

# RINGO3 polarimetry of very young ZTF supernovae

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## ABSTRACT

The early phases of the observed evolution of the supernovae (SNe) are expected to be dominated by the shock breakout and ‘flash’ ionization of the surrounding circumstellar medium. This material arises from the last stages of the evolution of the progenitor, such that photometry and spectroscopy of SNe at early times can place vital constraints on the latest and fastest evolutionary phases leading up to stellar death. These signatures are erased by the expansion of the ejecta within  $\sim 5$  d after explosion. Here we present the earliest constraints, to date, on the polarization of 10 transients discovered by the Zwicky Transient Facility (ZTF), between 2018 June and 2019 August. Rapid polarimetric follow-up was conducted using the Liverpool Telescope RINGO3 instrument, including three SNe observed within  $< 1$  d of detection by the ZTF. The limits on the polarization within the first 5 d of explosion, for all SN types, is generally  $< 2$  per cent, implying early asymmetries are limited to axial ratios  $> 0.65$  (assuming an oblate spheroidal configuration). We also present polarimetric observations of the Type I Superluminous SN 2018bsz and Type II SN 2018hna, observed around and after maximum light.

**Key words:** techniques: polarimetric – supernovae: general.

## 1 INTRODUCTION

With the advent of deep, wide-field, and high-cadence surveys, it has been possible to discover new supernovae (SNe) within hours of explosion. Rapid follow-up observations of these SNe has provided a new insight into the stellar origins of these explosion. Early photometric observations of the emergence of the explosion shock has, in a number of cases (e.g. Soderberg et al. 2008; Ofek et al. 2010; Garnavich et al. 2016; Rubin & Gal-Yam 2017; Bersten et al. 2018), revealed a behaviour that cannot be explained by just the shock breaking out of the stellar surface (Waxman & Katz 2017). Instead, the early, rapid rise in brightness can be greatly affected by the presence of circumstellar material (CSM).

For a subset of Type Ia SNe, early photometric observations have revealed an ultraviolet excess that has been interpreted as the shock interaction between the ejecta and a companion star (Cao et al. 2015; Hosseinzadeh et al. 2017). Early spectroscopy has revealed the presence of ‘flash ionized’ species, such as He II and N IV, which correspond to the very quick ionization of the surrounding CSM

(Khazov et al. 2016; Kochanek 2019). In the case of SN 2013cu, early observations by Gal-Yam et al. (2014) were able to reveal the presence of a Wolf–Rayet-like wind immediately around the progenitor (Crowther 2007). Using flash spectroscopy, Yaron et al. (2017) showed that some SN progenitors may exhibit enhanced levels of mass loss as pre-supernova instabilities become important in the final years before explosion. Bruch et al. (2020) report that at least 30 per cent of hydrogen-rich SNe exhibit such features in very early observations, consistent with elevated mass loss just prior to explosion. Despite the power of early-time observations to place important constraints on the nature of the progenitor system, the observable signatures disappear by  $\sim 5$  d, after the ejecta overrun the immediately surrounding dense CSM (Gal-Yam et al. 2014). Early-time observations are therefore very sensitive to the very last phases of stellar evolution.

The early-time optical observations of very young SNe have, predominantly, been in the form of photometry and spectroscopy. Polarimetry has been established as a sensitive probe of the presence of departures from spherical symmetry in SN explosions (Wang & Wheeler 2008). Observations around maximum light and at later epochs, as the photosphere recedes into the ejecta with time, have revealed important clues to the physics of the explosions responsible

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through their imprint on the geometry of the ejecta. In general, core-collapse (CC) SNe show increasing degrees of polarization with time indicating the ejecta becoming more asymmetric the closer to the origin of the explosion. On the other hand, Type Ia SNe exhibit the opposite behaviour appearing to become progressively more spherical closer to origin of the explosion (Wang & Wheeler 2008).

The application of polarimetry to SNe at very early times, however, has been limited, with the earliest spectropolarimetric observation of a Type Ia SN published to date, occurring at only  $\sim 5$  d after explosion (Yang et al. 2020). Given the power of polarimetry to probe the 3D structures of these events, at early times it has the potential to probe the shape of the progenitor system, including the nature of the mass loss prior to explosion. Indeed, as shown by Mauerhan et al. (2014) and Reilly et al. (2017), in the context of SN 2009ip, polarimetry can provide constraints on both the 3D physical characteristics of the explosions and their interaction with the CSM that are not accessible with ordinary photometric and spectroscopic observations.

A major difficulty with conducting rapid polarimetric follow-up of young supernovae is establishing the connection between the discovery surveys and facilities with an appropriate polarimetric observing capability. Here, we report a pathfinder campaign, using the 2.0 m *Liverpool Telescope* (Steele et al. 2004) and the RINGO3 polarimeter (Arnold et al. 2012), to acquire early-time observations of explosive transients; in particular, those discovered by the Zwicky Transient Facility (ZTF; Bellm et al. 2019; Graham et al. 2019; Gal-Yam 2019b). The RINGO3 instrument and its predecessors (Steele et al. 2006, 2010) were designed for the rapid follow-up of the Gamma Ray Burst afterglows (Mundell et al. 2007; Steele et al. 2009; Mundell et al. 2013), exploiting the flexibility to rapidly observe new targets afforded by the robotic nature of the *Liverpool Telescope*. Given the location of ZTF at *Palomar Observatory*, California, USA, it is possible to trigger polarimetric follow-up observations with the *Liverpool Telescope* (La Palma, Canary Islands, Spain) within  $< 24$  h of discovery.

## 2 OBSERVATIONS

### 2.1 Data acquisition & reductions

All observations of the target SNe were conducted with the RINGO3 instrument<sup>1</sup> mounted on the *Liverpool Telescope*. RINGO3 operates with three separate channels, each with its own camera, with the light split by wavelength using two dichroics. The three channels are: ‘d’ covering  $7700 - 10000\text{\AA}$ , ‘e’ covering  $3500 - 6400\text{\AA}$ , and ‘f’ covering  $6500 - 7600\text{\AA}$ ; following the nomenclature of Jermak (2017), we will refer to these band passes as  $r^*$ ,  $b^*$ , and  $g^*$ , respectively. Each camera has a slightly different plate scale: 0.43, 0.44, and 0.49 arcsec for the  $b^*$ ,  $g^*$ , and  $r^*$  bands, respectively. Each RINGO3 channel has its own  $512 \times 512$  px electron multiplying CCD which have, for the type of observations considered here, negligible noise associated with readout.

RINGO3 uses a rotating polaroid ( $\sim 0.4$  Hz) to sample all of the components of the Stokes parameters at eight separate polaroid position angles. Each camera produces, therefore, eight exposures in 2.3 s (i.e. 24 exposures total for the three cameras). The RINGO Data Reduction Pipeline (Arnold 2017)<sup>2</sup> creates a series of mean stacked images: one for the entire duration of the observation and a series

of mean stacked images for each minute of the observation. For this study, we only consider the mean stacked image constructed from all exposures acquired at a given epoch.

The reduced images of the science targets and the zero-polarization and highly-polarized standard calibration stars were retrieved from the *Liverpool Telescope* Archive.<sup>3</sup>

### 2.2 Data analysis workflow

In order to analyse the data, we created a bespoke package to process all the observations and, ultimately, derive the linear Stokes parameters for the science targets. The approach to the analysis follows those presented by Jermak (2017) and Maund et al. (2019).

