



Erratum: “Cloud Structure of Three Galactic Infrared Dark Star-forming Regions from Combining Ground and Space-based Bolometric Observations” (2017, *ApJ*, 840, 22)

Yuxin Lin^{1,2}, Hauyu Baobab Liu³, James E. Dale⁴, Di Li^{1,5}, Gemma Busquet⁶, Zhi-Yu Zhang^{3,7}, Adam Ginsburg³, Roberto Galván-Madrid⁸, Attila Kovács⁹, Eric Koch¹⁰, Lei Qian^{1,5}, Ke Wang³, Steve Longmore¹¹, Hwei-Ru Chen¹², and Daniel Walker¹¹

¹ National Astronomical Observatories, Chinese Academy of Sciences, China

² Max-Planck-Institut für Radioastronomie, D-53121 Bonn, Germany; ylin@mpifr-bonn.mpg.de

³ European Southern Observatory (ESO), Karl-Schwarzschild-Str. 2, D-85748 Garching, Germany

⁴ Centre for Astrophysics Research, University of Hertfordshire, College Lane, Hatfield, AL10 9AB, UK

⁵ Key Laboratory of Radio Astronomy, Chinese Academy of Sciences, China

⁶ Institut de Ciències de l’Espai (IEEC-CSIC), Campus UAB, Carrer de Can Magrans S/N, 08193, Barcelona, Catalunya, Spain

⁷ Institute for Astronomy, University of Edinburgh, Royal Observatory, Blackford Hill, Edinburgh EH9 3HJ, UK

⁸ Instituto de Radioastronomía y Astrofísica, UNAM, Apdo. Postal 3-72 (Xangari), 58089 Morelia, Michoacán, México

⁹ California Institute of Technology 301-17, 1200 E California Boulevard, Pasadena, CA 91125, USA

¹⁰ Department of Physics, University of Alberta, 4-181 CCIS, Edmonton, AB T6G 2E1, Canada

¹¹ Astrophysics Research Institute, Liverpool John Moores University, 146 Brownlow Hill, Liverpool L3 5RF, UK

¹² Institute of Astronomy and Department of Physics, National Tsing Hua University, Hsinchu, Taiwan

Received 2017 June 2; published 2017 July 14

1. Introduction

Due to a mistake in proofreading, there is a sub-figure of Figure 11 missing in the paper, which corresponds to the second caption therein. The right figure and caption is below.

We also note that the reference empirical line of Molinari et al. (2008) in our Figure 7 is not accurately plotted, the slope of which should be ~ 1.13 – 1.27 . While we do not use it to discuss our results since it is derived in a generally lower range of mass and luminosity of small-scale star-forming regions, we refer to the original work for a detailed description of this reference line.

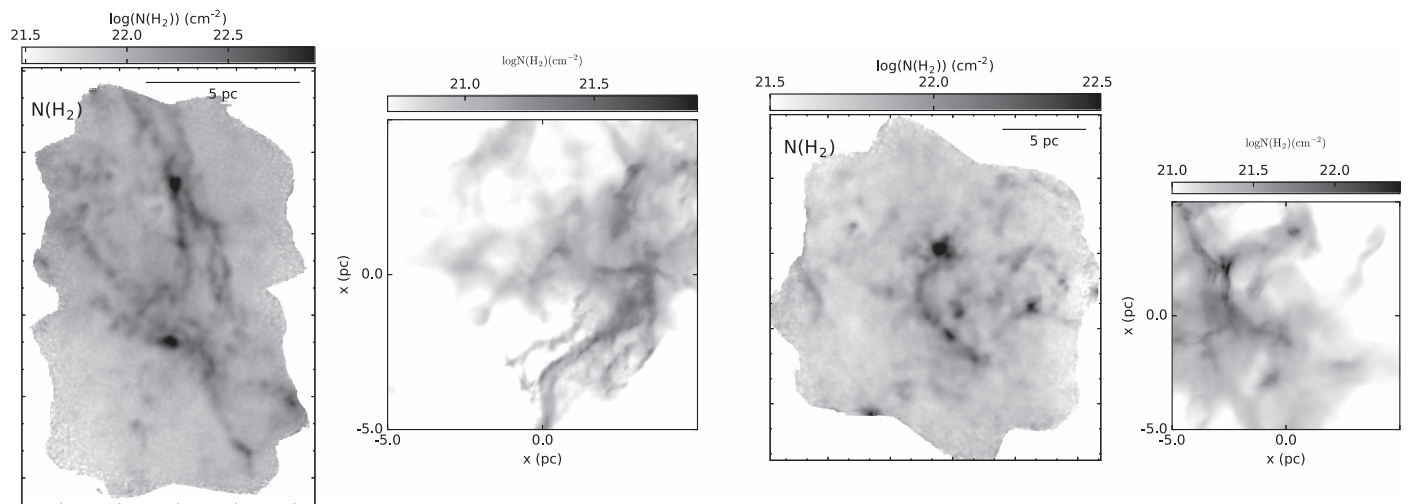


Figure 1. Left panel: G14.225–0.506 column density map and simulated image of time epoch 3.9 Myr. Right panel: G28.34+0.06 column density map and simulated image of time epoch 2.4 Myr. In each panel, the spatial scale of the simulated image is adjusted to be the same as the corresponding source.

Reference

Molinari, S., Pezzuto, S., Cesaroni, R., et al. 2008, *A&A*, 481, 345