

A *Fermi* GAMMA-RAY BURST MONITOR SEARCH FOR ELECTROMAGNETIC SIGNALS COINCIDENT  
WITH GRAVITATIONAL-WAVE CANDIDATES IN ADVANCED LIGO'S FIRST OBSERVING RUN

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A. VON KIENLIN<sup>11</sup>

*Fermi* GAMMA-RAY BURST MONITOR

B. P. ABBOTT,<sup>12</sup> R. ABBOTT,<sup>12</sup> T. D. ABBOTT,<sup>13</sup> F. ACERNESE,<sup>14,15</sup> K. ACKLEY,<sup>16</sup> C. ADAMS,<sup>17</sup> T. ADAMS,<sup>18</sup>  
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S. BISCOVEANU,<sup>16</sup> A. BISHT,<sup>20,21</sup> M. BITOSSO,<sup>40,32</sup> M. A. BIZOUARD,<sup>37</sup> J. K. BLACKBURN,<sup>12</sup> J. BLACKMAN,<sup>57</sup> C. D. BLAIR,<sup>17</sup>  
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M. J. HART,<sup>55</sup> C.-J. HASTER,<sup>124</sup> K. HAUGHIAN,<sup>55</sup> J. HEALY,<sup>69</sup> A. HEIDMANN,<sup>81</sup> M. C. HEINTZE,<sup>17</sup> H. HEITMANN,<sup>76</sup>  
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D. HUET,<sup>37</sup> B. HUGHEY,<sup>47</sup> M. HULKO,<sup>12</sup> S. HUSA,<sup>114</sup> S. H. HUTTNER,<sup>55</sup> T. HUYNH-DINH,<sup>17</sup> A. IESS,<sup>42,43</sup> N. INDIK,<sup>20</sup>  
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K. JANI,<sup>87</sup> P. JARANOWSKI,<sup>135</sup> D. S. JOHNSON,<sup>22</sup> W. W. JOHNSON,<sup>13</sup> D. I. JONES,<sup>136</sup> R. JONES,<sup>55</sup> R. J. G. JONKER,<sup>24</sup>  
L. JU,<sup>74</sup> J. JUNKER,<sup>20,21</sup> C. V. KALAGHATGI,<sup>46</sup> V. KALOGERA,<sup>100</sup> B. KAMAI,<sup>12</sup> S. KANDHASAMY,<sup>17</sup> G. KANG,<sup>49</sup>  
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E. KATSAVOUNIDIS,<sup>5</sup> W. KATZMAN,<sup>17</sup> S. KAUFER,<sup>20,21</sup> K. KAWABE,<sup>56</sup> N. V. KEERTHANA,<sup>28</sup> F. KÉFÉLIAN,<sup>76</sup> D. KEITEL,<sup>55</sup>  
A. J. KEMBALL,<sup>22</sup> R. KENNEDY,<sup>119</sup> J. S. KEY,<sup>137</sup> F. Y. KHALILI,<sup>73</sup> B. KHAMESRA,<sup>87</sup> H. KHAN,<sup>38</sup> I. KHAN,<sup>26,43</sup> S. KHAN,<sup>20</sup>  
Z. KHAN,<sup>118</sup> E. A. KHAZANOV,<sup>138</sup> N. KIJBUNCHOO,<sup>34</sup> CHUNGLEE KIM,<sup>139</sup> J. C. KIM,<sup>140</sup> K. KIM,<sup>104</sup> W. KIM,<sup>66</sup> W. S. KIM,<sup>141</sup>  
Y.-M. KIM,<sup>142</sup> E. J. KING,<sup>66</sup> P. J. KING,<sup>56</sup> M. KINLEY-HANLON,<sup>134</sup> R. KIRCHHOFF,<sup>20,21</sup> J. S. KISSEL,<sup>56</sup> L. KLEYBOLTE,<sup>44</sup>  
S. KLIMENKO,<sup>59</sup> T. D. KNOWLES,<sup>50</sup> P. KOCH,<sup>20,21</sup> S. M. KOEHLLENBECK,<sup>20,21</sup> S. KOLEY,<sup>24</sup> V. KONDRASHOV,<sup>12</sup> A. KONTOS,<sup>5</sup>  
M. KOROBKO,<sup>44</sup> W. Z. KORTH,<sup>12</sup> I. KOWALSKA,<sup>84</sup> D. B. KOZAK,<sup>12</sup> C. KRÄMER,<sup>20</sup> V. KRINGEL,<sup>20,21</sup> B. KRISHNAN,<sup>20</sup>  
A. KRÓLAK,<sup>143,144</sup> G. KUEHN,<sup>20,21</sup> P. KUMAR,<sup>145</sup> R. KUMAR,<sup>118</sup> S. KUMAR,<sup>29</sup> L. KUO,<sup>98</sup> A. KUTYNIA,<sup>143</sup> S. KWANG,<sup>30</sup>  
B. D. LACKEY,<sup>48</sup> K. H. LAI,<sup>104</sup> M. LANDRY,<sup>56</sup> R. N. LANG,<sup>146</sup> J. LANGE,<sup>69</sup> B. LANTZ,<sup>60</sup> R. K. LANZA,<sup>5</sup>  
A. LARTAUX-VOLLARD,<sup>37</sup> P. D. LASKY,<sup>16</sup> M. LAXEN,<sup>17</sup> A. LAZZARINI,<sup>12</sup> C. LAZZARO,<sup>63</sup> P. LEACI,<sup>108,45</sup> S. LEAVEY,<sup>20,21</sup>  
C. H. LEE,<sup>102</sup> H. K. LEE,<sup>147</sup> H. M. LEE,<sup>139</sup> H. W. LEE,<sup>140</sup> K. LEE,<sup>55</sup> J. LEHMANN,<sup>20,21</sup> A. LENON,<sup>50</sup> M. LEONARDI,<sup>20,21,123</sup>  
N. LEROY,<sup>37</sup> N. LETENDRE,<sup>18</sup> Y. LEVIN,<sup>16</sup> J. LI,<sup>92</sup> T. G. F. LI,<sup>104</sup> X. LI,<sup>57</sup> S. D. LINKER,<sup>120</sup> T. B. LITTENBERG,<sup>148</sup>  
J. LIU,<sup>74</sup> X. LIU,<sup>30</sup> R. K. L. LO,<sup>104</sup> N. A. LOCKERBIE,<sup>36</sup> L. T. LONDON,<sup>46</sup> A. LONGO,<sup>149,150</sup> M. LORENZINI,<sup>26,27</sup>  
V. LORLETTE,<sup>151</sup> M. LORMAND,<sup>17</sup> G. LOSURDO,<sup>32</sup> J. D. LOUGH,<sup>20,21</sup> C. O. LOUSTO,<sup>69</sup> G. LOVELACE,<sup>38</sup> H. LÜCK,<sup>20,21</sup>  
D. LUMACA,<sup>42,43</sup> A. P. LUNDGREN,<sup>20</sup> R. LYNCH,<sup>5</sup> Y. MA,<sup>57</sup> R. MACAS,<sup>46</sup> S. MACFOY,<sup>36</sup> B. MACHENSCHALK,<sup>20</sup>  
M. MACINNIS,<sup>5</sup> D. M. MACLEOD,<sup>46</sup> I. MAGAÑA HERNANDEZ,<sup>30</sup> F. MAGAÑA-SANDOVAL,<sup>53</sup> L. MAGAÑA ZERTUCHE,<sup>95</sup>  
R. M. MAGEE,<sup>97</sup> E. MAJORANA,<sup>45</sup> I. MAKSIMOVIC,<sup>151</sup> N. MAN,<sup>76</sup> V. MANDIC,<sup>54</sup> V. MANGANO,<sup>55</sup> G. L. MANSSELL,<sup>34</sup>  
M. MANSKE,<sup>30,34</sup> M. MANTOVANI,<sup>40</sup> F. MARCHESONI,<sup>61,52</sup> F. MARION,<sup>18</sup> S. MÁRKA,<sup>112</sup> Z. MÁRKA,<sup>112</sup> C. MARKAKIS,<sup>22</sup>  
A. S. MARKOSYAN,<sup>60</sup> A. MARKOWITZ,<sup>12</sup> E. MAROS,<sup>12</sup> A. MARQUINA,<sup>111</sup> F. MARTELLI,<sup>82,83</sup> L. MARTELLINI,<sup>76</sup>  
I. W. MARTIN,<sup>55</sup> R. M. MARTIN,<sup>122</sup> D. V. MARTYNOV,<sup>5</sup> K. MASON,<sup>5</sup> E. MASSERA,<sup>119</sup> A. MASSEROT,<sup>18</sup> T. J. MASSINGER,<sup>12</sup>  
M. MASSO-REID,<sup>55</sup> S. MASTROGIOVANNI,<sup>108,45</sup> A. MATAS,<sup>54</sup> F. MATICHARD,<sup>12,5</sup> L. MATONE,<sup>112</sup> N. MAVALVALA,<sup>5</sup>  
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S. C. MCGUIRE,<sup>152</sup> J. MCIVER,<sup>12</sup> D. J. MCMANUS,<sup>34</sup> T. MCRAE,<sup>34</sup> S. T. MCWILLIAMS,<sup>50</sup> D. MEACHER,<sup>97</sup>  
G. D. MEADORS,<sup>16</sup> M. MEHMET,<sup>20,21</sup> J. MEIDAM,<sup>24</sup> E. MEJUTO-VILLA,<sup>19</sup> A. MELATOS,<sup>107</sup> G. MENDELL,<sup>56</sup>  
