

THE LOCAL HOSTS OF TYPE IA SUPERNOVAE

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ABSTRACT

We use multi-wavelength, matched aperture, integrated photometry from *GALEX*, SDSS and the RC3 to estimate the physical properties of 166 nearby galaxies hosting 168 well-observed Type Ia supernovae (SNe Ia). The ultra-violet (UV) imaging of local SN Ia hosts from *GALEX* allows a direct comparison with higher redshift hosts measured at optical wavelengths that correspond to the rest-frame UV. Our data corroborate well-known features that have been seen in other SN Ia samples. Specifically, hosts with active star formation produce brighter and slower SNe Ia on average, and hosts with luminosity-weighted ages older than 1 Gyr produce on average more faint, fast and fewer bright, slow SNe Ia than younger hosts. New results include that in our sample, the faintest and fastest SNe Ia occur only in galaxies exceeding a stellar mass threshold of $\sim 10^{10} M_{\odot}$, leading us to conclude that their progenitors must arise in populations that are older and/or more metal rich than the general SN Ia population. A low host extinction sub-sample hints at a residual trend in peak luminosity with host age, after correcting for light-curve shape, giving the appearance that older hosts produce less-extincted SNe Ia on average. This has implications for cosmological fitting of SNe Ia and suggests that host age could be useful as a parameter in the fitting. Converting host mass to metallicity and computing ^{56}Ni mass from the supernova light curves, we find that our local sample is consistent with a model that predicts a shallow trend between stellar metallicity and the ^{56}Ni mass that powers the explosion, but we cannot rule out the absence of a trend. We measure a correlation between ^{56}Ni mass and host age in the local universe that is shallower and not as significant as that seen at higher redshifts. The details of the age – ^{56}Ni mass correlations at low and higher redshift imply a luminosity-weighted age threshold of ~ 3 Gyr for SN Ia hosts, above which they are less likely to produce SNe Ia with ^{56}Ni masses above $\sim 0.5 M_{\odot}$.

Subject headings: galaxies: evolution – supernovae: general

1. INTRODUCTION

A fundamental motivation for studying Type Ia supernovae (SNe Ia) is to arrive at a physical explanation for the observed fact that these objects have well-behaved and calibratable explosions (Phillips 1993). This property gives cosmographers a tool for measuring universal expansion and provided the first direct observa-

tional evidence for cosmic acceleration (Riess et al. 1998; Perlmutter et al. 1999). Knowing the physical process that generates SNe Ia provides strong constraints on the nature of the SN Ia progenitor. This, in turn, allows us to predict the evolution of the SN Ia population (e.g., Howell et al. 2007). This results in tighter control of the systematic errors arising from population evolution with redshift, an important uncertainty in measuring cosmological parameters (e.g., Astier et al. 2006; Conley et al. 2009). A well-constrained physical model for SNe Ia would have further utility in tracing high-redshift star formation and for predicting the effects of SNe Ia on the chemical enrichment of their host galaxies (e.g., Kobayashi & Nomoto 2007).

In order to calibrate SNe Ia and make them useful for cosmology, their intrinsic variation must be measured and accounted for. The peak absolute magnitude of a given SN Ia is a strong function of its initial decline rate (Phillips 1993), and a weaker function of its peak color (Riess et al. 1996; Tripp 1998; Tripp & Branch 1999; Parodi et al. 2000). These empirical correlations reduce the intrinsic variation of ~ 1 magnitude in the B-band to ~ 0.1 magnitudes and thus provide the accurate luminosity distance estimates required for measuring cosmological parameters (e.g., Riess et al. 1996; Tonry et al. 2003; Guy et al. 2005; Prieto et al. 2006; Jha et al. 2007; Conley et al. 2008). The goal of

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SN Ia progenitor studies is to use these measures of SN Ia intrinsic variation to explore trends with host properties that could shed light on the underlying population that produces SNe Ia (Phillips 1993; Hamuy et al. 1995, 2000; Howell 2001; Gallagher et al. 2005; Sullivan et al. 2006; Gallagher et al. 2008).

Since SN Ia explosions are powered by the radioactive decay of ^{56}Ni , it is widely accepted that the intrinsic variation in SN Ia brightness and decline rate is primarily driven by the amount of ^{56}Ni present in the SN explosion, with more luminous and slower declining explosions being powered by more ^{56}Ni (Truran et al. 1967; Colgate & McKee 1969). One avenue of exploration would be to connect this theoretical idea to the progenitor population of SNe Ia by comparing host galaxy properties with the measures of SN Ia light curve variation. This requires a mechanism that varies the amount of ^{56}Ni as a function of some property of the host galaxy.

Timmes et al. (2003) presented a model relating progenitor metallicity to produced SN Ia ^{56}Ni mass for a constant (Solar) O/Fe ratio. This model was expanded and tested by Howell et al. (2009, hereafter H09) using an intermediate redshift ($0.2 < z < 0.75$) sample of SNe Ia. They used spectral energy distribution (SED) fitting of optical integrated host photometry to infer the host stellar masses for their sample. These masses were used as a proxy for host metallicities through the Tremonti et al. (2004) mass-metallicity relation. They estimated the peak bolometric luminosity and rise time from the SN Ia photometry to calculate the required mass of ^{56}Ni to power the explosion and compared the calculated host metallicity versus Ni mass trend to the Timmes et al. (2003) model. Their data are consistent with the model, although with a higher scatter than predicted by the theory. They also found that varying the O/Fe ratio according to thin or thick-disk models produced no substantial change in their results.

If the dependence of ^{56}Ni mass on host metallicity is real, then fainter and faster declining SNe Ia should be associated with higher metallicity, and thus more massive, host galaxies. If a decline in the rate of fast fading, fainter SNe Ia with redshift could be detected, it would be evidence for the relation between Ni mass and metallicity because we expect the more distant universe to consist of lower-mass and lower-metallicity objects on average when compared with the local universe. At this time, however, Malmquist biases complicate an accurate determination of the rate evolution for low-luminosity SNe Ia (Foley et al. 2008; Gonzalez-Gaitan et al. 2009). We therefore look to host properties of low-luminosity SNe Ia in the local universe, where they are easier to detect and apparently more abundant. We use this larger sample to see if the fainter SNe do appear in higher mass, higher metallicity galaxies on average.

To compare host properties with SN Ia variation, we construct a low-redshift sample of SNe Ia with well observed light curves that provide a light curve shape parameter, called stretch, the observed maximum light color, and the peak luminosity of the SN in the rest-frame B-band (Conley et al. 2008). We characterize the hosts of these SNe using newly generated and catalog galaxy integrated UV and optical magnitudes measured within matched apertures as inputs to a SED fitting program

that estimates host stellar mass, luminosity-weighted age and star formation rate (SFR). We first use these results to explore the relation between the observed light curve parameters and the derived host properties and compare our results to studies at higher redshift using the same methods (Sullivan et al. 2006, H09). We isolate a subset of hosts with low internal extinction in order to explore the relationship between SN peak brightness and host properties where the source of line-of-sight color for the SNe is minimized. We then extend the work of H09 by applying their techniques to our local sample to derive SN Ia ^{56}Ni -mass, from the light-curve photometry, and compare these with host metallicities, from our derived host stellar masses, and host luminosity-weighted ages. Throughout this paper we express host masses in M_{\odot} and SFR in $M_{\odot} \text{ yr}^{-1}$ and assume a Hubble constant of $H_0 = 73 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2. SAMPLE SELECTION AND ANALYSIS

Here we outline the properties of our low-redshift SN Ia sample and describe the methods we use to characterize the SN and host galaxy properties.

2.1. The Local SN Sample and Light Curve Fitting

For our local sample we start with the low- z literature sample summarized in Conley et al. (2008). We supplement this with more recent SN photometry from Hicken et al. (2009a). The initial pool of nearby SNe Ia are required to have modern CCD photometry with errorbars, are required to have phase coverage beyond 3 days after maximum light and prior to 7 days after maximum light. They are also required to have one rest-frame B-band magnitude between 8 days before and 13 days after maximum light and one restframe U or V-band magnitude in the same phase range. In addition the wavelength range of the filters must be between 2700 and 7200 Å.

To be consistent with H09 we adopt the SN Ia light curve fitting method from Conley et al. (2008). This method fits the SN light curve data to produce the stretch, s (Perlmutter et al. 1997), a peak restframe B-band apparent magnitude, m_{Bmax} , and a SN color, \mathcal{C} , which is somewhat like $(B - V)_{Bmax}$ in the sense that redder SNe have higher \mathcal{C} values. Briefly, the stretch parameter is a time-axis scale factor that is applied to the light curve that aligns its shape with a ‘canonical’ SN Ia light curve. The sense of the stretch parameter is such that faster, fainter SNe Ia have low stretch values and brighter, slower SNe Ia have higher stretch values. The sense of the SN color parameter is such that bluer SNe at maximum are brighter. The corrected peak brightness of the SN Ia used for cosmology is determined by a linear combination of s , \mathcal{C} , and m_{Bmax} (Astier et al. 2006; Conley et al. 2008). Conley et al. (2008) present a much more detailed description of this method including comparisons with other popular fitting routines. These comparisons demonstrate that, at least for this study, our results are not dependant on the fitting technique.

Figure 1 shows distributions of the fitted SN light curve properties for all the low-redshift SNe passing the previously described criteria in hosts with sufficient integrated host photometry to provide good SED fits (168 SNe Ia, filled histograms). This is the base sample we will examine in this work. The figure shows the distributions in

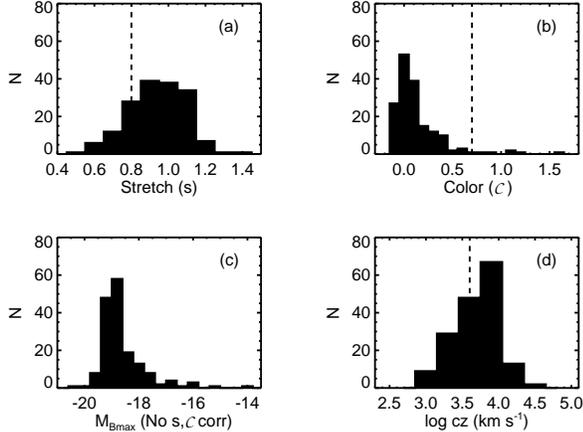


FIG. 1.— Frequency diagrams of light curve fit parameters for the local SNe Ia having enough observations to determine a stretch value and having sufficient host photometry to allow good fits to their SED: (a) stretch, (b) color, (c) M_{Bmax} (uncorrected for stretch-color), (d) recession velocity in km s^{-1} . The vertical dashed lines show various limits described in the text: (a) the limit between ordinary and low-luminosity SNe Ia ($s = 0.80$), (b) the peak color limit above which the fitting becomes less accurate ($C = 0.7$), and (d) the Hubble flow limit adopted here ($cz = 4000 \text{ km s}^{-1}$, $z = 0.013$). There are 168 SNe Ia in the sample with well-fit host SEDs.

stretch, peak color, absolute B magnitude (uncorrected for stretch-color), and recession velocity. The light curve fit properties of this sample are given in Table 1 which lists the SN name, the host galaxy, the recession velocity, the light curve stretch, the fitted apparent B magnitude at maximum light, and a status column which is described below. More details of the light-curve fits for these SNe can be found in Conley et al. (2009).

Low-luminosity SNe Ia are generally better represented in lower redshift samples for multiple reasons. The spectra and the stretch-luminosity and color-luminosity relations for low stretch SNe Ia differ enough from higher-stretch SNe (Garnavich et al. 2004; Taubenberger et al. 2008) to increase the errors in fitting their light curves beyond the error-budget requirement for cosmological parameter determination. This, combined with their faintness, makes low-stretch SNe less appealing targets for the spectroscopic followup required for cosmological surveys. As a result, the SN Ia training set derived for cosmology in Conley et al. (2008) contains no SNe with $s < 0.70$, and very few below $s = 0.75$. At lower redshifts where even faint SNe are relatively easy to followup, the unusual objects tend to attract attention and acquire spectroscopy.

The break in the stretch-luminosity relation appears to be at $s \sim 0.80$ (Gonzalez-Gaitan et al. 2009). For this paper, we define the set of low-luminosity SNe Ia based on stretch to be those with $s < 0.80$ (to the left of the dashed vertical line in Figure 1a). The training set from Conley et al. (2008) allows us to extend this limit slightly to $s = 0.75$ when we apply corrections for stretch-luminosity.

Our local sample contains 29 SNe with $s < 0.80$ of which 12 are below $s = 0.70$. Since we are using a light curve method developed for cosmology, this means our fitting can only tell that an object with $s < 0.70$ is low-stretch, but cannot accurately measure the stretch. We

have increased the error bars by a factor of three for objects with $s < 0.70$. The sample used in H09 was derived from the Supernova Legacy Survey (SNLS, Astier et al. 2006) which was optimized for cosmology. Either due to the fact that such objects are very rare beyond $z = 0.2$, or because the spectroscopic followup was biased against them or both, there are no objects with $s < 0.75$ in the H09 sample (see also Bronder et al. 2008). The low stretch SNe Ia in our sample are indicated with an ‘L’ in the status column in Table 1.

A similar situation exists for the peak color, C . SNe Ia with colors greater than $C \sim 0.7$ (see Figure 1b) are also unappealing objects for cosmological surveys because they are fainter and because they have larger fitting errors. These red SNe Ia should not be excluded when examining host properties, however. Our sample contains six SNe Ia with $C > 0.7$. These SNe are indicated with an ‘R’ in the status column of Table 1. We have increased the peak color errors by a factor of ten for these SNe to reflect the higher uncertainty in this parameter.