The data were first sorted into discrete data sets, containing all 24 files corresponding to one individual observation. For each data set and for each camera, source detection was conducted on the image at the first rotor position. These positions were then used to conduct photometry on the images at all eight rotor positions. Due to extreme vignetting for all three cameras, source detection was not conducted in the four corners of the images (corresponding to areas of  $128 \times 128$  px). Aperture photometry was conducted using the PYTHON PHOTUTILS<sup>4</sup> package. For the bright standards, we used a fixed aperture of radius 8 pixels. For the science targets, aperture sizes were selected to match the full width at half-maximum to balance possible contamination from nearby stars or enhanced background (e.g. host galaxy), but maximize the signal-to-noise. In the event that it was not possible for the *Liverpool Telescope* Data Reduction Pipeline to confidently establish the World Coordinate System for each image, the target (or targets) of interest in each image were selected by hand. For all targets, the intensity ( $I$ ) and the normalized linear Stokes parameters ( $q$  and  $u$ ) were calculated from the measured fluxes  $f_i$  at each rotor position  $i$ . Following the prescription of Clarke & Neumayer (2002), the intensities corresponding to each of the three Stokes parameters is given by

$$S_I = \sum_{i=1}^8 f_i$$

$$S_q = f_2 + f_3 + f_6 + f_7$$

$$S_u = f_1 + f_2 + f_5 + f_6 \quad (1)$$

for which the normalized, instrumental Stokes parameters are

$$q_{inst} = \pi \left( \frac{1}{2} - \frac{S_q}{S_I} \right)$$

$$u_{inst} = \pi \left( \frac{S_u}{S_I} - \frac{1}{2} \right). \quad (2)$$

In order to correctly propagate the photometric uncertainties, we used Monte Carlo sampling to create  $N = 10000$  samples from the distribution  $N(f_i, (\Delta f_i)^2)$  and carried these distributions through equations (1) and (2), to derive the corresponding distributions for  $q_{inst}$  and  $u_{inst}$ . Observations of zero and highly-polarized standards were used to remove any instrumental polarization signature. The standard stars observed as part of the *Liverpool Telescope* RINGO3 standard calibration plan are derived from Schmidt, Elston & Lupie (1992). The instrumental polarization offset ( $q_0, u_0$ ) was calculated using the observations of the zero-polarization standard stars, such

<sup>1</sup><https://telescope.livjm.ac.uk/TelInst/Inst/RINGO3/>

<sup>2</sup><https://telescope.livjm.ac.uk/TelInst/Inst/RINGO3/#pipeline>

<sup>3</sup>[https://telescope.livjm.ac.uk/cgi-bin/lt\\_search](https://telescope.livjm.ac.uk/cgi-bin/lt_search)

<sup>4</sup><https://pypi.org/project/photutils/>

**Table 1.** Polarized standards from Schmidt et al. (1992) in the RINGO3 channels.

Standard	$b^*/e'$		$g^*/f'$		$r^*/d'$	
	$p(\text{per cent})$	$\theta (^{\circ})$	$p(\text{per cent})$	$\theta (^{\circ})$	$p(\text{per cent})$	$\theta (^{\circ})$
BD +25 727	5.99	31.2	6.13	31.5	5.22	31.7
BD +59 389	6.40	98.2	6.27	98.2	5.45	98.2
BD +64 106	5.48	96.9	5.00	96.8	4.89	96.7
HD 155528	4.80	91.9	4.80	91.9	4.80	91.8
HD 215806	1.80	66.6	1.74	69.0	1.40	70.8
Hiltner 960	5.61	55.2	4.98	54.3	4.19	53.5
VI Cyg 12	8.42	115.6	8.42	115.6	8.42	115.6

that

$$q' = 1.14(q_{\text{inst}} - q_0)$$

$$u' = u_{\text{inst}} - u_0, \quad (3)$$

where the factor 1.14 corrects for elliptical distortion of the polarization circle of a constant polarization source in the  $qu$  plane (Arnold 2017).

The polarization properties of the highly-polarized standards in the appropriate RINGO3 wavelength channels are given in Table 1, as previously used by Maund et al. (2019). We used reported polarization values for seven standards from the ultraviolet to the infrared (in the *UBVRIJHK* passbands Schmidt et al. 1992), employing a fourth order polynomial to calculate the brightness-weighted polarization over the wavelength ranges corresponding to the three RINGO3 channels.

For highly-polarized stars, the observed polarization angle, in the instrument coordinate frame, is given as

$$\theta_{\text{obs}} = \frac{1}{2} \arctan \left( \frac{u'}{q'} \right) \quad (4)$$

from which the rotation offset of the instrument can be determined by

$$K = \theta_0 - (ROTSKYPA - \theta_{\text{obs}}), \quad (5)$$

where  $\theta_0$  is the previously determined polarization angle for the standard star in the reference catalogue and *ROTSKYPA* is the instrument rotation angle.  $K$  is therefore the relative offset of the polaroid positions, with respect to the standard astronomical definition of the Stokes parameter coordinate system ( $+q$  aligned with North and a polarization angle of  $0^{\circ}$ ), without any rotation of the instrument. The total degree of observed polarization was calculated as

$$p_{\text{obs}} = \sqrt{q'^2 + u'^2}. \quad (6)$$

Through comparison with previously catalogued values of the polarization for the highly-polarized standard stars ( $p_0$ ), the degree of instrumental ‘depolarization’ was derived as

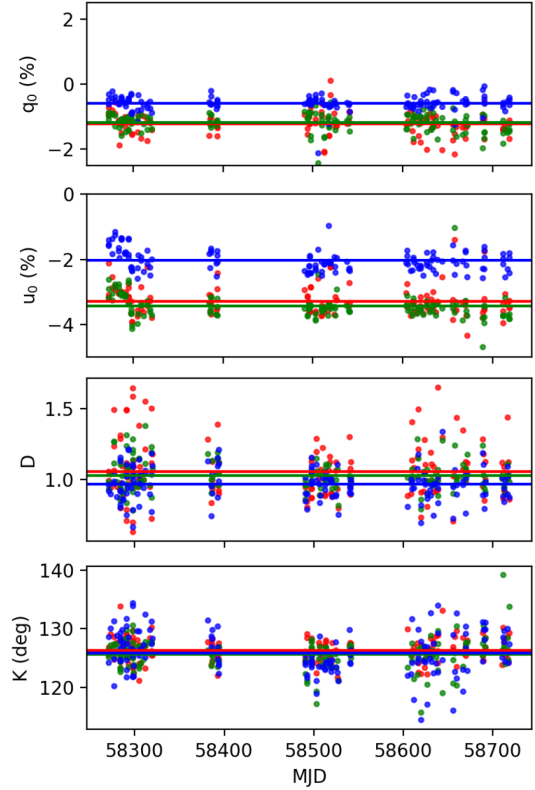
$$D = \frac{p_0}{p_{\text{obs}}}. \quad (7)$$

We note that this definition of the instrumental depolarization is the inverse to that used by Słowikowska et al. (2016); however, the two approaches are otherwise equivalent. For the science targets,  $q'$  and  $u'$  were used to calculate the intrinsic polarization angle

$$\theta_0 = K + ROTSKYPA - \frac{1}{2} \arctan \left( \frac{u'}{q'} \right) \quad (8)$$

and the intrinsic degree of polarization

$$p_0 = D \times p_{\text{obs}}. \quad (9)$$



**Figure 1.** The instrumental polarization parameters derived from each zero- and high-polarization standard observed over the survey period. The points are colour-coded according to wavelength channel (blue =  $b^*$ , green =  $g^*$ , and red =  $r^*$ ). Horizontal lines in each panel indicate the mean derived instrumental parameter (as summarized in Table 2).

The degree of polarization was further corrected for bias using the Modified ASymptotic (MAS) estimator of true polarization  $p_{\text{MAS}}$  (Płazczynski et al. 2014). We follow Higgins et al. (2019) by characterizing polarization measurements with  $p_{\text{MAS}}/\sigma_p < 3$  as non-detections, and quote the 95 per cent upper limit.

### 2.3 The stability of the RINGO3 instrument

To establish a baseline calibration for each science observation we utilized the zero- and high-polarization standards, observed as part of the standard RINGO3 calibration plan, from the night of and the nights before and after the science observation. As a single-beam polarimeter (Arnold et al. 2012), in which orthogonal polarization components are measured in series, instrumental and background effects may be additive and not completely removed through the calculations presented in equations (1) and (2). The determination of the instrumental polarization calibration parameters  $q_0$ ,  $u_0$ ,  $D$ , and  $K$  may also be limited by the level of the sky background, the seeing, and the throughput of each individual channel (Słowikowska et al. 2016). RINGO3 also uses two dichroics (to separate the three separate wavelength channels) and a depolarizing Lyot prism, for which it is estimated the minimum total systematic uncertainty is  $\sim 0.5$  per cent (Jermak 2017).

The measured instrumental polarization parameters, for each polarization standard star observed, are shown in Fig. 1. The mean and standard deviation of the instrumental polarization parameters of RINGO3, for the survey period, are summarized in Table 2. We also calculated the intra-night standard deviation of the instrumental

**Table 2.** Average properties of the RINGO3 instrumental polarization between 2018 June and 2019 August.

Channel		$\overline{q_0}$	$\sigma(q_0)$	$\overline{\sigma_N(q_0)}$	$\overline{u_0}$	$\sigma(u_0)$	$\overline{\sigma_N(u_0)}$	$\overline{D}$	$\sigma(D)$	$\overline{\sigma_N(D)}$	$\overline{K}$	$\sigma(K)$	$\overline{\sigma_N(K)}$
$b^*$	e	-0.58	0.24	0.10	-2.02	0.39	0.15	0.97	0.11	0.07	125.92	3.55	2.19
$g^*$	f	-1.18	0.28	0.14	-3.44	0.38	0.11	1.03	0.10	0.05	125.62	2.70	1.16
$r^*$	d	-1.22	0.40	0.15	-3.28	0.44	0.19	1.06	0.20	0.11	126.38	2.14	1.17

$\sigma$  = scatter (standard deviation) of the observed parameter across all observations.

