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Supplementary Material for
**The multiple merger assembly of a hyperluminous obscured quasar at
redshift 4.6**

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Materials and Methods

Keck/LRIS spectrum

We observed W2246–0526 multiple times with the dual-beam Low Resolution Imaging Spectrometer (LRIS) (28) on the Keck I telescope. All observations used the 400 lines mm^{-1} grating on the red arm of the spectrograph (blazed at 8500 Å; resolving power $R \sim 750$), with the 5600 Å dichroic, and the 1.5" wide long slit. W2246–0526 is not detected shortward of the dichroic wavelength, and so we do not discuss the data from the blue arm further.

Our first spectrum, obtained on Universal Time (UT) 2010 November 8, entailed two 600-s exposures at a position angle (P.A.) = 49.9° (east of north). A relatively narrow ($\sim 600 \text{ km s}^{-1}$), asymmetric Lyman- α ($\text{Ly}\alpha$) emission line was detected (see Figure S1), with some faint $\text{Ly}\alpha$ nebulosity. Combined with a modest continuum break across the $\text{Ly}\alpha$ emission line and a low signal-to-noise ratio detection of a matching broad ($\sim 4000 \text{ km s}^{-1}$) C IV emission line, these data yielded a redshift of $z = 4.6$.

We obtained a deeper follow-up spectrum on UT 2013 October 4 consisting of three 1200-s exposures at a P.A. = 38° in photometric conditions. The deeper spectrum confirmed the features observed in the 2010 data, with higher signal-to-noise ratio (Figure S1). A $\text{Ly}\alpha$ redshift of 4.601, based on a Gaussian fit to the peak of the emission line, matches the ALMA-derived [C II] redshift of 4.601 ± 0.001 . The C IV line detected by LRIS yields a lower redshift of $z_{\text{CIV}} = 4.548$ (blue-shifted $\sim 2800 \text{ km s}^{-1}$ from the $\text{Ly}\alpha$ line) based on the peak of the highly asymmetric, broad feature. This is a larger offset than typically seen in lower redshift quasars (29), and more typical of the high redshift ($z > 6$) quasar population (30). Such large blue-shifts are suggestive of high radiative efficiencies and very strong winds (31).

We obtained a third spectrum of W2246–0526 on UT 2015 December 12 (Figure S1 inset), consisting of two 1200-s exposures and a single 462-s exposure at P.A. = 20° (Figure S2). Though the night was clear, there were 40-50 mph winds, which produced poor seeing due to wind-shake ($\sim 1.5''$). These data reproduce the features seen previously, namely narrow $\text{Ly}\alpha$ emission and blue-shifted, broad C IV emission. The 2D spectrum also showed a faint, slightly red-shifted $\text{Ly}\alpha$ emission feature extending $\sim 3''$ to the northeast of W2246–0526, spatially coincident with the companion galaxy C3 (see Figure S2), further supporting its detection and the derived redshift estimated from the [C II] ALMA data. Unfortunately, the low angular resolution of the optical spectrum does not allow the identification of $\text{Ly}\alpha$ emission lying between W2246–0526 and C3.

ALMA 212 μm Dust Continuum Data

The ALMA observations of W2246–0526 in the $\sim 212 \mu\text{m}$ dust continuum emission were obtained in Band 6 (211–275 GHz) during 3 execution blocks, two on 2016 June 20 and

one on 2016 July 13. The on-source integration time was 49 min per block, for a total of 2.5 hours. The number of antennas used in each run was 38, 41 and 39, respectively. The minimum and maximum baseline lengths were 15.1 and 704.1 m. The sources Pallas and QSO J2148+0657 were used for amplitude calibration, the source QSO J2232+1143 was used for pointing and band-pass calibrations, and the source PMN J2243–0609 was used for phase calibration.

We used the Common Astronomy Software Application (CASA; v.5.1.1) (32) to process and clean the ALMA products. All the execution blocks of the source were concatenated in a single measurement set for the analysis. The cleaning algorithm was run using the task *tclean* in parallelized process, with a *briggs* weighting scheme, a *hogbom* deconvolver, and a *robust* parameter set to = 2 (similar to natural weighting) for the u - v visibility plane. A single circular aperture of radius = 1" was used to mask the core of W2246–0526, cleaning each window down to a depth of 2 times the r.m.s.. The angular size of the restoring beam ranges from 0.52" \times 0.44" in the spectral window (SPW) 0 (central observed frequency $\nu_{\text{obs}} \sim 260.929$ GHz; rest wavelength $\lambda_{\text{rest}} \sim 205$ μm) to 0.56" \times 0.48" in the SPW 3 (central $\lambda_{\text{rest}} \sim 219$ μm), with an average of 0.54" \times 0.46" at an effective $\lambda_{\text{rest}} \sim 212$ μm , or $\sim 3.6 \times 3.1$ kpc at the redshift of W2246–0526. The P.A. of the beam is $\sim 80^\circ$. The average r.m.s. of the data cube is ~ 125 $\mu\text{Jy beam}^{-1} \text{ channel}^{-1}$, measured in 48 km s^{-1} channels (averaged over 5 original channels). To create the moment-0 map shown in Figure 1 and Figure 2S, we collapsed the dust continuum emission of all four SPWs after discarding bad channels at the edges of the SPWs, reaching a r.m.s. = 11.5 $\mu\text{Jy beam}^{-1}$.

Figure 1 shows the dust continuum map of the W2246–0526 merger system down to a lowest contour of $2.5 \times \sigma$, where σ is the r.m.s. of the background. This is a sufficiently high threshold for the map not to show a large number of regions with negative flux, implying that the fidelity of the positive emission detected is high. However, for reference, we also show in Figure S2A a map where an additional contour is plotted at $2 \times \sigma$ (as well as the complementary $-2 \times \sigma$ contour). While noisier, this figure shows that most of the positive emission recovered at the $2 \times \sigma$ level is located around the tidal tail and the system of galaxies in general.

To check the reliability of the extended emission in the dust map, we processed and analyzed each execution block independently, which were obtained at different periods of time, thus having different u - v coverage. We followed the same cleaning procedure used for the combined dataset described above. Albeit with a lower signal-to-noise, extended emission was detected in between W2246–0526 and the companion galaxy C2 in every execution block. We have also processed and analyzed the dataset of the “check source”. The check source is a bright, point-like object that is always observed at the end of an execution block, which can be used to double-check the calibration of the science target. The image of the check source does not show any trace of residual/artifact emission in the direction of the tidal tail detected in W2246–0526, which was observed with the same configuration.

We also cleaned the image of W2246–0526 using a uv -tapering resulting in a FWHM $\sim 1''$. The tidal tail is recovered with a higher significance at larger angular scales (Figure S2B).

