

## References

- André, P., & Montmerle, T. 1994, ApJ, 420, 837
- Bastien, P. 1982, A&AS, 48, 153
- Beckwith, S. V. W., & Sargent, A. I. 1991, ApJ, 381, 250
- Beckwith, S. V. W., & Sargent, A. I. 1993a, in Protostars and Planets III, edited by E. H. Levy  
and J. L. Lunine (University of Arizona Press, Tucson) p. 521
- Beckwith, S. V. W., & Sargent, A. I. 1993b, ApJ, 402, 280
- Beckwith, S. V. W., Sargent, A. I., Scoville, N. Z., Masson, C. R., Zuckerman, B. & Phillips,  
T. G. 1986, ApJ, 309, 755
- Beckwith, S. V. W., Sargent, A. I., Chini, R. & Güsten, R. 1990, AJ, 99, 924
- Cohen, M., & Kuhl, L. V. 1979, ApJS, 41, 743
- Cohen, M., Emerson, J. P. & Beichman, C. A. 1989, ApJ, 339, 455
- d'Antona, F., & Mazzitelli, I. 1993, ApJS, 90, 467
- de Geus, E. J., Bronfman, L., & Thaddeus, P. 1990, A&A, 231, 137
- Dutrey A., Guilloteau, S. & Simon, M. 1994, A&A, 286, 149
- Elias, J. 1978, ApJ, 224, 857
- Elsasser, H., & Staude, H. J. 1978, A&A, 70, L3
- Galli, D., & Shu, F. H. 1993, ApJ, 417, 220
- Ghez, A., Neugebauer, G. & Mathews, K. 1993, AJ, 106, 2005
- Hartmann, L., Hewett, R., Stahler, S., & Mathieu, R. D. 1986, ApJ, 309, 275
- Hayashi, M., Ohashi, N. & Miyama, S. M. 1993, ApJ, 418, L71
- Henning, T., Pfau, W., Zinnecker, H., & Prusti, T. 1994, A&A, 276, 129
- Herbig, G. H., & Bell, K. R. 1988, in Lick Observatory Bulletin No. 1111 (University of California)
- Herbig, G. H. 1977, ApJ, 214, 747
- Hirth, G. A., Mundt, R., Solf, J., & Ray, T. P. 1994, ApJ, 427, L99

- Leinert, Ch., Hass, M., Richi, A., Zinnecker, H., & Mundt, R. 1991, *A&A*, 250, 407
- Kawabe, R., Ishiguro, M., Omodaka, T., Kitamura, Y., & Miyama, S. M. 1993, *ApJ*, L63
- Kenyon, S. J., Calvet, N., & Hartmann, L. 1993, *ApJ*, 414, 676
- Koerner, D. W., Sargent, A. I., & Beckwith, S. V. W. 1993a, *Icarus*, 106, 2
- Koerner, D. W., Sargent, A. I., & Beckwith, S. V. W. 1993b, *ApJ*, 408, L93
- Nakajima, T., & Golimowski, D. A. 1995, *AJ*, in press
- Sargent, A. I., & Beckwith, S. 1987, *ApJ*, 323, 294
- Sargent, A. I., & Beckwith, S. V. W. 1989, in *Structure and Dynamics of the Interstellar Medium*,  
IAU Colloquium No. 120, 215
- Sargent, A. I., & Beckwith, S. 1991, *ApJ*, 382, L31
- Sargent, A. I., & Beckwith, S. 1993, in *Millimeter and Submillimeter Wave Interferometry*,  
edited  
by M. Ishiguro and R. Kawabe, (Bookcrafters, San Francisco) p. 232
- Scoville, N. Z., Sargent, A. I., Sanders, D. B., Claussen, M. J., Masson, C. R., Lo, K., &  
Phillips,  
T. G. 1986, *ApJ*, 303, 416
- Stahler, S. W., Korycansky, D. G., Brothers, M. J., & Touma, J. 1994, *ApJ*, 431, 341
- Stapelfeldt, K. R., Burrows, C. J., Krist, J., Hester, J. J., Trauger, J., Ballester, G.,  
Casertano, S.,  
Clarke, J., Crisp, D., Evans, R. W., Gallagher, J. S., Griffiths, R. E., Hoessel, J. G.,  
Holtzman, J. A., Mould, J. R., Scowen, P., Watson, A. M., & Westphal, J. A. 1994,  
*ApJ*, submitted
- Strom, K. M., Strom, S. E., Edwards, S., Cabrit, S., & Skrutskie, M. F. 1989, *AJ*, 97,  
1451
- Tamura, M., & Sato, S. 1989, *AJ*, 98, 1368
- Ungerechts, H., & Thaddeus, P. 1987, *ApJS*, 377, 510
- van Langevelde, H. J., van Dishoeck, E. F., & Blake, G. A. 1994, *ApJ*, in Press
- Weintraub, D. A., Masson, C. R., & Zuckerman, B. 1987, *ApJ*, 320, 336