$\overline{\sigma_N}$  = average scatter of the observed parameter measured on individual nights.

**Table 3.** Science targets for RINGO3 observations.

Target	Original Name	$\alpha_{J2000}$	$\delta_{J2000}$	Discovery date	Redshift	Type	Discoverer
SN 2018bsz	ASASSN-18km	16:09:39.1	-32:03:45.6	2018-05-17	0.027	SLSN-I (Anderson et al. 2018b)	ASASSN <sup>1</sup> (Stanek 2018)
SN 2018cnw	ZTF 18abauprj	16:59:05.0	+47:14:11.2	2018-06-15	0.028	SN Ia-91T-like (Miller 2018)	ZTF <sup>2</sup> (Fremling 2018a)
SN 2018cyg	ZTF 18abdbysy	15:34:08.5	+56:41:48.7	2018-06-30	0.011	SN II (Fremling & Sharma 2018)	LOSS <sup>3</sup> (Jeffers, Zheng & Filippenko 2018)
SN 2018dfi	ZTF 18abfflyqp	16:50:50.1	+45:23:52.5	2018-07-10	0.031	SN IIb (Bruch et al. 2020)	POSS <sup>4</sup> (Gagliano et al. 2018)
SN 2018eay	ZTF 18abgmcmv	18:16:13.1	+55:35:27.2	2018-07-15	0.018	SN Ia-91T-like (Yin et al. 2018)	ZTF <sup>2</sup> (Fremling 2018b)
SN 2018gep	ZTF 18abukavn	16:43:48.2	+41:02:43.4	2018-09-09	0.032	SN Ic-BL (Burke et al. 2018)	ZTF <sup>2</sup> (Ho et al. 2018)
SN 2018gvi	ZTF 18abyxwrf	02:55:36.0	+43:03:27.3	2018-09-24	0.021	SN Ia (Fremling, Dugas & Sharma 2018)	ZTF <sup>2</sup> (Fremling 2018c)
SN 2018hna	...	12:26:12.1	+58:18:50.8	2018-10-22	0.002	SN II (Leadbeater 2018)	K. Itagaki (Itagaki 2018)
SN 2019np	ZTF 19aacslb	10:29:22.0	+29:30:38.3	2019-01-09	0.004	SN Ia (Burke et al. 2019a)	K. Itagaki (Itagaki 2019)
SN 2019ein	ATLAS19ieo	13:53:29.1	+40:16:31.3	2019-05-01	0.008	SN Ia (Burke et al. 2019b)	ATLAS <sup>5</sup> (Tonry et al. 2019)
AT 2019hgp	ZTF 19aayejww	15:36:12.9	+39:44:00.6	2019-06-08	...	...	ZTF <sup>2</sup> (Bruch et al. 2019b)
SN 2019nmv	ZTF 19abqhobb	17:25:38.7	+59:26:48.3	2019-08-19	0.019	SN II (Hiramatsu et al. 2019)	ZTF <sup>2</sup> (Nordin et al. 2019)

<sup>1</sup> All-Sky Automated Survey for Supernovae (ASAS-SN; Shappee et al. 2014); <sup>2</sup> Zwicky Transient Facility (ZTF; Bellm et al. 2019); <sup>3</sup> Lick Observatory Supernova Search (LOSS; Li et al. 2000); <sup>4</sup> Puckett Observatory Supernova Search (POSS); <sup>5</sup> Asteroid Terrestrial-impact Last Alert System (ATLAS; Tonry et al. 2018)

polarization parameters  $\sigma_N$  (and in Table 2 we report the average over all nights). In general, we find that the limiting systematic precision of RINGO3 is consistent with previous estimates (Słowińska et al. 2016; Jermak 2017). Although the calibration of RINGO3 is relatively stable over the period of the survey, there is some structure present in Fig. 1 [e.g. around Modified Julian Date (MJD) 58300, which coincided with the cleaning of a mirror in the optical path] that requires applying calibrations derived over short time-scales (rather than an average derived over the entire length of the survey).

## 2.4 Science targets and observations

Science targets were observed as part of programmes PQ18A02, PL18A10, PL18B01, and PL19A16. Targets were selected, primarily from ZTF, for their brightness and location in the sky to be suitable for RINGO3 observations. The sample is composed of four Type II SNe, three Type Ia SNe, two 1991T-like Type Ia SNe, one Type Ic SN, one Type I superluminous SN (SLSN), and a single transient (AT2019hgp) of unknown classification. Details of the formal discovery and classification of these transients are shown in Table 3. We note, however, that in a number of cases the objects had ZTF observations prior to the date and time given in the formal discovery announcement.

During the ZTF observations, difference images are generated based on the image subtraction algorithm by Zackay, Ofek & Gal-Yam (2016) implemented in the real-time reduction pipeline (Masci et al. 2019). Only ZTF alert streams that are above a  $5\sigma$  threshold will generate alerts. Using the IPAC ZTF difference imaging pipeline, we performed forced point-spread function (PSF) photometry at the location of SNe discovered by ZTF following the procedure described in Yao et al. (2019). We applied a  $4\sigma$  threshold and inspected both the last non-detection limit and the first detection in both  $g$  and  $r$  bandpasses. To establish the phase of our observations, for each transient, we consider the time of explosion  $t_{\text{exp}}$  to be midway between the last ZTF non-detection and the first ZTF detection of the candidate (see Table 4). The estimated explosion time of the Type II SNe 2018cyg and 2018dfi are consistent with Bruch et al. (2020). The

discovery time of SN 2018gep agrees with the first  $r$ -band detection reported by Ho et al. (2019).

A log of the science observations is presented in Table 5 and the locations of the SNe, with respect to their host galaxies, are shown on Fig. 2. It was not possible to observe two of the science targets (SN 2018bsz and 2018hna) at early times, and these constitute outliers from the main targets of the early-time polarimetry survey. As these were observed alongside our other targets, and using the same Liverpool Telescope programmes, we include them here for completeness.

## 3 RESULTS & ANALYSIS

### 3.1 SN 2018bsz

Discovered by ASAS-SN on 2018 May 17 (Stanek 2018), it was temporarily classified as a young Type II SN (Hiramatsu et al. 2018a). It was later reclassified as a superluminous supernova (for a review see Gal-Yam 2019a), albeit a lower luminosity example (Anderson et al. 2018a). Our observations commenced 4 d after the photometric light curve maximum or 69 d after the explosion date proposed by Anderson et al. (2018a). The polarization measurements for SN 2018bsz are presented in Table 6 and the time evolution, with respect to the photometric light curve, is shown on Fig. 3. We derive limits on the polarization, strictest in the blue, at the general level of  $< 1 - 2$  per cent. We do, however, make one single detection of  $p(b^*) = 2 \pm 0.5$  per cent at 11.4 d after maximum (or MJD 58 267.5).

### 3.2 SN 2018cnw

Miller (2018) classified SN 2018cnw as being a ‘91T-like’ Type Ia SN. Five sets of observations were acquired in the course of a single night (around MJD 58 287.0). The observation of the SN was subject to poor seeing of  $\sim 3.2$  arcsec, with the point spread function appearing obviously elongated. We note that it was not possible, under these conditions, to detect the SN in the  $r^*$  observations. For the other filters, we derived upper limits on the polarization of  $b^* < 2.0$  per cent and  $g^* < 0.9$  per cent.



