



DISK DETECTIVE: DISCOVERY OF NEW CIRCUMSTELLAR DISK CANDIDATES THROUGH CITIZEN SCIENCE

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ABSTRACT

The Disk Detective citizen science project aims to find new stars with 22 μm excess emission from circumstellar dust using data from NASA's *Wide-field Infrared Survey Explorer (WISE)* mission. Initial cuts on the AllWISE catalog provide an input catalog of 277,686 sources. Volunteers then view images of each source online in 10 different bands to identify false positives (galaxies, interstellar matter, image artifacts, etc.). Sources that survive this online vetting are followed up with spectroscopy on the FLWO Tillinghast telescope. This approach should allow us to unleash the full potential of *WISE* for finding new debris disks and protoplanetary disks. We announce a first list of 37 new disk candidates discovered by the project, and we describe our vetting and follow-up process. One of these systems appears to contain the first debris disk discovered around a star with a white dwarf companion: HD 74389. We also report four newly discovered classical Be stars (HD 6612, HD 7406, HD 164137, and HD 218546) and a new detection of 22 μm excess around the previously known debris disk host star HD 22128.

Key words: catalogs – infrared: planetary systems – methods: data analysis – protoplanetary disks – stars: emission-line, Be – surveys

1. INTRODUCTION

All-sky mid-infrared surveys have revolutionized the science of planet formation by discovering populations of young stars and main-sequence stars with excess infrared radiation indicating the presence of dusty circumstellar disks. These disks, which include gas-rich protoplanetary disks around Young Stellar Objects (YSOs) and dusty debris disks around main-sequence stars, serve as the signposts of planet formation (e.g., Kenyon & Bromley 2002). They inform us about the timescales and the environment of planet formation (e.g., Rieke et al. 2005; Hernández et al. 2007; Jang-Condell et al. 2015), and the present-day locations and dynamics of planets (e.g., Thalmann et al. 2010; Muto et al. 2012; Quanz et al. 2013; Currie et al. 2015; Nesvold & Kuchner 2015).

The *IRAS* all-sky survey discovered the first extrasolar debris disks (Aumann et al. 1984) and provided a large sample of debris disks (e.g., Rhee et al. 2007). After *IRAS*, AKARI surveyed the whole sky at 9 and 18 μm with ~ 7 times better sensitivity than *IRAS*, finding many more new disks (Ishihara et al. 2010). Some disk discoveries have come from pointed studies, like the *Spitzer* Formation and Evolution of Planetary Systems (FEPS) survey (Carpenter et al. 2009). But many of the best-studied, most informative disks (like TW Hydra, Fomalhaut, etc.) are relatively isolated on the sky, requiring an all-sky survey to find them.

NASA's *Wide-field Infrared Survey Explorer (WISE)* is the most recent and sensitive all-sky mid-infrared survey (Wright et al. 2010), with a further factor of ~ 80 gain in sensitivity over AKARI in the mid-IR. Using a 16 inch mirror in a Sun-synchronous polar orbit, *WISE* scanned the sky at 3.4, 4.6, 12, and 22 μm (bands *W1*, *W2*, *W3*, and *W4*, respectively). The

¹⁵ Disk Detective.

¹⁶ <https://www.diskdetective.org/#/authors>.

WISE cryogenic mission, launched in 2009, lasted a little over 10 months and was followed by the first post-cryogenic mission, NEOWISE. The AllWISE catalog¹⁷ combines data from both phases, making it the most comprehensive mid-infrared multi-epoch view of the sky available today.

Previous infrared surveys for debris disks have provided target lists for exoplanet searches via direct imaging (Apai et al. 2008; Janson et al. 2013; Wahhaj et al. 2013; Meshkat et al. 2015). Debris disks found with *WISE* should provide crucial targets for upcoming generations of exoplanet searches. *WISE* could detect debris disks around main-sequence A stars to a distance of 300 pc and protoplanetary disks around T Tauri stars to 1 kpc.

Indeed, many teams have used the *WISE* data to find new debris disks, searching a vast catalog of >747 million *WISE* sources. McDonald et al. (2012) cross-correlated the *WISE* source list with the *Hipparcos* catalog, finding over 86,000 stars with suspected infrared excesses. Kennedy & Wyatt (2013), Wu et al. (2013), and Patel et al. (2014) performed more careful searches for debris disks in the *WISE* source list using the *Hipparcos* catalog and found 6, 70, and 108 new debris disk candidates, respectively. Other specific surveys for debris disks have focused on stars with ages determined from chromospheric activity Vican & Schneider (2014), white dwarfs (Debes et al. 2011, 2012), M dwarfs (Avenhaus et al. 2012; O'Donnell et al. 2013), G-K dwarfs (Cruz-Saenz de Miera et al. 2014), *Kepler* candidate exoplanet systems (Kennedy & Wyatt 2012; Lawler & Gladman 2012; Ribas et al. 2012), and other exoplanet catalogs (Morales et al. 2012).

Likewise, the *WISE* data on young clusters and star-forming regions have attracted much attention. Rebull et al. (2011), Esplin et al. (2014), and Liu et al. (2014) scoured the Taurus-Auriga Region. Koenig et al. (2012) searched 11 outer Galaxy massive star-forming regions and three open clusters. Other studies have examined smaller regions, like the Western Circinus molecular cloud (Liu et al. 2011), the young open cluster IC 1805 (Straizys et al. 2013), the H II region S155 (Huang et al. 2014), the Sco-Cen and η Cha associations (Luhman & Mamajek 2012; Rizzuto et al. 2012), nearby moving groups of young stars (Simon et al. 2012), and λ and σ Orionis (Koenig et al. 2015). Still others have attempted to take in the whole sky, using color cuts (Majaess 2013) or cross correlating with *IRAS* (Lundquist et al. 2014). Many of these searches were based on preliminary data releases with less sensitivity than the AllWISE release, but they have already uncovered thousands of candidate Class I, II, and III YSOs and transitional disks, helping fill in our picture of the timing and progression of star formation.

Unfortunately, because of its limited spatial resolution (12 arcsec at 22 μ m) contamination and confusion limit every search for disks with *WISE* (e.g., Kennedy & Wyatt 2012). Contamination sources include unresolved companion stars and other stars nearby on the sky, background galaxies, Galactic cirrus, and even asteroids and airplanes. For this reason, most recent searches include visual inspection of the *WISE* images (see e.g., Debes et al. 2011; Wu et al. 2013; Cruz-Saenz de Miera et al. 2014; Patel et al. 2014). Computer cuts alone can provide a first stage of vetting, but they generate catalogs riddled with false positives (Kennedy & Wyatt 2012). Color cuts and source quality flags can help (e.g., Koenig et al. 2012;

Majaess 2013; Davenport et al. 2014), but the color loci of disk candidates overlaps with the color loci of blended background galaxies (Koenig et al. 2012) and peaks in the Galactic dust emission. Kennedy & Wyatt (2012) used the *IRAS* 100 μ m level to discard many false-positive disk candidates contaminated with Galactic dust emission, but using this method prohibits searching many interesting star-forming regions.

Because of these challenges, many disks remain to be found with *WISE* data, even after all the efforts described above. The largest published study of debris disks (Wu et al. 2013) and the still larger *WISE* science team disk study (D. L. Padgett et al. 2016, in preparation) are based on the *Hipparcos* and Tycho catalogs. These catalogs are magnitude limited in the *V* band, so they omit a vast population of redder, late-type stars.¹⁸

Moreover, a vast solid angle in young clusters and star-forming regions remains to be properly searched with *WISE*—each candidate examined by eye and followed up with spectroscopy and higher resolution imaging. When Majaess (2013) ran the all-sky data through a novel color filter to search for YSO candidates (without vetting the candidates by eye), he found a total of ~10,000 objects of interest; the *WISE* studies of young clusters and star-forming regions described above (which mostly included visual vetting) yielded a total of ~4000 disk candidates. The difference between these numbers provides a minimal measure of what remains for us to study with *WISE*: \gtrsim 6000 objects with colors consistent with YSOs that have not yet been visually inspected.

Here we describe a new project to scour the *WISE* data for new debris disks and YSOs. The Disk Detective citizen science/crowdsourcing project classifies *WISE* sources via a website, diskdetective.org, where volunteers examine images from *WISE*, the Two Micron All Sky Survey (2MASS), the Digitized Sky Survey (DSS), and, when available, the Sloan Digital Sky Survey (SDSS) to check them for false positives. This approach should allow us to unleash the full potential of *WISE* for finding new disks, probing the cooler stars and isolated objects missed by previous debris disk searches, a catalog eight times the size of the large Wu et al. (2013) survey. We describe the online vetting process in Section 2, our small-telescope follow-up program in Section 3, and we present our first list of 37 disk candidates in Section 3.1.

2. CITIZEN SCIENCE APPROACH

Disk Detective is a new addition to the successful Zooniverse network of Citizen Science Alliance projects (Lintott et al. 2008). Visitors to the site (“users”) view “flipbooks” showing several images of the same source at different wavelengths. Figure 1 shows a sample screenshot from <https://www.DiskDetective.org> illustrating one frame in a flipbook.