VLA CO(2→1) Data

Observations of the redshifted $^{12}\text{CO } J = 2-1$ emission line ($\nu_{\text{rest}} = 230.538$ GHz; $\lambda_{\text{rest}} \sim 1.3$ mm) in the W2246–0526 system were obtained using the VLA in 5 observing runs between 2015 November 02 and 2015 December 07. The observations were taken in D-configuration with 27 antennas, using the Q-band receivers (tunable range: 40–50 GHz). At the redshift of W2246–0526 the CO(2→1) line is redshifted to 41.160 GHz.

The observations were performed with the Wideband Interferometric Digital ARchitecture.(WIDAR) correlator, with two basebands (AC and BD) of eight contiguous SPWs each. Baseband AC covered the SPWs 0–7 while baseband BD covered SPWs 8–15. Each SPW was set to have 64 channels and 2 MHz per channel resolution. Both basebands were configured to overlap, with center frequencies 41.128 GHz and 41.228 GHz, for AC and BD respectively, making an effective bandwidth of 1.120 GHz. With this setup, the SPWs 4 and 11 covered the redshifted CO(2-1) line, being tuned to 41.160 and 41.192 GHz, respectively.

The nearby quasar QSO J2229-0832 was used for gain and pointing calibration and the source 3C 048 served as flux and bandpass calibrator. The data were calibrated using the Astronomical Image Processing System (AIPS) (33) and CASA (32). Time ranges with poor visibilities as well as edge channels where the bandpass deteriorated (at each edge) were flagged. The data were imaged using the *tclean* algorithm in CASA. All images were primary beam corrected. We used a *briggs* weighting scheme, a *hogbom* deconvolver, and cleaned down to 2σ in a circular aperture of $\sim 15''$ -diameter around our target. Setting the *robust* parameter = 2 (similar to natural weighting), results in an angular resolution of $2.47'' \times 2.01''$ (P.A. = 10.4°). We tried other weighting schemes, but found that most of the emission was not recovered when using lower values for the robust parameter. The final cube is shown in Figure 2 and has a r.m.s. = $85 \mu\text{Jy beam}^{-1} \text{ channel}^{-1}$ measured in 60 km s^{-1} channels. The continuum emission under the CO line was not detected. The r.m.s. of the collapsed cube is $\sim 10 \mu\text{Jy beam}^{-1}$.

Figure S3 presents a comparison between the spectra of the CO(2→1) and [C II] 158 μm emission lines (16). Both lines display high velocity dispersions, FWHM $\sim 550 - 600 \text{ km s}^{-1}$, suggesting that both the neutral and molecular gas phases of the ISM are very turbulent, probably reflecting the energy and momentum that are being injected by the central AGN on its surrounding medium.

Cosmology

Throughout the paper, we adopt a cosmology with the following parameters: $\Omega_M = 0.28$, $\Omega_\Lambda = 0.72$ and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. At the redshift of W2246–0526, $z = 4.601$, the angular scale is $1'' = 6.68 \text{ kpc}$.

Star Formation Rate from the [C II] Emission Line

The [C II] 158 μm emission line has been proposed as a reliable SFR tracer in normal, Milky Way-like, star-forming galaxies as the [C II] luminosity correlates well with other SFR indicators (34). However, the [C II] luminosity of IR-bright sources fall below this trend, displaying a “deficit” of line emission with respect to the total IR emission (35-37), which scales linearly with the SFR in dust-obscured objects. The [C II] to IR luminosity ratio ($L_{[\text{C II}]} / L_{\text{IR}}$) in luminous infrared galaxies can be, in extreme cases, a factor of ~ 20 smaller than in normal star-forming galaxies, which show a typical value $L_{[\text{C II}]} / L_{\text{IR}} \sim 5 \times 10^{-3}$ (36). Thus, if no additional information other than the IR luminosity of the source is available, a detection of the [C II] emission can only provide a lower limit to the SFR of the galaxy. We use a calibration between L_{IR} and SFR (38), with $\text{SFR}_{\text{IR}} [\text{M}_\odot \text{ yr}^{-1}] = 1.5 \times 10^{-10} L_{\text{IR}} [L_\odot]$, based on the Starburst99 stellar evolution synthesis models (39), and assuming a constant SFR history over 100 Myr, a Kroupa initial mass function (IMF) (40), and that the entire Balmer continuum emitted by the starburst is absorbed and re-radiated by optically thin dust.

The [C II] luminosity of W2246–0526 measured with a $1''$ -diameter aperture is $L_{[\text{C II}]} = 6.1 \times 10^9 L_\odot$ (16). The line is spatially resolved, with up to $\sim 55\%$ of the emission arising from an extended component (measured in the same aperture). The extended emission thus accounts for $L_{[\text{C II}]} \sim 3.4 \times 10^9 L_\odot$. If we assume that the extended [C II] emission is powered entirely by star formation, then using the $L_{[\text{C II}]} / L_{\text{IR}}$ ratio upper limit for normal galaxies and the SFR calibration described above, we estimate that the lower limit to the SFR of the host is $\gtrsim 100 \text{ M}_\odot \text{ yr}^{-1}$. If the star formation properties of the underlying galaxy in W2246–0526 are instead similar to local, purely star-forming ULIRGs, its SFR could be as high as $\sim 1000 \text{ M}_\odot \text{ yr}^{-1}$. For reference, the SFR derived via SED fitting for the central $\sim 4''$ of W2246–0526 is $\sim 560 \text{ M}_\odot \text{ yr}^{-1}$ (see below).

The [C II] luminosities of the companion galaxies C1, C2 and C3 are $(9.1, 6.6, \gtrsim 2.3) \times 10^8 L_\odot$, respectively. Assuming that all the IR luminosity in these galaxies is associated with star formation, we estimate SFR lower limits of $\gtrsim 27, 20$ and $7 \text{ M}_\odot \text{ yr}^{-1}$, respectively, and upper limits in the range of $200\text{--}500 \text{ M}_\odot \text{ yr}^{-1}$, when using a $L_{[\text{C II}]} / L_{\text{IR}}$ ratio typical of ULIRGs.

Calculation of Dust and Gas Masses

To estimate the dust mass content of W2256–0526 and its neighboring galaxies, we use a modified blackbody function scaled to the continuum flux density of each source, as

measured at the average rest-frame wavelength of the ALMA data cube, $\lambda_{\text{rest}} \sim 212 \mu\text{m}$. We further assume that the dust emission is optically thin, with an emissivity index $\beta = 1.8$ (41). We adopt a dust mass opacity $\kappa_{\nu, \text{dust}}(850 \mu\text{m}) = 0.0484 \text{ m}^2 \text{ kg}^{-1}$, which assumes a gas-to-dust mass conversion factor typical of local solar-metallicity galaxies, $\delta_{\text{GDR}} = 100$ (41,42), consistent with the gas mass estimates based on the CO(2 \rightarrow 1) emission line. The adopted $\kappa_{\nu, \text{dust}}$ is close to the value $\kappa_{\nu, \text{dust}}(800 \mu\text{m}) = 0.04 \text{ m}^2 \text{ kg}^{-1}$ obtained by (43). Under these assumptions, the dust mass is $M_{\text{dust}} = (S_{\nu, \text{obs}} \times d_L^2) / (\kappa_{\nu, \text{rest}} \times B(T_{\text{dust}}) \times (1+z))$, where $S_{\nu, \text{obs}}$ is the measured flux density of the dust continuum at the observed frequency ν_{obs} , d_L is the luminosity distance to the source, $B(T_{\text{dust}})$ is the Planck function evaluated at the dust temperature, T_{dust} , and z is the redshift of the galaxy.