**Table 4.** Non-detection limits and first detection epochs for the observed SNe.

Target	Filter	Last ZTF non-detection (MJD <sup>a</sup> [S/N] <sup>b</sup> )	First ZTF detection (MJD <sup>a</sup> [S/N] <sup>b</sup> )
SN 2018bsz/ASASSN-18km	...	...	...
SN 2018cnw/ZTF18abauprj	<i>g</i>	58 282.385 [ $<0$ ]	58 283.283 [5.1]
	<i>r</i>	58 283.329 [3.2]	58 284.280 [19]
SN 2018cyg/ZTF18abdbysy	<i>g</i>	58 294.223 [ $<0$ ]	58 295.205 [5.2]
	<i>r</i>	58 294.242 [2.1]	58 294.257 [4.0]
SN 2018dfi/ZTF18abffypq	<i>g</i>	58 306.307 [ $<0$ ]	58 307.214 [63]
	<i>r</i>	58 306.201 [ $<0$ ]	58 307.186 [44]
SN 2018eay/ZTF18abgmcmv	<i>g</i>	58 311.345 [3.7] <sup>d</sup>	58 312.350 [11]
	<i>r</i>	58 311.198 [2.5]	58 311.222 [5.5]
SN 2018gep/ZTF18abukavn	<i>g</i>	58 369.254 [3.9] <sup>d</sup>	58 370.186 [19]
	<i>r</i>	58 370.141 [4.0] <sup>d</sup>	58 370.163 [7.3]
SN 2018gvi/ZTF18abyxwrf	<i>g</i>	58 384.319 [2.3]	58 385.413 [6.3]
	<i>r</i>	58 386.328 [ $<0$ ]	58 388.485 [21]
SN 2018hna/ZTF18acbwaxk	...	...	...
SN 2019np/ZTF19aacgslb	<i>g</i>	58 491.454 [ $<0$ ]	58 494.483 [179]
	<i>r</i>	58 491.530 [1.2]	58 492.445 [18]
SN 2019ein/ATLAS19ieo	cyan-ATLAS	58 602.267	58 604.474 <sup>c</sup>
AT 2019hgp/ZTF19aayejww	<i>g</i>	58 640.362 [ $<0$ ]	58 641.201 [3.5] <sup>d</sup>
	<i>r</i>	58 640.291 [ $<0$ ]	58 641.289 [4.9]
SN 2019nvm/ZTF19abqhobb	<i>g</i>	58 713.218 [ $<0$ ]	58 714.163 [71]
	<i>r</i>	58 713.242 [1.2]	58 714.185 [55]

<sup>a</sup>Modified Julian Date; <sup>b</sup>Signal-to-noise ratio for the forced difference image PSF-fit flux measurement;<sup>c</sup>Discovered by ATLAS (Tonry et al. 2019); <sup>d</sup>Target was marginally detected;

### 3.3 SN 2018cyg

SN 2018cyg was observed three times in a single night (MJD 58 296.0). The SN was marginally detected ( $S/N = 4.0$ ) in ZTF *r*-band observation on MJD 58 294.257, followed by the *g*- and *r*-band detections at MJD 58 295.205 ( $S/N = 5.2$ ) and MJD 58 295.246 ( $S/N = 16.0$ ), respectively. We consider the RINGO3 observation was therefore conducted at  $\sim 0.8$ – $1.8$  d after the first detection with ZTF, or  $\approx 2$  d after the explosion. (see Table 4). Fremling & Sharma (2018) later classified it as a Type II SN, with a short plateau. Given the faintness of the SN at this epoch and poor observing conditions, we could not establish a photometric detection of SN 2018cyg in the  $b^*$ -band and were only able to place limits on the degree of polarization of  $g^* < 15$  per cent and  $r^* < 22.0$  per cent.

### 3.4 SN 2018dfi

Two sets of observations of SN 2018dfi, consisting of two separate exposures each, were conducted on a single night (MJD 58 308.0). The RINGO3 observations commenced 0.7 d after the first detection by ZTF, which we estimate to correspond to 1.2 d post-explosion. The SN was classified (at 4.5 d post-explosion) as a Type II SN (Hiramatsu et al. 2018b), which was further refined to being Type IIb (Bruch et al. 2020). Combining the Stokes parameters determined for all four exposures, we derive limits on the degree of polarization of SN 2018dfi of  $p(b^*) < 2.3$  per cent,  $p(g^*) < 6.8$  per cent, and  $p(r^*) < 4$  per cent.

### 3.5 SN 2018eay

Forced PSF photometry of ZTF images at the location of SN 2018eay shows that the  $S/N$  measured in both *g*- ( $S/N = 3.7$  at MJD 58 311.345) and *r*- ( $S/N = 5.5$  at MJD 58 311.222) bands are obviously higher compared to the previous non-detections. The

$S/N$  derived based on an *r*-band image obtained earlier during the same night (MJD 58 311.198) yields 2.5. Therefore, we adopt an explosion epoch at MJD 58 311.2 for SN 2018eay, indicating that the first RINGO3 imaging polarimetry was acquired 7.7 d post the SN explosion (the night of MJD 58 319.0). Similarly to SN 2018cnw, SN 2018eay was classified as a ‘91T-like’ Type Ia SN (Yin et al. 2018). In both sets of observations, the SN was only weakly detected in all three channels and it was only possible to derive limits on the degree of polarization of  $p(b^*) < 2.5$  per cent,  $p(g^*) < 6.1$  per cent, and  $p(r^*) < 7.5$  per cent.

### 3.6 SN 2018gep

Two sets of observations of SN 2018gep were conducted, in sequence, on the night of MJD 58 385.0, corresponding to 14.8 d post-explosion. We note that SN 2018gep was discovered very close to the moment of explosion, with the last ZTF non-detection of the transient occurring only 0.02 d before the first detection (Ho et al. 2018). They conducted a second-order polynomial fit to the first 3 d of the *g*-band flux and defined  $t_0$  at 58 370.146 as the time at which the flux of SN 2018gep is zero. In fact, the transient exhibits pre-explosion emission extended  $\approx 1$  week prior to the rapid rise in the light curve (see Fig. 7 of Ho et al. 2019). An observation of SN 2018gep at 10.1 d yielded a classification for SN 2018gep as a broad-lined Type Ic supernova at around maximum light (Burke et al. 2018); although Pritchard et al. (2020) suggest the fast rise-time ( $< 6.2$  d) may imply that SN 2018gep may be more closely related to the family of Fast Blue Optical Transients (Drout et al. 2014). From our RINGO3 observations, we do not detect any significant polarization for SN 2018gep, instead deriving polarization limits of  $p(b^*) < 1.6$  per cent and  $p(g^*) < 7.0$  per cent and  $p(r^*) < 5.1$  per cent.

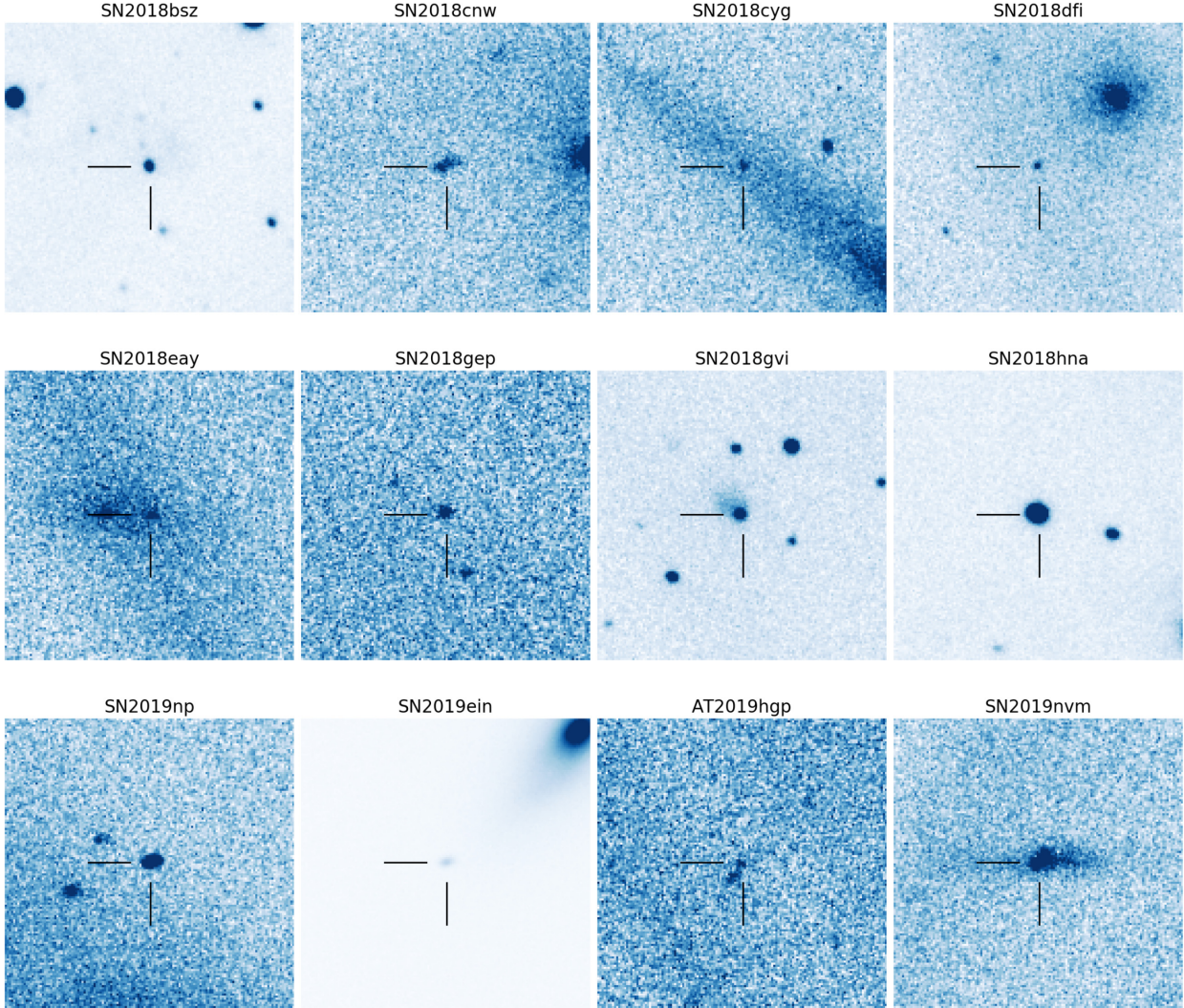
**Table 5.** RINGO3 observations of the target SNe. The exposure time is the total spent on the target across all eight polaroid rotator positions. The same exposure time is used for each of the three ( $b^*/g^*/r^*$ ) RINGO3 channels.

