

X-ray spectra reveal the reawakening of the repeat changing-look AGN NGC 1566

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ABSTRACT

We present simultaneous *XMM-Newton* and *NuSTAR* observations of the repeat changing-look AGN NGC 1566, which dramatically increased in brightness in the IR to X-ray bands in 2018. The broad-band X-ray spectrum was taken at the peak of the outburst and is typical of Seyfert 1 AGN. The spectrum shows a soft excess, Compton hump, warm absorption and reflection, ruling out tidal disruption as the cause of the outburst and demonstrating that a ‘standard’ accretion disk can develop very rapidly. The high resolution grating spectrum reveals that the outburst has launched a $\sim 500 \text{ km s}^{-1}$ outflow, and shows photoionised emission lines from rest-frame gas. We discuss possible mechanisms for the outburst, and conclude that it is most likely caused by a disk instability.

Key words: galaxies: active – galaxies: individual: NGC 1566 – galaxies: Seyfert – accretion, accretion discs

1 INTRODUCTION

Active Galactic Nuclei (AGN) exhibit large changes in flux across many bands (e.g. review by Uttley & Casella 2014). In X-rays in particular, many AGN are known to change in luminosity by orders of magnitude on timescales from months to a few hours (e.g. Boller et al. 1997; Komossa et al. 2017). We have an ongoing observing program studying AGN in anomalous flux states, which has been successful in revealing unusual reflection-dominated states (Schartel et al. 2007; Grupe et al. 2012), strong absorption events (Grupe et al. 2013; Parker et al. 2014), and bright outbursts (Parker et al. 2016). Of particular interest among AGN with extreme variability are the changing-look AGN, which change their Seyfert classification (e.g. Penston & Perez 1984) and are sometimes associated with a switch from Compton-thick to Compton-thin absorption in the X-ray band (Guainazzi 2002; Matt et al. 2003).

NGC 1566 is a local ($z = 0.005$) face-on Seyfert galaxy, which was observed to increase dramatically in flux in 2018. This activity in NGC 1566 was detected serendipitously by *INTEGRAL* (Ducci et al. 2018) and followed up with the

Neil Gehrels Swift Observatory (*Swift*), which found it to be a factor of ~ 15 brighter than archival observations in X-rays (e.g. Kawamuro et al. 2013) and nearly 3 magnitudes brighter in the UVW2 filter (Kuin et al. 2018; Grupe et al. 2018). The *ASAS-SN* optical and *NEOWISE* infra-red lightcurves show that the source has been brightening since September 2017 (Dai et al. 2018; Cutri et al. 2018), and a SAAO optical spectrum showed (Oknyansky et al. 2018b) much stronger broad emission lines, consistent with a change in Seyfert type to 1.2 from its typical quiescent Sy 1.9–1.8 type (Oknyansky et al. 2018a). Interestingly, these outbursts are recurrent: Alloin et al. (1986) identify four separate periods of activity between 1970 and 1985, each lasting for ~ 1300 days and with associated increases in broad-line strength causing the Seyfert type to move between Sy2 and Sy1. Another outburst in 2010 is visible in the *Swift* Burst Alert Telescope (BAT) 105-month lightcurve¹.

In this letter, we present broadband X-ray spectroscopy of the peak of the 2018 outburst of NGC 1566 with *XMM-Newton* and *NuSTAR*.

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¹ <https://swift.gsfc.nasa.gov/results/bs105mon/216>