After they view the flipbooks, users answer a question, “What best describes the object you see?,” by clicking on one or more of six buttons. The site then records the user’s choice (s) for interpretation by the Disk Detective science team and offers the user another source to classify. This approach is borrowed from another Zooniverse Project, Snapshot Serengeti (<http://www.snapshotserengeti.org>). Snapshot Serengeti shows users flipbooks of wildlife photographs, asking users to identify animal and bird species that appear in the images,

¹⁷ Available at <http://irsa.ipac.caltech.edu/Missions/wise.html>

¹⁸ The initial candidates listed in this paper are also all in *Hipparcos*, but our full data set does not cross-correlate with any stellar catalogs.

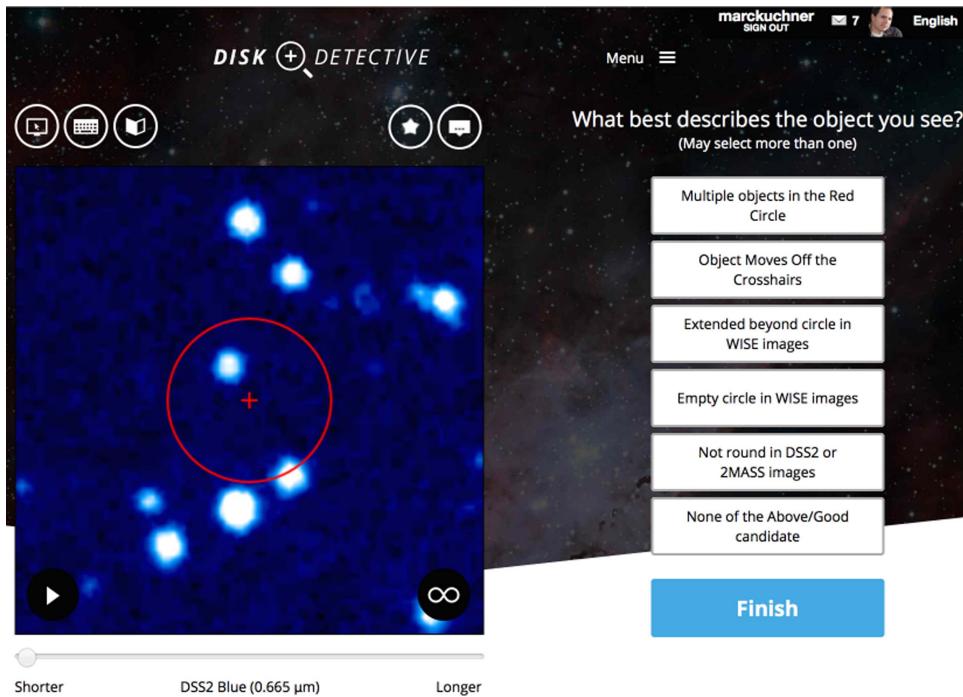


Figure 1. Screenshot from DiskDetective.org showing one frame from the flipbook on the left and the classification buttons on the right. This example clearly has multiple objects visible in the DSS2 Blue band that fall within the *WISE* 4 beam (the red circle); the Web site makes it easy for participants to discard this potential false positive by clicking on the appropriate button.

Table 1
Initial Cuts on the AllWISE Source Catalog

Criterion	Meaning
$w1\text{mpro} > 3.5$	<i>WISE</i> 1 magnitude >3.5
$w4\text{mpro} < (w1\text{mpro} - 0.25)$	<i>WISE</i> 4 excess of 0.25 magnitudes over <i>W1</i>
$w1\text{mpro} > 5 * \sqrt{w1\text{sigmpro} * w1\text{sigmpro} + w4\text{sigmpro} * w4\text{sigmpro}} + w4\text{mpro}$	The <i>WISE</i> 4 excess is significant at the $5-\sigma$ level.
$w4\text{sigmpro}$ is not null and $w4\text{rchi2} < 1.3$	<i>WISE</i> 4 profile-fitting yielded a fit with $\chi^2 < 1.3$
$w4\text{snr} >= 10$	<i>WISE</i> 1 profile-fit signal-to-noise ratio >10
$w4\text{nm} > 5$	Source detected at <i>WISE</i> 4 in at least 5 individual 8.8 s exposures with SNR > 3
$na = 0$ and $nb = 1$	The profile-fitting did not require active deblending.
$n_2mass = 1$	One and only one 2MASS PSC entries found within a $3''$ radius of the <i>WISE</i> 1 source position.
$cc_flags[1]$ not matches “[DHOP]”	No diffraction spike, persistence, halo, or optical ghost issues at <i>WISE</i> 1 or <i>WISE</i> 4.
$cc_flags[4]$ not matches “[DHOP]”	No 2MASS XSC source $<30''$ from the <i>WISE</i> source.
$xscprox$ is null or $xscprox > 30$	Photometry not contaminated by known 2MASS extended sources.
$ext_flg = 0$	

taking advantage of the human eye’s ability to spot moving objects camouflaged by noise (Swanson et al. 2015). At Disk Detective, instead of identifying animals in the grasses of the Serengeti, users identify clean point sources in a forest of astrophysical and instrumental contaminants: galaxies, ISM, artifacts, etc.

The flipbooks in Disk Detective generally consist of 10 images of each source: images in four *WISE* bands, three bands from the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006, 2MASS) and three bands from the Digital Sky Survey (DSS; Djorgovski et al. 1998). When possible, we use images from the SDSS seventh data release instead of DSS; these SDSS data cover about one-fourth of the sky (see Gunn et al. 2006; Abazajian et al. 2009). The images are independently scaled using IDL color Table 1 (BLUE/

WHITE), matching the color scale to the full range of the data. Hence, the flipbook approach works better than attempting to show all these bands in a single multi-color image, which would tend to be dominated by one or two bands. Overlaid on every flipbook is a circle with a radius of 10.5 arcsec, roughly the area we must ensure is free of contamination before we can trust the photometry. (The circle’s radius was chosen conservatively; the nominal resolution of *WISE* at $22\ \mu\text{m}$ is 12 arcsec FWHM.) Also, overlaid on each image is a cross marking the center of the image at the *WISE* $3.4\ \mu\text{m}$ band (*W1*) and indicating the expected uncertainty in position for an uncontaminated point source. The disk candidates are unresolved by *WISE*, so our search image is a tightly concentrated red object, possibly with diffraction spikes, that does not shift position from band to band.

Users view the flipbook by clicking on an arrow-shaped “play” button on the lower left of the screen, shown in Figure 1. They also have the option to scroll through the flipbook frame by frame using the scrub bar beneath it. The frames are each labeled with the wavelength and the name of the survey that yielded the image, e.g., “2MASS K (2.16 μm).” After a user has viewed the whole flipbook, he or she can then choose from among six classification buttons to click on, labeled “Multiple objects in the Red Circle,” “Object Moves off the Crosshairs,” “Extended beyond circle in *WISE* Images,” “Empty Circle in *WISE* images,” “Not Round in DSS2 or 2MASS images,” and “None of the Above/Good Candidate.” With the exception of the “None of the Above” option, the user can choose more than one description per flipbook. After at least one of these classification buttons is chosen, a button labeled “Finish” becomes active; clicking this button records the user’s choices and causes the next flipbook to appear.

Since its launch, Disk Detective has attracted a vast user community. Roughly 1.5 million classifications have been performed so far by roughly 28,000 volunteers. Roughly half of the classifications come from an enthusiastic group of “superusers.” Fifteen superusers have already classified $>10,000$ sources each; seven have classified $>30,000$. The superusers started their own email discussion group via Google groups to work together on Disk Detective issues. They now help train other users and form a crucial extension of the Disk Detective science team (see below).

Communication with Disk Detective users is aided by the new Zooniverse translation crowdsourcing tool. Using this online tool, volunteers have translated the site into Spanish, French, Russian, German, Hungarian, Polish, Bahasa, Romanian, Portuguese, Japanese, and Chinese (traditional and simplified characters); the translated sites are accessible via a link in the upper right corner of site.

The DiskDetective.org site is tied into the “TALK” social network common to Zooniverse sites. TALK has a page for each subject on Disk Detective that provides a simple Spectral Energy Distribution (SED) for the object composed of the 2MASS and AllWISE photometry and also a link to the SIMBAD page on the source if one exists. On TALK, users can create and maintain collections of their favorite subjects by clicking on a button labeled “collect.”

2.1. Pre-selection of *WISE* Sources

To choose sources to upload to the website, we performed some initial computer-based cuts on the *WISE* data, informed by published debris disk searches (see above) but not limited by the *Hipparcos* or Tycho catalogs, for example. We utilized signal-to-noise cuts (`w4snr`, `w1sigmpo`, `w4rchi2`, and `w4sigmpo`) and some of the AllWISE catalog flags (`cc_flags`, `xscprox`, `na`, `nb`, `n_2mass`, and `ext_flg`) to remove sources that were noisy or close to known extended sources. Though many searches have used sophisticated color cuts to focus on particular kinds of disks (e.g., Koenig et al. 2012), we kept our color cuts minimal to cast as broad a net as possible. To preselect sources with infrared excesses, we merely removed all sources with $[W4] > [W1]-0.25$. ($[W1]$ is magnitude in the *WISE* 3.4 μm band, and $[W4]$ is magnitude in the *WISE* 22 μm band.) This criterion corresponds to a 26% excess over a Rayleigh–Jeans slope between those two bands. We also required that a *WISE* 4 excess be significant at the 5- σ level. Table 1 summarizes all these initial cuts.

