



# An Upper Limit on the Linear Polarization Fraction of the GW170817 Radio Continuum

Alessandra Corsi<sup>1</sup>, Gregg W. Hallinan<sup>2</sup>, Davide Lazzati<sup>3</sup>, Kunal P. Mooley<sup>2,4,9</sup>, Eric J. Murphy<sup>5</sup>, Dale A. Frail<sup>4</sup>,  
Dario Carbone<sup>1</sup>, David L. Kaplan<sup>6</sup>, Tara Murphy<sup>7</sup>, Shrinivas R. Kulkarni<sup>2</sup>, and Kenta Hotokezaka<sup>8</sup>

<sup>1</sup>Department of Physics and Astronomy, Texas Tech University, Box 1051, Lubbock, TX 79409-1051, USA; [alessandra.corsi@ttu.edu](mailto:alessandra.corsi@ttu.edu)

<sup>2</sup>Caltech, 1200 East California Boulevard, MC 249-17, Pasadena, CA 91125, USA

<sup>3</sup>Department of Physics, Oregon State University, 301 Weniger Hall, Corvallis, OR 97331, USA

<sup>4</sup>National Radio Astronomy Observatory (NRAO), P.O. Box O, Socorro, NM 87801, USA

<sup>5</sup>National Radio Astronomy Observatory, 520 Edgemont Road, Charlottesville, VA 22903, USA

<sup>6</sup>Department of Physics, University of Wisconsin—Milwaukee, Milwaukee, WI 53201, USA

<sup>7</sup>Sydney Institute for Astronomy, School of Physics, University of Sydney, Sydney, New South Wales 2006, Australia

<sup>8</sup>Department Astrophysical Sciences, Princeton University, Peyton Hall, Princeton, NJ 08544, USA

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## Abstract

We present late-time radio observations of GW170817, the first binary neutron-star (NS) merger discovered through gravitational waves (GWs) by the advanced Laser Interferometer Gravitational-wave Observatory (LIGO) and Virgo detectors. Our observations, carried out with the Karl G. Jansky Very Large Array (VLA), were optimized to detect polarized radio emission, and thus to constrain the linear polarization fraction of GW170817. At an epoch of  $\approx 244$  days after the merger, we rule out linearly polarized emission above a fraction of  $\approx 12\%$  at a frequency of 2.8 GHz (99% confidence). Within the structured jet scenario (a.k.a. successful jet plus cocoon system) for GW170817, the derived upper limit on the radio continuum linear polarization fraction strongly constrains the magnetic field configuration in the shocked ejecta. We show that our results for GW170817 are compatible with the low level of linear polarization found in afterglows of cosmological long  $\gamma$ -ray bursts (GRBs). Finally, we discuss our findings in the context of future expectations for the study of radio counterparts of binary NS mergers identified by ground-based GW detectors.

*Key words:* gravitational waves – polarization – stars: neutron

## 1. Introduction

On 2017 August 17, the field of gravitational-wave (GW) astronomy reached the big leagues with a dazzling discovery. Only  $\approx 8$  days before the official end of their second observing run (O2), advanced Laser Interferometer Gravitational-wave Observatory (LIGO) and Virgo obtained their first direct detection of GWs from a binary neutron-star (NS) merger, an event dubbed GW170817 (Abbott et al. 2017a).

After the GW discovery, GW170817 gifted the astronomical community with an electromagnetic (EM) counterpart spanning all bands of the spectrum (e.g., Abbott et al. 2017b, and references therein). Only  $\approx 2$  s after the GW170817 merger, a short  $\gamma$ -ray burst (GRB) was detected by the *Fermi* and *Integral* satellites (Abbott et al. 2017c; Goldstein et al. 2017; Savchenko et al. 2017). The discovery of  $\gamma$ -rays from GW170817 was followed by the detection of a UV/optical/IR counterpart in NGC 4993, a lenticular galaxy located at  $\approx 40$  Mpc (Coulter et al. 2017), resulting in the nearest short GRB with a measured redshift (e.g., Fong et al. 2017). While it had long been thought that short GRBs are NS–NS mergers launching relativistic jets pointed directly at us, the close distance of GW170817 implied that its  $\gamma$ -ray counterpart was  $\sim 10^3$ – $10^4$   $\times$  less energetic than the previously known population of short GRBs (Fong et al. 2017; Goldstein et al. 2017). This formed the first piece of an intriguing puzzle.