We emphasize that while the local sample assembled here is appropriate for exploring the link between host and SN properties, many of the SNe in our sample are too local ($cz < 4000 \text{ km s}^{-1}$, $z < 0.013$, to the left of the vertical dashed line in Figure 1d) or have stretch errors or colors large enough to exclude them from use in cosmological fits. We indicate the cosmology status of our sample SNe Ia in the last column of Table 1 where an H indicates a SN Ia in the Hubble flow ($cz > 4000 \text{ km s}^{-1}$), and a C indicates a SN Ia that is eligible for cosmological fitting based on the light-curve fit quality and having parameters within the range where the light curve fitting has reasonably low errors.

2.2. Integrated Host Photometry and SED fitting

We assembled host integrated photometry in the ultraviolet (UV) and optical to characterize the SN Ia host SEDs. For our low-redshift host galaxy sample, the addition of the *GALEX* UV photometry improves the estimation of short-term star formation (e.g., Martin et al. 2005, Figure 1), and provides similar constraints on recent star formation when compared with the higher-redshift samples from the SNLS, whose blue optical photometry corresponds to the rest-frame UV. All host magnitudes are total magnitudes using matched elliptical apertures having the standard D25 diameter, to be comparable with the RC3. All host magnitudes are corrected for Milky Way (foreground) extinction using the dust maps of Schlegel et al. (1998).

In the UV, we generated host integrated magnitudes as part of the preliminary efforts to produce the *GALEX* Large Galaxy Atlas (GLGA, Seibert et al. 2009). This atlas will measure $\sim 20,000$ galaxies imaged by the *GALEX* UV-imaging satellite (Martin et al. 2005) having diameters in the UV greater than one arcminute. The *GALEX* imaging mode has two bandpasses, one in the far-UV (FUV, $\lambda_{eff} = 1539 \text{ \AA}$, $\Delta\lambda = 442 \text{ \AA}$) and another in the near-UV (NUV, $\lambda_{eff} = 2316 \text{ \AA}$, $\Delta\lambda = 1060 \text{ \AA}$). We use a technique similar to that used to generate the Nearby Galaxy Atlas (Gil de Paz et al. 2007) with which we cross-checked our integrated UV magnitudes. This method will be described in detail in Seibert et al. (2009), but to summarize, we perform sur-

face photometry in elliptical apertures on sky subtracted images that have had foreground point sources and background galaxies masked.

For optical photometry, we used the RC3 (de Vaucouleurs et al. 1991) and/or images obtained from the Sloan Digital Sky Survey (SDSS¹³, York et al. 2000) in the five SDSS bands: *ugriz*. Our sample required either SDSS coverage or an integrated magnitude from the RC3. The RC3 total integrated Johnson UB_V magnitudes were obtained directly from NED¹⁴. For the larger hosts in our sample, we found the SDSS catalog data to be inaccurate for a number of reasons. Some of these hosts spanned multiple image strips and many were broken up into sub-regions making the determination of a total flux problematic. In order to enforce consistency across wavelengths we decided to coadd and mosaic the SDSS image data for each SN host and derive the integrated photometry ourselves. To achieve this, we adapted our integrated photometry methods developed for the GLGA to SDSS image data. The major difference between GALEX and SDSS data is in the treatment of the sky background which is extremely low in GALEX data. We checked the consistency of our integrated SDSS magnitudes by comparing them with RC3 magnitudes and found them to agree within the errors. All UV and optical integrated galaxy magnitudes are presented in Seibert et al. (2009).

Sullivan et al. (2006) fit optical photometry to galaxy SED models produced with the PEGASE.2 SED galaxy spectral evolution code (Fioc & Rocca-Volmerange 1997; Le Borgne & Rocca-Volmerange 2002; Le Borgne et al. 2004). We adopted this method to allow a direct comparison with the higher redshift host studies of Sullivan et al. (2006) and H09. As with these studies, all redshifts are known from the hosted SNe. This has the benefit of eliminating redshift degeneracies in the SED fits. The average internal extinction was allowed to vary in the fits over the range $0.0 < E(B - V)_{HOST} < 0.7$ in increments of 0.05 magnitudes using a Calzetti et al. (1994) dust model (see Fioc & Rocca-Volmerange 1997). We point out the stellar mass derived from PEGASE.2 models is an estimate of the current mass in stars and is not the integral of the star formation history for a given galaxy (Sullivan et al. 2006). We also emphasize that our ages are strongly influenced by the flux produced by ongoing star formation yielding a fairly tight correlation between our derived host ages and specific star formation rates. Thus, we should keep in mind that our luminosity-weighted properties might be different from those produced with mass-weighted measurements. The details of the galaxy modeling and SED fitting method employed here can be found in §3 of Sullivan et al. (2006).

To select the sample with good SED fits, we first require that a unique solution be found. This reduced our initial host sample from 258 down to 209, mostly due to the rejected SED having too few photometric points. We next require SED coverage in both the UV and optical bands, further reducing our sample down to 174 hosts. At this point we examined the distribution of χ^2_{ν} and eliminated an additional six hosts that had extreme values ($\chi^2_{\nu} > 30$, see Figure 2 and

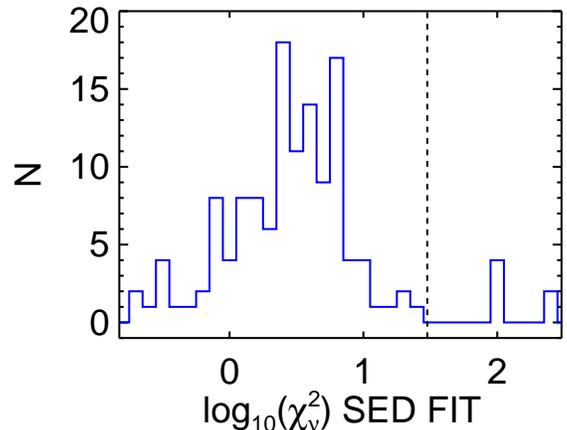


FIG. 2.— Frequency diagram of χ^2_{ν} of the host SED fits illustrating one criterion for eliminating unusable hosts: we require $\chi^2_{\nu} < 30$ for the fit. Large values of χ^2_{ν} arise because local hosts have smaller photometric errors than hosts at higher redshifts. The other criteria require a unique solution for the SED and require UV and optical photometric points in each SED. These criteria reduce the sample of SNe Ia with good light curves from 258 down to 168 and produce 166 unique host SED fits (two galaxies in our sample hosted two SNe Ia each, see text).

Le Borgne & Rocca-Volmerange 2002) leaving a total sample of 168. The 90 hosts that did not pass these criteria had no or limited optical photometry mostly due to not being members of any of the major galaxy catalogs (i.e., anonymous) and not being in the SDSS footprint which would allow us to derive the integrated flux from SDSS imaging. A subset of these also had no UV coverage because of proximity to a UV-bright star, which precludes GALEX observations due to detector safety. Table 2 lists the host T-type (de Vaucouleurs 1959) and the SED fit derived properties for the 166 unique hosts in our sample.

The number of unique hosts is two less than the number of SNe Ia because two galaxies hosted two SNe Ia each: NGC1316 hosted SN1980N and SN1981D, and NGC5468 hosted SN1999cp and SN2002cr. In NGC1316, the SNe Ia produced are low-stretch ($s \sim 8.5$) with values that differ by only 3%, while the two SNe Ia produced in NGC5468 are normal stretch ($s \sim 0.98$) and differ by 11%. This is certainly not a conclusive test of using integrated host magnitudes for comparison with stretch values, but it gives no cause to doubt this method, as would be the case if these pairs of SNe Ia differed by greater amounts.

Figure 3 shows the distributions of derived host properties for the local (blue histograms) and H09 (red histograms) samples. There are some interesting differences between the two samples. The discrepancy in the mass distributions (Figure 3b) can be explained because local SNe are discovered in host-targeted surveys which prefer more massive hosts, while the SNLS is an areal survey without a mass bias and therefore includes lower mass hosts. The discrepancy in the SFR distributions (Figure 3c) could be due to a luminosity bias from the host-targeted low-redshift surveys. Here the higher redshift areal survey does not prefer luminous, and therefore higher SFR, hosts. These differences are reinforced

¹³ <http://www.sdss.org>

¹⁴ <http://nedwww.ipac.caltech.edu/>

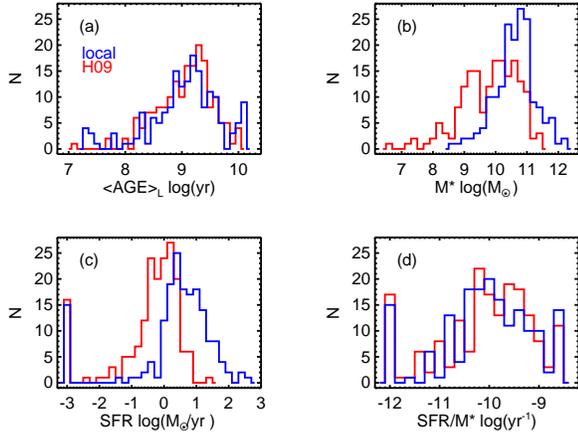


FIG. 3.— Frequency diagrams of host properties derived from SED fitting using the PEGASE.2 library (Le Borgne et al. 2004) with the SNLS sample from H09 as the red histograms and the local SN Ia light-curve fit sample as the blue histograms: (a) luminosity-weighted age, (b) stellar mass, (c) star formation rate, and (d) sSFR defined as SFR/M^* . Figure (c) illustrates our SFR threshold of $10^{-3} M_{\odot}/\text{yr}$. Figure (d) illustrates our sSFR threshold of 10^{-12} yr^{-1} . Hosts below these values are set to the threshold value. We see that the local sample is missing the low-mass and intermediate SFR peaks seen in the H09 sample. In sSFR, the two distributions are quite similar.

if there is a correlation between mass and SFR. In this case, the same bias that produces the deficit of low-mass hosts in the local mass distribution produces the deficit of lower SFR hosts in the local SFR distribution. Figure 3d shows that when considering the specific star formation (sSFR), defined as SFR divided by stellar mass, our distributions are fairly similar. Thus, from here on we consider sSFR in preference over SFR.

3. RESULTS

3.1. Host Properties versus SN Light Curve Stretch

We start by examining correlations in host properties with SN Ia light curve stretch. In the following discussion, it is good to keep in mind that higher stretch SNe Ia are brighter and lower stretch SNe Ia are fainter. In these plots, we highlight the transition between the low luminosity (i.e. low-stretch) and normal SNe with an (orange) dashed line at $s = 0.80$.

Figure 4 plots the log of the sSFR against hosted SN Ia light curve stretch. For this plot and plots following, the symbols indicate the host’s sSFR: filled circles (red) indicate hosts with no detected SF ($sSFR \leq 10^{-12} \text{ yr}^{-1}$), filled stars (blue) indicate strongly star-forming hosts ($sSFR > 10^{-9.5} \text{ yr}^{-1}$) and filled squares (green) indicate hosts with intermediate sSFR. Typically, one would expect elliptical galaxies to belong in the lowest sSFR group and spiral galaxies to belong to the intermediate and high sSFR groups. Our aim here is to use a metric that is more physical than morphology, hence we do not use morphological classifications to examine these SN Ia hosts.

We notice several interesting features in this diagram. We see that the sensitivity threshold of our models is just above $\sim 10^{-12} \text{ yr}^{-1}$. We set all hosts with no detected sSFR to this value and consider these to be ‘dead’ hosts. At the high end, we see the timescale limit which is determined by the lifetime of stars with SEDs that

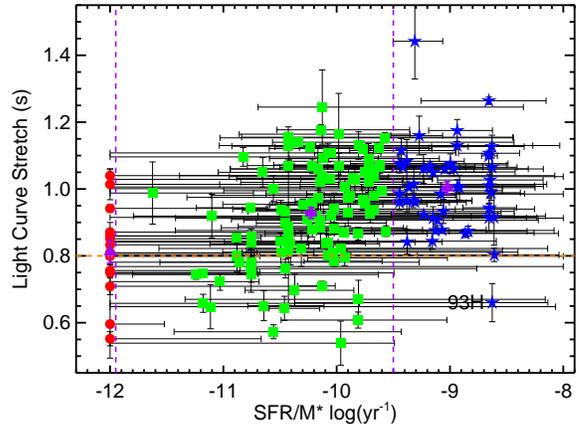


FIG. 4.— Specific star formation rate as derived from PEGASE.2 SED fits plotted as a function of light curve stretch for SN Ia hosts. Filled circles (red) indicate hosts with very low specific star formation ($sSFR \leq 10^{-12} \text{ yr}^{-1}$), filled stars (blue) indicate hosts with strong specific star formation ($sSFR > 10^{-9.5} \text{ yr}^{-1}$) and filled squares (green) indicate hosts with intermediate sSFR. Averages for the three sSFR bins (vertical dashed purple lines) are shown as the filled purple triangles with the error bars indicating the error in the mean. The division between normal and low-stretch SNe is indicated by the horizontal orange line at $s = 0.8$. The one high sSFR host of a low-stretch SN Ia (SN 1993H) is indicated. Stretch, on average, appears to increase with host sSFR.

peak in the UV; $\sim 3 \times 10^8 \text{ yr}$. Since all of our hosts are essentially at zero redshift, this limit appears as a line. The purple triangles show the average stretch in three sSFR bins indicated by the vertical dashed lines. Note that low stretch SNe are rare in high sSFR hosts (the one exception is SN1993H, as indicated on the plot), but common in intermediate and low sSFR hosts.