Weintraub, D. A., Masson, C. R., & Zuckerman, B. 1989, ApJ, 344, 915

Whitney, B. A., & Hartmann, L. 1992, ApJ, 395, 529

Whitney, B. A., & Hartmann, L. 1993, ApJ, 402, 605

Whittet, D. C. B. 1974, MNRAS, 168, 371

## **Imaging the Small-Scale Circumstellar Gas Around T Tauri Stars**

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### Abstract

We have detected circumstellar molecular gas around a small sample of T Tauri stars through aperture synthesis imaging of CO(2→1) emission at  $\sim 2\text{-}3''$  resolution. RY Tauri, DL Tauri, DO Tauri, and AS 209 show resolved and elongated gaseous emission. For RY Tau, the deconvolved, half-maximum radius along the direction of elongation,  $PA\sim 48^\circ$ , is 110 AU. Corresponding radii and orientations for the other sources are: DL Tau – 250 AU at  $PA\sim 84^\circ$ ; DO Tau – 350 AU at  $PA\sim 160^\circ$ ; AS 209 – 290 AU at  $PA\sim 138^\circ$ . RY Tau, DL Tau, and AS 209 show velocity gradients parallel to the elongation, suggesting that the circumstellar material is rotating. RY Tau and AS 209 also exhibit double-peaked spectra characteristic of a rotating disk. Line emission from DO Tau is dominated by high-velocity blue-shifted gas which complicates the interpretation. Nevertheless, there is in each case sufficient evidence to speculate that the circumstellar emission may arise from a protoplanetary disk similar to that from which our solar system formed.

## 1. Introduction

There is increasing evidence that the nebular environment assumed to have given rise to Solar System bodies is a common feature of the early stages of stellar evolution. Circumstellar dust disks are detected around a large fraction of T Tauri stars (e.g., Cohen, Emerson, & Beichman 1989; Strom *et al.* 1989; Beckwith *et al.* 1990; André & Montmerle 1994; Henning *et al.* 1994) and in at least one case, the velocity field of the associated circumstellar gas is consistent with that of a model of a disk in Keplerian rotation (Koerner, Sargent, & Beckwith 1993a, henceforth KSB). Disk properties, such as mass, size, and kinematics, are sufficiently similar to those expected for the early solar nebula (cf. Beckwith & Sargent 1993a) that continued study is likely to afford us the opportunity to constrain theories of planetary system formation. A better understanding of the detailed properties of these circumstellar environments may also narrow down the field of truly proto-planetary disks.

The small fraction of T Tauri stars for which circumstellar gas has been observed (cf. Sargent & Beckwith 1993) contrasts markedly with the 50% detection rate achieved by continuum surveys. This is largely because single telescope observations of the circumstellar dust continuum emission are readily obtainable. In general, unambiguous detection of circumstellar gas on small spatial scales requires molecular line millimeter-wave interferometer measurements; historically, these have been time consuming. Flattened structures of circumstellar gas with radii of order 1000 AU have been observed around the T Tauri stars, HL Tauri (Beckwith *et al.* 1986; Sargent & Beckwith 1987), T Tauri (Weintraub *et al.* 1987; 1989), DG Tauri (Sargent & Beckwith 1989; Ohashi *et al.* 1991), GG Tauri (Kawabe *et al.* 1993; Koerner, Sargent, & Beckwith 1993b; Dutrey, Simon, & Guilloteau

1994) and GM Aurigae (KSB). The earliest interpretations suggested that circumstellar material was centrifugally supported out to radial distances of 1000 AU (Sargent & Beckwith 1987; Weintraub *et al.* 1989). Recent observations indicate that gas at these large radii is still in a state of collapse (e.g., Hayashi *et al.* 1993; van Langevelde, van Dishoeck, & Blake 1994). However, it also appears that some of these objects may not represent most T Tauri stars in general. For example, Hubble Space Telescope observations imply that HL Tau is heavily embedded in its parent cloud and is therefore at a very early stage of pre-main sequence evolution (Stapelfeldt *et al.* 1994); the spectral energy distribution (SED) of T Tau, a binary object, can be understood only if there is a large component of non-planar dust (Kenyon *et al.* 1993), and GG Tau is part of an unusual quadruple star system and surrounded by a large torus (Leinert *et al.* 1991; Ghez *et al.* 1993; Koerner *et al.* 1993b; Dutrey *et al.* 1994).

To date, GM Aur seems the best example of a “typical” T Tauri star. It is relatively unobscured with  $A_V = 0.1$  (Strom *et al.* 1989) and is significantly older than the other stars which show circumstellar gas. It is also not a binary and affords a closer example of the early Solar System environment than the large fraction of T Tauri stars with a stellar companion (Ghez *et al.* 1993). Here, there is unambiguous evidence that the surrounding molecular gas is centrifugally supported out to 1000 AU, but emission beyond a 150 AU radius is tenuous compared to the other circumstellar structures observed. This star/disk system may well represent a planet-forming environment (KSB). Obviously, we would like to observe a range of circumstellar gaseous environments for a larger sample of T Tauri stars to determine detailed properties of the disks and generalize these to compare with the conditions in the early solar nebula. We are continuing a program of aperture synthesis imaging of the environments of T Tauri stars with a view to building up a statistical base

of information on the circumstellar gas. High resolution molecular line images of four more systems, RY Tauri, DL tauri, DO Tauri, and AS 209 are presented and discussed here.

## 2. Observations

RY Tau, DL Tau, DO Tau, and AS 209 are among the strongest sources of millimeter continuum emission in the nearest star-forming regions (Beckwith *et al.* 1990; André & Montmerle 1994). Their coordinates (Herbig & Bell 1988) and relevant properties are listed in Table 1. The spectral index,  $\alpha_{2.2-12\mu}$ , was derived from K-band magnitudes (Cohen & Kuhi 1979; Ghez 1993) and IRAS 12  $\mu\text{m}$  fluxes (Weaver & Jones 1992). The Taurus sources are located at distances of 140 pc and AS 209, near the Ophiuchi dark cloud, is at 160 pc (Elias 1978; Whittet 1974; de Geus *et al.* 1990).