Table 2
Classification Buttons, and the Distribution of Raw Classifications
as of 2015 August 25

Classification Button	Percentage of Times Selected	
	$J < 14.5$	$J > 14.5$
Multiple objects in the Red Circle	40.5	6.5
Object Moves Off the Crosshairs	6.5	6.9
Extended beyond circle in <i>WISE</i> Images	23.4	15.9
Empty circle in <i>WISE</i> images	0.6	0.4
Not round in DSS2 or 2MASS images	9.0	37.6
None of the Above/Good Candidate	20.0	33.0

Note. The distribution is broken down by J magnitude of the sources. The bright objects were more frequently classified as “Multiple Objects in the Red Circle”; the fainter ones as “Not Round in DSS2 or 2MASS Images.” We deactivated the $J > 14.5$ objects on 2014 May 30.

We launched the site on 2014 January 28 with a first batch of subjects covering only the Galactic latitudes $+30$ to $+40$, $+50$ to $+90$, and -40 to -90 , with no additional magnitude limit. At first, about 20% of the initial upload was made available on the site. We soon realized that most of the volunteer effort was being spent classifying faint, extragalactic sources, so we decided to impose a magnitude limit on the search. We chose a criterion of $J < 14.5$ because the subjects brighter than this magnitude were clearly concentrated in the Galactic plane, while fainter subjects appeared to be isotropically distributed on the sky. We uploaded a second batch of 272,022 subjects on 2014 May 30. This second batch covered the rest of the sky, but was limited to $J < 14.5$. We also deactivated all the previously uploaded and active $J > 14.5$ subjects on this date, so presently the $J > 14.5$ subjects currently make up a small subset of the data for which we have classifications.

2.2. Classification Data

We examine the classification data weekly to chart our progress. As of 2015 August 25, the selection of classification buttons had the distribution shown in Table 2. Clearly, the dominant false-positive rejected by the classification process is “Multiple objects in the Red Circle.”

Note that the “Empty Circle in *WISE* images” button exists mainly to allow users a reasonable response option if, for example, there were a network glitch. None of the data actually had empty circles in all the *WISE* images, though some subjects did have empty red circles in a single band, and some users chose this “Empty Circle in *WISE* images” classification for such subjects. The rarity of this situation is reflected in the very low (0.6%) classification rate for this button.

So far, we have investigated two basic algorithms for sorting the raw data into classifications: a plurality algorithm and majority algorithm. By the plurality algorithm, the classification with the most votes becomes the official classification. By the majority algorithm, the official classification is one with $\geq 50\%$ of the votes, if one exists.

To help test these algorithms, we prepared a “gold standard” data set of 500 subjects classified by the members of the science team. All of the subjects in this “gold standard” data set were classified by two or more science team members, and all tie votes were discarded. This process showed that science team members agree with one another at roughly the 82% level

as to whether an object is “good” or not. A challenge in creating this “gold standard” set was that sometimes science team members ruled out candidates based on their expertise/ outside information rather than strictly doing what the website asked for. There was also a range in terms of how conservative the science team members were in terms of crowded fields and ISM background. For example, the YSO experts (used to looking in the Galactic plane) on the science team tolerated background objects and “extended” objects more readily than the debris disk experts on the team.

The majority algorithm performed well at selecting objects deemed to be classified as “None of the Above/Good candidate.” On a randomly selected group of about 500 subjects, 94% of those classified as “None of the Above/Good candidate” by the majority algorithm agreed with the gold standard set rankings. For the other five classification buttons, the agreement between classifications selected via the plurality algorithm and the gold standard set varied: 52% for “Extended,” 54% for “Not round,” 24% for “Object Moves Off the Crosshairs,” and 7% for “Empty circle.” This lack of agreement suggests that users may sometimes neglect to click *all* of the relevant buttons when a subject is “bad” for more than one reason. However, since our desire is simply to rule out false positives, this lack of agreement on the precise nature of the false positive does not hamper our study.

Deciding on classifications via the plurality algorithm generally leads to larger disagreements with the gold standard set and the classification data. The exception to this rule seems to be the case of “multiple objects in the red circle”; subjects classified as such by the plurality vote agree with the gold standard set 96% of the time.

Table 2 shows the total selection distribution between the categories on the main site from the first ~ 1 million classifications. Users selected the “good” button nearly 20% of the time. Using the majority algorithm to determine a final classification, we get a 16% yield of “good” objects based on those classifications.

Some of the brighter Disk Detective subjects show diffraction spikes and noise from detector saturation, especially in DSS, Sloan, and *WISE 1* images. When we launched Disk Detective we did not explicitly explain these phenomena anywhere on the site, though we readily answered questions about them in the chat forum TALK. Nonetheless, many of our first users interpreted bright stars with spikes as “oval” or in some cases “extended.” On 2014 March 31, we edited the spotters guide and the tutorial at DiskDetective.org to include examples with diffraction spikes, which greatly reduced the problem, but some of these mistakes persist in our data. We expect this confusion over diffraction spikes to get better over time, as most of our classifications now come from participants who are well aware of the problem.

2.3. Disk Detective Objects of Interest (DDOIs)

With nearly 300,000 total objects from AllWISE to classify and 18%–25% yield described above, we estimate that the online classification scheme at DiskDetective.org will produce up to 75,000 good objects, still a daunting number to investigate. However, the automated online classification stage is just the beginning of our vetting process. The next stage of the process aims to harvest DDOIs, subjects that we consider deserving of additional follow-up observations.

For this next stage of vetting, we created a collection of Google spreadsheets that both the science team and the superusers could edit. We first populated the spreadsheets with subjects chosen as “None of the Above/Good Candidate” using the latest classification data. We chose subjects in the right portion of the sky for any upcoming follow-up observations and also populated the lists in order of agreement fractions and brightness (in *J* band), ensuring that the bright subjects, and those with high agreement percentages were looked at first. Then we invited the superusers to add their own favorite objects, which they collect using tools on the TALK social network. We also invited the entire Disk Detective user community to submit subjects automatically via a Google form (but not directly edit the spreadsheet).

We coached the superusers on how to research each source in SIMBAD and VizieR (and sometimes NED) to fill in information about spectral type, proper motion, variability, parallax, prior observations, and make comments on the SED, etc. Then we checked the superusers’ comments and selected the follow-up targets (DDOIs) from the list based on the following criteria.

1. SIMBAD object descriptions excluding post-AGB stars, carbon stars, novae, Cepheids, cataclysmic variables, high-mass X-ray binaries, eclipsing binaries, galaxies, Active Galactic Nuclei, planetary nebulae, reflection nebulae, rotational variables, symbiotic stars, or Wolf-Rayet stars. Note that we did keep sources with SIMBAD object descriptions Shell Star, Orion Variable, and White Dwarf.
2. No Long Period Variables, SR+L, Slow Irregular Variables, Miras, Semi-regular Variables, Semi-regular Pulsating Variables, or Carbon stars based on literature searches.
3. Including only spectral types B through M according to SIMBAD, when a type is available.

Only about half of the subjects have entries in SIMBAD, so we often relied on VizieR to help us search the relevant literature. Many of the subjects had unknown spectral types, and many were severely reddened. So at this stage of the vetting process, we often simply labeled sources as “late-type based on color.”

1. No sign that the *WISE 1* photometry drops out due to saturation, based on visual inspection of the SED.
2. No known companions within $16''$, except for spectroscopic binaries.
3. For M stars and other subjects with $V - J > 1$, we require $[W1] - [W4] > 0.9$. The peak of the thermal emission from cool star photospheres may lie at long enough wavelengths that the Rayleigh-Jeans limit no longer accurately describes the photosphere’s $[W1] - [W4]$ color even in the absence of circumstellar dust. Hence, we impose this more stringent requirement of 0.9 mag of $22 \mu\text{m}$ excess for these cool subjects. For some red subjects, however, we find that the SED clearly curves upwards at *W4*; we do not exclude these subjects.

In the process, we naturally rediscovered many known disks, so we added the following additional criterion:

1. No sources that have already been imaged by a pointed space mission (i.e., *Spitzer*, *Herschel*, *HST*) or 8 m class telescope (i.e., Keck, Gemini, VLT, *Magellan*; based on

the published literature), except those without good quality spectral types (SIMBAD quality C or higher).