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C. MESSENGER,<sup>55</sup> C. MESSICK,<sup>97</sup> R. METZDORFF,<sup>81</sup> P. M. MEYERS,<sup>54</sup> H. MIAO,<sup>70</sup> C. MICHEL,<sup>35</sup> H. MIDDLETON,<sup>107</sup>  
E. E. MIKHAILOV,<sup>153</sup> L. MILANO,<sup>89,15</sup> A. L. MILLER,<sup>59</sup> A. MILLER,<sup>108,45</sup> B. B. MILLER,<sup>100</sup> J. MILLER,<sup>5</sup> M. MILLHOUSE,<sup>113</sup>  
J. MILLS,<sup>46</sup> M. C. MILOVICH-GOFF,<sup>120</sup> O. MINAZZOLI,<sup>76,154</sup> Y. MINENKOV,<sup>43</sup> J. MING,<sup>20,21</sup> C. MISHRA,<sup>155</sup> S. MITRA,<sup>28</sup>  
V. P. MITROFANOV,<sup>73</sup> G. MITSSELMACHER,<sup>59</sup> R. MITTLEMAN,<sup>5</sup> D. MOFFA,<sup>94</sup> K. MOGUSHI,<sup>95</sup> M. MOHAN,<sup>40</sup>  
S. R. P. MOHAPATRA,<sup>5</sup> M. MONTANI,<sup>82,83</sup> C. J. MOORE,<sup>23</sup> D. MORARU,<sup>56</sup> G. MORENO,<sup>56</sup> S. MORISAKI,<sup>91</sup> B. MOURS,<sup>18</sup>  
C. M. MOW-LOWRY,<sup>70</sup> G. MUELLER,<sup>59</sup> A. W. MUIR,<sup>46</sup> ARUNAVA MUKHERJEE,<sup>20,21</sup> D. MUKHERJEE,<sup>30</sup> S. MUKHERJEE,<sup>116</sup>  
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- K. NAPIER,<sup>87</sup> I. NARDECCHIA,<sup>42,43</sup> L. NATICCHIONI,<sup>108,45</sup> R. K. NAYAK,<sup>158</sup> J. NEILSON,<sup>120</sup> G. NELEMANS,<sup>75,24</sup>  
T. J. N. NELSON,<sup>17</sup> M. NERY,<sup>20,21</sup> A. NEUNZERT,<sup>132</sup> L. NEVIN,<sup>12</sup> J. M. NEWPORT,<sup>134</sup> K. Y. NG,<sup>5</sup> S. NG,<sup>66</sup> P. NGUYEN,<sup>80</sup>  
T. T. NGUYEN,<sup>34</sup> D. NICHOLS,<sup>75</sup> A. B. NIELSEN,<sup>20</sup> S. NISSANKE,<sup>75,24</sup> A. NITZ,<sup>20</sup> F. NOCERA,<sup>40</sup> D. NOLTING,<sup>17</sup> C. NORTH,<sup>46</sup>  
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S. H. OH,<sup>141</sup> F. OHME,<sup>20</sup> H. OHTA,<sup>91</sup> M. A. OKADA,<sup>25</sup> M. OLIVER,<sup>114</sup> P. OPPERMANN,<sup>20,21</sup> RICHARD J. ORAM,<sup>17</sup>  
B. O'REILLY,<sup>17</sup> R. ORMISTON,<sup>54</sup> L. F. ORTEGA,<sup>59</sup> R. O'SHAUGHNESSY,<sup>69</sup> S. OSSOKINE,<sup>48</sup> D. J. OTTAWAY,<sup>66</sup> H. OVERMIER,<sup>17</sup>  
B. J. OWEN,<sup>93</sup> A. E. PACE,<sup>97</sup> G. PAGANO,<sup>31,32</sup> J. PAGE,<sup>148</sup> M. A. PAGE,<sup>74</sup> A. PAI,<sup>128</sup> S. A. PAI,<sup>72</sup> J. R. PALAMOS,<sup>80</sup>  
O. PALASHOV,<sup>138</sup> C. PALOMBA,<sup>45</sup> A. PAL-SINGH,<sup>44</sup> HOWARD PAN,<sup>98</sup> HUANG-WEI PAN,<sup>98</sup> B. PANG,<sup>57</sup> P. T. H. PANG,<sup>104</sup>  
C. PANKOW,<sup>100</sup> F. PANNARALE,<sup>46</sup> B. C. PANT,<sup>72</sup> F. PAOLETTI,<sup>32</sup> A. PAOLI,<sup>40</sup> M. A. PAPA,<sup>20,30,21</sup> A. PARIDA,<sup>28</sup>  
W. PARKER,<sup>17</sup> D. PASCUCCI,<sup>55</sup> A. PASQUALETTI,<sup>40</sup> R. PASSAQUIETI,<sup>31,32</sup> D. PASSUELLO,<sup>32</sup> M. PATIL,<sup>144</sup> B. PATRICELLI,<sup>160,32</sup>  
B. L. PEARLSTONE,<sup>55</sup> C. PEDERSEN,<sup>46</sup> M. PEDRAZA,<sup>12</sup> R. PEDURAND,<sup>35,161</sup> L. PEKOWSKY,<sup>53</sup> A. PELE,<sup>17</sup> S. PENN,<sup>162</sup>  
C. J. PEREZ,<sup>56</sup> A. PERRECA,<sup>121,106</sup> L. M. PERRI,<sup>100</sup> H. P. PFEIFFER,<sup>124,48</sup> M. PHELPS,<sup>55</sup> K. S. PHUKON,<sup>28</sup>  
O. J. PICCINI,<sup>108,45</sup> M. PICHOT,<sup>76</sup> F. PIERGIOVANNI,<sup>82,83</sup> V. PIERRO,<sup>19</sup> G. PILLANT,<sup>40</sup> L. PINARD,<sup>35</sup> I. M. PINTO,<sup>19</sup>  
M. PIRELLO,<sup>56</sup> M. PITKIN,<sup>55</sup> R. POGGIANI,<sup>31,32</sup> P. POPOLIZIO,<sup>40</sup> E. K. PORTER,<sup>39</sup> L. POSSENTI,<sup>163,83</sup> A. POST,<sup>20</sup>  
J. POWELL,<sup>164</sup> J. PRASAD,<sup>28</sup> J. W. W. PRATT,<sup>47</sup> G. PRATTEN,<sup>114</sup> V. PREDOI,<sup>46</sup> T. PRESTEGARD,<sup>30</sup> M. PRINCIPE,<sup>19</sup>  
S. PRIVITERA,<sup>48</sup> G. A. PRODI,<sup>121,106</sup> L. G. PROKHOROV,<sup>73</sup> O. PUNCKEN,<sup>20,21</sup> M. PUNTURO,<sup>52</sup> P. PUPPO,<sup>45</sup> M. PÜRRER,<sup>48</sup>  
H. QI,<sup>30</sup> V. QUETSCHKE,<sup>116</sup> E. A. QUINTERO,<sup>12</sup> R. QUITZOW-JAMES,<sup>80</sup> F. J. RAAB,<sup>56</sup> D. S. RABELING,<sup>34</sup> H. RADKINS,<sup>56</sup>  
P. RAFFAI,<sup>64</sup> S. RAJA,<sup>72</sup> C. RAJAN,<sup>72</sup> B. RAJBHANDARI,<sup>93</sup> M. RAKHMANOV,<sup>116</sup> K. E. RAMIREZ,<sup>116</sup> A. RAMOS-BUADES,<sup>114</sup>  
JAVED RANA,<sup>28</sup> P. RAPAGNANI,<sup>108,45</sup> V. RAYMOND,<sup>46</sup> M. RAZZANO,<sup>31,32</sup> J. READ,<sup>38</sup> T. REGIMBAU,<sup>76,18</sup> L. REI,<sup>71</sup> S. REID,<sup>36</sup>  
D. H. REITZE,<sup>12,59</sup> W. REN,<sup>22</sup> F. RICCI,<sup>108,45</sup> P. M. RICKER,<sup>22</sup> K. RILES,<sup>132</sup> M. RIZZO,<sup>69</sup> N. A. ROBERTSON,<sup>12,55</sup>  
R. ROBIE,<sup>55</sup> F. ROBINET,<sup>37</sup> T. ROBSON,<sup>113</sup> A. ROCCHI,<sup>43</sup> L. ROLLAND,<sup>18</sup> J. G. ROLLINS,<sup>12</sup> V. J. ROMA,<sup>80</sup> R. ROMANO,<sup>14,15</sup>  
C. L. ROMEL,<sup>56</sup> J. H. ROMIE,<sup>17</sup> D. ROSIŃSKA,<sup>165,65</sup> M. P. ROSS,<sup>166</sup> S. ROWAN,<sup>55</sup> A. RÜDIGER,<sup>20,21</sup> P. RUGGI,<sup>40</sup>  
G. RUTINS,<sup>167</sup> K. RYAN,<sup>56</sup> S. SACHDEV,<sup>12</sup> T. SADECKI,<sup>56</sup> M. SAKELLARIADOU,<sup>168</sup> L. SALCONI,<sup>40</sup> M. SALEEM,<sup>128</sup> F. SALEMI,<sup>20</sup>  
A. SAMAJDAR,<sup>158,24</sup> L. SAMMUT,<sup>16</sup> L. M. SAMPSON,<sup>100</sup> E. J. SANCHEZ,<sup>12</sup> L. E. SANCHEZ,<sup>12</sup> N. SANCHIS-GUAL,<sup>33</sup>  
V. SANDBERG,<sup>56</sup> J. R. SANDERS,<sup>53</sup> N. SARIN,<sup>16</sup> B. SASSOLAS,<sup>35</sup> B. S. SATHYAPRAKASH,<sup>97,46</sup> P. R. SAULSON,<sup>53</sup>  
O. SAUTER,<sup>132</sup> R. L. SAVAGE,<sup>56</sup> A. SAWADSKY,<sup>44</sup> P. SCHALE,<sup>80</sup> M. SCHEEL,<sup>57</sup> J. SCHEUER,<sup>100</sup> P. SCHMIDT,<sup>75</sup>  