The early UV/optical/IR emission from GW170817 was rather different from the non-thermal optical afterglows of short GRBs, and it was soon recognized to be dominated by a “kilonova,” a quasi-thermal transient powered by the

radioactive decay of *r*-process nuclei (e.g., Arcavi et al. 2017; Chornock et al. 2017; Cowperthwaite et al. 2017; Drout et al. 2017; Evans et al. 2017; Kasen et al. 2017; Kasliwal et al. 2017; Kilpatrick et al. 2017; Pian et al. 2017; Shappee et al. 2017; Smartt et al. 2017; Tanvir et al. 2017; Valenti et al. 2017). This kilonova detection solved a decades-old mystery of where most of the elements heavier than iron are synthesized (for a recent review, see e.g., Metzger 2017, and references therein).

About 10 days after the GW discovery of GW170817, an X-ray counterpart was detected by the *Chandra* satellite (Haggard et al. 2017; Margutti et al. 2017; Troja et al. 2017). A delayed radio afterglow was unveiled by the Karl G. Jansky Very Large Array (VLA) about two weeks after the merger (Hallinan, Corsi et al. 2017), and subsequently confirmed by the Australia Telescope Compact Array (ATCA). Radio and X-ray observations of GW170817 probe a completely different emission mechanism than the kilonova observed at optical-IR wavelengths, namely, non-thermal radiation from the fastest ejecta. Continued VLA and ATCA monitoring over the first  $\sim 100$  days since the merger revealed a steady increase of the optically thin synchrotron radio emission, followed by a turnover ( $t \gtrsim 150$  days since merger; Alexander et al. 2017, 2018; Dobie et al. 2018; Margutti et al. 2018; Mooley et al. 2018). Together with the weak  $\gamma$ -rays, these observations clearly set GW170817 apart from the previously known population of short GRBs with fast-fading afterglows (Fong et al. 2017). The relatively slow temporal rise of the radio flux (Hallinan et al. 2017; Dobie et al. 2018; Mooley et al. 2018), in particular, has ruled out the simplest scenario relating GW170817 radio counterpart to synchrotron emission from a uniform (“top-hat”) jet shocking the ISM (Margutti et al. 2017;

<sup>9</sup> Jansky Fellow (NRAO, Caltech).

Troja et al. 2017), which is usually invoked to explain cosmological GRB afterglows (e.g., Sari 1999).

Broadly speaking, two main scenarios have been proposed to explain GW170817 non-thermal emission: (i) A successful structured jet (a.k.a. successful jet-cocoon system) composed of an outflow with a narrow, highly relativistic core (similar to cosmological GRB jets) initially directed away from our line of sight (off-axis), plus slower-moving wings (e.g., Kasliwal et al. 2017; Lazzati et al. 2017a, 2017b, 2017c; Margutti et al. 2018); (ii) A choked-jet scenario where the jet is unable to break out of the neutron-rich dynamical ejecta, and the bulk of the energy is imparted to a radially stratified and quasi-spherical (QS; or wide-angle) mildly relativistic cocoon (Kasliwal et al. 2017; Nakar & Piran 2017, 2018; Gottlieb et al. 2018; Mooley et al. 2018). We note that a dynamical ejecta model where radio emission arises from the fast tail of the dynamical merger ejecta has also been proposed (e.g., Hotokezaka et al. 2018), but is somewhat unlikely given the sharp peak and fast decline of the radio light curve (e.g., Alexander et al. 2018; Dobie et al. 2018). Thus, in what follows, we will not discuss this third scenario further.

While scenarios (i) and (ii) above imply very different geometries for the outflow, there are sufficient free model parameters making it impossible to distinguish them based on the *total radio intensity* alone (e.g., Margutti et al. 2018). Fortunately, detecting a polarized radio signal from GW170817 can provide a useful discriminant. Indeed, the presence of a large degree ( $\approx 20\%$ ) of linear polarization in the radio continuum would be a “smoking gun” for a high degree of asymmetry, and hence favor the jet scenario (i) (e.g., Rossi et al. 2004; Lazzati et al. 2017c; Gill & Granot 2018; Nakar et al. 2018). In the absence of substantial linearly polarized emission, even though scenario (i) cannot be ruled out, strong constraints can be set on the structure of the post-shock magnetic field (e.g., Ghisellini & Lazzati 1999; Sari 1999; Rossi et al. 2004; Gill & Granot 2018; Nakar et al. 2018).