Our search includes many sources in and near the Galactic plane and many sources with no parallax measurements from *Hipparcos*. But since these sources require extra care, we have added two additional criteria for the purpose of this paper:

1. No sources within 5° of the Galactic plane.
2. Only sources with parallax measurements from *Hipparcos*.

Most of the subjects on the vetting spreadsheets do not meet these additional eight criteria. But sources that do meet all these criteria (plus the criteria in Table 1, of course) we label as DDOIs and place in our queue for follow-up observations. So far we have collected 770 DDOIs in total that have survived the above vetting by the science team and/or by multiple superusers. Of these, 517 have classification histories via the main Disk Detective online classification tool. The remaining DDOIs have minimal classification data, since they were submitted directly by volunteers, and some volunteers choose to flip through images on TALK rather than the main site.

The yield at this stage varies greatly depending on how we rank the spreadsheets. Higher galactic latitudes provide higher yields, as do sources that are brighter in *WISE* 4. Also, as our users have become more educated, they have become better at selecting subjects to place on the spreadsheets, which raises the yield. But as of now, the typical yield at this stage of vetting (DDOIs per source on the vetting spreadsheet) is about 12%.

Multiplying the size of the input catalog by the automated vetting yield (24.4%) and then by the DDOI vetting yield (about 12%) gives a final yield of about 3% and an estimate of the total number of DDOIs we ultimately expect to discover of ~ 8000 . This estimate suggests that our search is presently about 10% complete. However, it is worth noting that many of these criteria, while catching a lot of the aforementioned AGB stars that pollute our “good” candidate list, also eliminate a population of YSO disks and M stars (for example, YSOs can be irregular variables).

2.4. Using DDOIs as Quality Indicators

The DDOIs have all been carefully vetted by hand and researched in the literature by multiple scientists and/or well-trained superusers. So this list of sources serves as a reference set of subjects that we can use to make decisions about how to run the vetting.

One major consideration of any citizen science project is when to remove, or “retire,” subjects from classifications on the site. To make that decision, we need to know roughly how long it takes the user population to converge on an answer. Figure 2 shows how the standard deviation of the agreement for majority-algorithm-ruled “good” candidates in our DDOI set varies with number of total classifications. For these subjects, “agreement” is defined simply as the number of good classifications divided by the total number of classifications the subject has received. Between 10 and 15 classifications, the standard deviation of the agreement for this subset of “good” DDOI candidates levels off, suggesting that there would be minimal marginal benefit from requiring additional classifications beyond this point. Thus the current iteration of the site uses a conservative 15 total classifications as our benchmark for when to retire subjects.

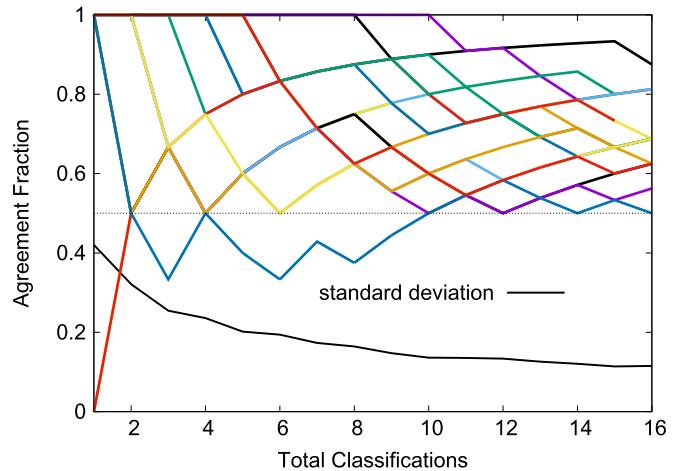


Figure 2. User agreement and standard deviation of the agreement vs. number of total classifications for DDOI candidates ultimately marked as “good.” The faint lines show how the running agreement for individual subjects in the set vary with the number of classifications and the dark bold line shows the running standard deviation of the entire set. The standard deviation flattens out between 10 and 15 classifications, so we use 15 classifications as our cutoff for the retirement of subjects from the site.

3. FOLLOW-UP OBSERVATIONS

We obtained and analyzed spectra for our DDOIs using the FAST spectrograph (Fabricant et al. 1998) on the Tillinghast 1.5 m telescope at Fred Lawrence Whipple Observatory during 2014 May–October, employing the 300 mm^{-1} grating and the $3''$ slit. These spectra cover $3800\text{--}7500 \text{ \AA}$ at a resolution of $\sim 6 \text{ \AA}$. We flux- and wavelength-calibrated the spectra using the Image Reduction and Analysis Facility (IRAF) software system. After trimming the CCD frames at each end of the slit, we corrected for the bias level, flat-fielded each frame, applied an illumination correction, and derived a full wavelength solution from calibration lamps acquired immediately after each exposure. The wavelength solution for each frame has a probable error of $\pm 0.5\text{--}1.0 \text{ \AA}$. We extracted sources and sky spectra using the optimal extraction algorithm within APEXTRACT. The absolute flux-calibration for each night relies on observations of two to five standard stars (Hayes & Latham 1975; Barnes & Hayes 1982; Massey et al. 1988) and has an uncertainty of $\pm 5\text{--}10\%$.

This follow-up spectroscopy has proven vital to our vetting of DDOIs. Roughly half of all DDOIs initially have no reliable spectral type. Additionally, for these objects the luminosity class is generally completely unconstrained, save for clues from parallax and proper motion measurements. This situation is more dire for red (or reddened) sources and late-type stars, which will be the subject of future work. M giants often produce their own dust, so disks around these stars are not of particular interest from the perspective of planet formation. FAST spectroscopy also allows us to screen for certain false positives, such as blended active galactic nuclei.

Though our initial candidate list in this paper is small, we plan to follow-up on the entire list of DDOIs and thus will have a large number of observed spectra. Manual spectral classification of every object is highly inefficient. To speed up the process, we used the semiautomatic quantitative spectral-typing code SPTCLASS¹⁹, an IRAF/IDL code based on the

¹⁹ <http://dept.astro.lsa.umich.edu/~hernandj/SPTclass/sptclass.html>.

methodology outlined in Hernández et al. (2004). We decided that in this first paper we would only publish DDOIs that are in *Hipparcos* catalogs, and consequently do have spectral types inferred from color and parallax, as a test of our pipeline.

SPTCLASS calculates spectral types of stars using spectral indices, comparing line fluxes of spectral features that are sensitive to effective temperature (e.g., Morgan et al. 1943; Pritchett & van den Bergh 1977; Reid et al. 1995; Coluzzi 1999; Stock & Stock 1999; Gray et al. 2001, and others). SPTCLASS uses three independent spectral-typing modules: indices characterizing early (OBA, 44 indices), intermediate (FG, 11 indices), and late (KM, 16 indices) spectral types. Each index is based on the equivalent width for the spectral feature, which is calculated by measuring the decrease in flux from the expected continuum due to line absorption. Indices measured by this procedure are generally insensitive to reddening, so long as each band's wavelength coverage is relatively small. The indices have been calibrated as a function of spectral type using O8-M6 main-sequence standards; this large extent for all indices ensures that degeneracy of determination is eliminated.

SPTCLASS first calculates three spectral types for each star, one for each module, by taking a weighted average of the indices used in that module, using weights estimated from the computed error for each index. It then calculates a weighted average of spectral types for all indices, using weights estimated by the computed error for each index from the measured index value and fit of the index to spectral type. The code then discards spectral indices if the spectral type they indicate deviates from the mean by $>3\sigma$, or if their computed error is more than six subtypes. This minimizes possible contamination by artifacts and emission lines. We report here the final SPTCLASS spectral type, which is the new average after discarding these deviant indices. The typical $1-\sigma$ formal uncertainty for these classifications is two subtypes.

3.1. First Batch of Disk Candidates

Table 3 lists our first batch of 50 DDOIs. Although some of the sources in Table 3 are Be stars or disk candidates that have been previously reported, 37 of these sources are new disk candidates from the Disk Detective project.

We have decided to present only stars in with *Hipparcos* parallaxes as a test of our spectral classification pipeline. Though Wu et al. (2013) performed a large search for stars in the *Hipparcos* catalog that have excess emission in $W4$, they used a $[K_s] - [W4]$ cutoff, and the All-Sky catalog. Our selection criteria $[W1] - [W4]$ and input catalogs (AllWISE) are different, so we find objects they missed, even in the *Hipparcos* catalog.

Table 3 lists the photometry of these disk candidates from 2MASS and SIMBAD (V band). It also lists the spectral types we derived from our FAST spectra, informed by photometry and *Hipparcos* distances, and *WISE* photometry. Eleven of the sources in Table 3 have previously been reported as disk candidates by Zuckerman et al. (2012), Kennedy & Wyatt (2012, 2013), Wu et al. (2013), Wahhaj et al. (2013), and Chen et al. (2014). We include these objects here as a check on our consistency with other searches. However, 37 of these sources have not been previously reported as disk candidates (and are not Be stars).

The *WISE* photometry in Table 3 is taken from the All-Sky data release and corrected for saturation according to the formulae in Patel et al. (2014). This corrected All-Sky *WISE*

photometry is more accurate for sources brighter than about the 8th magnitude in $W1$. No color correction is applied to the *WISE* photometry, since the sources have nearly Rayleigh-Jeans spectra. If we assumed instead that the emission were entirely from a 200 K blackbody disk, the correction to *WISE 4* would be -0.016 mag (Wright et al. 2010), negligible compared to our 0.25 mag *WISE 1-WISE 4* excess criterion for inclusion in Disk Detective.