This approach usually assumes that most of the dust and ISM mass in a galaxy is accounted by its coldest component (22). As such, instead of using a luminosity-weighted dust temperature in the equation above, $\langle T_{\text{dust}} \rangle_L$ (derived from fitting the IR spectral energy distribution (SED) of the galaxy, or the position of its far-IR emission peak), T_{dust} is normally replaced by a mass-weighted temperature, $\langle T_{\text{dust}} \rangle_M = 25 \text{ K}$ in normal star-forming galaxies, more representative of the bulk of their cold dust and ISM content. However, given the extreme luminosities and dust temperatures of Hot DOGs this assumption may not be appropriate, and could lead to overestimate their dust masses. Using a single-temperature modified blackbody model, the coldest component of the T_{dust} distribution in a sample of Hot DOGs is found to be in the range $\sim 60\text{--}120 \text{ K}$ (4). Thus, to calculate the dust mass in W2256–0526, we adopt a reasonable temperature range of $T_{\text{dust}} = 50\text{--}100 \text{ K}$ for the Rayleigh-Jeans tail emission, with the lower end providing the largest M_{dust} estimations. For the companion galaxies and extended emission in general, we use the narrower range of $T_{\text{dust}} = 25\text{--}50 \text{ K}$. A summary of the dust mass estimates can be found in Table S1. For a discussion of the possible heating mechanisms of the dust in the tidal tail, see below.

At the average rest-frame wavelength of $\sim 212 \mu\text{m}$, the observed flux density of the dust continuum emission within the central $\sim 7 \text{ kpc}$ ($1''$ -diameter) of W2246–0526 (i.e., without including its neighbors) is $2.3 \pm 0.1 \text{ mJy}$. This corresponds to a dust mass in the range $M_{\text{dust}} \sim 5.6\text{--}17 \times 10^8 \text{ M}_{\odot}$ ($T_{\text{dust}} = 100\text{--}50 \text{ K}$, respectively). The flux density of the extended structures and diffuse emission around W2246–0526 (including the neighbor galaxies as well as the tidal tail and bridges) is $2.6 \pm 1.3 \text{ mJy}$, and thus has a similar estimated dust mass. The companions C1, C2 and C3 have, respectively, individual flux densities of 0.27, 0.23 and 0.16 mJy in a $1''$ -diameter aperture, and account for $\sim 25\%$ of the extended emission. Using a gas-to-dust mass ratio $\delta_{\text{GDR}} = 100$, the total gas reservoir in W2246–0526 alone is $0.6\text{--}1.7 \times 10^{11} \text{ M}_{\odot}$, and the entire system could contain as much as $M_{\text{gas}} = 1.2\text{--}3.6 \times 10^{11} \text{ M}_{\odot}$.

Measuring the CO(2 \rightarrow 1) line flux in a circular aperture with a radius equivalent to the beam size of the observations ($\sim 4.4''$ -diameter, which includes W2246–0526 and C1, see Figure 1) and following (44), the luminosity of the line is $L_{\text{CO}(2\rightarrow 1)} = 3.3 (\pm 0.7) \times 10^7 \text{ L}_{\odot}$. Assuming that the CO is in local thermal equilibrium up to the J = 2–1 transition and using a $M_{\text{gas}}\text{-}L_{\text{CO}(1\rightarrow 0)}$ ratio, α_{CO} , typical of local normal, Milky-Way-like galaxies ($\alpha_{\text{CO, MW}} = 4.6 \text{ M}_{\odot} [\text{K km s}^{-1} \text{ pc}^2]^{-1}$) (41), which does not include the factor of 1.36 to account for the gas

mass in helium and heavier elements), we obtain a gas mass of $M_{\text{gas}} \sim 3.9 \times 10^{11} M_{\odot}$. If an α_{CO} conversion factor typical of nearby ULIRGs is used ($\alpha_{\text{CO,ULIRG}} = 0.8 M_{\odot} [\text{K km s}^{-1} \text{pc}^2]^{-1}$) (45), then $M_{\text{gas}} \sim 0.7 \times 10^{11} M_{\odot}$. The estimate based on the 212 μm dust continuum flux density (3.9 mJy) measured in the same aperture is in the range $\sim 1.0 - 2.8 \times 10^{11} M_{\odot}$ (using $\delta_{\text{GDR}} = 100$), which slightly favors an α_{CO} conversion factor closer to the ULIRG regime. The stellar mass of the galaxy measured in similar aperture ($r \sim 2''$) is $2.5 \times 10^{11} M_{\odot}$ (see below). Thus, the gas depletion time scale would be 125–700 Myr, depending on the α_{CO} employed.

The total gas mass measured in a 11''-diameter (~ 73 kpc) aperture that includes the bulk of the CO(2 \rightarrow 1) emission (i.e., the entire merger system: W2246–0526 and C1, as well as C2, C3, tidal tail and bridges) is $M_{\text{gas}} \sim 1.5 \times 10^{11} M_{\odot}$, using $\alpha_{\text{CO,ULIRG}}$. This is in agreement with the range obtained from the dust-based measurement ($M_{\text{gas}} = 1.2 - 3.6 \times 10^{11} M_{\odot}$). An $\alpha_{\text{CO,MW}}$ would yield a $M_{\text{gas}} \sim 8.6 \times 10^{11} M_{\odot}$.

Possible Heating Mechanisms for the Dust in the Tidal Tail

Physical mechanisms that could heat the dust in the tidal tail connecting W2246–0526 and the companion galaxy C2 are: (a) direct illumination from the AGN, and (b) in-situ star formation.

Scenario (a) assumes that dust in the tidal tail may be heated by the central AGN. Radiative transfer models can reproduce the spectral energy distribution of AGNs using spherical geometries where dust is smoothly distributed in a cloud that can be as large as ~ 15 kpc, surrounding the central energy source (46). The equilibrium temperature of dust heated by a central source (T_{dust}) can be calculated as a function of its luminosity and distance (47). For a bolometric luminosity equal to that of W2246–0526, dust can be heated up to $T_{\text{dust}} = 25$ K at distances up to ~ 7.5 and 100 kpc, for optically thick and thin material, respectively. That is, the central energy source is more likely to heat dust to a high temperature at distances $> 1''$ (see Figure 1) if the obscuring medium around W2246–0526 is effectively optically thin, such that the UV light is not quickly reprocessed by intervening material. This suggests that the dust geometry surrounding W2246–0526 would need to be clumpy, because virtually all the UV continuum is fully reprocessed (optically thick) along our line of sight, but the direction from the AGN towards the tidal tail would need to be effectively free of dust (mostly optically thin). If this were the case, we should see a gradient of temperatures along the tidal tail, which extends for ~ 35 kpc, related to the distance at which the AGN can heat the dust at a certain temperature. While the dust column density could compensate for the temperature gradient such that we would observe a uniform continuum flux density along the tail, such a scenario is rather complex.