Date (UT)	MJD	Phase (d) <sup>†</sup>	Data set	Target	Exposure Time(s)
2018-06-02 23:02	58 271.96	69.5	e_20180602_2.0	SN 2018bsz	598
2018-06-05 22:53	58 274.95	72.4	e_20180605_8.0	SN 2018bsz	598
2018-06-10 22:37	58 279.94	77.4	e_20180610_4.0	SN 2018bsz	895
2018-06-13 22:58	58 282.96	80.5	e_20180613_4.0	SN 2018bsz	895
2018-06-17 22:05	58 286.92	84.4	e_20180617_3.0	SN 2018bsz	1196
2018-06-20 22:01	58 289.92	87.4	e_20180620_4.0	SN 2018bsz	1197
2018-06-24 21:35	58 293.90	91.4	e_20180624_1.0	SN 2018bsz	1196
2018-06-28 21:33	58 297.90	95.4	e_20180628_3.0	SN 2018bsz	1198
2018-06-18 02:00	58 287.08	4.2	e_20180617_10.0	SN 2018cnw	948
2018-06-18 02:16	58 287.09	4.3	e_20180617_11.0	SN 2018cnw	948
2018-06-18 02:32	58 287.11	4.3	e_20180617_12.0	SN 2018cnw	948
2018-06-18 02:48	58 287.12	4.3	e_20180617_13.0	SN 2018cnw	945
2018-06-18 03:04	58 287.13	4.3	e_20180617_14.0	SN 2018cnw	945
2018-06-26 23:00	58 295.96	1.7	e_20180626_3.0	SN 2018cyg	1799
2018-06-26 23:30	58 295.98	1.7	e_20180626_4.0	SN 2018cyg	1796
2018-06-27 00:01	58 296.00	1.8	e_20180626_5.0	SN 2018cyg.	1796
2018-07-08 22:12	58 307.93	1.2	e_20180708_5.0	SN 2018dfi	1196
2018-07-08 22:32	58 307.94	1.2	e_20180708_6.0	SN 2018dfi	1196
2018-07-09 01:37	58 308.07	1.3	e_20180708_7.0	SN 2018dfi	1196
2018-07-09 01:58	58 308.08	1.3	e_20180708_8.0	SN 2018dfi	1196
2018-07-19 22:12	58 318.93	7.7	e_20180719_5.0	SN 2018eay	996
2018-07-19 22:29	58 318.94	7.7	e_20180719_6.0	SN 2018eay	999
2018-09-23 22:15	58 384.93	14.8	e_20180923_16.0	SN 2018gep	446
2018-09-23 22:23	58 384.93	14.8	e_20180923_17.0	SN 2018gep	446
2018-10-02 02:39	58 393.11	8.2	e_20181001_10.0	SN 2018gvi	1199
2019-01-10 05:07	58 493.21	82.4	e_20190109_4.0	SN 2018hna	476
2019-01-14 02:09	58 497.09	86.3	e_20190113_9.0	SN 2018hna	478
2019-01-20 01:44	58 503.07	92.3	e_20190119_7.0	SN 2018hna	537
2019-01-30 03:35	58 513.15	102.3	e_20190129_5.0	SN 2018hna	715
2019-02-04 02:54	58 518.12	107.3	e_20190203_5.0	SN 2018hna	957
2019-02-12 01:45	58 526.07	115.3	e_20190211_7.0	SN 2018hna	997
2019-01-11 03:46	58 494.16	2.2	e_20190110_4.0	SN 2019np	898
2019-01-12 03:48	58 495.16	3.2	e_20190111_4.0	SN 2019np	1197
2019-01-12 04:35	58 495.19	3.2	e_20190111_5.0	SN 2019np	1197
2019-01-13 01:13	58 496.05	4.1	e_20190112_9.0	SN 2019np	1198
2019-01-14 01:50	58 497.08	5.1	e_20190113_8.0	SN 2019np	998
2019-01-15 02:34	58 498.11	6.1	e_20190114_20.0	SN 2019np	717
2019-01-16 06:40	58 499.28	7.3	e_20190115_10.0	SN 2019np	598
2019-01-20 01:35	58 503.07	11.1	e_20190119_6.0	SN 2019np	417
2019-01-21 01:40	58 504.07	12.1	e_20190120_3.0	SN 2019np	417
2019-01-23 01:33	58 506.06	14.1	e_20190122_9.0	SN 2019np	717
2019-01-26 04:55	58 509.21	17.2	e_20190125_3.0	SN 2019np	717
2019-02-07 01:15	58 521.05	29.1	e_20190206_9.0	SN 2019np	996
2019-02-25 00:11	58 539.01	47.0	e_20190224_3.0	SN 2019np	1196
2019-05-03 21:27	58 606.89	3.5	e_20190503_3.0	SN 2019ein	1499
2019-05-04 21:46	58 607.91	4.5	e_20190504_3.0	SN 2019ein	1497
2019-05-04 22:13	58 607.93	4.6	e_20190504_4.0	SN 2019ein	1497
2019-05-05 21:32	58 608.90	5.5	e_20190505_3.0	SN 2019ein	1497
2019-05-06 21:56	58 609.91	6.5	e_20190506_3.0	SN 2019ein	1496
2019-05-07 22:47	58 610.95	7.6	e_20190507_3.0	SN 2019ein	896
2019-05-09 22:47	58 612.95	9.6	e_20190509_3.0	SN 2019ein	599
2019-05-11 00:04	58 614.00	10.6	e_20190510_3.0	SN 2019ein	598
2019-05-11 23:03	58 614.96	11.6	e_20190511_5.0	SN 2019ein	598
2019-05-14 00:45	58 617.03	13.7	e_20190513_26.0	SN 2019ein	598
2019-05-16 23:03	58 619.96	16.6	e_20190516_4.0	SN 2019ein	599
2019-05-19 21:40	58 622.90	19.5	e_20190519_1.0	SN 2019ein	598
2019-05-22 22:30	58 625.94	22.6	e_20190522_1.0	SN 2019ein	598
2019-05-27 22:20	58 630.93	27.6	e_20190527_3.0	SN 2019ein	596
2019-06-04 23:10	58 638.97	35.6	e_20190604_3.0	SN 2019ein	598
2019-06-24 23:24	58 658.98	55.6	e_20190624_8.0	SN 2019ein	1196
2019-07-05 22:23	58 669.93	66.6	e_20190705_3.0	SN 2019ein	1797

**Table 5** – *continued*

Date (UT)	MJD	Phase (d) <sup>†</sup>	Data set	Target	Exposure Time(s)
2019-07-26 21:09	58 690.88	87.5	e_20190726_3_0	SN 2019ein	1798
2019-06-09 22:56	58 643.96	3.2	e_20190609_3_0	AT 2019hgp	1797
2019-08-19 21:32	58 714.90	1.2	e_20190819_4_0	SN 2019nvm	1797

<sup>†</sup> Relative to calculated explosion epoch (see Section 2.4 and Table 4).