R. SCHNABEL,<sup>44</sup> R. M. S. SCHOFIELD,<sup>80</sup> A. SCHÖNBECK,<sup>44</sup> E. SCHREIBER,<sup>20,21</sup> D. SCHUETTE,<sup>20,21</sup> B. W. SCHULTE,<sup>20,21</sup>  
B. F. SCHUTZ,<sup>46,20</sup> S. G. SCHWALBE,<sup>47</sup> J. SCOTT,<sup>55</sup> S. M. SCOTT,<sup>34</sup> E. SEIDEL,<sup>22</sup> D. SELLERS,<sup>17</sup> A. S. SENGUPTA,<sup>169</sup>  
D. SENTENAC,<sup>40</sup> V. SEQUINO,<sup>42,43,26</sup> A. SERGEEV,<sup>138</sup> Y. SETYAWATI,<sup>20</sup> D. A. SHADDOCK,<sup>34</sup> T. J. SHAFFER,<sup>56</sup> A. A. SHAH,<sup>148</sup>  
M. S. SHAHRIAR,<sup>100</sup> M. B. SHANER,<sup>120</sup> L. SHAO,<sup>48</sup> B. SHAPIRO,<sup>60</sup> P. SHAWHAN,<sup>86</sup> H. SHEN,<sup>22</sup> D. H. SHOEMAKER,<sup>5</sup>  
D. M. SHOEMAKER,<sup>87</sup> K. SIELLEZ,<sup>87</sup> X. SIEMENS,<sup>30</sup> M. SIENIAWSKA,<sup>65</sup> D. SIGG,<sup>56</sup> A. D. SILVA,<sup>25</sup> L. P. SINGER,<sup>1</sup>  
A. SINGH,<sup>20,21</sup> A. SINGHAL,<sup>26,45</sup> A. M. SINTES,<sup>114</sup> B. J. J. SLAGMOLEN,<sup>34</sup> T. J. SLAVEN-BLAIR,<sup>74</sup> B. SMITH,<sup>17</sup>  
J. R. SMITH,<sup>38</sup> R. J. E. SMITH,<sup>16</sup> S. SOMALA,<sup>170</sup> E. J. SON,<sup>141</sup> B. SORAZU,<sup>55</sup> F. SORRENTINO,<sup>71</sup> T. SOURADEEP,<sup>28</sup>  
A. P. SPENCER,<sup>55</sup> A. K. SRIVASTAVA,<sup>118</sup> K. STAATS,<sup>47</sup> M. STEINKE,<sup>20,21</sup> J. STEINLECHNER,<sup>44,55</sup> S. STEINLECHNER,<sup>44</sup>  
D. STEINMEYER,<sup>20,21</sup> B. STELTNER,<sup>20,21</sup> S. P. STEVENSON,<sup>164</sup> D. STOCKS,<sup>60</sup> R. STONE,<sup>116</sup> D. J. STOPS,<sup>70</sup> K. A. STRAIN,<sup>55</sup>  
G. STRATTA,<sup>82,83</sup> S. E. STRIGIN,<sup>73</sup> A. STRUNK,<sup>56</sup> R. STURANI,<sup>171</sup> A. L. STUVER,<sup>172</sup> T. Z. SUMMERSCALES,<sup>173</sup> L. SUN,<sup>107</sup>  
S. SUNIL,<sup>118</sup> J. SURESH,<sup>28</sup> P. J. SUTTON,<sup>46</sup> B. L. SWINKELS,<sup>24</sup> M. J. SZCZEPAŃCZYK,<sup>47</sup> M. TACCA,<sup>24</sup> S. C. TAIT,<sup>55</sup>  
C. TALBOT,<sup>16</sup> D. TALUKDER,<sup>80</sup> D. B. TANNER,<sup>59</sup> M. TÁPAI,<sup>129</sup> A. TARACCHINI,<sup>48</sup> J. D. TASSON,<sup>103</sup> J. A. TAYLOR,<sup>148</sup>  
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K. A. THORNE,<sup>17</sup> E. THRANE,<sup>16</sup> S. TIWARI,<sup>26,106</sup> V. TIWARI,<sup>46</sup> K. V. TOKMAKOV,<sup>36</sup> K. TOLAND,<sup>55</sup> M. TONELLI,<sup>31,32</sup>  
Z. TORNASI,<sup>55</sup> A. TORRES-FORNÉ,<sup>33</sup> C. I. TORRIE,<sup>12</sup> D. TÖYRÄ,<sup>70</sup> F. TRAVASSO,<sup>40,52</sup> G. TRAYLOR,<sup>17</sup> J. TRINASTIC,<sup>59</sup>  
M. C. TRINGALI,<sup>121,106</sup> L. TROZZO,<sup>174,32</sup> K. W. TSANG,<sup>24</sup> M. TSE,<sup>5</sup> R. TSO,<sup>57</sup> D. TSUNA,<sup>91</sup> L. TSUKADA,<sup>91</sup>  
D. TUYENBAYEV,<sup>116</sup> K. UENO,<sup>30</sup> D. UGOLINI,<sup>175</sup> A. L. URBAN,<sup>12</sup> S. A. USMAN,<sup>46</sup> H. VAHLBRUCH,<sup>20,21</sup> G. VAJENTE,<sup>12</sup>  
G. VALDES,<sup>13</sup> N. VAN BAKEL,<sup>24</sup> M. VAN BEUZekom,<sup>24</sup> J. F. J. VAN DEN BRAND,<sup>85,24</sup> C. VAN DEN BROECK,<sup>24,176</sup>  
D. C. VANDER-HYDE,<sup>53</sup> L. VAN DER SCHAAF,<sup>24</sup> J. V. VAN HEIJNINGEN,<sup>24</sup> A. A. VAN VEGGEL,<sup>55</sup> M. VARDARO,<sup>62,63</sup>  
V. VARMA,<sup>57</sup> S. VASS,<sup>12</sup> M. VASÚTH,<sup>58</sup> A. VECCHIO,<sup>70</sup> G. VEDOVATO,<sup>63</sup> J. VEITCH,<sup>55</sup> P. J. VEITCH,<sup>66</sup>  
K. VENKATESWARA,<sup>166</sup> G. VENUGOPALAN,<sup>12</sup> D. VERKINDT,<sup>18</sup> F. VETRANO,<sup>82,83</sup> A. VICERÉ,<sup>82,83</sup> A. D. VIETS,<sup>30</sup>  
S. VINCIGUERRA,<sup>70</sup> D. J. VINE,<sup>167</sup> J.-Y. VINET,<sup>76</sup> S. VITALE,<sup>5</sup> T. VO,<sup>53</sup> H. VOCCA,<sup>51,52</sup> C. VORVICK,<sup>56</sup>  
S. P. VYATCHANIN,<sup>73</sup> A. R. WADE,<sup>12</sup> L. E. WADE,<sup>94</sup> M. WADE,<sup>94</sup> R. WALET,<sup>24</sup> M. WALKER,<sup>38</sup> L. WALLACE,<sup>12</sup>  
S. WALSH,<sup>30,20</sup> G. WANG,<sup>26,32</sup> H. WANG,<sup>70</sup> J. Z. WANG,<sup>132</sup> W. H. WANG,<sup>116</sup> Y. F. WANG,<sup>104</sup> R. L. WARD,<sup>34</sup> J. WARNER,<sup>56</sup>  
M. WAS,<sup>18</sup> J. WATCHI,<sup>109</sup> B. WEAVER,<sup>56</sup> L.-W. WEI,<sup>20,21</sup> M. WEINERT,<sup>20,21</sup> A. J. WEINSTEIN,<sup>12</sup> R. WEISS,<sup>5</sup>  
F. WELLMANN,<sup>20,21</sup> L. WEN,<sup>74</sup> E. K. WESSEL,<sup>22</sup> P. WESSELS,<sup>20,21</sup> J. WESTERWECK,<sup>20</sup> K. WETTE,<sup>34</sup> J. T. WHELAN,<sup>69</sup>  
B. F. WHITING,<sup>59</sup> C. WHITTLE,<sup>5</sup> D. WILKEN,<sup>20,21</sup> D. WILLIAMS,<sup>55</sup> R. D. WILLIAMS,<sup>12</sup> A. R. WILLIAMSON,<sup>69,75</sup>  
J. L. WILLIS,<sup>12,133</sup> B. WILLKE,<sup>20,21</sup> M. H. WIMMER,<sup>20,21</sup> W. WINKLER,<sup>20,21</sup> C. C. WIPF,<sup>12</sup> H. WITTEL,<sup>20,21</sup> G. WOAN,<sup>55</sup>  
J. WOELER,<sup>20,21</sup> J. K. WOFFORD,<sup>69</sup> W. K. WONG,<sup>104</sup> J. WORDEN,<sup>56</sup> J. L. WRIGHT,<sup>55</sup> D. S. WU,<sup>20,21</sup> D. M. WYSOCKI,<sup>69</sup>  
S. XIAO,<sup>12</sup> W. YAM,<sup>5</sup> H. YAMAMOTO,<sup>12</sup> C. C. YANCEY,<sup>86</sup> L. YANG,<sup>177</sup> M. J. YAP,<sup>34</sup> M. YAZBACK,<sup>59</sup> HANG YU,<sup>5</sup>  
HAOCUN YU,<sup>5</sup> M. YVERT,<sup>18</sup> A. ZADROŻNY,<sup>143</sup> M. ZANOLIN,<sup>47</sup> T. ZELENKOVA,<sup>40</sup> J.-P. ZENDRI,<sup>63</sup> M. ZEVIN,<sup>100</sup> J. ZHANG,<sup>74</sup>  
L. ZHANG,<sup>12</sup> M. ZHANG,<sup>153</sup> T. ZHANG,<sup>55</sup> Y.-H. ZHANG,<sup>20,21</sup> C. ZHAO,<sup>74</sup> M. ZHOU,<sup>100</sup> Z. ZHOU,<sup>100</sup> S. J. ZHU,<sup>20,21</sup>  
X. J. ZHU,<sup>16</sup> M. E. ZUCKER,<sup>12,5</sup> AND J. ZWEIZIG<sup>12</sup>