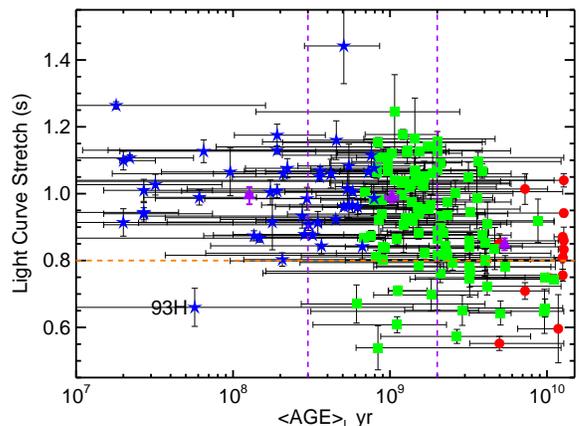


FIG. 5.— Luminosity-weighted stellar age as derived from PEGASE.2 SED fits plotted as a function of light curve stretch for SN Ia hosts. The host sSFR is coded as in Figure 4. The average in three age bins is shown at 3×10^8 and $2 \times 10^9 \text{ yr}$ (vertical dashed purple lines) is shown by the filled purple triangles. The horizontal dashed orange line indicates the division between normal and low-stretch SNe. The one low-stretch SN Ia with a luminosity-weighted age less than $3 \times 10^8 \text{ yr}$ (SN 1993H) is indicated. After a host exceeds 1 Gyr in age, it begins producing progressively lower stretch SNe Ia, on average.

Figure 5 plots the luminosity-weighted host age against stretch, with the same symbol coding as in Figure 4. This

plot shows that age and sSFR are well correlated with all the high sSFR galaxies having ages less than 1 Gyr and all the low sSFR hosts having ages greater than 4 Gyr as expected (see §2.2). The averages (purple filled triangles) were derived from three age bins (divided at 3×10^8 and 2×10^9 yr) and show a downward trend of stretch with age beyond 1 Gyr. Here we find the low-stretch SNe to be evenly distributed between host luminosity-weighted ages older than 5×10^8 yr, but rare below this age with SN 1993H again being the sole exception.

Figure 6 plots host stellar mass, M_* , against light curve stretch and shows that hosts of the lowest stretch SNe Ia tend to have stellar masses higher than $10^{10} M_\odot$. The higher stretch SNe Ia hosts have a much wider mass range (10^8 to $10^{12} M_\odot$).

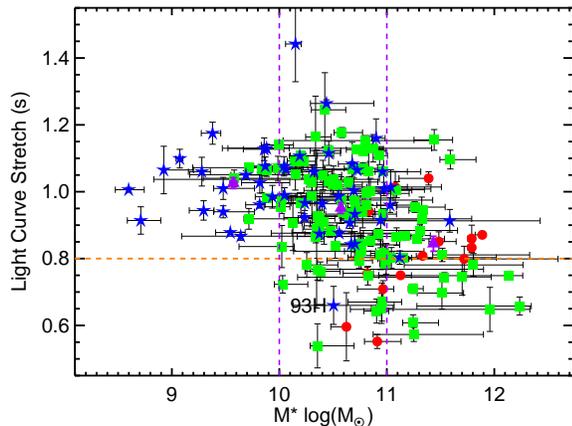


FIG. 6.— Stellar mass as derived from PEGASE.2 SED fits plotted as a function of light curve stretch for SN Ia hosts. The host sSFR is coded as in Figure 4. The averages of the sample divided at $\log M_* = 10$ and 11 (vertical dashed purple lines) are shown by the purple filled triangles. There are no hosts of low stretch SNe with masses less than $10^{10} M_\odot$. The exceptional SN 1993H is indicated (see text). This figure is very similar to Figure 4 in H09 derived from a higher redshift sample, but here we have a larger sample of low-stretch SNe Ia.

3.2. SN Ia Light Curve Color

There is a major challenge in interpreting SN Ia peak color because of the unknown mix of two potential sources of the color: intrinsic SN Ia color from the details of the explosion itself, and host galaxy line-of-sight reddening. In other words, the host extinction is not accounted for in the fitting process, which means that the peak color contains an intrinsic component and a host extinction component. This difficulty manifests itself when the color-luminosity relation is converted into a dust extinction law. In most cases the resulting R_V is much lower than what is found in the Milky Way or other galaxies (Tripp 1998). However, it is clearly not appropriate to assume Milky Way-like dust is responsible for SN Ia residual color (Conley et al. 2007).

There is potentially important information about the SN Ia explosion contained in the intrinsic color, if we could quantify and remove the line-of-sight, or host extinction part. Unfortunately, there are no observable signatures that allow these two sources to be disentangled, although extending the light curve data to the

near infra-red may minimize the problem (Kasen 2006; Wood-Vasey et al. 2008). Our host data show no correlations between host properties and SN peak color. The only feature worth noting, shown in Figure 7, is that all of the reddest SNe appear in hosts with intermediate age (~ 1 Gyr). The low-stretch SNe with $s < 0.80$, indicated in the plot with the large open (orange) circles, are not the reddest SNe as might be expected from the color-luminosity relation (Tripp 1998; Tripp & Branch 1999; Parodi et al. 2000).

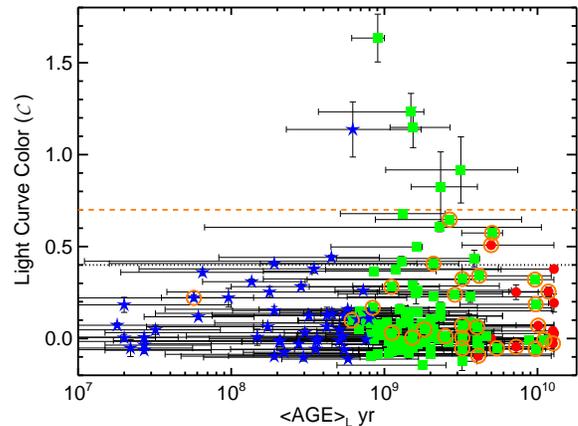


FIG. 7.— Luminosity-weighted age as derived from PEGASE.2 SED fits plotted as a function of SN color for the local SN Ia hosts. The host sSFR is coded as in Figure 4. The SNe with stretch less than $s = 0.80$ are overplotted with open (orange) circles. The two horizontal lines illustrate two color cuts applied to our sample: the orange dashed line at $C = 0.7$ is our cut for defining red SNe Ia, and the black dotted line at $C = 0.4$ is our cut for deriving Ni masses (see text). There is no obvious correlation, however all of the reddest SNe are hosted by galaxies with luminosity-weighted ages near 1 Gyr. It is also clear that the lowest stretch SNe are not the reddest SNe in our sample.

Another way to examine this issue is to look at SN color as a function of host extinction, as shown in Figure 8. We see an average trend in sSFR with host extinction such that hosts with higher sSFR also have higher host extinction, a finding that gives us encouragement that the extinction estimates from PEGASE.2 are robust. The average SN peak color (filled purple triangles) shows no trend with host extinction, and the reddest SNe appear in hosts covering a substantial range of host extinctions. This could be due to the clumpiness of host extinction which makes an average host extinction irrelevant for the specific line-of-sight to the SN.

3.3. A Subsample with Low Host Extinction

In an effort to mitigate the difficulty of untangling host extinction and intrinsic SN color, Gallagher et al. (2008) made a study of local SNe Ia in hosts that were morphologically classified as early-type and thus thought to be low-extinction hosts. They used spectroscopic line diagnostics to determine metallicities and ages. We have the opportunity to select hosts with low extinction inferred from our SED fits, without the assumption that early-type galaxies all have low extinction and without possible host classification errors.

To test the impact of host extinction on cosmological fitting, we look for residual trends in the cosmologically-

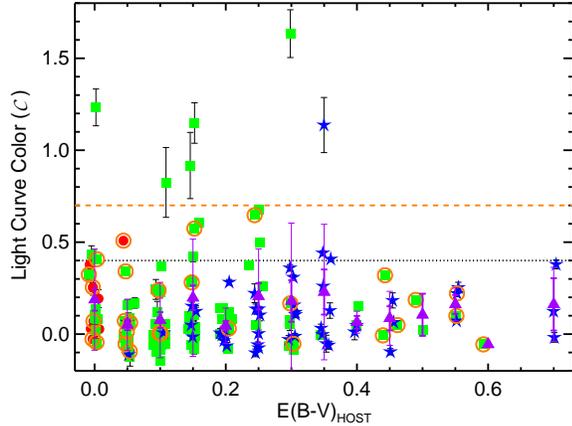


FIG. 8.— SN color as a function of average host color excess, $E(B - V)_{\text{HOST}}$, as derived from PEGASE.2 SED fits for the local SN Ia hosts. The host sSFR is coded as in Figure 4. The SNe with stretch less than $s = 0.80$ are overplotted with open (orange) circles, and the color cut lines are the same as in Figure 7. To better illustrate the distributions in each discrete bin of $E(B - V)_{\text{HOST}}$ (0.05 Mag), a small random offset was added to each point. The average SN color in each host extinction bin is shown by the filled (purple) triangles. Note that the sSFR goes up with host extinction, as expected, and that the reddest SNe appear in hosts with a range of host extinctions.

corrected peak SN Ia brightness as a function of host luminosity-weighted age. We calculate the peak absolute B-band brightness for our SNe Ia, corrected for stretch and color using:

$$M_B = m_{B_{\text{max}}} + \alpha(s - 1.0) - \beta C - \mu_B(z_{\text{spec}}), \quad (1)$$

where $\alpha = 1.2$, $\beta = 2.9$ (Sullivan et al. 2009) and $\mu_B(z_{\text{spec}})$ is the distance modulus based on the spectroscopic redshift (Conley et al. 2009). Since we are testing cosmological fitting, we eliminate low-stretch ($s < 0.75$) and red ($C > 0.7$) SNe Ia. The distance modulus, μ_B , is only accurate if the SNe Ia are in the Hubble flow so we also eliminate hosts with $z < 0.013$. After applying these cuts, we are left with a sample of 85 SNe Ia.

We plot the stretch-color-corrected absolute peak B magnitude against the luminosity-weighted host age for this sample in the bottom panel of Figure 9 using the previously defined symbols based on sSFR. In the top panel of Figure 9, we plot the same values for a low host extinction subsample ($N = 22$), with $E(B - V)_{\text{HOST}} \leq 0.05$ according to our SED fits. The bottom panel plot shows no obvious residual trend with age for the larger sample. The residual trend apparent in the low host extinction subset has a correlation of -0.68 . To determine the significance of this correlation, we selected 22 hosts at random from the full set of 85 hosts 10,000 times and ran the same linear regression analysis on each randomly selected sample. The low host extinction sample is a 2.1σ outlier when compared with the resulting distribution of random correlations.

3.4. Host Metallicity versus SN Ia ^{56}Ni Mass

Here we complement the higher-redshift ($0.2 < z < 0.75$) sample of H09 by using our nearby ($0.013 < z < 0.06$) sample of SNe Ia. This local sample, in addition to testing the Timmes et al. (2003) model in the nearby universe, will provide a test of the method developed in

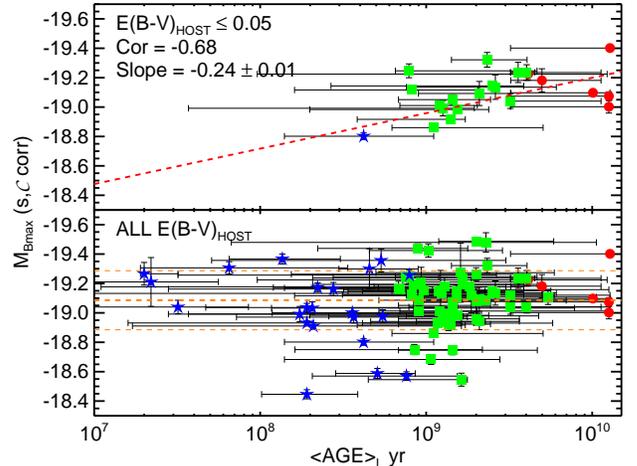


FIG. 9.— SN Ia stretch-color-corrected M_B as a function of host luminosity-weighted age for hosts in the Hubble flow ($z > 0.013$) plotted with the standard symbol coding. The host sSFR is coded as in Figure 4. Since we are applying stretch and color corrections, we cut the sample to remove SNe with $s < 0.75$ and $C > 0.7$. The *bottom* plot shows the Hubble flow subset ($N=85$) with no cut on host extinction, while the *top* plot shows a low host extinction subset ($N=22$) with $E(B - V)_{\text{HOST}} \leq 0.05$. We see no residual trend with host age for the larger set. The host extinction cut removes the youngest galaxies as expected. The residual trend in the low host extinction subset is indicated by the dashed (red) line in the top plot. This trend has a significance of 2.1σ when compared with a distribution of correlations created when drawing 22 data points at random from the total sample 10,000 times.

H09 for deriving Ni masses and host metallicities.

3.4.1. Host Metallicities from Host Masses

H09 point out the difficulties in using spectroscopic line indices (Hamuy et al. 2000; Gallagher et al. 2008) or line ratios (Gallagher et al. 2005) to derive host metallicities to compare with SN Ia intrinsic luminosity or ^{56}Ni mass. Instead, they use the mass-metallicity relation of Tremonti et al. (2004). There are drawbacks to using this trend, which uses gas-phase metallicity, on all galaxies regardless of gas content. However, as pointed out in H09, there are good reasons for assuming this method is accurate enough for the purposes here (see H09, §3.2). We also repeat the caution from H09, that gas-phase metallicity is not the same as stellar metallicity, although they should be correlated (Fernandes et al. 2005). The apparent scatter in the Tremonti relationship is 0.1 dex which we add in quadrature to our errors when we estimate host metallicities. For consistency, we use host masses derived using the same technique applied in H09 and outlined in Sullivan et al. (2006).

We test the use of the Tremonti et al. (2004) mass-metallicity relation for low- z SN hosts by using the SN host metallicity data from Prieto et al. (2008). Their metallicities come from the SDSS as were those used in Tremonti et al. (2004). A comparison of the two results should reveal any biases in the low- z SN host sample. We have sufficiently good photometry for 170 SN hosts in the Prieto et al. (2008) catalog to perform this comparison. These data are also from the GLGA and will be published in Seibert et al. (2009). We do not place any requirements on the SNe in this sample other than that they have metallicities from Prieto et al. (2008) and that we have enough host photometry for a good SED fit. Figure 10 plots our SED-fit stellar masses against the

metallicities from Prieto et al. (2008). The relation from Tremonti et al. (2004) is shown as the solid line and the low-metallicity extension from Lee et al. (2006) is shown as the dashed line for comparison with the host data. The hosts with masses in the range $9 < \log M_* < 11$ appear to follow the Tremonti relation well, while hosts with $\log M_* < 9$ appear to follow the Lee relation, albeit with a sparser sampling.