Motivated by the above considerations, here we present polarization observations of GW170817 radio counterpart. Our upper limit rules out strong linearly polarized GHz emission and, within the structured jet scenario, sets stringent constraints on the structure of the magnetic field within the shocked ejecta. This Letter is organized as follows. In Section 2 we describe our observations and data reduction; in Section 3 we discuss our results; in Section 4 we summarize and conclude.

## 2. Radio Observations and Data Reduction

We observed the field of GW170817 with the Karl G. Jansky VLA in its most extended (A) configuration on multiple epochs between 2018 March 02.321 UTC and 2018 May 12.168 UTC (via projects VLA/17B-397; PI: K. P. Mooley; and VLA/18A-457; PI: A. Corsi). These observations were carried out in the S-band, at a nominal central frequency of  $\approx 3$  GHz, and with a nominal 2 GHz bandwidth. We included bandpass, flux density, and polarization position angle calibration scans on 3C286. The unpolarized source J1407+2827 was observed to calibrate for polarization leakage. During all epochs, J1248–1959 was used as our phase calibrator.

The VLA data were first calibrated using the VLA automated calibration pipeline available in CASA, which is designed for Stokes  $I$  continuum calibration. Because our phase calibrator J1248–1959 is marginally resolved at the longest baselines of the VLA in its A configuration, we restricted the UV range for

this calibrator to  $\lesssim 200$  k $\lambda$  in all gain calibrator calls within the automated calibration pipeline. After the automated calibration, we set the polarization model for our polarized calibrator 3C286 (11.2% fractional polarization, and polarization angle of  $33^\circ$ ; Perley & Butler 2013). Then, polarization calibration steps were carried out using the automated pipeline calibration tables for pre-calibration. Specifically, we first solved for the cross-hand (RL, LR) delays due to the residual delay difference on the reference antenna used for the original delay calibration. Then, we solved for the instrumental polarization (the frequency-dependent leakage terms, also referred to as “D-terms”) using the unpolarized source J1407+2827. For all our observations, we found leakages of  $\lesssim 15\%$  for most antennas and spectral windows. Having calibrated the instrumental polarization, we carried out a frequency-dependent position angle calibration using the source 3C286. We found that the residual RL phase on the reference antenna (after taking out the cross-hand delays) spanned about  $10^\circ$ – $15^\circ$  across most spectral windows.

After calibrating and visually inspecting the data for any further flagging, we run the CLEAN algorithm (Högbom 1974) in interactive mode to image the fully calibrated Stokes  $IQUV$ , and derive single-epoch sensitivities. In the cleaning process, we used a natural weighting of the visibilities so as to maximize the map point source sensitivity. We applied the same source mask to all polarizations, and used a pixel of size  $0''.2$  (so as to oversample the synthesized beam, which was of  $\approx 0''.85 \times 1''.5$  in our observations). A summary of our results is reported in Table 1. For each observation, we give the central UTC, the epoch in days since GW170817 merger time (2017 August 17.528 UTC), the total duration of the observation (including calibration), the central observing frequency, and the rms sensitivity reached in Stokes  $Q$ ,  $U$  (we do not discuss Stokes  $V$  here, i.e., circular polarization, as no emission is expected or seen in this polarization state). In Table 1 we also report the sensitivity reached by co-adding our last four observations (which have comparable rms), and imaging the resulting data set following the same procedure described above for the single epochs. The rms sensitivity reached in Stokes  $Q$  and  $U$  after co-adding is  $\approx 1.7 \mu\text{Jy beam}^{-1}$ .

Using a circular region centered on the position of GW170817 ( $\alpha = 13^{\text{h}}09^{\text{m}}48^{\text{s}}.071$  and  $\delta = -23^\circ 22' 53''.37$ , J2000; e.g., Hallinan et al. 2017; Kasliwal et al. 2017) and with an area comparable to that of the FWHM synthesized beam, we calculate the peak brightness measured in Stokes  $Q$  and Stokes  $U$  at the various epochs, and in the co-added data set. In all cases we find that the measured Stokes  $Q$  and  $U$  peak brightness at the GW170817 location is below  $< 3 \times \sigma_{Q,U}$  where  $\sigma_{Q,U}$  is the map rms. Thus, all our polarization observations yielded non-detections in Stokes  $Q$  and  $U$ .