## THE LIGO SCIENTIFIC COLLABORATION AND THE VIRGO COLLABORATION

- <sup>1</sup>*NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA*
- <sup>2</sup>*Science and Technology Institute, Universities Space Research Association, Huntsville, AL 35805, USA*
- <sup>3</sup>*NASA Marshall Space Flight Center, Huntsville, AL 35812, USA*
- <sup>4</sup>*Harvard-Smithsonian Center for Astrophysics, 60 Garden St, Cambridge, MA 02138, USA*
- <sup>5</sup>*LIGO, Massachusetts Institute of Technology, Cambridge, MA 02139, USA*
- <sup>6</sup>*Department of Space Science, University of Alabama in Huntsville, Huntsville, AL 35899, USA*
- <sup>7</sup>*Center for Space Plasma and Aeronomic Research, University of Alabama in Huntsville, Huntsville, AL 35899, USA*
- <sup>8</sup>*Politecnico di Bari, Via Edoardo Orabona, 4, I-70126 Bari BA, Italy*
- <sup>9</sup>*Istituto Nazionale di Fisica Nucleare, Sezione di Bari, I-70126 Bari, Italy*
- <sup>10</sup>*Jacobs Space Exploration Group, Huntsville, AL 35806*
- <sup>11</sup>*Max-Planck-Institut für extraterrestrische Physik, Giessenbachstrasse 1, 85748 Garching, Germany*
- <sup>12</sup>*LIGO, California Institute of Technology, Pasadena, CA 91125, USA*
- <sup>13</sup>*Louisiana State University, Baton Rouge, LA 70803, USA*
- <sup>14</sup>*Università di Salerno, Fisciano, I-84084 Salerno, Italy*
- <sup>15</sup>*INFN, Sezione di Napoli, Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy*
- <sup>16</sup>*OzGrav, School of Physics & Astronomy, Monash University, Clayton 3800, Victoria, Australia*
- <sup>17</sup>*LIGO Livingston Observatory, Livingston, LA 70754, USA*
- <sup>18</sup>*Laboratoire d'Annecy de Physique des Particules (LAPP), Univ. Grenoble Alpes, Université Savoie Mont Blanc, CNRS/IN2P3, F-74941 Annecy, France*
- <sup>19</sup>*University of Sannio at Benevento, I-82100 Benevento, Italy and INFN, Sezione di Napoli, I-80100 Napoli, Italy*
- <sup>20</sup>*Max Planck Institute for Gravitational Physics (Albert Einstein Institute), D-30167 Hannover, Germany*
- <sup>21</sup>*Leibniz Universität Hannover, D-30167 Hannover, Germany*
- <sup>22</sup>*NCSA, University of Illinois at Urbana-Champaign, Urbana, IL 61801, USA*
- <sup>23</sup>*University of Cambridge, Cambridge CB2 1TN, United Kingdom*
- <sup>24</sup>*Nikhef, Science Park 105, 1098 XG Amsterdam, The Netherlands*
- <sup>25</sup>*Instituto Nacional de Pesquisas Espaciais, 12227-010 São José dos Campos, São Paulo, Brazil*
- <sup>26</sup>*Gran Sasso Science Institute (GSSI), I-67100 L'Aquila, Italy*
- <sup>27</sup>*INFN, Laboratori Nazionali del Gran Sasso, I-67100 Assergi, Italy*
- <sup>28</sup>*Inter-University Centre for Astronomy and Astrophysics, Pune 411007, India*
- <sup>29</sup>*International Centre for Theoretical Sciences, Tata Institute of Fundamental Research, Bengaluru 560089, India*
- <sup>30</sup>*University of Wisconsin-Milwaukee, Milwaukee, WI 53201, USA*
- <sup>31</sup>*Università di Pisa, I-56127 Pisa, Italy*
- <sup>32</sup>*INFN, Sezione di Pisa, I-56127 Pisa, Italy*
- <sup>33</sup>*Departamento de Astronomía y Astrofísica, Universitat de València, E-46100 Burjassot, València, Spain*
- <sup>34</sup>*OzGrav, Australian National University, Canberra, Australian Capital Territory 0200, Australia*
- <sup>35</sup>*Laboratoire des Matériaux Avancés (LMA), CNRS/IN2P3, F-69622 Villeurbanne, France*
- <sup>36</sup>*SUPA, University of Strathclyde, Glasgow G1 1XQ, United Kingdom*
- <sup>37</sup>*LAL, Univ. Paris-Sud, CNRS/IN2P3, Université Paris-Saclay, F-91898 Orsay, France*
- <sup>38</sup>*California State University Fullerton, Fullerton, CA 92831, USA*
- <sup>39</sup>*APC, AstroParticule et Cosmologie, Université Paris Diderot, CNRS/IN2P3, CEA/Irfu, Observatoire de Paris, Sorbonne Paris Cité, F-75205 Paris Cedex 13, France*
- <sup>40</sup>*European Gravitational Observatory (EGO), I-56021 Cascina, Pisa, Italy*
- <sup>41</sup>*Chennai Mathematical Institute, Chennai 603103, India*
- <sup>42</sup>*Università di Roma Tor Vergata, I-00133 Roma, Italy*
- <sup>43</sup>*INFN, Sezione di Roma Tor Vergata, I-00133 Roma, Italy*
- <sup>44</sup>*Universität Hamburg, D-22761 Hamburg, Germany*
- <sup>45</sup>*INFN, Sezione di Roma, I-00185 Roma, Italy*
- <sup>46</sup>*Cardiff University, Cardiff CF24 3AA, United Kingdom*
- <sup>47</sup>*Embry-Riddle Aeronautical University, Prescott, AZ 86301, USA*
- <sup>48</sup>*Max Planck Institute for Gravitational Physics (Albert Einstein Institute), D-14476 Potsdam-Golm, Germany*

- <sup>49</sup> *Korea Institute of Science and Technology Information, Daejeon 34141, Korea*
- <sup>50</sup> *West Virginia University, Morgantown, WV 26506, USA*
- <sup>51</sup> *Università di Perugia, I-06123 Perugia, Italy*
- <sup>52</sup> *INFN, Sezione di Perugia, I-06123 Perugia, Italy*
- <sup>53</sup> *Syracuse University, Syracuse, NY 13244, USA*
- <sup>54</sup> *University of Minnesota, Minneapolis, MN 55455, USA*
- <sup>55</sup> *SUPA, University of Glasgow, Glasgow G12 8QQ, United Kingdom*
- <sup>56</sup> *LIGO Hanford Observatory, Richland, WA 99352, USA*
- <sup>57</sup> *Caltech CaRT, Pasadena, CA 91125, USA*
- <sup>58</sup> *Wigner RCP, RMKI, H-1121 Budapest, Konkoly Thege Miklós út 29-33, Hungary*
- <sup>59</sup> *University of Florida, Gainesville, FL 32611, USA*
- <sup>60</sup> *Stanford University, Stanford, CA 94305, USA*
- <sup>61</sup> *Università di Camerino, Dipartimento di Fisica, I-62032 Camerino, Italy*
- <sup>62</sup> *Università di Padova, Dipartimento di Fisica e Astronomia, I-35131 Padova, Italy*
- <sup>63</sup> *INFN, Sezione di Padova, I-35131 Padova, Italy*
- <sup>64</sup> *MTA-ELTE Astrophysics Research Group, Institute of Physics, Eötvös University, Budapest 1117, Hungary*
- <sup>65</sup> *Nicolaus Copernicus Astronomical Center, Polish Academy of Sciences, 00-716, Warsaw, Poland*
- <sup>66</sup> *OzGrav, University of Adelaide, Adelaide, South Australia 5005, Australia*
- <sup>67</sup> *Dipartimento di Scienze Matematiche, Fisiche e Informatiche, Università di Parma, I-43124 Parma, Italy*
- <sup>68</sup> *INFN, Sezione di Milano Bicocca, Gruppo Collegato di Parma, I-43124 Parma, Italy*
- <sup>69</sup> *Rochester Institute of Technology, Rochester, NY 14623, USA*
- <sup>70</sup> *University of Birmingham, Birmingham B15 2TT, United Kingdom*
- <sup>71</sup> *INFN, Sezione di Genova, I-16146 Genova, Italy*
- <sup>72</sup> *RRCAT, Indore, Madhya Pradesh 452013, India*
- <sup>73</sup> *Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia*
- <sup>74</sup> *OzGrav, University of Western Australia, Crawley, Western Australia 6009, Australia*
- <sup>75</sup> *Department of Astrophysics/IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands*
- <sup>76</sup> *Artemis, Université Côte d'Azur, Observatoire Côte d'Azur, CNRS, CS 34229, F-06304 Nice Cedex 4, France*
- <sup>77</sup> *Physik-Institut, University of Zurich, Winterthurerstrasse 190, 8057 Zurich, Switzerland*
- <sup>78</sup> *Univ Rennes, CNRS, Institut FOTON - UMR6082, F-3500 Rennes, France*
- <sup>79</sup> *Washington State University, Pullman, WA 99164, USA*
- <sup>80</sup> *University of Oregon, Eugene, OR 97403, USA*
- <sup>81</sup> *Laboratoire Kastler Brossel, Sorbonne Université, CNRS, ENS-Université PSL, Collège de France, F-75005 Paris, France*
- <sup>82</sup> *Università degli Studi di Urbino 'Carlo Bo,' I-61029 Urbino, Italy*
- <sup>83</sup> *INFN, Sezione di Firenze, I-50019 Sesto Fiorentino, Firenze, Italy*
- <sup>84</sup> *Astronomical Observatory Warsaw University, 00-478 Warsaw, Poland*
- <sup>85</sup> *VU University Amsterdam, 1081 HV Amsterdam, The Netherlands*
- <sup>86</sup> *University of Maryland, College Park, MD 20742, USA*
- <sup>87</sup> *School of Physics, Georgia Institute of Technology, Atlanta, GA 30332, USA*
- <sup>88</sup> *Université Claude Bernard Lyon 1, F-69622 Villeurbanne, France*
- <sup>89</sup> *Università di Napoli 'Federico II,' Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy*
- <sup>90</sup> *Dipartimento di Fisica, Università degli Studi di Genova, I-16146 Genova, Italy*
- <sup>91</sup> *RESCEU, University of Tokyo, Tokyo, 113-0033, Japan.*
- <sup>92</sup> *Tsinghua University, Beijing 100084, China*
- <sup>93</sup> *Texas Tech University, Lubbock, TX 79409, USA*
- <sup>94</sup> *Kenyon College, Gambier, OH 43022, USA*
- <sup>95</sup> *The University of Mississippi, University, MS 38677, USA*
- <sup>96</sup> *Museo Storico della Fisica e Centro Studi e Ricerche "Enrico Fermi", I-00184 Roma, Italyrico Fermi, I-00184 Roma, Italy*
- <sup>97</sup> *The Pennsylvania State University, University Park, PA 16802, USA*
- <sup>98</sup> *National Tsing Hua University, Hsinchu City, 30013 Taiwan, Republic of China*
- <sup>99</sup> *Charles Sturt University, Wagga Wagga, New South Wales 2678, Australia*
- <sup>100</sup> *Center for Interdisciplinary Exploration & Research in Astrophysics (CIERA), Northwestern University, Evanston, IL 60208, USA*
- <sup>101</sup> *University of Chicago, Chicago, IL 60637, USA*
- <sup>102</sup> *Pusan National University, Busan 46241, Korea*