We will be looking for average trends in ^{56}Ni mass with metallicity, therefore we want to be sure to sample the range of the scatter in the Tremonti relation so we are not biased toward one side or the other by observational effects. The host mass range for our low- z SNe Ia with well-fit light curves is primarily above $\log M_* = 9$ (see Figure 6). The scatter in the mass-metallicity relation for SN hosts appears to be well sampled above this mass, thus we should be less subject to observational biases that could be in effect below this host mass. We also need not use the low-metallicity extension from Lee et al. (2006) for our base sample of SN Ia hosts.

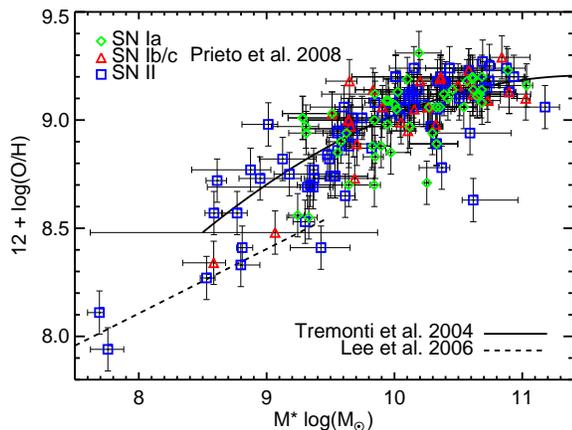


FIG. 10.— Metallicity plotted as a function of PEGASE.2 SED fitted stellar mass for the hosts from Prieto et al. (2008) for which there was sufficient integrated host photometry to derive good SED fits. The relation from Tremonti et al. (2004) is indicated as the solid thick line and the low-metallicity extension from Lee et al. (2006) is indicated by the dashed line. Open (green) diamonds indicate hosts of SNe Ia, open (red) triangles indicate hosts of SNe Ib/c, and open (blue) squares indicate hosts of SNe II. The Prieto sample is poorly sampled at masses below $\sim 10^{9.2} M_\odot$. This is most likely due to selection effects from host-targeted local SN surveys that rarely target low-mass hosts. Most of our local well fit SNe are in hosts with $M_* \gtrsim 10^9 M_\odot$ (see Figure 6), thus we should not be overly influenced by this selection effect.

3.4.2. SN Ia ^{56}Ni Masses

H09 present a technique for estimating the SN Ia ^{56}Ni mass using Arnett’s Rule (Arnett 1979, 1982) to convert bolometric luminosity and SN Ia rise time to ^{56}Ni mass. We use the identical technique on the well-sampled light curves for our local SNe Ia. Because the estimate of the bolometric luminosity assumes the luminosity distance is accurate, we use only the SNe in the Hubble flow (see Table 1).

The lowest stretch SNe Ia are also excluded from this experiment because they produce inaccurate ^{56}Ni masses using our technique. As previously stated, low luminosity SNe Ia deviate from the normal trend of lumi-

nosity versus light-curve shape (Garnavich et al. 2004; Taubenberger et al. 2008), thus the bolometric luminosity is inaccurate. The SN Ia spectral templates currently available are also not representative of SNe Ia with $s < 0.75$ (Conley et al. 2008). A higher-order relation between light-curve shape and luminosity that properly accounts for low-luminosity SNe Ia and better spectral templates would allow us to use these SNe to test this model (Gonzalez-Gaitan et al. 2009). They will be quite useful because the strongest effect from metallicity should occur in the faintest SNe. For our plots of Ni mass, we exclude objects with $s < 0.75$. To preclude inaccuracies due to excessive extinction we include only SNe Ia with $C < 0.4$. The final sample of 74 SNe that pass these cuts is presented in Table 3, where the host, and the derived ^{56}Ni masses and host metallicities are listed.

Figure 11 shows the SN Ia ^{56}Ni mass as a function of host metallicity for our final sample. We use the same symbol coding as in Figure 4. We have taken the average Ni mass in three bins of metallicity starting at $12 + \log(\text{O}/\text{H}) = 8.6$ and having a width of 0.2 and plot these as filled (purple) triangles. The model from Timmes et al. (2003) altered for thin disk O/Fe (H09) is overplotted as a thick dashed line. The average trend seen in H09 is shown by the open circles. Although our data are statistically consistent with no trend, they are also consistent with the H09 trend and the Timmes model.

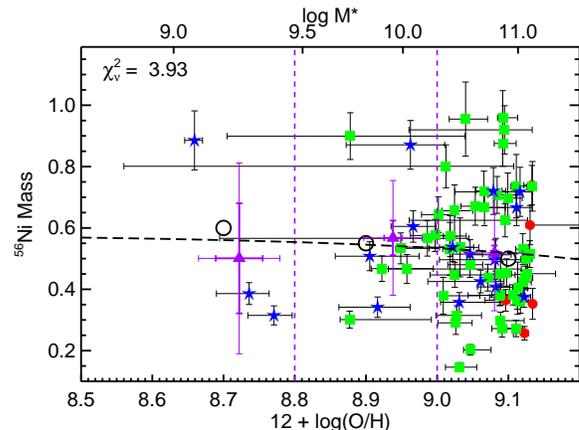


FIG. 11.— SN Ia ^{56}Ni mass as a function of metallicity. The host sSFR is coded as in Figure 4. The average ^{56}Ni masses in (O/H) bins of 0.2 dex starting at $12 + \log(\text{O}/\text{H}) = 8.6$ (vertical dashed purple lines) are plotted as filled (purple) triangles. The thin errorbars are the RMS in each bin, while the thick errorbars are the errors in the mean within the bin. The data from H09 is shown as the large open (black) circles. The expected trend with metallicity from the Timmes et al. (2003) model, altered for thin disk O/Fe, is plotted as the thick dashed line. The average trend is consistent with that predicted by Timmes et al. (2003) as was found by H09 for an intermediate redshift sample of SNe Ia ($0.2 < z < 0.75$).

3.4.3. Color Correction

H09 apply a correction for the intrinsic SN Ia light curve color to their derived Ni masses, after a color cut has been applied. Fainter SNe Ia are redder due to the intrinsic color-luminosity relation, or line-of-sight extinction, or both. Our Ni mass calculation requires the intrinsic luminosity, corrected for the dimming effects of

dust, but not corrected for color-luminosity. Since the correction is based on the observed and not intrinsic color of the SN there is bound to be some error in assuming that there is no dust dimming (uncorrected Ni-masses) or that the color-luminosity relation is purely due to dust (color corrected Ni-masses). Both will be inaccurate at some level. One way to examine this is to calculate and compare the χ^2/DOF or χ_ν^2 that results when comparing the corrected and uncorrected data to the Timmes model. When we apply the color correction, we find a mild improvement in the scatter as shown in Figure 12. When compared with the Timmes model, the χ_ν^2 is only marginally improved going from 3.93 (uncorrected) to 3.89 (color corrected).

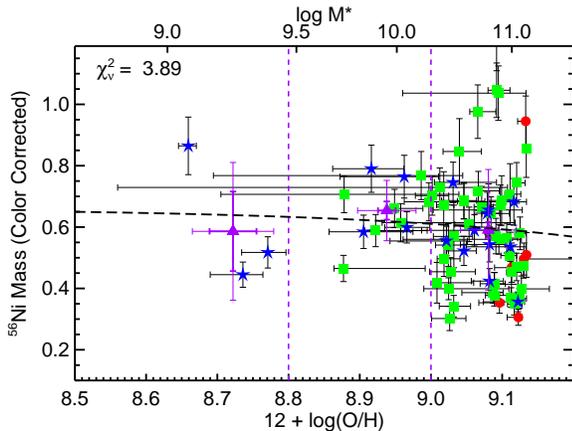


FIG. 12.— SN Ia ^{56}Ni mass as a function of metallicity with a correction for SN color applied (see text). The host sSFR is coded as in Figure 4. The Timmes model is plotted as the thick dashed line. The color correction was not as effective as in the H09 sample and only reduces the χ_ν^2 from 3.93 to 3.89, when comparing the data to the Timmes model.

These values of χ_ν^2 greater than one imply that our errors do not reflect the expected scatter of the data points around the Timmes model, assuming it is correct. H09 found that in order to achieve $\chi_\nu^2 = 1$, they had to add in quadrature an extra $0.16 M_\odot$ to their errors in Ni mass for their uncorrected masses. We also find we need an extra error of $0.16 M_\odot$ in our uncorrected ^{56}Ni masses to achieve $\chi_\nu^2 = 1$. Another measure of the effectiveness of the color correction is to see how much this extra error is reduced by the correction. For H09 it was significantly reduced from 0.16 (uncorrected) to 0.12 (color corrected) M_\odot . For our data there was only a marginal decrease in the extra scatter required for the color corrected ^{56}Ni masses (0.162 to 0.158 M_\odot).

One possible explanation of this difference is the homogeneous light curve photometry available for the H09 sample in contrast to the heterogeneous photometry available for our sample. The color cuts in H09 were made in a multi-color space that included wavelength regions that were not available for all the SNe in the local sample (see their §3.3). Consequently, our color cut here was in a single color (B-V) and may not have been as effective at removing outliers in the color corrected ^{56}Ni mass. We point out that there is an obvious reduction by the color correction in the scatter of Ni mass for lower

metallicity hosts and the biggest outliers in the color-corrected plot (Figure 12) are near the highest metallicity. The fact that the color correction provides any improvement at all is consistent with the idea that at least some part of the color-luminosity relation is due to host dust extinction.

3.4.4. SN Ia ^{56}Ni Mass in Three Stretch Bins

Figure 13 shows the uncorrected Ni mass versus metallicity divided in three stretch bins (see H09, Figure 10). We do not sample as large a range of metallicity as H09, nonetheless, we do see similar features. The range of host metallicity for the higher stretch SNe Ia is roughly double that for the lower stretch group. We expect the average ^{56}Ni mass to increase with stretch due to stretch-luminosity, but we also see that the scatter in the masses is lowest in the lowest stretch group. The stretch cut in the low-stretch bin undoubtedly limits the Ni mass in this bin (see H09, §4.5), but at the low Ni-mass end our local sample should be less subject to Malquist biases. Another effect is the color cut of $C < 0.4$, which precludes the reddest, faintest and therefore lowest Ni-mass SNe. Nonetheless, the trends are suggestive and support the idea that lower-stretch SNe Ia appear in higher metallicity hosts.

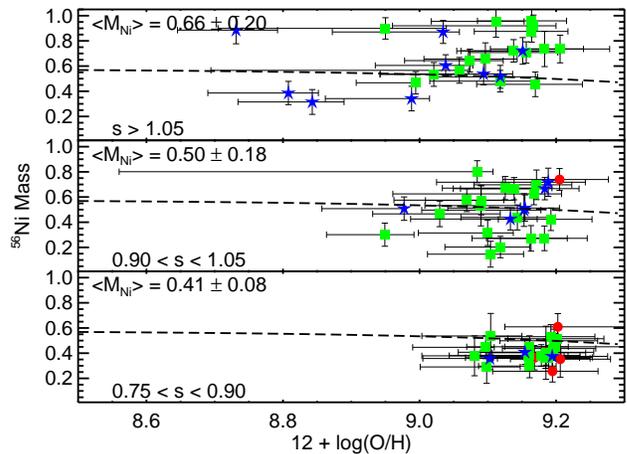


FIG. 13.— SN Ia ^{56}Ni mass (uncorrected) as a function of metallicity for three stretch bins, as annotated. The host sSFR is coded as in Figure 4. The Timmes model is plotted as the thick dashed line. The average ^{56}Ni mass increases with stretch across the three bins, as expected. The range of host metallicity and the scatter in Ni mass decreases toward lower stretch as was seen in the higher redshift sample from H09.

3.5. Host Age versus SN Ia ^{56}Ni Mass

Due to the age-metallicity degeneracy, we present plots of ^{56}Ni mass against luminosity-weighted age for the uncorrected (Figure 14) and color corrected (Figure 15) Ni masses. We use the same symbol scheme as previous plots. We use a linear regression analysis to measure and test the correlation between age and Ni mass. The color correction slightly steepens the slope of the fit (from -0.04 ± 0.02 to -0.06 ± 0.02), and does improve the correlation from -0.10 (uncorrected) to -0.16 (color corrected).

The age-Ni mass slopes reported in H09 for their sample are significantly steeper: -0.15 ± 0.03 (uncorrected)

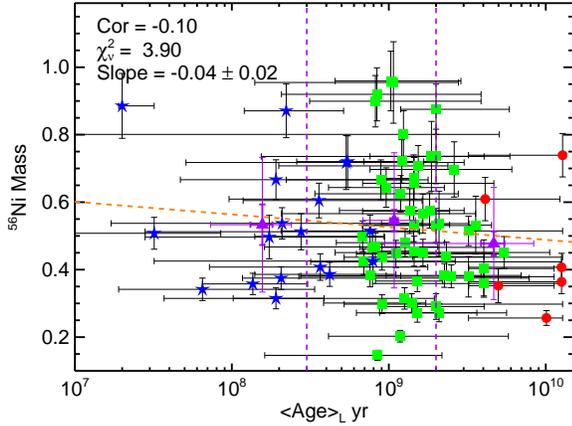


FIG. 14.— SN Ia ^{56}Ni mass, with no color correction, as a function of luminosity weighted host age. The host sSFR is coded as in Figure 4. The average ^{56}Ni mass in three bins dividing the sample at luminosity-weighted ages of 3×10^8 and 2×10^9 yr (vertical dashed purple lines) are plotted as filled purple triangles. The thin errorbars are the RMS in each bin while the thick errorbars are the error in the mean in each bin. A linear fit is shown as the dashed orange line and the results of the correlation are annotated on the plot. The correlation is weaker and the slope is shallower than what was seen in the H09 sample (see their Figure 8). The linear trend with age produces a χ^2_{ν} similar to the Timmes model. Combined with H09, Figure 8, we note that SNe Ia with ^{56}Ni masses greater than $\sim 0.5 M_{\odot}$ are rare in hosts with luminosity-weighted ages greater than $\sim 3 \times 10^9$ yr.