## 3. Results and Discussion

In our March 2 UTC observation, which had the shortest duration (Table 1), we measure  $p = \sqrt{Q^2 + U^2} / \sigma_{U,V} \approx 3.0$  (where  $\sigma_{U,V} = 4.5 \mu\text{Jy beam}^{-1}$ ; see Table 1) at 2.8 GHz. Accounting for Ricean bias, we thus set a 99% confidence upper limit of  $p < 5.2$  (Vaillancourt 2006). The Stokes  $I$  peak brightness measured at this epoch is  $(75.9 \pm 6.4) \mu\text{Jy}$  (this includes a 5% absolute flux density calibration error), and thus the corresponding upper limit on the linear polarization fraction is  $\Pi = \sqrt{Q^2 + U^2} / I \lesssim 31\%$  at  $\approx 197$  days since merger.

**Table 1**  
Sensitivity Reached in our VLA Polarization Observations of GW170817

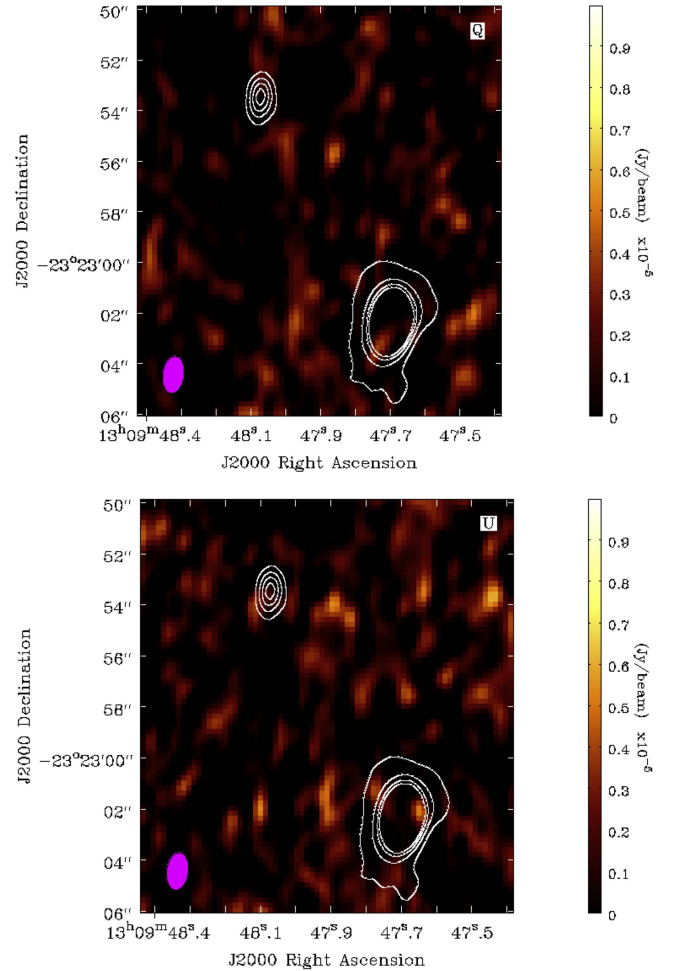
UTC	Epoch (days since 2017 Aug 17.528 UTC)	$\Delta T_{\text{obs}}$ (hr)	$\nu$ (GHz)	Stokes $Q$ rms ( $\mu\text{Jy beam}^{-1}$ )	Stokes $U$ rms ( $\mu\text{Jy beam}^{-1}$ )
2018 March 02.321	197	1.5	3.0	4.5	4.5
2018 March 25.344	220	3.5	2.8	3.3	3.3
2018 March 26.310	221	3.5	2.8	3.4	3.4
2018 May 11.167	267	3.5	2.8	3.5	3.4
2018 May 12.168	268	3.5	2.8	3.4	3.4
2018 March 25–May 12	$244 \pm 24$	$3.5 \times 4$	2.8	1.7	1.7

From the co-added map derived using our last four observations with comparable rms sensitivity (Table 1 and Figure 1), we get  $p = \sqrt{Q^2 + U^2}/\sigma_{U,V} \approx 1.7$  (where  $\sigma_{U,V} = 1.7 \mu\text{Jy beam}^{-1}$ ; see Table 1) at 2.8 GHz and at a mean epoch of  $\approx 244$  days since merger, which implies a 99% upper limit on  $p$  of  $p < 3.8$  (Vaillancourt 2006). The Stokes  $I$  peak brightness measured for GW170817 in the co-added image is of  $(51.9 \pm 3.3) \mu\text{Jy beam}^{-1}$  (fully consistent with the turnover trend identified by Dobie et al. 2018). Thus, our most stringent upper limit on the linear polarization fraction of GW170817 is of  $\Pi = \sqrt{Q^2 + U^2}/I \lesssim 12\%$  at  $\approx 244$  days since merger (Figure 2).