- <sup>103</sup> Carleton College, Northfield, MN 55057, USA
- <sup>104</sup> The Chinese University of Hong Kong, Shatin, NT, Hong Kong
- <sup>105</sup> INAF, Osservatorio Astronomico di Padova, I-35122 Padova, Italy
- <sup>106</sup> INFN, Trento Institute for Fundamental Physics and Applications, I-38123 Povo, Trento, Italy
- <sup>107</sup> OzGrav, University of Melbourne, Parkville, Victoria 3010, Australia
- <sup>108</sup> Università di Roma 'La Sapienza,' I-00185 Roma, Italy
- <sup>109</sup> Université Libre de Bruxelles, Brussels 1050, Belgium
- <sup>110</sup> Sonoma State University, Rohnert Park, CA 94928, USA
- <sup>111</sup> Departamento de Matemáticas, Universitat de València, E-46100 Burjassot, València, Spain
- <sup>112</sup> Columbia University, New York, NY 10027, USA
- <sup>113</sup> Montana State University, Bozeman, MT 59717, USA
- <sup>114</sup> Universitat de les Illes Balears, IAC3—IEEC, E-07122 Palma de Mallorca, Spain
- <sup>115</sup> University of Rhode Island
- <sup>116</sup> The University of Texas Rio Grande Valley, Brownsville, TX 78520, USA
- <sup>117</sup> Bellevue College, Bellevue, WA 98007, USA
- <sup>118</sup> Institute for Plasma Research, Bhat, Gandhinagar 382428, India
- <sup>119</sup> The University of Sheffield, Sheffield S10 2TN, United Kingdom
- <sup>120</sup> California State University, Los Angeles, 5151 State University Dr, Los Angeles, CA 90032, USA
- <sup>121</sup> Università di Trento, Dipartimento di Fisica, I-38123 Povo, Trento, Italy
- <sup>122</sup> Montclair State University, Montclair, NJ 07043, USA
- <sup>123</sup> National Astronomical Observatory of Japan, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan
- <sup>124</sup> Canadian Institute for Theoretical Astrophysics, University of Toronto, Toronto, Ontario M5S 3H8, Canada
- <sup>125</sup> Observatori Astronòmic, Universitat de València, E-46980 Paterna, València, Spain
- <sup>126</sup> School of Mathematics, University of Edinburgh, Edinburgh EH9 3FD, United Kingdom
- <sup>127</sup> University and Institute of Advanced Research, Koba Institutional Area, Gandhinagar Gujarat 382007, India
- <sup>128</sup> Indian Institute of Technology Bombay
- <sup>129</sup> University of Szeged, Dóm tér 9, Szeged 6720, Hungary
- <sup>130</sup> Tata Institute of Fundamental Research, Mumbai 400005, India
- <sup>131</sup> INAF, Osservatorio Astronomico di Capodimonte, I-80131, Napoli, Italy
- <sup>132</sup> University of Michigan, Ann Arbor, MI 48109, USA
- <sup>133</sup> Abilene Christian University, Abilene, TX 79699, USA
- <sup>134</sup> American University, Washington, D.C. 20016, USA
- <sup>135</sup> University of Białystok, 15-424 Białystok, Poland
- <sup>136</sup> University of Southampton, Southampton SO17 1BJ, United Kingdom
- <sup>137</sup> University of Washington Bothell, 18115 Campus Way NE, Bothell, WA 98011, USA
- <sup>138</sup> Institute of Applied Physics, Nizhny Novgorod, 603950, Russia
- <sup>139</sup> Korea Astronomy and Space Science Institute, Daejeon 34055, Korea
- <sup>140</sup> Inje University Gimhae, South Gyeongsang 50834, Korea
- <sup>141</sup> National Institute for Mathematical Sciences, Daejeon 34047, Korea
- <sup>142</sup> Ulsan National Institute of Science and Technology
- <sup>143</sup> NCBJ, 05-400 Świerk-Otwock, Poland
- <sup>144</sup> Institute of Mathematics, Polish Academy of Sciences, 00656 Warsaw, Poland
- <sup>145</sup> Cornell Universtiy
- <sup>146</sup> Hillsdale College, Hillsdale, MI 49242, USA
- <sup>147</sup> Hanyang University, Seoul 04763, Korea
- <sup>148</sup> NASA Marshall Space Flight Center, Huntsville, AL 35811, USA
- <sup>149</sup> Dipartimento di Fisica, Università degli Studi Roma Tre, I-00154 Roma, Italy
- <sup>150</sup> INFN, Sezione di Roma Tre, I-00154 Roma, Italy
- <sup>151</sup> ESPCI, CNRS, F-75005 Paris, France
- <sup>152</sup> Southern University and A&M College, Baton Rouge, LA 70813, USA
- <sup>153</sup> College of William and Mary, Williamsburg, VA 23187, USA
- <sup>154</sup> Centre Scientifique de Monaco, 8 quai Antoine 1er, MC-98000, Monaco
- <sup>155</sup> Indian Institute of Technology Madras, Chennai 600036, India
- <sup>156</sup> INFN Sezione di Torino, Via P. Giuria 1, I-10125 Torino, Italy

- <sup>157</sup> *Institut des Hautes Etudes Scientifiques, F-91440 Bures-sur-Yvette, France*
- <sup>158</sup> *IISER-Kolkata, Mohanpur, West Bengal 741252, India*
- <sup>159</sup> *Whitman College, 345 Boyer Avenue, Walla Walla, WA 99362 USA*
- <sup>160</sup> *Scuola Normale Superiore, Piazza dei Cavalieri 7, I-56126 Pisa, Italy*
- <sup>161</sup> *Université de Lyon, F-69361 Lyon, France*
- <sup>162</sup> *Hobart and William Smith Colleges, Geneva, NY 14456, USA*
- <sup>163</sup> *Università degli Studi di Firenze, I-50121 Firenze, Italy*
- <sup>164</sup> *OzGrav, Swinburne University of Technology, Hawthorn VIC 3122, Australia*
- <sup>165</sup> *Janusz Gil Institute of Astronomy, University of Zielona Góra, 65-265 Zielona Góra, Poland*
- <sup>166</sup> *University of Washington, Seattle, WA 98195, USA*
- <sup>167</sup> *SUPA, University of the West of Scotland, Paisley PA1 2BE, United Kingdom*
- <sup>168</sup> *King's College London, University of London, London WC2R 2LS, United Kingdom*
- <sup>169</sup> *Indian Institute of Technology, Gandhinagar Ahmedabad Gujarat 382424, India*
- <sup>170</sup> *Indian Institute of Technology Hyderabad, Sangareddy, Khandi, Telangana 502285, India*
- <sup>171</sup> *International Institute of Physics, Universidade Federal do Rio Grande do Norte, Natal RN 59078-970, Brazil*
- <sup>172</sup> *Villanova University, 800 Lancaster Ave, Villanova, PA 19085, USA*
- <sup>173</sup> *Andrews University, Berrien Springs, MI 49104, USA*
- <sup>174</sup> *Università di Siena, I-53100 Siena, Italy*
- <sup>175</sup> *Trinity University, San Antonio, TX 78212, USA*
- <sup>176</sup> *Van Swinderen Institute for Particle Physics and Gravity, University of Groningen, Nijenborgh 4, 9747 AG Groningen, The Netherlands*
- <sup>177</sup> *Colorado State University, Fort Collins, CO 80523, USA*

## ABSTRACT

We present a search for prompt gamma-ray counterparts to compact binary coalescence gravitational wave (GW) candidates from Advanced LIGO's first observing run (O1). As demonstrated by the multimessenger observations of GW170817/GRB 170817A, electromagnetic and GW observations provide complementary information about the astrophysical source and, in the case of weaker candidates, may strengthen the case for an astrophysical origin. Here we investigate low-significance GW candidates from the O1 compact-binary coalescence searches using the *Fermi* Gamma-Ray Burst Monitor (GBM), leveraging its all-sky and broad energy coverage. Candidates are ranked and compared to background to measure significance. Those with false alarm rates of less than  $10^{-5}$  Hz (about one per day) are used as the search sample for gamma-ray follow-up. No GW candidates were found to be coincident with gamma-ray transients independently identified by blind searches of the GBM data. In addition, GW candidate event times were followed up by a separate targeted search of GBM data. Among the resulting GBM events, the two with lowest false alarm rates were the gamma-ray transient GW150914-GBM presented in [Connaughton et al. \(2016\)](#) and a solar flare in chance coincidence with a GW candidate.

\* Deceased, February 2018.

† Deceased, November 2017.

## 1. INTRODUCTION

The first observing run (O1) of the Advanced LIGO detectors (Aasi et al. 2015) marked the dawn of gravitational wave (GW) astronomy with the ground-breaking discovery of merging black holes (BHs) (Abbott et al. 2016g,e) and Abbott et al. (2016b). The second observing run (O2) continued unveiling the population of binary black holes (Abbott et al. 2017b,c,d), saw the addition of the Virgo observatory to the detector network (Abbott et al. 2017d), and culminated in the spectacular multimessenger observations of a binary neutron star merger, summarized in Abbott et al. (2017e,f). Simultaneously observing the same astrophysical event in both gravitational and electromagnetic radiation will continue to uniquely enrich our understanding of sources. Because GWs negligibly interact with matter they are directly encoded with information about the central engines of the most violent, dynamic processes in the universe. Electromagnetic (EM) waves, on the other hand, are tightly coupled to matter thus providing information about the material and surrounding environment being affected by the central engine (Metzger & Berger 2012).

Several astrophysical transient phenomena are thought to produce GW and EM emission strong enough to be detected by current or proposed observatories including soft gamma-ray repeaters, rapidly rotating core collapse supernovae, binary neutron star mergers and gamma-ray bursts. Here we focus on short gamma-ray bursts (SGRBs), now directly confirmed to arise from the mergers of compact stellar remnants to which ground-based GW detectors are most sensitive (Abbott et al. 2017a). This paper is limited in scope to analysis of times during O1, and focuses on the follow-up searches of data from the *Fermi* Gamma-ray Burst Monitor (GBM) near in time to the GW search candidates. Results for the search in GW data for known GRBs which occurred during O1 can be found in Abbott et al. (2017g).