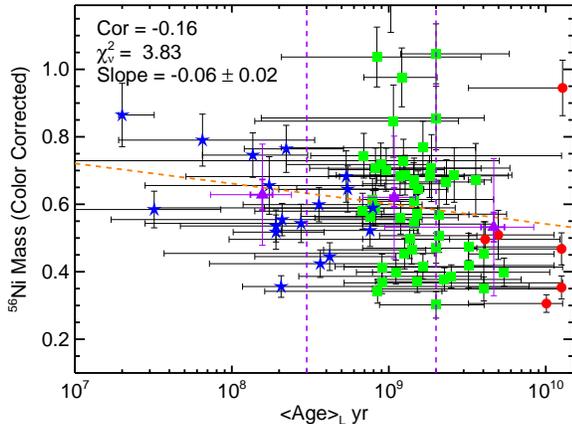


FIG. 15.— SN Ia ^{56}Ni mass, corrected for SN color, as a function of luminosity weighted host age. The host sSFR is coded as in Figure 4 and the bins for the averages are the same as in Figure 14. The color correction increases the correlation and marginally changes the slope. The rarity of higher Ni-mass SNe Ia in the oldest hosts is still apparent (see Figure 14 and H09, Figure 9).

and -0.11 ± 0.03 (color corrected). The correlations are also larger: -0.38 (uncorrected) and -0.37 (color corrected). Examining Figures 8 and 9 in H09, we see that a large part of the correlation seems to originate in hosts with luminosity-weighted ages beyond 3 Gyr which seem incapable of producing SNe with ^{56}Ni masses above $0.5 M_{\odot}$ in the uncorrected plot (their Figure 8) and above $0.6 M_{\odot}$ in the corrected plot (their Figure 9). Below 3 Gyr, there appears to be very little correlation between age and Ni mass. We see a similar feature in our data as

well (Figures 14 and 15), although not quite as strong. This feature implies a threshold that is reached at a luminosity-weighted age of ~ 3 Gyr, rather than a uniform trend with age.

4. DISCUSSION

The implications of these results must be tempered by the limitations of the techniques. In particular, the SED-fitting method employed was optimized for higher redshift galaxies and we do not employ infra-red data to constrain the SEDs of our low-redshift hosts. Our aim is to compare hosts at high and low redshifts using the same methods. We focus here on trends and this comparison and will pursue more precise determinations of host parameters with infra-red data and more detailed methods in a future paper.

We reiterate that the ages estimated with the PEGASE.2 fits are luminosity weighted and the masses are current stellar masses. The star formation rates only measure current activity based on UV and optical data and have not been modified by FIR data to account for dust. The global host extinctions are derived by extinguishing the entire SED and do not account for variable thermal components in the IR.

4.1. Host Properties at High and Low Redshift

Our results show features similar to those presented in higher redshift host studies using the same methods (Sullivan et al. 2006; Sullivan 2009, H09). Low stretch SNe Ia appear to be rare in high sSFR hosts and the average stretch of SNe Ia in the lowest sSFR hosts is smaller than in hosts with higher sSFR (see Figure 4). The average stretch of SNe Ia in hosts younger than ~ 2 Gyr is close to 1 and for older hosts the average stretch appears to drop below 1 (see Figure 5). H09 were the first to plot host stellar mass against stretch, but their sample does not include SNe Ia with $s < 0.75$ (see their Figure 4). The interesting feature in our plot (Figure 6) is the limited mass range for lower-stretch SNe. This feature is present in the H09 figure, but it is less pronounced due to having fewer low-stretch SNe in their sample.

The apparent stellar mass threshold of $10^{10} M_{\odot}$ for the hosts of low-luminosity SNe Ia is fairly robust against selection effects. Our sample includes many lower-mass hosts within which it is easier to detect low-luminosity SNe. That these SNe prefer higher-mass hosts supports the idea that they arise from an older progenitor population. Figures 4, 5 and 6 show that the one high sSFR host (E445-G66) of a SNe Ia with $s < 0.80$ (SN1993H) is also massive, implying that the recently formed populations are mixed with the older populations that could be responsible for the low-stretch SN Ia. The correlation between mass and metallicity implies that metallicity could play a role as well (see §3.4). We see no low-mass, high sSFR hosts of low-stretch SNe Ia and we also see few older, higher-mass hosts of the highest stretch SNe Ia. A method that can more accurately assess the star formation history of these hosts and the relative contributions from older and younger populations will produce a more definitive constraint on the progenitors of the low-luminosity SNe Ia (Schawinski 2009).

Our plot of SN peak color versus host age illustrates the difficulty of disentangling the source of the color (see

Figure 7) since it shows no correlation. We are also challenged in our comparison with higher redshift samples by the bias against SNe with very red peak colors intrinsic to cosmology surveys (see §2.1). Even in the local universe, our sample of six SNe Ia with $\mathcal{C} > 0.7$ is too small for definitive interpretation. We point out one common feature of our figure and Figure 3 in Sullivan et al. (2009): very red SNe appear to be rare in hosts less than 0.5 Gyr in luminosity-weighted age.

Based on Figure 11, we conclude that the model of Timmes et al. (2003), relating metallicity to produced ^{56}Ni mass, appears consistent with the derived host metallicities and SN Ia ^{56}Ni masses for SNe in the local universe as well as at higher redshifts (H09), although the local data are also consistent with no trend in these properties. The degeneracy between age and metallicity makes it difficult to decide which is the more important factor in determining the ^{56}Ni mass of hosted SNe Ia. The apparent transition at a luminosity-weighted host age of 3 Gyr (see Figure 14 and H09, Figure 8) could place interesting constraints on progenitor scenarios given that this implies a fairly long delay time for low ^{56}Ni mass, hence low-luminosity and low-stretch SNe Ia.

4.2. Host Properties and Cosmological Fitting

It is encouraging that the features in plots of light curve properties against host properties in local hosts are similar when compared with the SNLS sample out to $z = 0.75$ (Astier et al. 2006; Howell et al. 2007; Sullivan et al. 2009), although a detailed comparison using Hubble diagram residuals will be needed to determine if any significant evolution of the population can be detected (Conley et al. 2009; Sullivan et al. 2009). These trends do signal potential systematics for SN cosmology: as we extend the samples beyond redshift $z = 0.75$, we know the mix of host populations will move toward lower-mass, higher sSFR galaxies.

Our examination of low-extinction hosts suggests a trend in brightness with host age such that SNe Ia in older hosts appear brighter after correcting for stretch and color (see Figure 9). While the correlation we see is hardly significant (2.1σ), it does suggest that luminosity-weighted host age might be an interesting parameter to explore in cosmological fits of SNe Ia. A larger sample of SNe Ia from low-extinction hosts would allow a more definitive test of this correlation.

While there is no apparent trend in a similar plot in Gallagher et al. (2008, see their Figure 9), we see a trend that is marginally significant. There are several possible explanations for this difference. Since the hosts in the Gallagher et al. (2008) sample were selected morphologically, their sample may not be as free from dust extinction as originally thought. In addition, their corrections for light curve shape are derived with a method that assumes a Milky Way-like dust is responsible for residual color after a quadratic color-decline relation is subtracted (see comparison of fitting methods in Conley et al. 2008), which may not be appropriate for these hosts. The trend we see in Figure 9 is in the right sense to be caused by residual line-of-sight extinction since we would expect the SNe hosted by the oldest galaxies to have the least non-local extinction.

These data suggest that either the SNe Ia or the dust

produced in galaxies of different host properties are different or both. Hicken et al. (2009b) showed a weak trend in the Hubble residuals when using samples divided into three host morphology bins, after removing the reddest SNe. An exploration based on more physical host properties could provide a more accurate assessment of this effect (Sullivan et al. 2009). Further progress may be made by extending host photometry farther into the infra-red to help constrain the host dust content and by more detailed fitting of host star-formation history to constrain the most recent episode of star formation (e.g., Schawinski 2009).

5. SUMMARY

We gathered the host integrated photometry for 168 well-observed SNe Ia and correlated the derived host properties with the hosted SN Ia light curve properties. We used the derived average host extinction to isolate a set of SNe Ia arising in hosts with low extinction. We used the methods outlined in H09 to compare host metallicity with SN Ia ^{56}Ni mass.

With our low-redshift sample, we have corroborated trends in SN Ia light curve shape with host properties observed in higher-redshift SN Ia samples of similar size (Sullivan et al. 2006). Specifically, we find that the higher the sSFR of the host galaxy, the brighter and slower the SNe Ia that are produced, on average. We find that the typical stretch of the SNe Ia produced begins to drop after the host galaxy reaches a luminosity-weighted age of 1 Gyr. We find that low stretch ($s < 0.80$) SNe Ia in the local universe prefer hosts with stellar masses above $10^{10} M_{\odot}$. This is unlikely to be due to a Malmquist bias, since lower luminosity SNe are easier to detect against lower mass, lower luminosity hosts. We find that fainter, faster SNe Ia do indeed prefer more massive and presumably higher metallicity hosts, supporting the link between the production of ^{56}Ni and progenitor metallicity.

We suggest that it is likely that SN Ia cosmological fitting could be improved by adding a parameter characterizing the host properties to account for the impact of galaxy evolution on the evolution of SNe Ia with redshift. The trends in host properties with SN Ia light curve properties support this idea and suggest that perhaps host age would be a good property to use.

Our sample shows no obvious trend in host properties with SN Ia peak color, except that all the reddest SNe Ia appear in hosts with ages around 1 Gyr. We find that the reddest SNe Ia appear in hosts with a range of derived host extinctions.

Our low host extinction sample suggests a residual trend of SN peak absolute brightness and host age such that SNe Ia in older hosts appear brighter after their peak brightnesses are corrected for light-curve shape and color (i.e. stretch-color corrected). One possible explanation for this is that the color correction leaves a small residual brightness trend that is due to extinction and that older hosts contain less dust than younger hosts. This residual trend is consistent with the one reported in Hicken et al. (2009b), who find that SNe Ia in early-type galaxies are brighter, after shape and color correction, than SNe Ia appearing in late-type hosts.

For SNe Ia with $s > 0.75$ and $\mathcal{C} < 0.4$, we find local SNe Ia host metallicities and SN Ia ^{56}Ni masses are con-

sistent with the model of Timmes et al. (2003) altered for thin disk O/Fe as in H09, but with an additional scatter of $0.16 M_{\odot}$. The local data are also consistent with no trend. The failure of the SN color correction to significantly reduce the scatter in the Ni-masses is likely due to the heterogeneous photometry of the local SNe Ia. In spite of this, the trend of average ^{56}Ni mass and the scatter with stretch is very similar to that seen in H09.

We examined the trend in ^{56}Ni mass with host luminosity-weighted age and found a shallower slope and a weaker correlation than shown in H09 for a higher redshift sample ($0.2 < z < 0.75$). We point out what appears to be a threshold luminosity-weighted age of 3 Gyr apparent in our sample and the H09 sample, above which a host is less likely to produce SNe Ia with ^{56}Ni masses greater than $\sim 0.5 M_{\odot}$. Below this age threshold, there appears to be little correlation between ^{56}Ni mass and host age.