2 OBSERVATIONS AND DATA REDUCTION

2.1 XMM-Newton

Based on the detection of enhanced X-ray activity by *INTEGRAL*, we triggered a joint *XMM-Newton NuSTAR* target of opportunity (ToO) observation (*XMM-Newton* proposal ID 080084, PI Schartel). The observation length was 94 ks, taken on June 26 2018 (obs. ID 0800840201). We reduce the *XMM-Newton* data with the science analysis software (SAS) version 16.1.0. We reduce the Reflection Grating Spectrometer (RGS) data using the SAS task *rgsproc*. We filter the data for background flaring using a threshold of $0.2 \text{ counts s}^{-1}$. We combine the 1st and 2nd order spectra and the spectra from the two detectors into a single spectrum using *rgsfluxcombine*, and convert this to SPEX format using *rgsfmt*. We fit the RGS data from 8–35 Å. Due to the source's low redshift, the EPIC count rates are very high (40 s^{-1} in the pn), and the data is piled-up. To mitigate this, we use only the least affected EPIC-pn data, and use an annular source region. There are several ultra-luminous X-ray sources (ULXs) in NGC 1566 (Liu & Bregman 2005), but these are outside the window or obscured by the PSF of the AGN, which dominates the total count rate. We process the pn data using the *epproc* tool, and filter for background flares, leaving a clean exposure time of 65 ks. We extract source photons from a $30''$ radius annulus centered on the source, with an inner radius of $8''$, and a $40''$ background region, extracted from the furthest corner of the detector. We bin the EPIC spectrum to oversample the data by a factor of 3, and to a minimum signal to noise ratio of 6. We fit the pn data from 0.5–10 keV, excluding the 2–2.5 keV band where there is a calibration feature (see appendix of Marinucci et al. 2014).

We also reduce the data from the observation of NGC 1566 taken in 2015 when the source was faint, using the same procedure. A full analysis of these data will be presented in Tomás et al. (in prep.).

2.2 NuSTAR

A NuSTAR observation of 80 ks was taken simultaneously with the *XMM-Newton* exposure (obs. ID 80301601002). We reduced the *NuSTAR* data using the *NuSTAR* Data Analysis Software (NuSTARDAS) version 1.6.0. We extract source counts from a $60''$ circular extraction region, and background counts from a $90''$ circular extraction region on the same chip. We bin the spectra to a signal to noise ratio of 6, and to oversample the instrumental resolution by a factor of 3. We fit the FPMA and FPMB spectra separately, but group them in XSPEC for plotting purposes.

2.3 Swift

After the flare was detected *Swift* followup observations were immediately requested. The X-ray Telescope (XRT, Burrows et al. 2005) observations were mostly performed in Windowed Timing (WT) mode (Hill et al. 2005), however some initial observations were performed in photon counting (pc) mode. The XRT data were reduced using the task *xrt-pipeline*. Background and source events were extracted with XSELECT. For the WT data we used a 40 by 3 pixel box, rotated to match the spectrum orientation. Because the

first observations after the flare were in pc mode they were strongly affected by pileup, so we excluded the inner part of the PSF using an annular extraction region with inner and outer radii of $16.5''$ and $94.3''$. We used the latest (2013) response files, and created auxiliary response files (ARFs) with the FTOOL *xrtmkarf*. We binned the spectra with 20 counts per bin. Typical exposure times per spectrum were of the order of 1 ks.

The UVOT data of each observation were coadded in each filter. We extracted source counts from an extraction region with a radius of $3''$. The loss in the PSF was corrected with the command *uvotsource*. Count to flux density and magnitude conversion was performed based on the most recent calibration files (Poole et al. 2008; Breeveld et al. 2010). The UVOT data were corrected for Galactic reddening ($E_{B-V} = 0.025$; Schlegel et al. 1998).

3 RESULTS

3.1 RGS

We initially focus on the RGS spectrum to establish the extent of any absorption in the soft band. We fit the RGS data in SPEX (Kaastra et al. 1996) version 3.03.00. The spectrum shows several clear absorption lines from O, N and C, with an outflow velocity of $\sim 500 \text{ km s}^{-1}$, and emission lines at rest from OVII and NVI (Fig. 2). We fit the spectrum with a simple phenomenological powerlaw plus black-body continuum, two zones of absorption modeled with *xabs*, and three Gaussian emission lines. We also include a *hot* component to model Galactic absorption. This model gives a reasonable description of the data ($\chi^2/\text{dof} = 716/522$), although it misses some structure around 23 Å. As there are no strong emission lines in this region, this is likely either associated with the O edge, the shape of which can be modified by non-solar abundances or the presence of dust, or due to our simple model not perfectly modeling the continuum.

The entire RGS spectrum, including the emission lines, is far above the quiescent 2015 spectrum, shown in the lower set of points, so must be dominated by the AGN with negligible contribution from extended emission.

