A 492 GHz quasioptical SIS receiver for submillimeter astronomy

G. Engargiola  
Department of Astronomy, University of Illinois, Urbana, Illinois 61801  
J. Zmuidzinas  
Downs Laboratory of Physics, Caltech, Pasadena, California 91125  
K. Y. Lo  
Department of Astronomy, University of Illinois, Urbana, Illinois 61801

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A 492 GHz quasioptical superconductor-insulator-superconductor (SIS) heterodyne receiver for submillimeter astronomy has been constructed with a double-sideband noise temperature of 170±25 K, where the system response is an average over a 600 MHz i.f. bandwidth centered at 1.5 GHz. The excellent sensitivity is achieved through use of a planar twin-slot antenna feeding a pair of Nb–Al2O3–Nb junctions linked by an inductance. A main beam of f/15 has been measured. The performance of this receiver is comparable to that of the best 492 GHz waveguide receivers currently in operation.

I. INTRODUCTION

Because of their high sensitivity, superconductor-insulator-superconductor (SIS) tunnel junction receivers have found widespread use in millimeter-wave astronomy for nearly a decade. Only within the past few years have durable, low-noise SIS mixers become available in the 450–690 GHz frequency range, largely due to improved mixer designs and the use of Nb–Al2O3–Nb trilayers in the manufacture of junctions. Further efforts to improve the sensitivity of submillimeter SIS mixers continue, with the aim of eventually producing quantum-limited devices, where the mixer noise power is dominated by fluctuations in the vacuum state of the incident radiation field.

Low-noise SIS mixers generally fall into two design types, differing in how incident radiation is coupled onto the mixing element. In the waveguide mixer design, SIS mixers are suspended across waveguides into which radiation is coupled with a feedhorn; the impedance match to the mixing element is optimized with an E-plane tuner and a backshort. In the quasioptical mixer design, the mixing elements are connected by superconducting microstrip to a broad-band radiation shield at 51 K and an inner one at 15 K enclose a space modes is still at least 30%–40% less efficient than waveguide feedhorns. Another potential problem is the lack of a tunable element comparable to a waveguide backshort. However, integrated fixed-tuned rf matching circuits have been developed which significantly increase the quantum efficiency and bandwidth of quasioptical mixers.

In this paper, we describe the design and performance of a 492 GHz quasioptical SIS heterodyne receiver with sensitivity comparable to the best waveguide receivers currently operating at this frequency.

II. RECEIVER CONFIGURATION

A schematic drawing of the 492 GHz receiver is shown in Fig. 1. A 25 µm thick mylar beamsplitter mounted to the dewar injects local oscillator radiation, which is vertically polarized. A Teflon lens matches the LO to the f/15 beam of the receiver. Approximately 10% of the LO power (1–15 µW) is reflected into the dewar, with the remainder transmitted onto an Eccosorb termination.

The dewar is a hybrid design by Ellison. An outer radiation shield at 51 K and an inner one at 15 K enclose a work surface, which is mounted to a liquid helium reservoir. The shields are cooled by a two-stage CTI 350 helium refrigerator. Radiation enters the dewar through a vacuum window made from 25 µm thick Mylar sheet. Room-temperature infrared emission is filtered from the beam with a 1.27 mm thick z-cut quartz disk which covers the aperture in the 51 K shield; the quartz is antireflection coated with low-density polyethylene. The aperture in the 15 K shield is open. The mixer block and a charcoal adsorption pump are clamped onto the work surface at 4.2 K, and an NRAO 1.2–1.8 GHz HEMT amplifier, which provides 31 dB of gain, is mounted to the inside of the 15 K shield.

The i.f. chain is a standard design. Stainless steel coaxial cable carries the i.f. signal from the mixer block to the cooled L band amplifier via a Pamtech LTE 1310 K cryogenic isolator. The isolator, which terminates any power reflected back to the mixer from the amplifier, somewhat...
FIG. 1. A schematic drawing of the 492 GHz receiver. A Mylar beam splitter combines the LO with the signal. Microwave radiation injected into the dewar passes a Mylar pressure window and an IR blocking filter. A plastic lens and a quartz hyperhemisphere match the radiation to the planar antenna on the mixer substrate mounted inside the open structure mixer block.

improves mixer performance. The i.f. signal, after amplification at the 15 K stage, is fed into a room-temperature amplifier/filter/amplifier combination mounted to the outside of the dewar, which boosts the signal an additional 60 dB and defines a $-1 \text{ dB}$ passband of 600 MHz centered at 1.5 GHz. The shape of the i.f. passband is quite flat, varying about 1.0 dB pp; no padding is required to suppress standing waves on the $\sim 3.5$ ft of coaxial cable inside the dewar. For astronomical observations, the i.f. passband is upconverted from 1.5 to 2.0 GHz and analyzed with an acousto-optical spectrometer; for laboratory heterodyne measurements, a simple square-law detector (an HP 33331D planar doped barrier sensor) is used to measure total i.f. power.