Scenario (b) is that the dust in the tidal tail may be heated by in-situ star formation. This is disfavored by the fact that, aside from a faint detection of the knot K1, the dust emission connecting C2 and W2246–0526 as well as the bridge to C3 does not have associated UV emission detected in the HST image (Figure 2). However, the non-detection of UV

continuum may simply be caused by dust obscuration (e.g., 48), or the shallow depth of the HST image, which probes only down to a sensitivity of $\text{SFR} \sim 6 M_{\odot} \text{ yr}^{-1}$.

Alternatively, the dust could be cooling, without being exposed to any source of heating. However, this case is highly unlikely since dust is a very efficient coolant. The cooling rate of dust at $T_{\text{dust}} > 50 \text{ K}$ is extremely short ($\Lambda_{\text{dust}} \propto n T_{\text{dust}}^{4+\beta}$, where n is the number density of hydrogen atoms), reaching equilibrium with the cosmic microwave background (CMB) within years, thus making it effectively undetectable. In addition, while cooling, any gas clump in the tidal tail would undergo gravitational collapse. The free-fall time of a parsec-sized clump with a mass of $10^6 M_{\odot}$ is a few kyr, thus suggesting that star-formation should have already started to take place in the gas stream (if the shear is not too strong). We therefore favor scenario (b).

Estimation of the Global Stellar Mass and Star Formation Rate of the Host Galaxy Through SED Fitting

We have constructed the optical to far-IR spectral energy distribution (SED) of W2246–0526 (Figure S4) using photometry from the following telescopes: HST Wide Field Camera 3 at $1.6 \mu\text{m}$ with the F160W filter (16), Palomar200 at $2.2 \mu\text{m}$ with the Ks filter (10), Spitzer at 3.6 and $4.5 \mu\text{m}$ with the IRAC1 and IRAC2 filters (49), WISE at 12 and $22 \mu\text{m}$ with the W3 and W4 filters (3), Herschel at 250 , 350 and $500 \mu\text{m}$ with the Spectral and Photometric Imaging Receiver (SPIRE) (14), ALMA at $865 \mu\text{m}$ and 1.19 mm in Bands 7 and 6 (16, and see above), and VLA at 7.28 mm in Q-band (see above). In all cases we collected the photometry with the smallest aperture available to avoid including emission from the neighbor galaxies. While the angular resolution of HST and ALMA allows us to easily separate all companions, the Spitzer, WISE and Herschel angular resolutions are coarser, and thus the fluxes could be contaminated. We use the AllWISE catalog to investigate the mid-IR WISE photometry, which is extracted via a point-spread-function (PSF) fitting procedure that also provides the goodness of fit of the model. In all WISE bands W2246–0526 is unresolved (reduced $\chi^2 \leq 1$) and no deblending was needed. Therefore, any contamination of the extended companions to the mid-IR fluxes is negligible. Regarding the near-IR bands, the Spitzer photometry was obtained with a $6''$ -diameter circular aperture using SExtractor (49). While SExtractor performs a deblending process in crowded fields, extended emission close to C3 could contribute up to 30% of the total flux in IRAC1 and 15% in IRAC2. If we were to subtract these contributions from the Spitzer photometry, the estimated stellar mass (M_{\star}) of W2246–0526 (see below) would be reduced by $\sim 20\%$. Thus, even in this worst-case scenario, the change in M_{\star} would be well within the error of a factor of two quoted below. The Herschel photometry is not used for any fitting, just to guide the eye and compare them with the extrapolation to shorter wavelengths of the fit to the ALMA data, so we do not investigate contamination of the Herschel photometry.

To estimate the stellar mass of W2246–0526 we follow the same approach as (10), who modeled the optical to mid-IR SEDs of 96 Hot DOGs using the algorithm and SED templates for AGN and galaxies from (50). Briefly, these are four empirical SED templates

that span the wavelength range from 0.03 to 30 μm . An object is modeled as a non-negative combination of an old stellar population (E template), an intermediate stellar population (Sbc template), a young stellar population (Im template) and an AGN. We also fitted a reddening component to the AGN template, with a dust extinction law that is a combination of a Small Magellanic Cloud (SMC) reddening law at shorter wavelengths and a Milky Way reddening law at longer wavelengths. Intergalactic medium absorption is also considered. We do not include an extinction component for the host galaxy because the templates already include a small amount of intrinsic reddening (50) and the model does not require further obscuration. We discuss below the effects of unaccounted dust obscuration for the SED model of W2246–0526.

Fig S4 shows the best-fitting SED model for W2246–0526. The host galaxy dominates the rest-frame optical wavelengths while the AGN dominates at $\lambda_{\text{rest}} \gtrsim 1 \mu\text{m}$. We do not expect the optical spectrum of W2246–0526 to be contaminated by direct light from the AGN because i) the HST image shows that the emission from the core is resolved and ii) the considerable dust obscuration expected towards the central engine. Nevertheless, at $\sim 3000 \text{ \AA}$ there could be a contribution from scattered AGN light, as has been seen in a few Hot DOGS (51), which would result in a lower stellar mass. Figure S4 shows that the observed-frame Ks-band flux density is below the best-fitting SED model, and may indicate some obscuration toward the host galaxy. However, the IRAC bands are well fitted by the model, suggesting the host-galaxy obscuration is not enough to distort the redder parts of the SED, particularly the rest-frame near-IR. If we disregard the HST F160W band and allow for host reddening, then we find that the best-fitting SED still uses the same templates, and only requires a color excess $E(B-V) = 0.05$ for the host. The stellar mass estimate presented below would be only 5% lower, however, which is negligible. We therefore remove the AGN contribution from the best-fitting SED model and estimate the stellar mass using the rest-frame H-band luminosity of the host galaxy template. To do this, we follow (10) and derive the mass-to-light ratio of the host using the code *ezgal* (52) and the Galaxev stellar population SED models (53), which include a contribution from thermally pulsing-asymptotic giant branch (TP-AGB) stars, with a metallicity $Z = 0.008$ ($\equiv 0.4 Z_{\odot}$, where Z_{\odot} is the metallicity of the Sun) and a Chabrier IMF (54).

We consider exponentially declining SFR histories with different decaying timescales, τ . If we assume that the starburst started 100 Myr ago (i.e., at a redshift $z_f = 4.9$) and a $\tau = 1$ Gyr (similar to a constant SFR history), we estimate a stellar mass of $M_{\star} = 1.7 \times 10^{11} M_{\odot}$. A shorter decaying timescale of $\tau = 100$ Myr would increase the estimated stellar mass by only 12% to $M_{\star} = 1.9 \times 10^{11} M_{\odot}$. If we instead assume that the starburst started 500 Myr ago (i.e., $z_f = 6.7$) and a $\tau = 1$ Gyr, the derived stellar mass is $M_{\star} = 3.4 \times 10^{11} M_{\odot}$. We do not consider estimates for shorter τ , as the current SFR would be lower than a few hundred $M_{\odot} \text{ yr}^{-1}$, in disagreement with our measurements based on the IR luminosity (see below). Exponentially increasing SFR histories would yield smaller stellar masses than exponentially declining or nearly constant ones.

