



Supplementary Material for

HAT-P-26b: A Neptune-mass exoplanet with a well-constrained heavy element abundance

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Materials and Methods

We observed the transit of HAT-P-26b with HST STIS, and HST WFC3 over four new transit events, and analysed archival data from Spitzer IRAC 3.6 & 4.5 micron for two other transit events. HAT-P-26 is a quiet star with a consistent stellar activity level. This is shown by a consistent transit depth measured across all six observations with no indication of starspot occultations (Figure S1). If there are any unocculted spots the stellar spectrum is unlikely to mimic the observed water feature for starspots of any plausible temperature contrast (35). For our HST analysis we follow standard practice and discard the first orbit of each visit as it contains vastly different systematics to the subsequent orbits due to thermal settling of HST after initial target acquisition (e.g. 15-17, 35).

All the transit light curves were fitted with analytical transit models (36). We apply cosmic ray removal using IDL routines from (15), and account for stellar limb darkening using a 4-parameter limb-darkening law (37). We use marginalization across a series of stochastic models (38) to account for differing observatory and instrument configuration systematics. Each light curve is fit using a Lavenberg-Marquart (L-M) least-squares algorithm (39). The uncertainties on each data point were initially set to pipeline values dominated by photon and readout noise. After an initial fit the lightcurve, uncertainties are rescaled based on the standard deviation of the residuals, taking into account any underestimated errors calculated by the L-M reduction, which is then re-run to calculate the final fit to the data (16, 17). The L-M fit assumes the probability space around the best-fitting solutions are well described by a multivariate Gaussian distribution, as with previous studies (17, 40-42) we check this using a Markov Chain Monte Carlo (MCMC) analysis (43) and find excellent agreement with our data showing it is normally distributed. The transit depth uncertainty for each systematic model is then determined using the covariance matrix from the second L-M reduction. For each of the lightcurve fits we assume a normal likelihood function. We approximate the evidence-based weight based on the Akaike Information Criterion (AIC) and marginalize across all models to compute the desired light curve parameters and uncertainties (e.g. 16, 17, 40). Using marginalization across a grid of stochastic models allows us to therefore account for all tested combinations of systematics and obtain robust center of transit times from the band-integrated lightcurve, and transit depths for each spectroscopic lightcurve.

STIS observations and data analysis

We observed one transit of HAT-P-26b on 2016 January 25 using HST STIS. Our STIS observations were conducted using the G750L grating with a wide 52x2" slit to avoid slit losses. The STIS dataset was pipeline-reduced with the latest version of CALSTIS (12) and corrected for detector fringing using contemporaneous fringe flats. The spectral aperture extraction was done with the software package IRAF, using a 13-pixel-wide aperture with no background subtraction, which minimizes the

out-of-transit standard deviation of the band-integrated light curves. The extracted spectra were then Doppler-corrected to a common rest frame through cross-correlation, which helped remove sub-pixel wavelength shifts in the dispersion direction (16, 41). To correct for telescope and instrument systematics we use a grid of stochastic polynomial models accounting for a linear change in flux over the course of the whole observation (θ), orbit to orbit flux corrections in HST orbital period (ϕ) caused by thermal variations in the telescope, and linearly for the x and y detector position determined from linear spectral traces in IRAF. We use linear combinations of each parameter, θ , x , and y , and with ϕ up to a 4th order polynomial. This produces a grid of 80 stochastic models for the HST STIS systematics applied to each light curve which are then marginalized over to determine the transit parameters.

We divide the STIS wavelength range into multiple custom bins for different wavelength ranges and measure the R_p/R_* from each spectroscopic light curve following the same procedure as detailed for the band-integrated light curve. Figure S1a shows the spectroscopic light curves for seven custom bins offset for clarity and corrected by the highest weight systematic model, which contained corrections in θ , ϕ^3 , and y . We test a series of different bin widths and positions, however, due to the large scatter and relatively low resolution of the STIS observations we use broader bins in the red end of the wavelength regime to avoid increasing the uncertainty. The final scatter observed in the STIS transmission spectrum between 0.7-1.0 μm matches well with previously published optical data (19), which measures the spectrum at a higher resolution. Remaining outliers are likely the result of fringing which has been seen in various other datasets (e.g. 15, 16, 41). We show an expanded view of the optical portion of the HAT-P-26b transmission spectrum in Figure S2a, with the HST data, plotted with the previously observed ground-based results (19) for visual comparison.

WFC3 observations and data analysis

We observed HAT-P-26 over one visit with WFC3 G102 on 2016 August 16, and two visits with WFC3 G141, the first on 2016 March 12, the second on 2016 May 2. Observations for both spectroscopic grisms were conducted in forward spatial scan mode. Spatial scanning involves exposing the telescope during a forward slew in the cross-dispersion direction and resetting the telescope position to the top of the scan prior to conducting subsequent exposures. Scans with G102 were conducted at a scan rate of ~ 0.27 pixels per second with a final scan covering ~ 28 pixels in the cross-dispersion direction. For both G141 visits we used a scan rate of ~ 0.55 pixels per second with a final spatial scan covering ~ 62 pixels in the cross-dispersion direction on the detector.