The mixer design consists of a planar twin-slot antenna and a pair of $2.25 \mu m^2$ Nb-Al$_2$O$_3$-Nb junctions linked by an inductor. The mixer circuit is shown in Fig. 2. The dimensions of the antenna slots are $0.50 \lambda_0 \times 0.02 \lambda_0$ and the inter-slot spacing is $0.28 \lambda_0$, where $\lambda_0$ is the design wavelength; Ref. 5 contains detailed information about the twin-slot antenna characteristics. The mixers were fabricated at the Jet Propulsion Laboratory using optical lithography with three mask levels. The first mask defines the ground plane and twin-slot antenna, the second mask defines the junction area, and the third mask defines the wiring layer (contact pads, transmission lines, and radial stubs). The parallel junction circuit is fed at opposite ends by the two slots, which are $180^\circ$ out of phase. Since the junctions are driven antisymmetrically, an effective rf ground occurs midway between them at the center of the inductor. The inductance $L$ is therefore chosen so that $L/2$ resonates with the geometric capacitance of a single junction at 492 GHz. The junctions ($\sim 8 \Omega$) are matched to the antennas ($\sim 35 \Omega$) with two-section quarter-wave transformers; the transformers are rf coupled to the slots with radial stubs. Simulations with Touchstone and measurements with a Fourier transform spectrometer

FIG. 2. A photograph of the slot antenna mixer. The twin-slot antenna antisymmetrically feeds the two-junction array with a pair of two-section quarter-wave transformers.
FIG. 3. Two views of the quasioptical mixer block: (a) Front view, with the polyethylene lens removed, showing the quartz hyperhemispheric lens—the hyperhemisphere slows the f/0.5 beam of the twin-slot antenna to f/1.0. (b) Back view, with the electromagnet mounting plate removed. The mixer on a quartz mounting disk, the 15 Ω stripline, and discrete components of the mixer block circuit are visible. The dc lines, which are protected by 1 K series resistors and rf shorting capacitors, are attached to a Microtech connector (see Ref. 22). The bias tee and i.f. matching circuit are fabricated on pieces of RT/Duroid 6010 circuit board; the i.f. signal is brought out of the mixer block with an SMA connector.

demonstrate that this mixer design has a rf −3 dB bandwidth of ≈200 GHz.

Both front and inside views of the mixer block are shown in Fig. 3. In Fig. 3(a), the HDPE plano-convex lens which attaches to the front of the mixer block has been removed to show a quartz hyperhemisphere and the mixer substrate, which is visible behind it. The hyperhemisphere transforms the f/0.5 main beam of the planar antenna to approximately f/1.0, and the plastic lens further slows the beam to about f/15. The beam pattern of the receiver, which was measured by scanning a small 600 K blackbody source across the aperture, is shown in Fig. 4; the −10 dB width of the beam is 3.0°. The large f/# is convenient for optical matching to a telescope beam: the aperture coupling is less sensitive to small offsets relative to the beam axis, making it easier to align the LO and obtain the best coupling efficiency. Total optical losses, including vacuum window (0.32 dB), IR blocking filter (0.22 dB), plastic lens (0.41 dB), hyperhemisphere (0.66 dB), and antenna efficiency (1.55 dB), sum to 3.16 dB; these contribute to the total conversion loss of the mixer, which is discussed in the next section. In Fig. 3(b), the magnet assembly and mounting plate have been removed to show the mixer block interior. The i.f. matching network and mixer bias tee are fabricated on pieces of RT/Duroid 6010 circuit board attached to the inside front of the block. The mixer is wire bonded to a small 15 Ω stripline fabricated on a 0.25 mm thick piece of crystal quartz. The stripline, which feeds the i.f. matching network and carries the dc bias voltage, is epoxied to the quartz mixer mounting disk, on which a ground plane has been patterned.