ALMA observations showed that the dust continuum flux density at rest-frame $\sim 158 \mu\text{m}$ is 7.4 mJy (16). At least 72% of the flux is accounted for by a central point source, with the remaining 28% being in an extended component, and hence likely powered by star-

formation activity (16). If we assume that the SED of the host galaxy is well described by a starburst similar to that in M82 (55), the 8–1000 μm luminosity would correspond to a $\text{SFR} = 560 \text{ M}_{\odot} \text{ yr}^{-1}$ (56). If we instead assume the host galaxy SED corresponds to the Sd spiral template of (55), the implied SFR would be $190 \text{ M}_{\odot} \text{ yr}^{-1}$, although the compactness of the host galaxy and its turbulent nature (16) makes it very unlikely that there is no ongoing starburst. A $\text{SFR} = 560 \text{ M}_{\odot} \text{ yr}^{-1}$ is in rough agreement with all the SFR histories used above to estimate the stellar mass of the galaxy, which predict current SFRs $\sim 500 - 1500 \text{ M}_{\odot} \text{ yr}^{-1}$.

Considering that the uncertainties in stellar mass estimates are inherently high (57), all values derived above are roughly consistent. Hence, for simplicity, we adopt the following properties for the integrated emission ($r \sim 2''$) of W2246–0526: $M_{\star} = 2.5 \times 10^{11} \text{ M}_{\odot}$ and $\text{SFR} = 560 \text{ M}_{\odot} \text{ yr}^{-1}$, both with an uncertainty of a factor of 2. The SSFR of the system derived from these values is $\sim 2.2 \text{ Gyr}^{-1}$, equal to a stellar mass doubling-time of $\sim 450 \text{ Myr}$, both with an uncertainty of a factor of three.

Calculation of the Dynamical Mass and Accretion Rate

The velocity shear of the [C II] line is small (16) and so we can calculate, based on the [C II] kinematics, the dynamical mass of W2246–0526 by assuming that it is dispersion-supported. We obtain $M_{\text{dyn}} = 0.8 (\pm 0.4) 10^{11} \text{ M}_{\odot}$, for an intrinsic velocity width of the line $\text{FWHM} = 550 (\pm 100) \text{ km s}^{-1}$, and a deconvolved physical diameter of the detected emission of $d = 2.5 (\pm 1) \text{ kpc}$. While the choice of using a dispersion-supported scenario for the calculation of M_{dyn} is justified by the line’s large FWHM, having the system partially supported by rotation would allow for a larger M_{dyn} if the disk of the galaxy is inclined (face on).

The stellar mass of the system within a $4''$ -diameter ($\sim 27 \text{ kpc}$) is $M_{\star} \sim 2.5 \times 10^{11} \text{ M}_{\odot}$ (see above). If M_{\star} scales with aperture in the same way as M_{dust} , then $M_{\star} \sim 1.5 \times 10^{11} \text{ M}_{\odot}$ in W2246–0526 alone. If we use the gas mass derived from the CO(2 \rightarrow 1) observations within the central $\sim 4.4''$ and scale it down similarly, then $M_{\text{gas}} \sim 0.4 \times 10^{11} \text{ M}_{\odot}$ in W2246–0526 (assuming $\alpha_{\text{CO,ULIRG}}$). Thus, the total (baryonic) mass $M_{\text{bar}} = M_{\star} + M_{\text{gas}}$ agrees with M_{dyn} within a factor of ~ 2 , which is well within the uncertainties. The excess of M_{bar} over M_{dyn} puts constraints on the conversion factors used to obtain M_{gas} from the CO(2 \rightarrow 1) line and from the dust continuum in W2246–0526 alone, somewhat favoring the ULIRG α_{CO} conversion for CO(2 \rightarrow 1) and the higher end of the T_{dust} range (100 K) for the dust-based estimate. This agrees with the scenario of the dust within the central core of W2246–0526 ($r < \text{few kpc}$) being heated by the central AGN to $T_{\text{dust}} \sim 100 \text{ K}$ (4).

Assuming $M_{\text{bar}} = M_{\star} + M_{\text{gas}} \sim 4 \times 10^{11} \text{ M}_{\odot}$ for the entire merger system (see above), the free-fall time, t_{ff} , at the position of the most distant, spectroscopically identified companion galaxy (C3, at $\sim 35 \text{ kpc}$ from W2246–0526) is $t_{\text{ff}} \sim 170 \text{ Myr}$ ($\sim 100 \text{ Myr}$ if an $\alpha_{\text{CO,MW}}$ is used to derived the gas mass from the CO(2 \rightarrow 1) line). This is equivalent to an average inflow velocity of $\sim 200 \text{ km s}^{-1}$ (330 km s^{-1} for $\alpha_{\text{CO,MW}}$). Assuming the most conservative

value of $t_{\text{ff}} = 170$ Myr, and that only the gas mass contained in the tidal tail ($M_{\text{gas}} \sim 4.6 \times 10^{10} M_{\odot}$, assuming $T_{\text{dust}} = 50$ K, see Table S1) is funneled to the center of W2246–0526 while most of the gas within the central few kpc remains turbulent and buoyant as suggested by the kinematics of the [C II] and CO(2→1) lines (16; and Figure S3), we calculate that the average accretion rate could reach $dM_{\text{gas}}/dt \sim 270 M_{\odot} \text{ yr}^{-1}$. If the self-gravitating companion galaxies are also accreted (combined $M_{\text{gas}} \sim 4.8 \times 10^{10} M_{\odot}$), then $dM_{\text{gas}}/dt \sim 550 M_{\odot} \text{ yr}^{-1}$. The free-fall time is a lower limit to the true time-scale of accretion, but we have chosen the most conservative values for the estimations of the gas masses and the free-fall time itself. That is, these have been derived using an $\alpha_{\text{CO,ULIRG}}$ conversion factor and the upper bound of T_{dust} ranges. Instead, if an $\alpha_{\text{CO,MW}}$ is used to calculate t_{ff} , then dM_{gas}/dt can reach up to $\sim 900 M_{\odot} \text{ yr}^{-1}$, or even higher if the gas masses of the tidal tail and companion galaxies are obtained using the lower end of the dust temperature range.