We use the *IMA* (intermediate IR exposure) output files from the CalWF3 pipeline (13) which are calibrated using flat fields and bias subtraction. We extract the spectrum from each exposure by taking the difference between successive non-destructive reads. A top-hat filter is then applied around the target spectrum and all external pixels are set to zero which helps to remove cosmic rays (42). The image is then reconstructed by adding the individual reads back together. We extract the stellar spectrum from each exposure with an aperture of ± 14 pixels for G102 and ± 31 pixels for G141 around a centering

profile which was found to be consistent across the spectrum for each exposure for all three observations. For our WFC3 data analysis, we use a grid of stochastic models accounting for θ , ϕ , and shifts in the wavelength position of the spectrum caused by telescope pointing (δ_λ). We use combinations with and without θ , and with or without ϕ and δ_λ up to the 4th order. This results in a grid of 50 systematic models which we test against our data (see *17* for a table of systematic models).

We divide the WFC3 wavelength range for each grism into a series of bins and measure the R_p/R_* from each spectroscopic light curve following the same procedure as detailed for the band-integrated light curve. For each visit we test a range of bins widths and wavelength ranges for each visit and determine that the shape of the transmission spectrum is robust. Figure S1b shows each spectroscopic light curve for the G102 visit split into 7 bins ($\Delta\lambda \approx 0.048 \mu\text{m}$) between 0.8-1.1 μm and corrected using the highest weight systematic model, which corrects for θ , ϕ^3 and δ_λ^2 . We also show each spectroscopic light curve from both WFC3 G141 visits for 11 bins each ($\Delta\lambda \approx 0.046 \mu\text{m}$) between 1.13-1.67 μm , again corrected by the highest weight systematic model as defined by the data which includes systematic corrections for θ , ϕ^4 and δ_λ . We show the transmission spectrum measured from both visits with WFC3 G141 in Figure S2b with the combined transmission spectrum used for the final atmospheric interpretation.

Spitzer observations and data analysis

The Spitzer IRAC observations were conducted in subarray mode (32x32 pixels). Photometry was extracted from the basic calibrated data cubes, produced by the IRAC pipeline after dark subtraction, flat-fielding, linearization and flux calibration (see *15* and references therein). We performed outlier filtering for hot (energetic) or cold (low-count values) pixels in the data by examining the time series of each pixel and subtracted the background flux from each image. We measured the position of the star on the detector in each image incorporating the flux-weighted centroiding method using the background subtracted pixels from each image, for a circular region with a radius of 3 pixels centered on the approximate position of the star. We extracted photometric measurements from our data using both aperture photometry from a grid of apertures ranging from 1.5 to 3.5 pixels (in increments of 0.1) and time-variable aperture photometry, which resulted in the lowest white and random red noise correlated with the data points co-added in time for both channels. The best result was selected by measuring the flux scatter of the out-of-transit portion of the light curves for both channels after filtering the data for 5σ outliers with a width of 20 data points (*16*). To correct for systematic effects, we used a parametric models incorporating systematics associated with the x and y positions of the stellar centroid on the detector and linear trend in time (see *15* for details). We then use marginalization across all systematic models to calculate the resultant photometric transit depth for each channel. The photometric light curves for the 3.6 and 4.5 μm channels are shown in Fig. S1c binned in time by ~ 2 minutes. We again do not find a significant discrepancy between the results from this analysis in comparison to the Spitzer results presented in (*19*) with agreement at the 1.4σ and 1.8σ level for the 3.6 and 4.5 μm channels respectively.

Atmospheric retrieval model description

We use ATMO (20, 21) to compute the atmospheric metallicity, temperature and absorption species from the observations. ATMO is a code that computes the one dimensional (1D) T-P profile of an atmosphere in hydrostatic and radiative-convective equilibrium, and can also be used as a post-processing tool to compute the emission and transmission spectra from a given 1D atmospheric profile. The radiative transfer equation with scattering and irradiation is solved iteratively in its integral form, with correlated- k coefficients or line-by-line opacities, and the scheme has been benchmarked against the Met Office SOCRATES code (22). ATMO includes equilibrium chemistry with condensed species for a given elemental solar or non-solar composition (solved by minimization of the total Gibbs free energy), and out-of-equilibrium chemistry with both mixing and photo-chemistry (21). For the chemistry, we include 166 species with all gas phase species included in the C, H, N, O network (23) plus ~ 40 other gas/condensed phase species that are relevant for exoplanet and brown dwarf atmospheres. The code has been successfully applied to the study of Y, T, and L brown dwarfs and the atmosphere of exoplanets observed by direct imaging (20, 21).

The forward ATMO model was coupled to a Levenberg-Marquardt least squares minimizer (39) to initially find an optimal model fit, and a Differential Evolution MCMC analysis (43) was then used to measure the posterior distribution (Fig. S3 and Fig. S4) and determine the fit confidence intervals. We ran MCMCs with 10 to 12 chains each with 25,000 to 35,000 steps, and generally we found about 12,000 steps were needed for convergence which was monitored with the Gelman-Rubin statistic. The fit used k -coefficients with 500 bands, and confirmed with 5,000 bands uniformly spaced in wavenumber between 1 cm^{-1} and $50,000 \text{ cm}^{-1}$ for the best fit model. For the opacity sources, we include H_2 , He, H_2O , CO, CO_2 , CH_4 and NH_3 (see references in 20, 22). We also considered Na and K as potential sources of opacity, however the lack of evidence for their presence in the data (also see 19) only provided upper limits to their abundances.

To fit the data we consider a grid of 12 different atmospheric retrievals (labeled M1 to M12) using the ATMO model. For each model fit we assume a normal likelihood function on each datapoint. Each model requires the planetary radius to be a free parameter, in addition we fix or fit for the following parameters. The contribution of each to the number of free parameters is shown in parentheses.