Observations of RY Tau, DO Tau, DL Tau, and AS 209 in the 1.3 mm continuum and CO ( $2 \rightarrow 1$ ) line at 230.5 GHz were made between 1993, November, and 1994, March, using the five-element Owens Valley millimeter-wave array. Baselines up to 60 m E–W and 60 m N–S provided an approximately circular  $3''$  beam (FWHM) with area  $\Omega_B = 6.9 \times 10^{-10}$  sr for DO Tau and  $\Omega_B = 6.1 \times 10^{-10}$  sr for AS 209. More extended baselines, up to 100 m E–W and N–S, yielded spatial resolution of up to  $2''$  for RY Tau and DL Tau with  $\Omega_B = 3.6 \times 10^{-10}$  sr and  $\Omega_B = 6.0 \times 10^{-10}$  sr, respectively. Cryogenically-cooled SIS receivers on each telescope produced over-all system temperatures of 500–1000 K. The digital correlator was configured to have two bands of Hanning smoothed channels,  $32 \times 1$  MHz and  $96 \times 83$  kHz, yielding spectral resolutions 1.30 and 0.11 km s $^{-1}$ , respectively. Each band was centered at  $V_{LSR} \sim 6.0\text{--}7.5$  km s $^{-1}$  for the various Taurus sources (cf. Hartmann *et al.* 1986) and  $V_{HEL} = -4.0$  km s $^{-1}$  for AS 209. Continuum measurements were made simultaneously in a broad-band channel of width 1 GHz. The quasar, 0528+134, was

observed at 20-minute intervals to calibrate visibility phases for DO, DL, and RY Tau; NRAO 530 was used for AS 209. The absolute flux density scale relied on measurements of Uranus and 3C 273. Maps made with the NRAO AIPS software package are centered on the positions in Table 1. Uncertainties in absolute fluxes and positions are estimated to be 15% and  $0''.5$ , respectively, for peak map intensities, but may as high as 50% and one beamwidth ( $\sim 2\text{--}3''$ ) for features detected at  $3\text{--}5\sigma$  levels.

### 3. Results

Maps of the integrated CO(2 $\rightarrow$ 1) emission for each source are displayed in Fig. 1a–4a. Values for the integrated intensity are given in column (2) of Table 2. For RY Tau and DL Tau, integration was over the velocity range for which emission appeared above a level of  $300\text{ mJy beam}^{-1}$  ( $2\sigma$ ) at spectral resolution  $0.45\text{ km s}^{-1}$ . Since the linewidths for AS 209 and DO Tau exceeded the velocity range of the high-velocity-resolution band, emission was integrated over the wider  $1.30\text{ km s}^{-1}$  channels in which peak map intensity was above a level of  $50\text{ mJy beam}^{-1}$  ( $2\sigma$ ). Individual velocity ranges are listed in column (3) of Table 2.

#### 3.1 Morphology, Optical Depth, and Mass

To obtain the first approximate determination of the radial extent of the circumstellar gas, elliptical Gaussians were fit to the images shown in Fig. 1a–4a and deconvolved with the restoring beams from the CLEAN deconvolution process. The radius at half-maximum intensity,  $R_{HM}$ , in column (4) of Table 2 is the FWHM major axis of this Gaussian fit. If the expected power-law profiles in temperature and density prevail, emission from the outer parts of disks is drastically reduced compared to the central regions. Thus,  $R_{HM}$  is a *lower limit* to the full radial extent of the gas. All of the sources appear to

be resolved in  $\sim 2\text{--}3''$  beams, with  $100 \text{ AU} \lesssim R_{HM} \lesssim 400 \text{ AU}$ . These dimensions are closer to those anticipated for the early solar nebula and approximately an order of magnitude lower than measured for other T Tauri stars.

As shown in KSB, an upper limit for the size of a central optically thick region in a circumstellar disk can be derived from the brightness temperature of peak emission at the central velocity. Assuming a typical disk temperature of  $T_k = 30 \text{ K}$  (Beckwith & Sargent 1993b), an inclination angle of  $45^\circ$ , and taking into account the average beam size ( $\sim 3''$ ), a limiting radius for an optically thick region for sources at distances of 150 pc is given by:

$$R_{\tau>1} = 80[B]^{1/2} \text{ AU}$$

where  $B$  is the line-center intensity of peak emission (in  $\text{Jy beam}^{-1}$ ) in a  $1.30 \text{ km s}^{-1}$  wide velocity bin. This is a conservative upper limit, since contributions from optically thin emission are ignored, and higher temperatures will significantly lower the result. For objects where a sizeable fraction of circumstellar mass is still infalling, the average temperature is likely to be higher.

In column (5) of Table 2,  $R_{\tau>1}$  is listed for each source, taking into account individual beam sizes and subtracting any contribution from the peak continuum flux (never more than 20%). A comparison with  $R_{HM}$  reflects the extent to which the emission is optically thin. As will be evident, this is important in estimating the circumstellar mass from the  $\text{CO}(2 \rightarrow 1)$  fluxes.

The total mass of circumstellar gas is calculated from the integrated  $\text{CO}(2 \rightarrow 1)$  emission,  $\int S_\nu dv$ , using (cf. Scoville et al. 1976):

$$M_{H_2} = 1.42 \times 10^{-10} \frac{(T_x + 0.93)}{e^{-16.76/T_x}} \frac{\tau_{CO}}{(1 - e^{-\tau})} \frac{D_{kpc}^2}{X(CO)} \int S_\nu dv \quad M_\odot.$$

Here,  $T_x$  is the excitation temperature,  $\tau_{CO}$  is the optical depth in the CO line,  $D_{kpc}$  is the distance to the source in kpc, and  $X(CO)$  is the fractional abundance of CO. For optically thin emission,

$$\frac{\tau_{CO}}{(1 - e^{-\tau})} \approx 1,$$

and the mass is calculable without *a priori* knowledge of  $\tau_{CO}$ . If the magnitude of the optical depth is completely unknown, the above approximation allows a determination of a lower limit to  $M_{H_2}$ .