As discussed in Section 1, a successful structured jet (scenario (i)) and a choked-jet-cocoon system (scenario (ii)) can both explain the observed radio light curve of GW170817 (e.g., Hallinan et al. 2017; Kasliwal et al. 2017; Lazzati et al. 2017c; Gill & Granot 2018; Margutti et al. 2018; Mooley et al. 2018; Nakar & Piran 2018). Thus, polarization observations have been proposed as a way to break this degeneracy and discriminate between scenarios (i) and (ii) (Lazzati et al. 2017c; Gill & Granot 2018; Nakar et al. 2018). The predictions for the linear polarization near the peak of the radio light curve are indeed substantially different in these two cases. For a given magnetic field configuration, the successful jet scenario produces a larger polarization than that expected for a quasi-spherical (QS) outflow. However, for both outflow structures, the predicted polarization fraction also depends strongly on the configuration of the magnetic field (which is usually assumed to be completely tangled in the plane of the shock). Specifically, the degree of linear polarization is maximum for a magnetic field fully contained within the plane of the shock, and decreases with an increasing magnetic field component in the direction of the shock normal. This effect can be parametrized by the ratio  $b = 2\langle B_{\text{sn}}^2 \rangle / \langle B_{\text{sp}}^2 \rangle$ , where  $b = 0$  is for a magnetic field fully contained in the plane of the shock, while for  $b > 0$  the component of the field along the shock normal also contributes to the emission (Gill & Granot 2018).

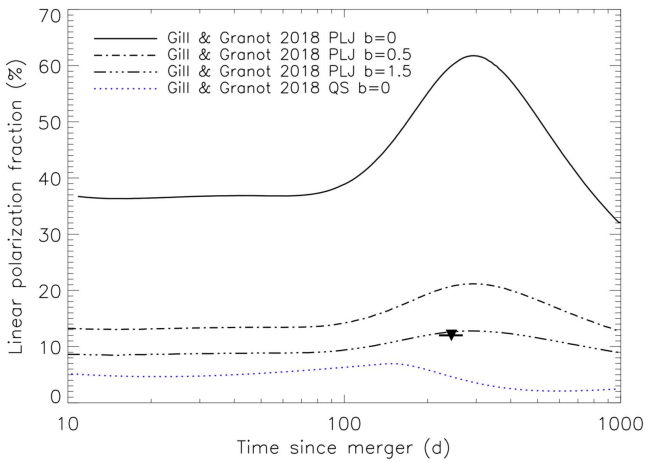
In Figure 2 we show the predictions by Gill & Granot (2018) for the linear polarization fraction of the radio continuum from a successful structured jet with a power-law distribution of energy and Lorentz factors (PLJ; black lines). Similar predictions for the case of a radially stratified QS ejecta are also shown (blue dotted line). These predictions strongly depend on the value of  $b$ , but for the QS ejecta case we only show  $b = 0$  as larger values of  $b$  would imply even smaller degrees of linear polarization.

As evident from Figure 2, our VLA linear polarization fraction upper limit (downward-pointing triangle) excludes a structured jet model with  $b = 0$ , and requires  $b > 0.5$ . We note that analogous magnetic field configurations are inferred



**Figure 1.** Top: Stokes  $Q$  intensity map of the co-added observations of the GW170817 field carried out in the S-band between March 25 and May 12 (see Table 1). Stokes  $I$  contours of GW170817 radio counterpart are also shown (white; 20%, 40%, 60%, and 80% relative emission contours). GW170817 radio counterpart is located at  $\alpha = 13^{\text{h}}09^{\text{m}}48^{\text{s}}.071$ ,  $\delta = -23^{\circ}22'53''.37$  (J2000; e.g., Hallinan et al. 2017; Kasliwal et al. 2017). The Stokes  $I$  intensity contours of the host galaxy of GW170817 are also overlaid (bottom-right portion of the panel). The FWHM synthesized beam ellipse is shown in magenta. Bottom: same as the top panel, but for the Stokes  $U$  intensity map.