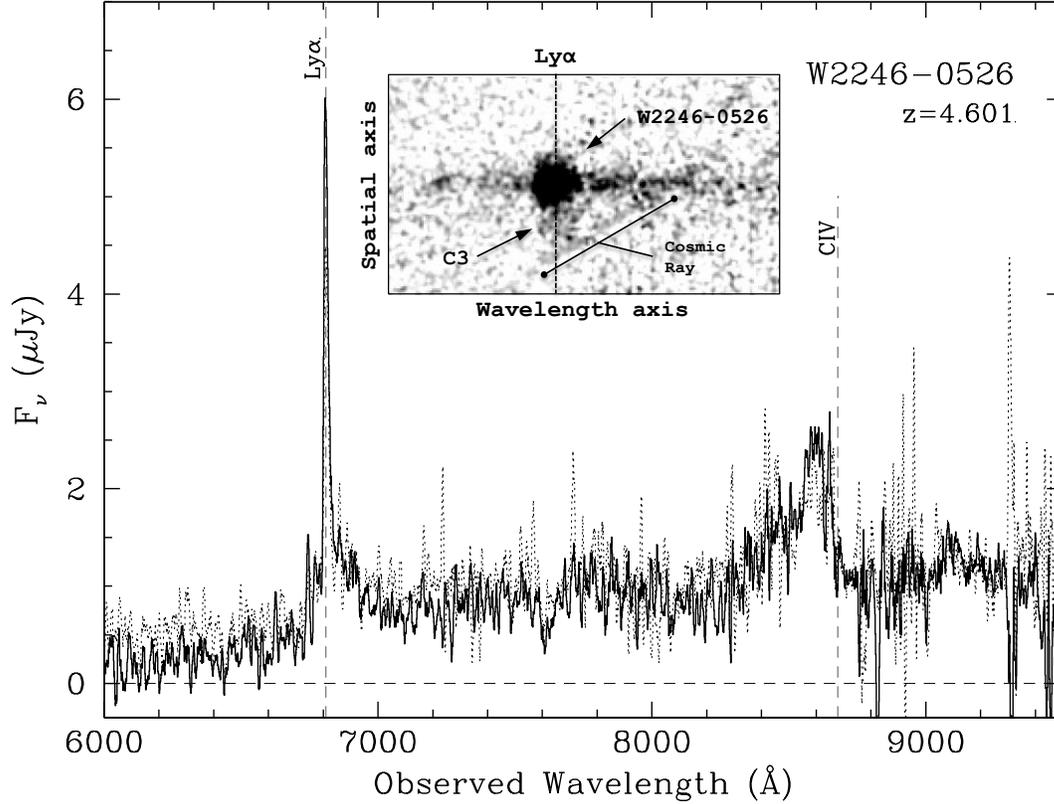


Figure S1: **Optical spectra of W2246–0526 obtained with Keck/LRIS.** F_ν is the flux density per unit frequency, which is shown as a function of the observed wavelength. The spectrum obtained in October 2013 is shown as a solid line, and the November 2010 spectrum is over-plotted as a dashed line. The inset figure shows the two-dimensional spectrum obtained in December 2015, smoothed to highlight the Ly α emission from the companion galaxy C3. The vertical axis corresponds to the slit spatial position and extends $\sim 20''$. The horizontal, wavelength axis is centered around 6800\AA . The narrow Ly α emission in the one-dimensional spectra of W2246–0526 is asymmetric, which is typical at high redshift, with a sharp edge to the blue side of the line due to foreground and associated absorption. The indicated ALMA-derived [C II] $158\ \mu\text{m}$ redshift matches the redshift derived from the peak of the Ly α line. The strongly asymmetric C IV line is much broader and is centered blue-ward of the Ly α and ALMA [C II] redshifts.

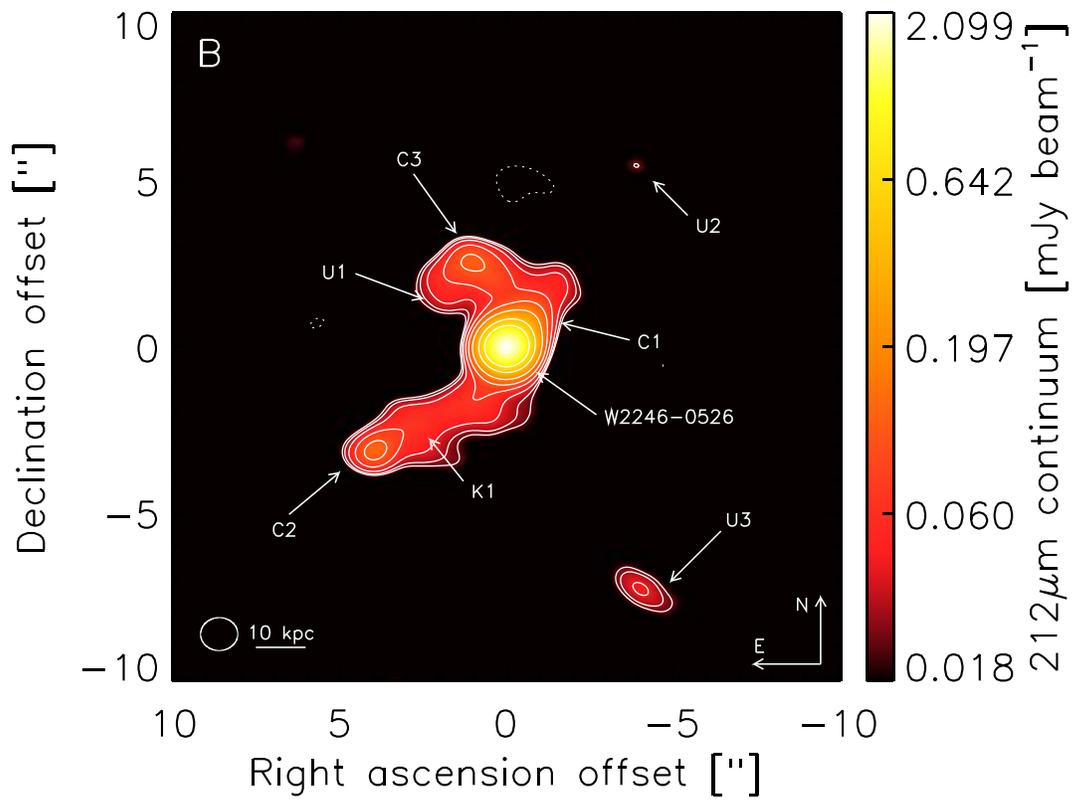
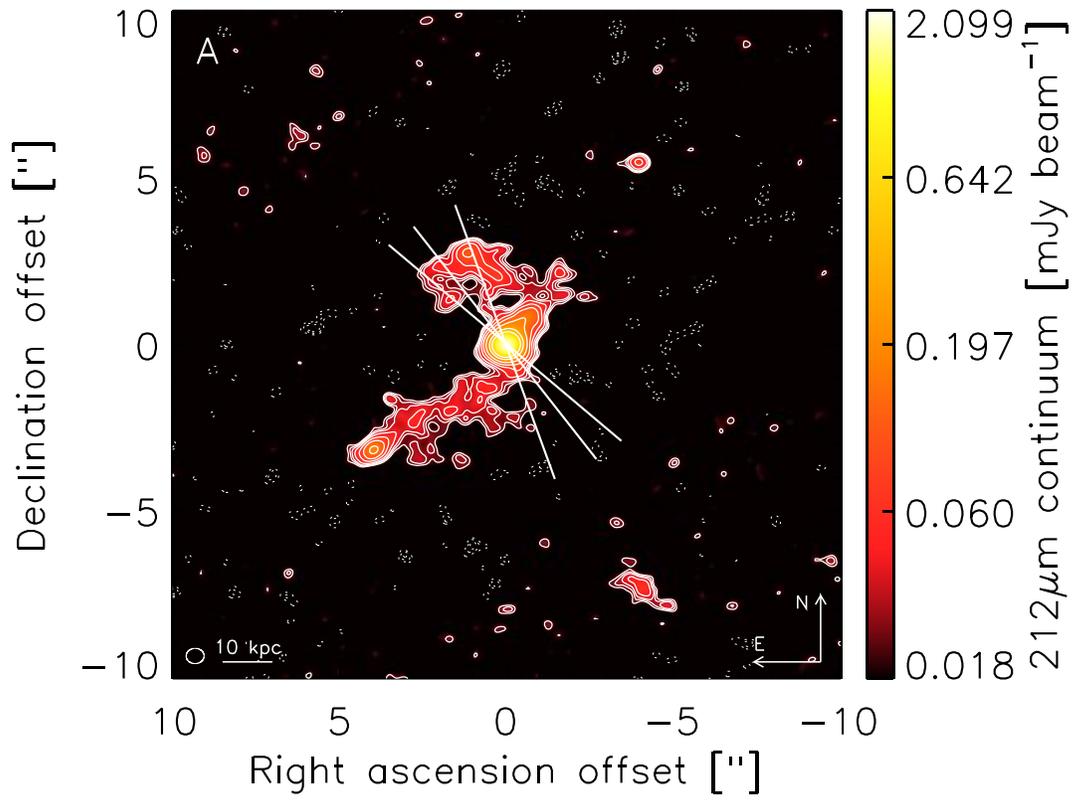