M1: C/O fixed, isothermal, uniform scattering cloud (+1), metallicity (+1)

M2: C/O fitted (+1), isothermal, uniform scattering cloud (+1), metallicity (+1)

M3: free-chemistry (+4), isothermal, uniform scattering cloud (+1)

M4: C/O fixed, TP model (+3), uniform scattering cloud (+1), metallicity (+1)

M5: C/O fitted (+1), TP model (+3), uniform scattering cloud (+1), metallicity (+1)

M6: free-chemistry (+4), TP model (+3), uniform scattering cloud (+1)

M7-M12 then consider each of these models without the cloud opacity included which removes one free parameter from each model listed above. In the free-chemistry models (M3, M6, M9, and M12) the abundances of CO, CO_2 , CH_4 , and H_2O were freely fit. The C/O ratio for the free-chemistry model fit is then estimated from the four fit molecules (CO, CO_2 , CH_4 , H_2O) and does not take into account any other species. We

apply a uniform scattering cloud which assumes the cloud is uniform throughout the atmosphere and fit for the opacity of the cloud such that it becomes optically thick at the altitude defined by the measured transmission spectra.

To determine the metallicity from using the information from all retrieval models applied to the data we use Bayesian Model Averaging (BMA) to combine the posterior distributions over all the reasonable models weighted by their evidence. We use the Bayesian Information Criterion (BIC) to approximate the evidence of fit E_q for each model S_q given by the probability of the data D given the model q and is often referred to as the marginal likelihood (38), such as,

$$\ln E_q = \ln P(D|S_q) \approx -\frac{1}{2}BIC. \quad (S1)$$

Each evidence value is then transformed into a weighting such that each retrieval is assigned a percentage of the overall probability. The weight W_q is calculated as

$$W_q = E_q / \sum_{q=0}^{N_q} E_q, \quad (S2)$$

where N_q is the number of retrieval models fit. The weighting for each model is then used to calculate the weighted mean of all parameters of interest α_q and their associated uncertainty σ_{α_q} are used to calculate the marginalized parameter α_m and weighted uncertainty $\sigma(\alpha)$ (17):

$$\alpha_m = \sum_{q=0}^{N_q} (W_q \times \alpha_q), \quad (S3), \text{ and}$$

$$\sigma(\alpha) = \sqrt{\sum_{q=0}^{N_q} (W_q [(\alpha_q - \alpha_m)^2 + \sigma_{\alpha_q}^2])}. \quad (S4)$$

We show the derived metallicity and uncertainty for each retrieval in Table S2 with the number of free parameters, the calculated BIC and W_q . Using BMA we determine that HAT-P-26b has an atmospheric metallicity of $4.8 \times$ solar and 1σ limits of 0.8 to $26 \times$ solar.

Atmospheric retrieval model C/O constraints

In the models where carbon is an unconstrained free parameter, we find the retrieval prefers solutions in which the C/O and abundances of non-H₂O species are very low. This has a direct effect on the atmospheric scale height, as it can substantially reduce the atmospheric mean molecular weight compared to solar-abundance scaled compositions. With comparatively lower molecular weights, the transmission spectra retrievals can accommodate lower temperatures near 400 to 600 K (Fig. S5), which is substantially lower than expected from radiative-hydrodynamic equilibrium (Fig. 2). Like clouds, low temperatures can reduce the amplitude of the water feature but do so by directly reducing the atmospheric scale height. At these low temperatures near 600K, we would expect substantial CH₄ signatures in the transmission spectra, but none are found. In addition, the mean molecular weight of a $30 \times$ solar composition model is about 3 atomic mass units (amu), but with very low C/O atmospheres, the H₂O abundance has to increase up to about $100 \times$ solar to produce a mean of 3 amu. Thus, additionally fitting for the C/O gives extremely low metallicities ($0.001 \times$ solar) and C/O ratios ($C/O < 0.0001$) which are both about three orders of magnitude below the solar value and has the effect of removing

practically all molecular species in the atmosphere apart from those containing H, He and O. The H₂O abundance is then greatly increased extending the tails of the posterior distribution for the models fitting for C/O. In Figure S5 we illustrate from M2 the effects on the retrieved temperature and H₂O abundance posterior distributions when placing a lower-limit prior constraint on the C/O. We find a lower C/O of 0.01 provides a reasonable range to allow the ratio to vary, compared to the solar value of 0.56, and restricts the impact on the temperature. Previous exoplanet H₂O abundance results for WASP-43b and HAT-P-11b (e.g., 5, 11) have fixed the C/O at solar, which we have also assumed in model cases M1, M4, M7 and M10, as no lower-limit constraints could be placed on carbon-bearing molecules.

Atmospheric retrieval model results

We summarize our ATMO retrieval fits in Table S2. Overall we find that models that do not incorporate clouds (M7-M12) are disfavoured and have low overall weights (<2%) in our BMA results. In addition, the cloudy models which freely fit for the chemistry or TP profile give similar H₂O abundances, but generally have lower weights (~1%) as the BIC penalizes the models due to their increased number of free parameters. The freely fit TP profiles (M4-6, M10-12) were consistent with the isothermal approximation. Our best fitting models have isothermal profiles, scattering clouds, and a freely fit radius and metallicity, with C/O which is fixed to solar (M1 - 42.5%) or fit using a prior of C/O > 0.01 (M2 - 44.9%).