Masses,  $M_{H_2}$ , were calculated for all the sources using the integrated line fluxes listed in Table 2, adopting an average value of  $T_x = 30\text{K}$ ,  $X(CO) = 10^{-4}$ ,  $D_{kpc} = 0.140$  kpc for the Taurus sources, and 0.160 kpc for AS 209. The results are shown in column (6) of Table 2. As optical depth decreases,  $M_{H_2}$  approaches the true mass;  $R_{\tau>1}/R_{HM}$  provides a qualitative assessment of the accuracy of the estimate. Table 2 shows that, for the entire sample, the upper limit to the radius of an optically thick region equals or exceeds the size of the present solar system (50 AU). Total masses much larger than the limits in Table 2, and still confined within a solar-system-sized radius, could easily be consistent with these observations. Indeed, masses typically 2 orders of magnitude larger are implied by 1.3 mm continuum measurements (Beckwith *et al.* 1990; André & Montmerle 1994). Further observations in optically thin isotopes of CO are needed to refine estimates of the mass of circumstellar gas.

### 3.2 Orientation and Kinematic Implications

The true orientation of the elongation axis of the circumstellar gas is critical to any kinematic analysis. In Fig. 1a–4a, emission is preferentially elongated along one axis, implying a flattened structure, but the degree of elongation is uncertain, since some objects are very small compared to the spatial resolution of the maps. However, the orientation of an associated optical jet can also be used to infer the PA of the sky-projected axis normal to the disk plane.

In Table 3, the PA of CO emission elongation,  $\Theta_{CO}$ , determined from the Gaussian fit, is listed for each of the sources. Uncertainties tabulated for  $\Theta_{CO}$  are derived from the goodness of fit and do not reflect the amplitude and position uncertainties in the map itself. Errors for RY Tau, where the deconvolved size is less than the FWHM size of the synthesized beam, are likely to be greater than the formal uncertainties listed in Table 3. In addition, irregular source features can confuse the determination of elongation orientation, especially when spatial resolution is limited. These features include azimuthal asymmetries in the circumstellar gas itself, molecular outflows, or ambient cloud material. Independent ways of constraining the orientation of the circumstellar material must be employed. In fact, the spatial distribution of the intensity-weighted average velocity (“first moment” with respect to velocity) can provide circumstantial evidence about disk orientation, if systematic velocity gradients are apparent. Maps of the first moment in velocity for the sample sources are shown in Fig. 1b–4b.

To a first approximation, rotation in a flattened disk will appear as a gradient in mean velocity parallel to the elongation axis of the disk. Outflow, or pure infall in the plane of a flattened envelope, on the other hand, will result in a gradient orthogonal

to this direction. Complex mean velocity structure will result from a combination of these effects, and the velocity gradient alone is not an unambiguous indicator of disk orientation, especially if the elongation itself is due to outflow. However, the molecular gas in maps shown here exhibits much smaller linewidths than typically observed for outflow sources (except for DO Tau) and is positioned close to the star. Therefore, for purposes of preliminary interpretation, the coincidence of a mean velocity gradient with the apparent CO elongation axis is considered as a significant clue to disk orientation. More detailed kinematic modeling like that carried out for GM Aur by KSB will be presented in a future paper. The PA of any systematic velocity gradient seen in Fig. 1b–4b,  $\Theta_{\delta V}$ , is given in column (3) of Table 3 for comparison with  $\Theta_{CO}$ . For RY Tau and AS 209,  $\Theta_{\delta V}$  was measured by taking the perpendicular to contours at the position of peak emission in Fig. 1a and 2a, since curvature was evident in the velocity gradient.

Other constraints on the orientation of a possible disk include the observation of an optical jet, presumably oriented perpendicular to the axis of elongation. The implied disk orientation is given as  $\Theta_{jet}$  in Table 3. Although  $\Theta_{jet}$  cannot be used alone to define orientation, agreement between several indicators strongly suggests the correct orientation of planar material. The PA of spatially averaged polarization vectors of scattered stellar light,  $\Theta_{pol}$ , is also a function of the orientation of the scattering material (Elsässer & Staude 1978; Bastien 1982). Monte Carlo models of light scattered by disks and envelopes indicate that  $\Theta_{pol}$  will be parallel to the elongation axis for flared disks and envelopes, but perpendicular to it for thin disks (Whitney & Hartmann 1992; 1993). In Table 3, the agreement between  $\Theta_{CO}$ ,  $\Theta_{\delta v}$ , and  $\Theta_{pol}$  is generally quite good and is supported by  $\Theta_{jet}$  in the case of DO Tau (Hirth *et al.* 1994). We claim this greatly strengthens an interpretation of the structures in Fig. 1–4 as thick disks or flattened envelopes. Thin

disks are ruled out, since  $\Theta_{pol}$  would be expected to differ from  $\Theta_{CO}$  and  $\Theta_{\delta v}$  by  $90^\circ$ .