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TABLE 1
 SN IA SAMPLE

SN	HOST	cz (log km s ⁻¹)	stretch (s) ¹	m_{Bmax}	Status ²
1980N	NGC1316	3.26	0.83±0.01	12.42±0.01	
1981B	NGC4536	3.26	0.88±0.02	11.95±0.01	C
1981D	NGC1316	3.26	0.86±0.04	12.54±0.05	C
1990N	NGC4639	3.00	1.09±0.01	12.70±0.02	C
1991T	NGC4527	3.24	1.07±0.02	11.44±0.03	
1991U	IC4232	3.98	1.01±0.05	16.33±0.08	H, C
1991ag	IC4919	3.64	1.10±0.03	14.43±0.08	H, C
1992A	NGC1380	3.27	0.81±0.01	12.54±0.01	C
1992P	IC3690	3.88	1.16±0.12	16.05±0.03	H, C
1992ag	E508-G67	3.88	0.98±0.05	16.27±0.05	H, C
1992bc	E300-G09	3.76	1.07±0.01	15.10±0.01	H, C
1992bl	E291-G11	4.12	0.78±0.03	17.31±0.05	L, H, C
1992bo	E352-G57	3.75	0.75±0.01	15.73±0.02	L, H, C
1993H	E445-G66	3.86	0.66±0.06	16.79±0.03	L, H, C
1993ae	UGC1071	3.76	0.77±0.03	16.24±0.05	L, H
1994M	NGC4493	3.84	0.80±0.03	16.21±0.04	H, C
1994Q	PGC0059076	3.95	1.07±0.07	16.36±0.08	H
1994S	NGC4495	3.66	1.04±0.04	14.78±0.02	H, C
1994ae	NGC3370	3.11	1.05±0.01	12.95±0.02	C
1995D	NGC2962	3.29	1.11±0.01	13.26±0.03	C
1995E	NGC2441	3.54	0.93±0.03	16.69±0.02	C
1995al	NGC3021	3.19	1.08±0.03	13.32±0.02	C
1995bd	UGC3151	3.66	1.03±0.01	15.28±0.20	H, C
1996C	M+08-25-47	3.91	1.07±0.04	16.63±0.03	H, C
1996X	NGC5061	3.31	0.85±0.02	12.99±0.03	C
1996Z	NGC2935	3.36	0.92±0.08	14.32±0.08	C
1996ai	NGC5005	3.00	1.11±0.03	16.90±0.01	R, C
1996bo	NGC673	3.72	0.87±0.01	15.84±0.03	H, C
1996bv	UGC3432	3.70	1.05±0.05	15.32±0.05	H, C
1997E	NGC2258	3.60	0.81±0.02	15.11±0.05	C
1997Y	NGC4675	3.68	0.91±0.07	15.28±0.07	H
1997bp	NGC4680	3.40	0.97±0.02	13.91±0.02	C
1997bq	NGC3147	3.45	0.88±0.02	14.33±0.04	
1997br	E576-G40	3.32	0.91±0.04	13.31±0.08	
1997cw	NGC105	3.72	1.13±0.04	15.99±0.06	H, C
1997do	UGC3845	3.48	0.94±0.03	14.27±0.04	C
1998V	NGC6627	3.72	1.00±0.04	15.08±0.08	H, C
1998ab	NGC4704	3.91	0.94±0.02	16.06±0.03	H, C
1998aq	NGC3982	2.97	0.95±0.01	12.31±0.01	C
1998bu	NGC3368	2.97	0.95±0.02	12.10±0.01	C
1998de	NGC252	3.70	0.57±0.02	17.36±0.03	L, H, C
1998dh	NGC7541	3.43	0.91±0.02	13.86±0.04	
1998dm	M-01-04-44	3.29	1.07±0.07	14.64±0.08	C
1998dx	UGC11149	4.21	0.80±0.04	17.55±0.04	L, H, C
1998ec	UGC3576	3.77	0.98±0.07	16.17±0.11	H
1998eg	M+01-57-14	3.87	0.92±0.06	16.11±0.05	H
1998es	NGC632	3.50	1.14±0.01	13.84±0.02	C
1999X	CGCG180-22	3.88	0.91±0.08	16.09±0.13	H, C
1999aa	NGC2595	3.65	1.13±0.01	14.73±0.02	H, C
1999ac	NGC6063	3.45	0.98±0.01	14.12±0.02	C
1999cc	NGC6038	3.97	0.78±0.02	16.78±0.01	L, H, C
1999cl	NGC4501	3.33	0.96±0.02	14.87±0.02	R, C
1999cp	NGC5468	3.45	1.01±0.03	13.95±0.02	C
1999da	NGC6411	3.57	0.55±0.02	16.60±0.03	L, C
1999dk	UGC1087	3.66	0.97±0.03	14.83±0.03	H, C
1999dq	NGC976	3.64	1.12±0.01	14.42±0.05	H, C
1999ee	IC5179	3.53	1.07±0.01	14.85±0.01	C
1999gd	NGC2623	3.74	0.92±0.06	16.93±0.03	H, C
2000E	NGC6951	3.12	1.06±0.01	12.85±0.15	C
2000ca	E383-G32	3.86	1.08±0.03	15.57±0.03	H, C
2000ce	UGC4195	3.69	1.02±0.03	17.06±0.04	H, C
2000cx	NGC524	3.38	0.87±0.01	13.06±0.04	
2000dk	NGC382	3.72	0.74±0.01	15.36±0.03	L, H, C
2000fa	UGC3770	3.80	1.03±0.03	15.86±0.05	H, C

TABLE 1 — *Continued*

SN	HOST	cz (log km s ⁻¹)	stretch (s) ¹	m_{Bmax}	Status ²
2001N	NGC3327	3.80	0.94±0.05	16.58±0.04	H, C
2001V	NGC3987	3.66	1.13±0.01	14.62±0.02	H, C
2001ay	IC4423	3.96	1.59±0.03	16.77±0.03	H
2001az	UGC10483	4.09	1.08±0.06	16.93±0.04	H, C
2001ba	M-05-28-01	3.95	1.01±0.02	16.21±0.03	H, C
2001da	NGC7780	3.71	0.89±0.24	15.48±0.09	H, C
2001el	NGC1448	3.07	0.98±0.01	12.78±0.01	C
2001en	NGC523	3.69	0.90±0.05	15.05±0.08	H, C
2001ep	NGC1699	3.59	0.88±0.02	14.90±0.02	C
2001fe	UGC5129	3.61	1.09±0.03	14.70±0.02	H, C
2001ie	UGC5542	3.97	0.80±0.05	16.78±0.05	L, H, C
2002bf	CGCG266-031	3.86	0.94±0.03	16.33±0.05	H, C
2002bo	NGC3190	3.11	0.94±0.01	13.96±0.02	C
2002cd	NGC6916	3.49	1.07±0.02	15.50±0.16	C
2002ck	UGC10030	3.97	1.06±0.06	16.28±0.12	H, C
2002cr	NGC5468	3.45	0.94±0.02	14.20±0.02	C
2002de	NGC6104	3.92	1.06±0.07	16.68±0.02	H, C
2002dj	NGC5018	3.45	0.94±0.01	13.93±0.04	C
2002dp	NGC7678	3.54	0.96±0.04	14.61±0.03	C
2002er	UGC10743	3.43	0.88±0.01	14.26±0.06	C
2002es	UGC2708	3.73	1.01±0.05	16.16±0.10	H
2002fk	NGC1309	3.35	0.99±0.03	13.13±0.03	C
2002ha	NGC6962	3.63	0.87±0.03	14.71±0.05	H, C
2002he	UGC4322	3.87	0.80±0.02	16.23±0.04	H, C
2002hu	M+06-06-12	4.04	1.04±0.02	16.63±0.02	H, C
2002hw	UGC52	3.72	0.76±0.03	16.65±0.05	L, H
2002jy	NGC477	3.77	1.12±0.04	15.74±0.03	H, C
2003U	NGC6365A	3.93	0.79±0.02	16.51±0.03	L, H, C
2003W	UGC5234	3.78	0.99±0.02	15.88±0.02	H, C
2003cg	NGC3169	3.09	0.95±0.01	15.79±0.02	R, C
2003du	UGC9391	3.28	1.01±0.01	13.47±0.01	C
2003fa	M+07-36-33	4.07	1.15±0.01	16.71±0.02	H, C
2003hu	A191131+7753	4.35	1.16±0.06	18.46±0.14	H, C
2003hx	NGC2076	3.33	0.84±0.06	14.86±0.05	C
2003ic	M-02-02-86	4.22	0.75±0.05	17.66±0.08	L, H
2003kc	M+05-23-37	4.00	0.84±0.04	17.14±0.05	H
2003kf	M-02-16-02	3.35	1.04±0.03	13.27±0.13	C
2004L	M+03-27-38	3.99	0.93±0.04	17.30±0.05	H, C
2004as	A112539+2249	3.97	1.06±0.04	16.96±0.02	H, C
2004eo	NGC6928	3.67	0.87±0.00	15.08±0.05	H, C
2004fu	NGC6949	3.44	0.87±0.01	14.25±0.16	C
2005am	NGC2811	3.40	0.70±0.05	13.66±0.03	L, C
2005eq	M-01-09-06	3.95	1.18±0.01	16.30±0.03	H, C
2005hc	M+00-06-03	4.14	1.02±0.03	17.36±0.02	H, C
2005hk	UGC272	3.59	0.88±0.00	15.95±0.01	
2005iq	M-03-01-08	4.01	0.87±0.02	16.82±0.02	H, C
2005ir	A011643+0047	4.36	1.44±0.11	18.42±0.03	H, C
2005kc	NGC7311	3.65	0.93±0.02	15.61±0.06	H, C
2005ke	NGC1371	3.16	0.64±0.04	14.80±0.03	L
2005ki	NGC3332	3.76	0.80±0.01	15.55±0.03	H, C
2005ls	M+07-07-01	3.80	1.13±0.03	16.25±0.04	H, C
2005mc	UGC4414	3.88	0.65±0.04	17.23±0.02	L, H
2005ms	UGC4614	3.88	1.06±0.02	16.16±0.02	H, C
2005mz	NGC1275	3.72	0.61±0.02	16.42±0.07	L, H
2006N	M+11-08-12	3.63	0.76±0.02	15.08±0.04	L, H, C
2006S	UGC7934	3.98	1.12±0.01	16.86±0.01	H, C
2006X	NGC4321	3.20	0.96±0.01	15.22±0.01	R, C
2006ac	NGC4619	3.83	0.87±0.02	16.18±0.02	H, C
2006ak	A110932+2837	4.05	0.84±0.04	17.24±0.10	H, C
2006al	A103929+0511	4.31	0.78±0.04	18.44±0.05	L, H, C
2006ar	M+11-13-36	3.83	0.92±0.03	16.48±0.01	H, C
2006ax	NGC3663	3.70	1.00±0.01	15.04±0.02	H, C
2006az	NGC4172	3.97	0.86±0.01	16.49±0.01	H, C
2006bk	M+06-33-20	4.17	1.10±0.03	17.00±0.06	H
2006bq	NGC6685	3.82	0.84±0.02	16.15±0.04	H, C

TABLE 1 — *Continued*

SN	HOST	cz (log km s ⁻¹)	stretch (s) ¹	m_{Bmax}	Status ²
2006br	NGC5185	3.87	0.81±0.04	18.95±0.03	R, H, C
2006bt	M+03-41-04	3.98	1.01±0.02	16.95±0.02	H, C
2006bw	A143356+0347	4.00	0.72±0.03	17.56±0.04	L, H, C
2006cc	UGC10244	3.99	1.03±0.01	17.81±0.01	H, C
2006cj	A125924+2820	4.31	1.25±0.11	18.14±0.03	H, C
2006cm	UGC11723	3.69	1.05±0.04	17.94±0.03	R, H, C
2006cp	UGC7357	3.82	1.07±0.02	15.99±0.02	H, C
2006ef	NGC809	3.73	0.84±0.04	15.16±0.17	H
2006ej	NGC191A	3.79	0.81±0.04	15.79±0.03	H, C
2006en	M+05-54-41	3.98	1.00±0.04	16.75±0.05	H, C
2006gj	UGC2650	3.93	0.65±0.07	17.61±0.05	L, H, C
2006gz	IC1277	3.85	1.26±0.01	15.84±0.03	H
2006hb	M-04-12-34	3.66	0.67±0.06	15.48±0.05	L, H, C
2006kf	UGC2829	3.80	0.71±0.03	15.83±0.10	L, H, C
2006le	UGC3218	3.72	1.11±0.01	14.80±0.17	H, C
2006mo	M+06-02-17	4.05	0.75±0.03	17.44±0.03	L, H, C
2006nz	A005629-0113	4.06	0.60±0.10	18.08±0.04	L, H
2006ob	UGC1333	4.25	0.71±0.01	18.23±0.02	L, H, C
2006on	A215558-0104	4.32	0.99±0.09	18.41±0.07	H, C
2006os	UGC2384	3.99	0.91±0.03	17.61±0.06	H
2006sr	UGC14	3.86	0.84±0.02	16.13±0.04	H, C
2006te	A081144+4133	3.98	1.03±0.05	16.51±0.04	H, C
2007F	UGC8162	3.85	1.07±0.01	15.90±0.01	H, C
2007O	UGC9612	4.03	0.93±0.04	16.78±0.05	H, C
2007R	UGC4008	3.96	0.82±0.02	16.65±0.05	H
2007S	UGC5378	3.62	1.13±0.01	15.82±0.01	H, C
2007ae	UGC10704	4.29	1.15±0.03	17.78±0.04	H, C
2007af	NGC5584	3.20	0.96±0.01	13.16±0.02	C
2007ap	M+03-41-03	3.67	0.54±0.07	15.86±0.03	L, H, C
2007au	UGC3725	3.79	0.66±0.03	16.51±0.03	L, H, C
2007bc	UGC6332	3.80	0.84±0.03	15.89±0.02	H, C
2007bd	UGC4455	3.97	0.82±0.01	16.57±0.02	H, C
2007bm	NGC3672	3.27	0.92±0.01	14.48±0.02	C
2007bz	IC3918	3.81	1.17±0.03	16.67±0.03	H, C
2007cg	E508-G75	4.00	0.82±0.05	18.28±0.07	H, C
2007ci	NGC3873	3.74	0.75±0.01	15.92±0.02	L, H, C
2007sr	NGC4038	3.21	0.99±0.02	12.76±0.04	C
2008af	UGC9640	4.00	0.85±0.03	16.78±0.08	H, C
2008bf	NGC4055	3.86	1.04±0.02	15.73±0.02	H, C

¹ stretch errors for SNe with $s < 0.7$ have been multiplied by 3.0² L - low stretch: $s < 0.80$, R - red: $C > 0.7$, H - in Hubble flow: $z > 0.0133$, C - eligible for cosmology fitting