Despite the consensus view that there would be no bright EM emission associated with stellar-mass BBH mergers, comprehensive observing campaigns with EM observatories were carried out. For example, twenty five participating teams of observers received and responded to notifications of the GW150914 detection, with follow-up observations taken from the radio to gamma-ray bands (Abbott et al. 2016f). In the follow-up analysis of time around the first BBH merger observation, GW150914, the *Fermi* Gamma-ray Burst Monitor found a weak transient signal (Connaughton et al. 2016), GW150914-GBM, though no corresponding signal was observed in other gamma-ray instruments (Hurley et al. 2016; Savchenko et al. 2016; Tavani et al. 2016), nor

was a similar signal found in relation to the other unambiguous GW detection in O1, GW151226 (Abbott et al. 2016e; Adriani et al. 2016; Racusin et al. 2017), though 85% of the LIGO localization for GW151226 was out of view by *Fermi*, occulted by the Earth (Racusin et al. 2017).

The case for a GW candidate too low in significance to be claimed as an unambiguous detection would be strengthened by an EM detection consistent with the putative GW event. To that end, this paper reports on the search of GBM data for potential EM counterparts to any GW binary merger candidates found in the O1 data set with false alarm rate (FAR) below  $10^{-5}$  Hz by the LIGO-Virgo analyses. Two GW candidates were identified in coincidence with gamma-ray signals exceeding the GBM background: a gamma ray transient found in the follow up of GW150914, previously reported in Connaughton et al. (2016), and a second GBM transient most likely due to a chance coincidence with a solar flare.

The paper is organized as follows. In Section 2 we describe the data analysis methods used to identify compact binary coalescence (CBC) triggers in GW data and to identify gamma-ray transients in the GBM data. Section 3 presents the search results. Section 4 summarizes the results and conclusions, and discusses the role for GBM and LIGO-Virgo joint analysis during future observing runs.

## 2. DATA ANALYSIS METHODS

### 2.1. *Advanced LIGO*

The LIGO observatories are two modified Michelson interferometers located in Livingston, Louisiana and Hanford, Washington. Each arm of the interferometer is formed by two mirrors separated by 4 km. Passing GWs produce a strain, typically denoted as  $h$ , which changes the separation of the mirrors, inducing a relative phase difference in the light when it returns to the beam-splitter which transmits an optical signal proportional to the incident strain. The large separation between the observatories reduces the false-alarm background due to coincident instrumental or environmental contamination, and differences in time and phase of arrival constrain the sky location of astrophysical sources. For a full description see Aasi et al. (2015).

The data are analyzed for continuous (Aasi et al. 2014b; Abbott et al. 2016h), stochastic (Abbott et al. 2016j), and transient astrophysical sources. Transient searches include “all sky” searches for new GW sources, and “targeted” searches for known astrophysical objects, e.g. GRBs (Abbott et al. 2017g) and supernovae (Abbott et al. 2016c). All sky sources can be



divided into two broad classes: Transients for which the gravitational-wave signal is well modeled, enabling the use of template waveforms in matched filtering searches for binary black holes (Abbott et al. 2016b), binary neutron stars and neutron star–black hole systems (Abbott et al. 2016i), and cosmic strings (Aasi et al. 2014a); and “unmodeled” sources for which excess power searches are used (Abbott et al. 2016a).

This work focuses on follow-up searches in GBM data of results from the template-based all sky search for CBC events, i.e. transient signals from the final stages of binaries containing compact stellar remnants (BHs and/or NSs) as gravitational-wave emission drives the objects to coalescence. Because the GWs from CBC sources are well modeled and computationally tractable (see for example Blanchet (2006); Buonanno & Damour (1999); Ajith et al. (2007)), data analysis strategies for such sources rely on template waveforms.

The candidate sources for GBM follow-up were acquired from two independently developed CBC search pipelines, PyCBC (Dal Canton et al. 2014; Usman et al. 2016; Nitz et al. 2017) and GstLAL (Privitera et al. 2014; Messick et al. 2017). Here we briefly summarize both searches. For a more detailed descriptions see Abbott et al. (2016d) and Abbott et al. (2016b). Both PyCBC and GstLAL use accurate models of the gravitational waveform from coalescing binaries to perform a phase-coherent matched filtering search. The mass and spin parameter space of plausible CBC signals is covered by a grid of  $\mathcal{O}(10^5)$  points and a *template waveform* is computed at each point. The pipelines cover a total mass range between  $2M_{\odot}$  and  $100M_{\odot}$  with mass ratio from 1 to  $\sim 100$ . The component spins are assumed to be parallel to the orbital angular momentum. Their dimensionless magnitudes can range from 0 to  $\sim 0.99$  for components heavier than  $2M_{\odot}$  and from 0 to 0.05 for components lighter than  $2M_{\odot}$ . Each template is correlated with the detectors’ strain data to calculate the signal-to-noise ratio (S/N). Data for which the same template produces a S/N above the pre-established detection threshold in each detector within 15 ms (slightly longer than the light travel time between detectors), are promoted to *coincident triggers* and ranked with a “network statistic” which, while different for PyCBC and GstLAL, is a function of the single-detector S/Ns and a measure of how consistent the data are with the template waveform. Both pipelines estimate the background distribution of the network statistic in the absence of GWs, which is then used to map the network rank of each coincident trigger to the FAR of the search. In the remainder of this work, the FAR from the GW searches will be referred to by  $\text{FAR}_{\text{GW}}$ . Candidates with

$\text{FAR}_{\text{GW}}$  less than one per month exceeded the threshold for alerting EM observing partners participating in the O1 follow-up program (e.g. Abbott et al. (2016f)). *Fermi* GBM has a unique data-sharing agreement with the LIGO Scientific and Virgo Collaborations, and has access to all CBC coincident triggers.

Because of their different assumptions and implementation details, PyCBC and GstLAL may produce different sets of triggers from the same data. Strong CBC signals are ranked by either pipeline to correspond to sufficiently low  $\text{FAR}_{\text{GW}}$ , although not numerically equal between pipelines. However, a weak signal can in principle be detected by only one pipeline. For this reason, here we combine the triggers produced by both pipelines into a superset of CBC candidates. For duplicate triggers found by both pipelines, we select the candidate with the lowest  $\text{FAR}_{\text{GW}}$ . The time range covered by the trigger set is from September 12, 2015 to January 19, 2016, with a total two-detector live-time of  $\sim 50$  days. Approximately  $10^4$  triggers were recorded by PyCBC and GstLAL. In keeping with the spirit of an “eyes-wide-open” search for gamma-ray counterparts, no selection cuts were made to select for CBC candidates favoring systems theorized to produce potential electromagnetic counterparts (i.e., binaries containing at least one NS and, in the case of NSBHs, a sufficiently comparable mass ratio to tidally disrupt the NS).

The CBC searches were repeated when improvements to the calibration of the detectors or the configuration of the pipelines were available. The triggers employed in the joint analyses were obtained from the first analysis using the offline CBC pipelines, and an intermediate calibration of the data. Data produced with the final calibration were not available at the onset of this study. The candidate list used here was compared to that produced using the final calibration, and the only differences between the two were consistent with random noise fluctuations around the detection threshold of the CBC searches.

## 2.2. Fermi GBM

The *Fermi* Gamma-Ray Burst Monitor is an all-sky instrument onboard NASA’s *Fermi* Gamma-ray Space Telescope. The GBM has an instantaneous field of view that subtends 70% of the celestial sphere (with the remaining 30% occulted by the Earth), with an 85% live-time, and energy coverage from 8 keV to 40 MeV. These capabilities make GBM an ideal partner in the searches for joint detections with ground-based GW observatories. The omnidirectional, broad energy coverage is accomplished through the use of 12 sodium iodide (NaI) detectors and two bismuth germanate (BGO) detectors.

The responses for the NaI detectors are strongly dependent on the incident angle to the source, while the responses for the BGO detectors are not. Additionally, responses from all detectors are dependent on the source spectrum, since the blockage or attenuation of photons by the spacecraft and scattering off spacecraft components are energy dependent. These properties of the response will cause a signal to appear with a different strength in each detector, and the relative observed strength between all detectors allows the localization of signals with an accuracy on the order of a few to tens of degrees (Connaughton et al. 2015). In practice, the expected photon rates at each detector are calculated over a grid of possible source spectrum templates and sky locations, and the observed relative rates are then compared to this expectation.

There are currently three distinct searches used on GBM data in the joint analysis of GW sources: The onboard *triggering* algorithm for automatic detection of strong GRBs (Meegan et al. 2009; von Kienlin et al. 2014; Bhat et al. 2016); the ground-based *untargeted* search for weak SGRBs; and the *targeted* search used to follow-up GW candidates (Blackburn et al. 2015). The onboard triggering and ground-based untargeted GBM searches are the means by which SGRBs are identified in the GBM data without any input about when or where a possible GW source may have occurred. The targeted search of GBM data uses as input the arrival time of a GW candidate and uses the angular and energy dependencies of the detector response to find spectrally- and spatially-coherent weak signals.

### 2.2.1. *Blind, Untargeted GRB searches*

The GBM flight software will trigger on board when counts rates in the time-binned data of two or more NaI detectors exceed the background count rates, determined by averages of  $\sim 15$  s of the previous data, by pre-specified thresholds (typically 4.5 sigma). The on-board trigger is activated by approximately 40 SGRBs/year. For a full description of GBM and the on-board triggering algorithm see (Meegan et al. 2009; von Kienlin et al. 2014; Bhat et al. 2016).

The untargeted, ground-based analysis searches primarily for SGRBs in continuous time-tagged event (CTTE) data. While a full description of the untargeted search will be in a forthcoming article we present a brief description here. The primary science goal of the search is to rapidly identify SGRBs missed by the triggering algorithm in order to maximize the potential for multimessenger observations, either through GW follow-up searches or to identify counterparts to astrophysical neutrinos. The analysis searches for signals

with durations ranging from 64 ms to 32 s. Candidates are identified when two NaI detectors have an excess number of counts, with one exceeding 2.5 standard deviations ( $\sigma$ ) relative to the background, and at least one other exceeding  $1.25 \sigma$ . Correcting for the number of time bins in a day, the Poisson probability of the excess in the two detectors must be less than  $10^{-6}$ . The untargeted search is currently finding approximately 80 SGRB candidates per year in addition to those found by the on-board triggering algorithm<sup>1</sup>.