**Figure 2.** RINGO3  $g^*$  images of the science targets. All images are centred on the science target, have dimensions of  $1' \times 1'$ , and oriented with North up and East to the left.

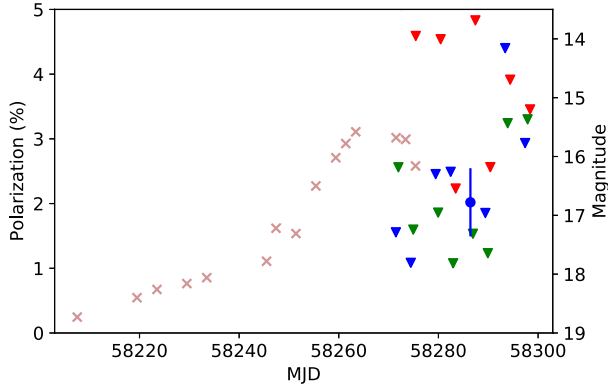
**Table 6.** RINGO3 polarization measurements of SN 2018bsz.

Epoch (MJD)	Phase (d)	$p(b^*)$ (per cent)	$p(g^*)$ (per cent)	$p(r^*)$ (per cent)
58 271.96	69.5	<1.56	<2.56	<5.07
58 274.95	72.4	<1.08	<1.60	<4.59
58 279.94	77.4	<2.46	<1.86	<4.54
58 282.96	80.5	<2.49	<1.08	<2.23
58 286.92	84.4	$2.02 \pm 0.53$	<1.53	<4.83
58 289.92	87.4	<1.85	<1.23	<2.56

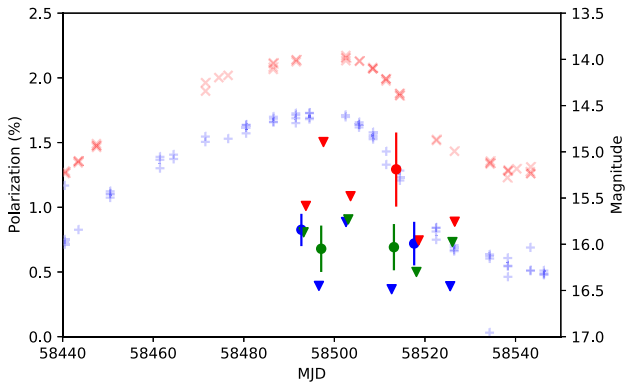
### 3.7 SN 2018gvi

We acquired a single observation of SN 2018gvi at 8.2 d post-explosion, or 7.7 d after the first detection by ZTF. The SN had earlier been classified by Fremling et al. (2018) as a Type Ia SN. We derive upper limits on the polarization of SN 2018gvi in all three channels to levels of  $p(b^*) < 1.6$  per cent,  $p(g^*) < 2.0$  per cent, and  $p(r^*) < 5.0$  per cent.





**Figure 3.** The evolution of the polarization of SN 2018bsz. The measurements are colour coded according to the RINGO3 channels  $b^*$  (blue),  $g^*$  (green), and  $r^*$  (red) and are composed of detections ( $\bullet$ ) and upper limits ( $\blacktriangledown$ ). Also shown is the ATLAS  $o$ -band photometry ( $\times$ ) reported by Anderson et al. (2018a).



**Figure 4.** The evolution of the polarization of SN 2018hna (see Table 7) using the same plotting scheme as Fig. 3. Also shown is ZTF  $g'$  (+) and  $r'$  ( $\times$ ) photometry.

**Table 7.** RINGO3 polarization measurements of SN 2018hna.

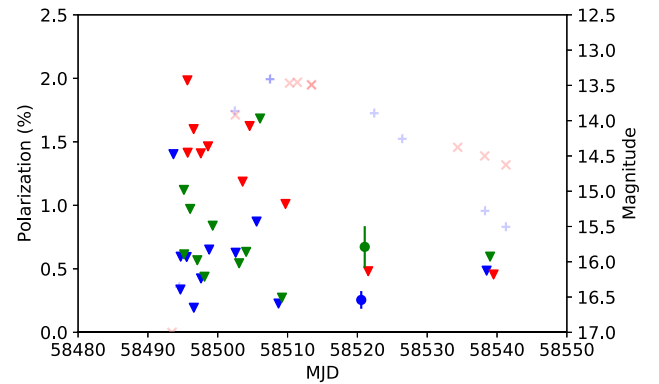
Epoch (MJD)	Phase (d)	$p(b^*)$ (per cent)	$p(g^*)$ (per cent)	$p(r^*)$ (per cent)
58 493.21	82.4	$0.83^{+0.12}_{-0.13}$	$<0.81$	$<1.01$
58 497.09	86.3	$<0.39$	$0.68 \pm 0.18$	$<1.50$
58 503.07	92.3	$<0.89$	$<0.91$	$<1.09$
58 513.15	102.3	$<0.37$	$0.69 \pm 0.18$	$1.29^{+0.28}_{-0.29}$
58 518.12	107.3	$0.72 \pm 0.17$	$<0.50$	$<0.74$
58 526.07	115.3	$<0.39$	$<0.73$	$<0.89$

### 3.8 SN 2018hna

SN 2018hna was observed with RINGO3 at six separate epochs, beginning  $\sim 5$  d before the  $V$ -band maximum. Given the long rise time to maximum light, the observations commenced  $\sim 82$  d post-explosion as shown in Fig. 4<sup>5</sup> (Singh et al. 2019). The polarization measurements at each of the six epochs are listed in Table 7. Overall, we constrain the polarization of SN 2018hna at around maximum light to be  $< 1.5$  per cent. In the  $b^*$ - and  $g^*$ -bands, we make four separate detections of significant polarization  $p \sim 0.5 - 1.0$  per cent, whilst at a single epoch we find  $p(r^*) = 1.3 \pm 0.3$  per cent. From Fig. 4, it

**Table 8.** RINGO3 polarization measurements of SN 2019np.

Epoch (MJD)	Phase (d)	$p(b^*)$ (per cent)	$p(g^*)$ (per cent)	$p(r^*)$ (per cent)
58 494.16	2.2	$<1.40$	$<3.01$	$<4.70$
58 495.16	3.2	$<0.34$	$<1.12$	$<1.98$
58 495.19	3.2	$<0.60$	$<0.61$	$<1.41$
58 496.05	4.1	$<0.59$	$<0.97$	$<1.60$
58 497.08	5.1	$<0.19$	$<0.57$	$<1.41$
58 498.11	6.1	$<0.42$	$<0.44$	$<1.46$
58 499.28	7.3	$<0.65$	$<0.84$	$<3.87$
58 503.07	11.1	$<0.63$	$<0.54$	$<1.19$
58 504.07	12.1	$<2.40$	$<0.63$	$<1.62$
58 506.06	14.1	$<0.87$	$<1.68$	$<2.40$
58 509.21	17.2	$<0.23$	$<0.27$	$<1.01$
58 521.05	29.1	$0.26 \pm 0.07$	$0.67 \pm 0.16$	$<0.48$
58 539.01	47.0	$<0.49$	$<0.60$	$<0.46$



**Figure 5.** The evolution of the polarization of SN 2019np (see Table 8) using the same plotting scheme as Fig. 3. Also shown is ZTF  $g'$  (+) and  $r'$  ( $\times$ ) photometry.

is clear that any polarization associated with SN 2018hna, despite its brightness ( $m_{\max}(r) \sim 14.0$  mags), is close to the systematic floor of the RINGO3 instrument.

### 3.9 SN 2019np

Polarimetric follow-up of SN 2019np commenced 2.2 d after explosion, or 1.7 d after the first detection by ZTF. In total there were 13 separate observations of the SN up to 47 d post-explosion. The polarization measurements for SN 2019np are presented in Table 8 and shown in Fig. 5 (alongside ZTF photometry<sup>6</sup>). The SN was discovered on the rise up to maximum light and polarization constraints, in particular in the blue, limit the polarization across the optical wavelength range to  $< 2.0$  per cent. At a later epoch, 28 d after discovery and 10 d after maximum, we detect significant polarization at the level of  $p(b^*) = 0.26 \pm 0.07$  per cent and  $p(g^*) = 0.67 \pm 0.16$  per cent consistent with the earlier limits on the polarization and the general level of polarization of this SN being low.