The vast majority of our DDOIs did not have previously identified luminosity classes. We assigned them using the approach of Patel et al. (2014), who separated dwarfs from giants in their catalog via a simple cut on an HR diagram, retaining only stars with $M_V > 6.0(B - V) - 1.5$. Figure 3 shows an HR diagram for our candidates. The stars on Figure 3 are color coded according to spectral type: blue = B, green = A, and red = F. Most of our disk candidates are probably dwarfs since they fall below the Patel et al. (2014) cut, which is shown by the dashed line.

If we conservatively assume that the quality of our spectra corresponds to the “C” quality flag in SIMBAD (“A” is the highest quality), then of the 50 targets with new spectral types that we present here, 48 show clear improvements in quality over the published literature, while the remaining 9 targets have spectral types of the same quality as the literature. The root-mean-square change in spectral types versus SIMBAD for all 50 targets was 2.64, with 27 objects shifted at least one subtype toward lower temperature and 12 objects shifted at least one subtype toward higher temperature.

Figures 4 and 5 summarize some more properties of the stars in Table 3: their distribution on the sky and their distance distribution. Their distance distribution peaks at roughly 110 pc; at this distance, a large telescope with adaptive optics like the Gemini Planet Imager could image an analog to the HR 8799 planetary system and comfortably detect at least two of the planets. For comparison, Patel et al. (2014) limited their study to sources with distances <120 pc. But the distance distribution of the DDOIs is grossly similar to that of the disk candidates in Wu et al. (2013), which peaks at about 90 pc. The distribution of our DDOIs also includes a long tail of objects beyond 200 pc; roughly one-fourth of our DDOIs are in this long tail, which consists mainly of B and early A dwarfs. Figure 5 also shows the distances to some well-known disks for references. The distribution of the DDOIs on the sky is shaped mainly by the range of declinations accessible to FAST and our decision to add sources to the Web site grouped by Galactic latitude.

Figure 6 compares our disk candidates with the Wu et al. (2013) color selection criteria. We did not impose such a criterion, but all of our disk candidates fall to the right of the dotted line in this figure, showing that they meet the Wu et al. (2013) criterion anyway. On average, the disk candidates in this paper are somewhat redder than those in Wu et al. (2013).

We fit simple models to the SEDs consisting of a stellar photosphere plus a blackbody dust component using a Levenberg–Marquardt algorithm. These fits yield constraints on the dust temperature and the fractional infrared luminosity, f , the total bolometric power emitted by our single-temperature blackbody disk model divided by the total bolometric power emitted by our blackbody stellar model for each source. For stars with excess at only *WISE 4*, these fits yield only an upper limit on the dust temperature and a lower limit on f .

Table 3
Observational Parameters for Disk Detective Objects of Interest

Zooniverse	Identifiers		Spectral Type	Distance (pc)	Photometry		Excess ($W1-W4$)	Notes
	HD Number	WISE			V^a	J		
AWI0005sz	7406	J011636.23+740136.6	B1III	549 ± 103	7.05	6.764 ± 0.056	0.619 ± 0.052	a, EW = -17.7064
AWI0005w41	218546	J230817.21+511146.3	B8III	794 ± 498	8.25	7.918 ± 0.020	0.784 ± 0.072	...
AWI0005bow	164137	J175911.27+135417.8	B8III	820 ± 544	8.02	7.768 ± 0.020	1.571 ± 0.050	...
AWI0005ae4	224098	J235449.26+742436.2	B8V	240 ± 20	6.6	6.388 ± 0.019	0.433 ± 0.050	...
AWI0000m2p	4670	J004848.00+181850.6	B9IV	386 ± 98	7.94	7.841 ± 0.020	0.898 ± 0.080	...
AWI0005sx	6370	J010652.55+743754.5	B9IV	286 ± 51	8.37	8.168 ± 0.029	0.919 ± 0.084	...
AWI0005yiz	6612	J010722.60+380143.9	B9V	316 ± 51	7.14	7.156 ± 0.020	0.453 ± 0.067	b, EW = 4.09265
AWI00062a8	9985	J013756.15+211539.9	B9V	202 ± 23	7.95	7.824 ± 0.018	1.993 ± 0.050	...
AWI0005ym7	23873	J034921.76+242251.0	B9V	122 ± 9	6.61	6.602 ± 0.017	1.138 ± 0.057	...
AWI0000uj2	138422	J153046.05+342756.4	B9V	108 ± 4	6.81	6.719 ± 0.021	0.542 ± 0.044	c
AWI00004o8	152308	J165204.85+145827.2	B9V	124 ± 7	6.5	6.504 ± 0.034	0.494 ± 0.048	...
AWI00062fh	207888	J215221.25-031028.9	B9V	184 ± 17	6.6	6.688 ± 0.026	1.235 ± 0.045	...
AWI0000bs0	2830	J003140.76-014737.3	A0V	107 ± 6	7.07	6.917 ± 0.019	1.015 ± 0.054	c
AWI0004ne5	9590	J013525.89+560237.3	A0V	162 ± 12	7.02	6.791 ± 0.018	0.655 ± 0.059	...
AWI0005yjp	14893	J022515.75+370707.9	A0V	192 ± 23	7.34	7.341 ± 0.018	1.037 ± 0.067	...
AWI0005ylw	22614	J033906.73+244209.8	A0V	120 ± 11	7.09	6.979 ± 0.024	0.614 ± 0.079	...
AWI00000wz	74389	J084546.93+485243.4	A0V	111 ± 7	7.48	7.296 ± 0.023	0.678 ± 0.067	...
AWI0000u8s	85672	J095359.12+274143.5	A0V	107 ± 7	7.58	7.215 ± 0.024	1.011 ± 0.056	c, d
AWI0000w9x	129584	J144313.04+014928.7	A0V	156 ± 20	7.33	7.260 ± 0.024	0.785 ± 0.052	...
AWI0000tz1	140101	J154030.20+370101.1	A0V	165 ± 13	7.17	7.109 ± 0.023	0.688 ± 0.051	c
AWI0000gjb	214982	J224206.62-032824.4	A0V	124 ± 9	7.16	7.101 ± 0.021	0.655 ± 0.065	...
AWI0000kg4	218155	J230533.05+145732.5	A0V	106 ± 5	6.77	6.712 ± 0.020	0.579 ± 0.059	...
AWI0000fyd	224155	J235537.71+081323.7	A0V	128 ± 6	6.82	6.770 ± 0.026	0.462 ± 0.066	...
AWI00062h1	224429	J235746.21+112827.6	A0V	95 ± 4	6.65	6.581 ± 0.023	0.426 ± 0.058	c
AWI0002mhd	173056	J184305.97+071626.4	A0.5V	203 ± 32	8.24	7.913 ± 0.027	1.504 ± 0.061	...
AWI0005yjn	14685	J022317.32+381509.7	A1IV	254 ± 35	7.14	7.037 ± 0.018	0.706 ± 0.062	e, EW = 4.01698
AWI0005abf	213290	J222753.27+704806.0	A1IV	326 ± 44	7.7	7.318 ± 0.023	0.627 ± 0.050	...
AWI0005mry	3051	J003412.66+540359.0	A1V	229 ± 32	7.6	7.350 ± 0.019	0.354 ± 0.073	...
AWI0000phh	18271	J025614.05+040254.2	A1V	116 ± 14	7.71	7.664 ± 0.032	0.974 ± 0.073	...
AWI0005zz5	25466	J040238.47-004803.7	A1V	113 ± 8	6.93	6.835 ± 0.020	0.606 ± 0.053	...
AWI0000uji	134854	J151147.67+101259.8	A1V	113 ± 8	6.87	6.799 ± 0.018	0.405 ± 0.065	...
AWI0005zx4	11085	J014928.21+244048.7	A2V	151 ± 22	8.31	8.013 ± 0.019	0.641 ± 0.088	...
AWI0004nfu	21375	J032853.67+490412.8	A2V	160 ± 21	7.47	7.186 ± 0.026	0.468 ± 0.068	e
AWI00062aj	12445	J020221.16+192323.6	A3V	227 ± 43	8.4	8.090 ± 0.024	0.944 ± 0.084	...
AWI0000tgc	84870	J094902.82+340506.9	A3V	88 ± 5	7.19	6.752 ± 0.032	0.398 ± 0.063	c, f
AWI0005bps	165507	J180533.55+182643.9	A3V	193 ± 33	8.17	7.952 ± 0.019	0.672 ± 0.099	...
AWI0005vyx	204829	J212959.78+413037.3	A3V	175 ± 13	7.34	6.744 ± 0.021	0.338 ± 0.067	...
AWI0005w29	212556	J222412.79+484918.7	A3V	150 ± 11	7.61	7.415 ± 0.029	0.469 ± 0.060	...
AWI0005a9r	208410	J215305.45+682955.0	A3V	189 ± 15	7.48	7.297 ± 0.027	0.539 ± 0.053	...
AWI0005ykd	21062	J032448.99+283908.6	A4V	102 ± 6	7.12	6.838 ± 0.018	0.612 ± 0.058	...
AWI0000v1z	138214	J152954.11+234901.6	A5V	138 ± 13	7.58	7.124 ± 0.024	0.322 ± 0.061	...
AWI0006222	201377	J210916.04-001405.6	A7V	101 ± 5	6.66	6.355 ± 0.032	0.574 ± 0.064	...
AWI0005w7o	20994	J032429.84+341709.9	A8V	245 ± 60	8.67	7.931 ± 0.027	1.191 ± 0.074	...
AWI0005c01	199392	J205143.50+730449.3	A8V	169 ± 18	8.28	8.004 ± 0.023	0.402 ± 0.071	...
AWI00002ms	71988	J083100.44+185806.0	F0V	83 ± 4	7.42	6.977 ± 0.032	0.358 ± 0.080	...
AWI00062bl	22128	J033337.91-072453.8	F0IV	166 ± 20	7.56	6.919 ± 0.024	0.853 ± 0.053	...
AWI0005yk3	19257	J030651.95+303136.8	F0V	79 ± 4	7.06	6.470 ± 0.020	2.093 ± 0.037	g
AWI000048c	87827	J100719.80-152718.9	F2V	107 ± 8	8.12	7.425 ± 0.020	0.968 ± 0.060	c
AWI0000hjr	221853	J233536.20+082256.9	F2V	68 ± 3	7.34	6.559 ± 0.019	1.335 ± 0.042	c, f, h
AWI0002yt	157165	J172007.53+354103.6	F6V	100 ± 7	8.27	7.496 ± 0.020	0.427 ± 0.064	...