### 2.2.2. *Follow-up Targeted GRB search*

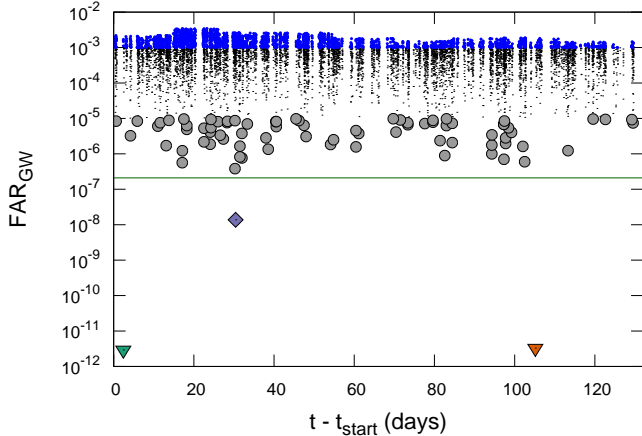
The targeted search was developed to examine GBM data for gamma-ray transients in coincidence with GW candidates from LIGO’s sixth science run (S6) and the concurrent second and third science runs (VSR2 and VSR3, respectively) of the Virgo 3 km interferometer in Cascina, Italy (Accadia et al. 2012). A description of how this method was applied to the S6+VSR2/3 data is presented in Blackburn et al. (2015).

The targeted search identifies candidate gamma-ray transients over a 30 s window centered on the GW trigger time. The 30 s time window is a conservative choice to not exclude precursor EM emission from the merger, or a delay between the compact objects’ coalescence and launching of the relativistic jet. Operating during O1, the targeted search utilized the daily continuous time (CTIME) data, with a nominal time resolution of 256 ms and 8 energy channels. The search was performed on binned data with bin-widths ranging from 256 ms to 8.192 s.

The targeted search uses three template spectra to coherently forward model the detector response. The templates are Band spectra (Band et al. 1993) comprised of two power-law components with spectral indices  $\alpha$  and  $\beta$  smoothly joined at the energy,  $E_{\text{peak}}$ . The definition of  $E_{\text{peak}}$ , and therefore the spectral model, are modified from the original definition in Band et al. (1993) and are explicitly shown in Connaughton et al. (2015). The templates are referred to as “soft,” “normal,” and “hard” spectra with Band function parameters  $\alpha$ ,  $\beta$ ,  $E_{\text{peak}} = (-1.9, -3.7, 70 \text{ keV})$ ,  $(-1, -2.3, 230 \text{ keV})$ , and  $(0, -1.5, 1 \text{ MeV})$  respectively. Each of the templates are utilized independently during the search, and then the most significant spectrum is identified for each time bin of the search.

The detection statistic  $\Lambda$  for the targeted search is derived from a log likelihood ratio formalism that measures the GBM signal strength relative to a polynomial fit to the background count rate. A detailed discussion of the

<sup>1</sup> [https://gcn.gsfc.nasa.gov/admin/fermi\\_gbm\\_subthreshold\\_announce.t](https://gcn.gsfc.nasa.gov/admin/fermi_gbm_subthreshold_announce.t)



**Figure 1.** Time series of CBC triggers relative to the start-time of the joint analysis (September 12, windows 2015). The horizontal line denotes the inverse live-time for the search sample ( $\sim 50$  days). The joint GBM/LIGO analyses used the blue points as the background sample, with  $\text{FAR}_{\text{GW}} > 10^{-3}$  Hz. The gray circles correspond to the search sample, with  $\text{FAR}_{\text{GW}} < 10^{-5}$  Hz, or approximately fewer than one per day. The green and orange downward-pointing triangles correspond to the upper limits on  $\text{FAR}_{\text{GW}}$  for GW150914 and GW151226, respectively. The purple diamond corresponds to LVT151012. These colors and marker style are preserved throughout the article.

search’s methodology can be found in Blackburn et al. (2015).

### 3. SEARCH RESULTS

#### 3.1. Advanced LIGO Sample

Figure 1 shows times relative to the start of the search sample (September 12, 2015) for the combined CBC triggers produced by the `GstLAL` and `PyCBC` pipelines. Intervals with no triggers correspond to times when one or both detectors were unable to produce science data, or when the data were vetoed due to well-understood instrumental or environmental disturbances. With the exception of the two high-confidence detections GW150914 (green triangle) (Abbott et al. 2016g) and GW151226 (orange triangle) (Abbott et al. 2016e), and the low-significance candidate LVT151012 (purple diamond) (Abbott et al. 2016b), no CBC triggers were more significant than the inverse live-time of the search  $\sim 2 \times 10^{-7}$  Hz (black horizontal line). For the targeted GBM search, the 81 subthreshold candidates with  $\text{FAR}_{\text{GW}} < 10^{-5}$  Hz ( $\lesssim 1 \text{ day}^{-1}$ ) marked by gray circles were used as the search sample, while the 2935 blue dots ( $\text{FAR}_{\text{GW}} > 10^{-3}$  Hz) served as the background sample. The candidates between the background and search samples (black dots) were discarded for two reasons. First, to avoid the scenario where the background

and search samples were sufficiently close in significance that the determination of which sample a candidate belonged in was up to random chance. Second, due to the long tails in the GBM search statistics, the threshold used for the background sample provided sufficient triggers to measure significance of anything in the search sample, making the inclusion of additional background triggers unnecessary.

#### 3.2. Triggering and Untargeted Search

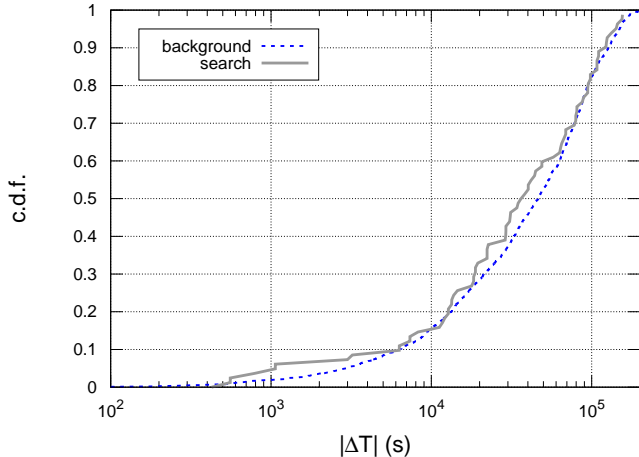
The simplest step in the joint analysis is to compare the list of times of known gamma-ray transients to the list of times of CBC candidates and search for a match. This comparison is complementary to the dedicated follow-up of known GRBs in GW data of GRBs, as reported in Abbott et al. (2017g) in two ways: first, here we do not impose a maximum time offset between the GW and gamma-ray candidates, and second, this study includes sub-threshold GBM candidates from the untargeted search which, in O1, were not incorporated in the LIGO-Virgo GRB follow-up searches.

For the joint analysis, events found in the GBM triggering and untargeted searches during O1 are combined and correlated with the combined list of CBC candidates, searching for coincident events in the two observatories. The GBM trigger list includes 115 GRBs, about half of which are SGRBs. The search statistic is the shortest absolute time offset between a GRB and the search sample of CBC triggers, allowing for the possibility that the EM signal can precede or follow the compact binary merger event.

Using CBC candidates with  $\text{FAR}_{\text{GW}} > 10^{-3}$  Hz as a background sample (blue dots in Fig. 1) and  $\text{FAR}_{\text{GW}} < 10^{-5}$  as the search sample, we find the 90<sup>th</sup> percentile of the absolute offset time from the closest GRB in our sample to be  $\sim 10^4$  s. Approximately 10% of the search sample are closer in time to a GRB, as expected for a null result based on the background sample. The agreement between cumulative distribution functions of the absolute time offset  $|\Delta T|$  for the background sample (blue dotted line) and the search sample (gray solid line) is shown in Figure 2. The shortest interval between the event times of a triggered GRB and a CBC candidate was  $|\Delta T| = 426.26$  s.

#### 3.3. Targeted Search

The targeted search follow-up of `PyCBC` and `GstLAL` CBC coincident triggers used the same selection criteria as described earlier, with an additional down-selection of triggers which appeared in both CBC pipelines within 1 s of each other, in which case the trigger with the highest  $\text{FAR}_{\text{GW}}$  was removed from the sample. By this procedure, there were 17 times that a trigger was found by



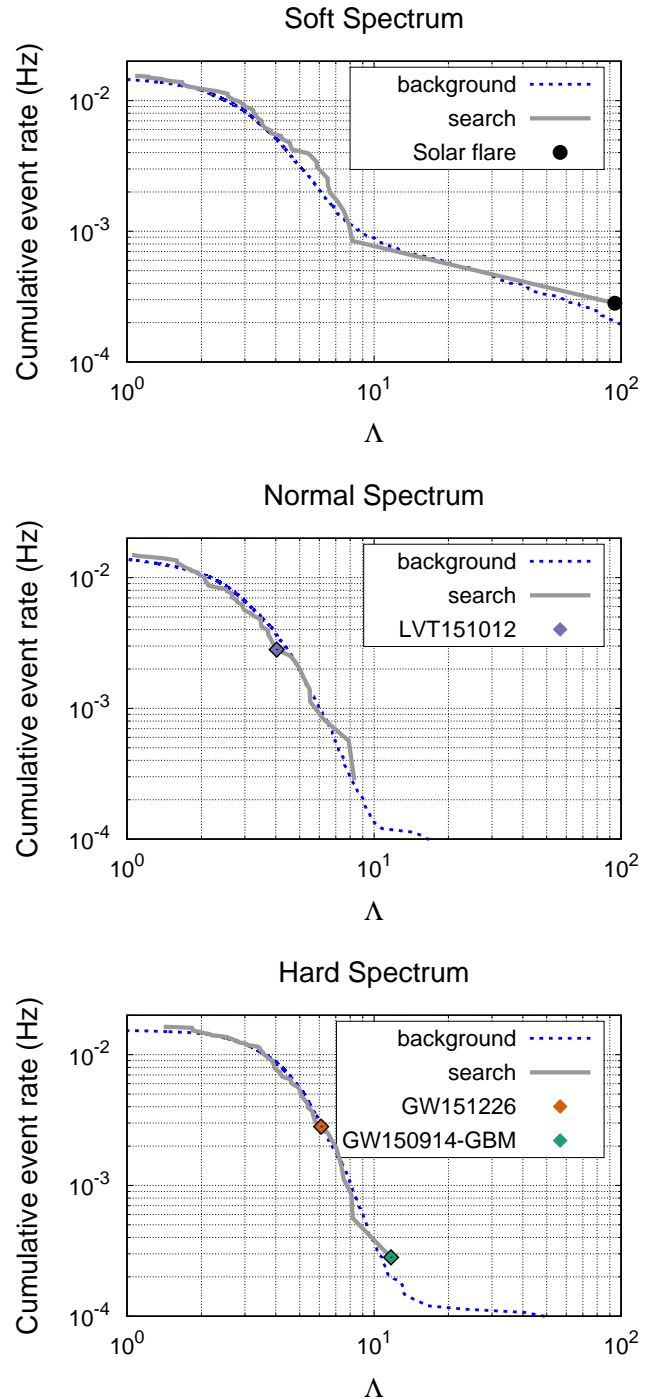
**Figure 2.** Cumulative distribution function of minimum absolute offset time between a GRB found by the triggered or untargeted GBM searches and a CBC candidate. The search sample (gray solid line) is consistent with the background distribution (blue dotted line), indicating that there are no significant coincidences between independently identified GBM and CBC triggers during O1.