### 3.10 SN 2019ein

RINGO3 observations of SN 2019ein started 3.5 d post-explosion or 2.4 d after the first ZTF detection. The Type Ia SN (Burke et al. 2019b) was heavily observed in the rise to maximum, and in general

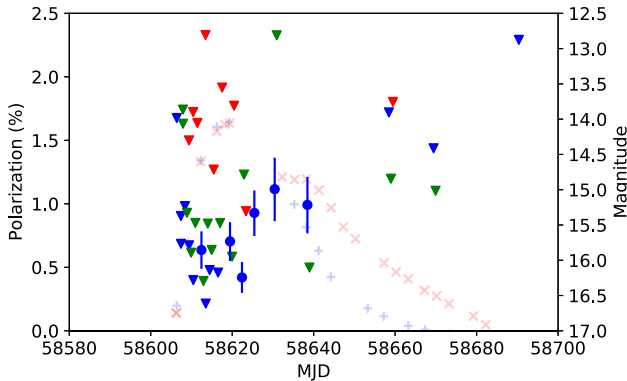
<sup>5</sup><https://lasair.roe.ac.uk/object/ZTF18acbwaxk/>

<sup>6</sup><https://lasair.roe.ac.uk/object/ZTF19aacgslb/>



**Table 9.** RINGO3 polarization measurements of SN 2019ein.

Epoch (MJD)	Phase (d)	$p(b^*)$ (per cent)	$p(g^*)$ (per cent)	$p(r^*)$ (%)
58 606.89	3.5	<1.68	<2.61	<8.55
58 607.91	4.5	<0.90	<1.74	<5.31
58 607.93	4.6	<0.68	<1.63	<3.99
58 608.90	5.5	<0.98	<0.93	<1.50
58 609.91	6.5	<0.67	<0.62	<1.72
58 610.95	7.6	<0.40	<0.85	<1.64
58 612.95	9.6	$0.64 \pm 0.15$	<0.39	<2.33
58 614.00	10.6	<0.21	<0.84	<2.54
58 614.96	11.6	<0.48	<0.64	<1.27
58 617.03	13.7	<0.46	<0.85	<1.92
58 619.96	16.6	$0.70 \pm 0.15$	<0.58	<1.77
58 622.90	19.5	$0.42 \pm 0.12$	<1.23	<0.94
58 625.94	22.6	$0.93 \pm 0.18$	<3.15	<6.83
58 630.93	27.6	$1.12 \pm 0.25$	<2.33	<4.64
58 638.97	35.6	$0.99 \pm 0.22$	<0.50	<8.31
58 658.98	55.6	<1.72	<1.20	<1.80
58 669.93	66.6	<1.44	<1.10	<2.75
58 690.88	87.5	<2.29	<2.74	<11.17

**Figure 6.** The evolution of the polarization of SN 2019ein (see Table 9) using the same plotting scheme as Fig. 3. Also shown is the ZTF  $g'$  (+) and  $r'$  (x) photometry.

we were only able to assess upper limits on the degree of polarization; however, we did measure significant levels of polarization in the  $b^*$ -band (see Table 9) that appear to increase with time around the period of the second light curve maximum that was observable at redder wavelengths (see Fig. 6; in conjunction with ZTF photometry<sup>7</sup>).

### 3.11 AT 2019hgp

AT2019hgp was discovered as a young transient at MJD 58 642.242 (Bruch et al. 2019a). An  $r$ -band detection at MJD 58 641.289 ( $S/N=4.9$ ) and a  $g$ -band signal at MJD 58 641.201 ( $S/N=3.5$ ) were recovered by forced PSF photometry. An early spectrum obtained at MJD 58 642.43 revealed emission lines of highly ionized species (Bruch et al. 2019a), suggesting that the candidate was a young and hot transient. We note, however, that no further specific classification of this transient has been recorded. Our RINGO3 observation was conducted 2.8 d after discovery or  $\sim 3.2$  d post-explosion. The transient was only photometrically detected at significant levels in the

$b^*$  observation, for which we derive an upper limit on the polarization of  $p(b^*) < 5.8$  per cent.

### 3.12 SN 2019nm

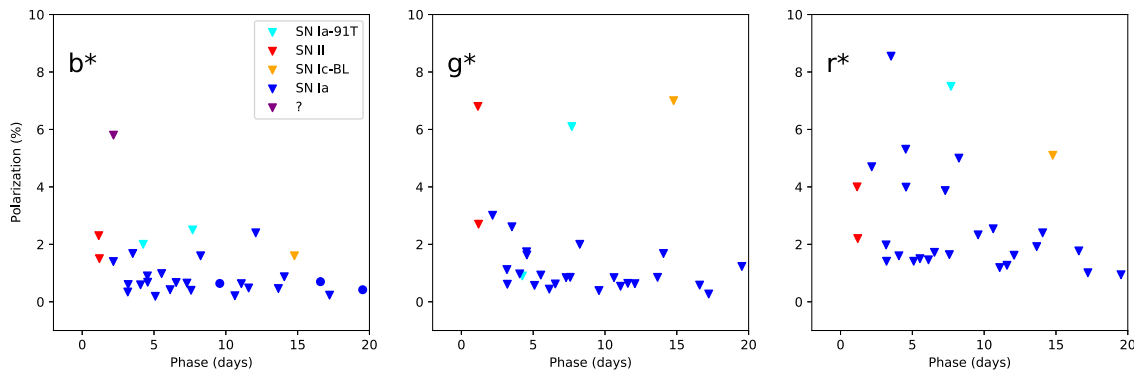
The single RINGO3 observation of SN 2019nm commenced 0.7 d after the SN was first detected by ZTF, or  $\sim 1.2$  d post-explosion. A spectrum, acquired  $\sim 16$  h after the RINGO3 observation revealed a young Type II SN showing ‘flash features’ (Yaron et al. 2017) of He II  $\lambda 4686$  and N IV  $\lambda 4537$  (Hiramatsu et al. 2019). This suggests that SN 2019nm was discovered very soon after explosion, potentially making this RINGO3 polarimetric observation the earliest ever acquired for a Type II SN. The observations were, however, conducted under poor seeing conditions ( $\approx 2.7$  arcsec), with an elongated point spread function possibly indicating the effect of wind on the telescope. SN 2019nm is located close to the nucleus in the edge-on galaxy UGC 10858 and, given the seeing conditions, it was not possible to accurately separate the SN and the host galaxy. For all three channels, we do not significantly detect a polarization signal, with upper limits on the degree of polarization of  $p(b^*) < 1.5$  per cent,  $p(g^*) < 2.7$  per cent, and  $p(r^*) < 2.2$  per cent.

## 4 DISCUSSION & CONCLUSIONS

For the SNe observed at early times we have, in general, only been able to place upper limits on the possible polarization. The evolution of the polarization constraints for all the early time observations, within 20 d of explosion, are shown in Fig. 7. The degree of the constraint on the early-time polarization is limited by two key factors: the brightness of the SN and the relatively high level of the systematic floor of the RINGO3 instrument (Słowikowska et al. 2016). In general, from our observations, the limits on the instrumental polarization means that the upper limits on the polarization are relatively high  $\sim 1$  per cent. As time progresses, we note that the polarization limits do become better, and this is correlated with increased levels of signal-to-noise as the target SNe rise to maximum light. In addition, we find that the throughput in the  $b^*$  channel is the best of the three RINGO3 channels and, from Fig. 7, the systematic floor for RINGO3 does increase towards the red. The  $r^*$ -band polarization limits are less constraining by a factor of  $\sim 2$ .

The constraints on the level of polarization in the  $b^*$ - and  $g^*$ -bands are around  $< 1$  per cent for Type Ia SNe. From Höflich (1991), we can place a constraint on the axial symmetry of the ejecta of Type Ia SNe at early times (assuming a spheroidal configuration) of  $> 0.9$  at  $t \sim 3$  d. These limits are consistent with the earliest spectropolarimetric observation for a Type Ia SN (2018gv) that exhibited a low level of continuum polarization ( $\lesssim 0.2$  per cent) at  $\sim 5$  d post-explosion (Yang et al. 2020). These limits imply that Type Ia SNe generally appear almost, if not completely, spherically symmetric at the earliest times. For the small number of CC SNe in this sample, the earliest constraints available could allow for the presence of significant asymmetries in the first few days; however, for these SNe the lack of subsequent follow-up observations at later epochs (see Fig. 7), as the SNe get brighter, means it is not possible to establish a baseline level of polarization [and some measure of the constant interstellar polarization (ISP) components] at maximum light as was done, for example, with SN 2019ein. For the Type II SNe, the early limits, in particular in the  $b^*$ -band, constrain the axial ratio to be  $> 0.65$  within  $\sim 1$  d of explosion. We note that these figures do not include corrections for the ISP, a constant source of polarization, across all epochs, arising from intervening dust along the line of sight that is independent of the evolution of the polarization of the SNe

<sup>7</sup><https://lasair.roe.ac.uk/object/ZTF19aatlmbo/>



**Figure 7.** All polarization measurements for all 10 transients with early-time observations for the first 20 d after the assumed epoch of explosion. Detections and limits on the polarization are colour-coded according to the type of the transient (according to the legend in the left-hand panel).

themselves (Serkowski 1973). We would expect, more generally, that the constraints on the intrinsic polarization of the SNe would be lower, if a correction of the ISP could be applied.