### 3.3 Velocity Structure of Circumstellar Gas

Maps of the first moment with respect to velocity for RY Tau, DL Tau, DO Tau, and AS 209, in Fig. 1b–4b, display velocity gradients parallel to the elongation axis of CO emission and strongly suggest the existence of rotating disks. This evidence is not sufficient to uniquely establish the kinematic structure of the gas, however, since detailed information about the velocity distribution of emission has been lost by averaging. In Fig. 5, spatially integrated CO(2→1) spectra provide a complementary presentation of the data which retains velocity information. The emission intensity in spectral line maps was integrated over a  $5''$ -radius region centered on the stellar position and plotted as a function of velocity to effectively separate the signature of circumstellar gas from that of the large ambient cloud.

The spectra in Fig. 5 may be used to refine preliminary interpretations of the circumstellar CO emission. Two of the sources, RY Tau and AS 209, have lineshapes which are approximately double-peaked, consistent with Keplerian rotation (Beckwith & Sargent 1993b). This interpretation is somewhat uncertain for RY Tau, since the peak-to-trough amplitude is only a few times that of the greatest channel-to-channel variations due to noise. However, it is bolstered by available observations of the stellar radial velocity; the molecular spectrum is centered exactly on the optical value given by Hartmann et al. (1986),  $V_{LSR} = 7.8 \pm 2.4 \text{ km s}^{-1}$ . The line shape for AS 209 is double-peaked about a central velocity  $V_{HEL} \approx -7 \text{ km s}^{-1}$ , consistent with a stellar velocity of  $V_{HEL} = -6 \pm 3 \text{ km s}^{-1}$  (Walter, F. M., private communication). The spectral line for DL Tau is centered at  $V_{LSR} \approx 7 \text{ km s}^{-1}$ , a typical velocity for T Tauri stars in the Taurus-Auriga molecular

dark cloud (cf. Hartmann et al. 1986). The linewidth is surprisingly narrow ( $\Delta V = 1.2 \text{ km s}^{-1}$ ) such that any kinematic signature is likely to be masked by thermal broadening.

The remaining source, DO Tau, shows an extremely asymmetric spectral line that is an order of magnitude stronger than the others. Most of the molecular emission is blue-shifted more than  $10 \text{ km s}^{-1}$  relative to the velocity of the ambient molecular cloud and the best available stellar velocity ( $V_{LSR} = 10 \pm 5 \text{ km s}^{-1}$ ; Herbig 1977). The dominant blue-shifted wing is most likely due to a molecular outflow, especially since a blue-shifted optical jet is known to emanate from the central star (Hirth *et al.* 1994). Redward of the peak emission, however, there is a narrow feature separated from the blue-shifted emission by a prominent dip at the systemic velocity of the local molecular cloud. These features may be circumstellar in origin.

A “double-peaked” line shape arises from a centrifugally supported disk as a result of the  $1/\sqrt{r}$  dependence of the Keplerian velocity, the  $r^2$  dependence of emitting areas with the same velocity, and the truncation of the disk at some outer radius. Peaks in the spectra largely reflect the Keplerian velocities at the outer radius of the disk, modulated by  $\sin(i)$ , where  $i$  is the angle of inclination to the line of sight. However, other spatial and kinematic configurations can yield a similar lineshape. Line-center absorption by a cold intervening molecular cloud can introduce a central dip in the lineshape, especially in  $^{12}\text{CO}$  observations where the emission is likely to be optically thick, and infall in a *flattened* structure also produces a double-peaked line shape, since infall velocities are also proportional to  $1/\sqrt{r}$ .

Although none of the spectra in Fig. 5 can be taken as unambiguous proof of the presence of a Keplerian disk, they do provide a consistency check. The Keplerian velocity

at the outer radius,  $\sqrt{GM/R_d}$ , can be calculated assuming the central mass is approximately the stellar mass,  $M \approx M_*$ , listed in Table 1, and that the outer disk radius is given by  $R_d \approx R_{HM}$  of Table 2. The derived Keplerian velocity can then be modulated over a range of acceptable inclination angles and compared with double-peak velocities suggested by Fig. 5. If the center-to-peak velocity difference is significantly greater than that predicted for an edge-on Keplerian disk of the observed size, Keplerian rotation can be ruled out as the dominant cause of the spectral lineshape.

Table 4 lists Keplerian velocities calculated in this way and center-to-peak velocity differences estimated from the spectra in Fig. 5. For RY Tau and AS 209, this was simply taken to be half the distance between the channels with the highest amplitudes. For DL Tau, where the narrowness of the line obscures any double-peaked structure, a nominal value of one quarter of the total line width was chosen. In the case of DO Tau, the distance between the spectral dip and the narrow velocity feature was used. None of the velocities estimated from the spectra exceed those of the Keplerian approximation. In fact, since free fall velocities exceed the Keplerian velocity by a factor of  $\sqrt{2}$ , this is more consistent with a rotational rather than an infall interpretation. The range of inclination angles required for agreement with Keplerian rotation is given in Column 4. The values are plausible and, in the case of the best-resolved disk, consistent with an estimate based on the aspect ratio of the major and minor axes of CO emission.

### 3.4 Combined Evidence for Rotating Disks around Individual Stars

None of the independent diagnostics presented thus far demonstrate conclusively that the objects observed are protoplanetary disks. However, their combined effect provides useful criteria with which to test the simplest possible interpretations. These include

a rotating disk, infall envelope, bipolar outflow, or some combination of these; more complicated explanations can perhaps be constructed, but their consideration is unwarranted unless these hypotheses are ruled out. With these restrictions in mind, it is useful to briefly summarize the diagnostics for each individual star in order to assess the likelihood that observed CO emission arises from a protoplanetary disk.