3.2 Broad-band

We now fit the broad-band 0.5–79 keV EPIC-pn and *NuSTAR* spectrum. We fit these data in XSPEC (Arnaud 1996) version 12.9.1p. We include Galactic absorption using *tbnew* (Wilms et al. 2000), with the column fixed at the value of $9 \times 10^{19} \text{ cm}^{-2}$ (Kalberla et al. 2005) consistent with the RGS spectrum. To account for the warm absorption identified in section 3.1, which cannot be directly constrained using the EPIC-pn spectrum, we write the SPEX model to a text file and convert it to an XSPEC table model (with no free parameters) using the *flux2tab* FTOOL. We also add a narrow Gaussian line at 0.566 keV to account for the OVII emission lines, which are unresolved.

Preliminary fitting from 3–10 keV with a power-law plus distant reflection (modelled with *xillver*, García et al. 2013) leaves some residuals around the Fe line, which are likely due to relativistic reflection from the accretion disk. We therefore fit the broad-band spectrum with the *relxill* relativistic

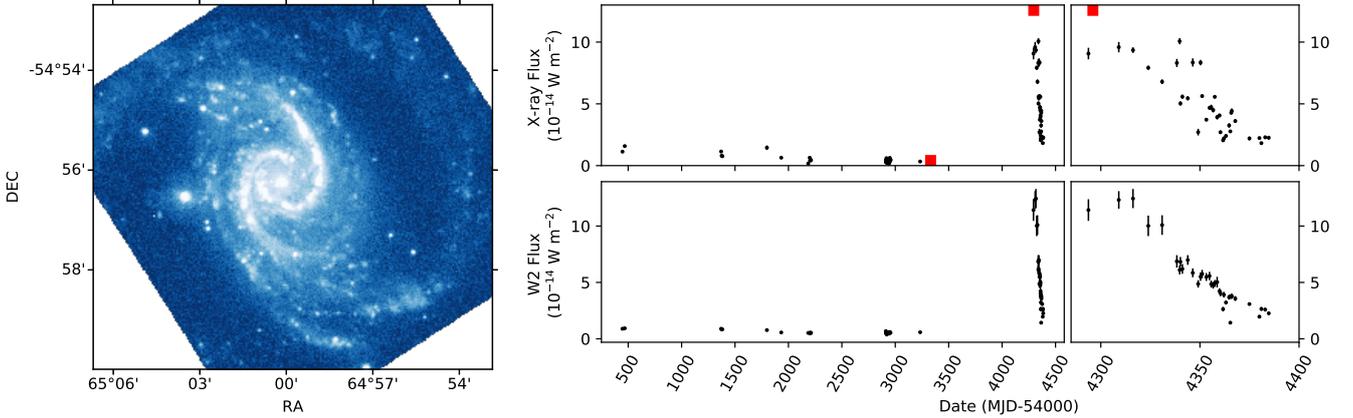


Figure 1. Left: Optical monitor (OM) UVW1 image. Right: Long-term *Swift* XRT and UVW2 (2246 Å) lightcurves. The 2015 and 2018 *XMM-Newton* EPIC-pn fluxes are marked with red squares. Observations started roughly at the peak of the optical outburst and have well sampled the peak and decay. At the time of writing, X-ray flux is around 1/5 of the peak flux, and has been stable for ~ 1 month.