While the carbon-bearing molecules are not fully constrained by our data, the upper limits on the abundance of CO₂ gives a further indication that the atmosphere does not have a high metallicity. CO₂ is well known to be highly sensitive to the atmospheric metallicity (e.g., 25, 44). Over the wavelength range of our data, CO₂ has by far its strongest signature in the Spitzer 4.5 micron bandpass (Figure 1). CO also has a strong spectral signature at those wavelengths; because the two are unresolved in the spectra, the CO₂ abundance is highly degenerate with CO. However, the measurements can still rule out strong CO₂ signatures as would be expected at high metallicities near that of Neptune. We find CO₂ would be expected from equilibrium chemistry to be abundant at volume mixing ratio (VMR) of 2×10^{-3} , if the overall metallicity were $100 \times$ solar. However, our posterior distribution of the best-fitting free chemistry model (M3) limits CO₂ to several orders of magnitude below that value with a 68.2% upper bound limit on the CO₂ VMR of 5×10^{-6} , which is consistent with compositions near solar (Figure S4).

We additionally test that our results are independent of the STIS data, and the possibility of a wavelength dependence of the cloud deck, by fitting the transmission spectrum without the STIS measurements (i.e. fitting 0.8-5.0 μ m, WFC3+Spitzer). To address the wavelength dependence of the cloud opacity we assume Rayleigh scattering. Using the highest weight model, we retrieve a water abundance which is consistent with fitting the whole transmission spectrum. The red-most point of the WFC3 G141 measurements is better fit and is within 1σ of the model. This suggests that any cloud likely becomes transparent at these wavelengths, as is expected (e.g. 45, 46). However, as the measured water abundance is consistent with fitting the whole transmission spectrum

with a uniform cloud, described by a single parameter, a more complex wavelength dependent cloud model is not justified in this case.

In summary, using the BIC as an approximation of the evidence we show that M1 and M2 fit the data equally well and are the best fitting models. Using BMA we incorporate the derived metallicities of all tested models to produce a marginalised metallicity value and uncertainty coherent across all models. We additionally test this result using the SCARLET retrieval code (5) and find a 1σ bound of $1-25\times$ solar, which agrees well with the ATMO results and final marginalised value. The wide wavelength coverage we obtain with HST and Spitzer, compared with the single H_2O absorption feature detection using WFC3 G141 alone, acts to further constrain the clouds and place strong constraints on the metallicity by the models.

Atmospheric profile and cloud models

The best fitting atmospheric models for the measured transmission spectrum suggest the presence of an absorbing cloud deck at optical wavelengths, which we expect to be produced by a cloud composed of relatively large particles which scatter the light uniformly at multiple wavelengths, thereby hiding any atomic features. To estimate the possible cloud absorbers in the atmosphere of HAT-P-26b, we also calculate the three-dimensional (3D) temperature structure using the SPARC/MITgcm. The SPARC/MITgcm couples the MITgcm, a finite-volume code that solves the 3D primitive equations on a staggered Arakawa C grid (47) with a radiative transfer code SPARC that is a two-stream adaptation of a multi-stream radiative transfer code for solar system planets (48). The radiative transfer code employs the correlated-k method with 11 bands optimized for accuracy and computational efficiency. The opacities are calculated assuming local thermodynamic and chemical equilibrium (44). This code has been used extensively to model the atmospheric circulation of hot Jupiters, hot Neptunes and super Earths (e.g., 26-28, 49). We use $1\times$ solar abundance values and utilize the system parameters listed in Table 1. We compare the T-P profiles to condensation curves (49) for various species between 500 and 1500 K to determine the most likely condensation cloud species responsible for the observed transmission spectrum (Figure 2).

To calculate the vertical extent of potential Na_2S clouds on HAT-P-26b, we utilize the cloud code of (50), updated for Na_2S cloud formation (29). We denote the cloud base as the region where the limb-averaged TP profile intersects the condensation curve of Na_2S (Figure 2) and the cloud top as the region of the limb where the slant optical depth equals one (51). Previous studies have shown that Na_2S is highly scattering (29, 46) with a high enough abundance to form optically thick clouds in exoplanet atmospheres (29). The pressure extent of Na_2S condensate derived from the cloud model is in agreement with the pressure range probed by transmission spectral observations. Additionally, as the most likely cloud to condense at this pressure is Na_2S , it is likely that all of the Na is locked up the condensed phase and would not present as atomic lines produced by sodium in the gas phase.