comparing optical data of long-duration GRBs with various predictions for the expected optical polarization fraction. Analytic models of long GRBs (Ghisellini & Lazzati 1999; Sari 1999) predict polarization values peaking at  $\sim 10\%$ – $15\%$  (assuming  $b = 0$ ) for the most probable observing configuration (with the observer at the edge of the jet). Optical observations, on the other hand, have provided polarization values of long GRBs optical afterglows clustered around a few percent (Covino & Gotz 2016), even for bursts with multiple



**Figure 2.** VLA upper limit on the linear polarization fraction  $\sqrt{Q^2 + U^2}/I$  of the GHz radio flux of GW170817 (downward-pointing triangle) compared with different theoretical predictions for the power-law structured jet model (PLJ; black), and for a quasi-spherical ejecta (QS; blue). These predictions are by Gill & Granot (2018). For the models here plotted,  $b = 0$  represents the case of a magnetic field completely contained in the plane of the shock, while  $b > 0$  is for a magnetic field with a component in the direction of the shock normal also contributes. See the text for discussion.

observations spanning the entire afterglow evolution (e.g., Greiner et al. 2003; Wiersema et al. 2012). Finally, we note that our upper limit on the radio linear polarization fraction of GW170817 cannot constrain the QS ejecta scenario, even in the most optimistic case of a magnetic field fully confined in the plane of the shock ( $b = 0$ ).

#### 4. Summary and Conclusion

We have presented the first observational constraint on the  $\approx 3$  GHz linear polarization fraction of GW170817 radio continuum. Thanks to the excellent point source sensitivity of the VLA, we are able to unambiguously rule out the most optimistic predictions for the linearly polarized radio flux expected within a successful structured jet scenario. We have also shown that, under the hypothesis that a successful structured jet did indeed form in GW170817, the magnetic field behind the shock cannot be fully contained within the plane of the shock. Instead, a significant component of the magnetic field in the direction perpendicular to the shock plane is required to reconcile theoretical predictions with our observational upper limit.

Even though GW170817 is a relatively nearby event ( $d_L \approx 40$  Mpc), its faint radio emission ( $\sim 10^{26}$  erg s $^{-1}$  Hz $^{-1}$  peak spectral luminosity density at  $\approx 3$  GHz) does not allow us to constrain the degree of linear polarization of the radio continuum down to a level that could probe the QS ejecta formed in a choked-jet scenario. More stringent constraints on this scenario may be achieved via direct Very Long Baseline Interferometry (VLBI) imaging of the merger ejecta.

With the advanced LIGO and Virgo detectors now scheduled to start their third observing run (O3; Abbott et al. 2016), we expect to soon have more opportunities for probing the possible variety of radio afterglows from NS–NS (or black hole–NS) mergers. The luminosity of the radio counterparts of these events will depend on several factors, including the interstellar medium (ISM) density, the total energy in the fastest ejecta, and the viewing geometry. If a compact binary merger with a radio afterglow twice as bright as GW170817 were to be

discovered, the VLA could probe linear polarization fractions below  $\approx 10\%$ . At this level, the absence of linear polarization is likely to challenge strongly the successful structured jet hypothesis (assuming an outflow geometry and viewing angle similar to that of GW170817).

Finally, looking further into the future, when advanced LIGO and Virgo will be reaching their nominal sensitivities and discovering NS–NS mergers up to  $\approx 120$ – $190$  Mpc (Abbott et al. 2016), detecting radio counterparts as faint as GW170817 will require radio arrays  $\approx 10\times$  more sensitive than the VLA. In this respect, the next-generation Very Large Array (ngVLA; Murphy 2017) may offer us a unique opportunity to probe the different ejecta structures via radio polarimetry.

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#### ORCID iDs

Alessandra Corsi <https://orcid.org/0000-0001-8104-3536>  
 Gregg W. Hallinan <https://orcid.org/0000-0002-7083-4049>  
 Davide Lazzati <https://orcid.org/0000-0002-9190-662X>  
 Kunal P. Mooley <https://orcid.org/0000-0002-2557-5180>  
 Eric J. Murphy <https://orcid.org/0000-0001-7089-7325>  
 Dario Carbone <https://orcid.org/0000-0002-6575-4642>  
 David L. Kaplan <https://orcid.org/0000-0001-6295-2881>  
 Tara Murphy <https://orcid.org/0000-0002-2686-438X>  
 Shrinivas R. Kulkarni <https://orcid.org/0000-0001-5390-8563>  
 Kenta Hotokezaka <https://orcid.org/0000-0002-2502-3730>

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