The mixer block incorporates an electromagnet which suppresses Josephson tunneling currents in the mixer. The magnet coil, which is made from 0.1 mm Nb–Cu wire wound onto a cylindrical steel core, is mounted on the back plate of the mixer block. Steel yoke pieces attached to the ends of the coil penetrate the mixer block and make contact at right angles with flat pole pieces, which are clamped on opposite sides of the mixer mounting disk. The tapered ends of the pole pieces, which are suspended ~1.0 mm over the mounting disk, closely bracket the mixer. The concentrated field which develops between the pole tips adequately suppresses the Josephson effect when a few ×10 mA is applied to the superconducting magnet windings.

III. RECEIVER PERFORMANCE

All heterodyne tests of the receiver were performed with a free running local oscillator nominally set to 492 GHz. The local oscillator consists of a continuously tunable 88–100 GHz Gunn oscillator and a frequency quintupler. The SIS mixer current as a function of dc bias voltage is shown in Fig. 5 for three different conditions: (a) with no local oscillator or magnetic field applied; (b) with a small applied magnetic field, inadequate to suppress the Josephson tunneling currents completely, but enough to enhance a rf tuning structure resonance; (c) with the Josephson currents completely suppressed by the magnetic field and the junctions pumped.
by the local oscillator. Figure 5(a) shows the exceptional nonlinear characteristics of Nb-Al$_2$O$_3$-Nb junctions. The transition from superconducting to normal state occurs over a narrow bias voltage interval (~0.2 mV) centered at 2.8 mV. The conductance below the gap voltage appears to be somewhat high (~0.08 when normalized by the normal-state conductance), with a leakage current of 47 µA at 2.0 mV. Nonetheless, this device should show excellent quantum mixing, in accordance with theory. In Fig. 5(b), a resonance appears on the I-V curve at 0.7 mV, driven by ac Josephson currents of frequency 340 GHz, where the complex out-of-band impedance of the planar antenna resonates with the impedance of the two-junction tuning circuit. Circuit simulations predict an analogous resonance. In Fig. 5(c), the device is fully pumped and the Josephson currents are visibly suppressed. The characteristic step-like structure in the I-V curve is due to photon-assisted tunneling of quasiparticles (electrons). The width of the photon step (2.0 mV) corresponds to the electric potential (hν/e) of a 492 GHz photon. The SIS current at the bias voltage corresponding to the middle of the first photon step is the sum of contributions from both leakage and photon-assisted tunneling.

Figure 6 shows an oscilloscope trace of the i.f. signal power as a function of bias voltage for three situations: the receiver with a 290 K blackbody at the input (top curve), the receiver with an 80 K blackbody at the input (middle curve), and the receiver without LO power (bottom curve). A magnetic field has been applied to suppress Cooper pair tunneling. The straight line at the bottom indicates the zero level of the square-law detector. The broad feature in the i.f. power curve on the first photon step arises from quasiparticle mixing. The jagged features appearing at bias voltages less than 1.5 mV are noise power resulting from chaotic Josephson currents. It is clear that the receiver sensitivity varies with bias voltage across the first photon step. A measure of this sensitivity is the Y-factor, which is the ratio of i.f. power for hot/cold thermal loads. A Y-factor can be converted to a noise temperature with the standard expression,

\[ T_{\text{rec}} = \frac{[T_h - Y \cdot T_c] / (Y - 1)}{1}. \]  

The receiver is most sensitive when biased at a voltage where the i.f. power is roughly a maximum. The best noise temperatures (150–200 K) were obtained for the interval of voltages between the two i.f. subpeaks at 1.7 and 2.1 mV. Within this interval, the most stable bias voltage settings cor-
respond to extremum in the i.f. power; at these settings, Y-factors are highly repeatable, with less than 0.1 dB variations over several hours.

Tuning to obtain maximum sensitivity requires the adjustment of three parameters—the bias voltage, the local oscillator power, and the applied magnetic field—all of which are coupled, since the linked junction pair in the mixer forms a superconducting quantum interface device (SQUID) loop. Varying the bias voltage alters the SIS current, and hence slightly changes the magnetic flux in the SQUID loop. Likewise, varying the applied LO power will alter the SIS current with the same result. The SQUID effect is observed in the Josephson branch of the IV curve as a series of oscillations which accompany changes in the applied B-field. The SQUID effect is also evident by the varying number and amplitude of jagged noise features appearing in the i.f. power curve.