Notes. (a) H α partially in emission. (b) H α in emission. (c) Appears in Wu et al. (2013). (d) Appears in Wahhaj et al. (2013). (e) Appears in Zuckerman et al. (2012). (f) Appears in Chen et al. (2014). (g) Appears in Kennedy & Wyatt (2013). (h) Appears in Kennedy & Wyatt (2014).

^a From SIMBAD. Typical uncertainty 0.01 mag.

Figure 7 shows SEDs for six of the new disk candidates, together with a simple model for the flux. Our SEDs employ WISE 1 and WISE 2 fluxes from the All-Sky survey catalog corrected for saturation using the formulae in Patel et al. (2014). For these models, we supplemented the WISE and 2MASS photometry with *UBVRI* photometry from SIMBAD when it was available. Dashed lines show blackbody models fit

to the star and to the disk component; the solid line shows the total model flux.

Table 4 summarizes our SED modeling of all the 39 new disk candidates reported in Table 3. The temperature uncertainties are $1-\sigma$ uncertainties from the shape of the two-parameter χ^2 surface near the minimum. The fractional infrared excesses, f , listed in this table should all be considered lower

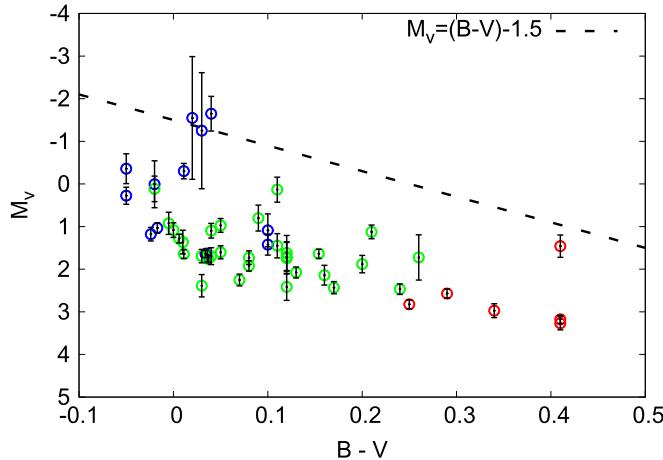


Figure 3. Most disk candidates in this paper easily fall below the limit on the HR diagram used by Patel et al. (2014) to separate giants from dwarfs. The three stars that lie above this limit (HD 7406, HD 164137, and HD 2185461) are Be stars. The points are color coded according to spectral type: blue = B, green = A, and red = F.

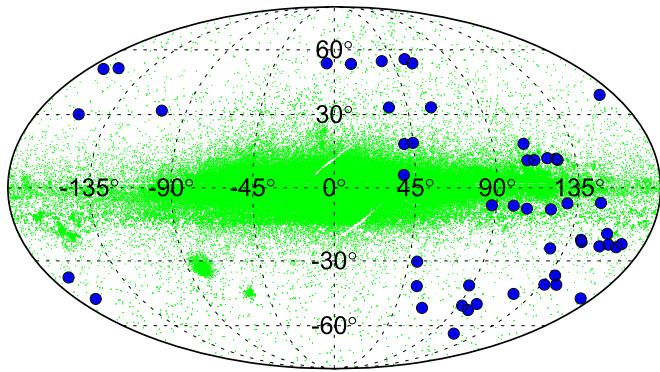


Figure 4. Galactic latitude and longitude distribution of the disk candidates in Table 1 (blue points) and all of the subjects pre-selected based on the criteria in Table 1 plus the $J < 14.5$ criterion (green points). The distribution of the DDOIs on the sky is shaped mainly by the range of declinations accessible to FAST and our decision to add sources to the Web site grouped by Galactic latitude.

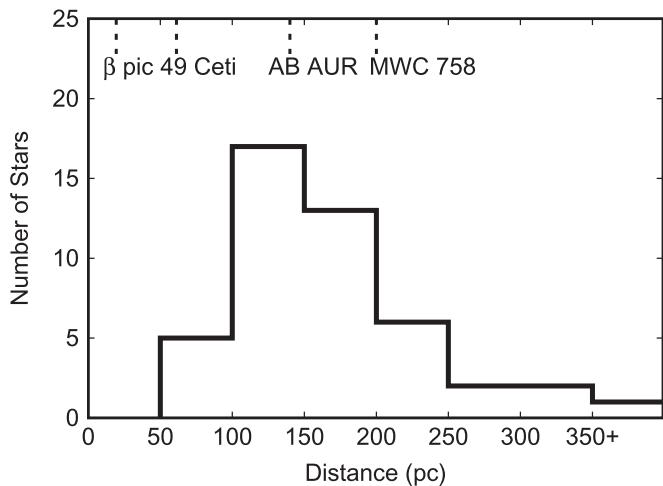


Figure 5. Distribution of distances to the DDOIs listed in Table 3 from Hipparcos parallax measurements. The distance distribution of the DDOIs is grossly similar to that of the disk candidates in Wu et al. (2013), but also includes a long tail of objects beyond 200 pc, which consists mainly of B and early A dwarfs.

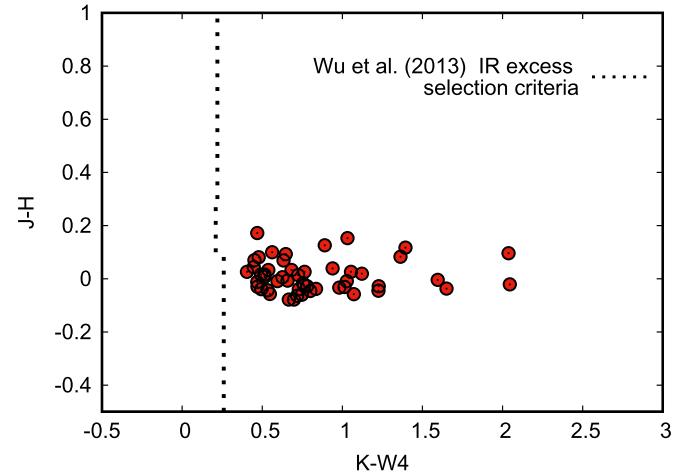


Figure 6. $J-H$ vs. $K-W4$ for the sources presented in Table 3. J , H , K magnitudes are from 2MASS, and $W4$ is the $22\ \mu m$ magnitude from AllWISE. The faint dashed line shows the criterion used by Wu et al. (2013) for selecting disk candidates with $22\ \mu m$ excess (see their Figure 2). Even though we selected our subjects based on $W1-W4$, our disk candidates all satisfy their criterion as well.

limits. When the temperature listed is an upper limit, the listed fractional infrared excess corresponds to a blackbody disk with the listed upper limit temperature.

3.2. $W3$ Excesses

All of the DiskDetective.org subjects were pre-selected to have $[W1]-[W4]$ excesses significant at the $5-\sigma$ level or better, based on the AllWISE catalog. We also checked the All-Sky catalog of *WISE* photometry for these sources, correcting this photometry via the correction factors in Patel et al. (2014). All of the sources in Table 4 still have $[W1]-[W4]$ excesses significant at the $5-\sigma$ level or better using the corrected All-Sky photometry.

Additionally, eight of our DDOIs also turned out to have significant ($>3\sigma$) $W3$ excess based on corrected All-Sky *WISE* photometry, as you can read in Table 5: HD 4670, HD 7406, HD 14685, HD 19257, HD 20994, HD 164137, HD 173056, and HD 218546.