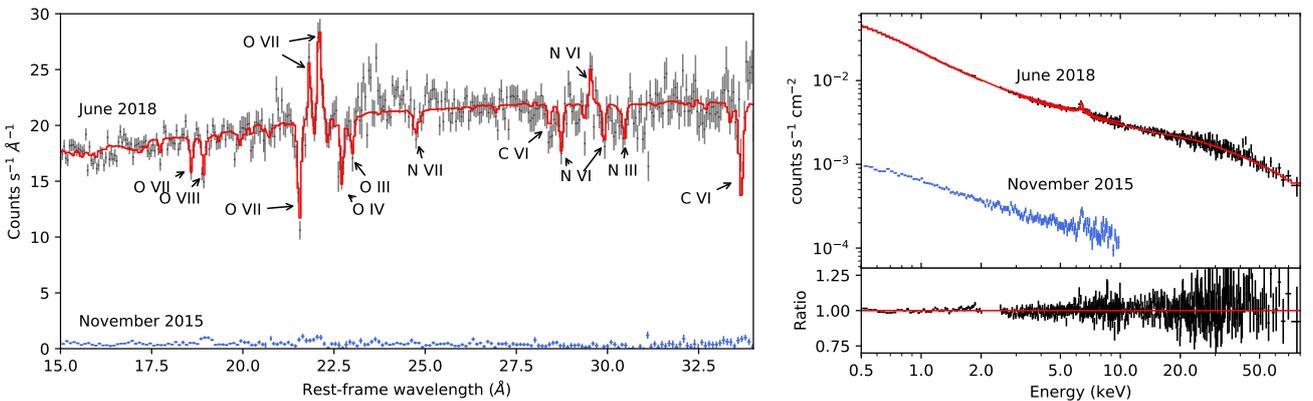


Figure 2. Left: RGS spectrum of NGC 1566, fit with two zones of warm absorption, and three emission lines. The lower set of points is the corresponding spectrum from 2015, when the source was quiescent. Right: Broad-band *XMM-Newton* and *NuSTAR* spectrum, corrected for the effective area of the instrument (but not unfolded from the instrumental resolution). The 2018 data is fit with our best-fit model (given in Table 2), and the data/model ratio is shown in the lower panel.

reflection model (García et al. 2014). As this component is weak, we fix the emissivity index to the classical value of 3. We tie the parameters of *xillver* to those of *relxill*, and include an additional soft excess component modelled with *nthcomp* (Zdziarski et al. 1996; Życki et al. 1999). Finally, we add a Gaussian line at ~ 6.9 keV to account for a narrow residual. We allow the photon index to vary between the three instruments, and include a constant multiplicative offset between them (the difference between *XMM-Newton* and *NuSTAR* is large, due to the annular extraction region for the pn). This model gives a good overall fit ($\chi^2/\text{dof}=818/785$), with no strong residuals. This is shown in Fig. 2, and the parameters are given in Table 2.

We estimate Eddington ratios for the 2015 and 2018 X-ray spectra using the 2–10 keV fluxes and a mass of $\sim 10^7$ (Woo & Urry 2002) and assuming a bolometric correction factor of 20 Vasudevan et al. (2009). This gives Eddington ratios of $\sim 0.2\%$ for 2015 and $\sim 5\%$ for 2018.

4 DISCUSSION

Overall, the X-ray spectrum of NGC 1566 is not unusual for a Sy 1 galaxy, showing standard spectral components. The ionization and velocity of the outflow are well within the normal range seen in other AGN (Laha et al. 2014). The column density is lower than generally seen, but this is likely due to detection bias: in more distant sources, a low column density outflow is unlikely to be detected. The rapid appearance of this spectrum after a period of quiescence is very unusual, and the increase in brightness by a factor of 30–40 within a short time period is extreme. We do not have X-ray data covering the rise period, but the ASAS-SN V-band lightcurve shows that the source flux began to rise around September 2017, peaking around the time of our observation (Dai et al. 2018). This means the outburst took ~ 9 months to reach the peak. This is longer than previous outbursts: Alloin et al. (1986) report a typical rise time of ~ 20 days, and the outburst visible in the 105 month *Swift* BAT catalog reaches the peak in 3 months. Given the lack of

Table 1. Best fit parameters for the model fit to the RGS data, shown in Fig. 2.

Comp.	Par.	Value	Description
<i>hot</i>	N_{H}	$(1.2 \pm 0.3) \times 10^{20} \text{ cm}^{-2}$	Galactic column
	kT	$0.25 \pm 0.03 \text{ keV}$	Temperature
<i>bb</i>	kT	$0.099 \pm 0.002 \text{ keV}$	Temperature
<i>pow</i>	Γ	$2.14^{+0.02}_{-0.03}$	Photon index
<i>xabs₁</i>	N_{H}	$(2.5 \pm 0.4) \times 10^{20} \text{ cm}^{-2}$	Column
	ξ	$10^{-0.7 \pm 0.1} \text{ erg cm s}^{-1}$	Ionisation
<i>xabs₂</i>	v_{RMS}	$65^{+32}_{-18} \text{ km s}^{-1}$	RMS velocity
	v	$541^{+96}_{-85} \text{ km s}^{-1}$	Outflow velocity
	N_{H}	$(2.3 \pm 0.3) \times 10^{20}$	Column
<i>xabs₃</i>	ξ	$10^{1.2 \pm 0.1} \text{ erg cm s}^{-1}$	Ionisation
	v_{RMS}	$185^{+35}_{-28} \text{ km s}^{-1}$	RMS velocity
	v	$472^{+43}_{-39} \text{ km s}^{-1}$	Outflow velocity
<i>gauss₁</i> *	λ	$21.83 \pm 0.01 \text{ \AA}$	Wavelength
	norm	$8 \pm 1 \text{ s}^{-1}$	Photon flux
<i>gauss₂</i> *	λ	$22.19 \pm 0.01 \text{ \AA}$	Wavelength
	norm	$18 \pm 2 \text{ s}^{-1}$	Photon flux
<i>gauss₃</i> *	λ	$29.69 \pm 0.02 \text{ \AA}$	Wavelength
	norm	$5 \pm 2 \text{ s}^{-1}$	Photon flux