*RY Tau:* Circumstellar CO emission around RY Tau is compact, located at the stellar position, and exhibits a symmetric double-peaked lineshape centered at the radial stellar velocity. The linewidth is consistent with that from a disk inclined  $25^\circ$  from face on with radius  $\sim 100$  AU and in Keplerian rotation about a central mass dominated by that of the star. Indicators of orientation – the PA’s of CO elongation, velocity gradient, and polarization – also support a rotating-disk interpretation and are aligned within  $20^\circ$  of each other. The observed values are somewhat uncertain due to the small radial extent of CO and variability of polarization, but their validity is confirmed by the orientation of a reflection nebulosity associated with RY Tau in R- and I-band coronagraphic images (Nakajima & Golimowski 1995). Furthermore, the compact size, lineshape, and linewidth argue that the CO is truly circumstellar and not part of an outflow. It must be conceded, however, that an element of infall is still possible, and a compact low-velocity outflow is not strictly excluded.

*DL Tau:* Orientation indicators for circumstellar gas around DL Tau are in exact agreement and strongly support the idea that CO emission represents a 250-AU-radius circumstellar disk. The spectral line is centered at a typical velocity for T Tauri stars in the Taurus-Auriga molecular dark cloud and is surprisingly narrow. A nearly face-on disk orientation is required to match it, given the stellar mass estimate of  $0.56 M_\odot$ . However,

pre-main sequence stellar mass estimates are most uncertain for stars of very low mass (cf. d’Antona & Mazzitelli 1993), and DL Tau is the least massive of the sample. It is also possible that the observed linewidth has been partially narrowed by absorption from the ambient molecular cloud. In either case, the narrow linewidth and compact morphology argue against an outflow interpretation and the orientation diagnostics overwhelmingly support a disk-like configuration. If this is truly the case, infall is strictly excluded by the orientation of the velocity gradient with respect to that of the elongation axis. In these respects, the rotating disk interpretation is robust for DL Tau. To confirm this, imaging at higher spatial resolution is needed to resolve velocity-dependent structure for comparison to a kinematic model.

*DO Tau:* The orientation of circumstellar gas around DO Tau is well established by the morphology of CO emission and observations of its optical jet and polarization. However, a single systematic velocity gradient is not apparent in the first moment velocity map. The lineshape is much broader than expected for rotation and asymmetrically blue-shifted to velocities higher than would be expected for Keplerian rotation in a disk that size. Its association with an asymmetric optical jet virtually establishes an outflow origin for the high-velocity molecular emission. In contrast, a disk-like appearance and agreement between three orientation indicators argue for a large circumstellar component of rotating and/or infalling gas. Spectral features near the systemic velocity are consistent with Keplerian rotation, but detailed study of the emission morphology at these velocities is required to identify any rotating disk. Clearly, some combination of hypotheses is needed to explain the morphology and kinematics of DO Tau’s CO emission. To obtain the correct interpretation, modeling of emission in individual spectral line maps is being carried out and will be presented in a future paper.

*AS 209*: The map of CO emission reveals an elliptical shape with somewhat tenuous extensions perpendicular to the elongation axis. The dominant velocity gradient and polarization vector are parallel to this axis, but the velocity gradient is not as clear as for DL Tau and RY Tau. The lineshape is double-peaked about a central velocity that agrees approximately with estimates of the stellar velocity, and the linewidth is consistent with a Keplerian disk inclined  $\sim 25^\circ$  from edge on. Taken together, these facts suggest that the emission is dominated by a rotating disk. However, the perpendicular extensions, irregular velocity gradient, and a slight asymmetry in the lineshape could easily be accounted for by an additional outflow component to the molecular emission. As for DO Tau, modeling of high spatial- and velocity-resolution spectral line maps is required to confirm this picture and will be the subject of a future paper.

#### 4. Summary and Discussion

Four T Tauri stars, RY Tau, DL Tau, DO Tau, and AS 209, have been imaged in the CO(2 $\rightarrow$ 1) line at  $\sim 2\text{--}3''$  resolution. In all cases, emission is resolved with deconvolved half-maximum radii of 100 to 350 AU. The PA's of optical linear polarization and, in one case, an optical jet show remarkable agreement with disk orientation implied by the CO elongation and by velocity gradients. From the peak strength of line emission, upper limits on the size of any optically thick inner region are found to be 50–150 AU, the approximate size of the solar system. Masses of circumstellar H<sub>2</sub> range over 2 orders of magnitude, from  $1.3 \times 10^{-6}$  to  $1.4 \times 10^{-4} M_\odot$ , but these values are lower limits and much less than masses estimated from the dust continuum emission (Beckwith *et al.* 1990; André & Montmerle 1994). Except for DO Tau, all sources show evidence of a velocity gradient along the elongation axis and spectral lineshapes and linewidths which are consistent with

Keplerian rotation about the estimated stellar mass.

The strongest case for a centrifugally supported disk can be made for circumstellar gas around DL Tau. With RY Tau, it shows the most distinct velocity gradient in first moment velocity maps. For both AS 209 and RY Tau, however, the uncertainty in orientation admits a possible infall or outflow component to circumstellar gas. Nevertheless, the PA of elongation is not likely to be off by as much as  $45^\circ$ , indicating that the *dominant* velocity component of the molecular gas is rotational. Spectral lineshapes for all three objects are also consistent with that of a Keplerian disk and strengthen the case for a rotational interpretation of the velocity gradients observed in moment maps. If the possible kinematic interpretations are limited to rotation, outflow, or infall, evidence presented here weighs heavily in favor of predominantly rotation of circumstellar gas for all three of these sources.