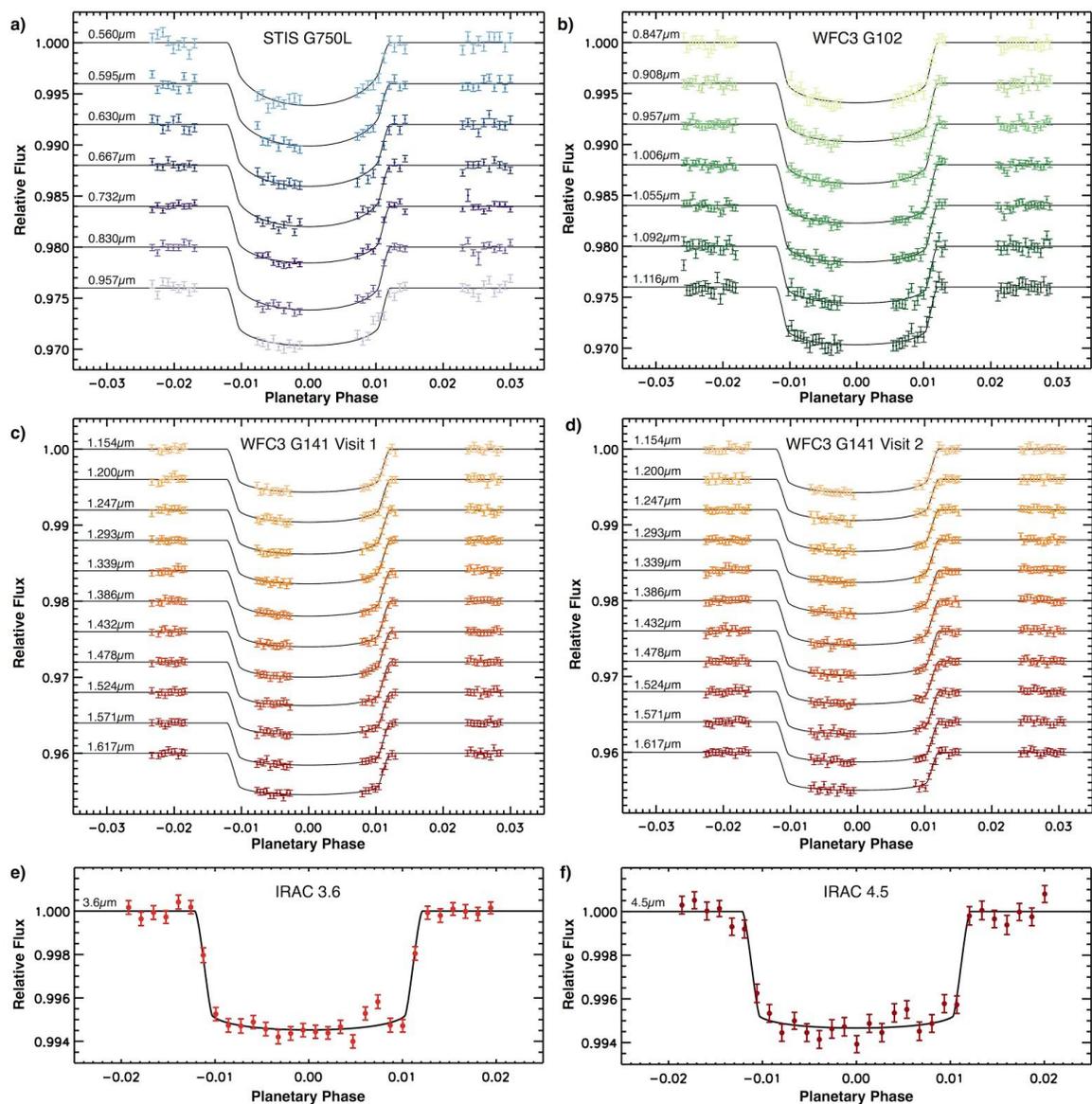


Fig. S1. Spectroscopic light curves of all six transits of HAT-P-26b. a) HST STIS, b) HST WFC3 G102 c) HST WFC3 G141 visit 1, d) G141 visit 2, e) Spitzer IRAC 3.6, and f) IRAC 4.5. Normalized and systematics-corrected data (colored points) are shown using the highest weighted systematic model for 7 spectroscopic channels spread in STIS from 0.5-1.0 μm , 7 channels for WFC3 G102 from 0.8-1.15 μm , 11 channels for WFC3 visit 1 & 2 from 1.1-1.7 μm , and Spitzer 3.6 & 4.5 μm photometric channels. Spitzer photometric light curves have been binned in time by ~ 2 minutes. Each light curve is offset and labeled for clarity.