Figure S2: **ALMA 212 μm dust continuum emission maps of the W2246–0526 merger system.** (A) Same as Figure 1 but including contour levels at $[+2, -2] \times \sigma$, where σ is the measured r.m.s. of the background. In addition, solid straight lines indicate the position angle of the slits used to obtain the optical spectra shown in Figure S1: 20° , 38° and 49.9° (east of north). There are more regions with negative flux than at $-2.5 \times \sigma$, but most of the positive $2 \times \sigma$ emission is recovered around the tidal tail and dusty bridges connecting W2246–0526 with the companion galaxies. (B) Same as Figure 1 but with a uv -tapering of $1''$ applied to the data. The tidal tail is detected at larger angular scales. These two maps confirm the extended emission detected in Figure 1.

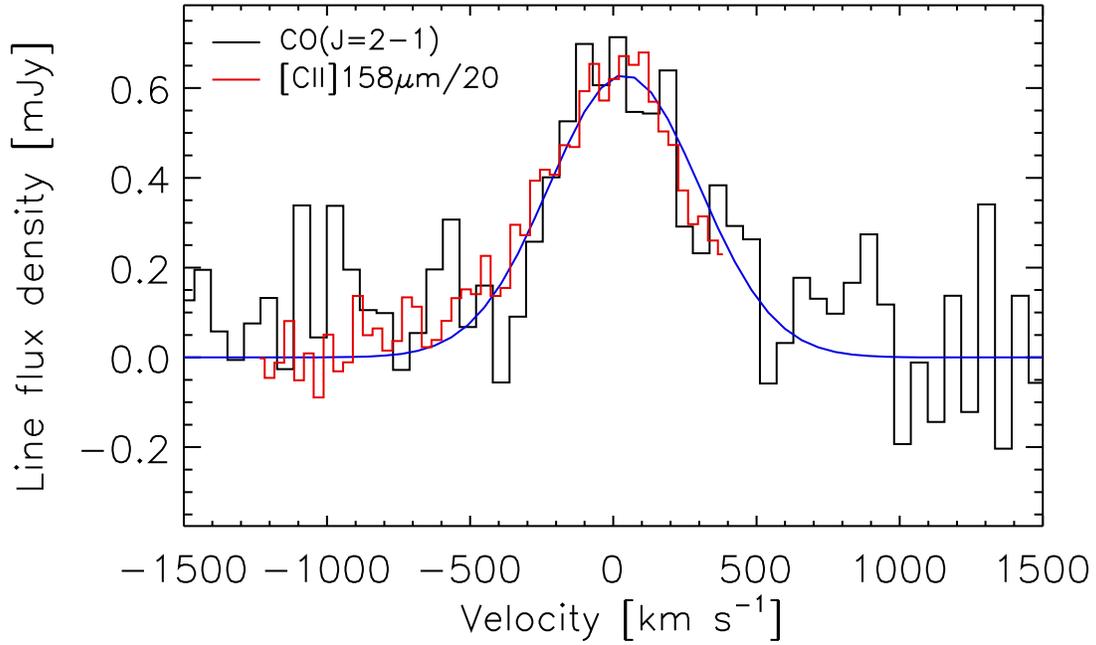


Figure S3. **The W2246–0526 spectra of the CO(2→1) (black) and [C II] 158 μm (red) emission lines.** The spectra were extracted using circular apertures of $r = 2.2''$ (similar to the beam size of the VLA observations) and $r = 0.35''$ (similar to the beam size of the ALMA observations), respectively, centered on W2246–0526. The [C II] spectrum has been scaled down in flux density by a factor of 20. The blue line is a Gaussian model fitted to the CO(2→1) spectrum, which results in a FWHM = 615 km s^{-1} . The FWHM of the [C II] line is $\sim 550 \text{ km s}^{-1}$ (16) within the central few kpc of W2246–0526. The cold molecular gas traced by the CO(2→1) line is as turbulent as the neutral gas traced by [C II], and may be also affected by the strong feedback from the central AGN.

