. This assumption is consistent with independent results on blazar and GRB jets by Nemmen et al. (2012). They found that the total jet power for both classes is 10 times the radiative power P_{rad} , i.e. similar to what we find in this work (see Table 3). Different proton-to-relativistic electron ratios were explored

Table 3. Logarithm of the jet power in different forms. Col. (1): source name. H indicates the higher state, L the lower; Col. (2): jet power in the form of radiation; Col. (3): jet power connected to Poynting flux; Col. (4): jet power in form of bulk motion of electrons; Col. (5): jet power in form of bulk motion of protons (assuming one cold proton per emitting electron).

Source (1)	$\log P_{\text{r}}$ (2)	$\log P_{\text{B}}$ (3)	$\log P_{\text{e}}$ (4)	$\log P_{\text{p}}$ (5)
S5 0014+813	46.4	47.2	44.6	47.3
B0222+185 H	46.5	46.0	45.5	48.2
B0222+185 L	46.1	46.3	45.3	48.1

by Sikora & Madejski (2000), who found that the relativistic pairs must be less than 10–20 per proton. With this combination, the total jet power can result equal to or even less than the radiative power, that instead is only a part of the total power carried by the jet, and hence should be a lower limit to the total P_{j} (Ghisellini 2012; Ghisellini et al. 2014). Therefore, assuming one proton per

relativistic electron is reasonable to explain the observed jet features and its physics.

The emitting regions of both sources are located within the broad-line region (and the infrared torus). In this way, the energy density of photons from the broad-line region feed the inverse Compton process, together with photons from the torus. The inverse Compton process is dominated by EC instead of SSC, as expected in FSRQs. Tagliaferri et al. (2015) obtained the same result for two other $z > 2$ blazars observed by *NuSTAR*: the emitting regions of both S5 0836+710 and PKS 2149–306 are located between the BLR and IR torus. Their results were obtained through SED fitting, and were also confirmed on the basis of the variability time-scales obtained with two *NuSTAR* observations per source.

5 DISCUSSION

Blazars are characterized by their prominent relativistically boosted jet emission. They usually show prominent high-energy emission, which results in high γ -ray luminosities, well detected by instruments like *Fermi*/LAT. In some cases, though, blazars are not detected in such energy bands. This is the case for S5 0014+318 and B0222+185.⁴ Even if lacking a high-energy detection, the Compton bump can be observed in the X-ray frequency range, but a detection in the soft X-rays usually is not enough to determine the relativistic jet features of a blazar, nor its orientation. S5 0014+318 and B0222+185 were detected by *Swift*/BAT, but these data were not precise enough to derive exact estimates of the bulk Lorentz factors and viewing angles. Figs 1 and 3 show that the *Swift*/BAT data do not have enough precision to constrain the hard X-ray slope. *NuSTAR*, however, provides a broad-band, precise measurement for both sources, confirming that both blazars are seen at small viewing angles, i.e. their jets are directed along our line of sight. Both sources also host massive central black holes (both with $M_{\text{BH}} > 10^9 M_{\odot}$).

5.1 Variability

Looking in detail at the broad-band SEDs of these two objects, we see some interesting differences. Both sources show the two humps, i.e. the signature of aligned jet emission, while in the IR-optical-UV band, the accretion disc emission dominates over the non-thermal jet emission. Comparing the two panels of Fig. 3, it can be noticed that S5 0014+318 and B0222+185 also show different variability behaviours. They were both observed in two epochs separated by ~ 1 month. S5 0014+318 does not show flux or spectral variation between the two observations, and the new data are consistent with archival data. Only the *Swift*/BAT detection could suggest a different state of the source, but due to the large uncertainty we cannot draw any strong conclusion.

B0222+185, instead, shows a clear variation of the X-ray flux between the two *Swift*/XRT + *NuSTAR* observations: in 2014 December the source was in a higher state, compared to both 2015 January and archival data. The different hard X-ray spectra showed

⁴Tavecchio et al. (2007) predicted with a previous modelling that B0222+185 would have been detected by *Fermi*/LAT in its first year of operation, but this did not happen. The authors did not have any IR data, though, to constrain the torus emission and hence the synchrotron component. Their analysis thus lead to a flatter synchrotron and EC peak, in principle detectable by *Fermi*/LAT. The steeper slope we now observe thanks to IR data, instead, is consistent with non-detection in the γ -rays.

in the two B0222+185 observations (see right-hand panel of Fig. 1) suggest that the peak of the high-energy hump is at the same or higher frequency when in the higher state, compared to the lower state. This would be opposite to the general trend displayed by the blazar sequence, but missing a higher frequency detection, this speculation is not conclusive. If real, such behaviour would not be uncommon in rapidly varying FSRQs: although they follow the blazar sequence when considering different sources, an individual object can behave opposite to the sequence itself while varying. According to the model shown in Fig. 3, B0222+185 variability can be described by a variation in the injected power (see second and third lines of Table 2), accompanied by γ_b increasing in the high state.