TABLE 2
 HOST PROPERTIES

HOST	T^1	Age $-$	$\langle \text{Age} \rangle_L$ (log yr)	Age $+$	M^*-	$\langle M^* \rangle$ (log M_\odot)	M^*+	sSFR $-$	$\langle \text{sSFR} \rangle$ (log yr $^{-1}$)	sSFR $+$	$E(B-V)_H$ (Mag)
A005629-0113	...	9.81	10.07	10.09	10.48	10.62	10.66	-12.00	-12.00	-12.48	0.00
A011643+0047	...	8.45	8.71	8.93	10.06	10.15	10.20	-9.50	-9.31	-9.06	0.20
A081144+4133	...	8.83	9.13	9.39	10.26	10.31	10.43	-10.32	-9.84	-9.41	0.10
A103929+0511	...	9.51	9.51	10.05	10.23	10.26	10.55	-12.00	-10.76	-10.71	0.05
A110932+2837	...	8.86	8.97	9.56	10.51	10.52	10.75	-10.77	-10.09	-9.75	0.30
A112539+2249	...	8.15	8.62	9.05	9.15	9.28	9.36	-9.58	-9.23	-8.79	0.05
A125924+2820	...	8.15	9.03	9.44	10.30	10.42	10.62	-10.69	-10.13	-8.35	0.20
A143356+0347	...	9.51	9.62	10.09	9.95	10.03	10.27	-12.00	-11.03	-10.69	0.05
A191131+7753	...	8.47	8.66	8.95	10.86	10.90	10.98	-9.52	-9.27	-9.07	0.25
A215558-0104	...	9.51	9.55	9.65	10.25	10.30	10.38	-12.00	-11.62	-10.73	0.00
CGCG180-22	...	9.04	9.09	9.28	10.10	10.13	10.21	-10.41	-10.14	-10.06	0.15
CGCG266-031	3.0	9.15	9.37	9.61	10.55	10.62	10.74	-10.83	-10.49	-10.17	0.05
E291-G11	1.0	8.90	9.73	10.03	11.16	11.81	12.27	-11.56	-10.85	-9.27	0.30
E300-G09	5.0	8.50	8.91	9.59	9.19	9.72	10.42	-10.56	-9.62	-8.50	0.10
E352-G57	-1.5	9.16	9.99	10.11	10.93	12.13	12.27	-12.00	-11.18	-9.49	0.60
E383-G32	4.5	8.08	8.35	8.71	9.70	10.04	10.27	-9.35	-9.00	-8.49	0.25
E445-G66	1.9	7.00	7.76	9.15	10.13	10.51	11.07	-10.37	-8.63	-8.14	0.55
E508-G67	5.0	8.81	8.90	9.67	9.92	10.02	11.06	-11.11	-9.88	-8.95	0.00
E508-G75	3.9	7.83	9.36	10.03	10.13	10.76	11.35	-11.57	-10.31	-8.14	0.15
E576-G40	7.0	7.00	7.30	8.00	8.59	8.71	8.90	-8.97	-8.66	-8.33	0.45
IC1277	6.0	7.00	7.26	8.21	10.29	10.44	10.88	-9.26	-8.66	-8.15	0.55
IC3690	4.0	8.83	9.16	9.67	10.24	10.34	10.48	-10.86	-9.98	-9.35	0.15
IC3918	4.4	8.01	8.28	8.59	9.27	9.38	9.45	-9.19	-8.94	-8.62	0.15
IC4232	3.8	7.71	8.73	9.48	10.34	11.04	11.51	-11.03	-9.32	-8.15	0.40
IC4423	4.1	8.88	9.18	9.72	10.49	10.62	10.77	-10.89	-10.02	-9.49	0.15
IC4919	7.9	7.28	7.30	7.51	9.04	9.07	9.10	-8.72	-8.66	-8.58	0.25
IC5179	4.0	8.68	8.86	8.97	10.21	10.72	10.83	-10.50	-9.41	-8.86	0.35
M+00-06-03	-2.0	8.88	9.13	9.68	10.45	10.54	10.77	-10.85	-10.15	-9.42	0.15
M+01-57-14	6.0	8.15	9.94	10.05	10.57	11.32	11.79	-11.79	-11.10	-8.16	0.20
M+03-27-38	5.0	7.57	9.10	9.29	10.20	10.35	10.56	-10.49	-9.90	-8.55	0.05
M+03-41-03	-1.0	8.78	8.92	10.09	10.29	10.35	10.69	-12.00	-9.96	-9.49	0.30
M+03-41-04	0.0	8.50	8.84	10.10	11.02	11.09	11.35	-12.00	-9.93	-8.54	0.40
M+05-23-37	3.0	8.54	8.82	9.10	10.53	10.67	10.76	-9.99	-9.38	-9.07	0.25
M+05-54-41	5.0	7.45	8.24	8.97	10.23	10.69	11.23	-9.78	-8.92	-8.25	0.45
M+06-02-17	4.2	9.51	9.51	9.86	10.80	10.82	10.99	-12.00	-10.75	-10.71	0.05
M+06-06-12	4.8	7.00	9.09	9.70	8.83	10.27	10.93	-10.95	-9.68	-8.00	0.15
M+06-33-20	-3.1	8.76	9.56	10.09	11.49	11.59	11.83	-12.00	-10.82	-9.46	0.10
M+07-07-01	4.2	7.28	7.81	8.53	9.66	9.86	10.04	-9.16	-8.63	-8.35	0.30
M+07-36-33	...	8.32	8.92	9.59	10.03	10.81	11.42	-11.03	-9.57	-8.44	0.20
M+08-25-47	...	8.83	9.16	9.39	9.96	9.99	10.16	-10.37	-9.98	-9.38	0.05
M+11-08-12	-2.4	9.51	10.10	10.11	10.59	10.82	10.88	-12.00	-12.00	-10.71	0.00
M+11-13-36	3.3	8.58	9.15	9.24	9.67	9.72	10.18	-10.69	-10.18	-8.81	0.00
M-01-04-44	6.0	7.54	7.98	8.15	8.91	8.93	9.20	-8.91	-8.63	-8.41	0.25
M-01-09-06	6.0	8.84	9.08	9.31	10.51	10.58	10.77	-10.53	-10.14	-9.36	0.05
M-02-02-86	-2.0	9.51	9.51	10.09	11.67	11.70	11.97	-12.00	-10.76	-10.69	0.10
M-02-16-02	3.0	7.00	9.03	10.05	9.29	9.58	10.50	-11.39	-10.13	-8.07	0.05
M-03-01-08	2.5	8.79	9.05	9.70	10.17	10.34	10.95	-10.63	-9.81	-9.20	0.05
M-04-12-34	-2.6	8.43	8.79	9.61	10.64	10.95	11.14	-11.34	-9.81	-8.16	0.55
M-05-28-01	3.7	7.67	8.28	9.15	10.48	10.98	11.41	-10.32	-8.94	-8.17	0.45
NGC105	1.5	7.04	9.11	9.59	9.47	10.87	11.42	-11.04	-9.81	-8.03	0.15
NGC1275	-1.6	8.51	9.05	9.36	11.15	11.24	11.48	-10.49	-9.81	-9.03	0.15
NGC1309	4.0	7.82	8.47	8.84	9.78	9.94	10.64	-10.27	-9.08	-8.00	0.25
NGC1316	-1.9	10.07	10.11	10.11	11.75	11.79	11.80	-12.00	-12.00	-13.75	0.00
NGC1371	1.0	9.51	9.70	9.77	10.79	10.90	11.02	-10.60	-10.46	-10.15	0.15
NGC1380	-1.9	10.05	10.10	10.10	11.27	11.33	11.36	-12.00	-12.00	-13.27	0.05
NGC1448	5.9	8.95	9.37	9.40	10.22	10.69	10.73	-10.19	-9.90	-9.40	0.20
NGC1699	3.0	8.71	9.59	9.62	10.12	10.37	10.41	-10.46	-10.22	-9.24	0.00
NGC191A	-2.0	9.51	10.10	10.11	10.76	11.00	11.02	-12.00	-12.00	-10.70	0.00
NGC2076	-0.6	9.18	9.67	9.77	9.62	10.02	10.14	-11.12	-10.41	-9.62	0.15
NGC2258	-2.0	8.43	8.91	9.48	11.43	11.52	12.33	-11.01	-10.03	-8.40	0.50
NGC2441	3.1	8.71	9.12	9.40	10.31	10.66	10.87	-10.04	-9.70	-9.28	0.25
NGC252	-1.0	8.94	9.43	9.90	10.87	11.25	11.90	-11.44	-10.56	-9.43	0.25
NGC2595	4.5	8.30	9.19	9.38	10.62	10.72	10.96	-10.55	-10.23	-8.84	0.00
NGC2623	PEC	8.21	8.93	9.34	10.27	10.38	10.52	-10.50	-10.05	-8.43	0.10

TABLE 2 — *Continued*

HOST	T ¹	Age−	<Age> _L (log yr)	Age+	M*−	<M*> (log M _⊙)	M*+	sSFR−	<sSFR> (log yr ^{−1})	sSFR+	E(B−V) _H (Mag)
NGC2811	1.0	8.90	9.26	9.56	11.28	11.51	11.71	-10.81	-10.37	-9.85	0.45
NGC2935	3.2	7.58	8.25	10.05	10.68	10.94	11.31	-12.00	-8.61	-8.23	0.55
NGC2962	-1.0	8.63	8.93	10.11	10.10	10.20	10.50	-12.00	-10.05	-8.95	0.35
NGC3021	4.3	8.71	8.90	9.19	9.76	9.87	9.97	-9.88	-9.45	-8.23	0.30
NGC3147	3.9	8.98	9.34	9.63	10.75	11.31	11.66	-11.02	-10.15	-9.41	0.15
NGC3169	1.2	9.04	9.19	9.43	10.64	10.70	10.79	-10.59	-10.23	-10.07	0.15
NGC3190	1.0	9.51	10.11	10.11	10.55	10.83	10.89	-12.00	-12.00	-10.62	0.00
NGC3327	3.0	8.79	9.18	9.75	10.66	10.77	10.96	-10.92	-10.02	-9.36	0.15
NGC3332	-3.0	8.72	8.96	10.11	11.11	11.15	11.51	-12.00	-10.07	-9.29	0.30
NGC3368	2.0	8.80	9.22	9.83	11.09	11.26	11.54	-11.01	-10.30	-9.59	0.25
NGC3370	5.3	8.00	8.55	9.08	9.55	9.69	9.83	-9.97	-9.15	-8.63	0.25
NGC3663	3.5	8.53	9.08	9.44	10.04	10.81	11.21	-11.00	-9.60	-8.54	0.30
NGC3672	5.0	7.92	8.65	8.97	10.04	10.23	10.34	-9.52	-9.23	-8.58	0.35
NGC382	-5.0	9.13	10.05	10.05	10.17	11.54	11.56	-11.80	-11.24	-9.63	0.45
NGC3873	-5.0	9.51	10.00	10.11	10.93	11.13	11.23	-12.00	-12.00	-10.68	0.00
NGC3982	3.0	8.69	9.25	9.40	9.86	10.02	10.10	-10.12	-9.75	-9.20	0.10
NGC3987	3.0	9.08	9.30	9.77	10.69	10.78	10.97	-10.93	-10.43	-10.08	0.25
NGC4038	8.8	7.43	7.79	8.15	9.85	10.05	10.17	-8.95	-8.64	-8.38	0.30
NGC4055	-5.0	9.51	10.11	10.11	11.11	11.39	11.42	-12.00	-12.00	-10.66	0.00
NGC4172	2.7	9.51	9.51	10.11	11.26	11.28	11.54	-12.00	-10.76	-10.72	0.10
NGC4321	4.4	8.57	9.17	9.26	10.74	10.81	11.09	-10.57	-10.22	-9.01	0.00
NGC4493	-4.0	8.90	9.61	10.11	10.93	11.04	11.23	-12.00	-10.88	-9.87	0.10
NGC4495	1.9	8.92	9.16	9.19	10.39	10.50	10.53	-9.72	-9.68	-9.54	0.30
NGC4501	3.0	8.36	8.79	9.24	10.91	11.03	11.23	-10.53	-9.45	-8.86	0.35
NGC4527	3.8	8.62	9.30	9.56	10.60	10.78	10.91	-10.79	-10.43	-9.03	0.20
NGC4536	4.5	9.00	9.18	9.36	10.42	10.47	10.68	-10.45	-10.02	-9.38	0.05
NGC4619	3.1	8.44	8.88	9.08	10.87	10.92	11.11	-9.97	-9.56	-9.01	0.15
NGC4639	3.8	9.16	9.32	9.46	10.07	10.14	10.24	-10.42	-10.12	-9.63	0.05
NGC4675	3.0	8.56	9.62	9.77	10.14	10.42	10.43	-10.93	-10.26	-9.10	0.10
NGC4680	PEC	8.18	8.76	9.05	10.11	10.24	10.33	-9.92	-9.32	-8.81	0.30
NGC4704	3.5	8.35	8.95	9.25	10.55	10.59	10.82	-10.31	-9.67	-8.95	0.15
NGC477	5.0	8.31	8.88	9.19	10.32	10.46	10.57	-9.89	-9.43	-8.95	0.20
NGC5005	4.0	8.79	8.96	9.00	10.90	10.93	11.01	-10.19	-10.07	-9.61	0.30
NGC5018	-5.0	9.51	9.51	9.61	11.21	11.34	11.37	-11.24	-10.76	-10.60	0.10
NGC5061	-5.0	9.51	9.61	9.61	10.78	10.85	10.87	-10.92	-10.88	-10.67	0.05
NGC5185	3.0	9.18	9.37	9.61	10.89	10.94	11.06	-10.83	-10.49	-10.20	0.10
NGC523	4.2	8.95	9.32	9.51	10.23	10.38	10.53	-10.71	-10.45	-9.56	0.00
NGC524	-1.0	10.05	10.10	10.10	11.82	11.89	11.91	-12.00	-12.00	-13.82	0.05
NGC5468	6.0	7.18	7.43	7.87	9.39	9.48	9.77	-9.04	-8.65	-8.25	0.20
NGC5584	6.0	8.27	8.71	8.92	9.69	9.82	9.87	-9.49	-9.31	-8.94	0.15
NGC6038	5.0	8.91	9.18	9.34	10.94	10.99	11.03	-10.29	-10.02	-9.60	0.10
NGC6063	6.0	8.56	8.95	9.21	9.85	9.89	10.11	-10.11	-9.67	-9.10	0.10
NGC6104	...	8.92	9.16	9.19	10.71	10.83	10.86	-9.90	-9.68	-9.51	0.20
NGC632	-1.5	7.89	9.24	9.35	9.83	10.00	10.13	-10.57	-10.34	-8.43	0.00
NGC6365A	5.9	8.43	9.40	9.59	10.44	10.74	10.87	-10.26	-9.93	-9.02	0.05
NGC6411	-5.0	9.51	9.70	10.11	10.79	10.91	11.13	-12.00	-12.00	-10.67	0.05
NGC6627	3.0	8.88	9.42	9.78	10.68	10.85	11.08	-12.00	-10.56	-9.44	0.05
NGC6685	-3.0	8.72	8.96	10.11	10.73	10.76	11.12	-12.00	-10.07	-9.31	0.30
NGC673	5.0	7.81	8.13	8.48	10.24	10.37	10.78	-9.39	-8.84	-8.32	0.30
NGC6916	4.0	8.69	9.59	9.62	10.43	10.68	10.72	-10.46	-10.22	-9.22	0.00
NGC6928	2.0	8.15	8.83	8.97	11.05	11.17	11.40	-10.37	-9.93	-8.26	0.40
NGC6949	5.0	7.00	8.17	9.37	9.09	9.64	10.42	-10.02	-8.86	-8.04	0.10
NGC6951	4.0	7.88	8.63	9.61	10.84	10.96	11.32	-10.94	-9.05	-8.18	0.35
NGC6962	2.0	8.83	9.30	9.56	10.96	11.09	11.23	-10.79	-10.43	-9.67	0.10
NGC7311	2.0	8.94	9.32	9.56	10.84	10.97	11.09	-10.78	-10.45	-9.94	0.10
NGC7541	4.0	7.18	8.48	9.70	10.29	10.66	10.98	-10.82	-9.11	-8.22	0.35
NGC7678	5.0	8.67	8.74	9.05	10.04	10.40	10.75	-10.37	-9.38	-8.81	0.15
NGC7780	2.0	8.94	9.30	9.61	10.22	10.35	10.48	-10.87	-10.43	-9.93	0.15
NGC809	-1.0	9.51	9.51	9.61	10.66	10.70	10.78	-10.99	-10.75	-10.69	0.05
NGC976	5.0	8.66	9.01	9.46	10.72	10.78	10.99	-10.60	-9.77	-9.16	0.15
PGC0059076	...	8.64	8.98	9.13	9.80	9.84	10.00	-10.08	-9.72	-9.14	0.05
UGC10030	3.0	8.88	9.27	9.77	10.89	10.97	11.19	-10.94	-10.16	-9.47	0.10
UGC10244	3.7	8.62	9.07	9.76	10.41	10.46	10.65	-10.88	-9.86	-9.15	0.25
UGC10483	3.8	8.41	8.74	9.22	10.55	10.68	10.82	-10.40	-9.38	-8.87	0.35
UGC10704	3.5	8.19	9.30	9.61	11.25	11.44	11.61	-10.84	-10.43	-8.44	0.10