3.3. Binary Companions and Variability

Five of the stars with newly reported excesses have known companions or secondaries. None of these is apparent in the *WISE* images or likely to be cool enough to create the observed *WISE* infrared excesses.

HD 7406. The separation of this binary is $\sim 1'$ (Schmeidler 1987), wide enough that the companion probably does not contaminate the *WISE* photometry.

HD 74389. HD 74389 has several nearby companion stars that could conceivably contaminate the SED. This star has a white dwarf companion discovered by Sanduleak & Pesch (1990), type DA1.3, $V = 14.62$ located $20''.11$ (2230 au projected separation) to the west (Holberg et al. 2013), which is readily visible in the DSS Blue, Red, and IR images. Holberg et al. (2013) also report a nearby M dwarf companion to the west and a nearby sdB companion to the north. Indeed, the same DSS Blue, Red, and IR image bands also show a second background object $\sim 13''$ to the north of the star, probably the sdB.

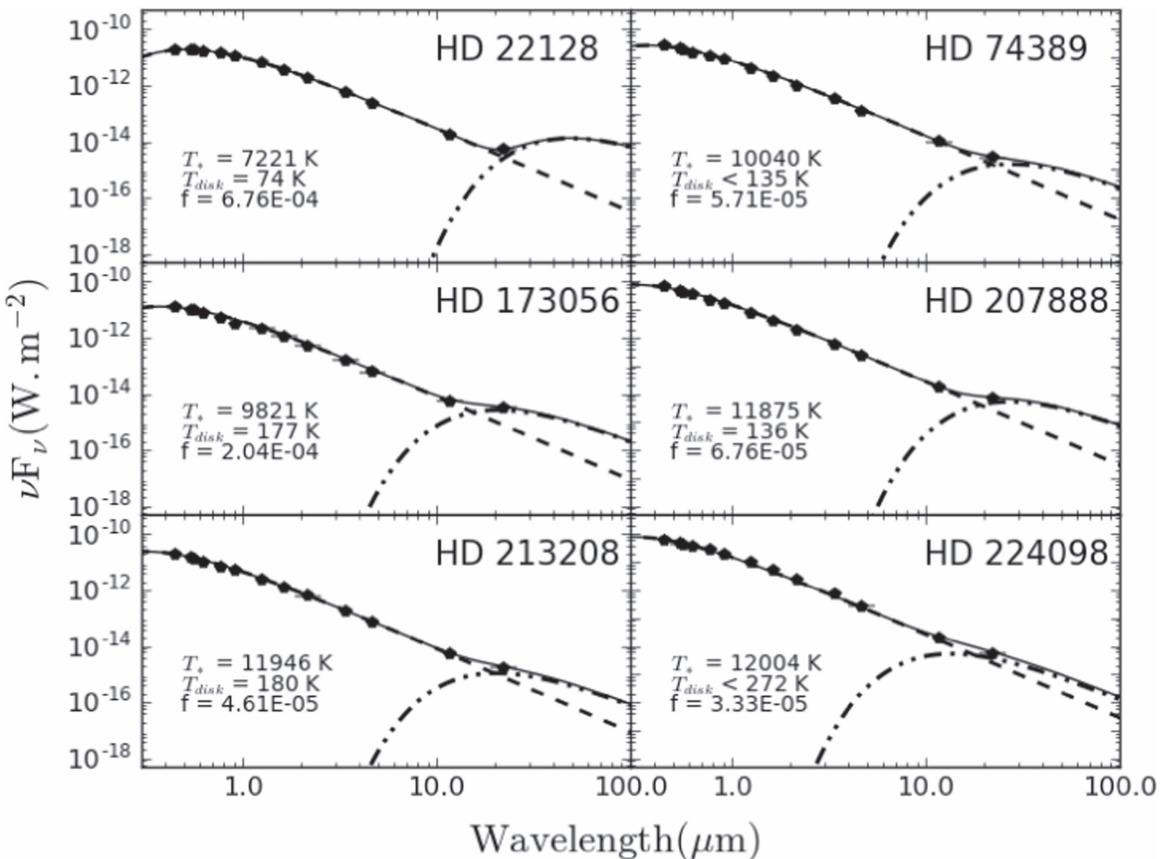


Figure 7. Sample spectral energy distributions (SEDs) for six new disk candidates. Dashed lines show blackbody models fit to the star and to the disk component; the solid line shows the total model flux.

However, none of these objects is visible in the *WISE* or even the 2MASS images, which look like unsaturated clean point sources, probing a resolution at *K* band of $<4''$. Moreover, we only see a significant excess at *W4*, but not at *W3*. The *W3*–*W4* color of an M dwarf or sdB photosphere is ~ 0 since the SEDs of these objects are in the Rayleigh–Jeans limit beyond $10\text{ }\mu\text{m}$, so the lack of a resolved companion at the *K* band rules out a significant contribution to the *W4* excess from the companion’s photospheric emission. An M dwarf or sdB unresolved by 2MASS might still be cause for concern if it were itself dusty. But since the A star is the most luminous object in the system, it seems most likely that we are observing reprocessed light from dust around the A star.

See below for further discussion of this interesting object.

HD 23873 is a member of the Pleiades, and listed as a binary in the *Hipparcos* Input Catalog (Turon et al. 1993). The separation of the binary is 11000 au according to Makarov & Robichon (2001), corresponding to roughly $1''$.

HD 207888 was identified as a visual binary by Scardia (1988). The separation is wide enough ($\sim 1'$) that the companion probably does not contaminate the *WISE* photometry.

HD 22128 is a double-lined spectroscopic binary system with a semimajor axis of roughly nine solar radii. We found a spectral type of F0IV for the combined light, but in fact, the two components are A stars; see Folsom et al. (2013) for a detailed spectral analysis. This star was first identified as a debris disk host star by Silverstone (2000) based on *Infrared Space Observatory* photometry (60 – $100\text{ }\mu\text{m}$). We are the

first to report excess emission from this source at *WISE 4*, excess at the 8σ level. Wyatt et al. (2007) considered the object to be an “anomalous system” based on the high luminosity of its inferred disk given the star’s age. We note that the system appears slightly extended in 2MASS K, but not in other bands. No $18\text{ }\mu\text{m}$ flux measurement for this object appears in the AKARI catalog (Ishihara et al. 2010).

HD 224098. This B8V star is listed as having faint optical secondary, type F0V, in Lindroos (1985). No data on the separation of the secondary is provided.

Two of our new disk candidates have known stellar variability of ≤ 0.04 magnitudes in *V* band. This degree of variability is too small to affect our selection process.

HD 152308 is classified as an $\alpha 2$ CVn variable star by Dubath et al. (2011) and as type A0 Cr Eu according to Renson & Manfroid (2009). The variability is 0.04 mag in the *Hipparcos* system, with a period of 0.94 days. Our classification yielded B9V.

HD 218155. This star varies over a range of ~ 0.03 magnitudes in the *V* band (Watson 2006).

3.4. Be Stars

Aside from normal main-sequence stars, our program reveals four new Be stars: *HD 6612*, *HD 7406*, *HD 164137*, and *HD 218546*. None of these stars appears in the Be Star Spectra (BeSS) database (Neiner et al. 2011) or in the Kohoutek & Wehmeyer (1999) catalog of H α emission-line stars, which

Table 4
Disk Parameters for New Disk Candidates

HD Number	Disk Temperature (K)	f
3051	<248	$<4.7 \times 10^{-5}$
4670	156 +21/-19	5.0×10^{-5}
6370	<168	$<6.9 \times 10^{-5}$
6612	<189	$<2.1 \times 10^{-5}$
9590	<200	$<5.8 \times 10^{-5}$
9985	<136	$<2.8 \times 10^{-4}$
11085	<200	$<7.2 \times 10^{-5}$
12445	<157	$<8.9 \times 10^{-5}$
14893	<117	$<3.4 \times 10^{-5}$
18271	<104	$<8.7 \times 10^{-5}$
20994	187 +25/-19	2.5×10^{-4}
21062	<207	$<8.0 \times 10^{-4}$
22128	<74	$<6.8 \times 10^{-4}$
22614	<187	$<6.8 \times 10^{-5}$
23873	<131	$<7.3 \times 10^{-5}$
25466	<107	$<3.3 \times 10^{-5}$
71988	<221	$<1.1 \times 10^{-4}$
74389	<136	$<5.7 \times 10^{-5}$
129584	<170	$<4.2 \times 10^{-5}$
134854	<251	$<4.6 \times 10^{-5}$
138214	<216	$<9.6 \times 10^{-5}$
152308	<228	$<2.6 \times 10^{-5}$
157165	<245	$<1.4 \times 10^{-4}$
158419	<244	$<4.1 \times 10^{-5}$
165507	<205	$<7.2 \times 10^{-5}$
173056	177 +20/-18	2.0×10^{-4}
199392	<252	$<8.4 \times 10^{-5}$
201377	<131	$<1.7 \times 10^{-5}$
204829	<263	$<5.3 \times 10^{-5}$
207888	<136	$<6.8 \times 10^{-5}$
208410	<212	$<7.4 \times 10^{-5}$
212556	<225	$<6.2 \times 10^{-5}$
213290	<198	$<7.3 \times 10^{-5}$
214982	<193	$<6.3 \times 10^{-5}$
218155	<193	$<6.3 \times 10^{-5}$
224098	<272	$<3.3 \times 10^{-5}$
224155	<242	$<4.4 \times 10^{-5}$

cataloged stars with galactic latitude $|b| < 10^\circ$. All show prominent H α emission and the strong upper level Balmer absorption lines characteristic of classical Be stars (e.g., Porter & Rivinius 2003; Rivinius et al. 2013). Two of these stars, HD 7406 and HD 218546, show significant W3 excess, in addition to their W4 excess. The B1e star HD 7406 has additional strong He I absorption lines; we detect weak He I absorption in the other stars. Our high S/N spectra reveal no trace of other permitted or forbidden emission lines, confirming our classification of them as classical Be stars rather than pre-main-sequence Herbig Ae/Be stars (Raddi et al. 2013; Rivinius et al. 2013).