Another remarkable example of this kind of variation, very similar to the one showed by B0222+185, but much more pronounced, has recently been seen for S5 0836+710 (Ciprini et al. 2015; Giroletti et al. 2015; Vercellone et al. 2015) during its 2015 August γ -ray flare, which triggered observations at X-ray and radio frequencies. The amplitude of the flux variation was huge in the *Fermi*/LAT band (factor 65 greater than the average flux reported in the third *Fermi*/LAT catalogue of Acero et al. 2015) and rather modest in the high energy part of the *Swift*/XRT band. This implies that the X-ray flux had to change more at larger energies, to connect to the enhanced γ -ray flux, and that the peak frequency of the high-energy hump must be ‘bluer’ than what was displayed during the *NuSTAR* observation described in Tagliaferri et al. (2015). In other words: if *NuSTAR* had followed the 2015 August flare of S5 0836+710, it likely would have detected a clear flux and spectral variation, leading to a predicted shift of the high-energy hump towards higher frequencies even in the absence of γ -ray data. The right-hand panel of Fig. 1 shows that *Swift*/XRT would not have been able to discriminate between the two states of B0222+185, while *NuSTAR* distinguishes them clearly. We conclude that:

- (i) observations in the hard X-ray band are not only instrumental to discover the most powerful blazars (that requires a survey of a large portion of the sky), but also to detect large flux variations that occur around the MeV energy band and are not noticeable in the classic 0.3–10 keV band, nor at energies beyond 100 MeV, where the source could go undetected even during a large flare; and
- (ii) repeated hard X-ray observations on time-scales of a few weeks up to 1–2 months are efficient to spot variability in high-redshift blazars. We managed to see variability in B0222+185, with similar features as the already observed variable FSRQs S5 0836+710 and PKS 2149–306 discussed by Tagliaferri et al. (2015).

5.2 Jets and accretion of the two most powerful blazars

We now aim to frame S5 0014+318 and B0222+185 within the larger blazar picture. We consider them in the jet–accretion correlation scenario. Ghisellini et al. (2014) found that in blazars the jet power not only correlates with the accretion power, but it is even larger. This suggests that accretion is strongly related to jet power, implying a role in jet production. At the same time, the fact that jet power is larger than accretion power tells us that some other process must play a role in the jet launch and acceleration. Black hole spin is the best candidate to play such a role. This result was obtained by studying a sample of *Fermi*-detected blazars, for which Shaw et al. (2012, 2013) obtained optical spectra. Ghisellini et al. (2014) selected all the objects with broad emission lines, in order to have a proxy of accretion luminosity, and compared jet and

accretion power for the 226 blazars in this sample. However, this sample did not include the most extreme blazars known, leaving open the questions: how does the jet–accretion relation look in the case of the most powerful blazars? Does the power balance change when accretion or jet emission are extreme? These questions will guide us in the following discussion.

First, we add to the original blazar sample the sources expected to be the most powerful. To this aim, we select the $z > 2$ blazars detected by *Swift*/BAT and all known high-redshift ($z > 4$) blazars. The BAT sensitivity limit is not very deep, and at high redshift it can detect mainly the most powerful sources, whose high-energy components peak in the \sim MeV range. BAT detected 10 $z > 2$ blazars, including S5 0014+318 and B0222+185, that we add to the blazar sample of Ghisellini et al. (2014). We also include all the known blazars at $z > 4$, as listed in Ghisellini et al. (2015). Being the highest redshift blazars currently known, they are expected to be among the most powerful blazars. They are not present in the BAT blazar catalogue because their distance makes their hard X-ray flux too weak for a detection with BAT. Since most of them were selected starting from optical catalogues, they are likely very powerful in accretion luminosity.

Fig. 4 shows how these samples are located in the overall jet–accretion relation, along with S5 0014+318 and B0222+185. The total jet power (calculated as the sum of different jet power components listed in Table 3) is plotted as a function of the disc luminosity. The grey stripes show the best fit of the sample by Ghisellini et al. (2014). Note that the powerful blazars we add in this work are all located within the 2σ dispersion of the previous correlation. This means that they still follow the jet–accretion relation found by Ghisellini et al. (2014), even if they are among the most powerful sources in the set.

S5 0014+318 and B0222+185 can be considered the two most extreme sources: respectively, they are the blazars with the most luminous disc and the most powerful jet. Still, they are close enough to the known jet–accretion correlation, to be less than 2σ from the Ghisellini et al. (2014) result. Thus we conclude that even the most powerful blazars follow the same jet–accretion relation as the γ -ray-detected bulk sample. The second interesting conclusion that we can draw from this comparison is that *NuSTAR* is once again confirmed to be the most suitable telescope to study the most powerful blazars in our Universe.