*The three Gaussian lines correspond to the OVII intercombination and forbidden lines, and the NVII forbidden line, respectively.

Table 2. Best-fit parameters for the broad-band model shown in Fig. 2.

Comp.	Par.	Value	Description	
<i>nthcomp</i>	kT	$0.8 \pm 0.1 \text{ keV}$	Temperature	
	Γ	2.69 ± 0.02	Photon Index	
	norm	0.595 ± 0.001	Normalization	
<i>relxill</i>	a	< 0.25	Spin	
	i	$< 11^\circ$	Inclination	
	ξ	$10^{2.4 \pm 0.1} \text{ erg cm s}^{-1}$	Ionization	
	A_{Fe}	3.0 ± 0.2	Iron abundance	
<i>xillver</i>	R	$0.091^{+0.005}_{-0.004}$	Reflection fraction	
	Γ_{pn}	1.435 ± 0.003	Photon index	
	Γ_{FPMA}	1.624 ± 0.004	Photon index	
	Γ_{FPMB}	1.599 ± 0.004	Photon index	
	E_{cut}	$167 \pm 3 \text{ keV}$	Cutoff energy	
	norm	$(3.89 \pm 0.01) \times 10^{-4}$	Normalization	
	norm	$(7.6^{+0.4}_{-0.3}) \times 10^{-5}$	Normalization	
	<i>zgauss</i>	E	$6.85^{+0.04}_{-0.05} \text{ keV}$	Energy
		norm	$(1.6 \pm 0.3) \times 10^{-5}$	Normalization
	<i>const</i>	C_{FPMA}	0.820 ± 0.003	Constant offset
C_{FPMB}		0.844 ± 0.003	Constant offset	

X-ray coverage of previous outbursts, we cannot explain this difference from an X-ray perspective without further data.

Changing look events like this one involve flux changes in multiple wavebands on timescales far faster than a standard thin disk can evolve. There are several different mechanisms that are invoked to explain this phenomenon, such as obscuration, disk instabilities, and tidal disruption events (TDEs). Variable obscuration, where clumps of cold gas from the torus block the AGN emission (Matt et al. 2003), can be discounted in this case as the optical broad lines have responded directly to the increased optical and UV flux.

Large changes in accretion rate can produce large changes in the flux at all wavelengths, but for a ‘standard’

disk the timescales involved are far too long. Dexter & Begelman (2018) show that disks supported by magnetic pressure have much faster infall times, and can produce changes of a factor of 2–10 in optical to X-ray flux within 1–10 years. While promising for many changing-look events, this is likely not extreme enough to reproduce the rapid increase in flux seen in NGC 1566, which brightened by a factor of ~ 40 –70 and has had previous outbursts with rise times of less than a month.

In principle, TDEs can produce repeated events over many years (for example by repeated tidal stripping of a star), and this has been suggested as a possible explanation of repeat X-ray flares in IC 3599 (Campana et al. 2015). However, we consider this unlikely in this case. The theoretical rate of TDEs is low ($\sim 10^{-4}$ per Galaxy, per year), and to find one in such a nearby galaxy that already hosts an AGN would be very unusual. The similarity of this outburst to other changing look events, which are common in nearby galaxies (Runco et al. 2016, e.g.) at a rate far in excess of the predicted TDE rate, suggests a common non-TDE origin. Finally, the X-ray spectrum of NGC 1566 is a classic hard AGN spectrum, whereas X-ray spectra of TDEs are typically extremely soft (Komossa 2017, and references therein).