The procedure for tuning is iterative. First, the magnet current is slowly increased from zero to the lowest level (~40 mA) where the Josephson current branch of the IV curve nearly vanishes for some local SQUID minimum; care should be taken not to apply too high a magnetic field, which will suppress superconductivity in the mixer. The LO is applied and the power is adjusted until the heights of the two subpeaks on the i.f. power curve are about equal; this empirically gives the best sensitivity and conversion loss. The bias voltage is then fixed and the Y-factor is measured. Further small adjustments in the magnet current, bias voltage, and LO power may be necessary in order to optimize the Y-factor and to minimize interference from the Josephson effect.

Shot noise from an unpumped mixer can be used to evaluate the noise contribution from the i.f. chain referred to the input of the cooled amplifier. When biased above the gap voltage, the mixer emits power in a bandwidth Δν into the cooled amplifier input:

$$P_{\text{SIS}} = e I_{\text{bias}} R_h \Delta \nu.$$  \hspace{1cm} (2)

The output voltage of the square-law detector $V_{\text{det}}$ is linearly proportional to the net i.f. noise power:

$$V_{\text{det}} = a_{\text{calib}} (T_{\text{if}} + P_{\text{st}}/k_B \Delta \nu).$$  \hspace{1cm} (3)

where $T_{\text{if}}$ is the noise power of the i.f. chain referred to its input. By measuring $V_{\text{det}}$ for two bias settings and using Eqs. (2) and (3), we solve for $a_{\text{calib}}$, the calibration constant of the square-law detector output and $T_{\text{if}}$. The value we found for $T_{\text{if}}$, the combined noise due to all i.f. components and cables, is 12 K. With $a_{\text{calib}}$ and $T_{\text{if}}$ known, the mixer conversion loss can be derived from the hot/cold load tests. The power from the pumped mixer into the i.f. amplifier is

$$P_{\text{SIS}} = k_B [T_{\text{mix}} + T_{\text{load}}] \Delta \nu / L,$$  \hspace{1cm} (4)

where $L$ is the mixer conversion loss, $T_{\text{mix}}$ is the combined noise temperature of the mixer and dewar optics, referred to the rf input, and $T_{\text{load}}$ is the temperature of the blackbody load. Using Eqs. (3) and (4) evaluated for 295 and 80 K loads, we solve for $L$ and $T_{\text{mix}}$. The overall noise temperature of the receiver is related to $T_{\text{mix}}$ and $T_{\text{if}}$ by

$$T_{\text{rec}} = T_{\text{mix}} + T_{\text{if}} \approx 170 \, \text{K}.$$  \hspace{1cm} (5)

Note that $L$ includes the loss due to the rf optics.

Table I lists double side-band noise temperatures and conversion losses for several 492 GHz SIS receivers which have been recently constructed. These include the receiver described in this paper (QO1), an earlier generation quasioptical receiver (QO2), and two low-noise 492 GHz waveguide systems (WG1 and WG2).

The earlier generation quasioptical receiver, QO2, has a DBS receiver temperature more than twice that of QO1. Both receivers have similar rf optics, including a planar twin-slot antenna. However, the QO2 mixer has a single untuned 2.25 mm$^2$ junction coupled to the antenna with a tapered transmission line. The measured conversion loss of QO1 is nearly three times less than that of QO2, in agreement with our calculations, and demonstrating the improvement possible with an integrated matching circuit.

A comparison of QO1 to the waveguide receiver shows that it has comparable sensitivity, with a somewhat lower conversion loss. A reduction in the effective noise temperature of the i.f. chain in QO1 seems possible, given the lower values achieved for WG1 and WG2.

### IV. DISCUSSION

In this paper we have described the design and performance of a 492 GHz quasioptical SIS receiver with an integral twin-slot planar antenna matched to a pair of SIS junctions linked by an inductance. Our results show that quasioptical receivers can be constructed with conversion losses and noise temperatures comparable to those of the most sensitive waveguide receivers currently in operation.

Scans across the dewar aperture with a blackbody source indicate a main beam of f/15. Heterodyne tests show a DSB noise temperature of 170 K. Combined with measurements of shot noise from the unpumped mixer, these results imply a conversion loss of 5.4 dB, a mixer temperature of 128 K, and an i.f. chain noise temperature of 12 K.

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