For DO Tau, the first moment velocity map and associated spectrum eliminates the possibility that *all* of the observed emission radiates from rotating circumstellar gas. An outflow interpretation for some of the gas is supported by the presence of an optical jet (Hirth *et al.* 1994). It is somewhat surprising, therefore, that the morphology of molecular emission in Fig. 3 resembles a disk with orientation perpendicular to the jet. The disk-like morphology and complex velocity structure may indicate the presence of a radially infalling component to the circumstellar velocity structure (e.g. Hayashi *et al.* 1993; for theoretical justification of a *flattened* infall region, see Galli & Shu 1993 and Stahler *et al.* 1994). Molecular material around DO Tau, with  $R_{HM} = 350$  AU, is substantially more extended than that around the other sources, but is still more compact than for many earlier objects such as HL Tau. It may represent an intermediate evolutionary state in which much of the circumstellar material is centrifugally supported, but some is still in

a state of infall. Modeling of high-spatial-resolution spectral line maps of DO Tau with a kinematic model of infall and rotation is required to identify such a kinematic state (Koerner 1994) and will be presented in a future paper.

In the above respects, circumstellar CO(2→1) emission from the T Tauri stars surveyed shows characteristics which confirm the hypothesis that these stars are surrounded by protoplanetary disks like that which gave rise to our solar system. Further work is needed to secure this interpretation. Kinematic modeling of spectral line maps of  $^{13}\text{CO}(2\rightarrow 1)$  emission from GM Aurigae provided satisfactory evidence for Keplerian rotation of circumstellar gas by duplicating systematic patterns of emission seen in maps of the disk at discrete velocities (KSB). However, the circumstellar gas structures described here are much smaller, and similar analyses will require imaging at correspondingly higher spatial resolution.

Despite the inability to rigorously identify kinematic patterns, this work confirms the presence of molecular gas in the circumstellar environment of T Tauri stars for which dust continuum emission is detected at millimeter wavelengths. The radial extent of the gas is more in accord with the dimensions of the solar system than that of previously imaged T Tauri stars; molecular line emission is resolved in 2–3'' FWHM beams (corresponding to 300–450 AU) and optically thick beyond a radial distance corresponding to the semi-major axis of Pluto's orbit ( $\sim 50$  AU). Some evidence of centrifugal support is present for most of the sample. These facts strongly support the hypothesis that a large fraction of T Tauri stars are surrounded by protoplanetary disks similar to the early solar nebula.

## 5. Acknowledgments

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Table I. Star Properties

(1)	(2)	(3)	(4)	(5)	(6)	(7)
Source	$\alpha$	$\delta$	$F_{1.3mm}$	$\alpha_{IR}$	$M_*$	Age
	(1950)	(1950)	(mJy)	(2.2-12 $\mu$ m)	( $M_\odot$ )	( $\times 10^6$ yr)
RY Tauri	04:18:50.80	+28 19 35.0	229 <sup>a</sup>	0.05	1.69 <sup>a</sup>	0.21 <sup>a</sup>
DL Tauri	04:30:36.02	+25 14 24.0	230 <sup>a</sup>	-0.37	0.56 <sup>a</sup>	1.20 <sup>a</sup>
DO Tauri	04:35:24.18	+26 04 55.2	136 <sup>a</sup>	-0.33	0.72 <sup>a</sup>	0.60 <sup>a</sup>
AS 209	16:46:25.24	-14 16 56.5	300 <sup>b</sup>	-0.60	—	—

a) Beckwith *et al.* (1990)

b) André & Montmerle (1994)

**Table II. Global Properties Derived from CO Emission.**

(1)	(2)	(3)	(4)	(5)	(6)
Source	$\int S_\nu dv$	$\Delta V$	$R_{HM}$	$R_{\tau>1}$	$M_{H_2}$
	(Jy km/s)	(km/s)	(AU)	(AU)	( $\times 10^{-6} M_\odot$ )
RY Tauri	6.6	6.9	107	52	9.6
DL Tauri	0.9	1.2	250	53	1.3
DO Tauri	100.2	13.6	350	164	145
AS 209	7.4	10.4	287	115	10.7

**Table III. Diagnostics of Disk Orientation.**

(1)	(2)	(3)	(4)	(5)
Source	$\Theta_{CO}$	$\Theta_{\delta V}$	$\Theta_{pol}$	$\Theta_{jet} + 90^\circ$
RY Tauri	$48^\circ \pm 5^\circ$	$21^\circ$	$20^\circ \pm 20^\circ_a$	—
DL Tauri	$84^\circ \pm 5^\circ$	$84^\circ$	$85^\circ \pm 7^\circ_b$	—
DO Tauri	$160^\circ \pm 5^\circ$	—	$170^\circ \pm 1^\circ_a$	$160^\circ_c$
AS 209	$138^\circ \pm 2^\circ$	$120^\circ$	$134^\circ \pm 10^\circ_a$	—

a) Bastien (1982)

b) Tamura & Sato (1989)

c) Hirth *et al.* (1994)

**Table IV. Derived Keplerian Disk Parameters**

(1)	(2)	(3)	(4)
Source	$\Delta V_{spectra}$	$\Delta V_{Kepler}$	$i$
	(km/s)	(km/s)	(degrees)
RY Tauri	1.5	3.7	25
DL Tauri	0.3	1.4	12
DO Tauri	1.0	1.4	45
AS 209	1.5	1.8	56