Given the limited number of significant detections of polarization across the whole sample, the Type Ia SN 2019ein stands out as having a series of detections in the  $b^*$ -band with a possible increasing degree of polarization (see Fig. 6), reaching a maximum at around the second light curve peak observed in the ZTF  $r'$ -band. SN 2019ein has the highest recorded expansion velocities for a Type Ia SN at early times (Kawabata et al. 2020; Pellegrino et al. 2020). SN 2019ein also exhibited blue-shifted line profiles in early spectra (at -14 d relative to B-band maximum light), in particular for the strong Si II feature. The expansion velocities decreased as the SN approached the B-band light-curve maximum. Pellegrino et al. (2020) interpreted the peculiar high expansion velocities as indicative of an asymmetric off-centre explosion, in which intermediate mass elements were mixed into the highest velocity portions of the ejecta. The largest measured polarization of  $p(b^*) = 1.12 \pm 0.25$  per cent occurred  $\sim 11$  d after the B-band light-curve maximum. This measurement supports the presence of significant asymmetries in SN 2019ein, as suspected from the earlier spectroscopic observations; however, the time delay may indicate that the peculiar expansion velocities and the polarization are measuring the asymmetry in different ways. As we are conducting broad-band polarimetry, however, we are not sensitive to the key diagnostic feature of Type Ia SN asymmetries, namely the polarization profile of the strong Si II line (Wang, Baade & Patat 2007; Maund et al. 2010; Cikota et al. 2019). The lower polarization detections for SN 2019np ( $< 1$  per cent; see Fig. 5) are more generally consistent with the low levels of intrinsic continuum polarization usually seen for Type Ia SNe (Wang & Wheeler 2008).

SN2018bsz joins a small, but growing collection of SLSNe with multi-epoch polarimetric observations (Leloudas et al. 2015, 2017; Inserra et al. 2016; Maund et al. 2019). In general, Type I SLSNe are noted for having low levels of polarization; however, SN 2015bn exhibited a significant rise in polarization after  $\sim 20$  d post-maximum (Leloudas et al. 2017), with the evolution from the pre-maximum to post-maximum state clearly evident in spectropolarimetric observations (Inserra et al. 2016). Our detection of significant polarization for SN 2018bsz also occurs  $\sim 20$  d (rest-frame) after the time of the V-band light-curve maximum (Anderson et al. 2018a). Inserra et al. (2016) and Leloudas et al. (2017) both explained the behaviour of SN 2015bn as being due to a fundamental change in the asymmetry of the ejecta, with early-time emission arising from an almost spherical outer layer, whilst at later times the emission arise from a more aspherical interior (giving rise to the increase in polarization with

time). Due to the limitations of RINGO3, we were only able to obtain single detection of polarization at  $\sim 2$  per cent. At a similar epoch in the evolution of SN 2015bn, Leloudas et al. measured  $\sim 1$  per cent (although the degree of polarization later rose to  $\sim 1.54$  per cent by  $\sim 46$  d). It is interesting to note that Inserra et al. and Leloudas et al. showed that, although the interior of 2015bn was more aspherical than the outer layers, the orientation of the asymmetry (in the plane of the sky) was the same at both early (before the light-curve maximum) and later times (after the light-curve maximum). Future dense time series of polarimetric observations of Type I SLSNe will be able to confirm if the rise in polarization, coupled with the stability of the polarization angle, is a common feature for this class of SLSNe.

SN 2018hna appeared, photometrically and spectroscopically, similar to SN 1987A, with the light-curve peaks occurring at  $\sim 87.5$  and 86 d post-explosion, respectively (Singh et al. 2019). Our RINGO3 observations straddle the light-curve peak and the level of polarization observed ( $\sim 0.7$  per cent) is slightly higher than seen for SN 1987A ( $\sim 0.4$  per cent) at similar epochs (Jeffrey 1991). Given the brightness of SN 2018hna, if we had been able to conduct earlier observations of this SN the RINGO3 observation would have been sufficiently sensitive to detect the presence of a similar peaks in the polarization that were seen for SN 1987A (Jeffrey 1991), if they occurred.

The observation of SN 2019nvm is one of the earliest polarimetric observations of a SN, potentially just beating the first observation by Cropper et al. (1987) of SN 1987A. Leonard et al. (2000) reported early spectropolarimetry of the Type II SN 1998S 5 d after discovery (corresponding to  $\sim 5 - 11$  d post-explosion), reporting levels of continuum polarization  $\sim 2$  per cent. Unlike the early emission-line features of SN 2013cu (Gal-Yam et al. 2014), the spectrum of SN1998S persisted for upto 14 d post-discovery, before cooling and developing the classical P Cygni profiles of a Type II SN (Fassia et al. 2001; Wang et al. 2001). Although the spectral evolution of SN 1998S occurred over time-scales of weeks, rather than days normally associated with ‘flash’ observations (Gal-Yam et al. 2014; Yaron et al. 2017; Soumagnac et al. 2019), it does demonstrate that large levels of polarization, due to presence of an aspherical CSM, could be observed at very early epochs.

This RINGO3 survey has crucially demonstrated the feasibility of employing polarimeters on robotic telescopes and, coupled with the appropriate feeder survey, the potential for ‘flash polarimetry’. While it is expected that the CSM signature at early times would be erased by the expansion of the ejecta, in seven cases we have been able to successfully observe targets within this 5-d window.

On two occasions, we have been able to trigger *Liverpool Telescope* observations within a day of discovery by ZTF.

The quality of the data presented in this paper has been limited by instrumental errors caused by the single-beam design of RINGO3 making the cancellation of systematic errors difficult. RINGO3 was decommissioned on the *Liverpool Telescope* in 2020 January and replaced with a prototype of a new dual beam polarimeter (MOPTOP). This uses a dual sCMOS imaging system to record the ordinary and extra-ordinary rays from a polarizing beam-splitter (Jermak, Steele & Smith 2016, 2018). It therefore has higher throughput than the polaroid-based RINGO3 as well as allowing differential cancellation of polarization errors. Commissioning observations with MOPTOP (Shrestha et al. 2020) show uncorrected systematic errors reduced to <0.2 per cent and the sensitivity increased by a factor  $\sim 4\times$  compared to RINGO3. These combined improvements mean that MOPTOP has a polarization accuracy of <0.3 per cent for a source with  $R = 17$  in a 600-s exposure. From our sample, five targets would be observable to this polarization precision, given that exposure time, at the earliest epochs; including SNe 2018hna and 2019ein. MOPTOP will make tighter constraints or even detections of the polarization feasible for the types of targets we have observed so far with RINGO3.

The improved sensitivity of the MOPTOP instrument and the capability to observe in four wavelength bands, from the optical to the near-infrared, will potentially provide a better handle on the ISP. As demonstrated for SN 2019ein, changes in the level of polarization are perceptible, and relative changes in intrinsic polarization can be directly measured. A complete correction for the ISP could be derived, including inferring the wavelength dependence, if observations continued into later phases when a SN might be considered intrinsically unpolarized (see e.g. Jeffrey 1991) when, despite its faintness, it would still be accessible to MOPTOP.

Despite the constraints placed on the early-time explosion geometries with this RINGO3 programme, the limited sample size and only a limited number of detections of significant polarization means that ‘flash polarimetry’ is still *terra incognita*. This survey has demonstrated it is possible to conduct these types of observations with very fast turnaround times and, with the advent of MOPTOP on the *Liverpool Telescope*, it will soon be possible to directly and systematically measure the polarization of SNe at the earliest times.

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## DATA AVAILABILITY

The observational data presented here is available in the public archive of the *Liverpool Telescope* ([https://telescope.livjm.ac.uk/cgi-bin/lit\\_search](https://telescope.livjm.ac.uk/cgi-bin/lit_search))

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