6 CONCLUSIONS

The simultaneous X-ray observations of S5 0014+318 and B0222+185 performed with *Swift*/XRT and *NuSTAR* gave us an interesting insight into the jet emissions of these two sources. We confirmed their blazar nature, with a refined estimate of their bulk Lorentz factors and viewing angles, supported by more precise sets of parameters (Tables 2 and 3). The accretion disc fitting to a more complete data set gave us the possibility to refine our previous estimates of the black hole mass and accretion luminosity of these two sources, implying fast accreting objects with extreme masses of $> 10^9 M_{\odot}$.

The overall SED modelling allowed us to estimate the jet power and accretion disc luminosity, allowing a comparison of these two sources with the overall blazar jet–accretion relation. The two sources are among the most powerful blazars known, and they populate the highest disc luminosity and jet power part of the jet–accretion correlation (Ghisellini et al. 2014; Fig. 4). It is remarkable that a sample formed by the most powerful blazars known is still within 2σ from the correlation derived from a sample of

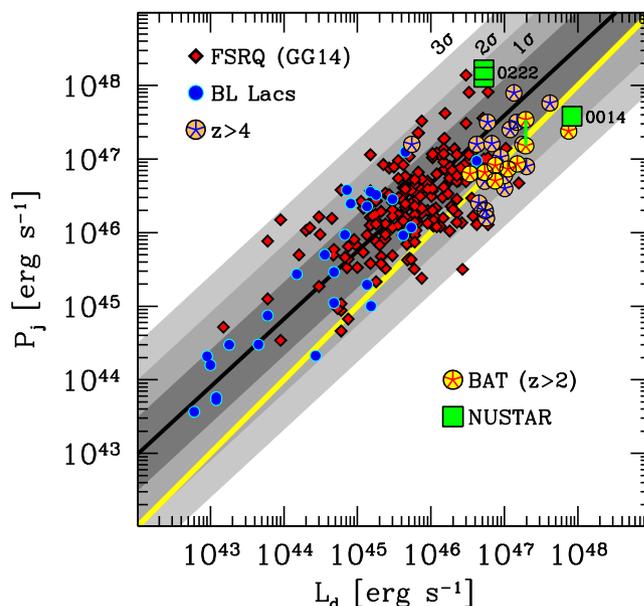


Figure 4. Total jet power as a function of accretion disc luminosity in blazars. Red (outlined in black) diamonds and blue filled circles are, respectively, FSRQs and BL Lacs from Ghisellini et al. (2014). Circled asterisks are high-redshift, non- γ -ray detected blazars: pink filled circles with blue asterisks are all the known $z > 4$ blazars (see Ghisellini et al. 2015), yellow filled circles with red asterisks are the $z > 2$ blazars detected by *Swift*/BAT. Green filled squares are S5 0014+318 and B0222+185 (both states) as labelled, observed by *NuSTAR* for this work. Note that the latter are the two blazars with the most luminous accretion disc (S5 0014+318) and the most powerful jet (B0222+185). The black line and grey stripes are the best-fitting relation of γ -ray detected blazars from Ghisellini et al. (2014), along with the 1σ , 2σ and 3σ dispersions. The yellow line is the one-to-one correlation. S5 0836+710 underwent a prominent γ -ray flare, and therefore it has been plotted with the jet powers of both its average and flaring states. The two states are connected with a green line. This suggests that part of the spread could be due to different states of the single sources, along with flaring episodes. The jet power is calculated as the sum of the different components listed in Table 3. Note that the very powerful blazars we added to the original sample by Ghisellini et al. (2014) are located within 2σ from the γ -ray detected jet–accretion relation. This means that also in the most powerful sources the jet power correlates with accretion luminosity, but it is larger than the accretion power, leading to an important role of black hole spin in jet launching.

blazars whose γ -ray flux was averaged over two years. The mechanisms governing the jet formation and evolution in the most extreme sources must not be different from the processes powering the more moderate objects. S5 0014+318 and B0222+185 themselves are consistent with the relation derived from the γ -ray blazars, even if they are the blazars with the most luminous accretion disc and the most powerful jet, respectively.

We found a different variability behaviour between the two sources: while S5 0014+318 did not vary in the two observation epochs separated by ~ 1 month, B0222+185 shows a clear variation, with an amplitude larger at larger frequencies. This last feature is the main reason why *NuSTAR* is such a crucial instrument to study high-redshift and powerful blazars. A soft X-ray telescope alone could not see a variability event like the one shown by B0222+185: *NuSTAR* instead observes at frequencies where the amplitude of a flaring activity is large enough to be seen. Were we lacking the γ -ray signature of a flare, *NuSTAR* could do the job.

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