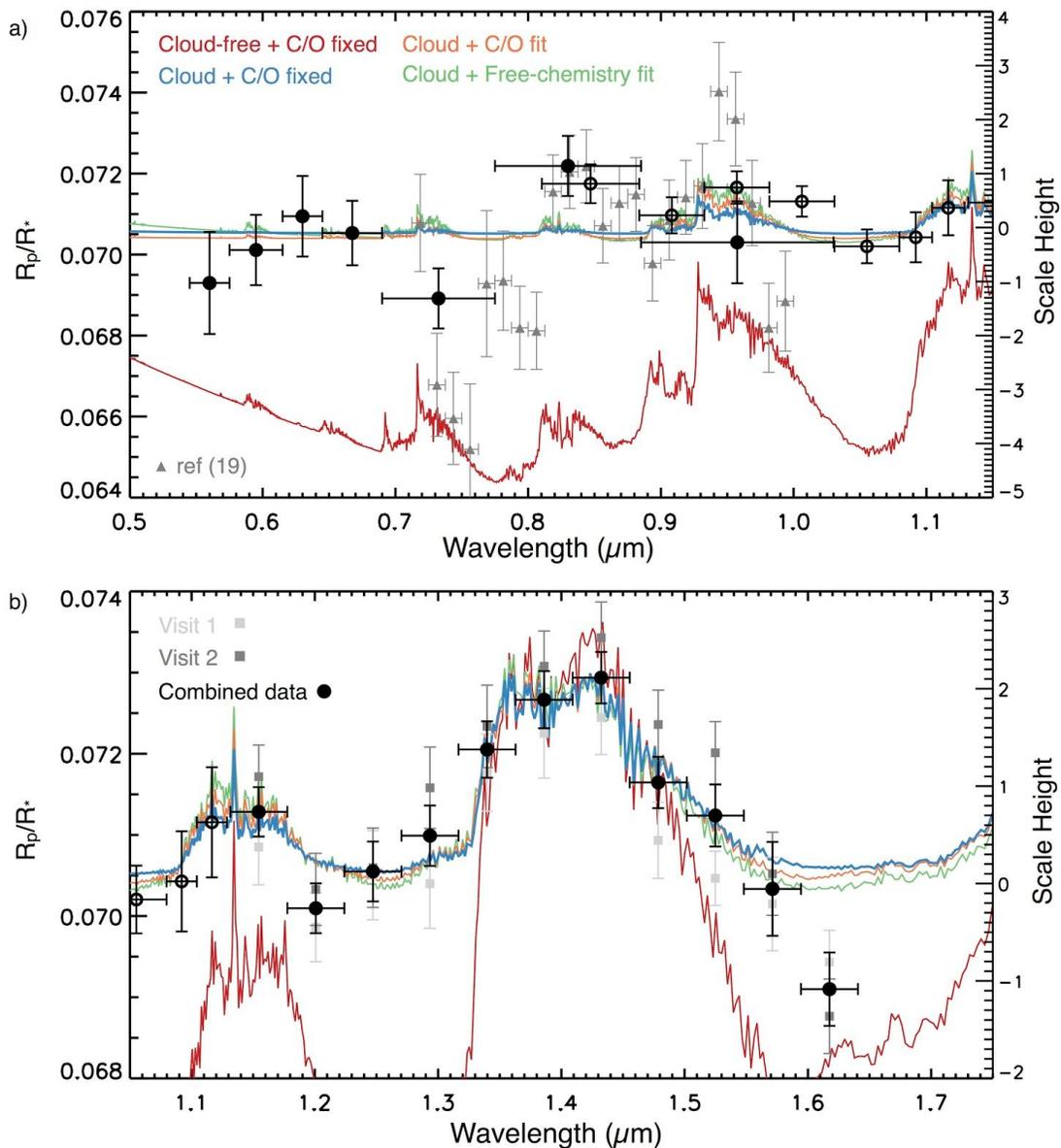


Fig. S2. Expanded view of HAT-P-26b transmission spectrum from each HST visit.

Transmission spectrum measured for HAT-P-26b expanded to show the optical and near-infrared regions of the spectrum separately. Each plot shows models for four of the 12 ATMO retrievals (M1-3, M7; see Table S2 for the statistics). a) Optical STIS and WFC3 G102 measurements plotted with previous ground-based data (19). b) Prominent water absorption measured with WFC3 G141 over two visits shown as squares, with the final combined spectrum as solid black points. The right-hand axis shows the corresponding scale of the atmospheric transmission in terms of planetary scale height (as in Figure 1).

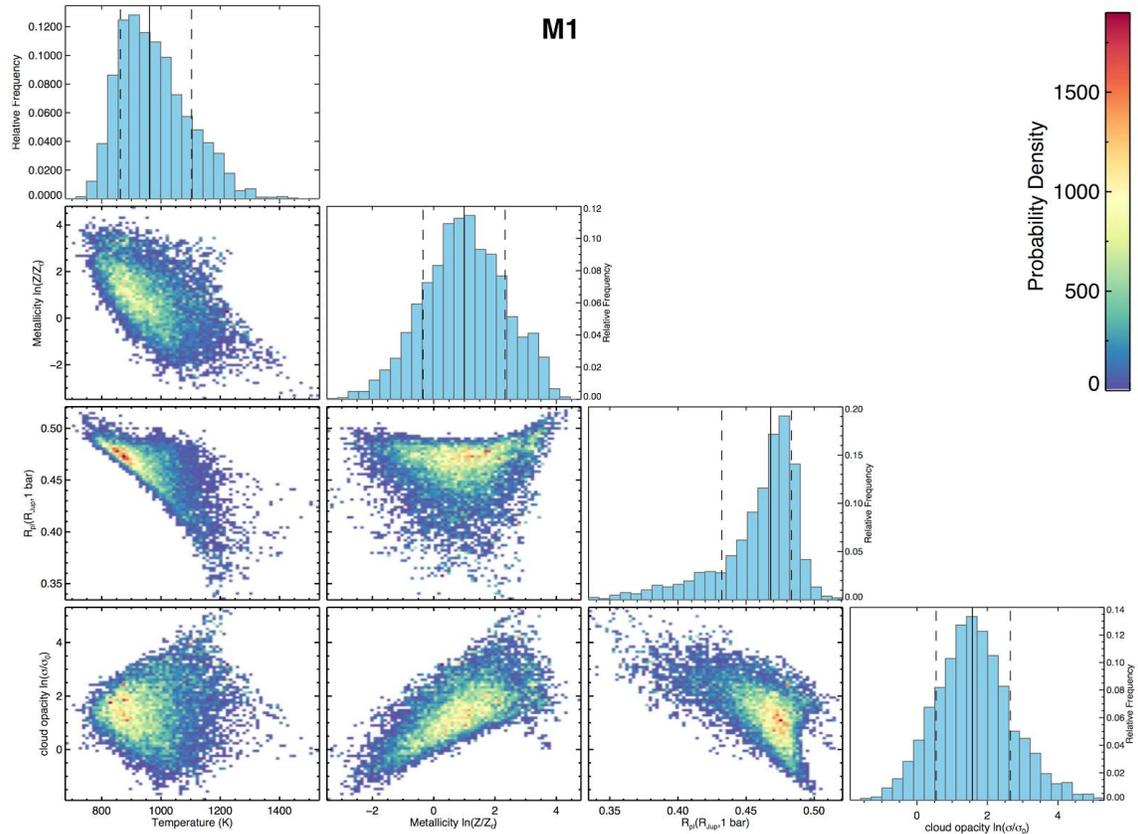


Fig. S3. Correlation plot for pairs of retrieved parameters from the ATMO retrieval. Correlation plot for model M1, which fits for the parameters, temperature, metallicity, radius ratio, and cloud opacity at a fixed solar C/O ratio and equilibrium chemistry. The posterior distributions show the probability density distribution of each fit. The vertical lines in the histograms indicate the fit value and uncertainty.