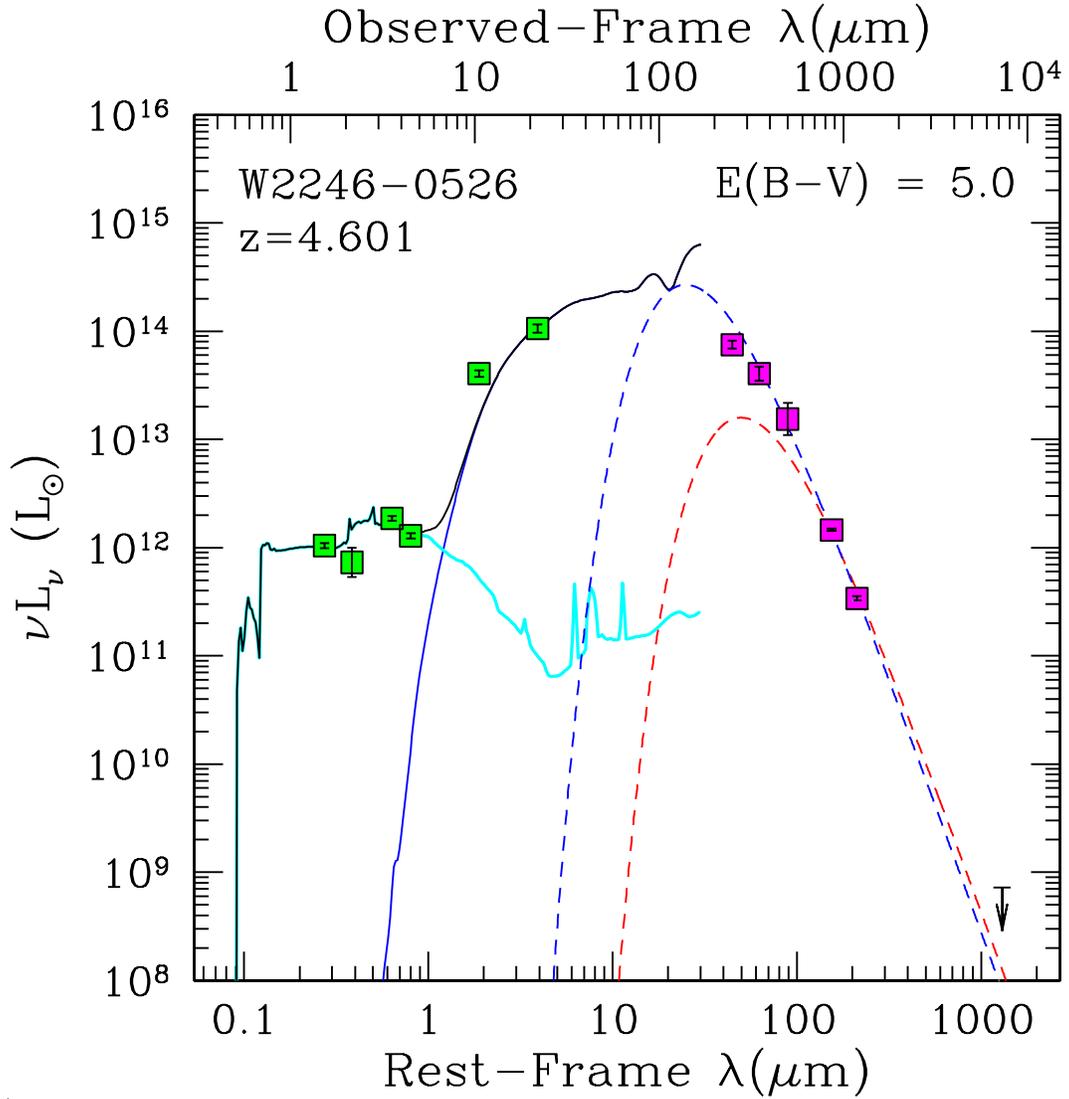


Figure S4. **Optical through far-IR SED of W2246-0526.** The mono-chromatic luminosity (νL_ν , where L_ν is the luminosity per unit frequency) is plotted as a function of the rest-frame wavelength. The green squares show, from left to right, the observed luminosities in the F160W HST (*I*6), Ks (*I*0), Spitzer 3.6 and 4.5 μ m (*I*9), and WISE W3 and W4 broad-band filters (*I*3). The black line shows the best-fitting SED using the templates and algorithm of (*I*47). The best fitting SED model is composed of an AGN template with reddening $E(B-V) = 5.0$ (dark blue solid line) and the Im galaxy template (cyan line). The magenta squares show, from left to right, the luminosities of SPIRE at 250, 350 and 500 μ m (*I*14), ALMA at 865 μ m and 1.19 mm, and the 3σ upper limit from the VLA at 7.28 mm. Two fits to the ALMA data-points are also shown, extrapolated to shorter and longer wavelengths, using an optically-thin modified black body with an emissivity index $\beta = 1.8$ and dust temperatures of 50 and 100 K (dashed red and blue lines respectively). These correspond to the two limiting temperatures used to calculate the dust mass in W2246-0526.

Table S1. Summary of the dust and gas mass estimates for each source in the W2246–0526 system, assuming different dust temperatures. Column 1: Source over which the photometry was performed. Column 2: Radius of the circular aperture used to obtain the flux densities. Column 3: Observed flux density of the dust continuum at $\lambda_{\text{rest}} \sim 212 \mu\text{m}$. Columns 4–6: Dust masses estimated using three different temperatures ($T_{\text{dust}} = 100, 50$ and 25 K), assuming optically thin modified black body emission with an emissivity $\beta = 1.8$, and a dust opacity coefficient $\kappa_{\nu, \text{dust}}(850 \mu\text{m}) = 0.0484 \text{ m}^2 \text{ kg}^{-1}$. Assuming that the dust and gas are well mixed, a standard gas-to-dust ratio typical of local, solar-metallicity galaxies (18) of $\delta_{\text{GDR}} = 100$ can be applied (multiplied) to obtain the gas mass of each component. Column 7: Gas mass estimates based on the CO(2→1) emission line assuming the gas is thermalized up to the $J = 2-1$ transition and using a $M_{\text{gas}}-L_{\text{CO}(1\rightarrow 0)}$ ratio typical of local ULIRGS $\alpha_{\text{CO, ULIRG}} = 0.8 M_{\odot} [\text{K km s}^{-1} \text{ pc}^2]^{-1}$. If a ratio typical of normal star-forming galaxies were used, all values in Column 7 would be multiplied by $\sim 5-6$. The radius of the aperture used to calculate the total gas mass based on the CO(2→1) line is $r = 5.5''$.

Source	Aper. radius ["]	f_{ν} (212 μm) [mJy]	M_{dust} (T=100K) [M_{\odot}]	M_{dust} (T=50K) [M_{\odot}]	M_{dust} (T=25K) [M_{\odot}]	M_{gas} CO(2-1) $\alpha_{\text{CO, ULIRG}}$
W2246	0.5	2.3 ± 0.1	$5.6 - 17 \times 10^8$		–	–
C1	0.5	0.27 ± 0.02	–	$2.0 - 9.6 \times 10^8$		–
C2	0.5	0.23 ± 0.02	–	$1.7 - 8.2 \times 10^8$		–
C3	0.5	0.16 ± 0.02	–	$1.2 - 5.7 \times 10^8$		–
K1	0.5	0.12 ± 0.02	–	$0.9 - 4.3 \times 10^8$		–
Tail	1.6	0.63 ± 0.05	–	$4.6 - 22 \times 10^8$		–
CO beam	2.2	3.9 ± 0.4	$9.5 - 28 \times 10^8$			$0.7 \times 10^{11} M_{\odot}$
Total	4.5	4.9 ± 0.6	$12 - 36 \times 10^8$			$1.5 \times 10^{11} M_{\odot}$

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