Figure 1: a) A map of CO(2→1) emission toward RY Tau integrated over the velocity range,  $V_{lsr} = 4.2$  to  $11.1$  km s<sup>-1</sup>. The contour interval is 400 mJy beam<sup>-1</sup> km s<sup>-1</sup> ( $1\sigma$ ). Contours start at the  $2\sigma$  level. Peak emission is 3.9 Jy beam<sup>-1</sup> km s<sup>-1</sup> with an integrated intensity of 7.3 Jy beam<sup>-1</sup> km s<sup>-1</sup>. The synthesized beam,  $2''.6 \times 1''.9$  FWHM at PA  $-54.27^\circ$  with area  $\Omega_B = 3.6 \times 10^{-10}$  sr, is displayed as a hatched region. The apparent elongation to the SE is largely a beam effect. The FWHM deconvolved size of a Gaussian fit to the emission is  $2''.3 \times 1''.9$  oriented at PA  $48.6^\circ$ , but this value is highly uncertain as an indicator of source size and orientation, since it is of order the beam size. b) A map of the mean velocity variations across the structure seen in a). Grey-scale values correspond to the velocity scale shown to the right. The contour interval is 0.5 km s<sup>-1</sup>. A systematic velocity gradient is evident along the axis of elongation. At the position of peak emission, the gradient is directed along PA  $\sim 21^\circ$ .

Figure 2: a) A map of CO(2→1) emission toward DL Tau integrated over the velocity range,  $V_{lsr} = 6.5$  to  $7.7$  km s<sup>-1</sup>. The contour interval is 73 mJy beam<sup>-1</sup> km s<sup>-1</sup> ( $1\sigma$ ). Contours start at the  $2\sigma$  level. Peak emission is 594 mJy beam<sup>-1</sup> km s<sup>-1</sup> with an integrated intensity of 1.1 Jy beam<sup>-1</sup> km s<sup>-1</sup>. The synthesized beam,  $2''.9 \times 2''.8$  at PA  $-85^\circ$  with area  $\Omega_B = 6.0 \times 10^{-10}$  sr, is displayed as a hatched region. The FWHM deconvolved size of a Gaussian fit to the emission is  $3''.5 \times 2''.5$  oriented at PA  $84^\circ$ , b) A map of the mean velocity variations across the structure seen in Fig. 2a. Grey-scale values correspond to the velocity scale at the top. The contour interval is 0.1 km s<sup>-1</sup>. A gradient is evident along the axis of elongation, at PA  $\sim 85^\circ$ .

Figure 3: a) A map of CO(2→1) emission toward DO Tau integrated over the velocity range,  $V_{lsr} = -4.8$  to  $8.2$  km s<sup>-1</sup>. The contour interval is 2.0 Jy beam<sup>-1</sup> km s<sup>-1</sup> ( $2\sigma$ ). Contours start at the  $2\sigma$  level. Peak emission is 34.5 Jy beam<sup>-1</sup> km s<sup>-1</sup> with an integrated intensity of 100.2 Jy beam<sup>-1</sup> km s<sup>-1</sup>. The synthesized beam,  $3''.2 \times 2''.9$  FWHM at PA  $-54.27^\circ$  with area  $\Omega_B = 6.9 \times 10^{-10}$  sr, is displayed as a hatched region. The FWHM deconvolved size of a Gaussian fit to the emission is  $5''.1 \times 3''.9$  oriented at PA  $160^\circ$ . The aspect ratio of the major and minor axes implies an inclination angle of  $40^\circ$  for a circularly symmetric disk. b) A map of the mean velocity variations across the structure seen in a). Grey-scale values correspond to the velocity scale shown to the right. The contour interval is 0.5 km s<sup>-1</sup>. A systematic velocity gradient is not easily identified. This may reflect a combination of kinematic effects, such as infall, rotation, and outflow.

Figure 4: a) A map of CO(2→1) emission toward AS 209 integrated over the velocity range,  $V_{hel} = -11.1$  to  $-2.1$  km s<sup>-1</sup>. The contour interval is 350 mJy beam<sup>-1</sup> km s<sup>-1</sup>,  $1\sigma$ . Contours start at the  $2\sigma$  level. Peak emission is 3.9 Jy beam<sup>-1</sup> km s<sup>-1</sup> with an integrated intensity of 7.4 Jy beam<sup>-1</sup> km s<sup>-1</sup>. The synthesized beam,  $3''.6 \times 2''.3$  FWHM at PA  $-18.6^\circ$  with area  $\Omega_B = 6.1 \times 10^{-10}$  sr, is displayed as a hatched region. The FWHM deconvolved size of a Gaussian fit to the emission is  $3''.1 \times 2''.2$  oriented at PA  $137.9^\circ$ . Low level contours are also elongated nearly perpendicular to this axis, as might be expected for CO entrained in an outflow.

b) A map of the mean velocity variations across the structure seen in a). Grey-scale values correspond to the velocity scale shown to the right. The contour interval is  $0.25 \text{ km s}^{-1}$ . There is some indication of a velocity gradient along the CO elongation axis. Deviations from this trend may be due to outflow or infall components.

Figure 5: CO(2→1) spectra from the T Tauri stars, RY Tau, DL Tau, DO Tau, and AS 209. The integrated intensity from circumstellar emission in high-spatial-resolution spectral line maps is plotted as a function of  $V_{LSR}$  ( $V_{HEL}$  for AS 209). Variations in velocity resolution are the result of limited bandwidth coverage for the high velocity resolution channels. Molecular lineshapes are uniformly plotted over a width of  $15 \text{ km s}^{-1}$  for easy comparison. The continuum flux level, evident as a perfectly flat line for DL Tau, is shown outside the region of spectral coverage. The flux scale, in Jy, is not uniform between spectra, but ranges over an order of magnitude to facilitate comparison of lineshapes.

**Figure 1**

**Figure 2**

**Figure 3**

**Figure 4**

**Figure 5**