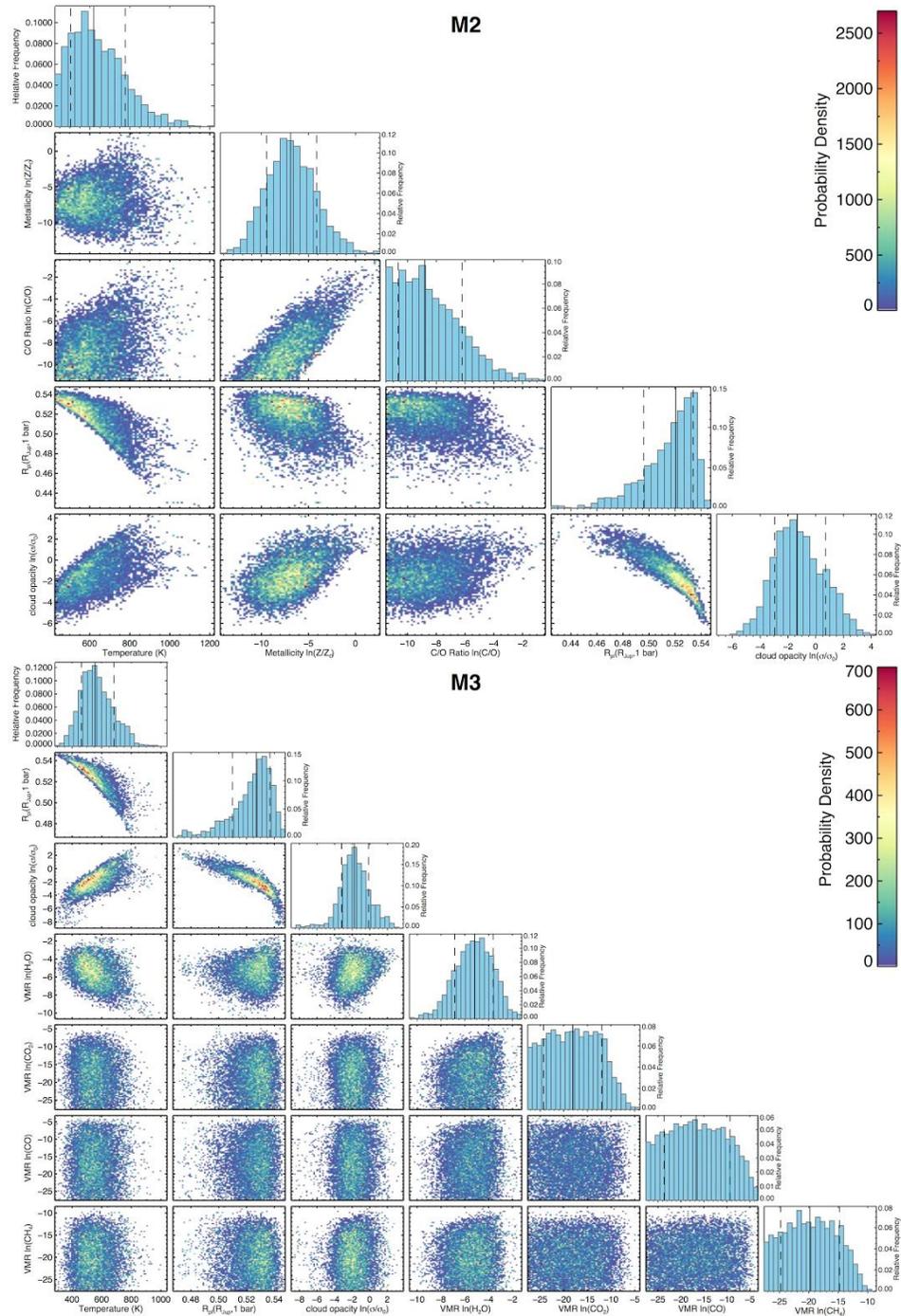


Fig. S4. Correlation plot for pairs of retrieval parameters for models M2 and M3. As in Fig S3 but for M2 and M3. M2 is the same as M1 but also fits C/O and M3 is a free-chemistry, isothermal model using the ATMO retrieval.