TABLE 2 — *Continued*

HOST	T ¹	Age−	<Age> _L (log yr)	Age+	M*−	<M*> (log M _⊙)	M*+	sSFR−	<sSFR> (log yr ^{−1})	sSFR+	E(B−V) _H (Mag)
UGC1071	...	8.21	9.51	9.62	10.29	10.35	11.90	-12.00	-10.76	-8.00	0.00
UGC10743	1.0	7.88	8.51	9.77	9.93	10.56	11.18	-11.42	-9.13	-8.08	0.70
UGC1087	5.0	8.69	9.26	9.48	10.04	10.20	10.29	-10.14	-9.77	-9.25	0.10
UGC11149	...	7.98	9.61	10.09	11.17	11.72	12.59	-12.00	-12.00	-8.00	0.05
UGC11723	3.0	9.01	9.50	9.87	10.22	10.41	10.58	-11.05	-10.65	-9.62	0.15
UGC1333	3.0	8.99	9.05	9.24	11.20	11.25	11.34	-10.34	-10.13	-10.01	0.20
UGC14	5.5	7.86	8.56	9.25	10.52	10.70	10.87	-10.20	-9.16	-8.59	0.35
UGC2384	3.9	7.40	8.54	9.74	10.66	11.59	12.43	-11.42	-9.15	-8.00	0.70
UGC2650	2.0	8.50	9.98	10.11	10.66	11.96	12.33	-12.00	-11.11	-8.07	0.45
UGC2708	-2.0	8.95	9.86	10.00	10.81	11.08	11.19	-12.00	-12.00	-9.98	0.00
UGC272	6.5	8.00	8.46	8.99	9.41	9.54	9.68	-9.69	-9.07	-8.68	0.20
UGC2829	-2.0	8.99	9.86	10.01	10.73	10.97	11.06	-12.00	-12.00	-10.09	0.00
UGC3151	5.0	8.70	9.21	9.77	10.53	10.65	10.92	-12.00	-10.06	-9.16	0.10
UGC3218	3.0	7.00	7.34	7.75	10.09	10.19	10.35	-8.93	-8.65	-8.39	0.35
UGC3432	6.0	8.32	9.22	9.55	9.17	10.15	10.45	-10.65	-9.76	-8.43	0.20
UGC3576	3.0	8.81	9.14	9.77	10.09	10.57	11.43	-11.32	-10.16	-8.94	0.10
UGC3725	-3.0	9.49	9.99	10.10	11.28	12.24	12.34	-12.00	-11.18	-9.80	0.50
UGC3770	10.0	7.00	7.51	7.93	9.65	9.82	10.10	-9.11	-8.65	-8.18	0.40
UGC3845	3.9	7.00	7.43	7.65	9.16	9.30	9.39	-8.89	-8.65	-8.41	0.20
UGC4008	0.0	7.57	9.30	9.58	10.85	10.98	11.12	-10.80	-10.43	-8.54	0.10
UGC4195	3.0	8.65	9.21	9.25	10.35	10.50	10.54	-9.97	-9.72	-9.23	0.25
UGC4322	-5.0	8.07	8.31	9.89	10.64	11.12	12.02	-12.00	-8.61	-8.04	0.70
UGC4414	0.0	9.09	9.46	9.56	10.86	10.95	11.06	-10.78	-10.64	-10.11	0.10
UGC4455	1.0	8.66	9.22	9.67	10.61	10.76	10.90	-10.78	-9.96	-9.17	0.15
UGC4614	...	7.23	8.32	9.21	9.96	10.32	10.49	-9.94	-8.97	-8.31	0.30
UGC5129	1.0	8.62	8.98	9.35	10.11	10.22	10.34	-10.59	-9.72	-9.09	0.15
UGC52	5.0	8.19	9.32	9.49	10.24	10.38	10.56	-10.74	-10.45	-8.76	0.00
UGC5234	5.0	7.51	8.90	9.42	10.15	10.55	10.80	-10.62	-9.44	-8.34	0.25
UGC5378	3.0	7.20	8.28	9.21	9.52	9.88	10.09	-10.19	-8.94	-8.27	0.35
UGC5542	-5.0	9.51	9.61	10.05	10.92	10.99	11.20	-12.00	-10.88	-10.68	0.05
UGC6332	1.0	8.98	9.35	9.84	10.66	10.76	10.96	-12.00	-10.49	-9.64	0.10
UGC7357	5.0	8.21	8.91	9.31	9.82	9.88	10.19	-10.01	-9.62	-8.71	0.00
UGC7934	2.8	8.67	9.10	9.36	10.29	10.46	10.57	-10.33	-9.63	-9.19	0.20
UGC8162	6.0	8.07	8.56	9.21	9.93	10.06	10.21	-9.75	-9.18	-8.78	0.15
UGC9391	8.0	8.33	8.76	9.22	8.47	8.60	8.74	-9.78	-9.36	-8.96	0.05
UGC9612	5.0	7.45	8.44	9.20	10.43	10.70	10.89	-10.12	-9.06	-8.43	0.30
UGC9640	-5.0	9.51	9.70	10.05	11.39	11.48	11.65	-12.00	-12.00	-10.72	0.05

¹ Numerical type according to de Vaucouleurs (1959)

TABLE 3
 ^{56}Ni MASS AND HOST METALLICITY

SN	HOST	^{56}Ni Mass (M_{\odot})	Corrected ^{56}Ni Mass (M_{\odot})	Mass-Metal $12 + \log(\text{O}/\text{H})$	Prieto et al. 2006 $12 + \log(\text{O}/\text{H})$
1999gd	NGC2623	0.15±0.02	0.34±0.04	9.10	...
2006cc	UGC10244	0.20±0.02	0.68±0.06	9.12	...
2007ci	NGC3873	0.26±0.02	0.31±0.03	9.19	...
2001N	NGC3327	0.27±0.03	0.57±0.06	9.16	...
2005kc	NGC7311	0.27±0.03	0.51±0.05	9.18	...
2001da	NGC7780	0.29±0.04	0.30±0.04	9.10	...
2006bq	NGC6685	0.30±0.03	0.41±0.04	9.16	...
2006ar	M+11-13-36	0.30±0.03	0.47±0.04	8.95	9.09
2004L	M+03-27-38	0.31±0.03	0.45±0.05	9.10	...
2007bz	IC3918	0.31±0.03	0.52±0.05	8.84	...
2005ls	M+07-07-01	0.34±0.03	0.79±0.08	8.99	...
2008af	UGC9640	0.35±0.05	0.51±0.07	9.21	...
1996bo	NGC673	0.36±0.03	0.75±0.07	9.10	...
1994M	NGC4493	0.36±0.04	0.35±0.04	9.19	...
2006N	M+11-08-12	0.36±0.04	0.35±0.03	9.17	...
1999cc	NGC6038	0.37±0.03	0.37±0.03	9.18	...
2002he	UGC4322	0.37±0.03	0.36±0.03	9.19	...
2006al	A103929+0511	0.38±0.06	0.42±0.07	9.08	...
2003U	NGC6365A	0.38±0.03	0.39±0.03	9.16	...
2006ac	NGC4619	0.38±0.03	0.56±0.05	9.18	...
2007bc	UGC6332	0.38±0.04	0.38±0.03	9.16	...
2004as	A112539+2249	0.39±0.04	0.44±0.04	8.81	...
2001ie	UGC5542	0.40±0.09	0.45±0.10	9.19	...
2006ej	NGC191A	0.41±0.04	0.47±0.05	9.19	...
2006sr	UGC14	0.41±0.04	0.42±0.04	9.15	...
2006bt	M+03-41-04	0.42±0.04	0.74±0.07	9.19	...
2003W	UGC5234	0.43±0.04	0.59±0.05	9.13	...
2005ki	NGC3332	0.44±0.04	0.37±0.03	9.20	...
2002bf	CGCG266-031	0.44±0.04	0.67±0.07	9.14	...
2005iq	M-03-01-08	0.45±0.04	0.40±0.04	9.10	...
2007bd	UGC4455	0.45±0.04	0.42±0.04	9.16	...
1992bl	E291-G11	0.45±0.05	0.40±0.04	9.20	...
2002de	NGC6104	0.45±0.04	0.67±0.07	9.17	...
1992ag	E508-G67	0.47±0.05	0.61±0.06	9.03	...
2006cp	UGC7357	0.47±0.04	0.59±0.05	8.99	...
2006S	UGC7934	0.48±0.04	0.69±0.06	9.12	...
2006en	M+05-54-41	0.50±0.06	0.66±0.09	9.15	...
2004eo	NGC6928	0.50±0.05	0.58±0.05	9.20	...
2000fa	UGC3770	0.51±0.05	0.58±0.05	8.98	...
2007O	UGC9612	0.51±0.05	0.54±0.06	9.15	...
2002jy	NGC477	0.51±0.05	0.52±0.05	9.12	...
2006az	NGC4172	0.51±0.04	0.47±0.04	9.20	...
2002ha	NGC6962	0.53±0.05	0.47±0.04	9.19	8.94
2006on	A215558-0104	0.53±0.09	0.67±0.11	9.09	...
1996C	M+08-25-47	0.53±0.05	0.66±0.06	9.02	...
2001en	NGC523	0.54±0.10	0.57±0.10	9.10	...
2005ms	UGC4614	0.54±0.05	0.55±0.05	9.09	9.06
1996bv	UGC3432	0.57±0.06	0.77±0.08	9.06	...
2006te	A081144+4133	0.57±0.07	0.50±0.06	9.09	9.10
1999dk	UGC1087	0.58±0.05	0.68±0.06	9.07	...
2007F	UGC8162	0.60±0.05	0.60±0.05	9.04	9.05
1998dx	UGC11149	0.61±0.06	0.50±0.05	9.20	...
2006ax	NGC3663	0.62±0.05	0.56±0.05	9.17	...
2001fe	UGC5129	0.64±0.06	0.70±0.06	9.07	...
1992P	IC3690	0.66±0.08	0.54±0.07	9.10	...
1998ab	NGC4704	0.67±0.06	0.72±0.06	9.14	9.16
2001ba	M-05-28-01	0.67±0.06	0.54±0.05	9.18	...
1994S	NGC4495	0.67±0.06	0.61±0.06	9.13	...
1998V	NGC6627	0.70±0.08	0.69±0.08	9.17	...
1999aa	NGC2595	0.71±0.06	0.64±0.05	9.16	...
1991U	IC4232	0.72±0.08	0.68±0.08	9.19	...
2001az	UGC10483	0.72±0.08	0.65±0.07	9.15	...
2005eq	M-01-09-06	0.72±0.06	0.98±0.09	9.14	...

TABLE 3 — *Continued*

SN	HOST	^{56}Ni Mass (M_{\odot})	Corrected ^{56}Ni Mass (M_{\odot})	Mass-Metal $12 + \log(\text{O}/\text{H})$	Prieto et al. 2006 $12 + \log(\text{O}/\text{H})$
2007ae	UGC10704	0.74 ± 0.08	0.86 ± 0.09	9.21	...
2002ck	UGC10030	0.74 ± 0.10	0.71 ± 0.10	9.18	9.23
2008bf	NGC4055	0.74 ± 0.06	0.94 ± 0.08	9.21	...
2002hu	M+06-06-12	0.80 ± 0.07	0.73 ± 0.06	9.08	...
2000ca	E383-G32	0.87 ± 0.08	0.76 ± 0.07	9.03	...
2001V	NGC3987	0.88 ± 0.07	1.05 ± 0.09	9.16	...
1991ag	IC4919	0.89 ± 0.10	0.86 ± 0.09	8.73	...
1992bc	E300-G09	0.90 ± 0.08	0.71 ± 0.06	8.95	...
2003fa	M+07-36-33	0.92 ± 0.08	1.04 ± 0.09	9.17	...
2006cj	A125924+2820	0.95 ± 0.12	0.85 ± 0.11	9.11	...
1999dq	NGC976	0.96 ± 0.09	1.22 ± 0.11	9.16	...