both pipelines: Nine triggers with no offset ( $|\Delta T| < 0.01$  s), seven triggers with offset of  $|\Delta T| \sim 0.01$  s, and one occurrence with an offset of  $|\Delta T| \sim 0.02$  s. Figure 3 shows the cumulative event rate of the GBM background and search samples as a function of the detection statistic  $\Lambda$ . The distribution of the search samples is largely consistent with that of the background.

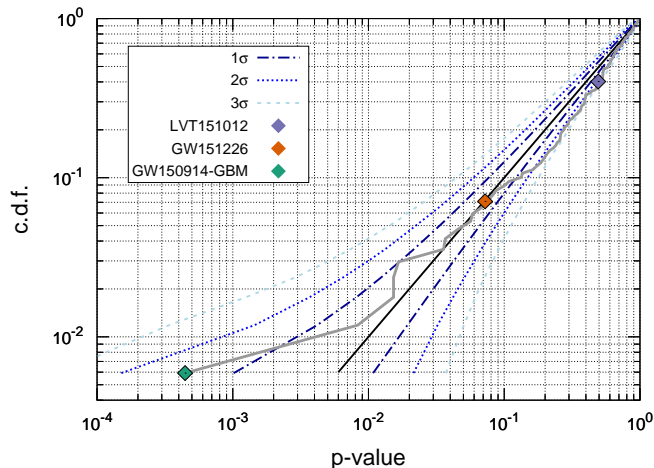
The GBM FAR,  $\text{FAR}_{\text{GBM}}$ , for coincidences in the targeted follow-up of CBC candidates with  $\text{FAR}_{\text{GW}} < 10^{-5}$  Hz can be mapped to a false alarm probability  $p$ -value. Here  $p$ -value is similar to the quantity defined as the false alarm probability in Connaughton et al. (2016), which used an analytic approximation to estimate the significance of a candidate.

However, the significance calculation used in this work differs from that in Connaughton et al. (2016). For a given candidate found in the GBM follow-up search, the False Alarm Probability (FAP) is intended to report the probability of detecting a false associated signal (either statistical fluctuation or real unassociated signal) with a ranking that is at least as significant as the candidate.

As in Connaughton et al. (2016), search candidates are first ranked by  $R = 1/(\text{FAR}_{\text{GBM}}|\Delta T|)$  where  $\text{FAR}_{\text{GBM}}$  is the cumulative event rate of the background distribution at the candidate's value for the search statistic (see Fig. 3), and  $|\Delta T|$  is the absolute value of the time offset between the closest GW trigger time and the GBM candidate. We impose a minimum  $\Delta T_{\text{min}}$  equal to the bin-width of the GBM data, which is 256 ms.



**Figure 3.** Cumulative event rates of the background (blue dotted lines) and search (gray solid lines) samples for the GBM targeted follow-up of CBC triggers as a function of the detection statistic  $\Lambda$ . The panels in the figure from top to bottom correspond to results from the soft, normal, and hard template spectra. The distribution of the search samples is largely consistent with that of the background. The two most significant events from the search distribution are the GBM transient identified in the follow-up to GW150914 (green diamond) and a soft, long-duration, transient found 26 s after a low-significance GW trigger. The long, soft transient is likely due to solar activity, and in a chance coincidence with a GW trigger.



**Figure 4.** Cumulative distribution function of empirically-determined of  $p$ -values for targeted search. The black solid line corresponds to the null hypothesis that the search sample is consistent with background. The blue dashed lines envelop the 1-, 2-, and 3- $\sigma$  confidence intervals. The highest ranking (lowest  $p$ -value) event is GW150914-GBM (green diamond) which, for this search, has a significance of  $\gtrsim 1.5\sigma$ .

For this analysis we empirically measure the FAP by using background noise events from the GW searches to seed GBM follow-up. Background events are ranked using the same statistic  $R$ . The FAP of a candidate event is the  $p$ -value of the background distribution with that candidate’s ranking statistic,  $R$ .

Figure 4 shows the cumulative distribution of  $p$ -values for all of the CBC triggers follow-up candidates. GW150914-GBM had the lowest  $p$ -value of any follow-up event, due to the fact that it had both the lowest  $\text{FAR}_{\text{GBM}}$  and had the shortest time offset to its GW trigger time. The  $p$ -value for GW150914-GBM is an upper bound, because although there were more significant events in the background sample as determined by detection statistic  $\Lambda$ , the short time offset from the GW trigger caused it to have a higher ranking statistic,  $R$ , than any of the events in the background distribution. The blue dashed lines envelope the expected 1- (dark blue) 2- (blue) and 3- $\sigma$  (light blue) confidence intervals assuming that the search distribution is statistically identical to the background distributions based on the number of events analyzed. From this it is clear that GW150914-GBM is a  $\gtrsim 1.5\sigma$  event. This is a lower significance than reported in [Connaughton et al. \(2016\)](#) owing to the large number of CBC events in the search sample, due to the low threshold chosen for trigger selection. This was intentional, as the point of this study was to search for sub-threshold joint-detections and we treat all CBC candidates in the search sample as equally worthwhile of follow-up.

Also note that the second lowest false alarm rate ( $\text{FAR}_{\text{GBM}} = 2.6 \times 10^{-4}$  Hz) of the search occurred in the analysis of the GBM transient detected at 2015-09-29 12:15:43.6 UTC. The event was at the edge of the search window, 26 s before the GW trigger time that prompted the GBM analysis, resulting in  $p$ -value  $\sim 1$ . The event was found using the soft template and the longest (8.192 s) timescale of the search. A search for other possibly associated astrophysical transients at that time revealed that the GBM event occurred during the exponential decay tail of an M-class solar flare (e.g. GBM triggers 150929346 and 150929553). The GBM localization for this candidate is consistent with the Sun and inconsistent with the LIGO sky map, thus we conclude this GBM event is due to solar activity.

#### 4. CONCLUSIONS AND DISCUSSION

The GBM follow-up analyses searched for gamma-ray transients associated with a list of CBC candidates during O1. These CBC candidates were well below the standards for the GW trigger alone to be considered likely due to a compact merger, but significant enough that if a gamma-ray transient was found by GBM in coincidence, it would support an astrophysical origin of the GW transient.

There were no coincidences between GBM transients found by the triggered or untargeted searches and the LIGO CBC candidates. For the targeted follow-up of GW candidates, GW150914-GBM is the most significant event of the search. The analysis presented here was not designed to revisit the significance of the GW150914-GBM association; additional observations of BBH mergers with GBM will be needed to establish or rule out the astrophysical nature of GW150914-GBM. No other significant transients were found in the GBM targeted search, ruling out the possibility of a low-significance gravitational-wave event being identified by its association with a gamma-ray transient during O1.

The significance of GW150914-GBM reported by the search described in this paper ( $\gtrsim 1.5\sigma$ ) is conservative in part because of the choice to not include in the candidate ranking the significance of the GW triggers. By not using the GW significance in our assessment of the validity of a GBM counterpart, the targeted search is hampered by implicitly placing marginal GW events on equal footing with more plausible ones. Joint GW/EM events may prove to be rare, and the search strategy as adopted is to cast a wide, model-independent, net when looking for possible sub-threshold coincidences in both observatories. Deriving a joint significance which includes information about the likelihood of the GW candidate being astrophysical is subtle; having a high-

significance GW event does not necessarily imply a high likelihood for an associated gamma-ray transient. Incorporating the significance of the GW candidate into the GBM follow-up analysis is an area of investigation for future joint observing campaigns (Ashton et al. 2017).

For the targeted search, the long tails in the background distribution out to high values of the ranking statistic limited the sensitivity of the search in O1. Further analysis have revealed that many ( $> 20\%$ ) of the high likelihood events in the background can be attributed to large background fluctuations due to high particle activity associated with *Fermi*'s proximity to the South Atlantic Anomaly (SAA). For analyses of later observing runs, the background fitting procedure and the vetoing of data from entry and exit of the SAA have been improved (Goldstein et al. 2016).

The multimessenger observations of the binary neutron star merger associated with a short GRB, and the resulting kilonova and GRB afterglow observations, demonstrated a best-case scenario of a bright, nearby transient which was identified unambiguously gravitationally and electromagnetically. However, the event had the closest recorded redshift for any GRB despite its relatively low brightness (Goldstein et al. 2017) while the binary neutron star system was well within the detection horizon for the LIGO-Virgo network (Abbott et al. 2017e). It is therefore not difficult to imagine scenarios where sub-threshold searches would have been critical for identifying the association. Joint analyses that can elevate the significance of candidate events are needed to best exploit the rich datasets from gravitational-wave and gamma-ray survey instruments. The searches reported here, while not producing any unambiguous counterparts, will serve as a foundation for future joint analysis.

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