These stars illustrate the promise of Disk Detective for identifying new classical Be stars. Although Be stars have large infrared excesses due to free-free emission (e.g., Gehrz et al. 1974; Dougherty et al. 1994; Bragg & Kenyon 2002), most have been identified from optical spectroscopic surveys (e.g., Slettebak 1985; Bragg & Kenyon 2002; Porter & Rivinius 2003; Raddi et al. 2013, 2015; Rivinius et al. 2013) optical photometric surveys (e.g., Wisniewski & Bjorkman 2006), and IR spectroscopic surveys (e.g., Chojnowski et al. 2015). Uniform selection based on infrared excess from *WISE* data

provides a flux-limited survey fairly independent of interstellar reddening.

With only four Be stars, it is premature to analyze spectroscopic properties and statistics for the *WISE* sample of classical Be stars. However, it is worth noting that the frequency of late-type Be stars (three B8–B9 stars and one B1 star) is somewhat larger than expected from current samples where B4 and earlier stars are much more common (e.g., Bragg & Kenyon 2002; Rivinius et al. 2013). We plan a detailed analysis of a larger sample in a future paper (A. S. Bans et al. 2016, in preparation).

3.5. HD 74389: A Star With a Candidate Debris Disk and White Dwarf Companion

One of our new candidate debris disks appears to orbit HD 74389 A, an A0V star with a white dwarf companion discovered by Sanduleak & Pesch (1990). It also has a possible M dwarf companion and sdB companion reported by Holberg et al. (2013). There are three known planetary systems with white dwarfs as distant companions: Gl 86 (Queloz et al. 2000; Mugrauer & Neuhäuser 2005), HD 27442 (Butler et al. 2001; Chauvin et al. 2006; Raghavan et al. 2006; Mugrauer et al. 2007), and HD 147513 (Porto de Mello & da Silva 1997; Mayor et al. 2004). However, the HD 74389 system appears to contain the first debris disk around a star with a white dwarf companion.

The white dwarf companion, HD 74389 B, has V mag 14.62, and is located 20.11 arcsec to the west (Holberg et al. 2013). It is readily visible in the DSS Blue, Red, and IR images. These bands also show a background object $\sim 13''$ to the north of the star, possibly the M dwarf companion reported by Holberg et al. (2013).

It is interesting to ponder the origin of the disk around HD 74389 A and how the post-main-sequence evolution of HD 74389 B may have affected it. Though the white dwarf may presently have a projected separation of 2230 au, it could have been two to four times closer in when it was on the main sequence, thanks to stellar mass loss. The disk may have merely survived the evolution of the higher-mass star, or it may represent a signature of dynamical changes to the system, like planet exchange (e.g., Kratter & Perets 2012). It could even have been built from mass loss by the companion, via the process described by Perets & Kenyon (2013).

4. DISCUSSION

We have outlined and demonstrated a novel process for identifying new candidate circumstellar disks in the *WISE* survey data. This paper reports only results from the first 10% of the search, so it might be premature to try to derive any statistically meaningful inferences about the population of debris disk from this limited sample. But our list of 37 new, well-vetted disk candidates demonstrates the utility of crowdsourcing analysis of *WISE* images. One of our disk candidate systems appears to contain the first debris disk discovered around a star with a white dwarf companion: HD 74389. We also report four newly discovered classical Be stars (HD 6612, HD 7406, HD 164137, and HD 218546) and a new detection of 22 μ m excess around previously known debris disk host star HD 22128.

We decided to only publish in this paper candidates that are in the *Hipparcos* catalog. Since the Wu et al. (2013) cross-

Table 5
Disk Detective Objects of Interest with $W3$ Excess

Zooniverse ID	HD Number	WISE ID	Spectral Type	Distance	$WI-W3$
AWI00055sz	7406	J011636.23+740136.6	B1III	549 ± 103	0.292 ± 0.042
AWI0005w41	218546	J230817.21+511146.3	B8III	794 ± 498	0.102 ± 0.029
AWI0005bow	164137	J175911.27+135417.8	B8III	820 ± 544	0.646 ± 0.029
AWI000m2p	4670	J004848.00+181850.6	B9IV	386 ± 98	0.115 ± 0.030
AWI0002mhd	173056	J184305.97+071626.4	A0.5V	203 ± 32	0.114 ± 0.030
AWI0005yjn	14685	J022317.32+381509.7	A1IV	254 ± 35	0.205 ± 0.035
AWI0005w7o	20994	J032429.84+341709.9	A8V	245 ± 60	0.174 ± 0.030
AWI0005yk3	19257	J030651.95+303136.8	F0V	79 ± 4	0.718 ± 0.031

correlated the *WISE* archive with the *Hipparcos* catalog, they could conceivably have identified all of the candidates that we are announcing. However, Wu et al. (2013) used a [2MASS]–[$W4$] color criterion, as opposed to our [WI]–[$W4$] color criterion, and used the *WISE* All-Sky data rather than the *WISE* ALLWISE data. Yet all of the candidates presented in this paper would have also been selected by the Wu et al. (2013) color criterion. More importantly, while we examined candidates by eye to discard objects potentially contaminated by nearby stars and galaxies, Wu et al. (2013) used a statistical likelihood-ratio (LR) technique to accomplish this goal. Perhaps their statistical technique was more conservative than our more labor intensive approach, leaving these candidates unidentified.

We have several further improvements to the Disk Detective project underway, which we will describe in upcoming papers, including the following:

1. New ways to retire the sources after fewer classifications.
2. Spectroscopy of Southern Hemisphere DDOIs via the CASLEO telescope in Argentina.
3. Imaging follow-up of DDOIs with the Robo-AO (Baranec et al. 2014) instrument at the Palomar Observatory 60 inch telescope to check for background contaminants located closer to the star than DSS can probe.

With its high sensitivity and angular resolution in the mid-infrared, we expect that the *James Webb Space Telescope* (*JWST*) will be an important tool for following up disks discovered via Disk Detective, so we aim to have the project mostly completed by the time *JWST* launches in the fall of 2018.

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The DSS was produced at the Space Telescope Science Institute under U.S. Government grant NAG W-2166. The images of these surveys are based on photographic data obtained using the Oschin Schmidt Telescope on Palomar Mountain and the UK Schmidt Telescope. The plates were processed into the present compressed digital form with the permission of these institutions. This work has made use of the BeSS database, operated at LESIA, Observatoire de Meudon, France: <http://basebe.obspm.fr>.

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APPENDIX ADDITIONAL NOTES ON INDIVIDUAL DISK CANDIDATES

HD 9985. All four *WISE* images appear slightly ($\approx 2''$) offset from the DSS and 2MASS images to the SW, a common feature of slightly saturated images in AllWISE.

HD 14893. Skiff (2014) lists this star as a B9.5 V. We find type A0V.

HD 22614. Possible member of the Pleiades cluster (White et al. 2001).

HD 173056. Grenier et al. (1999) found a spectral type of A1V; we find A0.5 V.

HD 213290. DSS Blue, Red, and IR show an unresolved background object roughly 14 arcsec north of the star.

HD 25466. Our spectral classification (A1V) matches that of Paunzen et al. (2001) for this star.

HD 134854. The SDSS images of this star are highly distorted.

HD 201377. Using higher resolution spectra ($R = 42,000$), Prugniel & Soubiran (2001) find that this star is type A3 with $\log g = 3.93$. (We find A7V.)

HD 20994. Azimlu et al. (2015) identified this star in the Perseus OB 2 association as an AGB star based on its *K-WISE* 4 color alone. However, our spectral typing (A8V) shows that this object is more likely to be a debris disk.

HD 19257. The *WISE* 3 and *WISE* 4 excesses of this star have been reported by Kennedy & Wyatt (2013). We are the first to report a luminosity class for it; we find a spectral type of F0V, which disagrees with SIMBAD but is consistent with Kennedy & Wyatt (2013).

HD 87827. All four *WISE* images appear slightly ($\approx 2''$) offset from the DSS and 2MASS images to the SW, a common feature of slightly saturated images in AllWISE.

HD 221853. We find a spectral type F2V for this well-known debris disk host star, consistent with the age estimate of 100 Myr in Mittal et al. (2015).

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