In our view, the most likely interpretation of this behaviour is an instability in the accretion disk. Grupe et al. (2015) discuss this for IC 3599, a Sy 1.9 AGN that has undergone at least two large outbursts. Saxton et al. (2015) and Grupe et al. (2015) explore the Lightman & Eardley (1974) instability, where the inner disk is quiescent until radiation pressure exceeds gas pressure, at which point the disk rapidly switches on. This mechanism produces variability on around the right timescales but requires that the rise time be longer than the decay time, a condition which has not been met by previous outbursts in NGC 1566. Additionally, the repeat time is set by the viscous time at the truncation radius which is typically decades, much longer than observed in NGC 1566 (Alloin et al. 1986). Ross et al. (2018) explain the changing look of a $z \sim 0.4$, $M \sim 10^{8.8} M_{\odot}$ quasar with a cooling front that propagates away from the ISCO, followed by a returning heating front over 20 years. This is similar in timescale after scaling for the mass (\sim a few months), but predominantly affects the flux at short wavelengths, so does not explain the uniform flux increase and decline in NGC 1566. Noda & Done (2018) propose that the drop in flux by a factor of 10 in Mrk 1018 and associated change from Sy 1.9 to Sy 1 is caused by a combination of the H instability, which produces the overall drop in luminosity, and evaporation of the inner disk, which causes the associated spectral hardening. These processes work in reverse, so a heating front caused by the H instability could propagate through the disk, causing an outburst. Interestingly, Noda & Done suggest that sources crossing a few per cent of Eddington should go through a changing look along with strong soft excess variability, and no soft excess was observed in earlier quiescent observations of NGC 1566 (Kawamuro et al. 2013). The timescales of the changing look events in these two sources are similar when the higher mass of Mrk 1048 ($M \sim 10^{7.8} M_{\odot}$, see Noda & Done 2018) is taken into account (9 months $\times 10 \sim 8$ years), although the variability amplitude is greater in NGC 1566. Given these similarities between the

outbursts, we consider it likely that they are due to the same mechanism.

The inclination measured from the relativistic Fe line is very low ($< 11^\circ$), which is consistent with the accretion disk being aligned with the face-on galaxy. While this is expected, we note that Middleton et al. (2016) found that there is not a strong correlation between host and disk inclinations. We find only an upper limit on the spin, of < 0.25 . A poor constraint is to be expected, given the weak feature and low inclination, which gives a narrow, hard to measure line. The low spin is interesting, and may be indicating some truncation of the accretion disk (although there are other explanations, e.g. Parker et al. 2018). While these results are intriguing, it is not possible to come to robust conclusions because of the limited signal available.

5 CONCLUSIONS

In 2018, the nearby Seyfert galaxy NGC 1566 went through a major outburst, reaching a peak X-ray flux in excess of 70 times the minimum observed with the *Swift* XRT, before slowly decaying by a factor of ~ 5 at the time of writing. This coincided with a change in Seyfert type to 1.2. We triggered a joint *XMM-Newton*/*NuSTAR* observation at the peak of the outburst, to obtain an X-ray perspective, and have uncovered several key results.

- The high resolution RGS spectrum shows several absorption lines from ionized gas, outflowing at $\sim 500 \text{ km s}^{-1}$. This outflow was likely launched or accelerated by the large increase in radiation pressure from the AGN outburst.
- There are several emission lines from ionized O, N, and Fe at rest in the *XMM-Newton* data, produced by photoionization of cold gas somewhere in the AGN system, such as the outer disk, torus, or BLR clouds. These lines are strong, which may indicate a significant solid-angle of cold gas.
- The broad-band *XMM-Newton*/*NuSTAR* spectrum requires a weak contribution from relativistic reflection off the accretion disk. The inclination is low ($< 11^\circ$), consistent with the face-on inclination of the galaxy.

Overall, the X-ray spectrum is not unusual for a Sy 1 AGN, however this is itself noteworthy given the extreme nature of the outburst. This implies that a ‘standard’ accretion disk can develop very quickly. We discuss possible mechanisms for the outburst, and conclude that the most likely scenario is a disk instability that rapidly and uniformly increases the emission from the AGN. NGC 1566 is an excellent candidate for understanding the changing-look phenomenon, due to its proximity and repeat outburst behaviour.