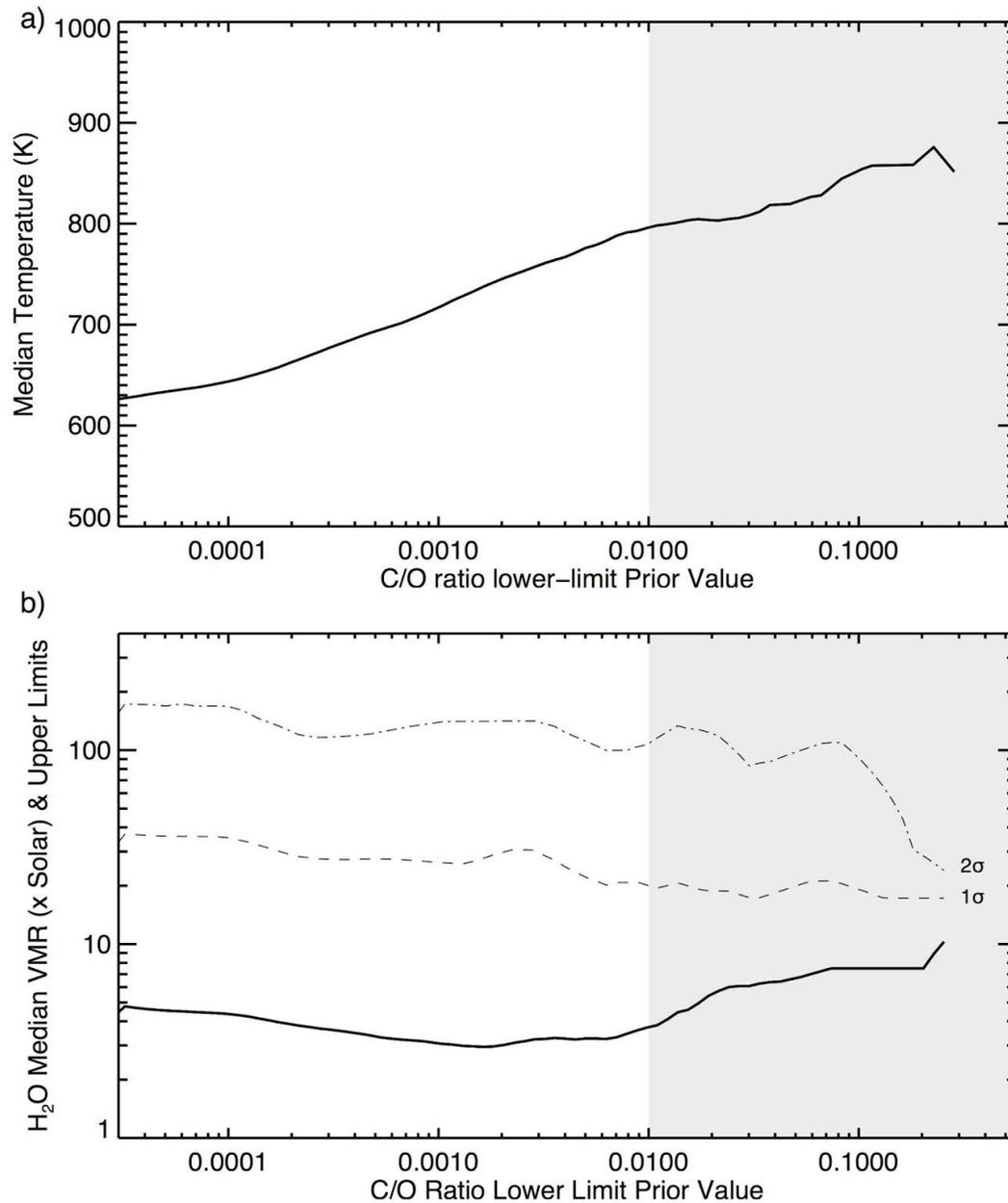


Fig. S5. C/O prior impacts on the temperature and volume mixing ratio. From model M2, a) the impact of a C/O prior on derived atmospheric temperature. b) impact of the C/O prior on the derived water abundance with the 1- and 2- σ upper limits. The grey shaded region in both plots shows the parameter space explored with the adopted prior constraint of C/O > 0.01.

Table S1. HAT-P-26 system parameters fixed and derived. The radius (R_*) and mass (M_*) of the star are in terms of solar radius and mass. T_{eff} is the effective temperature of the star. We list the planetary radius (R_p) and mass (M_p) are in terms of Earth radius and mass planetary with gravity (g_p), and the derived inclination and semi-major axis in terms of stellar radius (a/R_*).

Parameter	Value
Stellar Parameters	
R_* (R_{sun})	$0.788(-0.043 +0.098)^\dagger$
M_* (M_{sun})	$0.816 (\pm 0.033)^\dagger$
T_{eff} (K)	$5011(\pm 55)^\dagger$
[Fe/H]	$0.01(\pm 0.04)^\dagger$
Planet Parameters	
R_p (R_{Earth})	6.33^\dagger
M_p (M_{Earth})	18.6^\dagger
Period (days)	4.2345^\dagger
g_p (ms^{-2})	4.47^\dagger
eccentricity	0.124^\dagger
inclination ($^\circ$)	$88.09 (\pm 0.553)$
a/R_*	$11.89 (\pm 0.417)$
Derived Planetary Metallicity	
$\ln[M/H]$	$1.566 (\pm 1.7034)$
Measured center of transit time (BJD_{TBD})	
STIS G750L	$2457413.432836 (\pm 0.000172)$
WFC3 G102	$2457616.690103 (\pm 0.000011)$
WFC3 G141 Visit 1	$2457460.013266 (\pm 0.000016)$
WFC3 G141 Visit 2	$2457510.827100 (\pm 0.000016)$
IRAC 3.6	$2456545.361830 (\pm 0.000030)$
IRAC 4.5	$2456405.623110 (\pm 0.000070)$

† . system parameters adopted from (1)

‡ . Updated stellar parameters from (52)

Table S2. Retrieval parameters and fits for all 12 model retrievals. For each model we list the number of free parameters, BIC, and weight, which are used to calculate the marginalised metallicity of the atmosphere. Metallicity is listed as natural logs, where solar metallicity equals $\ln(1)$.

Model	ln(Metallicity)	1σ Uncertainty	# of free parameters	BIC	W_q
M1	1.00	1.3391	4	58.32	0.425
M2	1.72	1.7231	5	58.21	0.449
M3	2.81	1.5995	7	62.13	0.063
M4	2.14	1.0764	6	65.03	0.015
M5	2.86	1.5952	7	65.43	0.012
M6	2.84	1.5463	9	68.85	0.002
M7	6.05	0.3080	3	77.09	0.000
M8	-0.30	1.2775	4	67.28	0.005
M9	4.65	1.2242	6	64.05	0.024
M10	6.08	0.3447	5	80.48	0.000
M11	-0.27	0.9195	6	69.19	0.002
M12	3.80	1.9600	8	69.30	0.002
BMA	1.566 (4.8\times solar)	1.7034			

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