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REFERENCES

- Alloin D., et al., 1986, *ApJ*, 308, 23
- Arnaud K. A., 1996, in Jacoby G. H., Barnes J., eds, *Astronomical Society of the Pacific Conference Series Vol. 101, Astronomical Data Analysis Software and Systems V*. p. 17
- Boller T., et al., 1997, *MNRAS*, 289, 393
- Breeveld A. A., et al., 2010, *MNRAS*, 406, 1687
- Burrows D. N., et al., 2005, *Space Sci. Rev.*, 120, 165
- Campana S., et al., 2015, *A&A*, 581, A17
- Cutri R. M., et al., 2018, *The Astronomer’s Telegram*, 11913
- Dai X., et al., 2018, *The Astronomer’s Telegram*, 11893
- Dexter J., Begelman M. C., 2018, preprint, ([arXiv:1807.03314](https://arxiv.org/abs/1807.03314))
- Ducci L., et al., 2018, *The Astronomer’s Telegram*, 11754
- García J., et al., 2013, *ApJ*, 768, 146
- García J., et al., 2014, *ApJ*, 782, 76
- Grupe D., et al., 2012, *ApJS*, 199, 28
- Grupe D., et al., 2013, *AJ*, 146, 78
- Grupe D., Komossa S., Saxton R., 2015, *ApJ*, 803, L28
- Grupe D., Komossa S., Schartel N., 2018, *The Astronomer’s Telegram*, 11903
- Guainazzi M., 2002, *MNRAS*, 329, L13
- Hill J. E., et al., 2005, in Siegmund O. H. W., ed., *Proc. SPIE Vol. 5898, UV, X-Ray, and Gamma-Ray Space Instrumentation for Astronomy XIV*. pp 325–340, doi:10.1117/12.618026
- Kaastra J. S., Mewe R., Nieuwenhuijzen H., 1996, in Yamashita K., Watanabe T., eds, *UV and X-ray Spectroscopy of Astrophysical and Laboratory Plasmas*. pp 411–414
- Kalberla P. M. W., et al., 2005, *A&A*, 440, 775
- Kawamuro T., et al., 2013, *ApJ*, 770, 157
- Komossa S., 2017, *Astronomische Nachrichten*, 338, 256
- Komossa S., et al., 2017, in Gomboc A., ed., *IAU Symposium Vol. 324, New Frontiers in Black Hole Astrophysics*. pp 168–171, doi:10.1017/S1743921317001648
- Kuin P., et al., 2018, *The Astronomer’s Telegram*, 11786
- Laha S., et al., 2014, *MNRAS*, 441, 2613
- Lightman A. P., Eardley D. M., 1974, *ApJ*, 187, L1
- Liu J.-F., Bregman J. N., 2005, *ApJS*, 157, 59
- Marinucci A., et al., 2014, *ApJ*, 787, 83
- Matt G., Guainazzi M., Maiolino R., 2003, *MNRAS*, 342, 422
- Middleton M. J., et al., 2016, *MNRAS*, 457, 1568
- Noda H., Done C., 2018, *MNRAS*, 480, 3898
- Oknyansky V. L., et al., 2018a, *MNRAS*, p. sty3133
- Oknyansky V. L., et al., 2018b, *The Astronomer’s Telegram*, 11915
- Parker M. L., et al., 2014, *MNRAS*, 445, 1039
- Parker M. L., et al., 2016, *MNRAS*, 461, 1927
- Parker M. L., Miller J. M., Fabian A. C., 2018, *MNRAS*, 474, 1538
- Penston M. V., Perez E., 1984, *MNRAS*, 211, 33P
- Poole T. S., et al., 2008, *MNRAS*, 383, 627
- Ross N. P., et al., 2018, *MNRAS*, 480, 4468
- Runco J. N., et al., 2016, *ApJ*, 821, 33
- Saxton R. D., et al., 2015, *MNRAS*, 454, 2798
- Schartel N., et al., 2007, *A&A*, 474, 431
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
- Uttley P., Casella P., 2014, *Space Sci. Rev.*, 183, 453
- Vasudevan R. V., et al., 2009, *MNRAS*, 399, 1553
- Wilms J., Allen A., McCray R., 2000, *ApJ*, 542, 914
- Woo J.-H., Urry C. M., 2002, *ApJ*, 579, 530
- Zdziarski A. A., Johnson W. N., Magdziarz P., 1996, *MNRAS*, 283, 193
- Życki P. T., Done C., Smith D. A., 1999, *MNRAS*, 309, 561

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